

# A Catalog of Point Sources Towards NGC 1333

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## ABSTRACT

I present a catalog of point source objects towards NGC 1333, resolving a wide variety of confusion about source names (and occasionally positions) in the literature. I incorporate data from optical to radio wavelengths, but focus most of the effort on being complete and accurate from  $J$  ( $1.25\ \mu\text{m}$ ) to  $24\ \mu\text{m}$ . The catalog encompasses  $52^\circ < \text{RA} < 52.5^\circ$  and  $31^\circ < \text{Dec} < 31.6^\circ$ . Cross-identifications include those from more than 25 papers and catalogs from 1994-2014, primarily those in wide use as origins of nomenclature. Gaps in our knowledge are identified, with the most important being a lack of spectroscopy for spectral types or even confirmation of youth and/or cluster membership. I fit a slope to the spectral energy distribution (SED) between 2 and  $24\ \mu\text{m}$  for the members (and candidate members) to obtain an SED classification, and compare the resulting classes to those for the same sources in the literature, and for an SED fit between 2 and  $8\ \mu\text{m}$ . While there are certainly differences, for the majority of the sources, there is good agreement.

*Subject headings:* stars: pre-main sequence – stars:protostars – catalogs

## 1. Introduction

NGC 1333 is one of the youngest and most well-studied star forming regions, in part because it is located at only  $\sim 235$  pc (Hirota et al. 2008, 2011). Its stars are thought to have an average age of 1-2 Myr (e.g., Bally et al. 2008), but it also contains several Class 0 objects, objects in the earliest stages of star formation (see, e.g., Sadavoy et al. 2014 or Sandell & Knee 2001).

Because this region is very young and relatively nearby, it has been a subject of intensive study for decades. However, there has not yet been a published, comprehensive merging of all of the large catalogs in this region. A summary including each of the prior investigations of NGC 1333 is beyond the scope of this paper; see Walawender et al. (2008) for a recent review. In this paper, my primary goal is to merge all of the available data from the (relatively) large-field surveys in this region, and assemble one master catalog with all of the names from the various surveys reconciled. We are studying NGC 1333 as part of the Young Stellar Object VARIability (YSOVAR) project (Rebull et al. 2014, hereafter R14); this catalog was assembled originally as part of that effort. Because my original goal in taking on this task was to focus on the sources for which we have light curves in YSOVAR, some of the data reduction relevant to this paper is described in the YSOVAR overview paper (R14). Some of the detailed investigation described in the present paper is focused on the region mapped by YSOVAR; the region mapped by YSOVAR is the heart of the cluster. The NGC 1333 YSOVAR data are discussed in detail in Rebull et al. (2015, in preparation).

The catalog is somewhat artificially limited to being within  $52^\circ < \text{RA} < 52.5^\circ$  and  $31^\circ < \text{Dec} < 31.6^\circ$ . Data are available over a larger region for some surveys, but this region should include most of the objects

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actually belonging to NGC 1333, and it entirely includes the region monitored for YSOVAR, as per the original goals in assembling this catalog.

The primary reason I compiled data from several different sources was in order to assemble spectral energy distributions (SEDs). While, of course, many of the targets vary significantly with time, single-epoch archival data can help define the SED such that, in some cases, the assembled SED can reveal the underlying nature of the source, or at least help narrow the possibilities for the nature of the source.

In this paper, I first review the large surveys that I included in our catalog (Sec. 2), and then list the papers from which I drew data, and explain why I associated or dissociated (or removed) sources across wavelengths (Sec. 3). I describe obvious gaps in the literature (Sec. 4). I place the YSO candidates into SED classes based on the slope between 2 and 24  $\mu\text{m}$ , and compare them to other schemes from the literature (Sec. 5).

## 2. Large Datasets

In this section, I discuss the largest area surveys I used; they are listed, along with the smaller catalogs from the next section, in Table 1.

All of the large-area catalogs described here were merged by position with a catalog-dependent search radius (usually  $\sim 1''$ ). Many sources, especially those in regions where many sources are close together on the sky, were individually inspected and matched by hand. SEDs were constructed as an additional check on the source matching; objects with particularly strange initial SEDs were also individually inspected and matched by hand, which often resolved any SED issues.

### 2.1. 2MASS and 2MASS 6 $\times$

The Two-Micron All Sky Survey (2MASS; Skrutskie et al. 2006) catalog provided the original nucleus of the catalog, to which all other catalogs were merged by position, typically with a  $\sim 1''$  search radius.

NGC 1333 was in the original 2MASS survey, of course, and it was also included in the long exposure 6 $\times$  2MASS program. The original 2MASS data were obtained in 2000. I included the main 2MASS catalog and the deeper 6 $\times$  catalog near-IR (NIR)  $JHK_s$  data into the database.

### 2.2. Cryogenic-era Spitzer Archival Data

Early in the Spitzer mission, NGC 1333 was observed by both the guaranteed time observations (GTO) and the original Cores-to-Disks (c2d) Legacy program (Evans et al. 2003, 2009). For both the Infrared Array Camera (IRAC; Fazio et al. 2004) and Multiband Imaging Photometer for Spitzer (MIPS; Rieke et al. 2004) data, the observations were obtained at multiple epochs separated by at least a few hours to allow asteroids to move and thus be identified as asteroids (as opposed to embedded objects in NGC 1333). The IRAC observations are at 3.6, 4.5, 5.8, and 8  $\mu\text{m}$ ; the MIPS observations are at 24, 70, and 160  $\mu\text{m}$ , but the 160  $\mu\text{m}$  data is very difficult to interpret in this region and are not included here. The first IRAC observation was part of GTO program 6, obtained on 2004-02-10; a second group of IRAC observations was part of c2d, program 178, on 2004-09-08. The first MIPS observations were part of GTO program 58, obtained on

Table 1. Overview of Studies and Data Included

Dataset	Year published (or obtained)	Band(s)	notes
ASR	1994	<i>JHK</i>	many coordinate issues; see text
LAL	1996	<i>JHK</i>	many coordinate issues; see text
Preibisch	1997	X-ray (ROSAT)	two coordinate issues; source 2 may be spurious
VLA	1999	3.6, 6 cm	hard to find short-wavelength counterparts
2MASS	(2000)	<i>JHK<sub>s</sub></i>	data obtained in 2000 as part of 2MASS all-sky survey and subsequently the 2MASS 6× survey
Getman (et al.)	2002	X-ray (Chandra/ACIS)	two inconsistencies
Preibisch	2003	X-ray (XMM)	most sources match Chandra sources
Rebull (et al.)	2003	MIR	primary information included from this work is source multiplicities
Aspin	2003	NIR	spectroscopy; spectral types included
MBO	2004	<i>JHK</i>	duplicate MBO number
c2d	2006, 2007	Spitzer: 3.6 to 70 $\mu$ m	included limits and 70 $\mu$ m sources
Hatchell (et al.)	2007	450, 850 $\mu$ m	cross ids retained
Greissl (et al.)	2007	<i>JH</i>	NICMOS brown dwarf search
Gutermuth (et al.)	2008	Spitzer: 3.6 to 24 $\mu$ m	only YSO candidates reported
Oasa (et al.)	2008	<i>JHK<sub>s</sub></i>	brown dwarf search
Scholz (et al.)	2009, 2012ab	<i>i'z'JK</i>	SONYC; brown dwarf search
Winston (et al.)	2009, 2010	Spitzer/IRAC+Chandra/ACIS	some potentially confusing source numbers in literature; fixed here
Itoh (et al.)	2010	optical	spectroscopy of brown dwarf candidates
WISE	(2010)	3.5, 4.6, 12, 22 $\mu$ m	data obtained in 2010 as part of All-Sky Survey; AllWISE reduction used, but only in certain very specific cases (not broadly nor blindly used).
Arnold (et al.)	2012	MIR	spectroscopy with Spitzer IRS
3XMM-DR4	2013	X-ray (XMM)	all-sky catalog from XMM
Foster (et al.)	2015	NIR	spectroscopy with APOGEE
Sadavoy (et al.)	2014	FIR	Class 0 Herschel cross-identifications
YSOVAR	2014	3.6, 4.5 $\mu$ m	means from YSOVAR monitoring included; variables not yet recognized as YSOs are included here

2004-02-03; an additional 3 epochs were part of c2d program 178, obtained on 2004-09-19 and 2004-09-20.

As discussed in R14, the cryogenic data were combined and reduced identically to the YSOVAR monitoring data, except using cryogenic calibrations and combining the two observations into a single effective epoch (rather than maintaining separate measurements for each epoch). The apertures we used were  $2.4''$ . The data were bandmerged across Spitzer bands by position, and then to the 2MASS catalog, within a search radius of  $1''$ .

Gutermuth et al. (2008a, 2009, 2010) present methodology for identifying YSOs from the cryogenic catalog. The details of the selection process appear in those papers, but in summary, multiple cuts in multiple color-color and color-magnitude diagrams are used to identify YSO candidates, as distinct from, e.g., extragalactic and nebular contamination. This color selection process was run anew on the re-reduced data.

Spitzer data are also available from the c2d program final data delivery. The data used for the final delivery are typically the same BCDs as were used in the cryogenic data that we re-reduced above. As such, then, they are not independent measurements, and these data were only used to supplement the cryogenic-era catalog if a band was missing. There is more information on what I extracted from the c2d catalogs below.

### 2.3. Chandra ACIS

Chandra X-ray Observatory Advanced CCD Imaging Spectrometer for wide-field imaging (ACIS-I) observations of NGC 1333 were first reported in Getman (2002) and then Winston et al. (2009, 2010). There are three pointings in NGC 1333, with obsids 642, 6436, and 6437, with a total exposure time of 119.3 ks.

As was described in R14, we re-reduced the Chandra data in a self-consistent way across most of the YSOVAR clusters. Source detection was performed using CIAO (Chandra Interactive Analysis of Observations; Fruscione et al. 2006). Sources, even faint ones, were retained if they had a counterpart in the cryogenic IRAC catalog. Sources from Chandra were matched to the rest of the catalog with a position-dependent search radius; see R14.

Cross-IDs from these X-ray papers are included; see below. However, the X-ray flux measurement data are included in their entirety in Rebull et al. (2015), the paper on the NGC 1333 YSOVAR data, and do not play a role in the SEDs, so they are not explicitly included here. As R14 describes, we identify some candidate cluster members by looking for objects with star-like SEDs and an X-ray detection. There are only two new X-ray candidate cluster members that are introduced as part of this process (J032913.47+312440.7 & J032837.85+312525.3); all of the other so-identified members were identified in the literature as (candidate) members already.

### 2.4. WISE

The Widefield Infrared Survey Explorer (WISE; Wright et al. 2010) surveyed the whole sky at 3.5, 4.6, 12, and  $22\ \mu\text{m}$ ; all of the available WISE data taken between 2010 Jan and 2011 Feb were incorporated into the AllWISE catalog. WISE has lower spatial resolution than Spitzer, and is on average less sensitive. I do not generally include the AllWISE data, since NGC 1333 is often a complicated region with high surface

brightness, and because we have extensive higher spatial resolution Spitzer data. However, WISE provides a band at  $12\ \mu\text{m}$  that is not available from Spitzer. I have incorporated WISE data for certain individual sources, where viable photometry is available from WISE and the photometry from Spitzer is incomplete or results in an unusual SED shape. (For example, in some cases, the [24] point seemed unphysical in the context of the SED, but the [22] point is well-matched to it.) Comparison of the WISE images with the Spitzer images was also useful in certain circumstances, such as for investigating the influence of image artifacts – artifacts change between Spitzer and WISE, but sources on the sky should not.

We used the AllWISE data release in nearly all cases, but we now note an exception. In one very crowded region (sources that are components of IRAS 7), the WISE All-Sky catalog does a better job of separating the sources than the AllWISE catalog. AllWISE seems to have inferred that there was significant proper motion of one of the sources in the clump, and by inspection of the images, this is not correct. Thus, WISE flux densities for our catalog sources J032910.70+311820.9 and J032911.24+311831.8 are taken from the All-Sky catalog, not the AllWISE catalog.

### 3. Literature data and source reconciliation

Many studies have been made specifically of NGC 1333 – it is one of the most well-studied star forming regions, with >200 refereed publications in ADS. It is difficult to include data from every single paper, especially since so many papers focus on just one or a few objects in the region, or on just extended objects (e.g., Raga et al. 2012). I endeavored to include in the catalog the most recent catalogs of point sources and/or those that had the largest repositories of data, and/or those that were the origin of some source names still in common use today. The majority of the information actually included in the database from the earliest studies is the cross-identifications (cross-IDs) with the literature; additional broadband photometry was included where possible, and not superseded by subsequent reprocessing of the data. In the process of assembling the literature catalog of sources, I reconciled many ambiguities and inconsistencies in the literature. I provide below descriptions of the more complicated issues. All of the reconciled cross-identifications in NGC 1333 are included in Table 2, including an indication of whether or not the corresponding survey identified the object as a YSO.

All of the aggregate  $J$ -[24] single-epoch photometry appears in Table 3. There are nearly 7000 objects in the catalog, about 300 of which are identified in the literature as YSO candidates.

IAU standards recommend not renaming previously-identified sources, but as one of the purposes of the present catalog is to sort out inconsistencies and inaccuracies in existing catalogs, assigning a new, coordinate-based name seems appropriate. The coordinate based names presented here (in Tables 2 and 3) should be preceded by ‘R15-NGC1333.’

Figure 1 gives a rough indication of the various larger surveys included here. The footprints from Chandra and from the YSOVAR monitoring are shown close to their actual coverage. No attempt is made to capture complex polygons of coverage for the other surveys, but just squares encompassing all sources are shown; in other words, there are no data from a given survey outside of its square, but there may be incomplete spatial coverage inside of it.

### 3.1. On Coordinate Accuracies

Integral to the process of source matching across catalogs is an understanding of the systematic and statistical errors present in the positions of the objects in the catalogs. 2MASS provides a very high quality coordinate system,  $<0.1''$  with respect to the International Celestial Reference System (ICRS) reference frame for bright sources, over the whole sky. WISE and Spitzer coordinates are fundamentally tied to this 2MASS coordinate system. WISE positional uncertainties are typically  $<0.2''$ , often much less. Spitzer/IRAC positional uncertainties are comparable at  $<0.2''$ .

I present the prior studies I integrated in this section in roughly chronological order. However, in practice, I iteratively (and very often manually) compared each of the sources in the older catalogs to the 2MASS catalog and images, adjusting or correcting coordinates as needed, before merging all the catalogs together. In many cases, it was a simple shift within  $\lesssim 5''$  from a place on the sky without a 2MASS source to the location of a relatively bright 2MASS source. I note that the Infrared Science Archive (IRSA) tool FinderChart<sup>1</sup> was extremely helpful for this process. FinderChart uses data from WISE, 2MASS, and the Digitized Sky Survey (DSS), which is a digitization of the photographic sky survey plates from the Palomar (Palomar Observatory Sky Survey) and UK Schmidt telescopes. FinderChart provides a thumbnail image of the sky at multiple wavelengths, making comparisons relatively straightforward.

For any catalog, the coordinate accuracy depends on there being sufficient numbers of point sources in the field of view (or mosaic) in order to anchor the coordinate system. This has two ramifications for the catalogs considered in this section. First, for some early NIR observations, arrays were very small (for example, Aspin et al. 1994 had a single pointing field of view of just under an arcminute on a side), resulting in relatively few point sources per pointing, making astrometry very hard to do accurately. Moreover, then as now, astrometry was bootstrapped to prior observations, but at that time, this bootstrapping had to occur without the reliable all-sky anchor provided by 2MASS. Secondly, for long-wavelength observations such as Spitzer/MIPS and Herschel maps, there are many fewer point sources that can be linked to the 2MASS coordinate system, so the positional uncertainty can be worse, and sometimes astrometry relative to one or a few point sources is the best available. It can be difficult to make a clear correspondence between the shorter wavelength and longer wavelength sources, not just because of the coordinates, but also because the emission may not be coming from the same location in/near the object. I have attempted to make these matches here, being mindful of the fact that these long wavelength sources are often in regions that are extremely complex, with high surface brightness (complicating both photometry and astrometry). Moreover, the source of the long wavelength emission may just be by chance aligned with an emitter of short wavelength emission. I believe what I have done is correct, but I have provided descriptions below of what I have done in the event that subsequent investigators disagree.

Another source of positional uncertainty is the space motions of these objects. Over the time baselines considered here, could the sources in NGC 1333 be moving significantly? Karchenko et al. (2013) report that for NGC 1333 the average  $\mu_{\text{RA}}, \mu_{\text{Dec}} = 5.51, -10.28 \text{ mas yr}^{-1}$ , so over 50 years (longer than considered here), a typical object could move  $\sim 0.5''$ . However, the spatial resolution for both 2MASS and IRAC are both  $\sim 1''$ , and these catalogs are the input catalogs that establish the bulk of the master NGC 1333 catalog. I search for counterparts between catalogs with a typical tolerance of  $\sim 1''$ . Therefore, most sources will not have enough proper motion over the baselines considered here to affect our catalog merging. Assessing the motion of these objects based on these data is theoretically possible, but beyond the scope of this work.

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<sup>1</sup><http://irsa.ipac.caltech.edu/applications/finderchart/>

Table 2. Contents of Cross-ID Catalog<sup>a</sup>

Bytes	Format)	Units	Label	Explanations
1- 18	A18	—	cat	Catalog name (HHMMSS.ss+DDMMSS.s; J2000); should be preceded by ‘R15-NGC1333 J’
19- 23	I5	—	ASR	ASR number
24- 27	I4	—	ASR YSO?	Did ASR tag it as a YSO? 1=yes, 0=no, -9 no information
28- 32	I5	—	LAL	LAL number
33- 36	I4	—	LAL YSO?	Did LAL tag it as a YSO? 1=yes, 0=no, -9 no information
37- 41	I5	—	VLA	VLA number
42- 45	I4	—	VLA YSO?	Did Rodriguez et al. tag it as a YSO? 1=yes, 0=no, -9 no information
46- 50	I5	—	Preibisch	Preibisch number
51- 70	A20	—	2MASS num	2MASS name
71- 75	I5	—	Getman num	Getman number
76- 79	I4	—	Getman var	Did Getman tag it as <i>variable</i> ? 1=yes, 0=no, -9 no information
80- 84	I5	—	MBO num	MBO number
85-103	A19	—	c2d num	c2d name
104-107	I4	—	c2d YSO?	Did c2d tag it as a YSO? 1=yes, 0=no, -9 no information
108-117	A10	—	Greissl num	Greissl et al. (2007) number
118-123	A6	—	Hatchell num	Number from Hatchell et al. (2007)
124-128	I5	—	Oasa num	Number from Oasa et al. (2008) number
129-132	I4	—	Oasa YSO?	Did Oasa et al. tag it as a member? 1=yes, 0=no, -9 no information
133-137	I5	—	G08 num	Number from Gutermuth et al. (2008)
138-141	I4	—	G08 YSO?	Did Gutermuth et al. tag it as a member? 1=yes, 0=no, -9 no information
142-146	I5	—	S09 num	Number from Scholz et al. (2009)
147-151	I5	—	Winston num	Number from Winson et al. (2010)
152-161	A10	—	S12 num	Number from Scholz et al. (2012)
162-165	I4	—	S12 YSO?	Did Scholz et al. tag it as a member? 1=yes, 0=no, -9 no information
166-170	I5	—	A12 num	number from Arnold et al. (2012)
171-174	I4	—	A12 YSO?	Did Arnold et al. tag it as a YSO? 1=yes, 0=no, -9 no information
175-179	I5	—	Foster num	number from Foster et al. (2015)
180-207	A19	—	YSOVAR	YSOVAR name from Rebull et al. (2015)
208-211	I4	—	YSOVAR YSO?	Did Rebull et al. tag it as a YSO? 1=yes, 0=no, -9 no information
212-317	A105	—	other	Any other names or cautions for this object

<sup>a</sup>Entire data table available online at the journal.

Table 3. Contents of Single-Epoch 1-70 $\mu$ m Catalog<sup>a</sup>

Bytes	Format)	Units	Label	Explanations
1- 18	A18	—	cat	Catalog name (HHMMSS.ss+DDMMSS.s; J2000)
19- 28	F10.6	deg	RA	RA, J2000, decimal degrees
29- 38	F10.6	deg	Dec	Dec, J2000, decimal degrees
39- 42	A4	—	l_J	Limit flag on J
43- 49	F6.2	mag	J	J band magnitude
50- 55	F6.2	mag	e_J	Uncertainty in J
56- 59	A4	—	l_H	Limit flag on H
60- 66	F6.2	mag	H	H band magnitude
67- 72	F6.2	mag	e_H	Uncertainty in H
73- 76	A4	—	l_Ks	Limit flag on Ks
77- 83	F6.2	mag	Ks	Ks band magnitude
84- 89	F6.2	mag	e_Ks	Uncertainty in Ks
90- 93	A4	—	l_[3.6]	Limit flag on [3.6]
94-100	F6.2	mag	[3.6]	Spitzer/IRAC 3.6 $\mu$ m band magnitude
101-106	F6.2	mag	e_[3.6]	Uncertainty in [3.6]
107-110	A4	—	l_[4.5]	Limit flag on [4.5]
111-117	F6.2	mag	[4.5]	Spitzer/IRAC 4.5 $\mu$ m band magnitude
118-123	F6.2	mag	e_[4.5]	Uncertainty in [4.5]
124-127	A4	—	l_[5.8]	Limit flag on [5.8]
128-134	F6.2	mag	[5.8]	Spitzer/IRAC 5.8 $\mu$ m band magnitude
135-140	F6.2	mag	e_[5.8]	Uncertainty in [5.8]
141-144	A4	—	l_[8.0]	Limit flag on [8.0]
145-151	F6.2	mag	[8.0]	Spitzer/IRAC 8.0 $\mu$ m band magnitude
152-157	F6.2	mag	e_[8.0]	Uncertainty in [8]
158-161	A4	—	l_[24]	Limit flag on [24]
162-168	F6.2	mag	[24]	Spitzer/MIPS 24 $\mu$ m band magnitude
169-174	F6.2	mag	e_[24]	Uncertainty in [24]
175-178	A4	—	l_[70]	Limit flag on [70]
179-185	F6.2	mag	[70]	Spitzer/MIPS 70 $\mu$ m band magnitude
186-191	F6.2	mag	e_[70]	Uncertainty in [70]
192-204	A12	—	SpTy	Spectral type
205-215	A10	—	SpTySrc	Origin of spectral type (literature)
216-221	I4	K	Teff	$T_{\text{eff}}$ from Foster et al. (2015)
222-227	F6.2	—	oursedslope24	our SED slope from 2 to 24 $\mu$ m if YSO candidate
228-233	A6	—	oursedclass24	our SED class using the slope from 2 to 24 $\mu$ m if YSO candidate, else 'notY'
234-239	F6.2	—	oursedslope8	our SED slope from 2 to 8 $\mu$ m if YSO candidate
240-245	A6	—	oursedclass8	our SED class using the slope from 2 to 8 $\mu$ m if YSO candidate, else 'notY'
246-252	A7	—	G08class	SED class from Gutermuth et al. (2008)
253-259	A7	—	A12class	SED class from Arnold et al. (2012)
260-266	A7	—	H07class	SED class from Hatchell et al. (2007)
267-274	A7	—	S14class	SED class from Sadavoy et al. (2014)

<sup>a</sup>Entire data table available online at the journal.



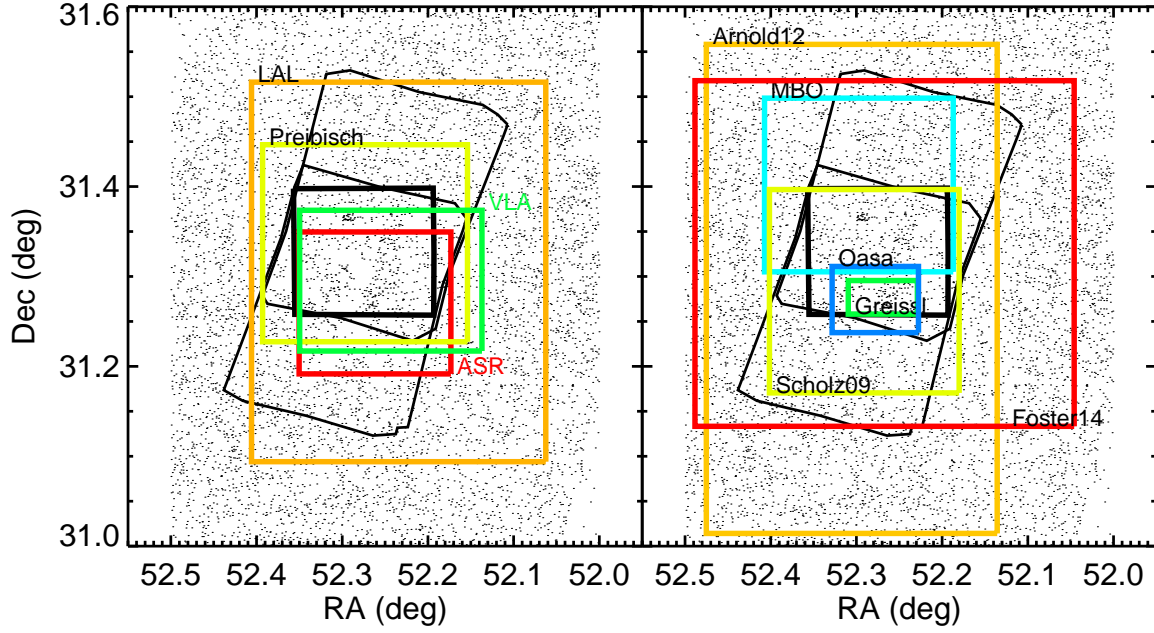


Fig. 1.— Approximate spatial coverage of the various catalogs included in this survey, distributed over two plots simply for clarity. Each small black dot corresponds to an object in the final catalog assembled here. The black polygons in both plots are the coverage of the 3.6 and 4.5  $\mu\text{m}$  YSOVAR monitoring regions, and the thick black square in both plots is the Chandra coverage. Colored squares correspond roughly to most of the other surveys included here, as indicated in the corners of the squares. For these colored squares, note that what is shown is the min/max of the RA/Dec; no attempt is made at polygon representation of the other survey regions, because this figure is simply meant to give a rough indication of the regions of the cluster observed in these respective surveys, e.g., a relatively small region has been probed with NICMOS (the green Greissl et al. square).

### 3.2. ASR Catalog

Aspin et al. (1994; ASR) was one of the first wide-field NIR surveys of NGC 1333 reported in the literature. They published a full *JHK* catalog (i.e., not just the objects they thought were cluster members) for 134 objects to  $K = 16.2$ . I incorporated their full catalog into our database, keeping track of the objects tagged by ASR as likely YSOs. Numbers from this catalog are preceeded by “ASR.” *JHK* measurements from this catalog were not retained because there were 2MASS measurements available. Moreover, the photometry on average is well-matched, typically within  $\sim 0.2$  mag (often less); while most of the stars in the catalog are unlikely to be members, young stars are expected to be intrinsically variable at levels greater than a few tenths of a magnitude.

Since this is the first (and oldest) literature catalog I consider here, I report some additional information on it as representative of the difficulties inherent in this process, and my approach to it. While an individual field of view was just under an arcminute on a side, ASR mapped  $\sim 10' \times 10'$ . They established astrometry by comparison to Herbig & Jones (1983) for 30 objects, estimating uncertainties of  $\sim 1''$  as a result. The original ASR catalog was reported in B1950 coordinates, which I precessed and overlaid on the images and compared to the 2MASS and IRAC catalogs. ASR was a relatively shallow  $K$ -band observation, comparable to 2MASS classic (as opposed to 2MASS  $6\times$ ). These sources (particularly the NGC 1333 members) are intrinsically variable. However, while I expected that variability at  $K$  could make the object appear or disappear out of a relatively shallow survey, I still expected it to show up at the 2MASS  $6\times$  depth or at least the IRAC maps, not necessarily gone entirely. The IRAC observations reach at the very least the same sources as 2MASS, for most SEDs; even relatively shallow IRAC observations commonly reach sources fainter than are detected by 2MASS. For nearly 90% of the 134 ASR sources, I could find counterparts within  $5''$  in the 2MASS catalog, with a median positional offset of  $\sim 1.7''$ , and larger offsets in RA than in Dec; see Figure 2. For most of the sources, particularly the larger positional offsets, images from 2MASS were compared to the positions provided in the original ASR to ensure that the match was appropriate. In most cases, the match was readily apparent.

Unlike most of the rest of the cluster, the region containing the trio of ASR 31/32/33 is quite confusing, as can be seen in Figure 3. I now discuss this region in detail.

ASR 33 is the easiest of this triplet, and is probably the same as MBO 171 (which is the same cross-identification provided by Wilking et al. 2004). It is catalog source R15-NGC1333 J032911.51+311828.6.

VLA 27 has been tied (including by myself in Rebull et al. 2003) to one of the ASR 31/32/33 sources. But note that there is an IRAC source coincident with the position reported for VLA 27, offset (by  $\sim 4.5''$ , a significant amount by IRAC standards) from ASR 31/32/33. And, note that there is no 2MASS source (in the shallower or deeper 2MASS data). The original ASR survey could not have been enough deeper than 2MASS  $6\times$  to detect a source not recovered by either of the 2MASS surveys; intrinsic variations of this level in this source are possible but unlikely. I have tied VLA 27 to the IRAC source (R15-NGC1333 J032700.47+313725.9), but not the ASR sources. The VLA positions should be very good, so leaving this source affiliated with the IRAC source with which it is coincident seems appropriate.

There is a rough arc of 2MASS sources with the same rough concavity as the arc of ASR 31/32/33 sources. But there are 5 total 2MASS sources, and only three ASR sources. Could the northern three 2MASS sources be the three ASR sources? Wilking et al. (2004), a.k.a. the MBO source numbers below, linked ASR 31 to MBO 161. Given the catalog sources overlaid in Fig. 3, MBO 161 is *not* the same as ASR 31.

The original ASR astrometry has some uncertainties, but not as bad as would be required to enforce the match of the two arcs of sources, especially over this small region. I obtained the original data from ASR (C. Aspin, priv. comm.), which had been updated with more recent astrometry; see Figure 3. There are two sources in the left corner of the ASR data bright enough to be in 2MASS classic. Note that the originally reported MBO, IRAC, and 2MASS source positions are a little to the south of the northern source, and a little to the north of the southern source, which should be representative of the net positional errors in this immediate region. The original ASR coordinates in this region also reflect this tiny distortion in the North-South direction, less so in the East-West direction. The local astrometry is probably correct (unless there is a tile boundary here, which does not seem to be the case in the ASR image; see also LAL coordinate matching below). So the differential astrometry should be high-quality among the local ASR sources. There is nothing approaching the magnitude of the distortions that would be required to map ASR 31 onto MBO 161, and ASR 32 onto the 2MASS source that would be in between the two other ostensible ASR matches.

I conclude that:

- ASR 31 should be matched to catalog source J032911.67+311832.3 (and c2d source J032911.6+311832).
- ASR 32 is likely spurious and cannot be tied to any source detected in these other surveys to date, so I have removed it from the catalog. It may be that it should be matched to VLA 27 (others have made this match in the literature), but I have not enforced this match here.
- ASR 33 should be matched to MBO 171 and catalog source J032911.51+311828.6 (and c2d source J032911.5+311828).
- VLA 27 should match to catalog source J032911.24+311831.8 (and c2d source J032911.3+311831).
- ASR 31 should *not* be matched to MBO 161 (which is different than what is found in the literature). MBO 161 is matched to the catalog source J032912.13+311832.1 (and c2d source J032912.1+311832).

Another more minor ASR issue pertains to ASR 50, which is clearly in the original ASR image, and can be seen in 2MASS  $6\times$ , but is on a diffraction spike and resembles an artifact enough that it was probably dropped by the 2MASS pipeline as a result. There is an IRAC source within  $2.5''$  of the Aspin source (J032855.08+311416.4) that is most likely the match.

There are several remaining ASR sources for which 2MASS-classic counterparts cannot easily be found. ASR 5=J032904.53+311554.6, oddly, can be seen clearly in the DSS images, but by 2MASS  $JHK_s$  is extremely faint, and is lost in the glare of a nearby very bright source by Spitzer and WISE bands. I have retained it with the original ASR coordinates, since it can be clearly identified in the DSS and is well-matched to that position. ASR 20, 21, and 22 are all in a region that becomes bright with extended emission at Spitzer bands. ASR 20 does not have a clear match in 2MASS, but ASR 21 is faintly seen in the 2MASS  $K_s$  images, and may be extended. There are two sources from the  $6\times$  catalog (both of which may be part of the extended emission); ASR 21 is matched to the closer one. At  $3.6\ \mu\text{m}$ , ASR 22 is close to a blob of extended emission which may have concentrated knots within it; I have tied it to 2MASS J03290842+3115284 as the brightest source in the blob. ASR 71 is entirely within another region that, at  $3.6\ \mu\text{m}$ , is a blob of extended emission that may have concentrated knots within it. The brightest portion of the blob is already identified with ASR 49. It is unclear to what object at 2MASS bands ASR 71 should be matched, so I have retained it with the original ASR coordinates. ASR 119 is almost exactly in between two 2MASS sources that are very far away, one  $\sim 15''$  north, and one  $\sim 18''$  south. Lacking a compelling reason to look this far away for a counterpart and make the association with one or the other, I have left ASR 119 alone at its original location despite

having no counterpart. ASR 133 is similarly more or less in between two sources  $\sim 15''$  north and south of this position; similarly, I have left it alone. ASR 134 does not have a 2MASS-classic counterpart, but can be seen faintly on the  $K_s$  images, and it is bright by IRAC bands. ASR 50, 75, 94, and 98 all have IRAC matches within  $5''$ , even if counterparts are not apparent in the 2MASS-classic images.

### 3.3. LAL Catalog

Lada et al. (1996; LAL) also imaged NGC 1333 in  $JHK$ , over a region 4 times larger in area than ASR. Their entire catalog was not included in the original article, but was obtained via private communication from C. Lada. Objects LAL thought were young were also indicated in this catalog. Numbers from this catalog are preceded by “LAL.”  $JHK$  measurements from this catalog were not retained because 2MASS measurements were available for all of the retained LAL sources. As for ASR above, the ensemble of  $JHK$  photometry is reasonably well-matched to 2MASS (better than  $\sim 0.2$  mag), though there can be excursions for individual sources.

The observations presented in LAL were assembled from many smaller pointings in the NIR, though unlike ASR, a single pointing field of view for LAL was relatively large, and at  $\sim 5'$ , comparable to a single IRAC field of view. LAL report that their astrometry is based on 5 stars from the Hubble Space Telescope Guide Stars Catalog, and that their estimated positional uncertainty is  $0.5 \pm 0.2''$ . However, based on my comparison to 2MASS, I strongly suspect that there are several astrometric problems. In most cases, the sources are recoverable; for 86% of the sources, I can find a counterpart in 2MASS within  $3''$ . However, in some cases, I could not find a counterpart. In one region, I noticed a pattern offset which I now describe.

As seen in Figure 4, I strongly suspect, based on relative positions, that several sources in the LAL catalog should all move  $54''$  to the southeast. These sources all appear, perhaps, to be in the corner of a component tile of the LAL final mosaic, and evidently the astrometry in this region (tile corner?) was not well-constrained. The consistency in the pattern of sources on the sky is compelling evidence that the sources should be shifted; a direct comparison of the  $JHK_s$  values on a source-by-source basis often supports this assertion.

The repercussions are that:

- LAL 81 moves from having no match to 2MASS J03284764+3124061, which is nominally close to the original LAL 93 position. There is a faint source that appears only at  $3.6 \mu\text{m}$  under this originally-reported location of LAL 81, but the SED for that object is different from the LAL-reported  $JHK$  by many orders of magnitude, which was what originally led me to suspect that this was not the correct match. The LAL-reported  $JHK$  matches the 2MASS photometry reasonably well, being offset by 0.16 mag, 0.12, and 0.02 mag in  $J$ ,  $H$ , and  $K$ , respectively.
- LAL 85 moves from having no match to 2MASS J03284947+3125066. The photometry matches between these sources to 0.14 mag in  $J$ , 0.08 mag in  $H$ , and 0.01 mag in  $K$ .
- LAL 87 moves from having no match to 2MASS J03284971+3124534. The magnitudes match to 0.24, 0.10, and 0.06 mag in  $J$ ,  $H$ , and  $K$ .
- LAL 93 moves from 2MASS J03284764+3124061 (where it is offset from a bright star, and to which it has sometimes been tied in the literature) to 2MASS J03285097+3123479. The photometry matches

Table 4. Summary of LAL Modifications

LAL number	notes
27	nothing at this location and nearby sources already have matches; removed
35	nothing at this location and nearby sources already have matches; removed
39	nothing at this location and nearby sources already have matches; nearby bright source 2M J03283695+3123121 does not have match, but that is nearly an arcminute away and thus is an unlikely match; removed
44	nothing at this location and nearby source already has match; removed
47	nothing at this location, no nearby sources; removed
50	nothing at this location, and nearby sources already have matches; removed
52	this one, plus 59 and 64, make an arc of 3 sources; this shape is not matched to anything in the 2MASS or IRAC images, and nearby sources have matches; removed
59	see 52; removed
61	near a faint (at 3.6 $\mu\text{m}$ ) extended source in [3.4] (WISE) and [3.6], but not in 2MASS; morphology suggests jet; removed
64	see 52; removed
71	nothing at this location, no nearby sources; removed
81	one of the “54” to the southeast”, now 2MASS J03284764+3124061
84	nothing at this location, no nearby sources; removed
85	one of the “54” to the southeast”, now 2MASS J03284947+3125066
87	one of the “54” to the southeast”, now 2MASS J03284971+3124534
93	81 moved close to original position of this one as part of the “54” to the southeast”; this one moves from 2MASS J03284764+3124061 (where it is offset from a bright star, and to which it has sometimes been tied in the literature) to 2MASS J03285097+3123479
98	one of the “54” to the southeast”, now identical to LAL 120, matched to 2MASS J03285216+3122453
108	one of the “54” to the southeast”, now identical to LAL 139, matched to 2MASS J03285521+3125223
116	nothing at this location, and nearby sources already have matches; removed
118	nothing at this location, and several nearby sources all already have matches; removed
120	could justifiably move to match LAL 147, but left it matched to 2MASS J03285216+3122453 (see “54” to the southeast” discussion); no net changes
122	nothing at this location, and nearby sources already have matches; removed
124	one of the “54” to the southeast”, now 2MASS J03285630+3124432
130	nothing at this location, nor at a place offset in the same direction and size as the “54” to the southeast”; removed
134	one of the “54” to the southeast”, now identical to LAL 165, matched to 2MASS J03285842+3122567
135	nothing at this location, and nearby sources already have matches; removed
149	in original LAL catalog, within 0.6” of 148, though photometry is different; retained 148 (and removed this one)
174	nothing at this location, and nearby sources already have matches; removed
242	in original LAL catalog, within 0.13” of 241, and only band available ( $K$ ) matches to 0.02 mag; retained 241 (and removed this one)
244	nothing at this specific location, but on edge of complex extended region with many sources; unclear what should match it; removed
259	in original LAL catalog, 258 and 259 have identical positions and matching photometry; retained 258 (and removed this one)
271	270-271 are 1.03” apart, and are likely two components of an object that is multiple in the 2MASS catalog, but both of the LAL sources are much closer to 2M J03291433+3114441 than 2M J03291409+3114423; combined and 270 retained
281	nothing at this location and nearby source already has match; removed
284	nothing at this location and nearby sources already have matches; removed
289	in original LAL catalog, 289 and 290 have identical positions though different photometry (one is missing $J, H$ ; $K$ matches to 0.02 mag) identical to 290; retained 290 (and removed this one)
291	nothing at this location and nearby sources already have matches; removed
338	nothing at this location and nearby source already has match; removed
352	nothing at this location and nearby sources already have matches; removed

this latter source, 2MASS J03285097+3123479, better (0.16, 0.04, 0.01 in  $JHK$ ) than the former source, 2MASS J03284764+3124061 (0.17, 0.22, 0.41 in  $JHK$ ).

- There is nothing at the nominal location of LAL 98. If I apply the offset in the same direction as the other sources above, then LAL 98 is identical to LAL 120, which is already matched to 2MASS J03285216+3122453. However, the LAL reported magnitudes for 98 and 120 differ by  $\sim 1$  mag at each band. On the other hand, 2MASS J03285216+3122453 has measured  $JHK_s$  within 0.19 mag at all bands to LAL98, not LAL 120. Given this evidence, I have opted to combine LAL 98 and 120 into 2MASS J03285216+3122453.
- LAL 108 I strongly suspect to be identical with LAL 139, which is matched to 2MASS J03285521+3125223. LAL 139 itself does not seem to be offset; this must be the boundary between tiles in the original observation. The reported photometry for LAL 108 and 139 are nearly identical (matching at better than 0.08 mag, all bands).
- I investigated whether it was reasonable to shift LAL 120 to become identical to LAL 147, because such a shift would be in the same direction. However, it is right on top of 2MASS J03285216+3122453, so I left it at that location. Confusingly, LAL 120 has photometry that is much different at  $J$  from LAL 147, but matches to within 0.13 mag at  $K$ ; LAL 120 and 2MASS J03285216+3122453 have differing  $JHK$  magnitudes by  $>4$  mag. I have made no net changes as a result of this consideration.
- LAL 124 moves from having no match to 2MASS J03285630+3124432. LAL does not report  $J$  or  $H$ , but these sources match each other at  $K$  to 0.13 mag.
- LAL 134 I strongly suspect to be identical to LAL 165, which is well-matched to 2MASS J03285842+3122567. The two LAL sources have no  $J$  or  $H$ , but match each other in  $K$  to 0.09 mag; they match the 2MASS  $K_s$  to within 0.06 mags.

Just based on patterns, this seems to be the end of the sources I can (or need to) match in this region. Beyond the apparent corner of the tile, LAL 147, 166, 164, and 163 are all well-matched to 2MASS counterparts. I cannot recover LAL 130; there is nothing at that location in 2MASS or IRAC, nor at a place offset in the same direction and size as the offsets above.

These coordinate uncertainties have some more minor repercussions in the rest of the catalog – there are things I strongly suspect are duplicates, as well as sources without counterparts that I can find. As for ASR, even given intrinsic variability, the LAL observations are not so deep as to be likely to reach sources undetected in the 2MASS, 2MASS 6 $\times$ , IRAC, or even the WISE observations. If there was no source in the 2MASS, IRAC, or WISE images at the LAL position or within a ‘reasonable’ distance, I often dropped the LAL source; see Table 4.

In the end, I made changes to 37 LAL sources. Sources that are duplicates are explicitly indicated in the last column in Table 2 via notation like ‘==LAL44’ for LAL 38.

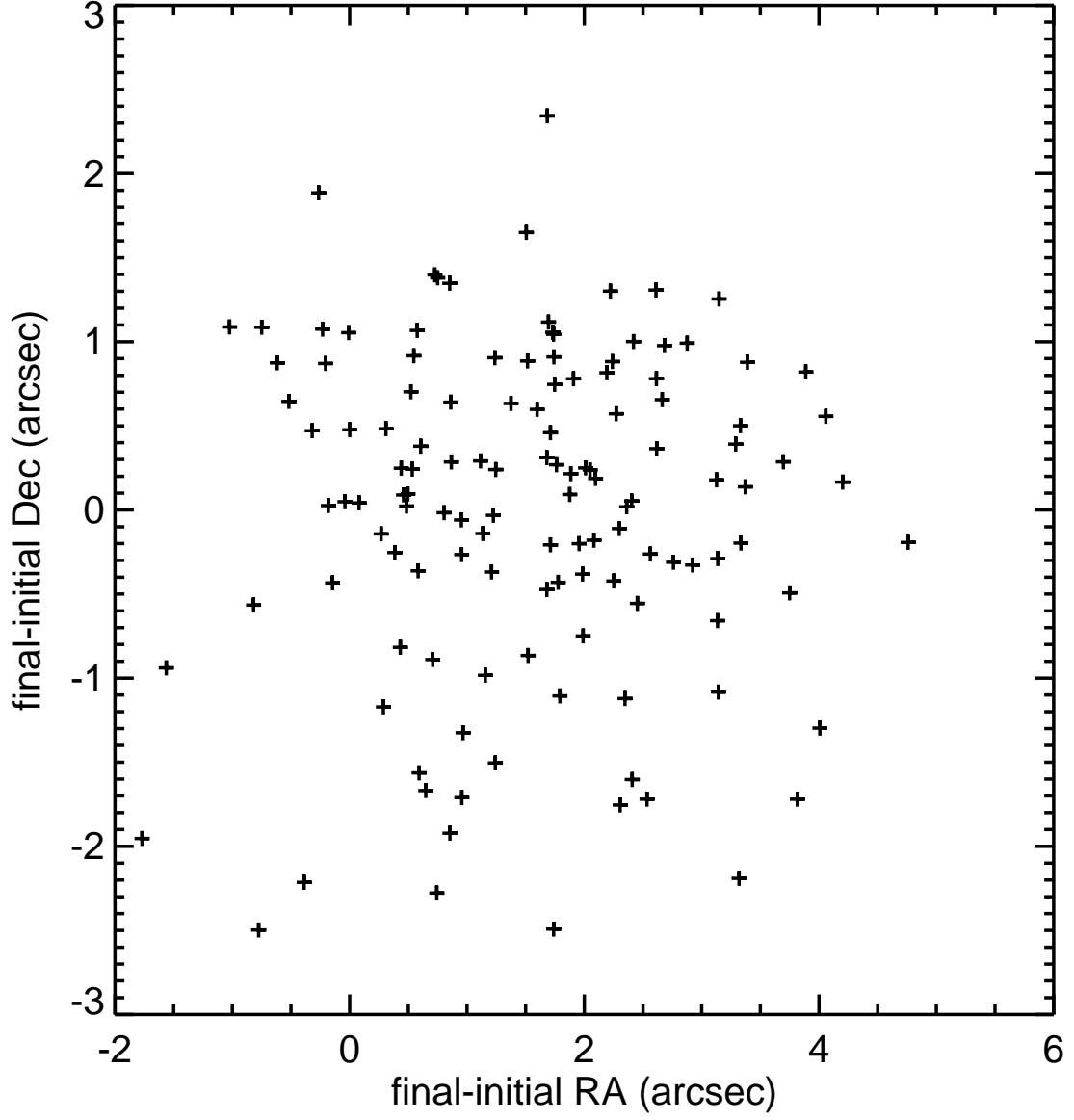


Fig. 2.— Comparison of offsets in RA and Dec (both in arcseconds) between the final, adopted position (matching 2MASS) and the initial position precessed from the B1950 coordinates provided in ASR. The offsets are symmetric in declination, and reveal a systematic offset in RA. The median net offset is  $\sim 1.7''$ .



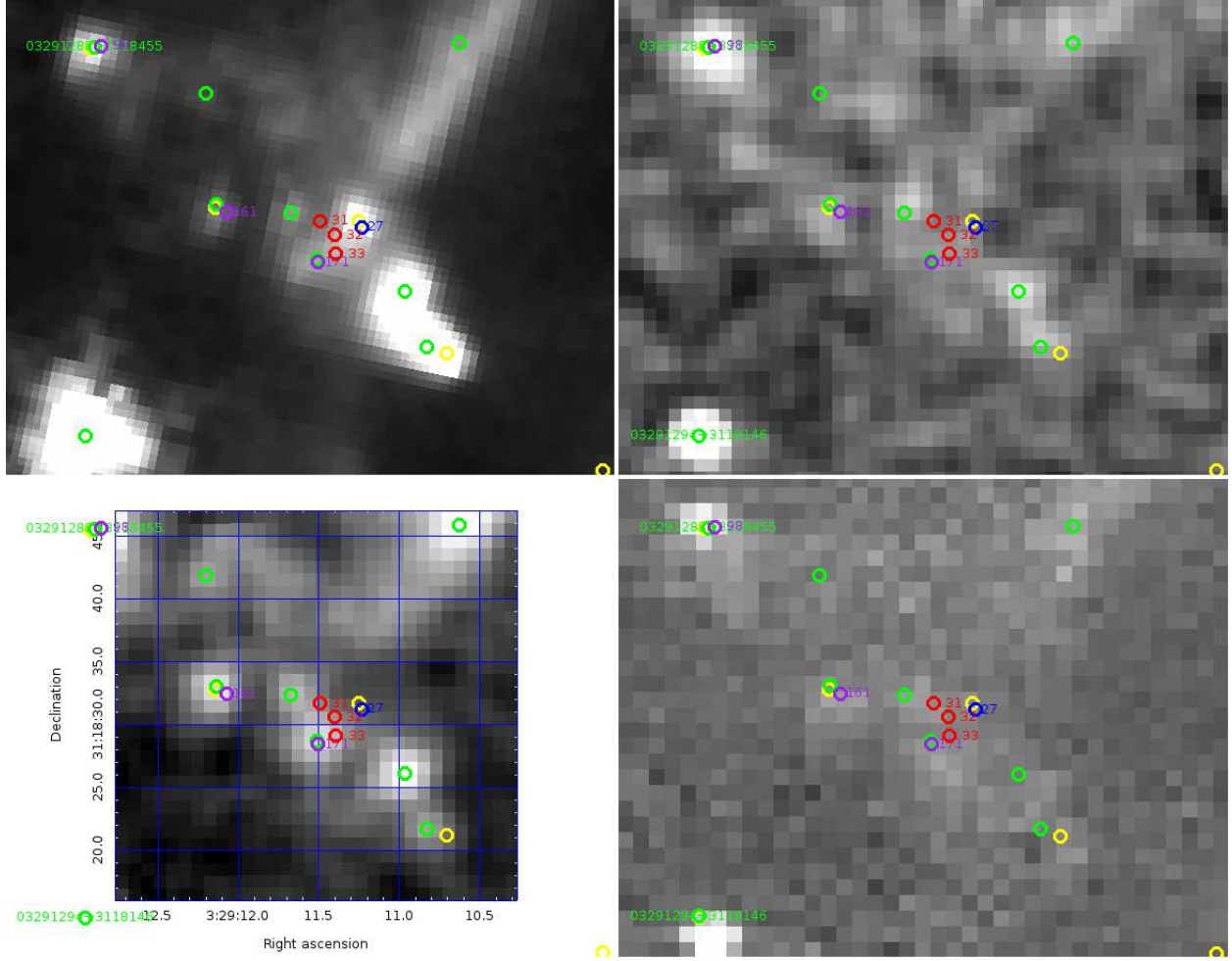
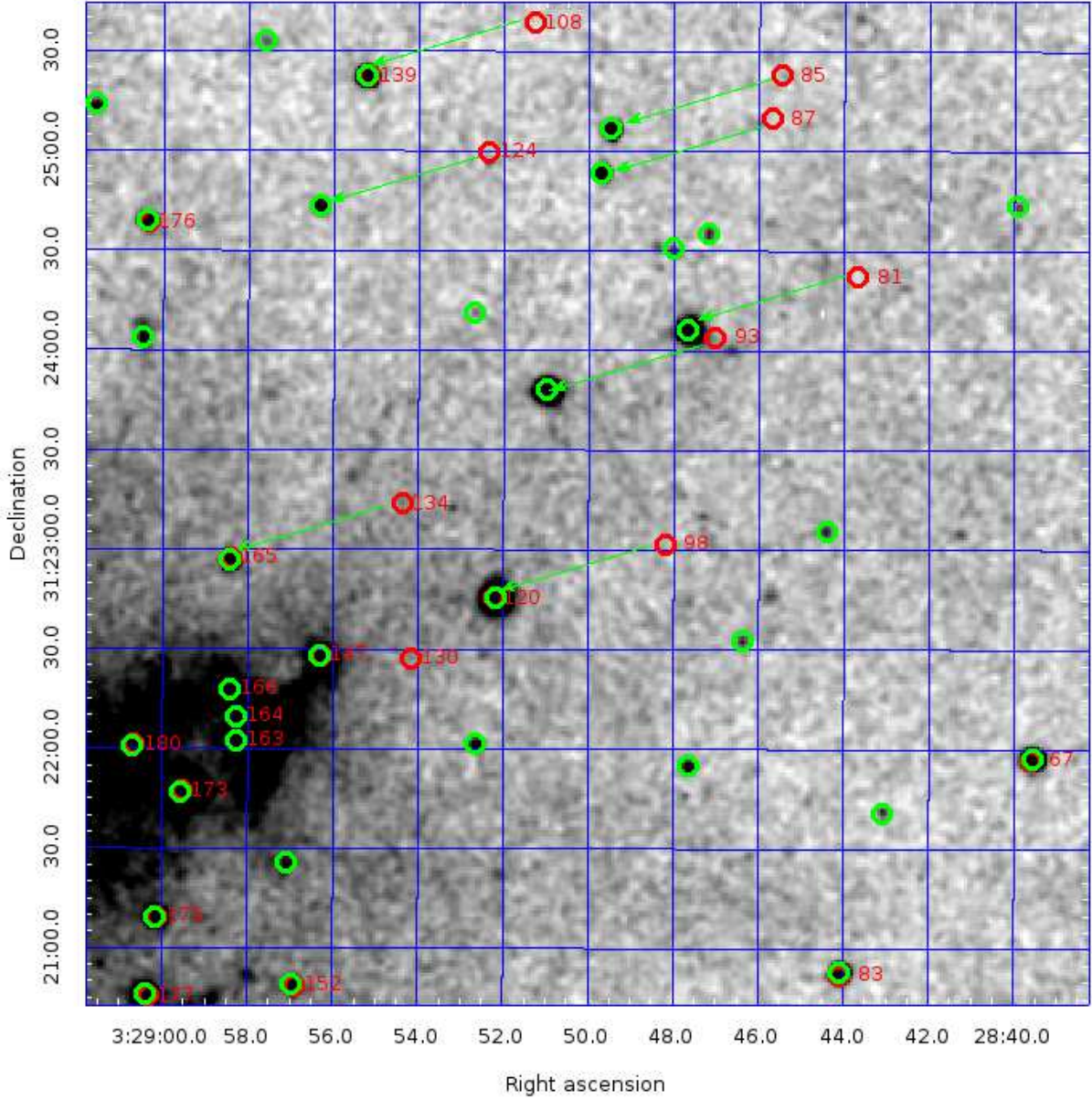


Fig. 3.— Resolving the ASR 31-32-33 confusion. The images shown are (left to right, top to bottom): IRAC-1 ( $3.6 \mu\text{m}$ ) from the cryogenic mission, 2MASS  $K_s$  from the classic survey, 2MASS  $K_s$  from the  $6\times$  deeper survey, and then the original ASR image from 1994, unfortunately with subsequently updated astrometry. The sky is shown on the same scale in each image; RA and Dec coordinates are indicated in the lower left for reference. Sources are overlaid as follows: red circles are ASR 31/32/33 (ordered top to bottom), with the original coordinates as reported and simply precessed to J2000 from 1950; green circles: 2MASS sources, where ones that have the long names indicated are from 2MASS classic and ones without numbers are from 2MASS  $6\times$ ; yellow circles: full cryogenic catalog from Gutermuth et al. (2008); purple circles: Wilking et al. (2004) MBO catalog; blue circle: Rodriguez et al. (1999) catalog (VLA 27). See text for detailed discussion of which sources are identified with each other in the ASR 31-32-33 region.





### 3.4. X-ray catalogs

Using a deep Roentgen Satellite (ROSAT) High Resolution Imager (HRI) observation, Preibisch (1997) detected 20 X-ray sources, 16 of which were taken to be likely cluster members. Getman et al. (2002) report on a Chandra observation of this region, detecting 127 sources, 95 of which were identified with known cluster members. I retained only the source numbers (not the fluxes) reported in Preibisch (1997). Getman et al. (2002) find matches to all Preibisch sources except sources 1 and 2. I was able to locate a counterpart at multiple other bands for source 1, and no counterpart at any band for source 2. I suspect that source 2 from Preibisch (1997) may be spurious (or extragalactic), and I have removed it from my catalog.

Some of the sources from Getman et al. (2002) were specifically tagged as variable sources (in X-rays). No new YSO candidates were identified in their catalog, though it is quite likely many of the newly identified sources are members. They reported their entire catalog, which I absorbed into our database for the cross-IDs, though I did not retain the X-ray fluxes or luminosities in favor of our own re-reduction of the Chandra data (R14).

I found two inconsistencies in the Getman et al. (2002) crossmatching between this catalog and LAL and ASR, which I corrected in the catalog. Many of the cross-matches in Getman et al. (2002) are correct. However, LAL 79 is much closer to ASR 126 than 127, which is not what is reported in Getman et al. (2002). LAL 93 moved as per the discussion in Section 3.3 above; therefore LAL 93 should not actually be matched to Getman source 15. Getman source 15 should be matched to 2MASS J03270047+3137259.

Preibisch (2003) reports on an XMM-Newton observation of NGC 1333. It covers a wider area than the Chandra data from Getman et al., but most of the sources identified in this paper are also identified in Getman et al.. There are 7 sources identified based solely on the XMM data. Preibisch finds counterparts in the optical or IR for 5 of the sources. Given the coordinates in Preibisch, 2 of the objects are within 1 arcsec of objects in our database, and the rest are within 2 arcsecond of objects in our database, including the two claimed not to have a counterpart. I have made these associations in our catalog. Names from this study are incorporated into ours, including for the two objects previously claimed to not have counterparts.

The XMM-Newton Serendipitous Source Catalogue 3XMM-DR4 was released in 2013, and consists of source detections over most of the XMM data taken as of 2012 December, which includes the NGC 1333 observations from Preibisch (2003). I find 46 matches to sources in this region, only 10 of which do not already have a Chandra detection in our reprocessed data. Names for these sources are included in the ‘other names’ field, preceded by ‘XMM.’

There are no sources for which I now have an X-ray measurement (from YSOVAR’s reprocessing or from 3XMM-DR4) that did not already appear in Preibisch (1997), Getman et al. (2002), Preibisch (2003), or Winston et al. (2009, 2010), described below.

### 3.5. Rodríguez et al. (VLA) Catalog

Rodríguez et al. (1999) report on Very Large Array (VLA) observations at 6 and 3.6 cm of an  $8' \times 8'$  region centered on the HH 7-11 region. The nomenclature for this catalog as established in the paper is not “RAC” (as one might assume) but rather “VLA.” Most of the objects were identified with a YSO counterpart in this paper; some were identified as variable. I retained nomenclature and positions from this catalog, because they are still in wide usage today.

Finding counterparts between 3.6 and 6 cm objects and near- to mid-IR sources can be quite complicated, as there is no assurance that the emitting source is the same. I have thus retained some VLA sources without shorter wavelength counterparts, and I have shifted some sources to match shorter wavelength counterparts. As mentioned above in Sec. 3.2, VLA 27 has been tied (including by Rebull et al. 2003) to one of the ASR sources. I have now tied VLA 27 to that source (R15-NGC1333 J032911.24+311831.8=c2d J032911.3+311831), but not any of the ASR sources. Also, I have tied VLA 43 to ASR 114 (=2MASS J03290149+3120208 = R15-NGC1333 J032901.53+312020.6) and VLA 42 to 2MASS J03290116+3120244 (R15-NGC1333 J032901.16+312024.4). Based on long-wavelength information from Herschel (in the context of the Sadavoy et al. (2014) data incorporation below), I have associated VLA 28 with R15-NGC1333 J032912.05+311301.4.

### 3.6. Other NIR and MIR data

Rebull et al. (2003) reported on ground-based MIR data, and has largely been superceded by Spitzer data. However, the source cross-identifications and the source multiplicities found there are useful, and have been retained in the catalog. For example, SVS 12 may be extended, and SVS 16 and ASR 107 are both multiple sources.

Aspin (2003) reported on NIR spectroscopy, including spectral types which I have included in our database. Some types are not particularly precise (“early K”) but for many stars, this is all that is available.

Wiling et al. (2004) revisited the cluster with *JHK* over a  $\sim 11' \times 11'$  region to  $K \sim 16$ . They also searched for brown dwarfs, reporting on spectral types for their 25 candidates. The catalog abbreviation for this study, as established in their paper, is “MBO,” standing for Mount Bigelow Observatory. In the published Wilking et al. (2004) catalog, there are two MBO 221s. Both have (separate) IRAC counterparts, only one has a 2MASS counterpart, and it is faint at that (it comes from the 2MASS  $6\times$  catalog). In consultation with B. Wilking (priv. comm.), the source on the west was manually added. MBO 222 is the highest number in the published catalog, so this source is now assigned to be MBO 223. Thus, MBO 221 is now matched to R15-NGC1333 J032847.19+311845.9 and MBO 223 is now matched to R15-NGC1333 J032847.27+312310.0. For completeness, and for the benefit of future users of the catalog, I note that there are also several MBO numbers missing entirely.

### 3.7. Cryo-era Spitzer Catalogs

As noted above, the first Spitzer data for the NGC 1333 region were taken were taken cooperatively between the GTOs and the c2d Legacy team. The Legacy team mapped the entire Perseus molecular cloud complex, including NGC 1333; Jørgensen et al. (2006) reported on the IRAC data, and Rebull et al. (2007) reported on the MIPS data. Notably, the c2d data delivery included “bandfilled flux densities,” meaning that if the source was not detected, an aperture was laid down at the location of the source to obtain an upper limit. I included flux densities and upper limits from the c2d catalog where measurements were not already present in our reprocessing of the cryogenic-era data. (This was the case where there was a low signal-to-noise ratio detection, or a limit.) Note that our cryo reprocessing does not include 70 or 160  $\mu\text{m}$  fluxes; where relevant (rarely, in this very crowded region), I obtained 70  $\mu\text{m}$  measurements from the c2d catalog. The c2d project also identified YSO candidates using a multiband color selection; I retained in the database an indication of this status.

Gutermuth et al. (2008) reported on the Spitzer data specifically for NGC 1333 (as opposed to the entire Perseus cloud as c2d did). Only the candidate young stars were reported in Gutermuth et al. (2008). Since these cryogenic observations were reprocessed for YSOVAR using the same approach as Gutermuth et al. (2008), I only retained identifications of YSO candidates from Gutermuth et al. (2008).

Jørgensen et al. (2007) reported on submillimeter data combined with the Spitzer c2d data. They included a list of embedded YSOs in Perseus. These cross-IDs are included in ‘other names’ as J07-xx.

Arnold et al. (2012) included data from a Spitzer Infrared Spectrograph (IRS) survey of objects in NGC 1333. From this article, I have retained cross-IDs, SED classes (for comparison to ours), and identifications of YSOs (and non-members).

### 3.8. Hatchell et al. (2007) catalog

Hatchell et al. (2007) observed Perseus at 450 and 850  $\mu\text{m}$ . For some sources, they compiled SEDs and made cross-identifications. I included the cross-identifications in our catalog, as well as made new associations between these sources and the short-wavelength counterparts.

Source number 49 is listed as a match to an IRAS source, as well as the 2MASS source J03283609+3113346. Looking at the 2MASS and WISE images, this is an extended source, and that 2MASS source is one of three near the core of it. The 2MASS point source closest to the core of the extended source is J03283681+3113326; the 2MASS extended source catalog source to which it should be matched is J03283630+3113329.

Source number 70 does not have a short-wavelength counterpart listed in Hatchell et al.. There is nothing obvious there in the 2MASS images, and only faintly is there a source in the WISE images. There is a source from Spitzer/IRAC at J032914.96+312031.7, which is 4.4'' away from the position given for source 70, within the range of positional uncertainties given by the Hatchell et al. cross-references to 2MASS. Hatchell et al. place this source as a SED Class 0. I have very few points between 2 and 25  $\mu\text{m}$  delineating J032914.96+312031.7 – I only have four IRAC bands, and it is falling, not rising. I place it in SED Class II. The Hatchell et al. source 70 may not be a good match to the source at R15-NGC1333 J032914.96+312031.7, but it is the best thing I can match to it at this time. A note about this uncertainty is included in the last column of Table 2.

### 3.9. Brown Dwarf Searches

There are several studies seeking brown dwarfs in this region. Oasa et al. (2008) is one of these brown dwarf searches, using  $JHK_s$  over a  $\sim 5' \times 5'$  region with a spatial resolution of  $\sim 1.5''$ . Their YSO candidate identifications were retained in my catalog. Matches were found within 2'' for 85% of their sources. The rest of the sources were retained in the catalog with the positions from Oasa et al., except for one, which I now describe.

I find that one of the sources reported upon in Oasa et al. (2008) is subject to source confusion issues. It is in the region associated with IRAS-2b. The source in question was assigned a position of 03:28:57.09 +31:14:21.4 by Oasa et al., who reported no other sources within 40'' of this location. In the IRAC images, two sources can be seen, J032857.37+311415.7 and J032857.20+311419.1, both of which are quite far away from the source from Oasa et al. (7.4'' and 3.4'', respectively). They have very different flux ratios – the more

northern one dominates at  $\sim 3 \mu\text{m}$  (and is closer to the original Oasa position), but by  $\sim 8 \mu\text{m}$ , the southern one dominates. By inspection of the available images (2MASS, IRAC, and WISE), I conclude that this Oasa et al. source should not be retained as a separate source, but instead be tied to one of the other two seen in the IR. Based on proximity, I have tentatively tied this source to R15-NGC1333 J032857.20+311419.1. However, Oasa et al. report that this source is a Class III, and not a YSO. This is not consistent with the rest of the information I have on J032857.20+311419.1, but there aren't any closer sources bright in the NIR. This is the best association I can make. A note about this uncertainty is included in the last column of Table 2.

Greissl et al. (2007) used the Hubble Space Telescope (HST) NICMOS instrument to search for brown dwarfs in NGC 1333. They reported  $JH$  photometry and spectral types from low-resolution grism spectroscopy for their objects of interest, both of which I included in the database. Greissl et al. source S1-6 is very close to other sources in the region; Figure 5 shows the region in  $3.6 \mu\text{m}$ . There is an ‘appendage’ off of a brighter source; that ‘appendage’ is what I matched to source S1-6=J032857.41+311536.9. The net SED for this source is largely optical/NIR data, from Greissl et al.; it is hard to apportion IR data correctly for this source from telescopes other than HST. I considered that this source may be incorrectly matched to an IR source, and that the optical points may correspond to a different source. The nearest one I can identify is the brighter source in Fig. 5 (J032857.15+311534.6), which is  $\sim 4''$  away; this is significantly larger than the IR positional uncertainties, or that of Greissl et al., so this is unlikely. Moreover, since they were working with HST data, it is likely that their astrometry is correct. I have left it in the catalog as-is. There are notes about this in the last column of Table 2 for S1-6=J032857.41+311536.9.

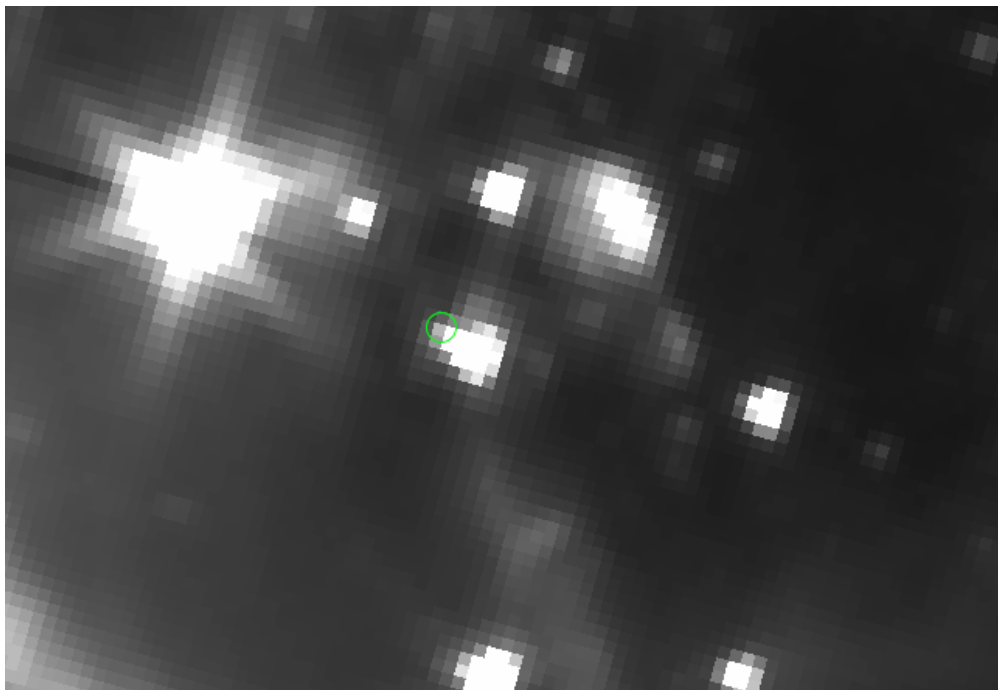


Fig. 5.— The region of sky in  $3.6 \mu\text{m}$  immediately around J032857.41+311536.9=Greissl S1-6, circled. North is up. The distinctive triangularly shaped IRAC point spread function of the point sources can be seen in several point sources in this field of view, but one has an ‘appendage’ extending to the left; the ‘appendage’ is matched to Greissl S1-6.



Scholz et al. (2009), Scholz et al. (2012a), and Scholz et al. (2012b) are all part of the SONYC (Substellar Objects in Nearby Young Clusters) survey searching for brown dwarfs. These investigators obtained very deep imaging in four bands ( $i'$ ,  $z'$ ,  $J$ ,  $K$ ) but only report photometry for their brown dwarfs, candidates, and rejects, which I incorporate into the catalog. I also included the spectral types and membership as reported in the SONYC papers, both for stellar and substellar objects. Scholz et al. (2012b) report a new very low mass member, which becomes S-45. However, it is within 0.2 arcsec of Sp105, a source they list as a previously confirmed member in Scholz et al. (2012a). I take these two to be referencing the same object. They have essentially identical NIR magnitudes and spectral types in the papers.

Itoh et al. (2010) is another search for brown dwarfs; spectral types and cross-identifications from this paper were incorporated into the database.

### 3.10. Winston catalog

Winston et al. (2009, 2010) report on a combined Spitzer and Chandra investigation of NGC 1333. Follow-up spectra were also included. A full X-ray catalog (of every source detected) was included in Winston et al. (2010). Spectral types and X-ray identifications were incorporated from this paper into our database, as well as whether Winston et al. identified the object as a YSO. The X-ray measurements from Winston et al. were superseded by our own re-reduction (R14).

The source numbers presented in Winston et al. (2009) refer to source numbers tied to coordinates (RA/Dec) in Winston et al. (2010), but there is some potential confusion as to whether the numbers are sequential IR or X-ray numbers, where the former come from Gutermuth et al. (2008), and the latter are assigned in the Winston work. The numbers in Winston et al. (2009) are not specified as being IR or X-ray numbers. Winston et al. (2010) reports both X-ray and IR numbers, but not all of them appear in Gutermuth et al. (2008). Via personal communication with E. Winston, I have verified that the numbers in Winston et al. (2009) are meant to be IR numbers. The IR numbers in Winston et al. (2009) and (2010) that are greater than 137 do not appear in Gutermuth et al. (2008). These are ones that Winston et al. assigned (not Gutermuth et al.) and are defined via positions (RA/Dec) in Winston et al. (2009). Thus, while Gutermuth et al. (2008) only assigns numbers to YSO candidates reported there, there can be a “Gutermuth08 number” assigned by Winston et al. (2009, 2010). In other words, Winston et al. (2009) reports those as YSO candidates selected from the IR, but Gutermuth et al. (2008) does not report those objects as YSO candidates. All of this has been resolved in my catalog.

### 3.11. Sadavoy et al. Herschel identifications

Sadavoy et al. (2014) report on candidate Class 0 objects in NGC 1333 using Herschel data. Because the shortest Herschel wavelength, 70  $\mu\text{m}$ , overlaps with the longest Spitzer wavelength I have here (also 70  $\mu\text{m}$ ), I wished to link the Herschel observations with the rest of my catalog. However, this is another example of the difficulty in finding clear correspondences between sources at shorter and longer wavelengths. In many cases, the corresponding source at 3.6 and/or 4.6  $\mu\text{m}$  is blended with nebular emission and may not represent flux solely from the source seen at long wavelengths. Some of the sources in Sadavoy et al. in this region are noted as not having Spitzer counterparts, but I have been able to find counterparts for most of them. All of these cross-identifications appear as part of the ‘other names’ in Table 2; where the association is uncertain, a question mark is added to the cross-identification. No flux densities are reported

in Sadavoy et al. (2014), so these points cannot be added to tables or SEDs. My assembled SEDs for these objects appear in Figure 6. Notes on specific sources follow.

Counterparts to Sadavoy et al. sources West162 and West33 are likely to have contamination from extended emission at 3.6 and 4.5  $\mu\text{m}$ . Moreover, for West33, the short-wavelength counterpart has a falling SED, but by 24  $\mu\text{m}$ , is very bright and rising fast. The source detected at 3-8  $\mu\text{m}$  may not be the same object seen at  $>24$   $\mu\text{m}$  in this case.

West19 is very complicated because it is in a very bright and confused region. There does seem to be an object (J032901.91+311541.4) at IRAC wavelengths close to that location; at 7.5'', it is among the largest positional offsets for all of the matches to the Sadavoy sources, but consistent with the positional uncertainties. At 24  $\mu\text{m}$ , the object at this location is very unfortunately placed in the MIPS image so as to be largely obscured by the Airy ring of a nearby extremely bright source. (Note that I have not tied West19 to the very bright 24  $\mu\text{m}$  source, since it is too far away!) This association is the best I can do at this time, despite the relatively large positional offset.

West40 and West87 are both in the same neighborhood as West19, though they are both farther from the very bright sources and extended emission that plague West19. By contrast, West40's match to J032904.07+311446.5 seems relatively straightforward, though the  $J$ -[24] SED is somewhat unusually shaped. West87's match at J032906.45+311534.4 does not have a [24] measurement because of the halo around the bright 24  $\mu\text{m}$  source. It has another strange SED, and the object responsible for the emission at the shortest wavelengths may very well not be the same as for the  $>20$   $\mu\text{m}$  wavelengths. Because it is missing SED points longer than 20  $\mu\text{m}$ , and because of the 2-8  $\mu\text{m}$  slope, I place this object in SED Class II.

For West6, there is a clear correspondence of the 24  $\mu\text{m}$  source (at J032910.49+311330.7) and the Herschel source. However, no clear 8  $\mu\text{m}$  source can be seen on the IRAC image (and no 12  $\mu\text{m}$  source can be seen on the WISE image). There is a faint source at 4.5  $\mu\text{m}$ , and perhaps something on the 5.8  $\mu\text{m}$  image, but nothing clear on the 3.6  $\mu\text{m}$  image. It is not apparent if the source at 4.5 (and 5.8)  $\mu\text{m}$  is the same as the 24  $\mu\text{m}$  source and West6, but all measurements have been retained in this associated source.

West14 is in another very complicated region. There are several sources in close proximity as well as extended nebular emission. I have tied it to R15-NGC1333 J032911.24+311831.8 on the basis of image morphology at Spitzer bands, but note that there could be contributions from other adjacent sources at nearly all bands. (For example, WISE may blend the two closest sources.)

West13 is another object that is difficult to match to short wavelength sources. The 70  $\mu\text{m}$  source seen by Herschel has a clear counterpart in the 24  $\mu\text{m}$  MIPS images. However, there may be more than one source in the IRAC images at this location. Based on the images, I associate West13 with R15-NGC1333 J032912.05+311301.4 as the most likely match. Based on this association, I move VLA 28 from its nominal published location to coincide with this source as well. There is a nearby source seen in 2MASS 6 $\times$ , but only at  $K_s$ . This source, however, is too far away to be associated with R15-NGC1333 J032912.05+311301.4, so I do not associate them.

West30 and West23 are both additional cases of a clear match among the 70, 24, 8, and 5.8  $\mu\text{m}$  images, but the measurements at 3.6 and 4.5  $\mu\text{m}$  may be affected by extended emission.

West37 is another very complicated, bright region. There is clearly an object in the 3.6, 4.5, and 5.8  $\mu\text{m}$  images, but if I had only those short wavelengths, I might have called it a dust clump based on image morphology. There is a source faintly seen in the 24  $\mu\text{m}$  image, but is on a nearby source's Airy ring. I have provisionally tied West37 to R15-NGC1333 J032918.88+312313.0.

### 3.12. Foster APOGEE data

Foster et al. (2015) report on the velocity dispersion of young stars in NGC 1333 based on near-IR spectra obtained with the Apache Point Observatory Galactic Evolution Experiment (APOGEE; Zasowsky et al. 2013). Target selection for that project included considerations based on the variability of objects in the YSOVAR data, so many of the objects in which we are interested in the context of YSOVAR also have measurements in Foster et al.. Through spectral fitting, Foster et al. derive several parameters for stars with sufficient signal-to-noise ratios in their APOGEE spectra, including  $T_{\text{eff}}$ ,  $\log g$ , and  $v \sin i$ . I included in our database cross-identifications, identifications of non-members,  $T_{\text{eff}}$ , and  $v \sin i$ , as well as other information from that paper.

### 3.13. YSOVAR data

Because I had originally undertaken this project in the context of analysis of our YSOVAR data, I have taken the liberty in this paper of including the means from those light curves (at 3.6 and 4.5  $\mu\text{m}$ ) in our database. As mentioned above, the NGC 1333 YSOVAR data will be described in detail in Rebull et al. (2015 in prep).

### 3.14. Two Very Bright, Confused Sources

In one of the brightest regions of NGC 1333, there are two sources very close to each other, at or near the spatial resolution of many of the surveys here. Early on, this pair of sources was identified as SVS 12=IRAS 6. However, in 2MASS, IRAC, and even in MIPS images, two sources can be distinguished by eye if the stretch is severe enough. Both of these objects have slightly different SEDs, certainly in part due to saturation and flux apportionment issues. The one tagged ASR 114=LAL 181=MBO 19=Gutermuth 27=c2d J032901.6+312021=R15-NGC1333 J032901.53+312020.6 made it onto most published lists as a YSO or candidate. For reasons probably due to saturation and/or spatial resolution and/or flux apportionment, the second did not; it is c2d J032901.2+312025=R15-NGC1333 J032901.16+312024.4. In the c2d catalog, one of these source gets all the MIPS 70  $\mu\text{m}$  flux density, though the other source probably should be allocated some of it. Many bands are listed as limits for these sources in various catalogs; they are probably meant as lower, not upper, limits, though they are often tagged upper limits. I have declared both of these sources to be literature-identified YSOs, and added a note about the 70  $\mu\text{m}$  flux apportionment to the last column in Table 2.

### 3.15. Jet Drivers

I consulted two papers to identify the sources most commonly thought of as driving outflows and/or jets in NGC 1333. Davis et al. (2008) used 2.122  $\mu\text{m}$  imaging to identify outflows; they identified 11 sources as having outflows of any sort, and their IDs were tied back to those from Jørgensen et al. (2007). Plunkett et al. (2013) reported on CARMA observations of outflows, identifying seven YSOs as the drivers of the outflows. Three sources are identified by both papers, resulting in 12 sources identified as driving jets or outflows. These sources, as well as some of their other names, are included in Table 5.



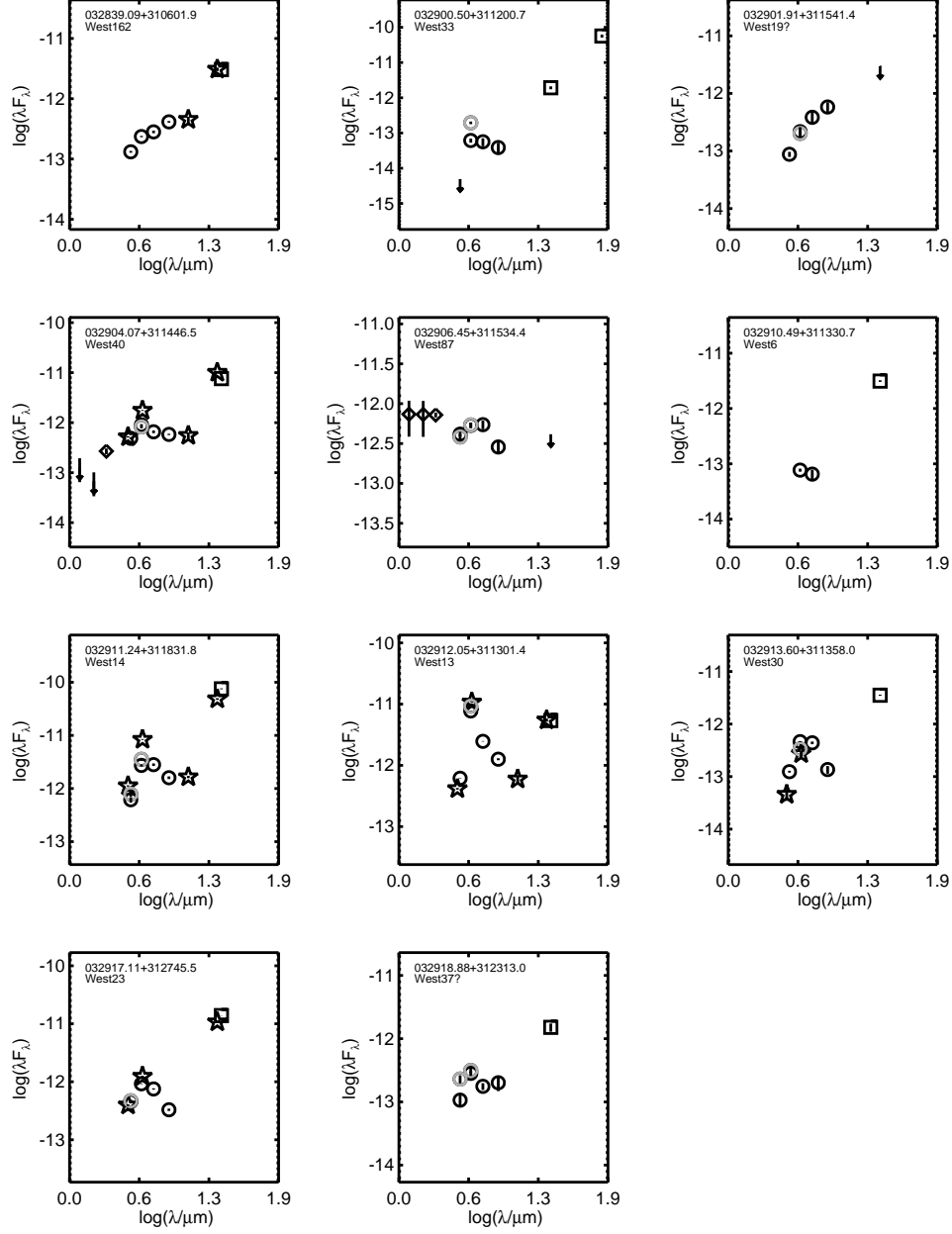


Fig. 6.— SEDs for the sources matched to the sources from Sadavoy et al. (2014). Plots are  $\log \lambda F_\lambda$  in cgs units ( $\text{erg s}^{-1} \text{cm}^{-2}$ ) against  $\log \lambda$  in microns. Diamonds= 2MASS, circles=IRAC (with grey circles being mean from YSOVAR campaign), squares=MIPS, stars=WISE, and downward arrows=upper limits. Small vertical lines in each point are the error bars. R15-NGC1333 catalog numbers appear in the plots. See the text for details of source matching, missing bands, etc.

Table 5. List of Jet/Outflow Driving Sources

Our ID (R15-NGC1333 J)	Reference(s)	Some of the Other IDs
032837.06+311331.0	Davis et al. (2008)	IRAS1, IRAS 03255+3103, J07-11
032845.30+310542.0	Davis et al. (2008)	IRAS 03256+3055, J07-14
032855.53+311436.3	Plunkett et al. (2013), Davis et al. (2008)	IRAS 2A, SK8, J07-15
032857.37+311415.7	Davis et al. (2008)	IRAS 2a, J07-16
032900.50+311200.7	Plunkett et al. (2013), Davis et al. (2008)	IRAS 4B1, SK1, Sadavoy2014-West33, J07-18
032903.39+311602.0	Davis et al. (2008)	J07-20
032903.75+311603.9	Plunkett et al. (2013)	SVS13, IRAS3, IRAS 03259+3105,
032904.07+311446.5	Plunkett et al. (2013)	IRAS5, Sadavoy2014-West40, SK14, J07-21
032910.49+311330.7	Plunkett et al. (2013), Davis et al. (2008)	IRAS4A, SK4, Sadavoy2014-West6,J07-22
032910.70+311820.9	Davis et al. (2008)	J07-23
032911.24+311831.8	Davis et al. (2008)	Sadavoy2014-West14,J07-24
032912.05+311301.4	Plunkett et al. (2013)	IRAS4B, SK3, Sadavoy2014-West13, J07-25

### 3.16. Studies That Are Not Included

There are several individual famous objects of particular interest in this region, some with many papers entirely of their own (e.g., the components that make up IRAS-2, IRAS-4, ...). I did not include additional data from these many individual projects in our catalog, but instead assume that users can match by cross-id to any few sources of interest (though I encourage image inspection via the IRSA tool FinderChart if nothing else). I also did not place an emphasis on matching to many surveys at wavelengths longer than about 50  $\mu\text{m}$ , because (a) there are relatively few true point sources at the longer bands, and (b) it is often hard to ensure a good match between the point sources from shorter wavelengths and those at the longer wavelengths. Some long-wavelength matches have been made in three cases – the VLA catalog is the source of some nomenclature still in wide use, and both the Hatchell et al. (2007) and Sadavoy et al. (2014) papers made specific efforts to find matches at the shorter wavelengths. A comprehensive set of matches to all of the sub-mm and mm sources in the literature is beyond the scope of this paper.

## 4. Obvious Gaps for Future Work

About half of the objects in this region that are identified as YSO candidates lack spectroscopy. This could be used not just for spectral classification, but for confirmation that they are young stars, or even just confirmation that they are not background galaxies (this is a very real concern; see, e.g., Rebull et al. 2010). Most of the spectral types found in the literature come from searches for brown dwarfs, so the set of objects for which there is spectral types is incomplete and highly skewed to late M and later. There are about 150 objects (out of  $\sim 300$  YSO candidates) that have some sort of estimate of spectral type in the literature, even a coarse one (e.g., “<M0” or “early K”). The  $T_{\text{eff}}$  values from Foster et al. (2015) can be used to constrain the spectral type, even though the  $T_{\text{eff}}$  values are much more uncertain for the hotter stars. However, there are only about a 20 objects for which there is a  $T_{\text{eff}}$  estimate but no published spectral type of any kind. Spectroscopy would help limit the non-member contamination and improve the inventory of members (and therefore knowledge of the mass function) in NGC 1333.

Multi-band broadband optical data could be very helpful over this entire region for delineating the Wien side of the SED. Even just  $r$  and  $i$  (or similar) would help show if the SEDs are, in fact, turning over for the YSO candidates. Moreover, short wavelengths such as  $u$ - and  $g$ -band data would be useful for constraining mass accretion rates in these young stars, but such observations would be quite challenging due to the high extinction towards most sources. High spatial resolution optical observations have been shown (e.g., Rebull et al. 2010) to be critical for distinguishing background star forming galaxies from nearby YSOs, since both kinds of objects overlap in IR color space with IR colors indicating star formation.

X-ray data can be very helpful for identifying young stars without disks. The existing Chandra and XMM data are limited in area, focused just on the heart of NGC 1333. X-ray data over a larger area could help identify additional less-embedded cluster members.

Some of the most famous YSOs in this region are those originally identified by the Infrared Array Satellite (IRAS) in 1983. However, most of these very bright sources have broken into pieces every time astronomers have looked with higher spatial resolution. Some sources are still clearly multiple and still have flux apportionment problems. Higher spatial resolution MIR and FIR observations will aid in correctly attributing source flux to the correct source. In some cases, the surface brightness is so high that higher spatial resolution observations could be very difficult.

## 5. Placement into SED Classes

After assembling all of the literature data, including the mean [3.6] and [4.5] measurement from our YSOVAR campaign, I constructed an SED for each object using all available data. Some objects have well-defined SEDs, with data from optical to 8, 24, or even 70  $\mu\text{m}$ ; others have far less-well-defined SEDs, for example, with only one point that is the mean of one channel’s light curve from our YSOVAR campaign.

As discussed in R14, I defined an internally consistent placement of the YSOVAR objects into SED classes as follows. In the spirit of Wilking et al. (2001), I define the near- to mid-IR (2 to 24  $\mu\text{m}$ ) slope of the SED,  $\alpha = d \log \lambda F_\lambda / d \log \lambda$ , where  $\alpha > 0.3$  for a Class I, 0.3 to  $-0.3$  for a flat-spectrum source,  $-0.3$  to  $-1.6$  for a Class II, and  $< -1.6$  for a Class III. For each of the YSOs and candidate YSOs in the sample, I performed a simple ordinary least squares linear fit to **all available photometry** (just detections, not including upper or lower limits) **as observed between 2 and 24  $\mu\text{m}$ , inclusive**. Note that formal errors on the IR points are generally so small as to not affect the fitted SED slope. Note also that the fit is performed on the observed SED, e.g., no reddening corrections are applied to the observed photometry before fitting. Classification via this method is provided specifically to enable comparison to other YSOVAR papers via internally consistent means; see discussion in R14. I can only perform this calculation for those objects with points at more than one wavelength in their SED between 2 and 24  $\mu\text{m}$ . Therefore, objects for which I have, e.g., one 3.6  $\mu\text{m}$  point from the cryogenic era and one 3.6  $\mu\text{m}$  YSOVAR point (the mean YSOVAR measurement at this band) cannot have a fitted SED slope.

The SED slopes and classes I calculate appear in Table 3. Slopes can be calculated for any of the objects with at least two distinct wavelengths in the SED between 2 and 24  $\mu\text{m}$ , but it is only meaningful if the source is a YSO, so I only include slopes and a class for objects tagged as YSOs by any of the references I included in this paper. Of the  $\sim 300$  objects in the catalog that are tagged as YSOs, where I can calculate the slope between 2 and 24  $\mu\text{m}$ , I find 55 Class Is, 38 Flat class, 103 Class IIs, and 81 Class IIIs. There are 25 with insufficient points in the SED, such that I cannot assign a class.

In R14, we asserted that the classes are generally well matched whether or not the 24  $\mu\text{m}$  point is included. To understand the influence of the 24  $\mu\text{m}$  point, in the cases where there is a 24  $\mu\text{m}$  point, I can compare the 2-24  $\mu\text{m}$  and 2-8  $\mu\text{m}$  SED slopes. Figure 7 compares the fitted slope for the 2-24  $\mu\text{m}$  and 2-8  $\mu\text{m}$  approaches, just for those sources detected at 24  $\mu\text{m}$ , and also identified as candidate YSOs. (The 24  $\mu\text{m}$  detections are strongly biased towards cluster members, so this figure includes most of the 24  $\mu\text{m}$  detections.) The vertical and horizontal lines indicate the divisions between SED classes as defined above. For  $\sim 65\%$  of this subset of objects, the resulting SED class is the same. (Table 6 has the total numbers.) The objects that do not match typically have a borderline slope. As expected, there is a bias such that inclusion of the 24  $\mu\text{m}$  point frequently pushes an object to more embedded SED classes. There are a few sources that are approximately photospheric until a sharp rise at 24  $\mu\text{m}$ ; those are difficult to classify correctly using our approach. Of the  $\sim 300$  things identified as YSOs (or candidates),  $\sim 150$  are detected at 24  $\mu\text{m}$ , and  $\sim 100$  have the same class even including the 24  $\mu\text{m}$  point.

In R14, we also asserted that our classes are generally well-matched to the G09 classes. Arnold et al. (2012) also report classes for YSO candidates. Table 6 has the numbers of objects in our catalog for which my classes match (or do not match) the classes obtained from these other approaches. The first thing to notice is that both G09 and Arnold et al. (2012) have different class bins than I have defined – G09 has no flat class, and a “II/III” class, and Arnold et al. have transition disk (TD) and pre-transition disk (PTD) classes. Despite this, the majority of the sources have the same class regardless of approach.

For the most embedded sources, there is no provision in our scheme for identifying Class 0s, and I do not use points at wavelengths longer than  $24\ \mu\text{m}$  to determine classes. I would expect, though, that those objects that others have identified as Class 0s would work out to be Class Is in my scheme. There are 10 sources with SED classifications from the submm in Hatchell et al. (2007) for which I have counterparts – in Hatchell et al., eight of them are Class 0 and two are Class I, and by my classification, nine are Class I and one is Class II. There are eleven Class 0 sources in the region identified in Sadavoy et al. (2014) from Herschel data – out of those 11, all but one are Class I, but that last one is Class II. Both of those discrepant sources (J032914.96+312031.7 for the Hatchell source and J032906.45+311534.4 for the Sadavoy source) have been discussed above. The matches may not be a good match to the source of the long wavelength flux. At any rate, in most cases, my method at least recovers Class I status for these very embedded objects.

Ideally, one would have a fully-populated SED, as well as a spectral type, and thus be able to attempt SED modelling to determine the proper classes for these objects, but this is not possible at this time.

Table 6. Comparison of Classes

Other class	Our Class I	Our Flat Class	Our Class II	Our Class III
Literature tagged YSOs: 2-8 $\mu$ m Class I	25	7	1	0
Literature tagged YSOs: 2-8 $\mu$ m Flat Class	9	12	5	0
Literature tagged YSOs: 2-8 $\mu$ m Class II	8	11	49	0
Literature tagged YSOs: 2-8 $\mu$ m Class III	4	0	8	8
Any matches: G09 Class I,I*	29	9	0	0
Any matches: G09 Class II,II*	5	19	67	3
Any matches: G09 Class II/III	0	0	4	0
Any matches: Arnold12 class I	12	2	0	0
Any matches: Arnold12 class FS	2	7	2	0
Any matches: Arnold12 class II	0	1	33	0
Any matches: Arnold12 class III	0	0	0	1
Any matches: Arnold12 class PTD	0	1	3	0
Any matches: Arnold12 class TD	0	2	3	0

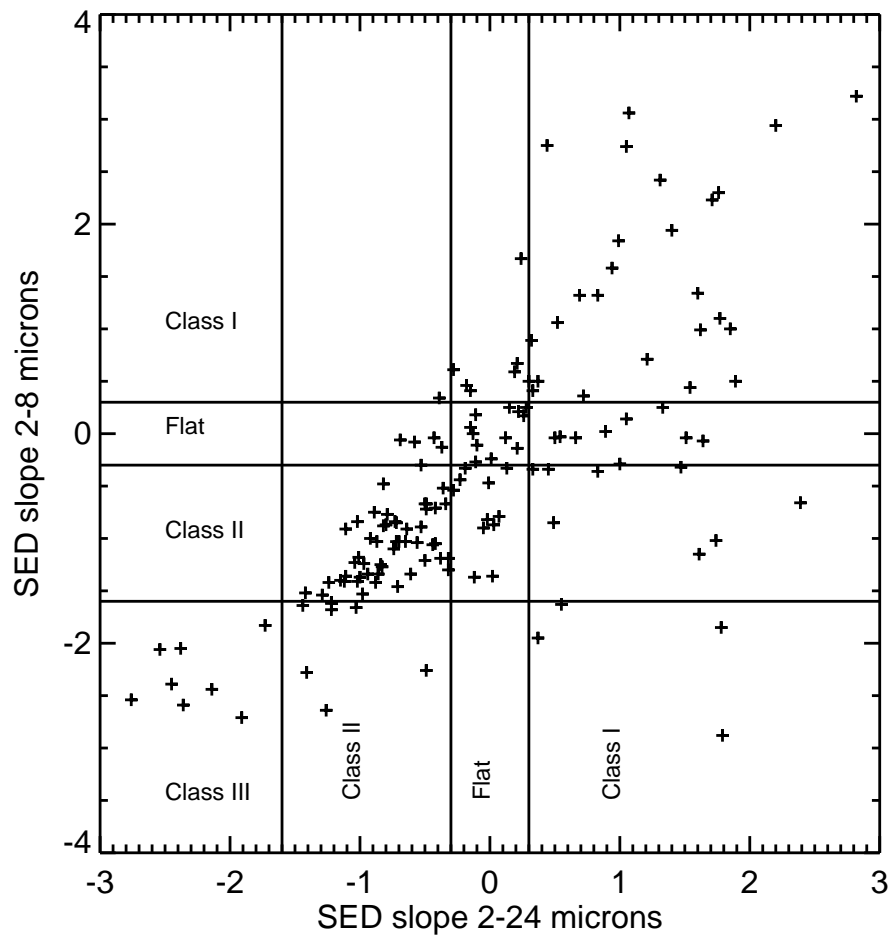


Fig. 7.— A comparison of the SED slopes calculated from the observed SED, just for those objects with  $24\ \mu\text{m}$  detections and identified somewhere in the literature as YSOs. Since the  $24\ \mu\text{m}$  detections are strongly biased towards cluster members, most of the points here are YSOs (or candidates). The vertical and horizontal lines indicate the divisions between SED classes as defined in Sec. 5. For 65% of the entire set of objects, the resulting SED class is the same. The objects that do not match typically have a borderline slope. As expected, there is a bias such that inclusion of the  $24\ \mu\text{m}$  point frequently pushes an object to more embedded SED classes.

## 6. Summary

I have presented a catalog of apparent point source objects towards NGC 1333, within  $52^\circ < \text{RA} < 52.5^\circ$  and  $31^\circ < \text{Dec} < 31.6^\circ$ . I have attempted to resolve a wide variety of confusion in the literature, ranging from mismatched sources to missing or duplicate IDs. I incorporated data from optical to radio wavelengths, but focused most of my effort on  $J$  ( $1.25\ \mu\text{m}$ ) to  $24\ \mu\text{m}$ . Cross-identifications include those from more than 25 papers and catalogs from 1994-2015.

I have also identified obvious holes in the accumulated literature, the most obvious one being spectroscopic confirmation of the many YSO candidates presented in the literature, along with spectral types.

While I have done the best that I can, the reliability of this catalog is likely lower for the longer wavelengths, and for the most bright and/or confused regions in the heart of NGC 1333. Image inspection proved invaluable in making many of these associations.

I have compared a few different methodologies for classifying objects by the SED shape and/or IR colors. While the methods agree in most cases, they can fail in identifying the most embedded sources, and in sources that are on the borderline between SED class slope definitions, and can more weakly depend on whether or not there is a detection at  $\sim 20\text{-}25\ \mu\text{m}$  to anchor the SED slope between 2 and  $24\ \mu\text{m}$ .

We will use this catalog as the basis for our upcoming work using YSOVAR data in NGC 1333.

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