

"Per aspera ad astra..."

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Introduction to the stars in high energy

Let your imagination soar. When sitting on the old rocker looking at the sky with a couple of good old whiskey you can easily start thinking about the universe. There are myriads of different kinds of cosmic objects, but what you see in the night sky are predominantly stars. Indeed, almost all the shiny dots in the sky are stars, most of them being exceptionally bright and close - just a tiny speck of the galaxy's total population. You may be able to spot a handful of other galaxies that are bright enough to be seen by the naked eye, however, all the exotic objects that we are imagining remain elusive. They are too faint to be observed easily, because they are not only far, far away, but they also usually shine on different wavelengths, not visible by the human eye.

Think about the distances in the universe. One of the most accurate explanations is the one from: Adams (1979) *"Space," it says, "is big. Really big. You just won't believe how vastly, hugely, mindbogglingly big it is. I mean, you may think it's a long way down the road to the chemist's, but that's just peanuts to space..."*

Consider this, sometimes you want to study processes in these extreme, very faint objects, but they are too faint and too far out in the universe. You are looking for a "laboratory" with similar processes, but located much closer to the observer. The X-ray binary stars can be considered as this kind of laboratories.

There are, of course, many interesting phenomena which could be studied in X-ray binaries or in non-binary X-ray stars. Several of them are mentioned in the motivation section.

I am mentioning about several various space objects in this work, but the main effort is made to study the post-shock region in the Intermediate Polars (IPs).

1.1 Motivation

We can easily find many reasons why to study stars in the high energy bands. We can consider the direct and the most common scientific applications like observations of the supernovae, black holes and neutron stars in X-ray binaries. But for the education purposes I prefer several other, very nice examples closer to the topic of this work.

- **Relativistic jet phenomena:** like it was proposed by Mirabel (2002) that the universal mechanism should be at work in all the relativistic jet sources in the universe. Better understanding of sources including: microblazars, AGNs and gamma-ray burst will help to gain a more comprehensive understanding of these phenomena. Microblazars can play a role of “space laboratories”, where interesting processes last on different timescales as is the case with AGNs or GRBs.

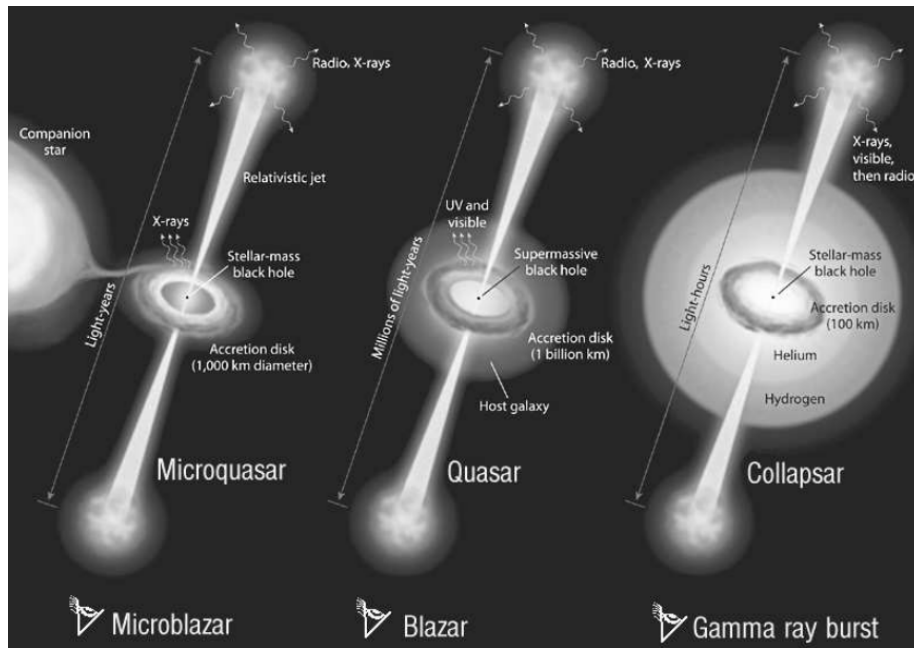


Figure 1.1: NOT in scale diagram, showing current ideas of micro-quasars, AGNs and gamma-ray bursts as space objects driven by the same, universal mechanism Mirabel (2002).

- **Galactic ridge X-ray emission (GRXE):** various physical processes contribute to brightness of GRXE in different bands, but several studies in 3-20 keV provide evidence that the diffuse X-ray radiation originates from a huge number of stellar X-ray sources, mostly coronally active stars and white dwarf X-ray binaries. In particular for the energies over 20 keV to 200 keV,

the spectrum is very similar to the spectrum of magnetic white dwarf binaries – e.g. Intermediate polars (IP) and polars (P). Krivonos et al. (2007)

- **White dwarfs' masses in Intermediate Polars (IP):** as was proposed in Rothschild et al. (1981), the temperature of the post shock region (PSR) depends on WD mass. Therefore the X-ray spectrum can be used for WD mass determination Suleimanov et al. (2005). The WD mass estimations in cataclysmic stars is in general complicated. Usually, the curve of radiation velocities can be used, but it is quite hard to construct. Therefore the X-ray spectrum method is very attractive for several reasons. This work is dedicated to this topic.

1.2 Aim of this work

To cover the whole topic: “stars in high energies” is far behind a capacity of a master thesis and because of that I have decided to concentrate on cataclysmic variable stars (CVs), especially on intermediate polars (IPs).

As it will be mentioned in the next sections closely, IPs are magnetized CVs where the compact, primary star is a white dwarf with $B \sim 10^6 - 10^7$ Gauss. The mass accretion is taking place from, mainly a low-mass, non-degenerate star through its Roche lobe. The accretion disk is in some distance from the WD surface destroyed by a strong magnetic field and the accretion continues through, the so called, accretion curtain across the magnetic force-field.

The falling material in some point creates a stationary shock near the WD surface where the kinetic energy is converted through thermal bremsstrahlung to radiation. The temperature of the created plasma is typically more than 10 keV with a low density. The optically thin hard X-ray¹ emission is taking place and heated gas creates the post-shock region (PSR) with temperature gradient. The hot gas then descends and cools by the X-ray emission while it hits the WD surface.

Because of the relatively high temperature of PSR are IPs very well observed in hard X-rays band. IPs comprise only a small fraction $\sim 15\%$ of all CVs, but they dominate in hard X-ray band over 10keV, at most $\sim 80\%$ of detected CVs are IPs Landi et al. (2009).

The temperature of the PSR depends in first order only on the WD mass, which is the most fundamental parameter of WDs. This means, that if we are able to find temperature from fitting the thermal bremsstrahlung model to a spectrum of IP, we are also able to establish the WD's mass.

The accreting WDs are very important for cosmology, because some of such objects probably cause Type Ia supernovae, when the WD mass reaches Chandrasekhar limit.

As is showed on fig.(1.2), IPs as NY Lup are well observed by INTEGRAL/IBIS detector which makes them interesting space laboratories to investigate the basic

¹In this case, hard X-rays means 10 - 120 keV region.

parameters of WDs. In some cases, can be also accretion stream studied if it is strong enough.

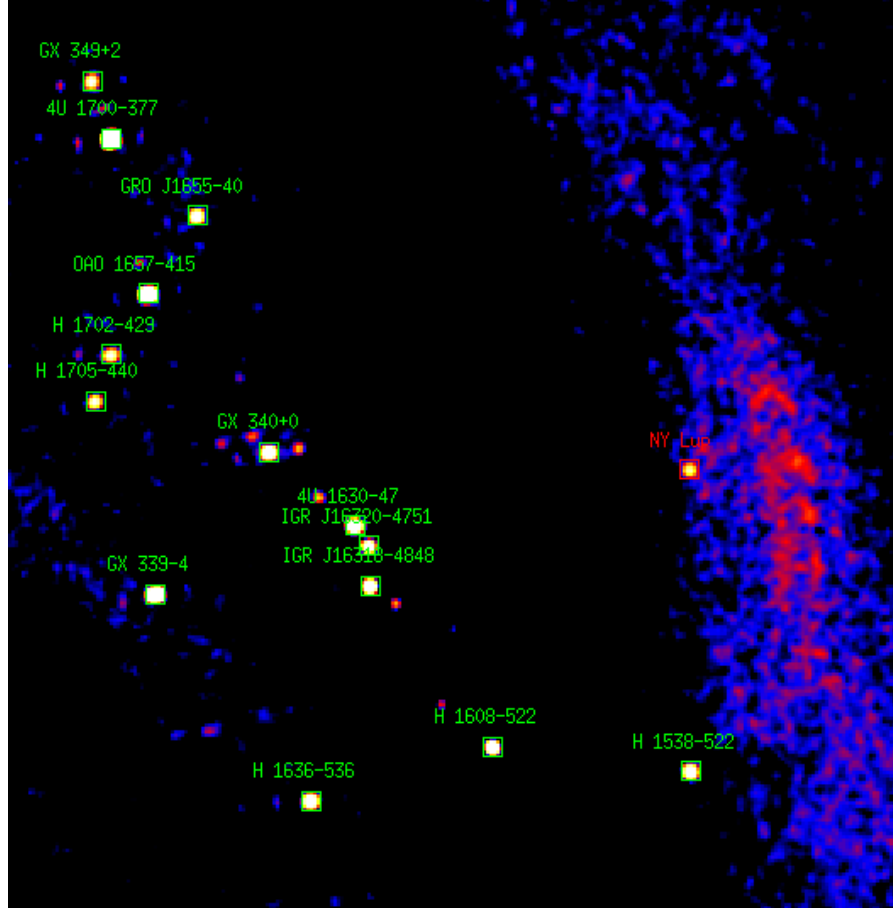


Figure 1.2: 1.2 Msec exposure of NY Lup region in 17-80 keV. The NY Lup is marked by red square. There are many others X-ray sources, mostly HMXBs or LMXBs.

1.3 Observations

Cataclysmic Variable stars (CVs) have been, in fact, observed as early as the ancient times. In the historical records of many civilizations we can find references for various astronomical events. Most of them are about objects that exhibits periodic variations of a some type, most notably: planets, Moon and Sun. However several of them are about comets and new stars. These new stars are in many cases novae and supernovae. In China, the records date back to 1500 BC.

Many records are saved from the medieval time, for example positions of Nova Vulpecula 1670 and Nova Cygni 1600 (now knows as P Cygni) in Hevelius maps.

With the advent of astronomical photography in the late 19th century started the era of continuous observations and with the development of first photo-multipliers in the mid-1940s CVs begun to be attractive targets because of their big variability in the different time scales.

The AAVSO has light curve of SS Cyg from 1896 up to date.

1.3.1 Optical and IR observations

The very first visual observation was followed by the photographic photometry and then spectroscopy, followed by the photo-multiplier photometry since mid-1940s. The binary nature of all CVs was confirmed. The flickering was discovered and was assumed that it is somehow connected to stars' duplicity Walker (1957), Warner (1995).

Statistical studies by Luyten and Hughes in mid-1960th showed, that novae remnants have $M_V \approx 4$ and dwarf novae at quiescence have $M_V \approx 7.5$. They conclude that the hot primary star in CVs must be WD or hot subdwarf Warner (1995).

The most important contribution of optical astronomy to this work is the discovery of large and variable circular polarization in several CVs. This helped to identify magnetic CVs, which were later divided to two categories, polars and intermediate polars.

There are more important discoveries in optical and IR bands in CVs subject. In the case of interest the Warner (1995) is te recommended book.

1.3.2 X-ray observations

The very first CV detected in X-rays was EX Hya observed by Uhuru X-ray space mission. Uhuru worked in 2.0 – 6.0 keV and in spite of its poor sensitivity the well-known 4U catalogue was created Forman et al. (1978) from its observations.

The NASA's HEAO² program followed with three space missions. As the X-ray detectors technology evolved, the number of detected CVs grewed linearly. EXOSAT provided long and uninterrupted data for many CVs during its operation from May 1983 until April 1986. Similar results were obtained from Soviet mission Kvant 1 and Japan's Ginga. The high hopes were invested into ROSAT which provided the all-sky survey in the 0.1 – 2.0 keV but the expected huge number of new CVs was not discovered.

The situation slightly changes with RXTE³ which after several years on orbit provided good data for several articles about WD masses Suleimanov et al. (2005). The data from RXTE are used in the new articles even ~ 15 years after its launch Butters et al. (2011).

Several others missions were launched in the last ten years period. Few of them carried several detectors where one was sensitive in X-rays, like SUZAKU/XIS and Swift/XRT. But for the X-ray astronomy the year of 1999 was the most important

²High Energy Astronomy Observatory, The HEAO 2 was also known as The Einstein Observatory

³Rossi X-ray Timing Explorer

so far. The two major big observatories were launched on the Earth's orbit. The Chandra X-ray Observatory flew onboard STS-93 space shuttle Columbia on July and the XMM-Newton was launched onboard ESA's Ariane 5 rocket.

That was the beginning of the X-ray astronomy's golden era. During the last decade the combination of Chandra and XMM provided enormous data archives which will be useful for astronomers in another decades.

Sadly, there is no big X-ray observatory planned for the next decade. One of the bigger space mission will be Japan's ASTRO-H with several X-ray and gamma ray detectors on-board to cover broad high energy bands. The future of big ESA & NASA space mission Athena (formerly: Constellation-X, XEUS, IXO) is questionable because of budget cuts in both space agencies.

Fortunately, there are several data archives with open acces for anybody interested. This is a big challenge mainly for young astronomers, who are not directly involved in any big space mission program but want to do science. In this case, they don't need any special hardware, even modern laptops are powerful enough.

1.3.3 Gamma ray observations

In last millennium several space mission observed few CVs in bands from tens of keV to TeV.⁴ The biggest breakthrough came with ESA's INTEGRAL space mission which was able to observed many CVs with its exceptional sensitivity and large field of view. Mostly intermediate polars. Only $\sim 2\%$ of all CVs are actually magnetic ones, but these ones are only visible in gamma rays. INTEGRAL/IBIS was been used to determine white dwarf masses by Landi et al. (2009).

Two others space missions have on-board detectors similar to INTEGRAL/IBIS with their sensitivity and coverage: the NASA's Swift/BAT and Japan's Suzaku/XRT. Both are widely used to study white dwarf masses in IPs Brunschweiler et al. (2009), Yuasa et al. (2010).

⁴The most studied CV from this era is AE Aqr (Meintjes 1990; Bowden et al. 1991)

White Dwarfs

White dwarfs born when normal mass stars die. WDs are degenerated, late type stars with typical mass $\sim 1M_{\odot}$. Their typical radius is about 5000 km and mean density around 10^6 g cm^{-3} (Shapiro & Teukolsky (2004)). They no longer burn nuclear fuel and if they don't have any other matter influx e.g. by accretion from close star, they slowly cool as they radiate away residual thermal energy.

WDs support themselves against gravity by the pressure of electron degenerate gas and their interior is in the local thermal equilibrium, except the thin atmosphere.

2.1 First look of white dwarfs interior

WDs are a class of the less compact objects among the possible endpoints of the stellar evolution. The mass of the star is the main factor determining whether the star ends up as a WD, neutron star or a black hole. The medium mass stars with masses $M \lesssim 4M_{\odot}$ ¹ in some point of late state of their evolution gently spreads mass forming planetary nebulae. The rest of the star become the white dwarf.

Let's use data from the table 2.1 and try to assume the pressure inside of WD. For very rough estimation, we can use equation of mechanic equilibrium to compare WD with the Sun:

$$dP = -G \frac{M \rho dr}{r^2}, m(r) = \int_0^r \rho 4\pi r^2 dr. \quad (2.1)$$

The ratio between variables P , M , ρ and r in the Sun case and P' , M' , ρ' , r' in the WD's case can be written as follows:

¹Also $M \lesssim 8M_{\odot}$ can be find in some literature, Padmanabhan (2001).

Table 2.1: Basic statistics of the compact objects Shapiro & Teukolsky (2004)

Object	Mass ^a [M]	Radius ^b [R]	Mean Density [r cm ⁻³]	Surface Potential [GM/Rc ²]
Sun	M _☉	R _☉	1	10 ⁻⁶
White Dwarf	≲ M _☉	~ 10 ⁻² R _☉	≲ 10 ⁷	~ 10 ⁻⁴
Neutron Star	~ 1 – 3M _☉	~ 10 ⁻⁵ R _☉	≲ 10 ¹⁵	~ 10 ⁻¹
Black Hole	Arbitrary	2GM/c ²	~ M/R ³	~ 1

^aM_☉ = 1.989 × 10³³g^bR_☉ = 6.9599 × 10¹⁰cm

$$\begin{aligned}
M' &= M, \\
r' &= 10^{-2}r, \\
\rho' &= 10^6\rho,
\end{aligned}$$

From easy calculation we get:

$$dP' = -G \frac{M' \rho' dr'}{r'^2} = -G \frac{M \cdot 10^6 \rho \cdot 10^{-2} dr}{r^2 (10^{-2})^2} = 10^8 dP \quad (2.2)$$

From such results is eminent that there is something wrong with the WDs interior in comparison with central reion of the normal star. As the ionization is increasing by higher temperatures, the bigger pressure otherwise helps recombination. However very big pressure actually increase ionization up to totally ionized atoms. Atoms without their electrons shells are closer to each others which explains very high density.

For densities in range $10^5 \text{ g cm}^{-3} \lesssim \rho \lesssim 10^9 \text{ g cm}^{-3}$ the WD is made of ideal nondegenerate gas of ions and a degenerate gas of electrons. The system will be degenerate if $T > T_c$ where $T_c \approx 3 \times 10^9 K (\rho/\rho_c)^{2/3}$

If the electrons are relativistic or not can be found by comparing $m_e c$ with a fermi momentum

$$p_F = (3\pi^2)^{1/3} \hbar n_e^{1/3} = (3\pi^2)^{1/3} \hbar (\rho/\mu_e)^{1/3} \quad (2.3)$$

where

$$\mu_e = (\rho/n_e m_p) = 2(1+X)^{-1} \quad (2.4)$$

is the mass per electron. The fermi momentum will be equal to $m_e c$ at the critical density:

$$\rho_c \equiv \frac{8\pi}{3} m + p \mu_e \frac{m_e c^3}{h} \approx 10^6 \mu_e \text{ g cm}^{-3} \quad (2.5)$$

If the density is higher than $\rho \gtrsim 10^9 \text{ g cm}^{-3}$, the electrons are combined with protons inside of nuclei and create different matter called neutron degenerated gas. This is how neutron stars are made.

2.2 Fermi energy

Now we can imagine interior of WDs as an area full of atoms nuclei very close to each other with free electrons around them. But electrons are fermions, which means that they must behave according to Pauli exclusion principle and it allows only at most one fermion per each quantum state.

If we imagine a normal, everyday gas at standard temperature and pressure, only one of every 10^7 quantum states is occupied by gas particles, so Pauli exclusion principle limitation is very insignificant. When energy is removed from the gas and it's temperature falls down, an increasingly large fraction of the particles been forced into the lower energy states.

For fermions gas only one particle can take the lowest energy state and others must take another, higher and higher states, thus only one particle per state is allowed. Even in limit $T \rightarrow 0$ pressure is produced by motions of electrons on excited positions.

At the zero temperature all of the lowest states and none of the higher states are occupied, this kind of fermion gas is called completely degenerated. The max energy of electron in completely degenerate gas at $T = 0\text{K}$ is known as Fermi energy.

For determining the limiting energy we can imagine 3D box where length of each of its sides will be L . The wavelengths of electrons trapped in the box in each dimension will be

$$\lambda(xyz) = \frac{2L}{N(xyz)} \quad (2.6)$$

where N_{xyz} are integer quantum numbers for each dimension. The momentum can be written through de Broglie wavelength

$$p(xyz) = \frac{hN(xyz)}{2L} \quad (2.7)$$

We can now write the total kinetic energy of the electron when $p^2 = p_x^2 + p_y^2 + p_z^2$ as follows

$$\varepsilon = \frac{p^2}{2m} \Rightarrow \frac{h^2 N^2}{8mL^2} \quad (2.8)$$

Total number of electrons is same as total number of unique quantum numbers N_x, N_y, N_z , multiply by factor two. The two factor comes from the fact that electrons are particles with spin, which can be $\pm 1/2$. This means that two electrons are allowed to have the same tree quantum numbers but different spin. The total number of electrons out to radius $N = \sqrt{N_x^2 + N_y^2 + N_z^2}$ will be

$$N_e = 2 \left(\frac{1}{8} \right) \left(\frac{4}{3} \pi N^3 \right) \Rightarrow N = \left(\frac{3N_e}{\pi} \right)^{1/3}. \quad (2.9)$$

By using Eq. (2.8) and simplifying it, we will get the Fermi energy as follows

$$\varepsilon_F = \frac{\hbar^2}{2m} (3\pi^2 n)^{2/3} \quad (2.10)$$

where m is the mas of the electron or any other fermion² and $n \equiv N_e/L^3$ is the electron count per volume. The average energy per electron at $T = 0\text{K}$ is $3/5\varepsilon_F$.

2.3 Degeneracy

Matter inside of WD has very symmetric spherical distribution, we can calculate the mass interior to radius r like³

$$\frac{dm(r)}{dr} = 4\pi r^2 \rho \quad (2.11)$$

In normal case the WD is in steady state and gravitation force is balanced by the pressure at every point. For deriving the hydrostatic equilibrium equation we need to consider an infinitesimal element laying between r and $r + dr$ with an area dA . The element is lying perpendicular to the radial direction, should looks like fig.(2.1).

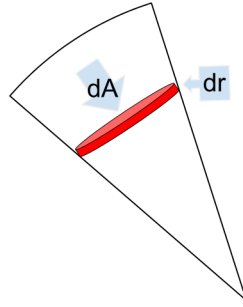


Figure 2.1: How could be infinitesimal fluid element laying between r and $r + dr$ with an area dA imagine

The gravitation attraction between mass $dm = \rho dA dr$ and $m(r)$ is the same as if $m(r)$ were only point at the center with the same mass. While the outside mass exerts no force on dm . Then the net outward pressure forced on dm is $-[P(r + dr) - p(r)]dA$. In equilibrium

$$\frac{dP}{dr} = -\frac{Gm(r)\rho}{r^2} \quad (2.12)$$

In such equilibrium, the gradient of degeneracy pressure is balanced by gravitation:

$$\nabla P = -\rho \nabla \Phi, \quad (2.13)$$

²Can by applies for any fermion, not only electrons

³More elaborate and precises approach can by find Shapiro & Teukolsky (2004), Padmanabhan (2001), Kleczek (1957), Camenzind (2007), Nauenberg (1972) and simpler explanation with less math in Carroll & Ostlie (2007)

where Φ means gravitation potential. Consequence of the Eq.(2.12) is the viral theorem. The gravitation potential energy of the star is then

$$W = - \int_0^R \frac{Gm(r)}{r} \rho 4\pi r^2 dr = \int_0^R \frac{dP}{dr} 4\pi r^3 dr = -3 \int_0^R P 4\pi r^2 dr \quad (2.14)$$

We can characterize the gas by an adiabatic equation of state, where K and Γ are constants

$$P = K \rho_0^\Gamma \quad (2.15)$$

Then Eq.(2.14) can be rewritten as follows

$$W = -3(\Gamma - 1)U, \quad (2.16)$$

Where U is the total star's internal energy

$$U = \int_0^R \epsilon' 4\pi r^2 dr \quad (2.17)$$

and ϵ' comes from:

$$\epsilon' \equiv \epsilon - \rho_0 c^2 \quad (2.18)$$

where

$$\epsilon = \rho_0 c^2 + \frac{P}{\Gamma - 1} \quad (2.19)$$

Now we can see that the energy density, of the gas (excluding the rest mass energy) can be also written as

$$\epsilon' = \frac{P}{\Gamma - 1} \quad (2.20)$$

Assuming adiabatic changes, the Eq.(2.20) follows from the first law of thermodynamics

$$d\left(\frac{\epsilon}{\rho_0}\right) = -P d\left(\frac{1}{\rho_0}\right). \quad (2.21)$$

The equation of state for ideal Fermi gas reduce to the simple polytropic form Eq.(2.15) in the two limiting cases Shapiro & Teukolsky (2004), Padmanabhan (2001):

- Nonrelativistic electrons, $\rho_0 \ll 10^6 g.cm^{-3}$, $x \ll 1$, $\Phi_{(x)} \rightarrow x^5/15\pi^5$

$$\Gamma = \frac{5}{3} \implies K = \frac{3^{2/3} \pi^{4/3}}{5} \frac{\hbar^2}{m_e m_u^{5/3} u_e^{5/3}} = \frac{1.0036 \times 10^{13}}{u_e^{5/3}} \text{cgs}. \quad (2.22)$$

- Extremely relativistic electrons, $\rho_0 \gg 10^6 g.cm^{-3}$, $x \gg 1$, $\Phi_{(x)} \rightarrow x^4/12\pi^2$

$$\Gamma = \frac{4}{3} \implies K = \frac{3^{1/3} \pi^{2/3}}{4} \frac{\hbar c}{m_u^{4/3} u_e^{4/3}} = \frac{1.2435 \times 10^{15}}{u_e^{4/3}} \text{cgs}. \quad (2.23)$$

This kind of the equilibrium configurations such these are called polytropes. We can use Eq.(2.13), take the divergence of it and by using $\nabla^2\Phi = 4\pi G\rho$ in spehrically symmetric case we get

$$\frac{1}{r^2} \frac{d}{dr} \left(\frac{r^2}{\rho} \frac{dP}{dr} \right) = -4\pi G\rho. \quad (2.24)$$

Using few trics, Eq.(2.24) can be rewrite and reduce to dimension less form

$$\frac{1}{\xi^2} \frac{d}{d\xi} \xi^2 \frac{d\theta}{d\xi} = -\theta^n \quad (2.25)$$

use Eq.(2.15) with $\Gamma \equiv 1 + \frac{1}{n}$ where n is called the polytropic index

$$\rho = \rho_c \theta^n \quad (2.26)$$

$$r = a\xi \quad (2.27)$$

$$a = \left[\frac{(n+1)K\rho_c^{(1/n-1)}}{4\pi G} \right]^{1/2} \quad (2.28)$$

where ρ_c is the central density, $\rho_c = \rho$ when $r = 0$.

The Eq.(2.25) is called the *Lane-Emden equation* for the structure of n index, the boundary conditions at the center such polytropic white dwarf are

$$\theta(0) = 1, \quad (2.29)$$

$$\theta'(0) = 0. \quad (2.30)$$

The condition (2.29) come dirrectly from Eq.(2.26) and condition (2.30) is derived from the fact that near the center of the star is $m(r) \approx 4\pi\rho_c(r^3/3)$. If we use Eq.(2.12), we get

$$\frac{dP(\rho)}{dr} = \frac{d\rho}{dr} = 0. \quad (2.31)$$

If we numerically integrate Eq.(2.25) using Eq.(2.29) and Eq. (2.30) as the boundary conditions, starting at $\xi = 0$. We will find that for $n < 5$ ($\Gamma > 6/5$) this solution decrease monotonically with a zero at finite value $\theta(\xi_1) = 0$, which corresponds to the star's surface where $P = \rho = 0$ and we get the equation for white dwarf radius

$$R = a\xi_1 = \left[\frac{(n+1)K}{4\pi G} \right]^{1/2} \rho_c^{\frac{1-n}{2n}} \xi_1. \quad (2.32)$$

The mass of the WD can be find as follows

$$\begin{aligned}
 M &= \int_0^R 4\pi r^2 \rho dr \\
 &= 4\pi a^3 \rho_c \int_0^{\xi_1} \xi^2 \theta^n d\xi \\
 &= -4\pi a^3 \rho_c \int_0^{\xi_1} \frac{d}{d\xi} \left(\xi^2 \frac{d\theta}{d\xi} \right) d\xi \\
 &= 4\pi a^3 \rho_c \xi_1^2 |\theta'(\xi_1)| \\
 &= 4\pi \left[\frac{(n+1)K}{4\pi G} \right]^{3/2} \rho_c^{\frac{(3-n)}{2\pi}} \xi_1^2 |\theta'(\xi_1)|,
 \end{aligned} \tag{2.33}$$

then ρ_c can be eliminated between Eq.(2.32) and Eq.(2.33), which gives the mass-radius relations for polytropes Shapiro & Teukolsky (2004):

$$M = 4\pi R^{\frac{3-n}{1-n}} \left[\frac{(n+1)K}{4\pi G} \right]^{\frac{n}{n-1}} \xi_1^{\frac{3-n}{1-n}} \xi_1^2 |\theta'(\xi_1)|. \tag{2.34}$$

2.4 The Chandrasekhar limit

What will happen if $\rho_c \rightarrow \infty$? This is one of the most important questions in the universe. Because how central density of WD increases by increase of the WD mass, the electrons become more and more relativistic throughout the star. As $R \rightarrow 0$ also the space for free electrons moving between atoms nuclei become smaller and smaller. With increasing mass must also increase the speed of electrons to support the degeneracy pressure. While the mass of WD is larger the radius become smaller, mass-volume relations implies that $\rho \propto M_{WD}^2$. But the electrons can't have speed larger than speed of light and in some point where $\rho \sim 10^9 \text{ g cm}^{-3}$ there will be no space left for moving electrons and they will be press into the nuclei. This will leads to explosion of such WD.

This all means, that there is a maximum mass for WDs. This mass is called *Chandrasekhar limit* ⁴ M_{ch} .

The classical theory of type 1A supernovae assumes explosions of such WDs reaching the M_{ch} , the very same mass of the WDs when they explode makes Type 1A supernovae perfectly suitable objects for measuring the long range space distances. That's why they are so important objects for cosmology and astronomers called them standard candles.

The Chandrasekhar limit for relativistic electron case is

$$M_{ch} = 1.457 \left(\frac{2}{\mu_e} \right)^2 M_{\odot}, \tag{2.35}$$

⁴Named after Indian physicist Subrahmanyan Chandrasekhar who received the Nobel Price in Physics in 1983 for his work on stellar structure and evolution. Chandrasekhar made his great discovery at the age of 21 in the 1931.

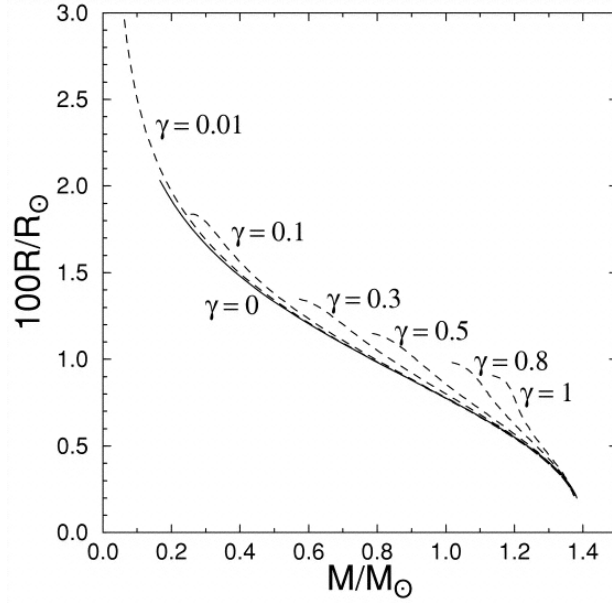


Figure 2.2: Relation between the mass M and radius R of a ^{12}C magnetic white dwarf for the indicated magnetic field strengths. The solid line denotes the Hamada & Salpeter (1961) model for nonmagnetic white dwarfs ($\gamma = 0$). The dashed lines are magnetic white dwarfs. Suh & Mathews (2000)

where μ_e is the average molecular weight per electron, which depends on the chemical composition of the star and is typically sets to $\mu_e = 2$.

The Eq.(2.35) can be obtained by solving Eq.(2.34) with parameters⁵:

$$\Gamma = \frac{4}{3}, \quad n = 3, \quad \xi_1 = 6.89685, \quad \xi_1^2 |\theta'(\xi_1)| = 2.01824. \quad (2.36)$$

Another, illustrativ aproach to estimating the Chandrasekhar limit is mention in Carroll & Ostlie (2007). An approximate value for M_{ch} may by obtain by setting the estimate of the central pressure

$$P_c \approx \frac{2}{3} \pi G \rho^2 R_{\text{wd}}^2 \quad \text{with} \quad \rho = M_{\text{wd}} / \left(\frac{4}{3} \pi R_{\text{wd}}^3 \right) \quad (2.37)$$

equal to electron degeneracy pressure equation

$$P = \frac{(3\pi^2)^{2/3}}{4} \hbar c \left[\left(\frac{Z}{A} \right) \frac{\rho}{m_H} \right]^{4/3}, \quad (2.38)$$

with $Z/A = 0.5$, the R cancels from the equation leving for the greatest possible mass:

$$M_{\text{ch}} \sim \frac{3\sqrt{2\pi}}{8} \left(\frac{\hbar c}{G} \right)^{3/2} \left[\left(\frac{Z}{A} \right) \frac{1}{m_H} \right]^2 M_{\odot}. \quad (2.39)$$

⁵Extensive list of polytropic parameters can be find in Chandrasekhar (1939)

It is important to notice that Eq.(2.39) contains three fundamental constants \hbar, c and G representing the combined effects of quantum mechanics, relativity and Newtonian gravitation on the WD structure.

The Chandrasekhar limit slightly depends upon the chemical composition (2.35), but in the literature is well known as $M_{ch} = 1.44M_{\odot}$.

The similar limit exists for neutron stars, where pressure of degenerate neutron gas support the stars against its own gravity. This limit is called Tolman — Oppenheimer — Volkoff limit (or TOV limit). But it is not so easy in neutron stars case. Their limit mass depends strongly on type of matter in star center. The limit varies from $0.7M_{\odot}$ up to $\sim 3M_{\odot}$ in extreme case, when inner section of neutron star is partially filled up by quark - gluon matter.

2.5 White dwarfs classes

Previous sections describe interior and fundamental parameters of white dwarfs, but it is also important to tell something about their classes. WDs occupy narrow sliver line in the left bottom corner of H-R diagram fig.(2.3), slightly parallel with the main sequence. Although WDs are typically whiter than normal stars, the name itself is not very comprehensive. WDs in fact come in all colors with surface temperatures from less than 4000K to even more than 80,000K. The *D* for *dwarf* spectral type has several subdivisions:

- **DA** white dwarfs is the largest group ($\sim 60\%$) including Sirius B, their spectra have only pressure-broadened hydrogen lines
- **DB** white dwarfs ($\sim 8\%$) have only helium absorption lines in their spectra
- **DC** white dwarfs ($\sim 14\%$) have no lines, but only continuum features
- **DQ** white dwarfs show carbon features in their spectra
- **DZ** white dwarfs show some evidence of metal lines

Very small radius makes WDs hard to observe in optical band. The brightest and also the most known WD is Sirius B with visual magnitude ~ 8.6 . Unfortunately it is located only 8" from its companion Sirius A, the brightness star of the night sky. Although the WDs are not common stars on the night sky, they are not rare in our Galaxy. From hundred closest stars to our Sun, eight of them are WDs.

WDs usually have very strong magnetic field which comes from weak surface magnetic field of progenitor's star which large surface area is conserved to their very small surfaces. The magnetic field intensities vary from about 1000 G up to 10^9 G in extreme cases.

WDs can be found alone in interstellar space, young ones can be found also in planetary nebulae, like one in center of the Ring nebula M57 or they can be found in cataclysmic variables.

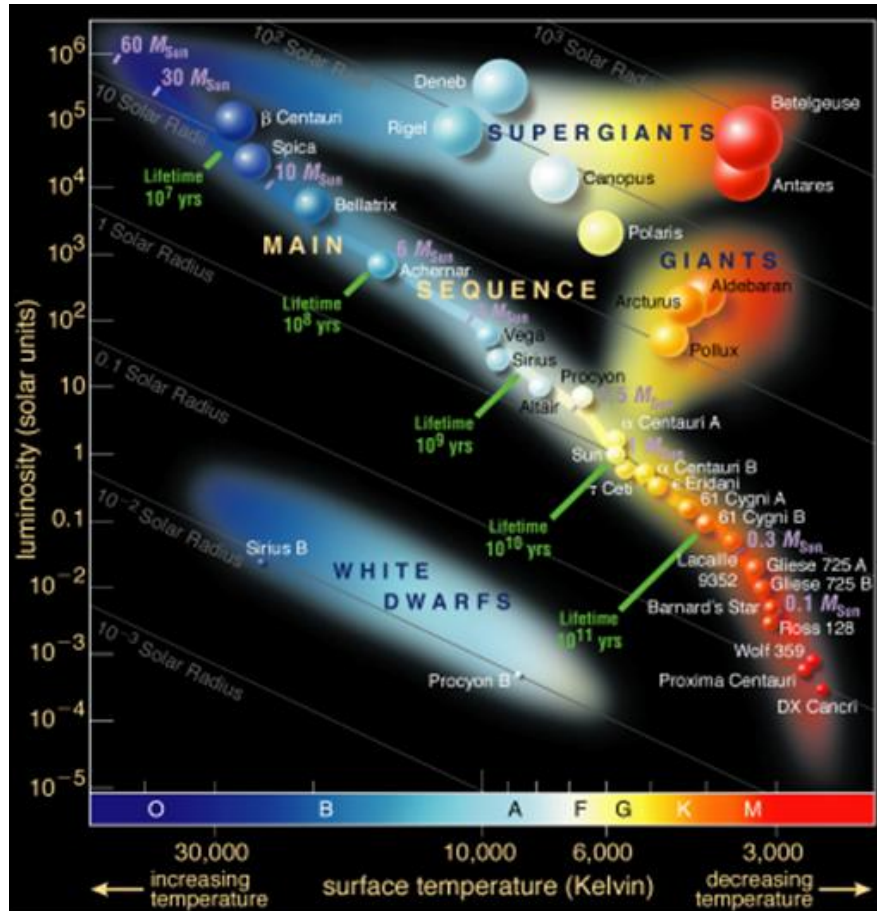


Figure 2.3: WD lies in right bottom corner of HD diagram, under main sequence, which exactly means, they are very small (dwarf) stars with high surface temperature.

Cataclysmic variable stars

Cataclysmic variable stars (*CVs*) received their name for their cataclysmic events, like outbursts, novae and supernovae eruptions. In this one name several very different mechanisms are used for creating the various cataclysmic events. Dwarf novae are common for their increase of brightness by factor ten, classical novae are increasing their brightness usually in factor 10^6 and type 1A supernovae in factor 10^{12} , typical visual absolute magnitude for them is $M_V = -19.3$ ¹.

What is common for all the CVs types is their binary star character with mass transfer through Roche lobe to primary companion, which is white dwarf star. Several thousands of CVs with even more candidates are known up to date. The mean mass of the primary star is $\sim 0.8M_\odot$, which is larger than the average of $\sim 0.6M_\odot$ for isolated, forever alone WDs. The secondary star is a main-sequence low mass star of G or later type, but usually with less mass than the primary WD.

Configuration like this make stars orbit each other with periods from around 20 minutes up to several days, however the vast majority of CVs have orbital periods from 1 to 12 hours².

The various classifications and approaches to study the CVs exists. They mostly depends on the different wavelength in which the CVs are observed. The biggest and also the well observed is optical band where we can see these types:

- **Classical novae** are characterize by large outbursts causes by thermonucler exlosion of accreted material on WD's surface
- **Recurrent novae** have small outbursts repeating every few years, typical example is RS Oph

¹About 5×10^9 times brighter than the Sun

²Interesting think is that there exists a "period gap" in CVs periods in range from 1.5 to 3.25 hours. The cause of this gap is not well understood yet.

- **DN - dwarf novae** use to have several smaller outbursts, have several sub-classes (*Z Cam*, *SU UMa*, *SS Cyg*³)
- **AM Canum Venaticorum** are extremely interesting space objects where the secondary star is also a compact object and the accretion disk is mostly composed from helium and they could be source of strong gravitation waves
- **Novae like systems** are possible nova remnants or stars with outburst behavior similar to novae but maybe miss-classify
- **Polars** are magnetic CVs with $B \sim 10^7 - 10^9$ G, they got the name because of strong polarization of their light in optical and IR bands
- **IP intermediate polars** are magnetic CVs with B less than polars $\sim 10^6 - 10^7$ G
- **Type 1A supernova** is sub-class of supernovae which become a result of thermonuclear explosion of whole WD when it reaches the M_{ch} limit

During last years CVs became more and more attractive target for modern astronomers mainly because of the progress in detectors technology in different then optical band and also for their role as a space laboratories for study accretion of matter on compact star, X-ray emission from shock regions, and even gravitation waves.

The CVs are complex and complicated space objects, for this reason is hard to make a good classification. If we look inside Warner (1995), we will find several fact which are not consider as correct now. With new huge sky surveys as GAIA⁴ or LSST⁵ the new and even more complex classification will need to take place. The number of known CVs systems will grow rapidly, also as the number of observations. Next decade will be the gold era for new scientific approaches like data-mining in huge amount of data-sets.

In following sections of this work we use the most broad CVs classification, for magnetic and non-magnetic systems. Because of the aim of this work on magnetic systems, we will take care of non-magnetic only roughly. Please also note, that if we want to make a very good classification of CVs to many classes, almost everyone of them will has its own class.

³SS Cygni had few outbursts, but according to latest observations, it seems to be more-likely intermediate polar

⁴Is space mission to chart a three-dimensional map of our Galaxy, will observed \sim one billion stars. Launch is planned for 2013.

⁵The Large Synoptic Survey Telescope (LSST) will be wide-field survey 8.4 m telescope which will photograph the available sky every three nights from his place on the El Peñón peak of Cerro Pachón, at 2682 m in Chile.

3.1 Non magnetic cataclysmic variables

The magnetic field plays essential role in all the CVs, even in non-magnetic ones, at-least weak magnetic field provides viscosity in accretion disk. We will called non-magnetic CVs when $B \lesssim 10^5$ m. In such system is movement of gas from the secondary companion to the WD determined predominantly by dynamical and hydro-dynamical flows Warner (1995).

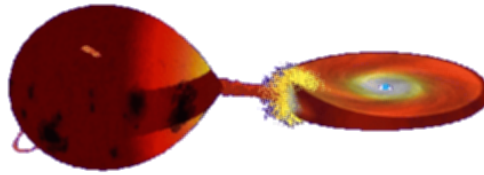


Figure 3.1: The illustration of cataclysmic variable star, showing the normal, secondary star on the right, accretion stream meeting the accretion disk and accretion disk itself around the white dwarf.

The non-magnetic CV could looks like fig.3.1 or the (A) part of fig.3.3. The secondary is a small main sequence star with mass about half of the Sun mass which is filling out its Roche lobe. The mass transfers from secondary to primary through the Lagrangian point L1 and in some point its creates the accretion disk. Streaming mass meets the accretion disk and if the mass stream is big enough and disk is dense enough, so called hot spot is created as we can see on fig.3.1.

3.1.1 Orbital periods

Orbital period P_{orb} is usually the most precisely know physical parameter of CV. Because of facts that the CVs orbital periods are relatively short and to take enough data-sets to create a phase light curve is then not that difficult task. P_{orb} of the system reveals how big the CV systems are. The most of the CVs with typical periods in range hours can fit inside Earth - Moon orbit.

Inspecting the Riter & Kolb catalogue Ritter & Kolb (2003) within the orbital period range from 1 hour to 1 day, without the AM CVn systems, it is found that almost half of all CVs in catalogue are the dwarf novae.

- 166 DNs 63% have $P_{orb} < 2h$
- 26 DNs 10% are found in the 2-3 h gap
- 70 DNs 27% are behind the gap with $P_{orb} > 3h$

The conventional idea about where the gap comes from points to internal structure changing in the secondaries. In some point of secondary star evolution it has a

radius corresponds to its mass in excess of the equilibrium value. This force secondary to detaches from Roche lobe for a while. The stellar wind is also very low during this period which cause disappearance of accretion disk. The period gap and the orbital period distribution per all the CVs and between the different CVs types is well shown on fig.3.2.

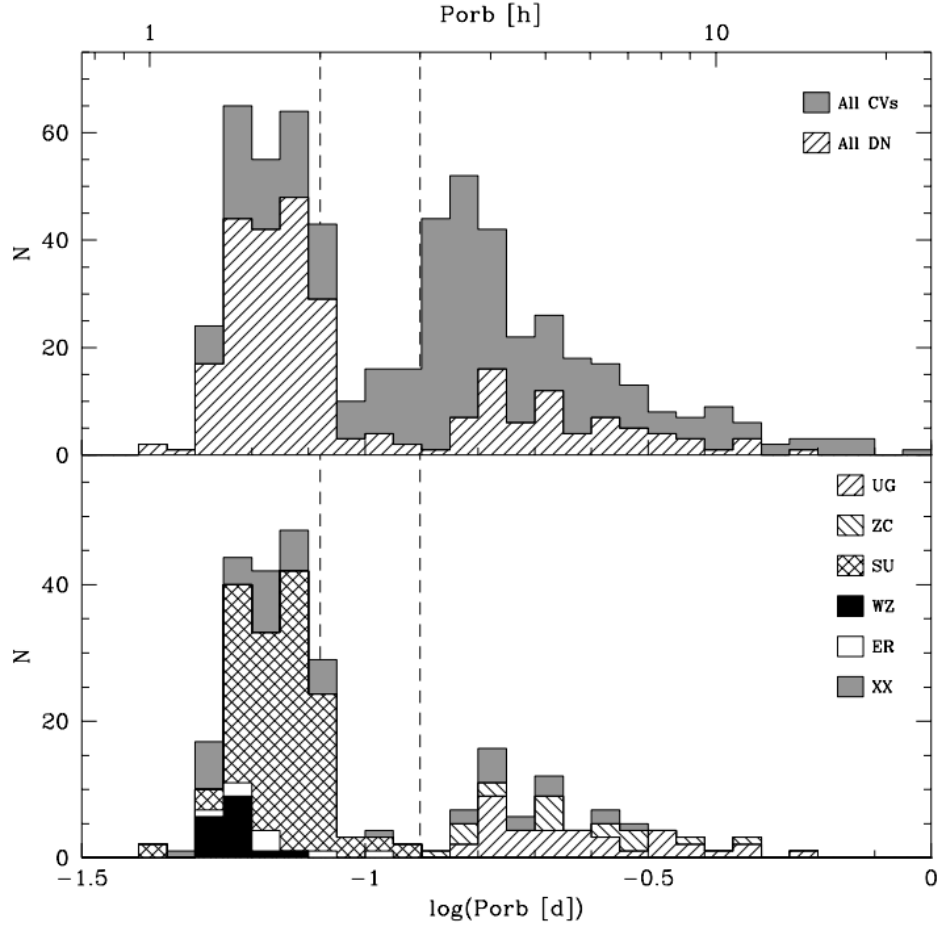


Figure 3.2: *Top panel:* the orbital period distribution of known CVs and dwarf novae from are shown in gray and shade Ritter & Kolb (2003). *Bottom panel:* the period distribution of known dwarf novae according to their subtypes, U Gem (UG), Z Cam (ZC), SU UMa (SU), WZ Sge (WZ), ER UMa (ER), and unclassified subtype (XX). The dashed lines represent the conventional 2-3h period gap Aungwerojwit (2007).

3.2 Magnetic cataclysmic variables

Magnetic CVs are these where the strong magnetic field significantly affects the accretion process. These systems are divided according to strength of their magnetic field to two categories:

- Polars, also called AM Herculis stars, $B_{wd} \sim 10^7 - 10^9$ G
- Intermediate polars $B_{wd} \sim 10^6 - 10^7$ G

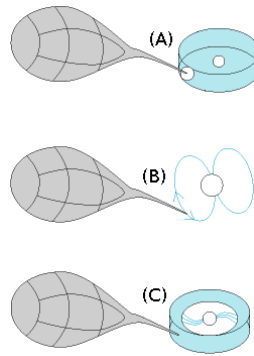


Figure 3.3: An illustration of difference between non-magnetic and magnetic CVs. (A) is an illustration of non-magnetic disk-type system, same as artist impression fig.3.1. (B) is a system with strongly magnetized WD called Polar where magnetic field prevent creations of accretion disk and the in-falling material forms an accretion stream following the magnetic field lines of WD. (C) is intermediate case between (A) and (B), the not such strongly magnetized WD as one in polars destroying only inner part of accretion disk by his B_{wd} creating accretion curtain following its magnetic field lines.

3.2.1 Polars

The (B) part of fig.3.3 is how schematically polars looks like. The material from secondary is falling via the inner Lagrangian point towards to the primary, which is the highly magnetized WD as was mention before. The gas in the stream is ionized by collisions and X-rays from the accretion region on the WD. In some point at the stream's trajectory the magnetic energy density of the WD's field is sufficient to divert the stream from its free fall trajectory and force it to follow the magnetic field lines. Accretion therefore occurs over a small area near both or one of the magnetic poles of the WD. The magnetic field is assumed to be bipolar Cropper (1990).

How material traveling at roughly the escape velocity of the WD it approaches to the surface of the WD a very strong shock region is formed in accretion flow. The

hard X-rays are emitted from post-shock region. Fraction of X-rays is reprocessed by the surface of WD and re-emitted as the softer X-rays component.

Free electrons close to the shock spiral around the magnetic field lines, therefore, emit strongly polarized cyclotron radiation. The magnetic field of WD is such strength that the cyclotron radiation is emitted in optical and near infra-red wavelengths.

In polars, the WD magnetic field is sufficiently strong to force the primary to rotate synchronously. The magnetic field therefore always present the same section to the incoming accretion stream. This is the principal distinguishing feature for polars (intermediate polars rotates asynchronously).

The principal defining characteristic of the polar CVs class is their strongly and variably circularly and linearly polarized optical emission. They are also characterized by high ratio of soft-to-hard X-ray luminosity and high excitation optical spectra with prominent He II 468.6 emission lines. For a review of the polars, see Cropper (1990) or Warner (1995).

3.2.2 Intermediate polars

The (C) part of fig.3.3 illustrates how can be intermediate polars imagined. They are CVs with not as strong magnetized WDs as in polars. This is causing a several important differences between polars and IPs. In IPs the other part of accretion disk exists, but at some distance of WD the kinetic energy density ρv^2 of gas in accretion disk is locally exceeded by the magnetic energy density $B^2/8\pi$ and within this radius the in-falling gas will be guided along magnetic field lines to accrete radially onto the WD creating kind of accretion curtain. For the spherical accretion is easy to calculate transition radius, because v is simply the free-fall velocity: $\sqrt{2MG/r}$, the radius is according Patterson (1994):

$$R_A = 3.7 \times 10^9 \dot{M}_1^{-2/7} M_{wd}^{-1/7} \mu_{32}^{4/7} \text{ cm}, \quad (3.1)$$

where R_A is the *Alfven radius*, \dot{M}_1 is the accretion rate in units of 10^{17} g s^{-1} , M_{wd} is the mass of WD in units M_\odot and μ_{32} is the magnetic moment of the WD in units of 10^{32} G cm^3 . For disk accretion, the magnetospheric radius should be slightly smaller, theoretical estimates suggest $R_{mag} \approx 0.5 R_A$ Patterson (1994).

As the gas almost freely falls onto WD channeled along magnetic field lines, in some distance near the WD surface a stationary shock stands to convert the kinetic energy of the gas into thermal energy, Yuasa et al. (2010). The temperature of such shock heated gas is typically over 10keV and the density is low. The hard X-rays are emitted via optically thin thermal emission. Then the heated gas forms a post shock region (PSR) with a temperature gradient. The gas descends while it is cooled via X-ray emission and finally hits the surface of WD with small energy by emitting UV light Aizu (1973). This property makes from IPs very good targets for hard X-ray space missions. The closer look to PSR will be taken in the next chapter.

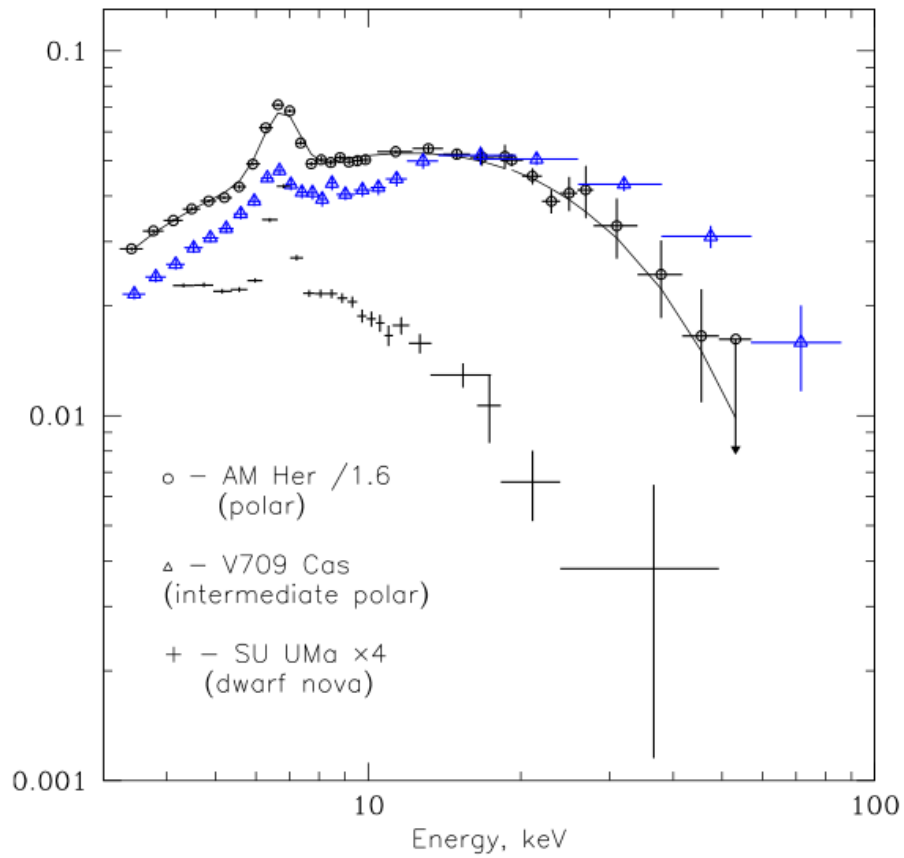


Figure 3.4: Typical broadband spectra of different CVs' classes, indicating that INTEGRAL observations in the hard X-ray energy band (17–60 keV) are biased toward detecting the hardest CVs – intermediate polars, Revnivtsev et al. (2008).

Masses of white dwarfs in intermediate polars

As it was mentioned before, the mass of a WD is the most fundamental parameter in CVs. It is also the main parameter characterizing the accretion flow, the emission from the accretion region and, most importantly, it is the parameter which governs the dynamics of orbital motion of the system. The knowledge of WDs' masses in CVs plays a fundamental role in understanding the binary evolution.

Previous attempts to determine the WDs' masses in IPs using X-ray measurements did not succeed, because the data from different instruments led to different mass estimates. In all cases, the masses determined by the X-ray measurements were substantially higher than masses derived from the optical observations. The most notable example is XY Ari, which is an eclipsing IP with a very small uncertainty in the inclination (Ramsay et al. (1998)).

- RXTE spectrum yields $1.22 M_{\odot}$
- ASCA data yields $1.27 - 1.4 M_{\odot}$
- Ginga results $1.3 M_{\odot}$
- the optical methods yields $0.78 - 1.03 M_{\odot}$

Observations in hard X-rays over 20 KeV are crucial as they provide a clean signal, not contaminated by other system components such as the accretion disc or the main-sequence star and they avoid complications due to cold or warm absorbers as well as the fluorescence emission (Brunschweiler et al. (2009)). However, this method is not suitable because of the weakness of the X-ray flux. In spite of the fact that the IPs are among the brighter sources in the hard X-ray sky, their photon count is relatively low. Thus a sensitive instrument must be used for this kind of observations. At present, there are three instruments which are widely used for such

observations: Swift/NASA Burst Alert Telescope, Hard X-ray Detector (HXD) on-board Suzaku/JAXA space mission and IBIS on-board INTEGRAL/ESA.

The idea how to study the post-shock region, the X-ray emission from IPs and then determine the mass of WD is to try to cover broad bands from soft to hard X-rays. Ideally, from 1 keV up to 100 keV. In the softer X-rays below 10 keV there are, however, several other features present, but the photon flux in this band is much higher than in bands over 15 keV.

In this case, it is necessary to combine two different detectors, one for soft and one for hard X-rays. The optimal combination would be XMM-Newton for soft X-ray in range 1–10 keV and INTEGRAL for hard X-ray from 20–100 keV.

This combination provides a good compromise. The photon flux in the soft X-ray is high enough to allow short pointings of XMM with a good s/n ratio. However, the fitting model for such spectrum can be complicated because of the presence of several other features e.g. reflections or the 6.4 keV emission line.

The hard X-rays are important for the localization of the spectral energy cut-off so a good fit of the model, either power-law or bremsstrahlung, can be applied. INTEGRAL/IBIS is, due to a relatively good sensitivity and large field of view (FOV), a suitable instrument for observations in this band. The large FOV allows a good coverage of the X-ray sky enabling long exposures covering multiple objects in this band.

4.1 Post shock region

Reason why IPs are the most suitable for measurements of the mass of their WD companion then others CVs is their magnetic field, which is not as low as in non-magnetic CVs $B < 10^6$ but still weaker than in polars $B \gtrsim 10^7$. As was mentioned before, in IPs is the accretion disk destroyed in some distance from WD and accretion continues through accretion curtain, where the gas follows magnetic force lines. The resulting configuration is often known as an accretion column. From the theory of spherical accretion (Frank et al. (2002)) we expect that accreting matter to be highly supersonic and therefore in free-fall above the WD's pole-caps.

The in-falling matter must be somehow decelerated to subsonic velocities in order to accrete, therefore we expect strong shock-wave in matter stream where the kinetic energy of infall

$$\frac{1}{2}v_{ff}^2 = GM_{wd}/R_{wd}, \quad (4.1)$$

will be turned into thermal energy. Because that $m_p \gg m_e$ the ions are those who bring the most of this kinetic energy supply. The kinetic energy of ions must be emitted, absorbed or re-emitted by WD. The luminosity corresponding to the typical accretion rate $\dot{M} \gtrsim 10^{16} \text{ g s}^{-1}$ will be of order $10^{33} \text{ erg s}^{-1}$ which greatly exceeds any intrinsic luminosity (typically $\ll 10^{31} \text{ erg s}^{-1}$) that the WD may have as a result of left-over thermal energy in its degenerated core (Frank et al. (2002)).

The accretion luminosity is according to Frank et al. (2002):

$$L_{acc} \lesssim 2 \times 10^{29} f_{-2} m_1^{18/7} R^{-1/7} \text{ erg s}^{-1}, \quad (4.2)$$

where $f_{-2} = 10^2 f$ and f is a surface fraction. Which is much lower than the bolometric luminosities of observed CVs. Therefore, we must conclude that the infalling matter goes through strong shock which slows it down before reaching the WD surface.

Now we know that the shock gas is hot with characteristic (shock) temperature

$$T_s = \frac{3}{8} \frac{GM\mu m_H}{kR_{wd}} = 3.7 \times 10^8 m_1 R^{-1} \text{ K}. \quad (4.3)$$

From previous chapters we already know that the main cooling mechanism is thermal bremsstrahlung, but let's try to determine this on our own a little bit.

In steady state in accretion column, the accretion energy released below the shock must be removed from the post-shock column at the same rate in which it is deposited. Such removal must be effected by a combination of the cooling mechanisms such as radiation and particle transport processes. It is also determined by distance of the shock-wave from WD surface.

If the cooling were inadequate the gas in PSR would simply expand adiabatically and raise the shock. On the other hand, if cooling became too effective the shock would move in towards the WD surface, reducing the volume of the PSR cooling region (Frank et al. (2002)).

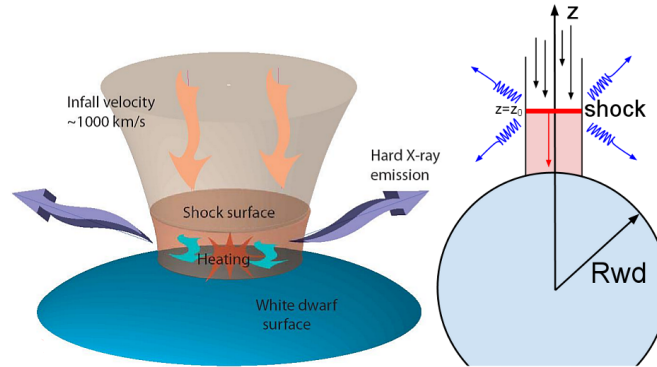


Figure 4.1: Both images are showing the same accretion column geometry for a IP. The shown accretion column is circular and accretion plasma is assumed to be uniform across the column, image is from the Krivonos & Revnivtsev (2011). The left figure is showing 3D illustration while the right one shows the schematically the geometry of the PSR.

For understanding and constructing appropriate accretion column model we must specify how the shock-heated gas in PSR is cooled. The estimation of which cooling mechanisms is important in IPs should be based on following ideas.

The gas velocity below the shock is according to Frank et al. (2002)

$$v = 1.3 \times 10^8 m_1^{1/2} R^{-1/2} \text{ cm s}^{-1}, \quad (4.4)$$

using the continuity equation

$$4\pi R^2 f(-\rho v) = \dot{M} \quad (4.5)$$

we can find a PSR gas density corresponding to the electron density

$$\begin{aligned} \rho_{PSR} &= 6.2 \times 10^{-10} \dot{M} f_{-2}^{-1} m_1^{-1/2} R^{-3/2} \text{ g cm}^{-2}, \\ N_e &= 5.9 \times 10^{14} \dot{M} f_{-2}^{-1} m_1^{-1/2} R^{-3/2} \text{ cm}^{-3}, \end{aligned} \quad (4.6)$$

assuming a fully ionized gas with standard cosmic abundance. With characteristic PSR gas temperature T_s is electron scattering only possible source of opacity. The shock distance from WD surface is lower the $R_w d$. The results of this is that the horizontal optical depth is relatively low, therefore the radiative cooling will be optically thin (unless the accretion rate per unit area will be extremely high).

Because of high T_s (Eq.(4.3)), radiative cooling of the PSR gas will be dominated by the bremsstrahlung (free—free) emission. The volume emissivity is

$$4\pi j_{\text{bremss}} \cong 2 \times 10^{-27} N_e^2 T_e^{1/2} \text{ erg s}^{-1} \text{ cm}^{-3}, \quad (4.7)$$

for a gas of cosmic abundance with electron temperature and density T_e, N_e .

The estimates from last paragraphs show that the PSR is hot, optically thin and that shock lies quite close to the WD surface. The falling gas is cooled by optically thin bremsstrahlung emission and then it hits WD surface, with much lower than free-fall speed, where black-body emission will take place with maximum in UV band.

We may briefly consider some of the extreme cases. If the WD is extremely massive close to the $M_c h$ therefore with very low $R_w d$ and the accretion stream will be low, the Compton scattering will become important as radiative cooling process. Also the cyclotron cooling, which is negligible in IPs, will take place in case of low accretion stream and higher B_{wd} .

4.1.1 Bremsstrahlung

Bremsstrahlung is a German word for breaking radiations, which is a radiations due to the acceleration of a charge in the Coulomb field of another charge also called as free-free emission. A full understanding of this process requires a quantum treatment, since photons of energies comparable to that of the emitting particle can be produced, Rybicki & Lightman (1979).

The electrostatic accelerations of the in its rest frame, parallel a_{\parallel} and perpendicular a_{\perp} to its direction of motion are:

$$\begin{aligned} a_{\parallel} = \dot{v}_x &= -\frac{eE_x}{m_e} \frac{\gamma Z_e^2 v t}{4\pi\epsilon_0 m_e [b^2 + (\gamma v t)^2]^{2/3}}, \\ a_{\perp} = \dot{v}_z &= -\frac{eE_z}{m_e} \frac{\gamma Z_e^2 b}{4\pi\epsilon_0 m_e [b^2 + (\gamma v t)^2]^{2/3}}, \end{aligned} \quad (4.8)$$

where Z_e is the charge of the nucleus, Longair (2011). Then we can take the Fourier transformation of the acceleration equations Eq.(4.8):

$$\begin{aligned} \dot{v}_x(\omega) &= \frac{1}{(2\pi)^{1/2}} \int_{-\infty}^{\infty} \frac{\gamma Z_e^2 v t}{4\pi\epsilon_0 m_e [b^2 + (\gamma v t)^2]^{2/3}} \exp(i\omega t) dt, \\ \dot{v}_z(\omega) &= \frac{1}{(2\pi)^{1/2}} \int_{-\infty}^{\infty} \frac{\gamma Z_e^2 b}{4\pi\epsilon_0 m_e [b^2 + (\gamma v t)^2]^{2/3}} \exp(i\omega t) dt. \end{aligned} \quad (4.9)$$

Changing variables to $x = \gamma v t / b$ will leads to:

$$\begin{aligned} \dot{v}_x(\omega) &= \frac{1}{(2\pi)^{1/2}} \frac{Z_e^2}{4\pi\epsilon_0 m_e} \frac{1}{\gamma b v} I_1(y), \\ \dot{v}_z(\omega) &= \frac{1}{(2\pi)^{1/2}} \frac{Z_e^2}{4\pi\epsilon_0 m_e} \frac{1}{b v} I_2(y), \end{aligned} \quad (4.10)$$

where $y = \omega b / \gamma v$. The integrals $I_1(y)$ and $I_2(y)$ are

$$I_1(y) = 2iyK_0(y), \quad I_2(y) = 2yK_1(y) \quad (4.11)$$

where K_0 and K_1 are modified Bessel functions of order zero and one, Longair (2011). The radiation spectrum of the electron encountering a charged nucleus therefore looks like:

$$\begin{aligned} I(\omega) &= \frac{e^2}{3\pi\epsilon_0 c^3} \left[|a_{\parallel}(\omega)|^2 + |a_{\perp}(\omega)|^2 \right] \\ &= \frac{Z^2 e^6}{24\pi^4 \epsilon_0^3 c^3 m_e^2 v^2} \frac{\omega^2}{\gamma^2 v^2} \left[\frac{1}{\gamma^2} K_0^2 \left(\frac{\omega b}{\gamma v} \right) + K_1^2 \left(\frac{\omega b}{\gamma v} \right) \right] \end{aligned} \quad (4.12)$$

where b is the collision parameter.

The radiation spectrum showing both parallel and perpendicular accelerations to the motion direction of the electron is shown in fig.4.2. The perpendicular impulse has higher contribution, even in non-relativistic case, where $\gamma = 1$. In addition, this component cause significant radiation at low frequencies.

It is very useful to examine the behavior of the asymptotic limits of $K_0(y)$ and $K_1(y)$:

$$\begin{aligned} y \ll 1 \quad K_0(y) &= -\ln y \quad K_1(y) = 1/y, \\ y \gg 1 \quad K_0(y) &= K_1(y) = (\pi/2y)^{1/2} \exp(-y). \end{aligned} \quad (4.13)$$

There is an exponential cut-off in high frequencies radiation spectrum

$$I(\omega) = \frac{Z^2 e^6}{48\pi^3 \epsilon_0^3 c^3 m_e^2 v^2} \frac{\omega}{\gamma v b} \left[\frac{1}{\gamma^2 + 1} \right] \exp\left(-\frac{2\omega b}{\gamma v}\right), \quad (4.14)$$

because of the duration of the relativistic collision, which is roughly $\tau = 2b/\gamma v$. According to this, the dominant Fourier corresponds to $v \approx 1/\tau = \gamma v/2b$ and therefore to $\omega \approx \pi v \gamma/b$. Because $\omega \approx \gamma v/b$ there is only a little power emitted at grater frequencies then ω .

The low frequency spectrum has the form:

$$I(\omega) = \frac{Z^2 e^6}{24\pi^4 \epsilon_0^3 c^3 m_e^2 v^2} \frac{1}{b^2} \left[1 + \frac{1}{\gamma^2} \left(\frac{\omega b}{\gamma v} \right)^2 \ln^2 \left(\frac{\omega b}{\gamma v} \right) \right], \quad (4.15)$$

In the limit case $\omega b/\gamma v \ll 1$, the second term in square brackets of Eq.(4.15) is negligible and hence the good approximation for low frequency intensity spectrum is

$$I(\omega) = \frac{Z^2 e^6}{24\pi^4 \epsilon_0^3 c^3 m_e^2 v^2} = K. \quad (4.16)$$

We could guessed that the low frequency bremsstrahlung spectrum would be flat if we consider that the momentum impulse is a delta function, thus the collision duration is much less then the waves period. The Fourier transform of a delta function is the flat spectrum $I(\omega) = \text{constant}$. The spectrum will be flat approximation up to the frequency $\omega = \gamma v/b$, above this, it will falls of exponentially. Note also that γ disappeared from Eq.(4.16), even in the relativistic case. This means that the momentum impulse is the same in the relativistic and non-relativistic cases, for detailed derivation see Longair (2011), Jackson (1999).

Now we can integrate over all the collision parameters which contribute to the radiation at frequency ω :

$$I(\omega') = \int_{b'_{min}}^{b'_{max}} 2\pi b' \gamma N v K db' = \frac{Z^6 e^6 \gamma N}{12\pi^3 \epsilon_0^3 c^3 m_e^2} \frac{1}{v} \ln \left(\frac{b'_{max}}{b'_{min}} \right), \quad (4.17)$$

note that if the electron is moving by relativistic speed, because of the relativistic length contractions, it observed more nuclei enhanced by a factor γ . $N' = \gamma N$, where N is nuclear density .

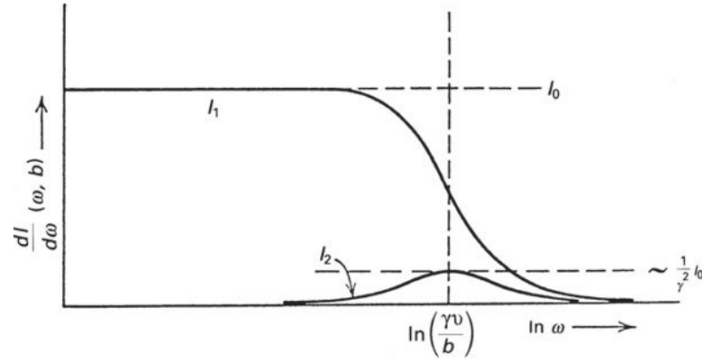


Figure 4.2: The bremsstrahlung spectrum from the acceleration of the electron parallel and perpendicular to its initial direction of motion as was displayed in Jackson (1999), or Rybicki & Lightman (1979).

4.1.2 Thermal bremsstrahlung

In previous sub-section we derive the bremsstrahlung radiation spectrum for one electron, but in real plasma at temperature T is necessary to integrate single partition spectrum over the collision parameters and over the Maxwell distribution of electron velocities

$$N_e(v)dv = 4\pi N_e \left(\frac{m_e}{2\pi kT} \right)^{2/3} v^2 \exp \left(-\frac{m_e v^2}{2kT} \right) dv. \quad (4.18)$$

The approximate expression for the spectral emissivity of a plasma of electron density N_e in the low frequency limit is

$$I(\omega) \approx \frac{Z^2 e^6 N N_e}{12\sqrt{3}\pi^3 \epsilon_0^3 c^3 m_e^2} \left(\frac{m_e}{kT} \right)^{1/2} g(\omega, T), \quad (4.19)$$

where $g(\omega, T)$ is known as Gaunt factor. It is important to realize that the low frequency thermal bremsstrahlung spectrum is almost independent of frequency, the only dependence upon ω comes from the slightly varying function in the Gaunt factor.

At the high frequencies the spectrum of thermal bremsstrahlung exponentially cuts off according to $\exp(-\hbar\omega/kT)$. This behavior is reflecting the exponential decrease in the electrons population in hard energy tail of Maxwell distribution.

Finally, we can get the total energy loss by the plasma by integrating the spectral emissivity over all possible frequencies, because of the exponential cut-off, the integration should be helds from 0 to $\omega = kT/\hbar$ which leads to

$$-\frac{dE}{dt} = (\text{constant}) Z^2 T^{1/2} \bar{g} N N_e. \quad (4.20)$$

Detailed calculations give the following results, in terms of frequency ν rather than the angular frequency ω . The spectral plasma emissivity according to Longair

(2011) is

$$K_\nu = \frac{1}{3\pi^2} \left(\frac{\pi}{6}\right)^{1/2} \frac{Z^2 e^6}{\epsilon_0^3 c^3 m_e^2} \left(\frac{m_e}{kT}\right)^{1/2} g(\nu, T) N N_e \exp\left(-\frac{h\nu}{kT}\right), \quad (4.21)$$

$$= 6.8 \times 10^{-51} Z^2 T^{-1/2} N N_e g(\nu, T) \exp\left(-\frac{h\nu}{kT}\right) \text{ W m}^{-3} \text{ Hz}^{-1},$$

where both: densities of electrons N_e and of nuclei N are in particles per m^3 . At frequencies $h\nu \ll kT$, the Gaunt factor has only a logarithmic dependence on the frequency. Therefore for radio and X-ray wavelengths are suitable following forms:

$$\begin{aligned} \text{Radio : } g(\nu, T) &= \frac{\sqrt{3}}{2\pi} \left[\ln \left(\frac{128 \epsilon_0^2 k^3 T^3}{m_e e^4 \nu^2 Z^2} \right) - \gamma^{1/2} \right], \\ \text{X-ray : } g(\nu, T) &= \frac{\sqrt{3}}{2\pi} \ln \left(\frac{kT}{h\nu} \right), \end{aligned} \quad (4.22)$$

where $\gamma = 0.577$ is the Euler's constant.

The total loss rate of the plasma is

$$-\left(\frac{dE}{dt}\right)_{\text{brems}} = 1.435 \times 10^{-40} Z^2 T^{1/2} \bar{g} N N_e \text{ W m}^{-3}. \quad (4.23)$$

Detailed calculation show that the average value of the Gaunt factor \bar{g} lies in the range 1.1 – 1.5, then $\bar{g} = 1.2$ will be a good approximation (Longair (2011)). To apply a suitable Gaunt factors in the thermal bremsstrahlung formula is nontrivial and complex task. Several useful papers about suitable Gaunt factors were published. One of the more recent is that from Sutherland (1998).

4.2 PSR model

By using previous results and assumptions, we can now find an explicit column solution. We specified that the cooling process in PSR is thermal bremsstrahlung and also that other mechanisms such Compton and cyclotron cooling are negligible in IPs case. We know that the electron and ion temperature will remain very stable. We also expect, that shock height (see fig. 4.1 where $z_0 - R_{wd} = D$ is shock distance from WD surface) is $D < R_{wd}$, therefore we can adopt one dimensional geometry with vertical coordinate z as is showed on the fig. 4.1.

The structure of the stationary PSR in plane-parallel one-dimensional geometry is described (see e.g. Suleimanov et al. (2005), Cropper et al. (1999) or Frank et al. (2002)) by the mass continuity equation:

$$\frac{d}{dz}(\rho v) = 0, \quad (4.24)$$

the momentum equation

$$\frac{d}{dz}(\rho v^2 + P) = -\frac{GM_{wd}}{z^2} \rho, \quad (4.25)$$

the energy equation

$$v \frac{dP}{dz} + \gamma P \frac{dv}{dz} = -(\gamma - 1)\Lambda, \quad (4.26)$$

and also by the ideal gas law:

$$P = \frac{\rho k T}{\mu m_H}. \quad (4.27)$$

where z is a spatial coordinate (see fig.4.1), v is the gas velocity, ρ is the gas density, T is the temperature, P is the gas pressure, $\mu = 0.62$ is the mean molecular weight of the fully ionized solar composition matter and $\gamma = 5/3$ is the adiabatic index. Λ is the cooling rate:

$$\Lambda = \left(\frac{\rho}{\mu m_H} \right)^2 \Lambda_N(T), \quad (4.28)$$

where the cooling function $\Lambda_N(T)$ for low densities plasma with $\mu = 0.62$ is tabulated by Sutherland & Dopita (1993).

The Eq.(4.24) can be interested with result:

$$\rho v = a, \quad (4.29)$$

where the a is local mass accretion rate of dimension $\text{g cm}^{-2} \text{s}^{-1}$. This is the free parameter of the model, but according to Suleimanov et al. (2005) the spectrum of emergent radiation in PSR is rather insensitive to this parameter with reasonable limits.

By using the substitution $z' = z_0 - z$ where z_0 is the shock-wave coordinate (see fig. 4.1), and with $g(z') = (GM_{wd})/(z_0 - z')^2$ we should rewrite the Eq.(4.25) and Eq.(4.26) with knowledge of Eq.(4.29):

$$\frac{dv}{dz'} = g(z') \frac{1}{v} - \frac{1}{a} \frac{dP}{dz'}, \quad (4.30)$$

$$\frac{dP}{dz'} = \frac{(\gamma - 1)\Lambda a + g(z')\gamma P a/v}{\gamma P - a/v}. \quad (4.31)$$

The equations (4.30) and (4.31) could be solved by using sufficient numerical method, e.g Suleimanov et al. (2005) and Cropper et al. (1999) used shooting method from $z = z_0 (z' = 0)$ to $z = R_{wd} (z' = z_0 - R_{wd})$ with boundary condition at $z = z_0$ with results:

$$\begin{aligned} v_0 &= 0.25 \sqrt{2GM_{wd}/z_0}, \\ \rho_0 &= a/v_0, \\ P_0 &= 3av_0, \\ T_0 &= 3 \frac{\mu m_H}{k} v_0^2. \end{aligned} \quad (4.32)$$

The equation which ties up the R_{wd} and M_{wd} is typically used that one from the Nauenberg (1972) article:

$$R_{wd} = 7.8 \times 10^8 \left[\left(\frac{1.44 M_{\odot}}{M_{wd}} \right)^{2/3} - \left(\frac{M_{wd}}{1.44 M_{\odot}} \right)^{2/3} \right]^{1/2} \text{ cm}. \quad (4.33)$$

The simple analytical model from the Frank et al. (2002), based on assumption of the constant pressure in the PSR region can by also used as follows:

$$\begin{aligned} T(z) &= T_0 \left(\frac{z - R_{wd}}{z_0 - R_{wd}} \right)^{2/5}, \\ \rho(z) &= \rho_0 \left(\frac{z - R_{wd}}{z_0 - R_{wd}} \right)^{-2/5}. \end{aligned} \quad (4.34)$$

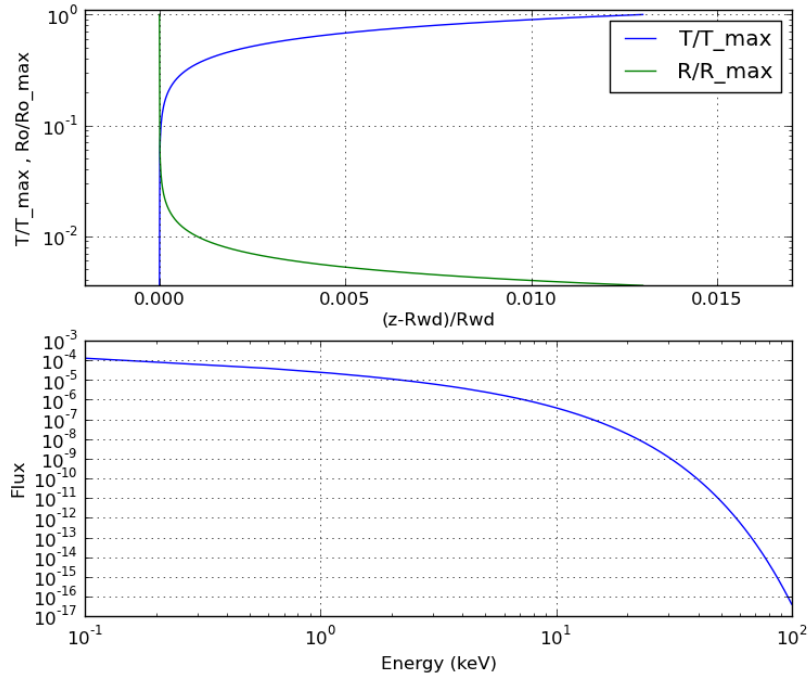


Figure 4.3: *Top*: the temperature and density profiles for the analytical model from Frank et al. (2002). The input parameters for script was: $M_{wd} = 0.7 M_{\odot}$ and shock distance from WD surface $0.013 R_{wd}$, $a = 1.0 \text{ cm}^{-2} \text{ s}^{-1}$, The corresponding WD radius is $\sim 7797 \text{ km}$, while the resulting $kT_{\max} = 29.71 \text{ keV}$. *Bottom*: the emergent model spectrum calculated by summing of the local bremsstrahlung spectra.

The emergent model spectrum will be calculated by summing of the local

bremsstrahlung spectra:

$$F_E = \int_{R_{wd}}^{z_0} j(z) dz, \quad (4.35)$$

where the local spectra were taken in form (Zombeck (1990), Suleimanov et al. (2005)):

$$j(z) = 9.52 \times 10^{-38} \left(\frac{\rho(z)}{\mu m_H} \right)^2 T(z)^{-1/2} \left(\frac{E}{kT(z)} \right)^{-0.4} \exp \left(-\frac{E}{kT(z)} \right). \quad (4.36)$$

4.2.1 WD mass estimations methods

$$kT_s = 16 \times \left(\frac{M}{0.5M_\odot} \right) \left(\frac{R}{10^9 \text{cm}} \right)^{-1} \text{keV} \quad (4.37)$$

$$\frac{M_{wd}}{M_\odot} = 1.44 \times \left[\frac{1}{2} \left(1 + \sqrt{1 + 4 \times \left(\frac{59}{kT_s} \right)^2} \right) \right]^{-3/4} \quad (4.38)$$

Data reduction and analysis

In this chapter we describe the results and data analyze threads for both XMM–Newton and INTEGRAL data. All data, software and online utilities necessary to finding the data, download and analyze are free and runs under *Linux*.

5.1 INTEGRAL

INTErnational Gamma-Ray Astrophysics Laboratory (INTEGRAL) is the European Space Agency’s satellite, launched in 2002 from Baikonur Cosmodrome on-board Proton-K launcher. INTEGRAL lies on orbit with period 72 hours where the large fraction of its orbit is behind radiation belts. The INTEGRAL service module is clone of the XMM-Newton one which significantly reduced construction costs. INTEGRAL has four scientific instruments on-board, the IBIS/ISGRI is used in this work.

- **IBIS** (Imager on-Board the INTEGRAL Satellite) observes in broad band 15 keV – 10 MeV.
- **SPI** (SPectrometer for INTEGRAL) observes in 20 keV – 8 MeV, part of the SPI is also ACS (AntiCoincidence Shield).
- **JEM-X** observed in 3 – 35 keV.
- **OMC** is an optical monitor sensitive in 500 – 580 nm.

5.1.1 IBIS/ISGRI

IBIS is a coded mask gamma ray telescope operating in very broad energy range, from the hard X-rays to gamma-rays, however its sensitivity strongly decreases

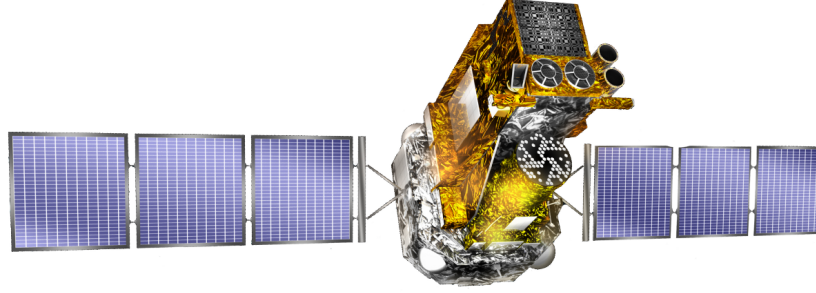


Figure 5.1: ESA/INTEGRAL spacecraft with four scientific instrument and 4 ton launch mass is one of the largest European satellites.

with the increasing energy bands. It is an imaging detector which uses coded mask. The mask itself its a special coded pattern, that chosen for the IBIS detector is based of a cyclic replication of MURA (Modified Uniform Redundant Array) of order 53, this detection technology is describes in Gottesman (1955).

Table 5.1: Scientific parameters of the IBIS instrument on-board INTEGRAL space mission

Operating energy range	15 keV – 10 MeV
Energy resolution (FWHM)	7% at 100 keV 9% at 1 MeV
Effective area	ISGRI: 960 cm ² at 50 keV PICsIT: 870 cm ² at 300 keV (single events) PICsIT: 275 cm ² at 1 MeV (multiple events)
FOV	8.3° × 8.0° (fully coded) 19° × 19° (partially coded, 50%)
Angular resolution (FWHM)	12'
Point source location accuracy 90% error radius	30'' at 100 keV < 5' at 1 MeV
Continuum sensitivity, photons cm ⁻² s ⁻¹ keV ⁻¹ (3σ detection, $\Delta E = E/2$, 10 ⁶ s integration)	3.8 × 10 ⁻⁷ at 100 keV 1 – 2 × 10 ⁻⁷ at 1 MeV
Narrow line sensitivity, photons cm ⁻² s ⁻¹ (3σ, 10 ⁶ s integration)	1.3 × 10 ⁻⁵ at 100 keV 4 × 10 ⁻⁵ at 1 MeV
Absolute timing accuracy (3σ)	ISGRI: 61 μs PICsIT: 0.976 – 500 ms (selected from the ground)

The IBIS is two layers detector, where the first layer is called ISGRI and second PICsIT. ISGRI is based on the cadmium telluride (CdTe), which is semiconductor operating at normal temperatures. The detector layer is made of 8 MDUs /footnote-Modular Detector Unit each having 32 × 64 pixels. The lower detector, PICsIT is based on cesium iodide (CsI) crystals.

The coded mask produced a shadowgram, the photon from the image are dis-

tributed across the entire detector, but cross-correlation techniques are used for taking a full reconstructed image in fully coded FOV. For reconstruct the image in partially coded FOV is necessary to use special cleaning techniques.

5.1.2 Data reduction

For the IBIS/ISGRI data reduction we used the 64-bit version of the Off-line Science Analysis (OSA 9.0) software which can be downloaded from INTEGRAL Science Data Center web-pages¹. It is necessary to download also the instrument calibration and response data and the latest objects reference catalogue. For suitable data analysis in high energy astrophysics is necessary to use also several typical software packages as: `ds9`, `fv`, `FTOOLS`². For the reduction itself we used following set of commands in the sequence:

1. `rbnrmf` was used for creating a rebinning response matrix for 10 spectral bins in 20–100 keV
2. `og_create` was used for creating observation groups
3. `ibis_science_analysis` package was used for image reconstruction and creating mosaics
4. `mosaic_spec` was used for the spectral extraction from the final mosaics

5.2 XMM-Newton

The European Space Agency's (ESA) X-ray Multi-Mirror Mission (XMM-Newton) was launched by an Ariane 504 on December 10th 1999. The XMM weighs 3.8 tons and carries three X-ray telescopes each contains 58 Wolter I grazing – incidence mirrors. The instruments on-board are three European Photon Imaging Cameras (EPIC) and two Reflection Grating Spectrometers (RGS).

Two of the cameras are Metal Oxide Semi-conductor (MOS) CCD arrays. Third camera is pn CCD camera. There is also an optical telescope on-board, the 30 cm diameter Ritchey—Chrétien observing in the UV/optical band.

In each MOS there are seven EEV type 22 front illuminated CCDs. The six surrounding CCDs are stepped towards because of the focal plane curvature while the middle one is at the focal point of the X-ray telescope. The MOS cameras are sensitive in the energy range 0.2 to 10 keV. All the EPIC CCDs operate in photon counting mode. In the case of MOS cameras, the central CCD can be operated separately.

¹<http://www.isdc.unige.ch/integral/>

²See <http://heasarc.nasa.gov/lheasoft/>

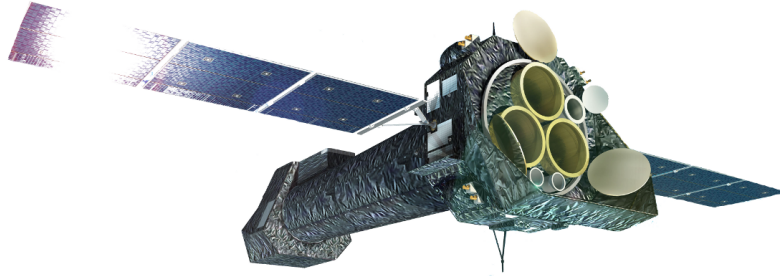


Figure 5.2: The XMM-Newton (X-ray Multi-Mirror Mission) is one off the most successful space mission of all time.

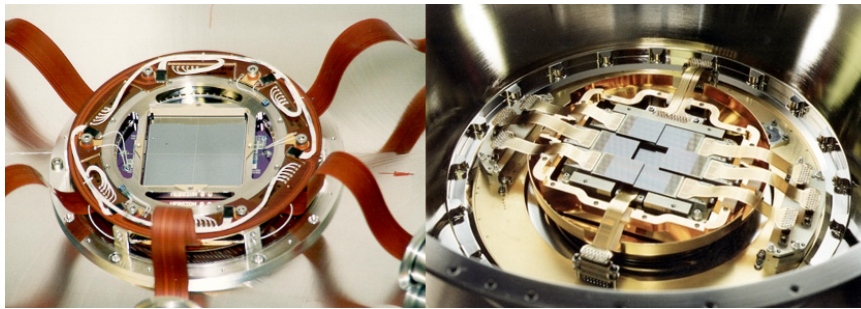


Figure 5.3: EPIC – European Photon Imaging Camera consists of (*left image*) pn ccd and (*right image*) two MOS CCDs.

5.2.1 Data reduction

For the data acquisition was used XMM-Newton Science Archive (XSA)³ and candidates was selected from the Koji Mukai online catalog of IPs and IP candidates⁴. From the catalogue of 92 IPs/IP candidates and 33 others systems was selected only strongly confirmed (in catalogue marked as ★★★★★) and confirmed (marked as ★★★★) systems.

This leads to selection of 25 sources:

³<http://xmm.esac.esa.int/xsa/>

⁴<http://asd.gsfc.nasa.gov/Koji.Mukai/iphome/catalog/alpha.html>

1RXS J005528.0+461143	1RXS J070407.9+262501	1RXS J180340.0+401214
AO Psc	BY Cam	EI UMa
EX Hya	FO Aqr	GK Per
HT Cam	IGR J00234+6141	IGR J15094-6649
IGR J17195-4100	IGR J17303-0601	IGR J21335+5105
MU Cam	NY Lup	PQ Gem
SDSS J233325	UU Col	V1223 Sgr
V2069 Cyg	V2400 Oph	V405 Aur
XY Ari		

For all the public available pointings for all the 25 sources the following data reduction procedures took place. The XMM-Newton Science Analysis System (SAS) version 11.0.0 for Feora 32-bit Linux was used and executed by simple bash skript:

```
#!/bin/bash
#this script should run simple xmm MOS spectral analysis for future spectral extraction
alias ldir="ls -l | egrep '^d'"
alias xmm32='. /home/mkocka/diplomka/soft/sas32/xmmsas_20110223_1801/setsas.sh'
export HEADAS=/home/mkocka/diplomka/soft/heasoft-6.11/x86_64-unknown-linux-gnu-libc2.14
export SAS_CCFPATH=/home/mkocka/diplomka/soft/sas/xmmsas_20110223_1801/ccf
alias heainit=". $HEADAS/headas-init.sh"
cd /home/mkocka/diplomka/work/xmm1
export LDR='pwd'
heainit
xmm32
cd $SAS_DIR
export LD_LIBRARY_PATH=$LD_LIBRARY_PATH:'pwd'/libextra
export PATH=$PATH:/home/koci/XMM/xmmsas_20100423_1801/bin
cd $LDR
for i in *
do cd $i
for j in `ldir | awk '{print $9}'`
do cd $j/odf
export SAS_ODF='pwd'
if [ -f *.SAS ]; then rm *.SAS; fi
cifbuild
export SAS_CCF='pwd'/ccf.cif
odfingest
cd ..
mkdir MOS
cd MOS
emchain
mos-filter
for k in `ls *clean*`
do evselect table=$k imagebinning=binSize imageset=image_$k \
withimageset=yes xcolumn=X ycolumn=Y ximagebinsize=64 yimagebinsize=64
done
cd ../../
done
cd ..
done
```

In the next step of the analysis thread was checked which of the 25 IPs observed by XMM has been also observed by INTEGRAL and have enough exposures for necessary spectral extraction. The seven sources was been selected for spectral extraction.

The spectra from all suitable MOS observations was taken by standard SAS data product extraction. The following commands took place:

1. `evselect` was used for extraction of the source and background spectrum.
2. `rmfgen` was used for creating a redistribution matrix for the previously extracted spectrum.
3. `arfgen` was used for generate an ancillary file.
4. `specgroup` was used for rebinig the spectrum and link associated files.

5.3 Data analysis

5.4 Results

5.5 Discussion

Table 5.2: Summary of the observations. Columns give the object name, coordinates of the source, XMM-Newton observation ID, total exposures times for MOS1 and MOS2, total exposures for IBIS/ISGRI, and period of INTEGRAL observations, the IPs are ordering according to their RA.

obj.	Ra [°]	Dec [°]	XMM		Exp. time [ks]		Exp. time [ks]		Obs. period
			obs. ID	obs. ID	MOS1	MOS2	IBIS/ISGRI	IBIS/ISGRI	
IGR J15094-6649	227.358	-66.823	0551430301	0551430301	31.61	31.62	592.46	1999-12-31 — 2008-07-24	
NY Lup	237.060	-45.479	0105460301	0105460301	21.40	21.40	1589.29	1999-12-31 — 2005-04-08	
V2400 Oph	258.152	-24.245	0105460101	0105460101	10.07	10.12	1368.89	1999-12-31 — 2009-03-07	
IGR J17195-4100	259.899	-41.014	0601270201	0601270201	33.62	33.62	878.12	1999-12-31 — 2008-10-23	
IGR J17303-0601	262.589	-5.992	0302100201	0302100201	13.28	13.28	248.62	1999-12-31 — 2008-04-21	
V1223 Sgr	283.759	-31.163	0145050101	0145050101	38.66	38.67	264.66	1999-12-31 — 2006-11-09	
IGR J21335+5105	323.375	51.092	0302100101	0302100101	16.59	16.40	1589.68	1999-12-31 — 2009-05-22	

Conclusions

Bibliography

- Adams, D. 1979, *The Hitchhiker's Guide to the Galaxy* (Great publishing house of Ursa Minor Beta)
- Aizu, K. 1973, *Progress of Theoretical Physics*, 49, 1184
- Arnaud, Smith, S. 2011, *Handbook of X-ray Astronomy* (Cambridge University Press)
- Aungwerojwit, A. 2007, PhD Thesis
- Brunschweiler, J., Greiner, J., Ajello, M., & Osborne, J. 2009, *A&A*, 496, 121
- Butters, O. W., Norton, A. J., Mukai, K., & Tomsick, J. A. 2011, *A&A*, 526, A77
- Camenzind, M. 2007, *Compact Objects in Astrophysics* (Springer)
- Carroll, B. W. & Ostlie, D. A. 2007, *An Introduction to Modern Astrophysics*, second edition (Pearson Inc.)
- Chandrasekhar, S. 1939, *MNRAS*, 99, 673
- Cropper, M. 1990, *Space Science Reviews*, 54, 195
- Cropper, M., Wu, K., Ramsay, G., & Kocabiyik, A. 1999, *MNRAS*, 306, 684
- Forman, W., Jones, C., Cominsky, L., Julien, P., Murray, S., Peters, G., Tananbaum, H., & Giacconi, R. 1978, *ApJS*, 38, 357
- Frank, J., King, A., & Raine, D. 2002, *Accretion Power in Astrophysics Third Edition* (University Press, Cambridge)
- Gottesman, S. R., F. E. F. 1955, *Imaging in High Energy Astronomy*, 177

- Hamada, T. & Salpeter, E. E. 1961, *ApJ*, 134, 683
- Jackson, J. D. 1999, *Classical Electrodynamics* (New York: John Wiley)
- Kitchin, C. R. 2009, *Astrophysical Techniques* fifth edition (Taylor & Francis Group)
- Kleczek, J. 1957, *Interior of Stars* (Press of Czechoslovak Academy of Science)
- Krivos, R. & Revnivtsev, M. 2011, *Galactic Difuse Emission slides from presentation* (The Extreme and Variable High Energy Sky, Workshop)
- Krivos, R., Revnivtsev, M., Churazov, E., Sazonov, S., Grebenev, S., & Sunyaev, R. 2007, *A&A*, 463, 957
- Landi, R., Bassani, L., Dean, A. J., Bird, A. J., Fiocchi, M., Bazzano, A., Nousek, J. A., & Osborne, J. P. 2009, *MNRAS*, 392, 630
- Longair. 2011, *High Energy Astrophysics Third Edition* (University Press, Cambridge)
- Middleton, M. J., Cackett, E. M., Shaw, C., Ramsay, G., Roberts, T. P., & Wheatley, P. J. 2012, *MNRAS*, 419, 336
- Mirabel, I. F. 2002, *ASP Conference Series*
- Nauenberg, M. 1972, *ApJ*, 175, 417
- Padmanabhan, T. 2001, *Theoretical Astrophysics Volume II: Stars and Stellar Systems* (Cambridge University Press)
- Patterson, J. 1994, *Publications of the Astronomical Society of the Pacific*, 106, 209
- Ramsay, G., Cropper, M., Hellier, C., & Wu, K. 1998, *MNRAS*, 297, 1269
- Revnivtsev, M., Sazonov, S., Krivos, R., Ritter, H., & Sunyaev, R. 2008, *A&A*, 489, 1121
- Ritter, H. & Kolb, U. 2003, *A&A*, 404, 301
- Rothschild, R. E., Gruber, D. E., Knight, F. K., Matteson, J. L., Nolan, P. L., Swank, J. H., Holt, S. S., Serlemitsos, P. J., Mason, K. O., & Tuohy, I. R. 1981, *ApJ*, 250, 723
- Rybicki, G. B. & Lightman, A. P. 1979, *Radiative Processes in Astrophysics* (John Wiley & Sons, Inc.)
- Sazonov, S., Revnivtsev, M., Gilfanov, M., Churazov, E., & Sunyaev, R. 2006, *A&A*, 450, 117

- Shapiro, S. L. & Teukolsky, S. A. 2004, Black holes, white dwarfs, and neutron stars: The physics of compact objects, ed. Shapiro, S. L. & Teukolsky, S. A.
- Suh, I.-S. & Mathews, G. J. 2000, *ApJ*, 530, 949
- Suleimanov, V., Revnivtsev, M., & Ritter, H. 2005, *A&A*, 435, 191
- Sutherland, R. S. 1998, *MNRAS*, 300, 321
- Sutherland, R. S. & Dopita, M. A. 1993, *ApJS*, 88, 253
- Warner, B. 1995, *Cataclysmic Variable Stars* (University Press, Cambridge)
- Yuasa, T., Nakazawa, K., Makishima, K., Saitou, K., Ishida, M., Ebisawa, K., Mori, H., & Yamada, S. 2010, *A&A*, 520, A25+
- Zombeck, M. V. 1990, *Handbook of Astronomy and Astrophysics* (Cambridge University Press)

Appendix

this will be the appendix

Table 1: Estimated WD masses from previous reports ...

System	Suzaku XIS+HXD M_{WD}	Swift BAT M_{WD}	RXTE PCA+HEXTE M_{WD}	RXTE PCA M_{WD}	Ginga LAC M_{WD}	ASCA SIS M_{WD}	This work XMM & Integral M_{WD}
NY Lup							
V2400 Oph							
V1223 Sgr							
IGR J17195-4100							
IGR J21335+5105							
IGR J15094-6649							
IGR J17303-0601							