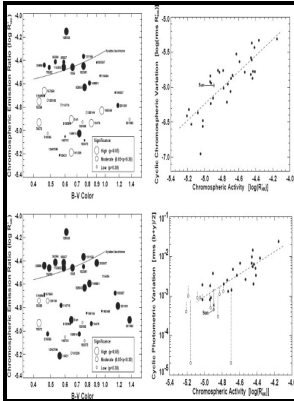


# Consequences of a chromospheric temperature gradient on the width of H [alpha] in late-type giants

Big Bear Solar Observatory, California Institute of Technology - Atmospheric Electrification in Dusty, Reactive Gases in the Solar System and Beyond



Description: -

-  
Animals -- Fiction  
Circus -- Fiction  
Temperature gradients.  
Optical properties.  
Chromosphere.  
Giant stars.  
Stellar chromospheres. Consequences of a chromospheric temperature gradient on the width of H [alpha] in late-type giants  
-  
[NASA contractor report] -- NASA CR-173359. Consequences of a chromospheric temperature gradient on the width of H [alpha] in late-type giants  
Notes: Microfiche. [Washington, D.C. : National Aeronautics and Space Administration], 1984. 1 microfiche.  
This edition was published in 1984



Filesize: 12.61 MB

Tags: #The #Astrophysical #Journal, #Volume #865, #Number #1, #2018 #September #20

**The Astrophysical Journal, Volume 865, Number 1, 2018 September 20**

The particle collisions are mediated by gravity acting on large particles and by turbulent convection within the cloud. Small wonder, then, that researchers have struggled with the apparently simple question, what is a chromosphere? The HETG spectrum will allow us to discriminate between X-rays from radiative shocks in the winds of one or both O stars and colliding wind shock emission in HD 150136, which is known to be a O + O spectroscopic binary.

## Solar and stellar photospheric abundances

In addition, close binaries lead to far more complex stellar evolution scenarios where the events expected for an isolated star may happen sooner, or never take place. The shell phases coincide with variation in the double emission peak separations, and both forms of variability are likely caused by a two-armed spiral density perturbation in the Be disk. The presence of the Regener—Pfitzer maximum results from a balance between the energy of the incoming particles, and the density of the atmosphere.

## Solar and stellar photospheric abundances

Since these phenomena were seen in H-alpha, we believe that magnetic reconnection in the lower solar atmosphere plays a causal role in solar flare eruptions. Its main instrument will be a for studying the dynamics of the solar corona.

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Richard Altmann Title: Long-Term Variation of the Rotation of the Solar Corona Authors: Altmann, R. However, the observed RMs are not as large as predicted by simplified analytic models that include substantial amplification of the Galactic magnetic field within the shell.

**The Astrophysical Journal, Volume 865, Number 1, 2018 September 20**

The bursty component is made up of the U type as well as the Reverse Drift RS bursts. The corona is the extended atmosphere of the Sun, which has a volume much larger than the volume enclosed by the Sun's photosphere. The absence of simultaneous measurements of filament and arcade topologies and fields leaves us with several observational pieces that do not readily fit together.

## **Sun**

With this somewhat revised view to a chromosphere, I review in Section the principal results from long-term observations of solar and stellar chromospheric activity.

### **Consequences of a chromospheric temperature gradient on the width of H [alpha] in late**

Model atmospheres Accurate absolute abundances in LTE and NLTE require realistic model atmospheres, hence spectrum synthesis based on 3D hydrodynamical models is expected to give superior results. The surface gravity of the Sun is 27.

### **The FIP and Inverse FIP Effects in Solar and Stellar Coronae**

Unfortunately, LASCO data were not available to confirm the existence of a halo CME. Vertical profile of the ionisation rate in the terrestrial atmosphere, as A originally obtained by Hess 7th August 1912 , with ionisation at each height shown relative to the measured surface ionisation and, B from a series of balloon flights colours used to identify individual flights made from Reading, UK, during 2013.

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