

INTRODUCTION

Above Earth's atmosphere are the Van Allen radiation belts, a complex and dynamic plasma environment that can effect our technology-driven society. These effects include: a higher radiation dose for astronauts and cosmonauts, higher chance of spacecraft failure due to single event upsets that can lead to catastrophic latchups, cumulative degradation of silicon (changing the silicon doping) from an extended radiation dose that can degrade a transition to the point where it no longer function as a switch, and the degradation of the ozone layer due to the chemical production of NOX and HOX. With these effects in mind, it is no surprise that the radiation belts have been extensively studied since their discovery in the 1960s.

A topic of interest in the space physics community is wave-particle intersections that as we will see later in the introduction, can accelerate particles and scatter them into the atmosphere.

The goal of this dissertation is to study the wave-particle mechanism that scatters microbursts into Earth's atmosphere. This goal will be achieved by first introducing single charged particle motion in electric and magnetic fields, the major particle populations and how they couple in the magnetosphere, and then describe the history and current state of the fields relating to microbursts and wave-particle scattering

Charged Particle Motion in Electric and Magnetic Fields

A charged particle trapped in the magnetosphere will experience three types of periodic motion in Earth's nearly dipolar field. The three motions are ultimately due to the Lorentz force that a particle of momentum \vec{p} , charge q , and velocity \vec{v}

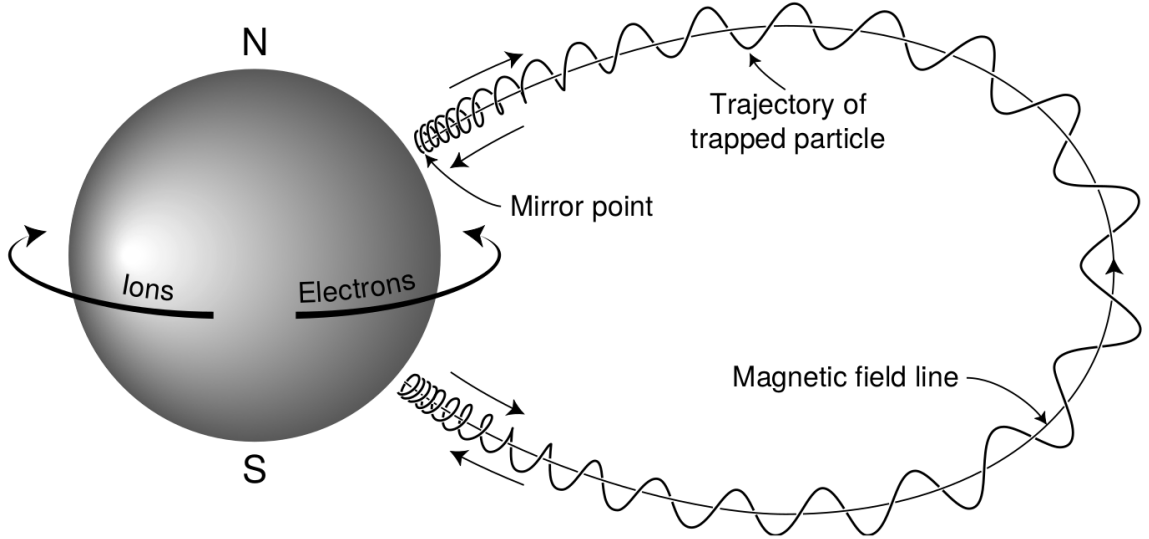


Figure 1.1: A diagram of the three motions that a charged particle experiences in Earth's dipole magnetic field. These motions are: gyration about the magnetic field line, bounce motion between the magnetic poles, and azimuthal drift around the Earth. Figure from (Baumjohann and Treumann, 1997).

experiences in an electric field \vec{E} and magnetic field \vec{B} and is given by

$$\frac{d\vec{p}}{dt} = q(\vec{E} + \vec{v} \times \vec{B}). \quad (1.1)$$

In the magnetosphere, the three periodic motions in decreasing frequency are gyration, bounce, and drift and are schematically shown in Fig. 1.1. Each of periodic these motions have a corresponding conserved quantity i.e. an adiabatic invariant.

The highest frequency periodic motion is on the order of a kHz in the magnetosphere and is gyration about a magnetic field of magnitude B . This motion is circular with a Larmor radius of

$$r = \frac{mv_{\perp}}{|q|B} \quad (1.2)$$

where m is the mass and v_{\perp} the particle velocity perpendicular to \vec{B} . This motion

has a corresponding gyrofrequency

$$\Omega = \frac{|q|B}{m} \quad (1.3)$$

in units of radians/second. The corresponding adiabatic invariant can be found by integrating the particle's canonical momentum around the particle's path

$$J_i = \oint (\vec{p} + q\vec{A}) \cdot d\vec{l} \quad (1.4)$$

where J_i is the i^{th} adiabatic invariant and \vec{A} is the magnetic vector potential. This integral is carried out by integrating the first term over the circumference of the gyro orbit and integrating the second term using Stokes theorem to calculate the magnetic flux enclosed by the gyro orbit. With suitable integration, $J_1 \sim v_\perp^2/B$ that is conserved as long as a driving force on the particle has a frequency much less than Ω .

The second highest frequency periodic motion is the bounce motion that is due to a parallel gradient in \vec{B} . This periodic motion naturally arises in the magnetosphere because the Earth's magnetic field is stronger near the poles, and artificially in the laboratory in magnetic bottle machines. First we need to define the concept of pitch angle α as the angle between the magnetic field and particle's velocity vectors. This is schematically shown in Fig. 1.2a. The pitch angle relates v with v_\perp , and v_\parallel (the component of the particles velocity parallel to \vec{B}). As shown in 1.2b and c, a larger α will tighten the particle's helix trajectory and vice versa.

Assuming that the particle's kinetic energy is concerned, the conservation of J_1 implies that given a particle's $v_\perp(0)$ and $B(0)$ at the magnetic equator, we can calculate its $v_\perp(s)$ along the particle's path with $B(s)$ from magnetic field models.

The perpendicular velocity is related via

$$\frac{v_{\perp}^2(0)}{B(0)} = \frac{v_{\perp}^2(s)}{B(s)} \quad (1.5)$$

which can be rewritten as

$$\frac{v^2 \sin^2 \alpha(0)}{B(0)} = \frac{v^2 - v_{\parallel}^2(s)}{B(s)} \quad (1.6)$$

and re-arranged to solve for $v_{\parallel}(s)$

$$v_{\parallel}(s) = v \sqrt{1 - \frac{B(s)}{B(0)} \sin^2 \alpha(0)} \quad (1.7)$$

which will tend towards 0 since $B(s) > B(0)$ in a dipole magnetic field.

The location where $v_{\parallel} = 0$ is called the mirror point where a particle stops and reverses direction. Since Earth's magnetic field is stronger towards the poles, the particle will execute periodic bounce motion between the two mirror points. The corresponding adiabatic invariant, J_2 is

$$J_2 = \oint p_{\parallel} ds \quad (1.8)$$

where ds describes the particle path between the mirror points in the northern and southern hemispheres (see Fig. 1.1). Substituting Eq. 1.7 into Eq. 1.8 and using the fact that at the mirror point the magnetic field strength is B_m and $\alpha(m) = 90$, J_2 can be written as

$$J_2 = 2p \int_{S \text{ mirror}}^{N \text{ mirror}} \sqrt{1 - \frac{B(s)}{B(m)}} ds \quad (1.9)$$

where $N \text{ mirror}$ and $S \text{ mirror}$ are the northern and southern mirror points, respectively. The bounce period can be estimated (e.g. Baumjohann and Treumann,

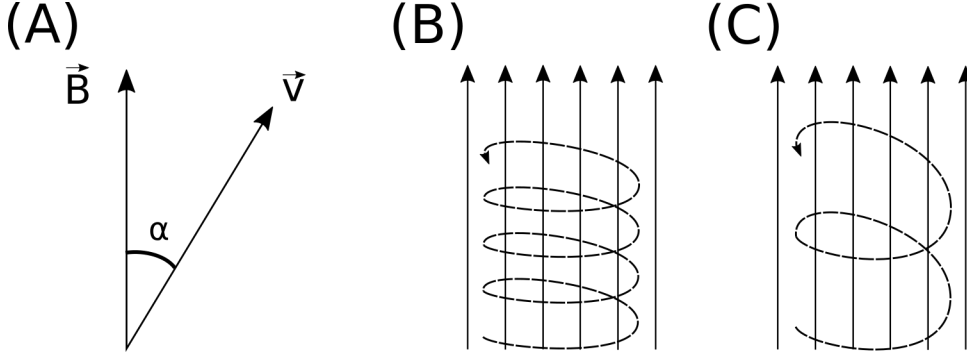


Figure 1.2: Charged particle motion in a uniform magnetic field \vec{B} . Panel (A) shows the geometry defining the pitch angle, α . Panel (B) and (C) show two helical electron trajectories assuming a large and small α (corresponding to a small and large parallel velocity v_{\parallel}), respectively.

1997) to be

$$t_b \approx \frac{LR_e}{\sqrt{W/m}}(3.7 - 1.6 \sin \alpha_0) \quad (1.10)$$

where L is the L -shell which describes the distance from the Earth's center to where a particular magnetic field line crosses the magnetic equator, in units of Earth radii, R_e . W is the particle's kinetic energy, and α_0 is the equatorial pitch angle.

At this stage it is instructional to introduce the notion of the loss cone angle, α_L . α_L is defined such that if a charged particle has a $\alpha < \alpha_L$ it will mirror at or below ≈ 100 km altitude in the atmosphere. A particle at those altitudes will encounter Earth's atmosphere and has a significant probability of Coulomb scattering with atmospheric particles and be removed from the magnetosphere.

The slowest periodic motion experienced by charged particles in Earth's magnetic field is azimuthal drift around the Earth. This drift results from a combination of a radial gradient in \vec{B} and the curvature of the magnetic field. The radial gradient drift can be easily justified because Earth's magnetic field is stronger near the Earth where the particle's gyroradius radius of curvature will be smaller as it gyrates towards

the Earth, and larger when it gyrates outward. The overall effect is the particle gyro orbit does not close on itself and negatively charged particles drift East and positively charged particles drift West. The drift due to the centrifugal force that a particle experiences as it bounces on curved field lines also increases the drift rate. The drift adiabatic invariant is found by integrating the J over the entire particle's orbit (drift shell) around the Earth. The first term is negligible, and the second term is the magnetic flux enclosed by the drift shell, Φ_m i.e. $J_3 \sim \Phi_m$.

Up until now the three drift motions conserve energy due to the absence of electric fields, \vec{E} . If \vec{E} is present, a particle's center of gyration (found by averaging over the gyro orbit) will drift in a direction perpendicular to both \vec{E} and \vec{B} . The drift velocity can be solved directly from Eq. 1.1 and is

$$\vec{v}_E = \frac{\vec{E} \times \vec{B}}{B^2}. \quad (1.11)$$

Lastly, detailed derivations of these motions can be found in (Baumjohann and Treumann, 1997; Schulz and Lanzerotti, 1974; Tsurutani and Lakhina, 1997).

Particle Populations and Their Interactions in the Magnetosphere

The single-particle motion described in the previous section was a prerequisite to understanding how these particles organize into macroscopic structures in the magnetosphere. The structure of the outer magnetosphere is shown in Fig. 1.4 and inner magnetosphere in Fig. 1.5. In this section we will discuss the macroscopic magnetosphere structures and how they couple.

The sun and its solar wind are ultimately the source of energy input into the magnetosphere. The solar wind at Earth is a plasma traveling at supersonic speeds with an embedded interplanetary magnetic field (IMF). When the solar wind



Figure 1.3: Contours of constant gyration, bounce, and drift frequencies for electrons and protons in a dipole field. Figure from Schulz and Lanzerotti (1974).



Figure 1.4: The large scale structures in the outer magnetosphere. The solar wind with its frozen-in interplanetary magnetic field is shown on the left and is traveling supersonically to the right. The solar wind envelops Earth's magnetic field to create the magnetosphere cavity. Since the solar wind is traveling supersonically, it creates a bow shock up stream. The shocked solar wind plasma in the magnetosheath flows around the magnetopause, a boundary between the solar wind and magnetosphere. Figure from Baumjohann and Treumann (1997).

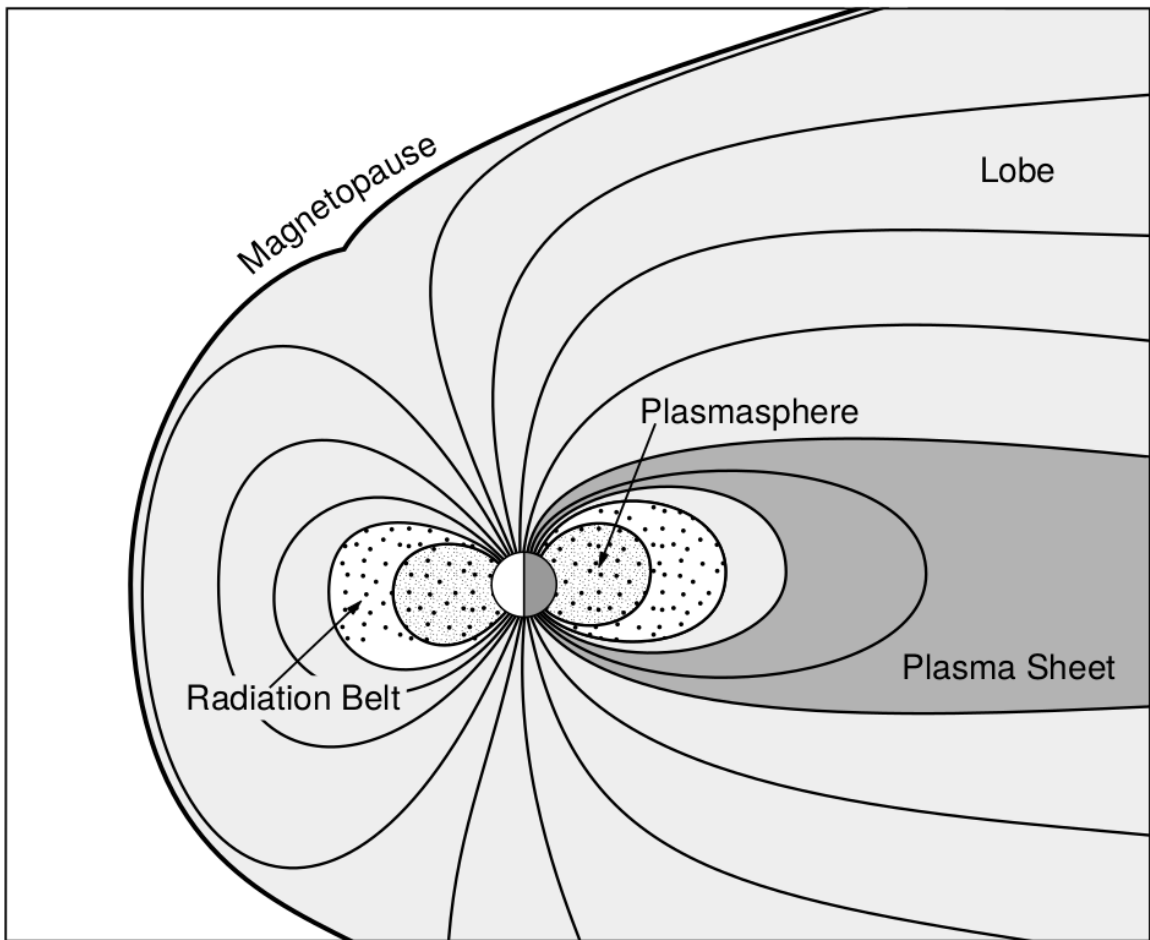


Figure 1.5: The large scale structures in the inner magnetosphere most relevant to this dissertation. The plasmasphere, and the radiation belts are shown and ring current is co-located there as well. Figure from Baumjohann and Treumann (1997).

encounters Earth's magnetic field the plasma can not easily penetrate into the magnetosphere, rather it drapes around the magnetosphere forming a cavity in the solar wind that is roughly shaped as shown in Figs. 1.4 and 1.5. Because the solar wind is supersonic at 1 AU, a bow shock exists upstream of the magnetosphere. The solar wind plasma, after it is shocked by the bow shock, then flows around the magnetosphere inside the magnetosheath. The surface where the solar wind ram and Earth's magnetic pressures balance is termed the magnetopause, which can be thought of as a boundary between the solar wind and Earth's plasma populations. This is a slightly naive description of the magnetopause, but is nonetheless an instructive conceptual picture. The shocked plasma then flows past the Earth where it carves out the magnetotail where the solar wind magnetic pressure balances Earth's magnetic field pressure in the lobes. The magnetotail extends on the order of $100 R_E$ downstream of Earth, and the tailward stretching of magnetic field lines creates the plasma sheet which exists in the region of low magnetic field strength near the magnetic equator. The plasma sheet flows from dusk to dawn (out of the page in Figs. 1.4 and 1.5) and this current is connected to a variety of other currents in the magnetosphere which is beyond the scope of this dissertation.

The idea of the magnetopause as a barrier between the solar wind and the magnetosphere is not entirely accurate due to the presence of reconnection. Reconnection was first conceived by Dungey (1961) who described the convection of Earth's magnetic field between the bow and tail regions of the magnetosphere. This process is known as the Dungey cycle and is most effective when the IMF is negative and pointing southward as is shown in 1 in Fig. 1.6. As the IMF contacts Earth's magnetic field it reconnects with it so that Earth's magnetic field is directly connected to the IMF. Then as the solar wind flows tailward the IMF drags Earth's magnetic field towards the magnetotail as shown in 2-6 in Fig. 1.6. As more and

more magnetic field lines are draped in the magnetotail, magnetic pressure squeezes the plasma sheet until Earth's magnetic field reconnects as is shown in 7 in Fig. 1.6. Lastly as 8 in Fig. 1.6 shows, the newly merged magnetic field line and the plasma on it moves Earthward under the magnetic tension force to become more dipolar. This is called a dipolarization of the magnetic field, and the plasma frozen on these field lines can be observed as injections (e.g. Turner et al., 2015).

Then since there can not be a depletion of Earth's magnetic field in the bow region, the rate of merging on the bow side must equal the rate of reconnection in the magnetotail. Reconnection in the magnetotail is shown in 7 in Fig. 1.6.

The Earthward movement of near perfectly conducting plasma to an observer on Earth will appear as a dawn-dusk convection electric field which with the co-rotation electric field, drive the dynamics for low energy particles.

Inner Magnetosphere Coordinates

Radiation Belts

The Van Allen radiation belts were discovered by Van Allen (1959) and Vernov and Chudakov (1960) during the Cold War.

No one had predicted the existence of Earths radiation belts nested tori of energetic particles trapped by the planets magnetic field

Particle Acceleration

Adiabatic Heating

Wave Resonance Heating

Particle Losses

Electromagnetic Ion Cyclotron Wave Driven

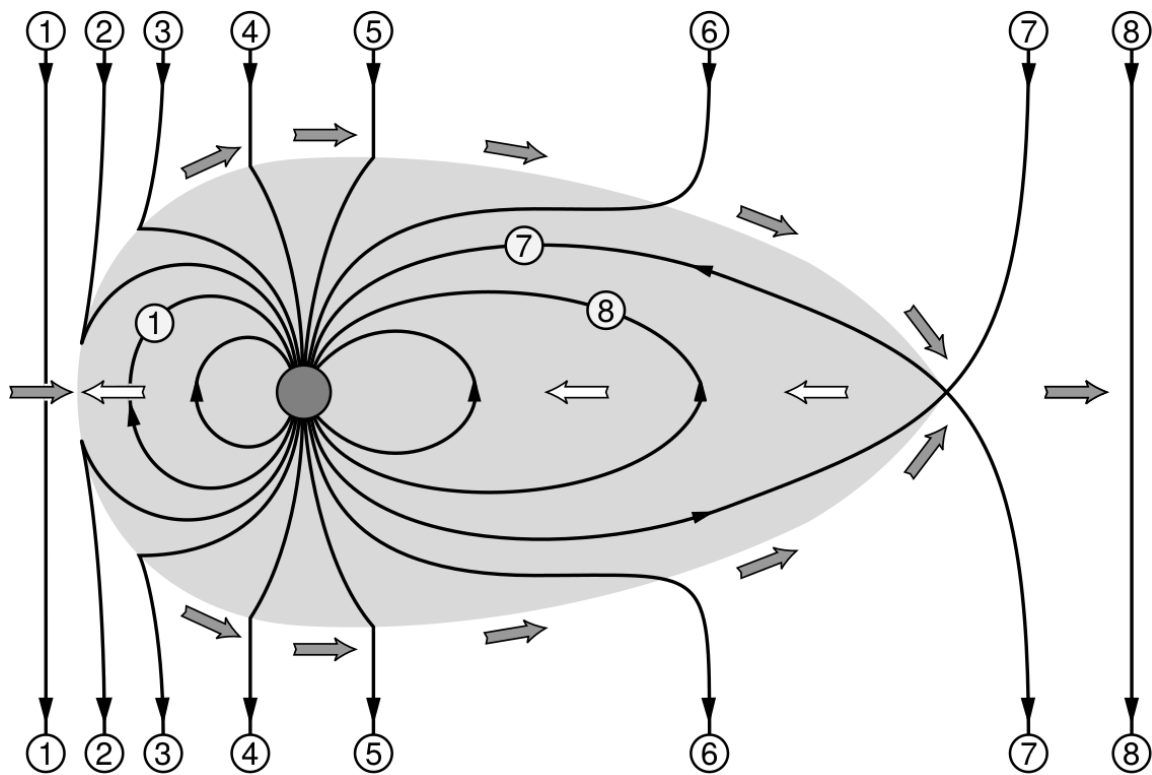


Figure 1.6: The series of steps involved in magnetic reconnection with a southward IMF. Figure from Baumjohann and Treumann (1997).

Whistler Mode Chorus Wave Driven

Scope of Reserach

This dissertation furthers our understanding of the microburst scattering mechanism and is organized into the following chapters. Chapter **X** will describe the spacecraft missions used to study microburst precipitation and wave-particle scattering. Then Ch. **Y** will describe a microburst scattering event observed by NASA's Van Allen Probes and the quasi-linear diffusion model that was developed. Next, Ch. **Z** will describe a bouncing packet microburst observation made by MSU's FIREBIRD-II mission where the microburst's lower bound longitudinal and latitudinal scale sizes were estimated. Chapter **ZZ** then expands the case study result from Ch. **Z** to a statistical study of microburst sizes and the microburst size models developed to interpret the data. Lastly, **ZZZ** will summarize the dissertation work and make concluding remarks about research to be done.

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