

Cosmological Hydrodynamic Simulations of Preferential Accretion in the SMBH of Milky Way Size Galaxies

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ABSTRACT

Using the cosmological hydrodynamic simulation of a Milky Way-type (MW-type) Galaxy, we explore how a merger-rich assembly history affects the mass budget of the central SMBH. We examine a MW-mass halo characterized by several major mergers. We attempt to identify the importance of merger history on black hole accretion using a new high resolution simulation and advanced SPH code. This study is an extension of Bellovary et. al. 2013, which analyzed the accretion of high mass, high redshift galaxies and their central black holes, and found that the gas content of the central black hole is proportional to that accreted by the host galaxy halo. Contrarily, in this study, we find that a merger-rich galaxy will have a central SMBH preferentially fed by merger gas. Moreover, we find that nearly 20% of total mass budget of the SMBH enters the galaxy through the main merger. Through an investigation of the angular momentum of the gas entering the host and its SMBH, we determine that merger gas enters the galaxy with lower angular momentum compared to smooth accretion, partially accounting for the preferential fueling witnessed in the SMBH. In addition, the presence of mergers, particularly major mergers ($q > 0.3$), also helps funnel the low angular momentum gas more readily to the center of the galaxy. Our results imply that galaxy mergers play an important role in the feeding of the SMBH in merger-rich Milky Way-type galaxies.

Subject headings: Black hole physics – Galaxies: spiral – Galaxies: kinematics and dynamics – Methods: Numerical

1. Introduction

Supermassive black holes (SMBHs) are thought to exist in almost all massive galaxies (Kormendy & Ho 2013). In the canonical picture of BH growth, these black holes may become active galactic nuclei (AGN) during periods of high accretion and wane in periods of quiescence (Alexander et al. 2005; Papovich et al. 2006; Volonteri 2012). The host galaxy’s size, star formation rate, and other environmental effects may help to influence the growth of the black hole residing at its center; however, there are still uncertainties concerning the relationship between these SMBHs and their much larger host galaxies, as well as how they grow and evolve together (Haehnelt & Kauffmann 2000; Di Matteo et al. 2005; Hopkins et al. 2006; Fu & Stockton 2008; Sijacki et al. 2009; Silverman et al. 2009; Mullaney et al. 2012).

The $M-\sigma$ relation, which relates the SMBH’s mass and the velocity dispersion of the host galaxy’s central stellar population, gives some insight into the complex interplay between these objects (Ferrarese & Merritt 2000). A prominent trend appears, as SMBHs tend to scale with the velocity dispersion of the host galaxy bulge. The tightness of the relation is significant and can be seen over several orders of magnitudes in velocity dispersion and black hole mass (Merritt & Ferrarese 2001b; Graham et al. 2011; McConnell & Ma 2013; Kormendy & Ho 2013). Scatter exists among the low mass galaxies and a deviation may appear at the high mass end, where overmassive BHs may reside (Van Den Bosch et al. 2007; Moster et al. 2010; Natarajan 2011). However, scatter in less massive galaxies may imply that there are several channels of black hole growth at play in the low mass end of the relation (Micic et al. 2007; Volonteri & Natarajan 2009; Reines et al. 2013; Graham & Scott 2014). One standard explanation for the $M-\sigma$ relation lies in galaxy mergers, which build up galaxies, feed SMBHs, and assemble bulges (Di Matteo et al. 2005; Shen et al. 2008). Major mergers are thought to supply gas to the central SMBH resulting

in feedback which quenches star formation and affects the structure of the galaxy.

The black hole mass-bulge luminosity relation was implied by the work of [Dressler & Richstone \(1988\)](#) and first illustrated by [Kormendy \(1993\)](#). A wealth of evidence continues to relate these two characteristics, and follow up examinations using HST observations have determined the current, best estimates yield a $M_{\text{BH}}/M_{\text{bulge}}$ fraction between about 0.0013 and 0.0023 ([Merritt & Ferrarese 2001a](#); [McLure & Dunlop 2001](#); [Marconi & Hunt 2003](#)).

Major mergers between massive galaxies are thought to be efficient fueling mechanisms for bright AGN. The large influx of material due to tidal torques from the merger causes bursts of star formation and helps funnel gas directly into the center where the SMBH resides (e.g. [Richards et al. 2006](#); [Reddy et al. 2008](#); [Hopkins & Quataert 2010](#)). Additionally, the most massive, highest-luminosity AGN (i.e. quasars) reside in incredibly luminous infrared galaxies where star formation is abundant, signifying that major mergers may have recently occurred ([Treister et al. 2012](#)). Distorted morphologies are often characteristics of quasar hosts, and companions can also be present around quasars, both of which are evidence that strengthen the possibility of a recent merger having affected their lifetimes.

In many less massive and less luminous AGN, however, there is a clear lack of distorted morphologies, close neighbors, and/or other obvious merger evidence ([Ryan et al. 2007](#); [Schawinski et al. 2011](#); [Ellison et al. 2013](#); [Hicks et al. 2013](#)). It is also important to note that many of these AGN exist in spiral galaxies, which are unlikely to have been recently disturbed by major mergers ([Springel & Hernquist 2005](#); [Kocevski et al. 2011](#)). Nevertheless, some evidence suggests ([van Gorkom & Schiminovich 1997](#); [Governato et al. 2009b](#)) that disturbed galaxies may reform a disk quickly, even after a major merger as long as it is gas-rich. More recently, [Treister et al. \(2012\)](#) has suggested that only the highest luminosity AGN require fueling via major mergers; $\sim 90\%$ of AGN across all redshifts are fueled by various other mechanisms which may include minor mergers, flybys, and smooth

accretion, whereby gas is directly accreted via large filaments from the ambient intergalactic medium (Cox et al. 2006; Bellovary et al. 2013; Sinha & Holley-Bockelmann 2012).

Smooth accretion, in particular, may play an important role in fueling these low mass galaxies. Halos less than $10^{11} M_{\odot}$ can accrete filaments of unshocked gas; thereafter, gas will shock heat to the virial temperature of the halo (Keres et al. 2005). Even for massive halos, unshocked gas may still penetrate shocked regions to fuel the galaxy (Brooks et al. 2007; Dekel et al. 2009; Nelson et al. 2013). In addition, SMBH feedback, the depositing of energy and angular momentum back into the gas reservoir during accretion, also affects the overarching structure of the host galaxy (Governato et al. 2009a). Secular processes, including bar formation and disk instabilities, may also be prominent forms of accretion for these SMBHs (Kormendy & Ho 2013).

It is clear that galaxy hosts grow through a variety of channels that depend on mass, environment, and interaction history. Therefore, we want to understand how these different galaxy evolutionary paths translate into SMBH fueling mechanisms, and see how they affect the fueling gas flowing into the SMBH itself. Bellovary et al. (2013) compared simulations of three high mass, high redshift galaxies and found that while mergers and smooth accretion both efficiently build up galaxies, no particular method was more adept at feeding the SMBH. Using a similar method as Bellovary et al. (2013), this work examines the SMBH and galaxy fueling mechanisms of a MW-mass galaxy with a rich merger history. MW-type galaxies host SMBHs on the order of $10^6 M_{\odot}$, which are likely the most common type of massive black hole, yet little is known about them or how they may grow (Kormendy & Ho 2013). Through this examination, we hope to better understand the coevolution of SMBHs and their hosts in this class of galaxy.

In this study, we analyze the Milky Way-type galaxy, h258 (Figure 1), which has a history characterized by major mergers. Since this galaxy is similar to the MW in

virial mass, stellar mass, and circular velocity, without a deeper examination, we may not recognize the turbulent history from which it results. We will examine the origins of gas entering the SMBH and halo to look for clues about SMBH fueling within this galaxy. By examining its assembly, and its SMBH’s fueling, we can determine the accretion rate and gain further understanding about how SMBHs grow in galaxies like our own.

2. Simulation Parameters

Using the smoothed particle hydrodynamics (SPH) N-body tree code: Charm N-body GrAvity solver, ChaNGa, an initial DM-only, uniform resolution 50 comoving Mpc box determined what halo would be selected for zoom-in examination, including the halo examined in this paper. The DM-only simulation assumed a WMAP Year 3 cosmology (Spergel et al. 2007) with the following specifications: $\Omega_m = 0.24$, $\Omega_\Lambda = 0.76$, $H_0 = 73$ km/s, and $\sigma_8 = 0.77$. The halo h258 was chosen for its Milky Way-mass, $8.6 \times 10^{11} M_\odot$, at $z=0$ and its active merger history. The halo has a virial mass of $M_{\text{vir}} = 7 \times 10^{11} M_\odot$ defined relative to a critical density, ρ_c , where $\rho/\rho_c = 100$. A recent major merger characterizes the h258 halo at $z=1.2$. This “zoom-in” high resolution simulation was run on this galaxy and includes gas and star particles using the volume renormalization of Katz & White (1993), resimulating only a few virial radii from the main halo at the highest resolution. The simulation was run from $z=150$ to $z=0$.

While a previous, low resolution version of h258 was run using the Gasoline code, this newer, high resolution version of h258 includes the following properties: a spline force softening length of 174 pc and initial gas particle masses of $2.7 \times 10^4 M_\odot$. Star particles are created with 30% of their parent gas particle mass, allowing a maximum initial mass of $8100 M_\odot$. Each galaxy contains about 5 million DM particles inside the virial radius at $z=0$ and over 14 million DM, star, and gas particles total. The resolution of both force

and mass in these simulations is comparable to the “Eris” simulation which has one of the highest resolutions for an N-body+SPH cosmological simulation of a Milky Way-mass galaxy so far produced. At $z=9$, a uniform UV background is applied to simulate the cosmic reionization energy in a variation of [Haardt & Madau \(2012\)](#). Gas can reach a minimum temperature of ~ 100 K, in the absence of cooling via molecular hydrogen or metals. The simulation includes stochastically-modeled star formation, and once the density threshold and temperature satisfy conditions for star formation (10.0 amu cm^{-3} ; $T < 10^4$ K), gas particles are eligible to form stars with an efficiency of $\epsilon^* = 0.1$. Star particles form along the Kroupa initial mass function ([Kroupa 2001](#)). Supernova feedback releases 10^{51} ergs of thermal energy and affects a “blastwave” radius determined by the equations of [Ostriker & McKee \(1988\)](#). In the affected region, cooling turns off for a time relative to the expansion phase of the SN remnant also determined by the blastwave equation. SN Ia and II from [Thielemann et al. \(1986\)](#) and [Woosley & Weaver \(1986\)](#) are adopted, respectively, and implemented through the [Raiteri et al. \(1996\)](#) method, which uses the stellar lifetime calculations of the Padova group ([Alongi et al. 1993](#); [Bressan et al. 1993](#); [Bertelli et al. 1994](#)) to describe stars with varying metallicities. Both the supernova “blastwave” radius and supernova (Ia and II) prescriptions are described in detail by [Stinson et al. \(2006\)](#). A low-temperature extension to the cooling curve is used to trace metals ([Bromm et al. 2001](#)). Simulated galaxies are shown to conform with the observed Tully-Fisher relation ([Governato et al. 2009b](#)), the size-luminosity relation ([Brooks et al. 2011](#)), and the mass-metallicity relation ([Brooks et al. 2007](#)), in addition to having realistic matter distributions and baryon fractions ([Governato et al. 2009a](#); [Guedes et al. 2011](#)). Parameter and resolution choices described above allow the galaxies to adhere to the stellar-mass-halo-mass relation at $z=0$ and maintain a realistic period of star formation ([Moster et al. 2010](#); [Munshi et al. 2013](#); [Brooks et al. 2007](#); [Maiolino et al. 2008](#)). Given the strict adherence of this simulated galaxy with observations, we are confident that it reasonably represents growth in the galaxy and

its SMBH. We exclude AGN feedback in these small to moderate mass-galaxies; however, we have determined that has little effect on the global properties of the galaxy and SMBH.

Since there are uncertainties in the formation of black holes “seeds,” we implement a BH seeding method that is broadly consistent with several theories of direct collapse black holes (Couchman & Rees 1986; Abel 2002; Bromm & Larson 2004) and Population III stellar remnants (Loeb & Rasio 1994; Eisenstein & Loeb 1995; Koushiappas et al. 2004; Begelman et al. 2006; Lodato & Natarajan 2006). While this method allows the BH formation process to remain physically motivated, BH seeds form if their parent gas particle matches the criteria required for star formation and also maintains zero metallicity (2.5 amu cm^{-3} ; $T < 10^4 \text{ K}$; $Z = 0$) (Stinson et al. 2006). A probability of $\chi_{\text{seed}} \sim 0.01$ is applied to determine whether a gas particle (with the above specifications) will become a BH seed with a mass of $M_{\text{BH}} = 2.28 \times 10^5 M_{\odot}$, the same mass as its parent gas particle. This probability was chosen to match the predicted occupation fraction of BH seeds at $z \sim 3$ (Volonteri et al. 2008).

The requirement that BH seeds must form from zero metallicity gas particles also causes BH formation to be constrained in areas of early star formation bursts, where the earliest and most massive halos are expected to form in the simulation. BH formation is dependent only on local environment, neglecting any large-scale properties of the host halo. Black holes are not fixed within the center of their host, allowing them to be dynamically affected by mergers and other perturbations within the galaxy. Nevertheless, BHs remain near their host centers by choosing dark matter particle masses in high-resolution regions to be on the same order as gas particle masses, minimizing two-body interactions (Bellovary et al. 2011).

Black hole mergers occur when they are separated by less than twice the softening length, and must be bound or satisfy $(1/2)\delta v^2 < \delta a \cdot \delta r$, where δv and δa are the velocity

and acceleration differences between the two black holes and δr is the distance separating them. In addition to gaining mass via merger, black holes gain mass through Bondi-Hoyle gas accretion:

$$\dot{M} = \frac{4\pi\alpha G^2 M_{\text{BH}}^2 \rho}{(c_s^2 + v^2)^{3/2}}, \quad (1)$$

where α is a constant equal to 1, ρ is the density of the surrounding gas, c_s is the sound speed, and v is the black hole’s relative velocity to the gas. Feedback is applied to surrounding gas with an energy boost determined by the accreted mass as follows: $\dot{E} = \epsilon_r \epsilon_f \dot{M} c^2$ where \dot{M} is the accreted mass, and $\epsilon_r = 0.1$ and $\epsilon_f = 0.03$ are assumed for the radiative efficiency and feedback efficiency, respectively. This energy is distributed as thermal energy to the 32 nearest particles via a kernel probability function. Though other groups use a higher value for this efficiency, $\epsilon_f = 0.05$ (Sijacki et al. 2007; Di Matteo et al. 2008), we find that $\epsilon_f = 0.03$ in our code produces MBHs in better agreement with MBH-host galaxy scaling relations. However, as our main concern is in the relative proportion of gas in various phases, our results are not sensitive to our choices of ϵ_r or ϵ_f .

Compared to the previous h258 simulation, the ChaNGa simulated h258 scales better and includes a new improved SPH formalism (Keller et al. 2014). The hydrodynamic treatment now includes a geometric density average— $(P_i + P_j)/(\rho_i \rho_j)$ rather than $P_i/\rho_i^2 + P_j/\rho_j^2$ where P_i and ρ_i are the particle’s pressure and density—in the force expression, in addition to the standard SPH density estimator (Ritchie & Thomas 2001). Adjusting the force expression diminishes tensions in the numerical surface due to shear flows, such as Kelvin-Helmholtz instabilities. We also apply a consistent and entropy-conserving energy equation to account for the modified force expression and correctly model strong shocks, such as Sedov blasts.

3. Reduction Method

The Amiga Halo Finder identifies and sets the virial radii in a simulation using an overdensity criterion for a flat universe (Knebe et al. 2001; Knollmann & Knebe 2009; Gill et al. 2004). We select the primary halo to be the most massive galaxy at $z=0$ and locate the central SMBH. For this simulation, the primary halo in h258 had a final mass on the order of $10^{12} M_{\odot}$ and a formation redshift of $z \sim 4$.

In this analysis, we retrace each gas particle that was accreted by the galaxy or SMBH, following the gas back through its journey in the galaxy and recording its host halo and time of accretion (Brooks et al. 2009). The particles are then classified into types by their method of entrance into the primary halo. In particular, gas particles that exist in the primary halo at the first time step are classified as “early” gas. Gas that belonged to a different halo than the primary prior to accretion is classified as “clumpy,” entering the primary halo through mergers. All other gas is classified as “smooth” accretion, and is then subdivided into two categories: “unshocked” and “shocked.” Unshocked gas will usually flow into the halo via large-scale, dark matter filaments (Bellovary et al. 2013). It’s possible for unshocked gas to be dense enough to pierce an already developed shock, allowing it to funnel into the galaxy core where it can be accreted onto the SMBH.

However, as we discussed in Section 1, if the galaxy halo is $\gtrsim 10^{11} M_{\odot}$, the gas is known to shock heat up to the virial temperature of the halo. We identify shocked particles through an increase in entropy and temperature using the following criteria:

$$T_{\text{shock}} \geq 3/8 T_{\text{vir}}, \quad (2)$$

where T_{vir} is the virial temperature of the halo and T_{shock} is the temperature of the gas particle, and

$$\Delta S \geq S_{\text{shock}} - S_0, \quad (3)$$

where S_0 is the initial entropy of the gas particle, and

$$S_{\text{shock}} = \log_{10}[3/8 T_{\text{vir}}^{1.5}/4\rho_0], \quad (4)$$

where ρ_0 is the gas density prior to the shock. Since our halos are $\sim 10^{12} M_{\odot}$ by $z = 0$, we should expect to find more shocked gas entering the halo.

Once all the gas particles have been individually categorized, we can use these labels to determine which particles are accreting onto the SMBH, and we can better contrast the methods that feed the galaxy and its SMBH in MW-mass halos.

4. Results

The galaxy h258 is characterized by two major mergers; the first occurs at $z \sim 1.8$ and the second at $z \sim 1.2$. Despite its merger-rich history, gas accretion can be seen smoothly increasing the cumulative black hole mass throughout its evolution (Figure 2). Figure 2 shows the cumulative SMBH mass in h258 as a function of time (lower axis) and redshift (upper axis). The black dashed line indicates the total cumulative BH mass (including both mass from gas and BH mergers), while the black solid line indicates the total accreted gas mass. The blue dot-dashed line represents the gas mass accreted via unshocked gas, while the green solid line and red dashed line show the gas mass accreted through mergers and shocked gas, respectively. It is also worthwhile to point out that the largest part of the mass budget at high redshift is not gas at all, but other black holes that have merged with the SMBH seed. This has important implications for gravitational wave astronomy, increasing the event rate for SMBH assembly at high redshifts (Holley-Bockelmann et al. 2010). Aside from this early BH assembly, the largest gain in SMBH mass comes from gas associated with the major mergers between $z \sim 2$ and $z \sim 1$. It's clear from Figure 2, that unshocked gas makes up the majority of gas entering the galaxy at early times. However,

the transitions between when smooth, unshocked accretion and mergers dominate is clearly distinguished. While clumpy gas (green) dominates at the earliest time, unshocked gas (blue) overtakes it for a short time (~ 2 Gyr) before clumpy gas once again dominates by $z \sim 1.5$. This low redshift transition to a clumpy gas preference results in the large fraction of clumpy gas seen in the SMBH (Figure 3). Figure 3 depicts the fractions of total gas in the galaxy (a) and the SMBH (b) at $z=0$, again differentiated by gas origin. Blue, green, and red distinguish unshocked, clumpy, and shocked, respectively. The galaxy has a mass nearly half comprised of gas entering the galaxy through mergers (56 %), with 36 % of the gas entering through unshocked, smooth accretion. The smallest fraction of the total gas is comprised of shocked gas (8 %). Unlike the galaxy, nearly three quarters (74 %) of the gas accreted by the central SMBH was gas accreted via mergers, while only a quarter (24 %) is comprised of unshocked, smoothly accreted gas. Shocked gas makes up the last 2% of total gas entering the SMBH. **It is evident then that the SMBH more readily accretes gas gained through mergers.** This result is contrary to [Bellovary et al. \(2013\)](#) which found that the fractions of gas comprising the SMBH and its host were nearly the same. Through our results, we find that lower mass galaxies can readily employ the physical effects of mergers to feed their SMBH.

In particular, it is important to examine the results of the most massive merger, which occurs at $z \sim 1.2$, and its solitary effect on the SMBH mass budget. By tracing the gas particles in the second largest halo of the zoom-in simulation, we could determine what mass fraction of those particles would become accreted into the galaxy and, at later timesteps, which fraction of their mass would accrete onto the SMBH. In our first attempt at this method, we considered all the gas particles associated with the secondary halo in the timesteps leading up to the merger. From this first analysis, we discovered that nearly 25% of the total gas accreted by the SMBH originated in the secondary halo. However, due to the complex and messy nature of the merger, we examined additionally a single

timestep (during which the secondary halo and main halo had not had much interaction) and found that in this case 18% of the total gas budget of the SMBH was comprised of the secondary halo’s gas. In other words, at minimum a fifth of the total mass budget of the SMBH entered through this major merger. Considering that 74 % of the SMBH’s gas mass is composed of clumpy gas, this also implies that a quarter of the total clumpy gas in the SMBH comes from this main merger alone.

A previous study by [Fu & Stockton \(2007\)](#) supplies direct observational evidence that gas from a merger can be funneled directly into the vicinity of the SMBH (< 1 pc). Their study examined a collection of low-redshift quasars and the metallicity correlation between their extended emission-line regions (EELRs) and the broad-line regions (BLRs) at their centers. When luminous EELRs were found in the extended region of a quasar, the quasar was also found to host a low-metallicity BLR. The quasars without EELRs, however, hosted BLRs with metallicities above Z_{\odot} . The correlation between the presence of EELRs along with low-metallicity BLRs, and the low-metallicity discovered in the accompanying EELRs, allowed [Fu & Stockton \(2007\)](#) to determined a common external origin for the EELRs and the centrally located gas in their observations. An example of this external origin being a merger with another galaxy which is most likely gas-rich. Though our simulations did not examine similar galaxies or use a similar method to categorize the gas in the central SMBH of our main halo, the results still provide strong evidence for mergers that funnel in large amounts of gas into the central region of the galaxy.

To better understand the apparent preference for merger-accreted gas, we examine the angular momentum of each type of gas at the moment it enters the galaxy. [Figure 4](#) shows a cumulative distribution of the gas particles’ angular momentum as they enter the halo (solid lines) and SMBH (dashed lines). The gas is again distinguished by their origin (Clumpy, unshocked, shocked being green, blue, and red, respectively.) and end up

in the halo and/or SMBH by $z=0$. We find that the angular momentum of gas entering the SMBH is lower when comparing between types of gas. We also see that the lowest angular momentum gas is comprised of both clumpy and unshocked gas. This can be seen if we examine Figure 5 which shows the cumulative distribution of the angular momentum of the incoming gas particles at the time of the most massive merger at $z \sim 1.2$. (Colors and linestyles as in Figure 4.) Figure 5 explicitly shows that the gas ending up in the SMBH enters with the lowest angular momentum. There appears to be an influx of gas with low angular momentum feeding the black hole at this timestep that is not associated with the merger occurring at $z \sim 1.2$. This may be attributed to the prior, less massive merger at $z \sim 1.8$ (Figure 7) during and after which a large influx of low angular momentum gas enters the galaxy and SMBH. (Color and linestyles as in Figure 4)

We can also examine another similar cumulative distribution which includes specifically angular momentum of the gas that is associated with secondary halo (the second most massive halo in the zoom-in simulation). Figure 6 has the same angular momentum information for the main halo (solid lines), but the green dashed line now indicates the angular momentum gas *only* associated with the merging secondary halo. We note that a large portion of this gas enters with a higher angular momentum than the cumulative average, yet not nearly as high as the angular momentum of cumulative unshocked and shocked gas.

A lower resolution version of the galaxy h258 was run using the N-body code, Gasoline, like those simulations in [Bellovary et al. \(2013\)](#). Despite fundamental differences in the hydrodynamic implementation and gas physics included, an analysis of this low resolution h258 results in a SMBH with the same distinct preference for accreting gas. The broad consistency between the low and high resolution simulation of the same galaxy indicates that the large scale gravitational dynamics could be a main driver of the SMBH fueling

in this case. We also stress that while major mergers may not be the only physical mechanisms by which gas can be funneled into the centers of galaxies (the previous study being a strong example of this), mergers between galaxies clearly play an important role when considering the gas accretion of SMBHs.

5. Conclusion

This study examines the gas accretion onto the fully cosmological simulation of a Milky Way-size galaxy to redshift $z = 0$, with major mergers characterizing its past. We trace the gas into the SMBH at its center and differentiate the gas accreted onto the galaxy and SMBH by origin. Gas gained through mergers is classified as “clumpy” gas and smoothly accreted gas is separated into “shocked” and “unshocked” categories. Our goal is to determine what types of gas are primarily feeding the SMBH and the galaxies of this class, and to determine what effects the merger history of a galaxy may have on these processes.

A previous study by [Bellovary et al. \(2013\)](#) which analyzed high mass, high redshift galaxies and found the gas composition of the SMBHs mirror their host. Contrary to these previous results, when we examined a galaxy with an active merger history, we determined that the SMBH at the center more readily accretes gas gained through mergers. This remained true both in an older low resolution simulation of the same galaxy as well as this current iteration. **In both the low and high resolution cases, we see a significant increase in the clumpy gas accreted by the SMBH compared to its host.**

The angular momentum of the accreted gas as it enters the galaxy sheds some light on the mechanism driving this preferentially accreted clumpy gas. Smoothly accreted gas, which enters with galaxy with a wide range of angular momentum, may become additionally torqued by the already existing disk. Meanwhile, gas entering through mergers is restricted

by the direction of its entry and allows for a small range of low angular momentum gas to enter the main halo. This gives clumpy gas the advantage of falling more readily to the center and accreting onto the SMBH. Considering all origins of gas, we see a clear distinction wherein lower angular momentum gas preferentially feeds the SMBH.

However, because this result departs from the previous study, we consider that some of the physical processes affecting the galaxy with an active merger history accounts for the ready funneling of low angular momentum to the central region of the black hole.

While the examination of this extreme case of a galaxy with an active merger history depicts a class of galaxy with varying SMBH accretion methods, a further study of cases with varying merger histories is required to begin understanding the broad spectrum of Milky Way-mass galaxy accretion. It is clear through this study that the presence of major mergers can play an important role in the final compositions of central SMBHs, but the question of how important these mergers are remains to be seen.

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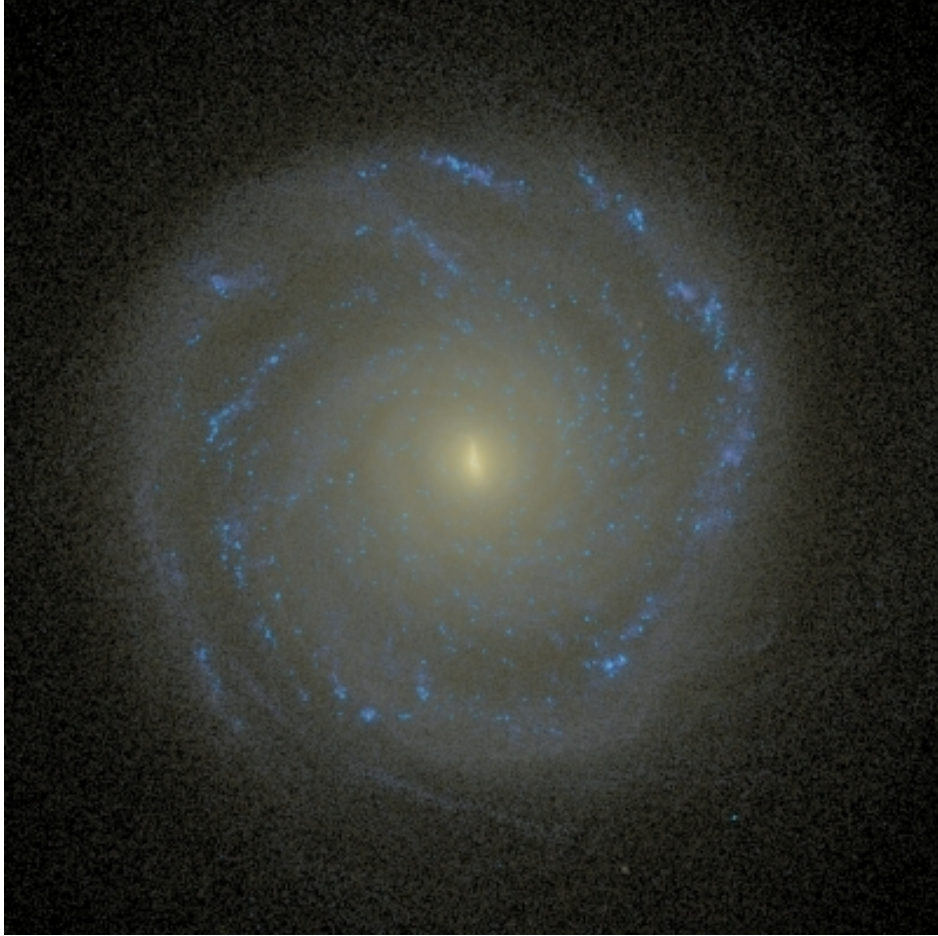


Fig. 1.— Face-on Sunrise images of our ChaNGa galaxy, h258, which has an active, merger-rich history. NEED TO REPLACE, this is low res h258

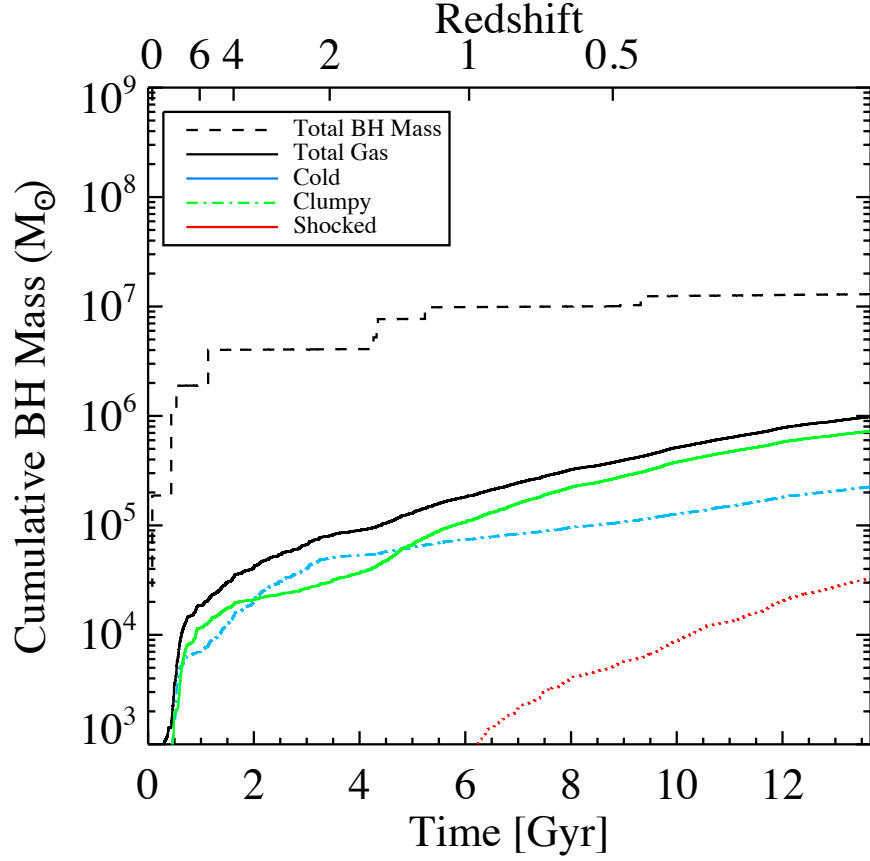


Fig. 2.— The central BHs cumulative mass as a function of time and redshift. The black dashed line indicates the total cumulative BH mass. The black solid line indicates the total gas mass. The blue dot-dashed line indicates the gas mass accreted via unshocked gas. The green solid line indicates the gas mass accreted through mergers. The red dashed line indicates gas mass that was shocked upon entry into the halo.

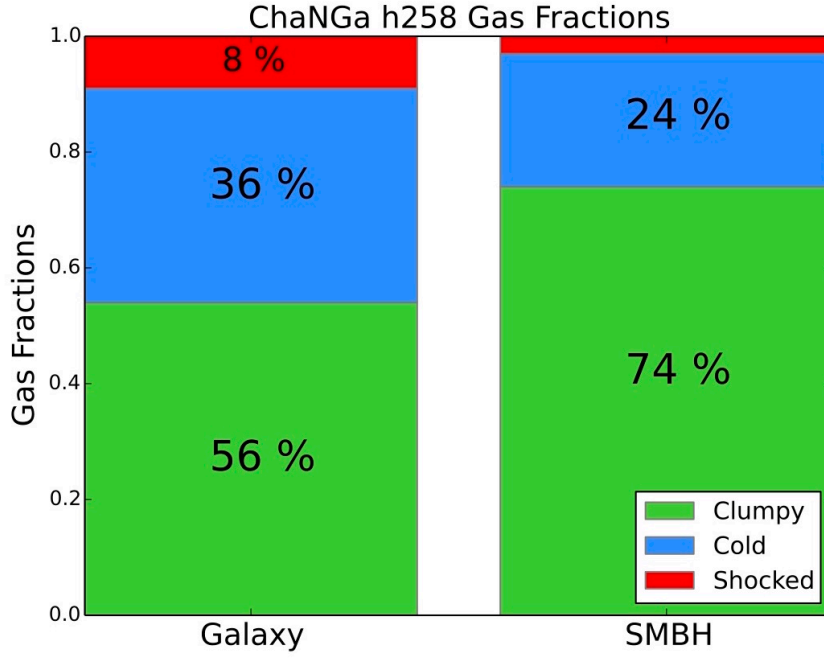


Fig. 3.— Gas fractions of the gas particles accreted by the high resolution ChaNGA h258 by the main halo (left) and the SMBH (right), distinguished by type. Blue, green, and red distinguish gas gained through unshocked gas, gained through mergers, and gas shocked upon entry, respectively. Yellow indicates gas that existed within the main halo upon formation; this “early” gas is negligible ($< 1\%$) within the SMBH.

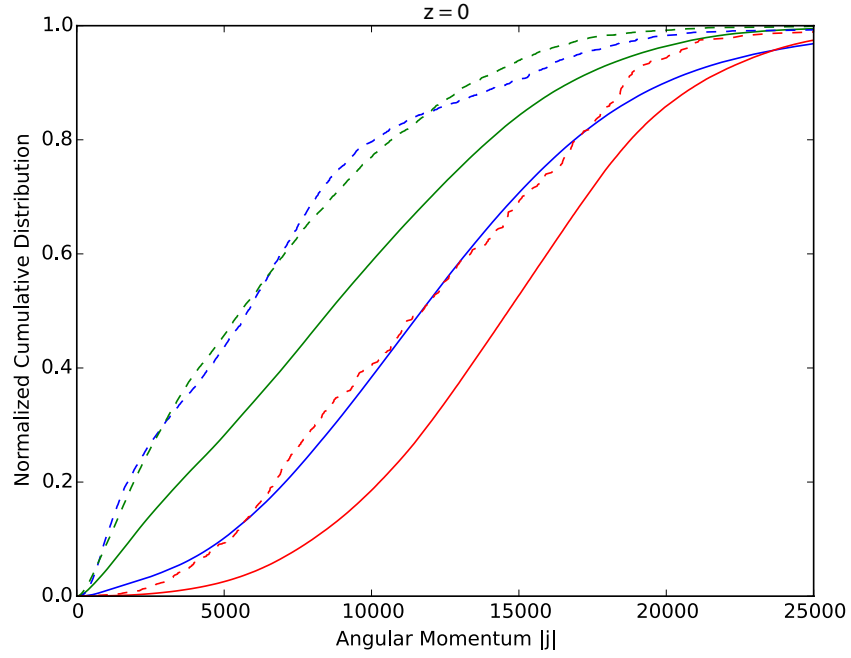


Fig. 4.— Cumulative distribution of angular momentum of the gas particles accreted onto the high resolution ChaNGa h258. Gas particles accreted onto the main halo (solid lines) and central black hole (dashed lines). The green, blue, and red lines indicate clumpy, unshocked, and shocked gas, respectively.

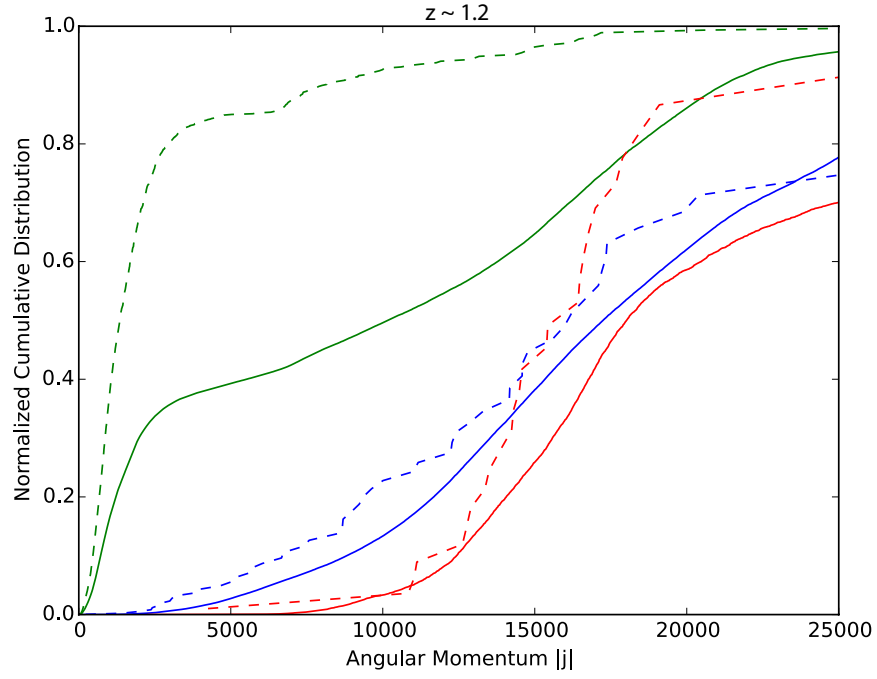


Fig. 5.— Cumulative distribution of angular momentum of the gas particles accreted onto the high resolution ChaNGA h258 galaxy at the time of the primary major merger ($z \sim 1.2$). There are about 450,000 gas particles accreted at this timestep, which gas particles accreted onto the main halo and central black hole are distinguished by solid and dashed lines. The green, blue, and red lines indicate clumpy, unshocked, and shocked gas, respectively.

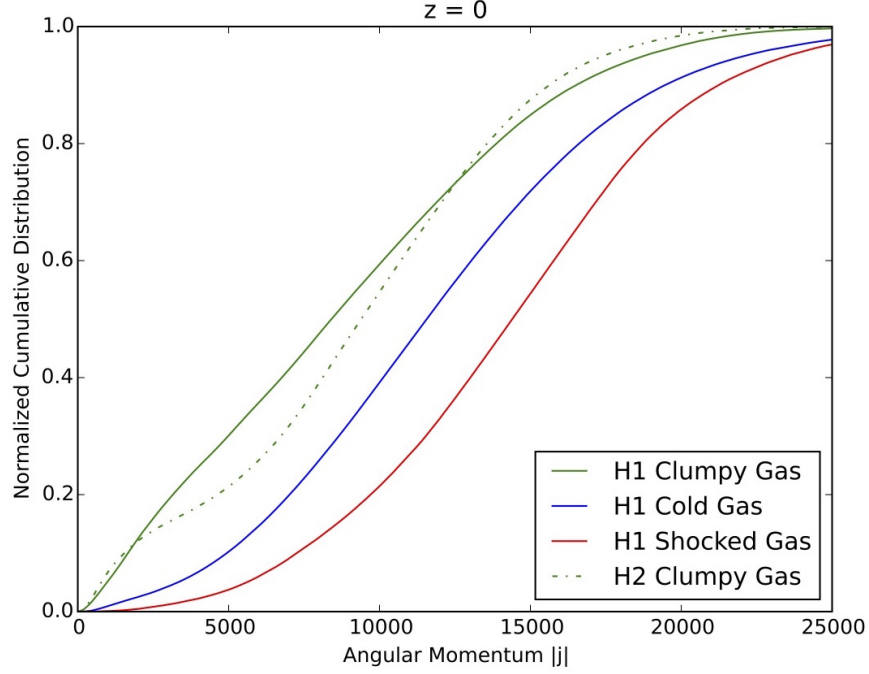


Fig. 6.— Cumulative distribution of angular momentum of the gas particles accreted onto the high resolution ChaNGA h258 galaxy at the time of the primary major merger ($z \sim 1.2$). The solid lines indicate the angular momentum of gas accreted onto the main halo (similar to Figure 4) and the dashed line indicates the angular momentum of gas entering from the secondary halo. As before, green, blue, and red lines indicate clumpy, unshocked, and shocked gas, respectively.

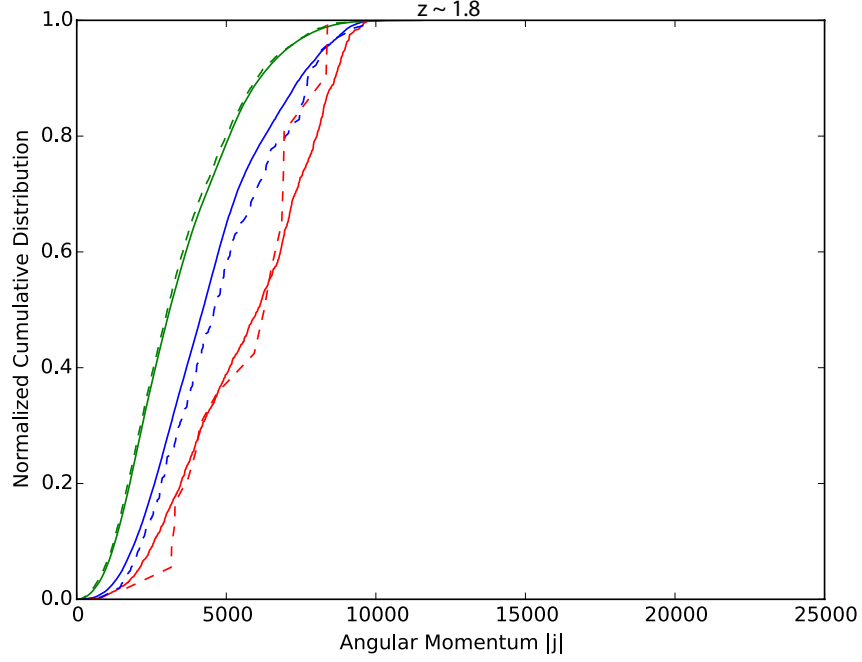


Fig. 7.— Cumulative distribution of angular momentum of the gas particles accreted onto the high resolution ChaNGA h258 galaxy at the time of the less massive merger ($z \sim 1.8$). There are about 900,000 gas particles accreted at this timestep, which gas particles accreted onto the main halo and central black hole are distinguished by solid and dashed lines. The green, blue, and red lines indicate clumpy, unshocked, and shocked gas, respectively.

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