Goup Project Descriptions

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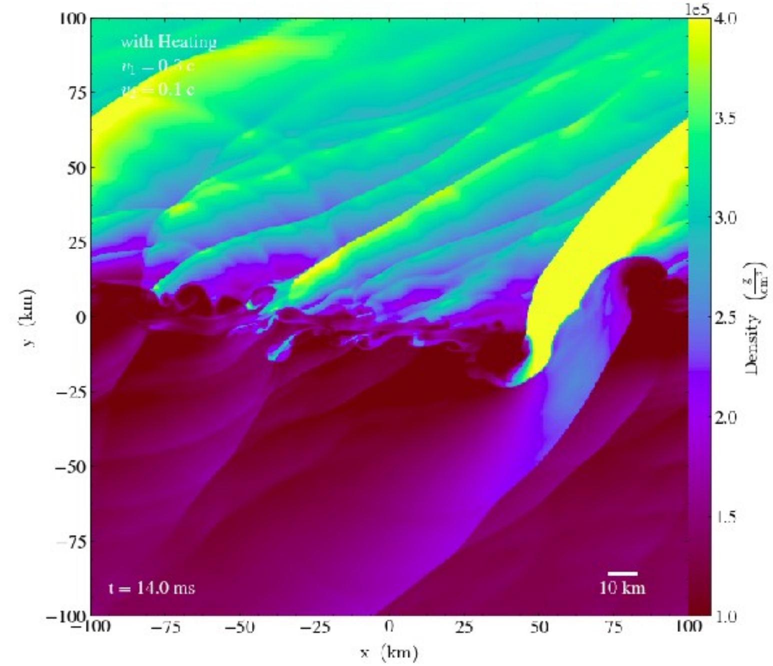
NTHU, Hsinchu, Taiwan, Sept 04–06, 2019

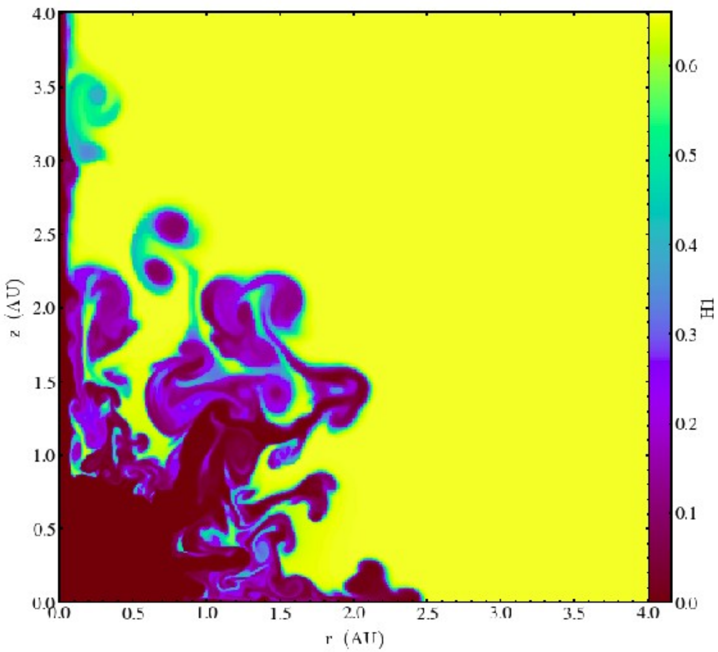
Three main projects

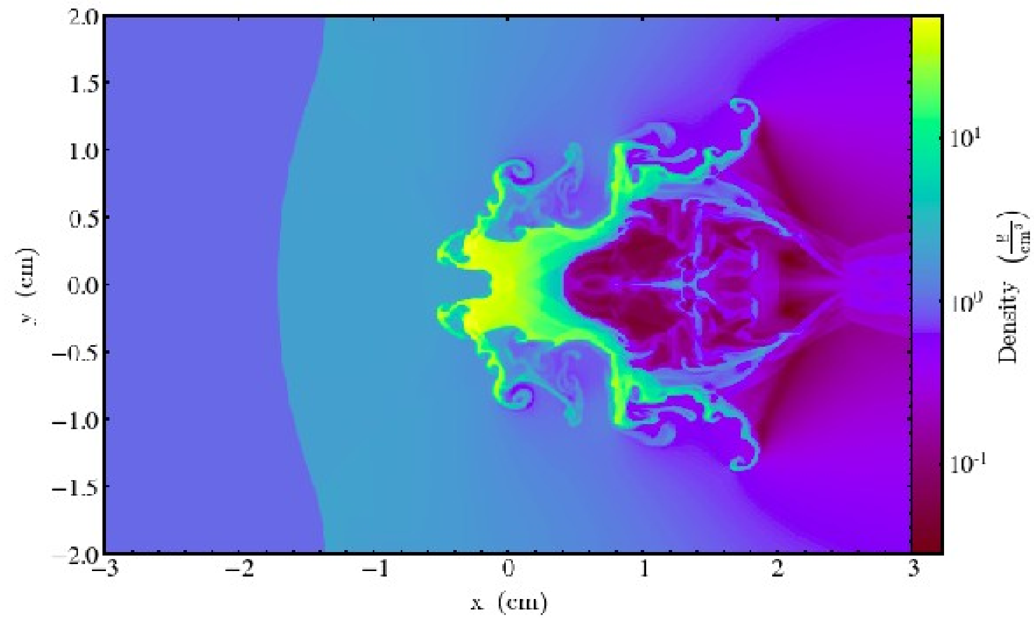
A: Mixing in neutron star merger ejecta

B: Shock Cloud Interaction

C: Exploding a star with a thermal bomb







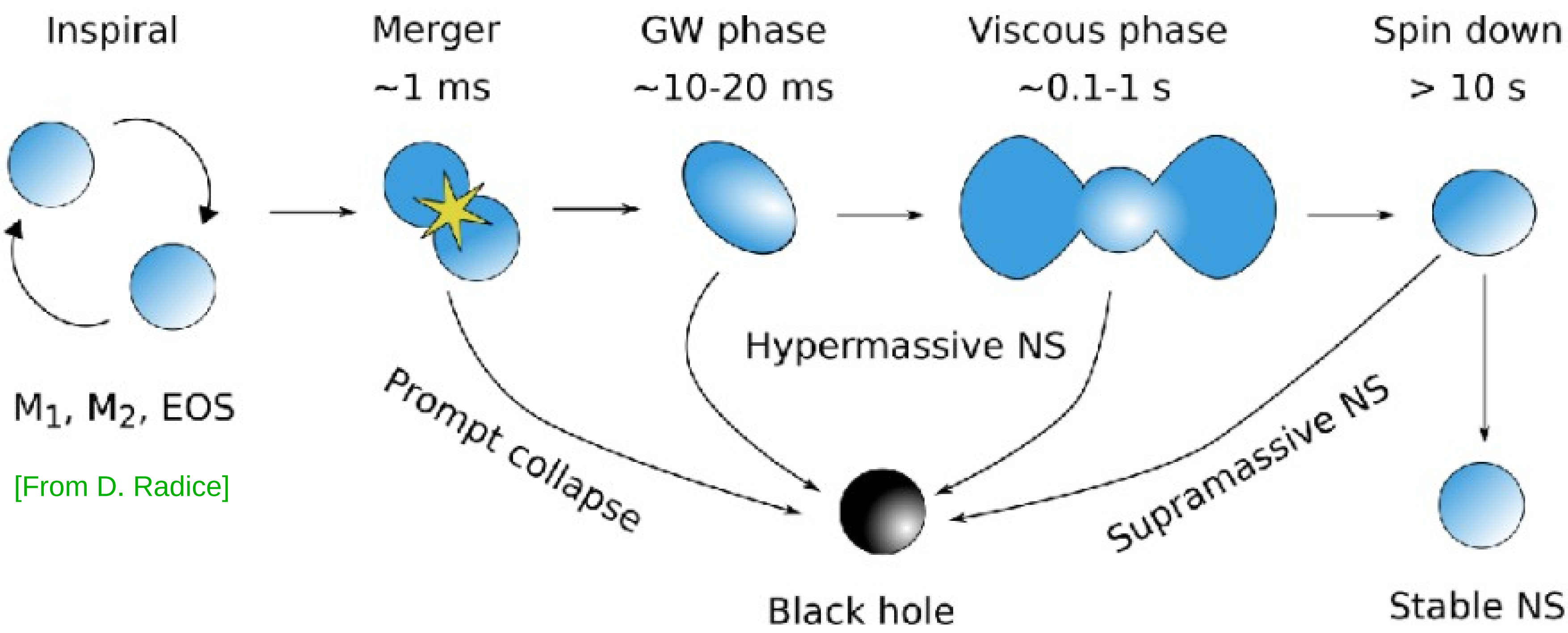
Project A: Mixing in neutron star merger ejecta

Project A: Mixing in neutron star merger ejecta

Mergers of two neutron stars can result in a variety of outcomes, depending on the initial state of the system and the yet-unknown nuclear EoS

$\rightarrow$ different post-merger GW signals & EM emissions

both were detected in GW170817!



[From D. Radice]

Project A: Mixing in neutron star merger ejecta

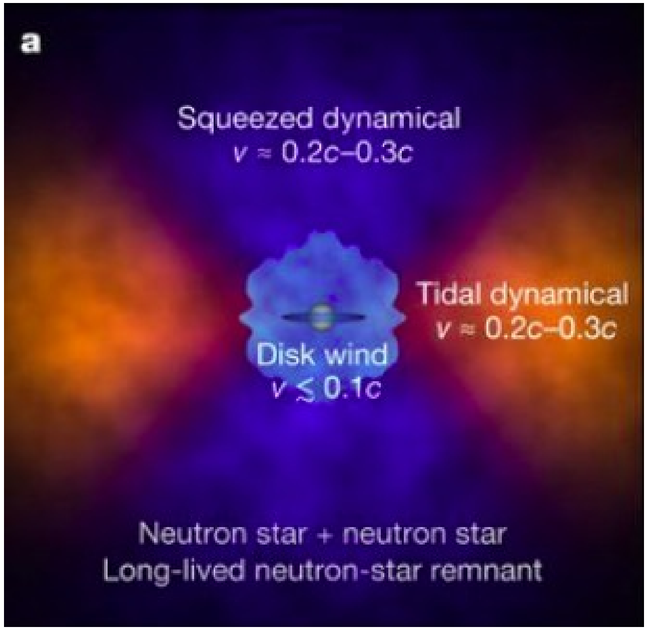
Different outflow components:

[Modified from Fujibayashi+ 2017]

Type of ejecta	Mass ($M_\odot$)	V_{ej}/c
Dynamical ejecta	$O(10^{-3})$	~ 0.2
Early viscosity-driven ejecta	$\sim 10^{-2}(\alpha_{vis}/0.02)$	$\sim 0.15 - 0.2$
Late-time viscosity-driven ejecta (polar)	$\sim 10^{-3}(t_\nu/\mathrm{s})$	~ 0.15
Late-time viscosity-driven ejecta (equatorial)	$\gtrsim 10^{-2}$	~ 0.05

Y_e	Direction	Duration
0.05–0.5	$\theta \gtrsim 20°$	$t - t_{\mathrm{merge}} \lesssim 10\,\mathrm{ms}$
0.2–0.5	$\theta \gtrsim 30°$	$t - t_{\mathrm{merge}} \lesssim 0.1\,\mathrm{s}$
0.4–0.5[a]	$\theta \lesssim 30°$	$t - t_{\mathrm{merge}} \sim t_\nu \sim 10\,\mathrm{s}$
0.2–0.4[a]	$\theta \gtrsim 30°$	$t - t_{\mathrm{merge}} \sim 1–10\,\mathrm{s}$

[Kasen+ 2017]

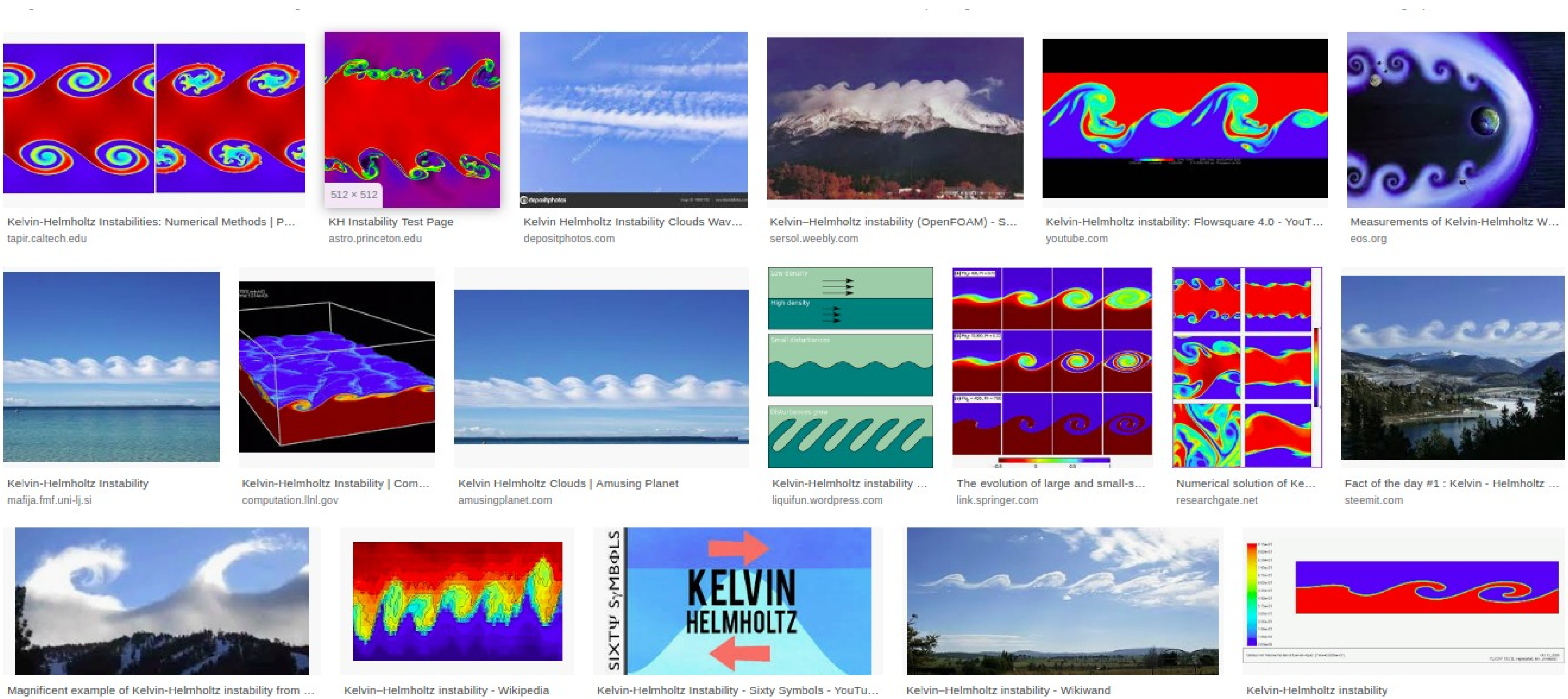


Y_e and v_{ej} can depend on the outflow drection

$(Y_e = n_e/n_b)$

Project A: Mixing in neutron star merger ejecta

Velocity shear drives instability at the interface of two fluids, which then generates fluid mixing $\rightarrow$ Kevin-Helmholtz instability



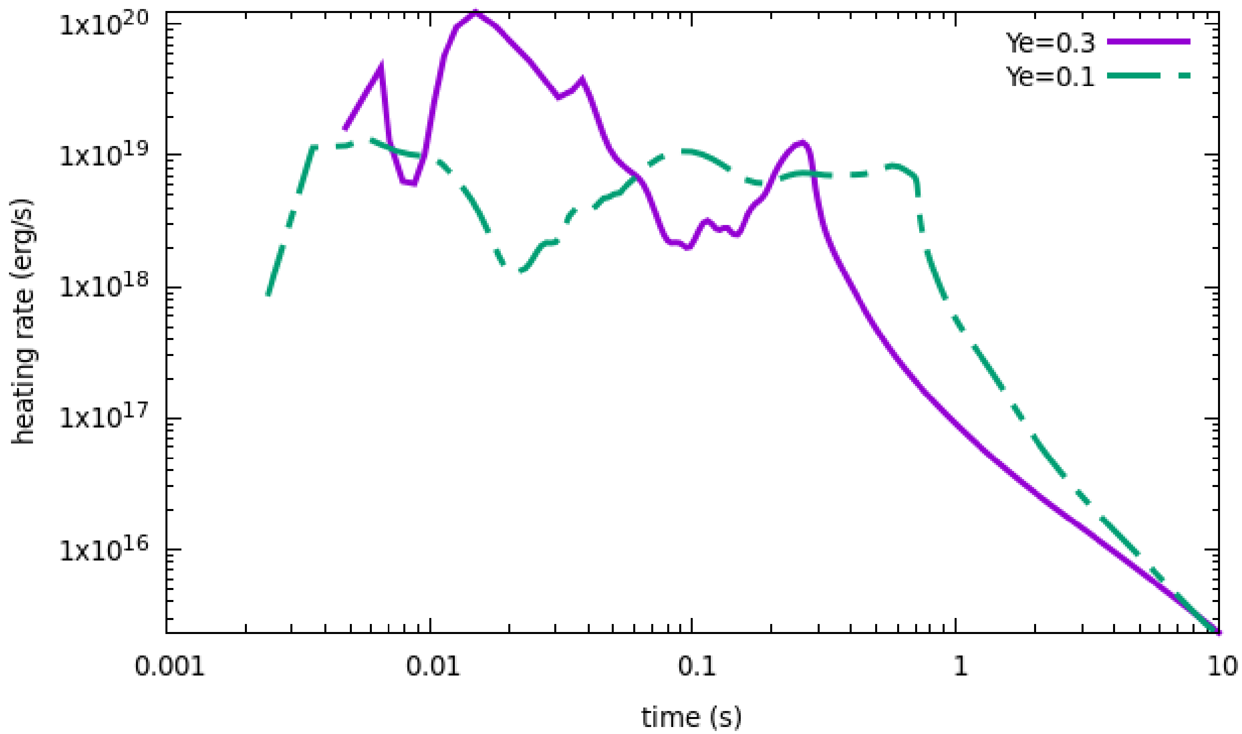
Project A: Mixing in neutron star merger ejecta

Merger ejecta not only simply expand and cool, but also receive energy inputs due to the decay of unstable heavy r-process nuclei (how the gold was made)

$\rightarrow$ composition-dependent nuclear heating

high Y_e: stronger earlier heating but last for a shorter time

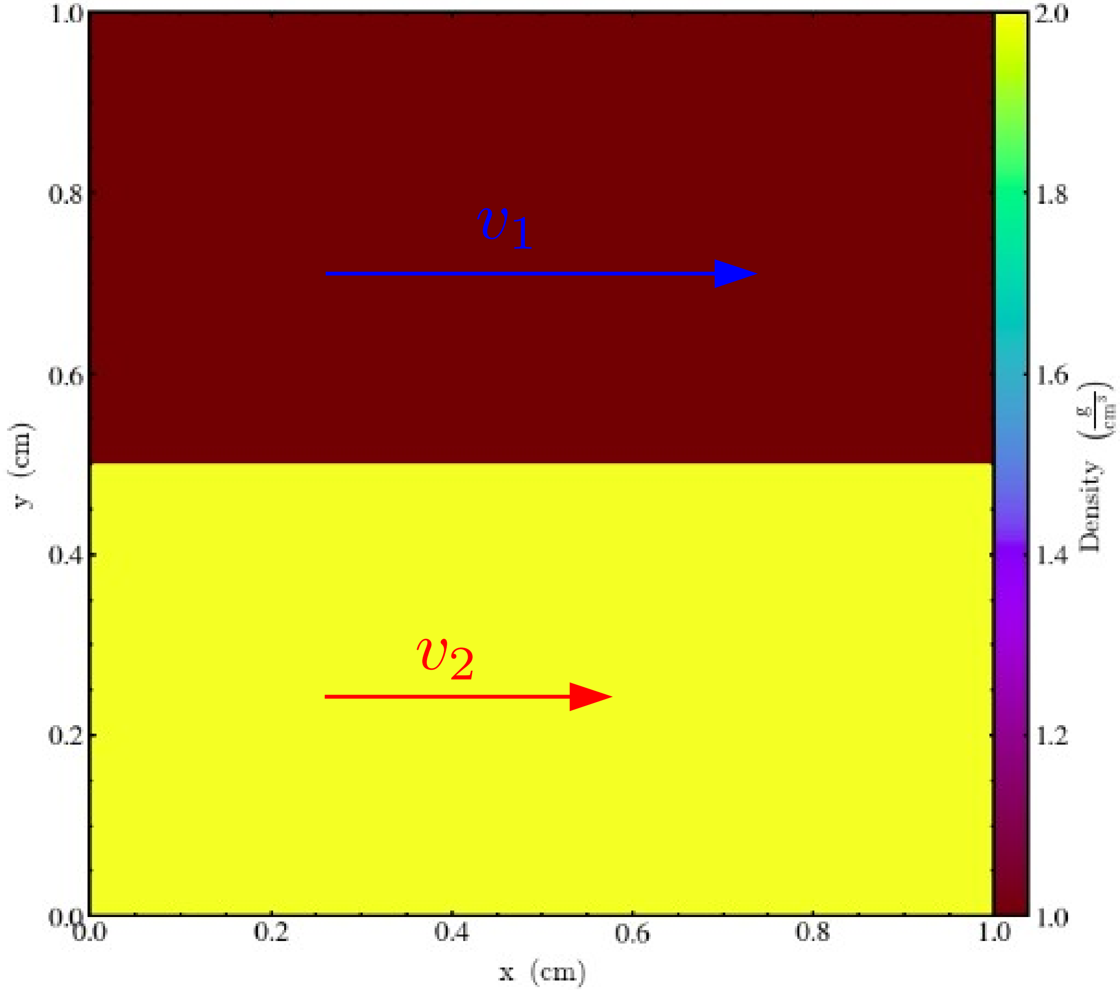
low Y_e: weaker earlier heating but last for a longer time



Does this affect the mixing?

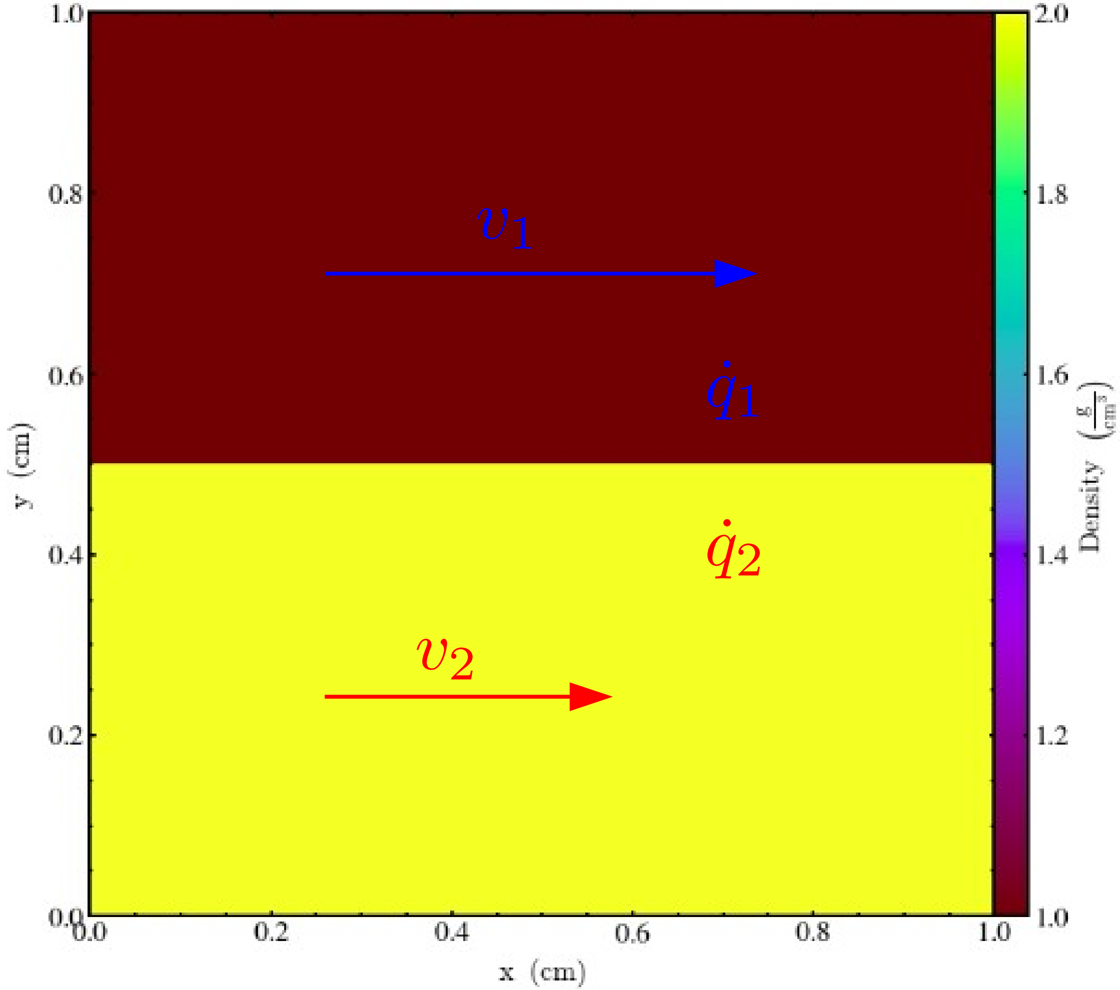
Project A: Mixing in neutron star merger ejecta

Toy model (I): 2D simulation of two fluids moving in positive x-direction with different velocities



Project A: Mixing in neutron star merger ejecta

Toy model (II): 2D simulation of two fluids moving in positive x-direction with different velocities, plus different heating rates



Project A: Mixing in neutron star merger ejecta

Things that you may explore:

- different velocities v_1 and v_2
- different heating rates $\dot{q}_1$ and $\dot{q}_2$
- different initial densities or temperatures
- inhomogeneous or time-dependent heating rate
- any set-up beyond two fluids!

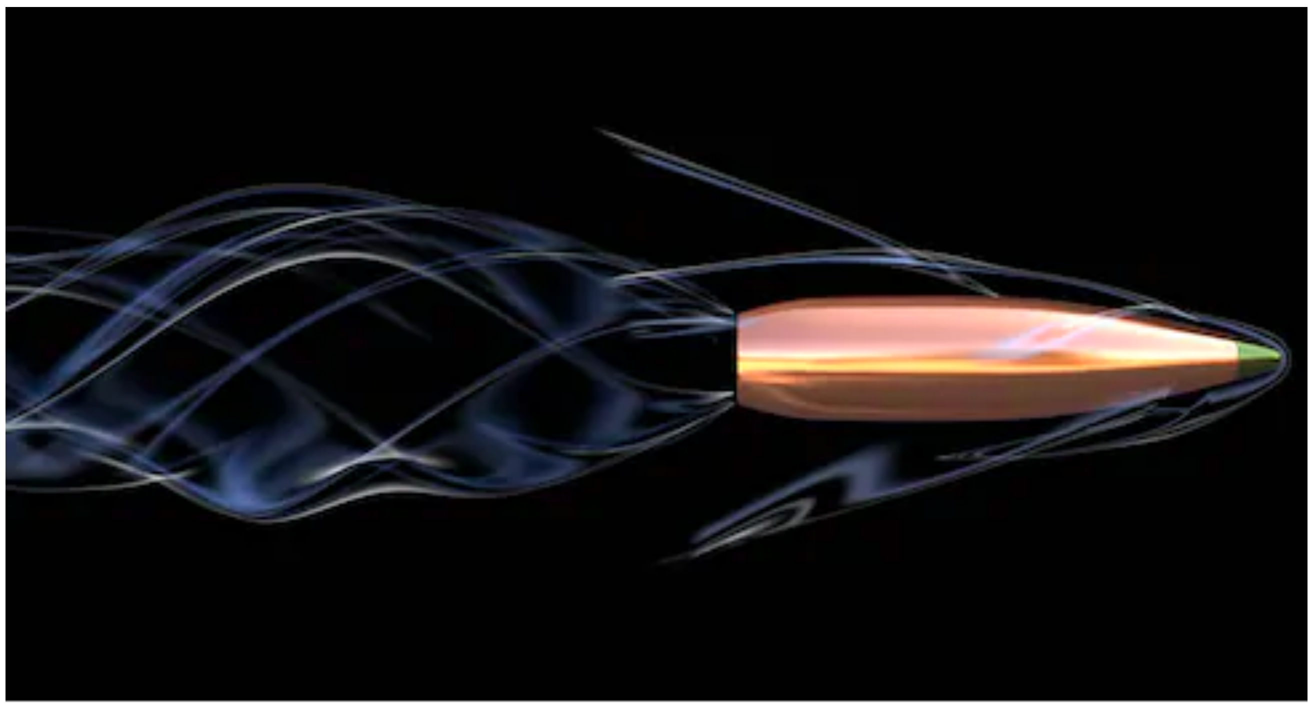
Project B

Project B: Shock Cloud Interaction

Interaction between shock and material is one of the most common astrophysical phenomena.

Shock: SN shock, outflow, jet, winds, …etc.
Cloud: Star, ISM, IGM

Project B: Shock Cloud Interaction



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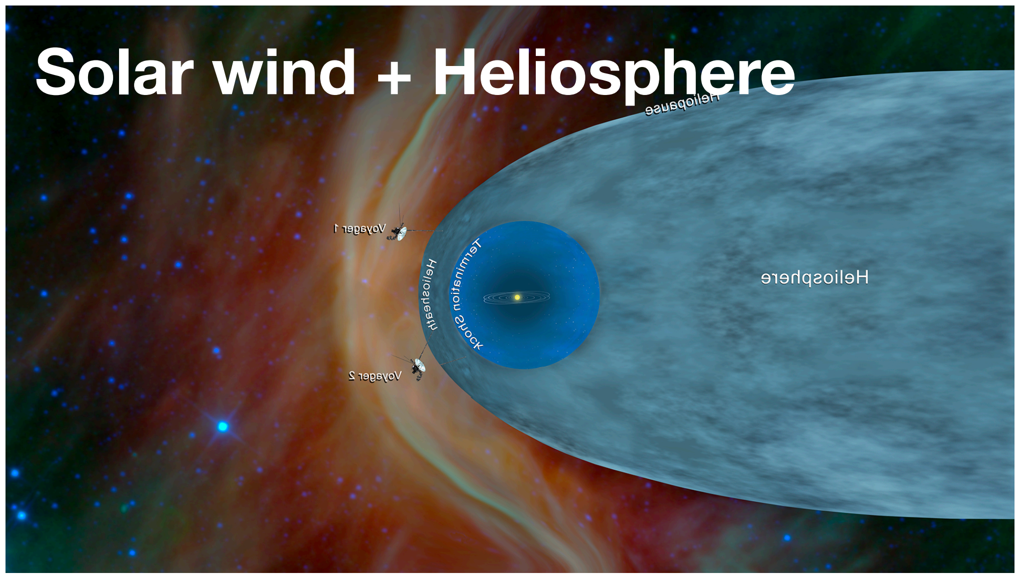
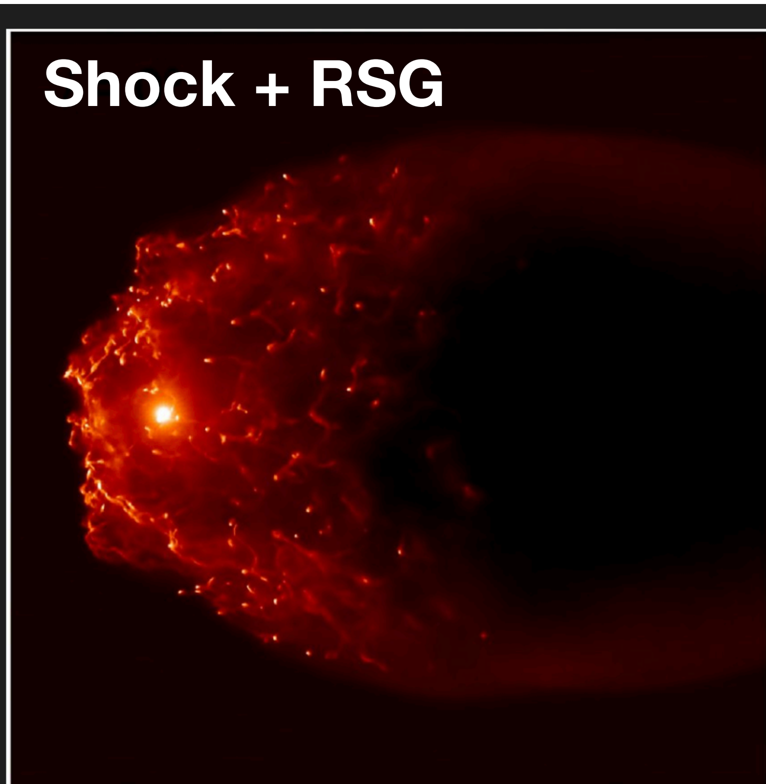
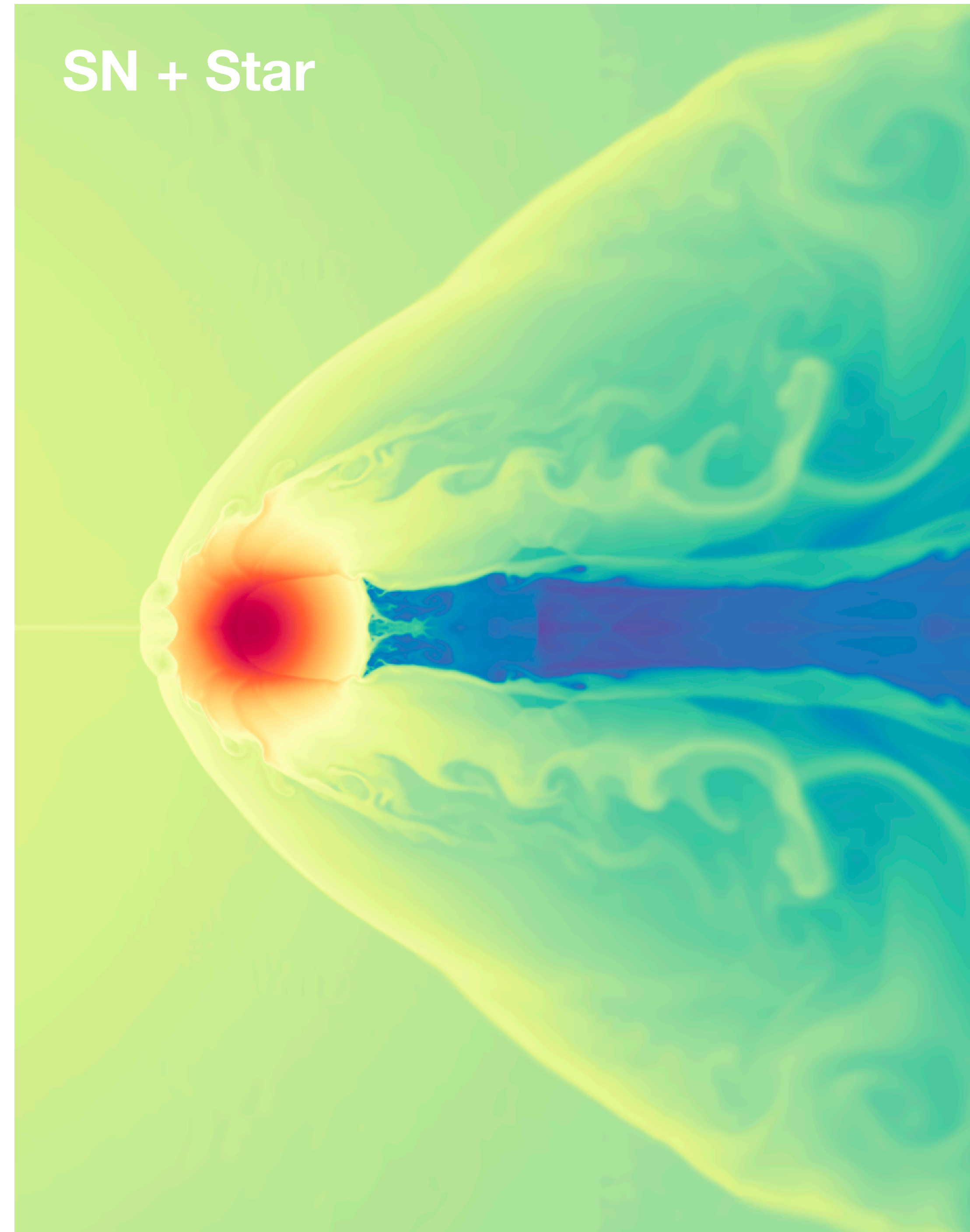
Project B: Shock Cloud Interaction



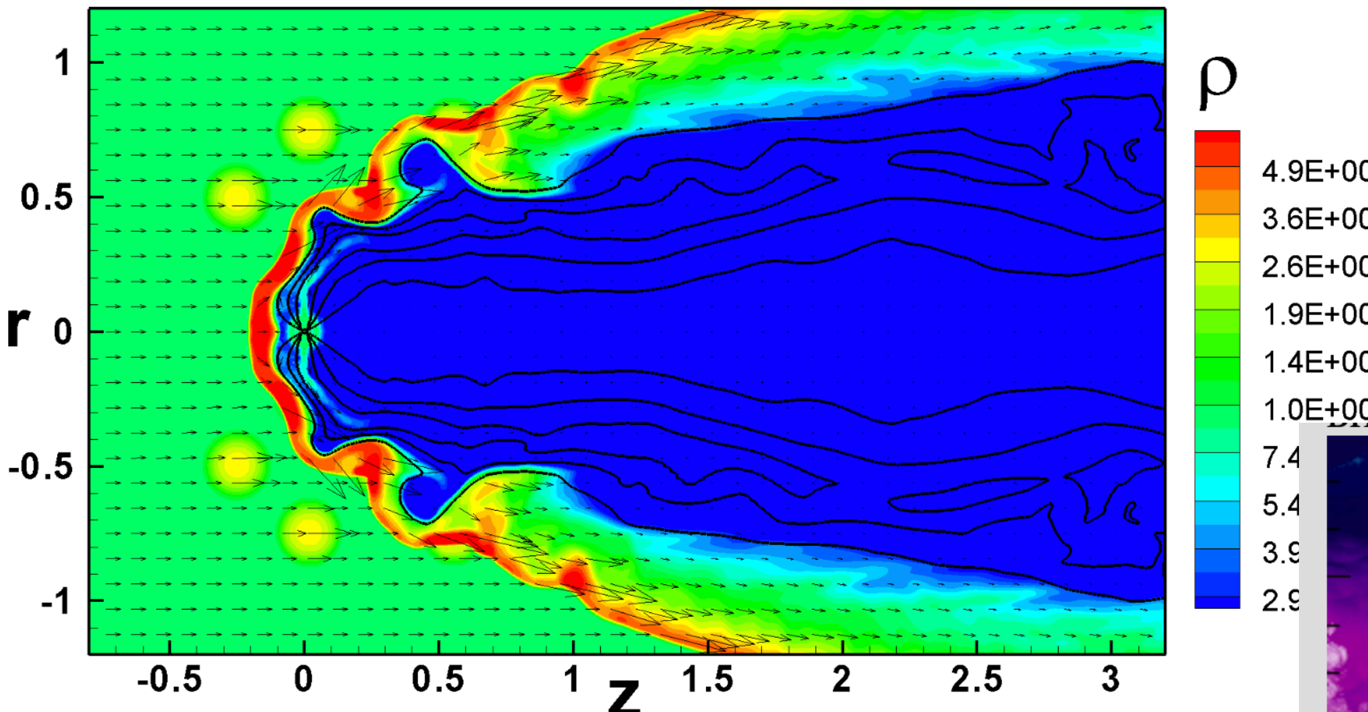


Simulations of the bow shock around the red supergiant star Betelgeuse (Mohamed et al. 2012), used as a cover image by Astronomy & Astrophysics.



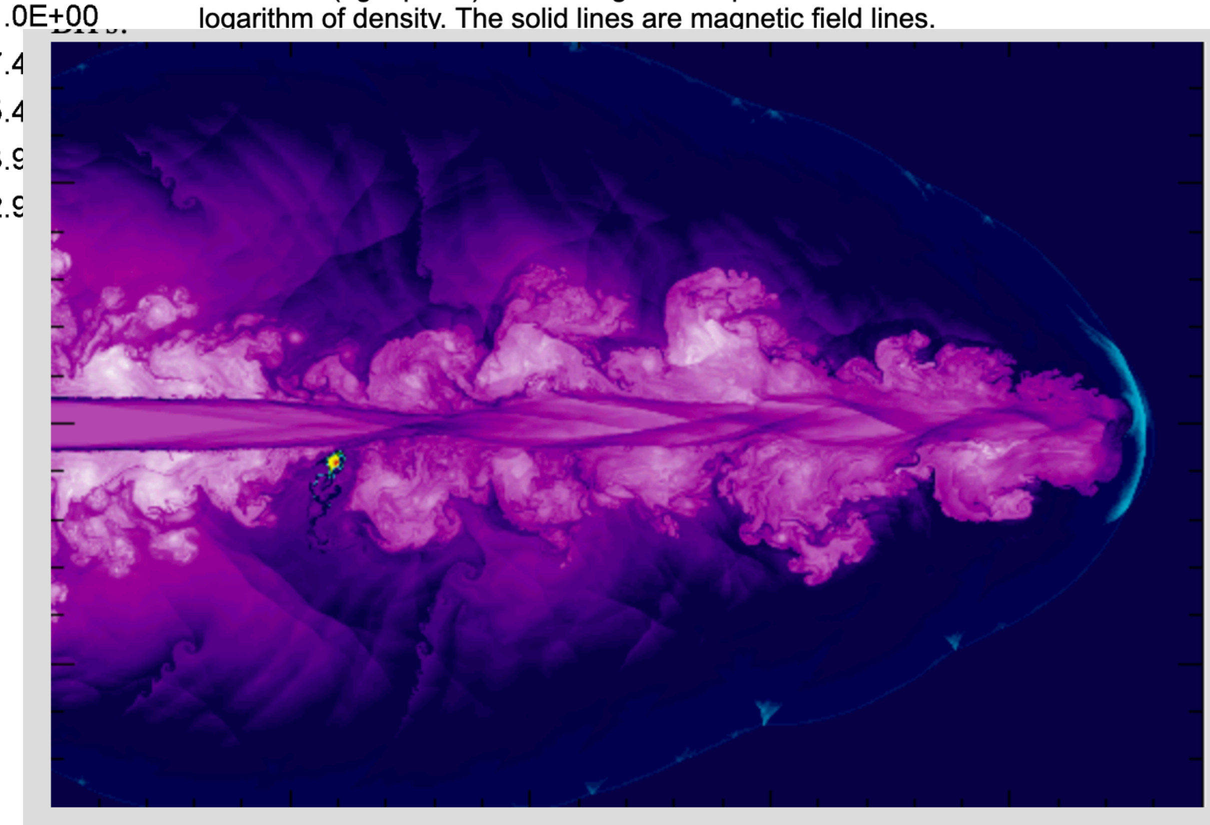
Project B: Shock Cloud Interaction

[Modelling the bow shock Pulsar Wind Nebulae propagating through a non-uniform ISM](#) - Toropina, O.D. *et al.* Mon.Not.Roy.Astron.Soc. 484 (2019) no.2, 1475-1486 arXiv:1803.06240 [astro-ph.HE]



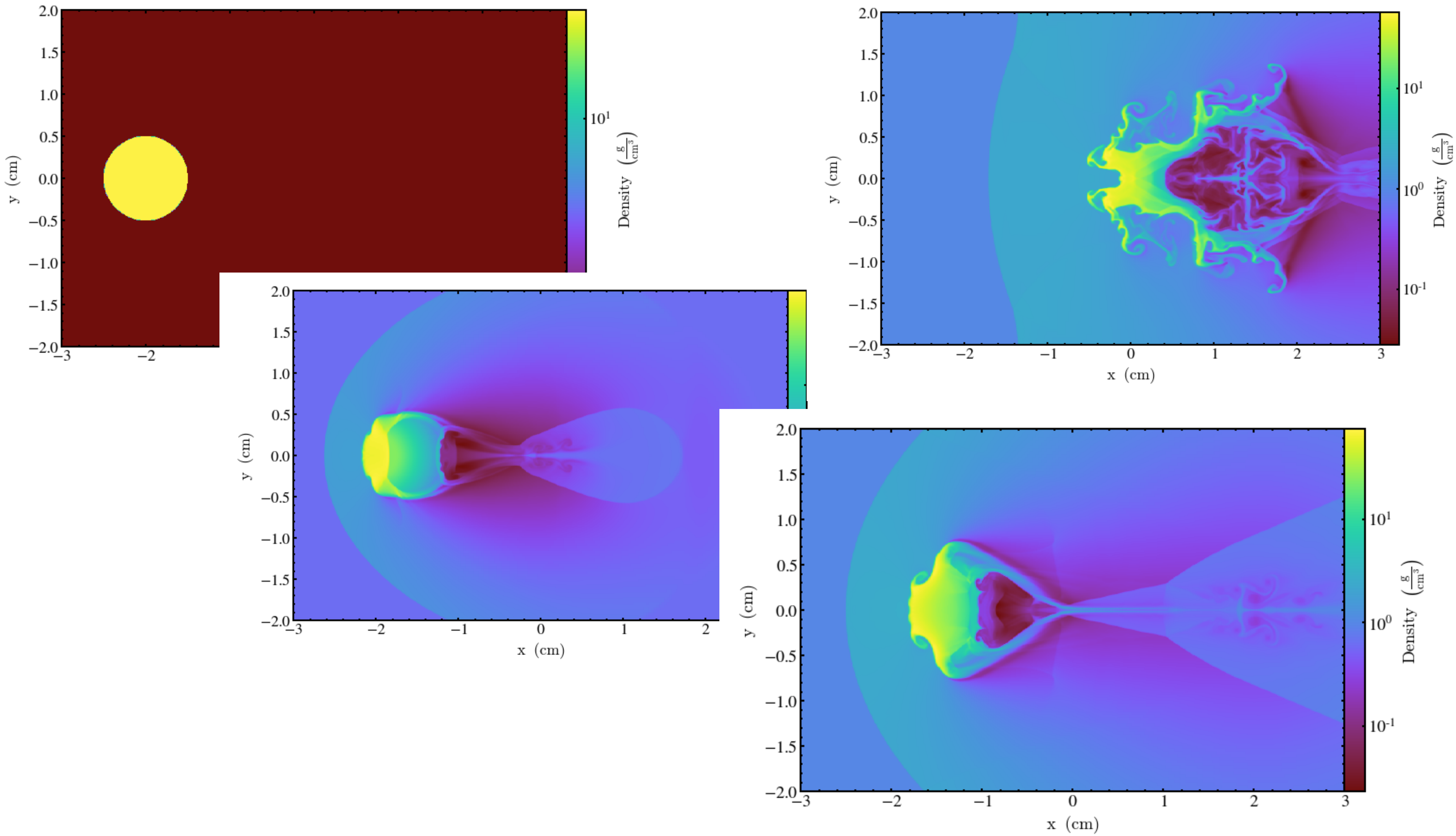
Pular wind + ISM

Interaction of the bow shock with small-scale clouds of maximum density $\rho_{\rm cloud} = 3\rho_0$ in the model $B1M20w50$ with low σ (left panel) and in the model $B5M20w50$ with medium σ (right panel). The background represents the logarithm of density. The solid lines are magnetic field lines.



Jet + IGM

Project B: Shock Cloud Interaction

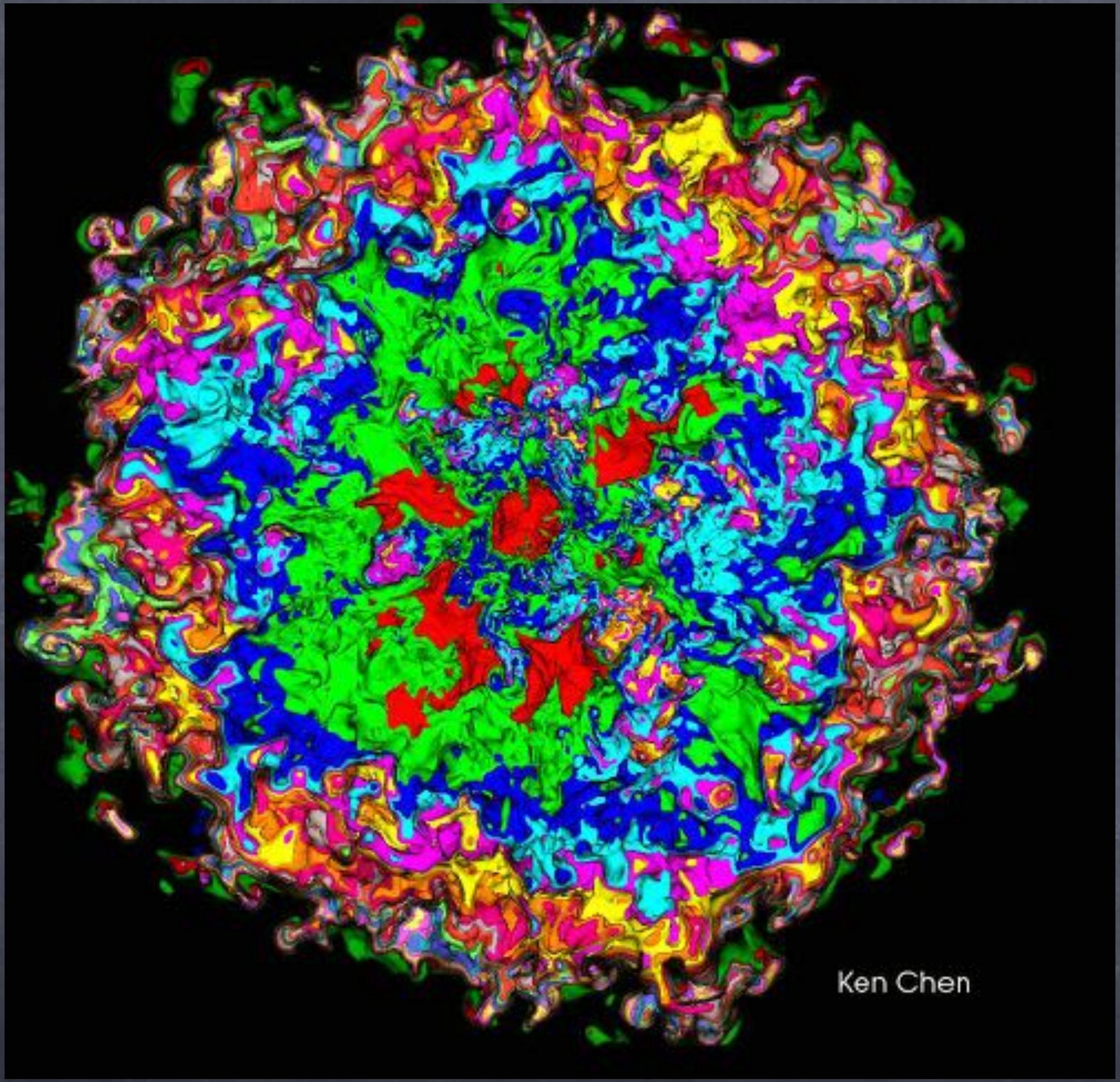


Project B: Shock Cloud Interaction

Things that you may explore:

- Different initial density distribution
- Explore different inflow speed
- Explore different geometry
- Time/Location dependent inflow
- Radiative cooling

Explode a star in computer



Ke-Jung (Ken) Chen
陳科榮
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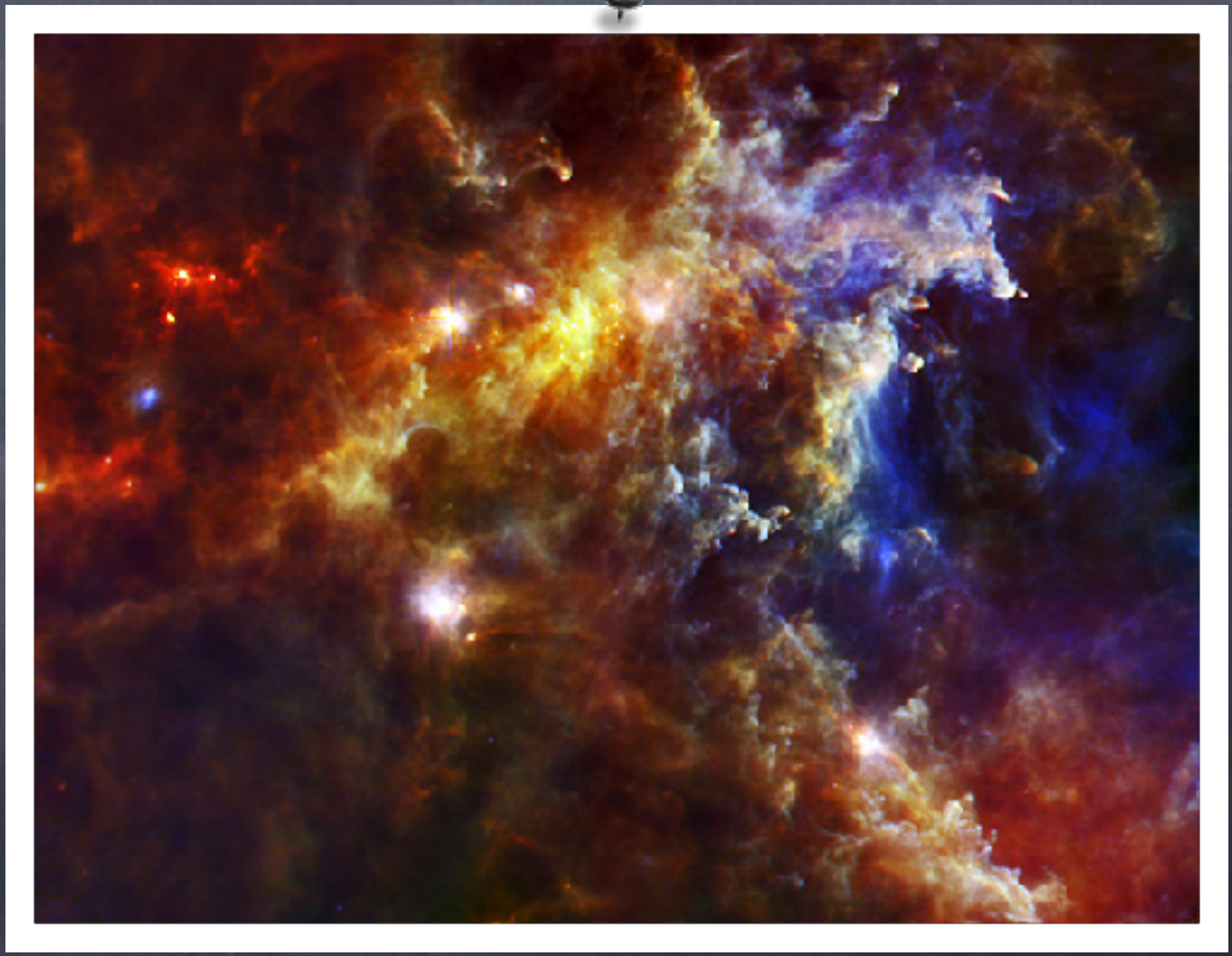
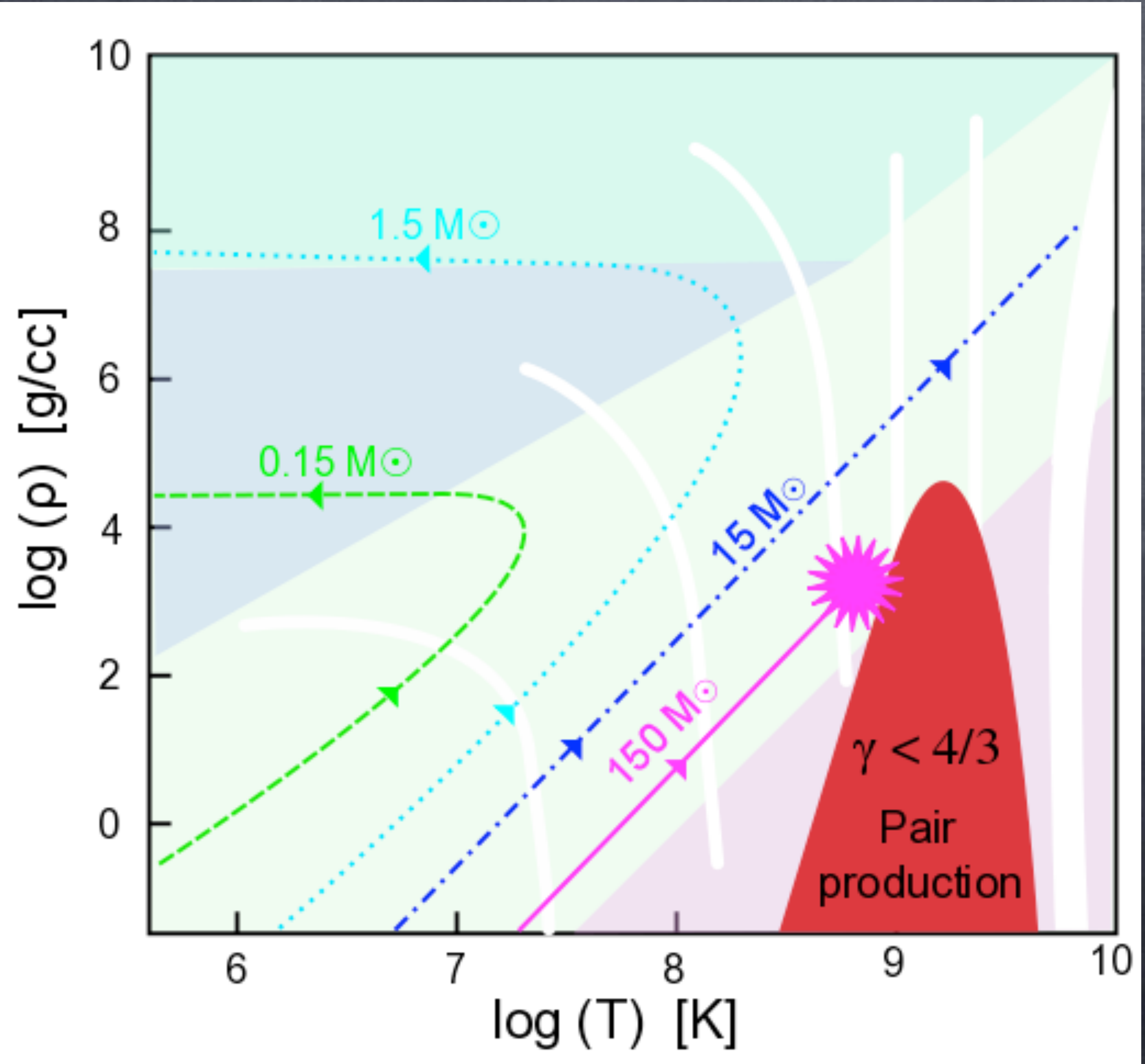
AFD School, NTHU, 09/05/2019

Stars are atoms of Cosmos



Temperature-Density Diagram of Stellar Evolution



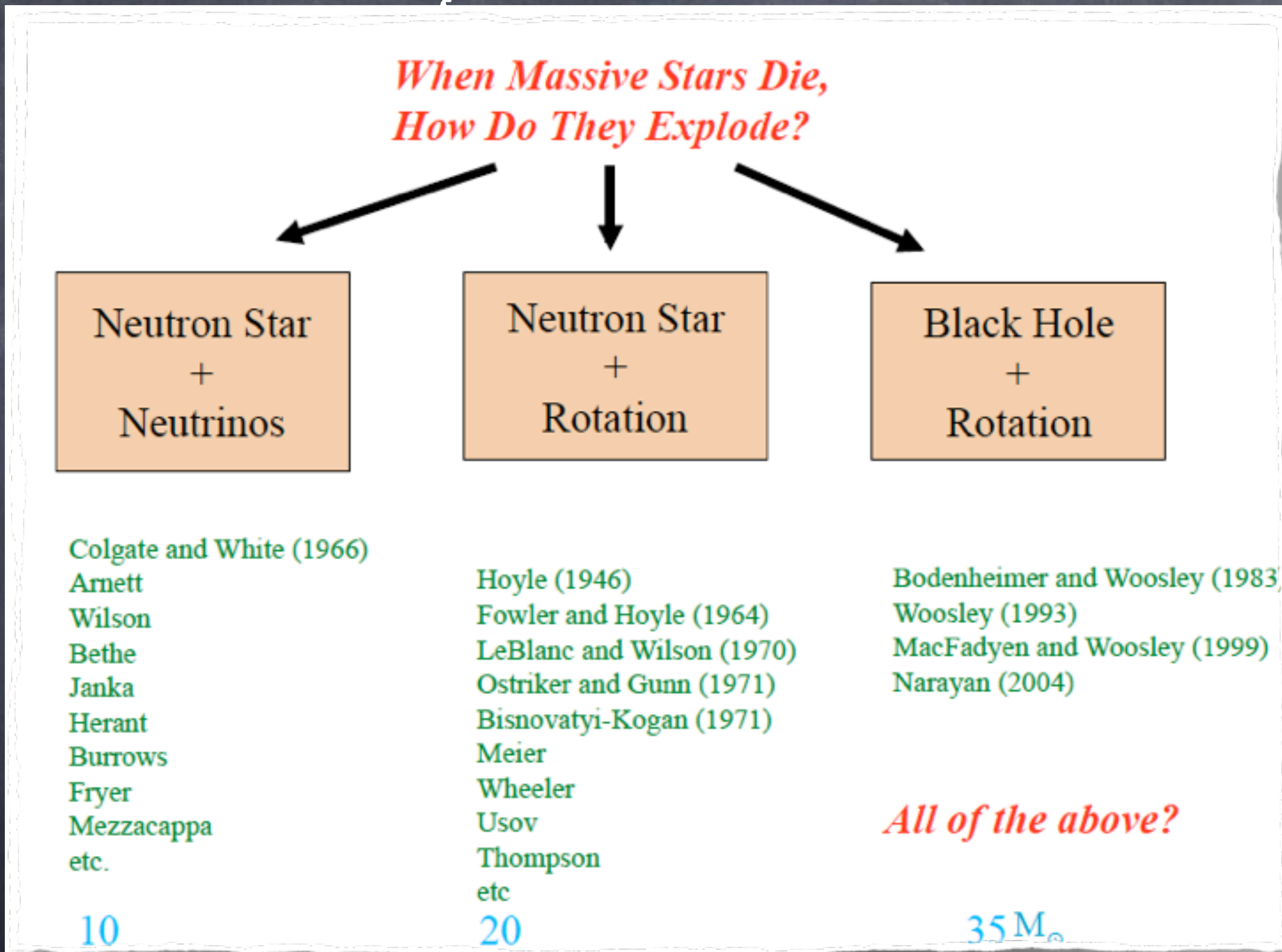
Life of Massive Stars

Advanced Nuclear Burning Stages
(e.g., 20 solar masses)

Fuel	Main Product	Secondary Products	Temp (10^9 K)	Time (yr)
H	He	^{14}N	0.02	10^7
He	C,O	^{18}O, ^{22}Ne s- process	0.2	10^6
C	Ne, Mg	Na	0.8	10^3
Ne	O, Mg	Al, P	1.5	3
O	Si, S	Cl, Ar K, Ca	2.0	0.8
Si	Fe	Ti, V, Cr Mn, Co, Ni	3.5	1 week

Woosley & Heger

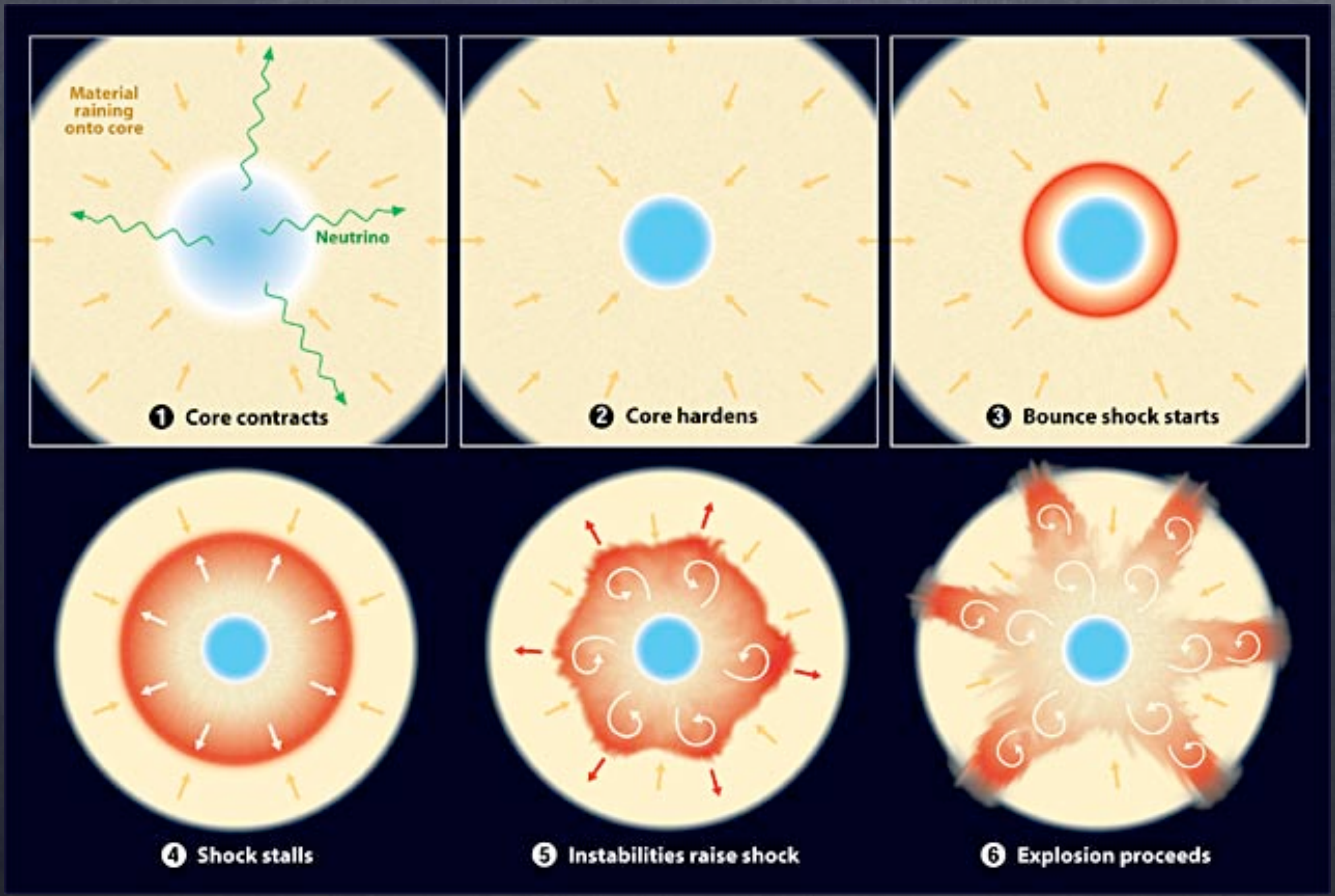
Death of Massive Stars



Woosley

Core-Collapse Supernova

(20 M⊙ > M* > 10 M⊙, E~1E51 erg)



Energy Source

<u>Neutrino Burst Properties:</u>

$$E_{tot} \sim \frac{3}{5}\frac{GM^2}{R} \qquad M = 1.5\, M_\odot$$

$$\boxed{\sim 3\times10^{53} \text{ erg}} \qquad R = 10 \text{ km}$$

emitted roughly equally in $\nu_e, \bar{\nu}_e, \nu_\mu, \bar{\nu}_\mu, \nu_\tau$, and $\bar{\nu}_\tau$

<u>Time scale</u>

$$\tau_{Diff} \sim \left(\frac{R^2}{l\ c}\right) \qquad l = \frac{1}{\kappa_\nu \rho}$$

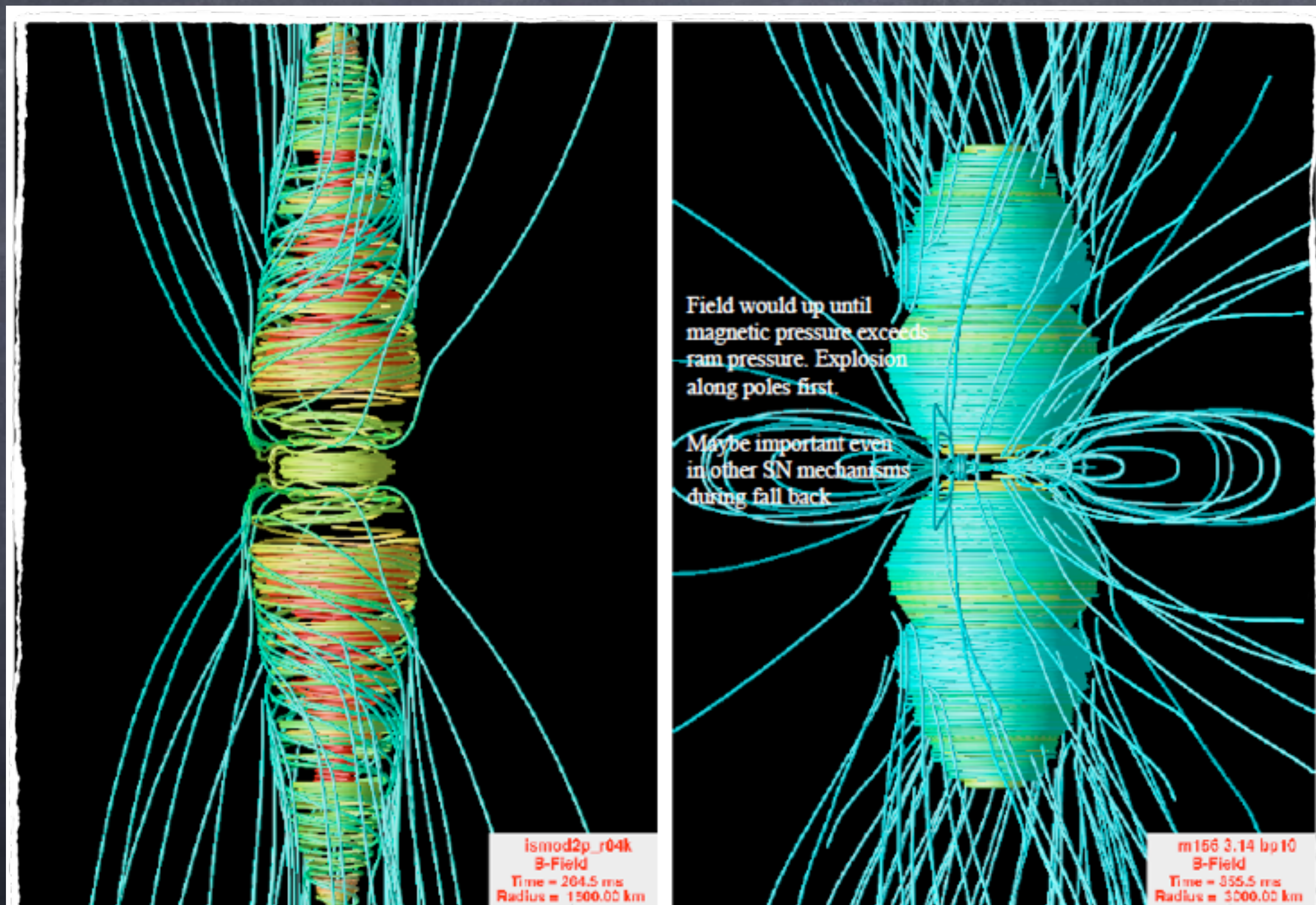
$$\kappa_\nu \sim 10^{-16} \text{ cm}^2 \text{ gm}^{-1} \text{ for } \varepsilon_\nu = 50 \text{ MeV (next page)}$$

$$\rho \sim 3\times10^{14} \text{ gm cm}^{-3} \Rightarrow l \sim 30 \text{ cm} \qquad R \sim 20 \text{ km}$$

$$\tau_{Diff} \sim \left(\frac{(2\times10^6)^2}{30 \cdot 3\times10^{10}}\right) \sim \boxed{5 \text{ sec}}$$

Very approximate

Hypernovae
(E > 1E52 erg)



Burrows et al 2007, *ApJ*, **664**, 416

Energy Sources

Assuming the emission of high amplitude ultra-relativistic MHD waves, one has a radiated power

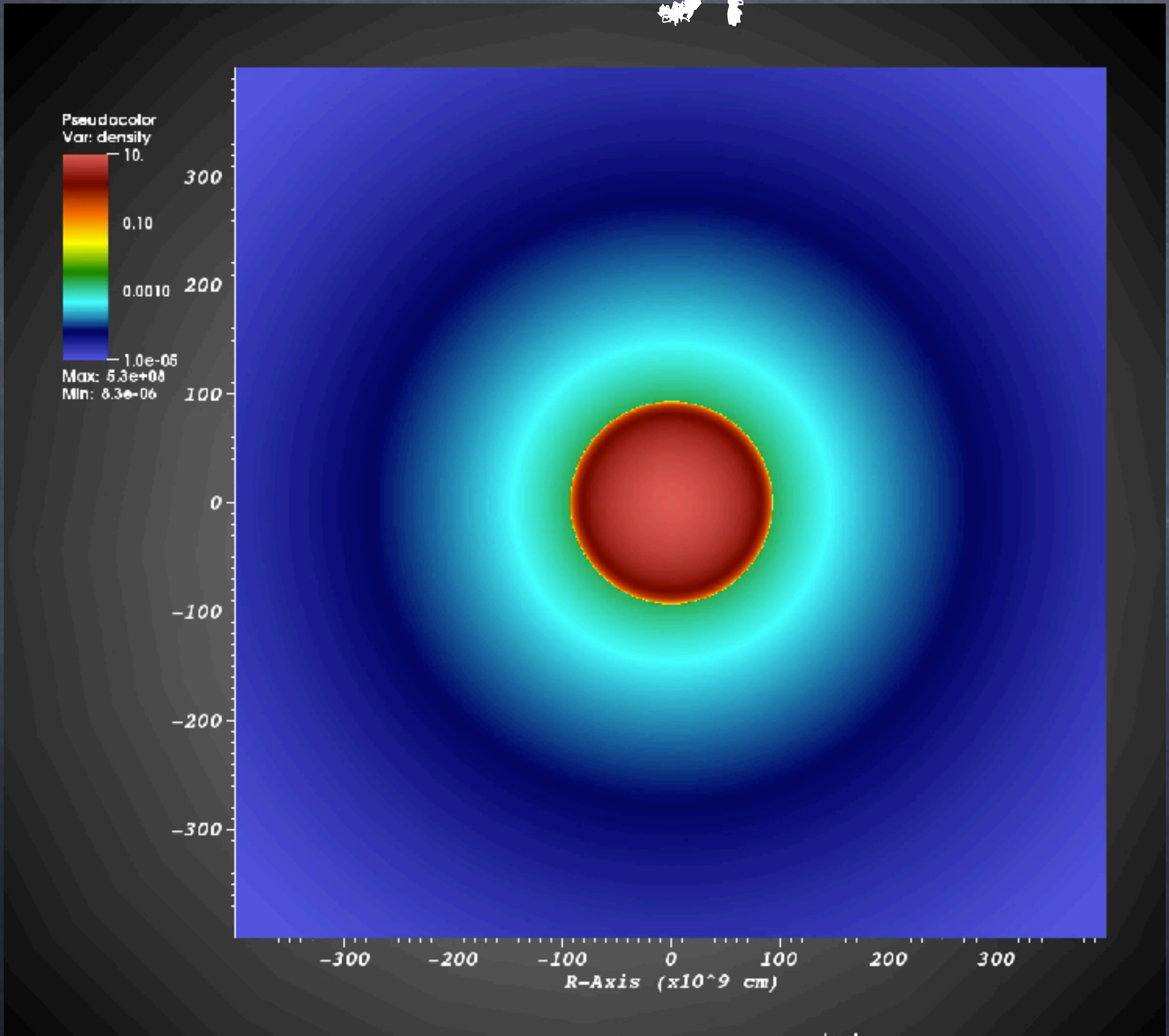
$$P \sim 6 \times 10^{49}\ (1\ \mathrm{ms/P})^4\ (B/10^{15}\ \mathrm{gauss})^2\ \mathrm{erg\ s^{-1}}$$

and a total rotational kinetic energy

$$E_{rot} \sim 4\ x\ 10^{52}\ (1\ \mathrm{ms/P})^2\ (10\ \mathrm{km/R})^2\ \mathrm{erg}$$

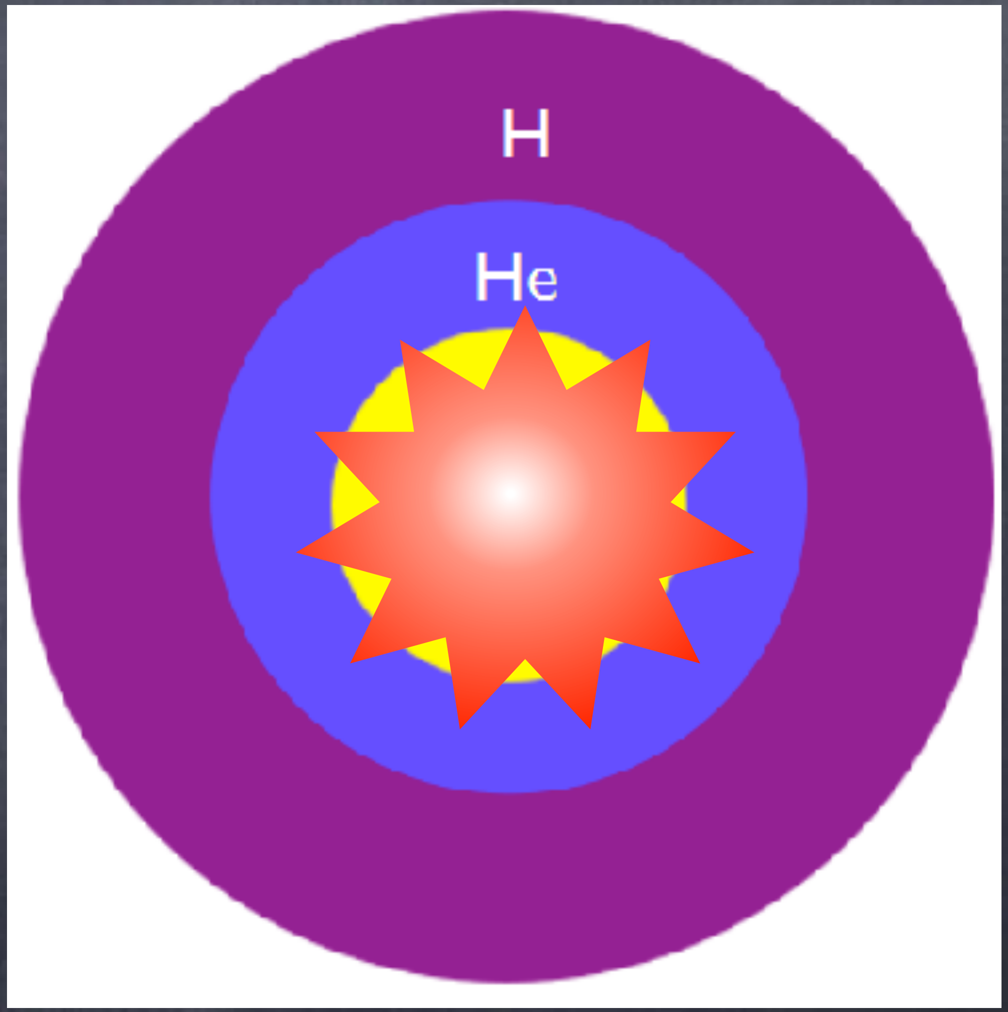
For magnetic fields to matter one thus needs magnetar-like magnetic fields and rotation periods (for the cold neutron star) of < 5 ms. This is inconsistent with what is seen in common pulsars. Where did the energy go?

GRB and Hypernova



Project C
Blowing up a star with Flash

- Explosion Energy E~1e50- 1e52 erg?
- Size and shape of the energy deposition area?
- Properties of surrounding ISM and its density/temperature/... profile?
- Striped Envelope SNe?



Have fun in blowing up Stars