

New asteroseismic rotation rates of *Kepler* dwarfs show strong agreement with weakened magnetic braking on the late-age main sequence

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Studies using asteroseismic ages and rotation rates from star-spot rotation have indicated that standard age-rotation relations may break down roughly half-way through the main sequence lifetime, a phenomenon referred to as weakened magnetic braking. While rotation rates from spots can be difficult to determine for older, less active stars, rotational splitting of asteroseismic oscillation frequencies can provide rotation rates for both active and quiescent stars, and so can confirm whether this effect really takes place on the main sequence.

We obtained asteroseismic rotation rates of 91 main sequence stars showing high signal-to-noise modes of oscillation. Using these new rotation rates, along with effective temperatures, metallicities and seismic masses and ages, we built a hierarchical Bayesian mixture model to determine whether the ensemble more closely agreed with a standard rotational evolution scenario [1], or one where weakened magnetic braking takes place [2]. The weakened magnetic braking scenario was found to be 98.4% more likely for our stellar ensemble, adding to the growing body of evidence for this stage of stellar rotational evolution. This work represents the largest catalogue of seismic rotation on the main sequence to date, opening up possibilities for more detailed ensemble analysis of rotational evolution with *Kepler*.

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Gyrochronology is the study of the relationship between a star’s rotation period and its age. As a star grows older along the main sequence, magnetic winds will cause it to lose angular momentum, slowing it down. Because the loss rate is related to temperature, the rotation period of a young star will rapidly settle on to a plane in age-colour-rotation space [3]. As a result, knowing the rotation and colour of a star provides an avenue to measure its age, to a precision of $\sim 10\%$ for stars most similar to the Sun [4]. Age can otherwise be difficult to come by, and so gyrochronology enables more in-depth studies of stellar populations and individual stars [5, 6].

Gyrochronology was previously calibrated on stellar clusters, which have well constrained ages, but only up to roughly 2.5 Gyr so far [4]. With the launch of the *Kepler* mission [7], ages of main sequence field stars (i.e. not in clusters) became more widely available through asteroseismology, the study of stellar variability [8]. When looking at these stars, disagreements were found between the observations and gyrochronology beyond the middle of stars’ main sequence lifetimes, which could not be reconciled with existing theories [9, 10, 11]. It was proposed that at some stage in a star’s evolution it undergoes *weakened magnetic braking* (WMB, [2]), where the efficiency of angular momentum loss rapidly drops. This would cause stars to keep fast rotation rates at old ages, which we would not expect from existing gyrochronology relations.

The mechanism by which weakened magnetic braking occurs has since been the subject of ongoing debate, and may be connected to changes in the magnetic field morphology [12, 13, 14, 15]. It may also be explained from an observational point of view. A large scale survey of stellar rotation rates measured using star-spots pointed out a lack of slowly rotating stars older than Sun [16]. As activity reduces with age, older stars with fewer star-spots are less likely to have their rotation measured [17]. The point at which the detection probability drops appears to lie at a similar level of activity as when the proposed weakened magnetic braking takes place, pitting these two possibilities against one another [18].

Determining whether weakened magnetic braking is a true phenomenon or a bias in star-spot observations requires new data, which can be provided by asteroseismology. A star’s rotation causes modes of oscillation to split into multiplets, making it possible to obtain an independent rate of rotation by measuring the splitting between oscillation frequencies [19]. Asteroseismic rotation is more difficult to obtain than spot rotation, but does not require spots to be visible, meaning that we can measure rotation for older, quiescent stars that would not have been present in existing rotation catalogues. This makes asteroseismic rotation rates the perfect avenue to explore whether weakened magnetic braking truly takes place on the main sequence. Here, we measured asteroseismic rotation rates of 91 main sequence stars across a broad range of ages, and evaluated this new ensemble to determine whether it agrees more closely with a classical rotational evolution scenario, or one where weakened magnetic braking takes place.

Asteroseismic main sequence targets: In order to obtain robust asteroseismic rotation rates, we required detections of multiples of both dipole (denoted as $\ell = 1$) and quadrupole ($\ell = 2$) oscillations for each star, of which the latter have significantly lower signal-to noise. Radial oscillations ($\ell = 0$), which are also visible, do not split. In this work, we studied a sample of 95 of the highest signal-to-noise targets observed by *Kepler*, combining the ‘Kages’

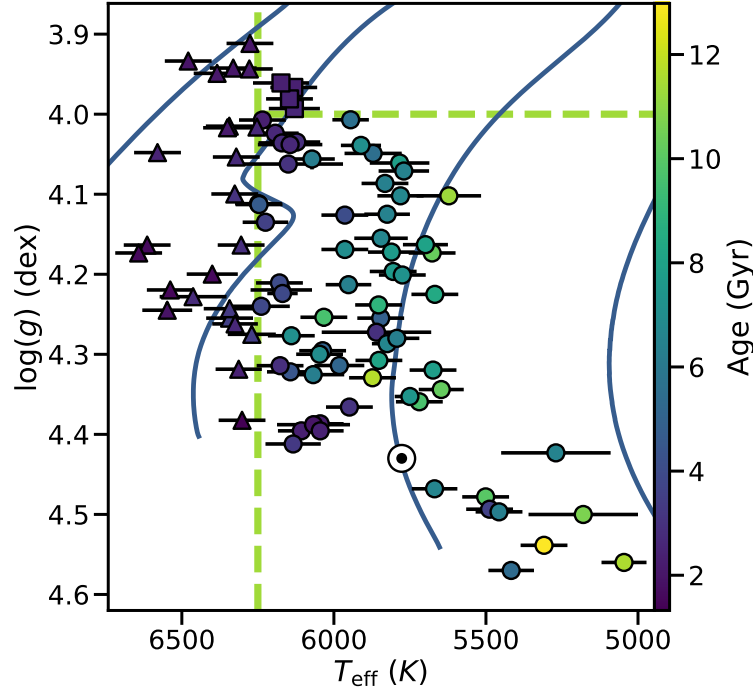


Figure 1: Our sample of 95 stars from the ‘Kages’ and LEGACY catalogues coloured by asteroseismic age [8, 22]. The dashed lines indicate our classification boundaries: main sequence (circles, $T_{\text{eff}} < 6250 \text{ K}$, $\log(g) < 4 \text{ dex}$), potential sub-giants (squares, $T_{\text{eff}} < 6250 \text{ K}$, $\log(g) > 4 \text{ dex}$), and ‘hot’ stars (triangles, $T_{\text{eff}} > 6250 \text{ K}$). The Sun is shown with the ‘ \odot ’ symbol, and has an age of 4.6 Gyr [26]. The solid lines are evolutionary tracks generated using MESA [27], for a metallicity of $Z = 0.01493$ and helium content of $Y = 0.26588$. Left to right, they represent masses of 1.5, 1.25, 1 and $0.75 M_{\odot}$. Vertical errorbars are too small to be visible [28].

[8, 20] and LEGACY [21, 22] catalogues*. Further details on these data can be found in **Methods**.

We separated our sample into three categories [23]: 64 main sequence stars (MS, $\log(g) > 4.0 \text{ dex}$, $T_{\text{eff}} < 6250 \text{ K}$); 4 potential sub-giant stars (SG, $\log(g) < 4.0 \text{ dex}$, $T_{\text{eff}} < 6250 \text{ K}$), which may have begun to evolve off the main sequence, complicating their rotational history[†]; and 24 ‘hot’ stars (H, $T_{\text{eff}} > 6250 \text{ K}$), which lie above the Kraft break, the point at which convective envelopes are tenuous, and angular momentum loss through magnetic winds becomes inefficient [25]. Our conclusions are largely insensitive to these categorical assignments, and they are intended solely to help explore results category-by-category. The sample, which can be seen in Figure 1, spans surface gravities of $3.8 \text{ dex} < \log(g) < 4.6 \text{ dex}$, temperatures of $5000 \text{ K} < T_{\text{eff}} < 6700 \text{ K}$, and ages from 1 to $\sim 13 \text{ Gyr}$.

*For these catalogues the mode extraction through frequency fitting [20, 21] and modelling using mode frequencies to obtain stellar parameters [8, 22] are covered in separate papers.

[†]It should be noted that none of these stars have identified mixed dipole modes, which indicate an evolved structure (see [24]).

Results of asteroseismic fitting: The stars in ‘Kages’ and LEGACY which form our stellar sample have been subject to detailed asteroseismic frequency fitting in the catalogue papers, but did not have reported asteroseismic rotation. Here, we repeated the mode frequency fitting with a Bayesian model that accounted for rotation in more detail, and improved the treatment of existing asteroseismic relations in frequency fitting techniques (see **Methods** and **Supplementary Information**). We fit our new model to *Kepler* power-spectra of 95 stars, which was successful in 94 cases. The summary statistics used throughout the rest of this paper were taken as the median of the posterior distribution on the model parameters, with the uncertainties being the 15.9th and 84.1st percentiles, meaning that some reported parameters (such as period) have asymmetric uncertainties.

We compared our results with both published and unpublished asteroseismic rotation rates [20, 21, 11, 10, 29], and found that the three stars with the lowest angles of inclination ($< 10^\circ$) were inconsistent with multiple other studies. To some degree this was expected, as for stars viewed near pole-on the split modes of oscillation have extremely low power, making them hard to detect reliably. We flagged our rotation rates for these stars, and did not use them in the next steps in our analysis, leaving a sample of 91 stars for which we considered our asteroseismic rotation measurements to be robust (for a full justification, see **Methods**).

Seismic versus surface rotation rates: Different methods of rotation measurement probe different depths of stars. Asteroseismology of main sequence stars probes internal rotation in the near surface layers, where the observed modes of oscillation are most sensitive [30]. Measurements of star-spot modulation instead probe the rotation rates of star-spots on the surface. Previous studies have shown that seismic rotation rates show no statistically significant deviation from spot rotation rates [10, 31]. We confirm that this holds true for our larger sample, in order to ensure that our seismic rotation periods can be reliably used to draw conclusions on the evolution of surface rotation.

Figure 2 shows a comparison between 48 stars from our sample for which spot rotation rates had previously been measured. Rotation rates from spot modulation are subject to measuring multiples of the true rotation rate, and also sometimes over-correct for this effect. This is not the case for asteroseismic rotation, and so some stars may appear at the 2:1 and 1:2 lines in the figure. We fit a line of the form $P_{\text{surf}} = m \times P_{\text{seis}}$, using the larger of the asymmetrical uncertainties on asteroseismic rotation, and excluding stars above the 1.8:1 line to avoid outliers (transparent in the Figure). Our fit found a value of $m = 0.96 \pm 0.03$, showing a close agreement ($< 2\sigma$) between asteroseismic and surface rotation rates on a population level. This overall agreement between spot and seismic rotation rates is in line with previous studies of Sun-like stars [32, 33, 10, 31], and so we concluded that our asteroseismic ensemble can be reliably used to draw inferences about gyrochronology.

Comparing gyrochronology models: To evaluate the implications of our seismic ensemble for gyrochronology, we compared our new asteroseismic rotation rates to two population models of rotational evolution [18]: a ‘standard’ model, which assumes a traditional angular momentum transport through magnetically driven stellar winds [35, 36], and one where weakened magnetic braking takes place (hereafter the ‘WMB model’) [2] The WMB model was identical to the standard model in its input physics, except for the condition that

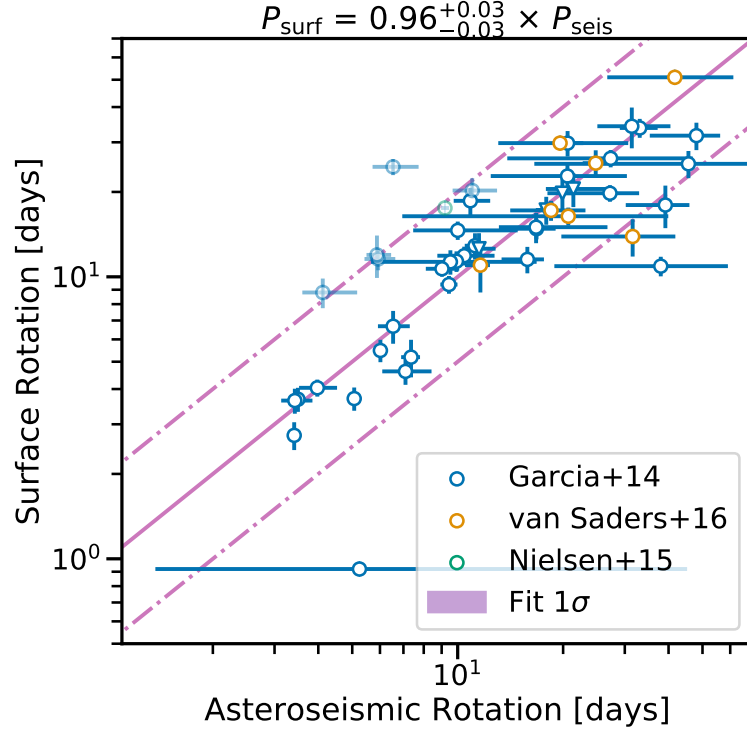


Figure 2: Comparisons between asteroseismic and photometric measures of stellar rotation, for 48 stars. Literature values were taken, in order of priority, from: [23] and [34] (i.e. if a value was reported in [34] and also in [23], the latter was used). Stars used in [2] are highlighted. The four triangles represent stars included in the original [10] work. The solid line indicates 1:1 line, while the dash-dotted lines represent 2:1 and 1:2 lines. The shaded region around the 1:1 line is the 1σ credible interval of the fit to the data, the result of which is shown as the title of the figure. Stars that were not included in the fitting process are transparent. The errorbars on the asteroseismic rotation represent the 15.9th and 84.1st percentiles on the measured posterior distributions.

angular momentum loss ceases above a critical Rossby number of $Ro_{\text{crit}} = 1.97^\dagger$

Instead of comparing our rotation rates directly to models of individual stars, we instead compared them to synthetic TRILEGAL populations based on what the *Kepler* field would look like under these two different braking laws [37, 18]. These synthetic populations were then altered to match the distribution of existing catalogues of stars in the *Kepler* field in order to replicate the *Kepler* selection effect [38], and evolved following a standard or WMB braking law. For the full details on how these models were constructed, see **Methods**.

In order to ensure our results were based on the most robust rotation measurements, we focused on the stars with the best convergence metrics in the mode frequency fitting. Excluding those with a number of effective samples $n_{\text{eff}} < 25\%$ and a Gelman-Rubin metric of $\hat{R} > 1.1$ [39, 40] (see **Extended Data Table 1**). The remaining sample of 73 stars contained 4 potential sub-giants, 22 ‘hot’ stars and 47 main sequence stars. Figure 3 shows these stars plotted over the top of both the standard and WMB population models in rotation-temperature space, alongside the stars not used in the gyrochronology analysis (shown without uncertainties).

We evaluated our ensemble against both population models simultaneously using a Bayesian mixture model, which treated the data as being drawn from a mixture made from Kernel Density Estimates (KDEs) of both populations, modulated by a mixing parameter P_s . In the limit that $P_s \rightarrow 1$, the data are most likely drawn from the standard model. In the limit $P_s \rightarrow 0$, the WMB model is more likely. This mixture model was evaluated in a five-dimensional parameter space, accounting for the population distributions in effective temperature (T_{eff}), metallicity ([Fe/H]), asteroseismic mass (M) and age (t) and our new rotation periods (P). For further detail, see **Methods**.

Results of comparing models: We fit our mixture model to each star individually, which yielded a posterior probability density for the value of P_s for that star. In order to assess the posterior probability for the full sample, we multiplied the individual posterior probabilities for P_s . This was done for the full ensemble of 73 stars, as well as for the three different stellar types. These joint probabilities can be seen in Figure 4 in the left-hand plot. The cumulative probability can be seen in the right-hand plot. In both of these plots, the left-hand side holds 98.4% of the probability, indicating that the weakened magnetic braking model is preferred. When only considering the 47 main sequence stars, for which we expect the effects of weakened magnetic braking to be strongest, this rises to 99.2%.

Both hot and sub-giant stars do not strongly prefer one model over the other (both crossing the $P_s = 0.5$ line at roughly 45%), which is expected given that the rotation rates of both these stellar types are not expected to be strongly influenced by weakened magnetic braking. When performing this analysis on the full ensemble of 91 stars for which seismic rotation was measured, this the total probability below $P_s = 0.5$ is 96.6%.

Verifying consequences for gyrochronology: In order to verify that our results presented a meaningful confirmation of weakened magnetic braking, we ensured that the same results were obtained when using different subsets of the ensemble, and when exploring poten-

[†]The Rossby number may be defined as $Ro \equiv P/\tau_{\text{cz}}$, where τ_{cz} is the convective turnover timescale, and P is the rotation. The Rossby number roughly scales inversely with stellar activity (and so increases as stars age).

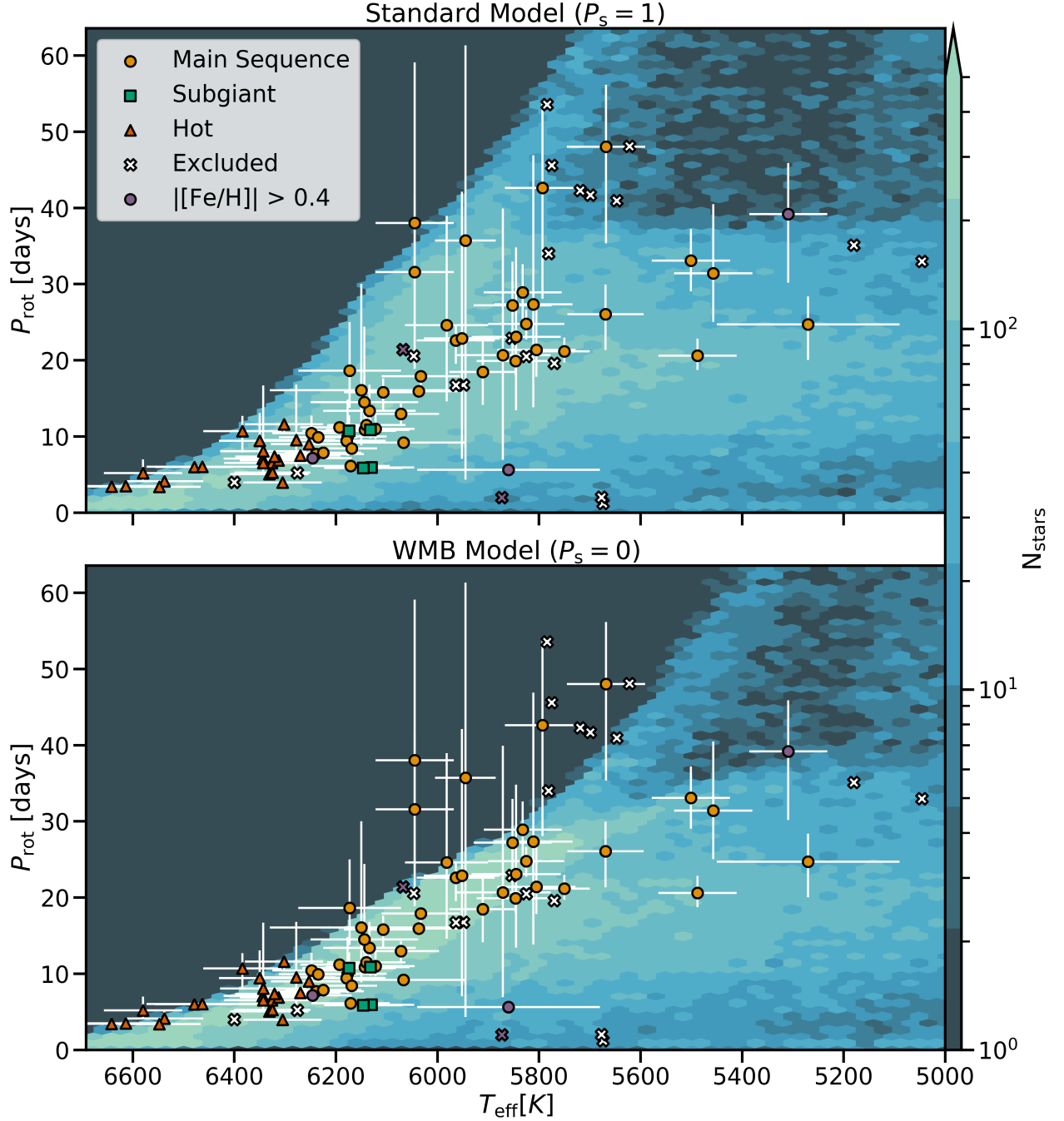


Figure 3: Stars for which rotation was measured in this work, plotted over TREILGAL [37] stellar population models evolved under standard (top) and weakened magnetic braking (WMB, bottom) prescriptions of angular momentum evolution. The 21 Stars that were excluded from the gyrochronology analysis are marked with crosses and without uncertainties (see text). Circles, squares and triangles denote main sequence, sub-giant, and ‘hot’ stars respectively. Stars with a metallicity of $|[\text{Fe}/\text{H}]| > 0.4$ are shown in purple, to indicate that they fall outside the functional range of the [18] stellar models. Note that only period-temperature space is shown here, but when evaluating between the two model prescriptions we also considered mass, age and metallicity. The errorbars on the asteroseismic rotation represent the 15.9th and 84.1st percentiles on the measured posterior distributions.

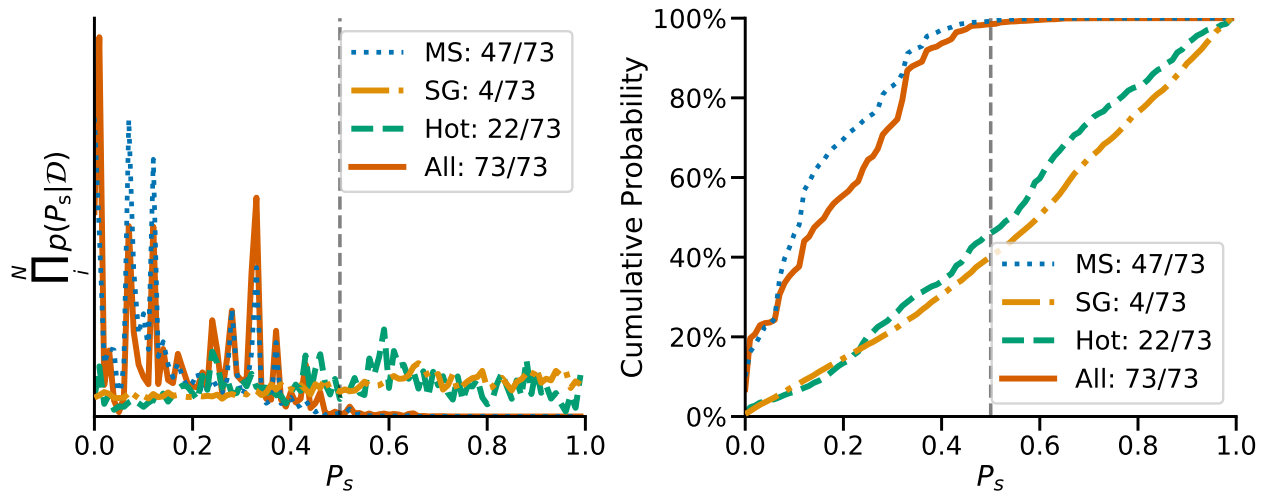


Figure 4: Posterior estimates of the mixture model parameter P_s , broken down by stellar classification. A value of P_s close to zero indicates that the data are more consistent with a rotational evolution that includes weakened magnetic braking, whereas a value close to 1 indicates that the data are more consistent with a standard rotational evolution scenario. *Left:* the product of histograms of the posterior estimates for 73 stars, or subsets thereof for different stellar classifications. *Right:* the cumulative posterior probability for 73 stars, or subsets thereof for different stellar types.

tial biases. We found that the joint posterior distribution did not change significantly when excluding stars with a metallicity outside the model range (purple in Figure 3), or those with extremely small uncertainties on asteroseismic mass. Stars in our sample with known binary companions that may disrupt their rotational evolution were not found to contribute strongly to the joint posterior probability of P_s .

We checked that our results remained the same if the ‘Kages’ and LEGACY catalogues had underestimated the uncertainties on the asteroseismic mass and age, which form part of our mixture model analysis. We reran our analysis after inflating the uncertainties on these properties to align with the expected systematic uncertainty from stellar models [8, 22]. The results from this echoed what we found for the unaltered ensemble. We also accounted for the possibility that asteroseismic methods systematically overestimated stellar age, and reran our analysis after shifting the seismic ages younger by their systematic uncertainty. This resulted in 95.8% of the posterior probability to lie below $P_s = 0.5$, compared to 98.4% for the unaltered ensemble. From this, we can infer that our conclusions are robust against potential issues in seismic analyses. For a full discussion of these tests and other biases, see **Methods**.

Weakened magnetic braking was first proposed by van Saders et al. (2016) [2, hereafter the ‘van Saders’ study] in response to stars with spot rotation rates faster than expected from gyrochronology at their asteroseismic ages. This discrepancy was also indicated at around the same time in other studies of asteroseismic ages of main sequence stars [10, 9, 11]. The theory of weakened magnetic braking has been both reinforced by recent studies [41], as well as disputed [42], at least at the critical Rossby number originally proposed.

A recent study using spectroscopic rotation measurements of 14 solar twins, compared

to grids of stellar models, found marginal evidence favouring a standard rotational evolution [42]. If weakened magnetic braking were to take place, this work proposed that it would most likely be outside the range of their sample, at $Ro_{\text{crit}} \geq 2.29$, compared to the value of 1.97 used in this work. Our sample does not overlap with theirs, so we can not directly compare to their results. Our work does find statistical agreement with a model of weakened magnetic braking at a critical Rossby number of 1.97 over one without any braking at all, indicating that traces of weakened magnetic braking do occur at Rossby numbers of around 2. A comparison to models of different Rossby numbers will be explored in future papers, using the LEGACY and ‘Kages’ rotation sample presented in this work.

As a sanity check that the evidence in favour of weakened magnetic braking presented in this work is independently robust, we recreated the joint posterior probability for P_s , this time looking only at MS stars, and excluding any of the 22 stars that were included in the original van Saders study used to detect weakened magnetic braking. Of the 47 MS stars used to distinguish between stellar models, 16 were included in the van Saders sample. Both the 16 stars in the original van Saders study as well as the remaining MS stars favoured the WMB model, although the van Saders stars did so more strongly, not including any stars that favoured the standard model. Specifically, when considering only the van Saders stars, 96.3% of the total joint probability lay below $P_s = 0.5$, compared to 91.7% when considering only those stars *not* included in the van Saders study. The stars initially used to propose the WMB model are still those most strongly in favour of it even when using asteroseismic rotation rates only, a result which is now further reinforced by an expanded sample that includes quiescent stars.

Conclusions: In this work, we obtained asteroseismic rotation rates for 91 stars using a hierarchical Bayesian model to fit for oscillation frequencies, convective background, and rotational splitting simultaneously. We validated these results against previously published and unpublished asteroseismic rotation measurements, making our new rotation catalogue the largest self-consistent sample of asteroseismic rotation in main sequence stars to date. We evaluated the stars in our sample against synthetic population models of the *Kepler* field evolved under either a ‘standard’ model of rotational evolution, or one where weakened magnetic braking takes place. This evaluation was done using a Bayesian mixture model to determine whether the ensemble preferred one model over another, and was done simultaneously in asteroseismic rotation, mass, age, and spectroscopic temperature and metallicity.

We leave the reader with the following conclusions:

1. We found that our ensemble was 98.4% more likely to be drawn from a weakened magnetic braking model than a standard model of rotational evolution [1], at a critical Rossby number of $Ro_{\text{crit}} = 1.97$. This work expands upon previous analysis [18] by including quiescent stars older than the Sun. This conclusion was found to be robust against change of ensemble members and potentially underestimated asteroseismic systematic uncertainties. A comparison to other braking laws [17] and other gyrochronology relations [43] using our asteroseismic ensemble will take place in future work.
2. We compared our new asteroseismic rotation rates with surface rotation measures from spot modulation. Our findings replicate those in similar studies [10, 31], in the sense

that we find no statistically significant difference between seismic and spot modulation measures of stellar rotation in our ensemble.

3. We present new techniques for asteroseismic ‘peak-bagging’ frequency analyses, specifically: expressing well known empirical asteroseismic relations in a hierarchical modeling structure to account for small-scale effects not included in our model (such as acoustic glitches), and replacing uninformative empirical relations with Gaussian Process priors, detailed in the **Supplementary Information**.

In the near future our new asteroseismic rotation catalogue will be further complemented with high-precision surface rotation measurements and atmospheric parameters from spectroscopic surveys such as LAMOST [44], 4MOST [45] and WEAVE [46], and new asteroseismic measurements of age from the K2 and *TESS* missions. With these surveys, large scale ensemble asteroseismology will continue to increase the possibilities for understanding gyrochronology.

All analysis performed as part of this work is stored in an online public repository[§].

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[§]<https://github.com/ojhall194/malatium>

implications of the gyrochronology results. All authors have contributed to the interpretation of the data and the results, and all discussed and provided comments for all drafts of the paper.

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Methods

Contents

1	Asteroseismic Methods	12
1.1	Asteroseismic Data	12
1.2	Mode Frequency Fitting	13
2	Gyrochronology Methods	13
2.1	Stellar Models	13
2.2	Bayesian Mixture Model	15
3	Verifying asteroseismic results	17
3.1	Priors on rotational parameters	17
3.2	Comparisons to previous studies	17
3.3	Seismic vs spectroscopic rotation	19
3.4	Asteroseismic detection biases	19
4	Verifying consequences for gyrochronology	20
4.1	Limits of our stellar models	20
4.2	Systematic uncertainties from asteroseismology	21
4.3	Binaries and Planet Hosts	22
5	Supplementary Information	23
5.1	Asteroseismic Model	23

1 Asteroseismic Methods

1.1 Asteroseismic Data

For our asteroseismic power spectrum data we used the unweighted power spectra from the KASOC pipeline [47][¶]. We did not apply any additional treatment to these data. For 16 Cyg A & B (KIC 12069424 and KIC 12069449) we used the KEPSEISMIC lightcurves [48]^{||}, which have significantly better signal-to-noise for these two stars.

For our asteroseismic ages, we used the ages obtained by **BASTA** [8, BAYesian STellar Algorithm] in the ‘Kages’ and LEGACY catalogues. These ages have been obtained by comparisons of measured oscillation properties to stellar models, accounting for an expanded range of metallicities. **BASTA** is thoroughly compared to four other seismic modelling techniques in [22]. While uncertainties found through **BASTA** are typically higher than for other techniques, only **BASTA** and **ASTFIT** [49, Aarhus STellar Evolution Code] recover the radius, mass and age of the Sun, when applied to solar data. Although the uncertainties on **ASTFIT** ages are overall lower, they are not published for the ‘Kages’ sample. In order to maintain an

[¶]Obtainable from the KASOC webpage.

^{||}Obtainable from MAST.

internally consistent stellar age sample, we used age results from **BASTA** for both the ‘Kages’ and LEGACY samples.

For our stellar masses we used asteroseismic model masses obtained by **BASTA** reported in ‘Kages’ and LEGACY, in order to maintain internal consistency with the age measurements. We note that age and mass posteriors from **BASTA** are correlated, but chose not to account for the unpublished correlations in this work.

As described in the catalogue papers, for ‘Kages’ stars atmospheric properties (T_{eff} and $[\text{Fe}/\text{H}]$) were measured through high-resolution spectroscopy reported in [50]. For LEGACY stars, atmospheric properties were similarly taken from [51] for most stars in the catalogue, and complemented by other values from the literature for the remaining stars [22, see Table 3].

Other asteroseismic properties used in this study, such as those used as first guesses on the free parameters in our asteroseismic model, were taken from the ‘Kages’ and LEGACY catalogues [52, 53, 54]. In cases where both catalogues contained the same target, we used the stellar parameters reported in LEGACY [55]. This is also the case for the masses and ages described above.

1.2 Mode Frequency Fitting

Todo: water down the contents of the methods here.

From our asteroseismic analysis, we report inclination angle, rotational splitting, and rotation period in **Extended Data** Table 1. The summary statistics on these parameters were taken as the median of the posterior distribution, with uncertainties being the 15.9th and 84.1st percentiles. For inclination angle ($i = \arccos(\cos(i))$) and rotation ($P = 1/\nu_s$, the full posterior samples were transformed before taking the summary statistics, as our model sampled in $\cos(i)$ and ν_s .

We flagged any sub-optimal conditions of the final fit. Of 94 stars, we flagged 5 for which the Gelman-Rubin convergence metric, \hat{R} , was greater than 1.01 and 2 stars for which it was greater than 1.1 on inclination angle or rotation, where a $\hat{R} = 1$ indicates a converged result [39]. We also performed visual checks of the sampled chains on all hyperparameters and of the best-fit model compared to both the raw and smoothed asteroseismic data. We found no issues in the visual investigation of the 94 stars. KIC 8478994 was excluded due to both a poor fit as well as high \hat{R} on rotational parameters. KICs 6603624, 8760414 and 8938364 are reported in **Extended Data** Table 1, but were excluded from the gyrochronology analysis due to strong disagreement with multiple studies in the literature (see main body of paper). Finally, we flagged any stars with fewer than 1000 effective samples (n_{eff}) of ν_s .

2 Gyrochronology Methods

2.1 Stellar Models

The braking models used in this work have several parameters that inform the rotational evolution. These are: a normalization factor to reproduce the solar rotation (f_k); a disk locking timescale (T_{disk}) and period (P_{disk}) which together regulate the stellar angular velocity during the pre-main sequence; the critical angular velocity that marks the transition from

saturated (rapidly rotating) to unsaturated (slowly rotating) regimes (ω_{crit}), and the critical Rossby number, above which stars conserve angular momentum (Ro_{crit}), mentioned above. T_{disk} , P_{disk} and ω_{crit} are calibrated to match the behavior in young open clusters, but have little impact on the rotational evolution beyond ≈ 1 Gyr in solar-mass stars. Both f_k and Ro_{crit} affect the late-time evolution. Both models adopt $\omega_{\text{crit}} = 3.4 \times 10^{-5} \text{ s}^{-1}$, $P_{\text{disk}} = 8.1$, $T_{\text{disk}} = 0.28$ and $f_k = 6.6$. In the weakened magnetic braking model, $Ro_{\text{crit}} = 1.97$. For further details, see [1] and [2, 18] for the standard and weakened magnetic braking cases respectively.

To construct a synthetic population of rotating stars, we expanded upon previous forward-models of the *Kepler* field with the purpose of studying weakened magnetic braking [18], with some important improvements. We started with a TRILEGAL [37] Milky-Way simulation of the *Kepler* field, using the simulation’s standard population values intended for this purpose. In order to replicate the *Kepler* selection effect, the TRILEGAL simulation was matched to the largest current catalogue of temperatures, luminosities and 2MASS *K*-band magnitudes of stars in the *Kepler* field [38]. This was done using a nearest-neighbours approach, based on the density of stars on the HR diagram. This ensured that the TRILEGAL population matched the density of stars actually observed by *Kepler*, replicating its selection effects, and improved upon previous efforts [18] by incorporating contemporary *Gaia* mission data [56].

We made further changes to account for possible binarity in the the matching sample [38]. If the first step is performed blindly, blended binaries in the sample cause an overestimation of the number of old stars. In order to overcome this, we:

1. blended the TRILEGAL stars with binary companions drawn from a flat mass-ratio distribution, using a known binary fraction [57],
2. recalculated the ‘observed’ luminosities and magnitudes assuming that each binary pair was blended, and
3. shifted these stars’ temperatures following a *g-K* magnitude relation [38].

This new distribution was used for the nearest-neighbour matching. Once drawn we dropped the binary companion and used the true TRILEGAL properties of those stars. For stars where the mass of the companion was $M < 0.4 M_{\odot}$, binary contributions were ignored. Every binary was assumed to result in a blend, regardless of separation. This results in slightly more young stars than reality, because young, blended binary systems contaminate regions of the HR diagram where one expects to find old stars, and the number of blends is overestimated by assuming every binary system is a blend.

Our asteroseismic sample of stars with short cadence observations are subject to additional selection functions not included in the creation of the model populations above. We did *not* explicitly account for these asteroseismic selection functions in our model, by design. Both the standard and WMB models contain stars with the same fundamental parameters (mass, radius, effective temperature, metallicity) but a different period based on the choice of rotational evolution prescription. Applying an asteroseismic selection function that depends on these fundamental parameters would apply an identical prior to both models, therefore providing no net effect on our posterior distribution [58]. Additionally, we expected any seismic selection function to be relatively flat (and therefore uninformative) on a star-by-star scale, on which we run our model analysis.

2.2 Bayesian Mixture Model

In order to determine whether weakened magnetic braking occurs on the main sequence, we compared our sample of seismic age and rotation, along with temperature, metallicity and mass, to the two stellar population models of the *Kepler* field adapted from [18], discussed above. Both stellar models were evaluated in a Bayesian framework, with the rationale of determining which of the two models (standard or WMB) is most likely to reproduce our observed data. Each model sample contained temperature (T_{eff}), mass (M), age (t), metallicity ($[\text{Fe}/\text{H}]$) and rotation (P) information.

In order to draw probabilistic inference about the models, we built a five-dimensional Kernel Density Estimate (KDE) of both model populations using the `statsmodels` package [59]. We used a band-width (setting the resolution of the KDE) of $0.02 M_{\odot}$ in mass, 10 K in T_{eff} , 0.01 dex in $\ln(t)$, 0.01 dex in $[\text{Fe}/\text{H}]$ and 0.01 dex in $\ln(P)$. Note that age and rotation were treated in log space, where the posterior estimates from asteroseismology more closely resemble normal distributions. This approach translates the population models to a probability distribution we can use in a Bayesian framework.

We evaluated our data against both models simultaneously by treating the data as being drawn from a mixture of both model KDEs. In this mixture model structure, the two KDEs were modulated by a weighting factor, P_s . In the limit $P_s \rightarrow 1$, the data are most likely drawn from the standard model. In the limit $P_s \rightarrow 0$, the data are most likely drawn from the WMB model.

The posterior probability of obtaining P_s and additional parameters θ given our data \mathcal{D} is $p(P_s, \theta | \mathcal{D})$. Using Bayes equation, we can express this as:

$$p(P_s, \theta | \mathcal{D}) \propto p(\mathcal{D} | \theta) p(\theta | P_s, \kappa_s, \kappa_{\text{WMB}}) p(P_s), \quad (1)$$

where κ_s and κ_{WMB} are the KDE functions for the standard and WMB models respectively, and θ here are parameters, $\theta = \{M, T_{\text{eff}}, \ln(t), [\text{Fe}/\text{H}], \ln(P)\}$. As done above for the asteroseismic model, the parameters θ may be referred to as latent parameters, as they form a step between the parameter we want to infer (P_s) and our data. Using this approach allowed our model to properly take into account the observational uncertainties on the data.

The second component on the right hand side of Equation 1 describes the probability of obtaining our latent parameters θ given our KDEs and the mixture model weighting parameter P_s , and is described by the mixture model

$$p(\theta | P_s, \kappa_s, \kappa_{\text{WMB}}) = P_s \times \kappa_s(\theta) + (1 - P_s) \times \kappa_{\text{WMB}}(\theta), \quad (2)$$

where all parameters are as described above. This probability function describes a distribution that is a mixture of both KDEs. While the KDEs are constant, P_s is a free parameter, and so the shape of this distribution can vary. The latent parameters θ are drawn from this distribution, and therefore from some combination of the two stellar models.

The first component in Equation 1 describes the likelihood of obtaining the parameters θ given our data and their observational uncertainty. It takes the form

$$p(\mathcal{D} | \theta) = \mathcal{N}(\mathcal{D} | \theta, \sigma_{\mathcal{D}}), \quad (3)$$

a normal distribution evaluating the latent parameters θ against the observations, with observational uncertainty $\sigma_{\mathcal{D}}$. This approach means that in each parameter space (such as age),

the age drawn from the stellar model mixture is entered into the likelihood equation with our static observations. The value of this equation (and thus the likelihood) will increase if θ is closer to the observations, and the mixture model will be modulated in a manner that maximises this probability, inferring whether one stellar model is more likely to produce our data than the other.

The final term, $p(P_s)$, represents the prior on the mixture model weight, which is uniform between 0 and 1. A visual representation of our model is shown in Figure 6.

Typically, this model would evaluate all stars in our sample against the stellar models simultaneously for a single posterior estimate of P_s . At 95 stars, in 5 parameter spaces, this totals 476 free parameters to marginalise over. This is not an issue for Hamiltonian Monte Carlo [60, HMC], however the use of KDE functions, over which a probabilistic gradient can not be measured, reduces HMCs effectiveness. Alternative Markov Chain Monte Carlo techniques [61, MCMC] can more efficiently sample the KDE functions, but can not treat the large number of hierarchical parameters. To overcome this, we fit our model to each star to obtain an independent individual posterior distribution for P_s , and multiplied these post hoc to obtain a combined posterior. This comes with the benefit of easily allowing us to calculate the combined posterior for different stellar classifications, at the expense of the ability to marginalise for a single value of P_s directly.

2.2.1 Fitting procedure

The parameter space of the stellar models populations were reduced before calculating the KDEs. These cuts were made in M , T_{eff} , $\ln(t)$ and $[\text{Fe}/\text{H}]$, removing any stars in the models that fell more than $3 \times \sigma_{\mathcal{D}}$ outside the observations. Our observables M , t and P have asymmetric uncertainties from the Bayesian asteroseismic analysis. We used the larger of the reported uncertainties on each parameter as $\sigma_{\mathcal{D}}$ in each parameter space.

KIC 6278762 was excluded from this stellar model analysis, because its age fell more than 3σ outside of the highest age in the stellar models**, and KICs 7106245 and 8760414 were excluded for the same reason due to low metallicities outside the functional range of the stellar models (-0.99 and -0.92 respectively). We also excluded any stars with $n_{\text{eff}} < 1000$ for rotational splitting, and those with $\hat{R} > 1.1$. The remaining sample of 73 stars contained 4 sub-giants, 22 ‘hot’ stars, and 47 main sequence stars.

We fit our model Equation 1 using `emcee` [61, 62], using 32 walkers for a total of 7500 samples per walker, of which the first 2500 were discarded as a burn-in.

After fitting, we took a normalised histogram of the posterior estimate of P_s for each star, using 100 bins. In this histogram, each bin approximated the value of the posterior function for P_s . The array of 100 bins for all stars were multiplied, resulting in an approximation of the joint posterior probability function for P_s given all stars in the ensemble.

**This is a metallicity issue, as the oldest stars have metallicities outside the range of the rotational model grids.

3 Verifying asteroseismic results

3.1 Priors on rotational parameters

In our Bayesian analysis, we have placed weakly informative priors on our sampled rotational parameters, $\nu_s \sin(i)$ and $\cos(i)$. The prior is especially important for the angle of inclination, which is hardest to infer from the data. We are able to validate the robustness of our asteroseismic results by confirming that their posterior distributions are data-dominated, and not prior-dominated. We can do so by comparing the 68% credible regions of the posterior estimates of $\nu_s \sin(i)$ and i against the 68% credible regions of their priors.

A comparison between prior and posterior is shown for 94 stars in $\nu_s \sin(i)$, i and P_{rot} in **Extended Data Figure ??**, arranged by age. In the Figure, results with means and error bars that closely resemble the prior distribution can be interpreted as prior-dominated (i.e. poorly informed by the data). The projected splitting, $\nu_s \sin(i)$, is overall well constrained, with only one star appearing prior dominated. This is expected, as the projected splitting is what we observe on the star before decoupling inclination and rotation. The angle of inclination i , sampled as $\cos(i)$, more closely follows the prior distribution in most cases. Combining the two, the rotation period P_{rot} has no stars directly corresponding to the effective prior on period, and globally follows a trend with increasing age. The three outliers with fast rotation at late ages (KICs 6603624, 8760414 and 8938364) are discussed in more detail in below.

The rotation rates as presented in this work are a product of our Bayesian sampling of both projected splitting and angle of inclination. As seen in the Figure there are instances where i or $\nu_s \sin(i)$ closely resemble the prior (and are therefore prior-dominated). There are no cases of this when looking at the resulting period measurements, as they will have been informed by at least one strongly data-driven parameter (judging from the Figure, commonly $\nu_s \sin(i)$). From this, we conclude that our ensemble of asteroseismic rotation is not strongly dominated by the priors imposed on projected splitting and inclination in our Bayesian analysis.

3.2 Comparisons to previous studies

In order to validate our results, we compared our rotational parameters to those obtained in the literature, as well as those resulting from the work presented in LEGACY and Kages, which were unreported and received through private communication by the authors of the catalogue papers.

Comparisons with LEGACY and ‘Kages’ are shown in **Extended Data Figure 7** for projected splitting, inclination angle, and rotation period. In all three cases we show the fractional difference between the values obtained in this work and those from LEGACY and Kages. On the right of the Figure, we show the distribution of the fractional differences for the three parameters. The projected splitting is in good agreement with both studies, however LEGACY finds slightly lower $\nu_s \sin(i)$ for the faster rotators, deviating from our work by over 1σ . Neither LEGACY nor ‘Kages’ used a spatially isotropic prior for the inclination angle in their analyses, instead opting for a uniform prior. As posterior estimates of inclination angle are only loosely data-driven, the introduction of an isotropic prior should result in our analysis reporting globally higher inclination angles. This effect is seen in the comparisons of both inclination angles and rotation rates for the LEGACY stars, where we find overall lower

rotation rates compared to LEGACY for stars at very similar $\nu_s \sin(i)$.

A number of stars are excluded from **Extended Data** Figure 7 and compared individually as extreme outliers: KICs 5094751, 6196457, 8349582, 8494142, 8554498, 105114430 and 11133306 all have fast rotation rates (< 5 days) in Kages, but are found to have a broader spread of rotation rates in this work. At similar values for $\nu_s \sin(i)$, ‘Kages’ found much lower inclination angles with highly asymmetrical uncertainties. Based on a comparison between the summary statistics of these stars, we conclude that the results found in this work have better marginalised over inclination angle, improving our measure of rotation.

Conversely, KICs 6603624, 8760414 and 8938364 have extremely slow rotation periods in LEGACY, but extremely *fast* (< 3 days) in this work. KICs 8760414 and 8938364 are excluded from the gyrochronology analysis below based on checks for \hat{R} and the number of effective samples. Both stars have ages greater than 10 Gyr, making their fast rotation rates highly unlikely under any model of rotational evolution. The posterior estimate of P for KIC 6603624 is well-defined in our analysis, but with an age of 7.8 Gyr, its measured rotation of 1.2 days is also highly unlikely under any model of rotational evolution. The LEGACY estimate of rotation is similarly extreme at 378 days. These three stars have the lowest inclination angles in our sample ($< 10^\circ$), at which point the power in the split components of the seismic modes is so low that it becomes difficult to probe the measure of splitting. The split components for these stars will have a height roughly 3% of the central mode. For comparison, for the next lowest inclined star at 17° , this rises to 10%. In cases such as these with lower signal-to-noise, spurious peaks may be interpreted as split components.

We also compared our asteroseismic estimates of stellar rotation with similar studies in the literature, shown in **Extended Data** Figure 8. These included: a study of the binary solar analogues 16 Cyg A & B [11], ; a study of surface and seismic rotation which our catalogue shares 5 stars with [10], ; and an asteroseismic study of differential rotation with which our catalogue shares 40 targets [29]. For the latter, we used their reported splitting value a_1 , which is equivalent to ν_s .

Overall, **Extended Data** Figure 8 shows no strong disagreements between our asteroseismic measurements for stellar rotation and those from the literature. The scatter of the fractional differences lies cleanly around the zero line, with a mean and spread of $0.0^{+16.4}_{-15.6}$ %. The increase in uncertainty with period is due to more slowly rotating stars being more difficult to constrain using asteroseismology.

It is of note that 16 Cyg A was found to be rotating slightly faster in our analysis compared to the literature [11] (deviating within 2σ), despite the fit being performed on the same data. We found an inclination angle that is slightly lower for 16 Cyg A but at a similar projected splitting, which would explain finding a lower value of rotation.

There are three outliers at low period in **Extended Data** Figure 8: KICs 6603624, 8760414 and 8938364. These are the same targets found to be outliers in a comparison to the LEGACY and ‘Kages’ measurements (see above), with anomalously fast rotation rates and low inclination angles. As these stars represent the lowest inclination angles in our sample, and were found to disagree with two independent studies, we opted to exclude them from the gyrochronology analysis, and to flag the rotation measurements for these three stars presented in this work.

3.3 Seismic vs spectroscopic rotation

A distinct difference between rotation rates obtained through different techniques may hold information about differential rotation (both latitudinal and radial) of near-surface layers, such as those we observe in the Sun [63]. A previous comparison of spectroscopic and seismic rotation rates, performed on a sample of 22 stars, found no significant radial differential rotation [29]. In this work they not only considered rotation rates from spots, but also spectroscopic measures of the projected surface rotation $v \sin(i)$, which they found to be more reliable.

With our expanded sample of asteroseismic rotation we can perform a similar analysis, to both validate our sample and probe radial differential rotation. **Extended Data** Figure 9 shows a comparison between spectroscopic $v \sin(i)$ measurements as listed in LEGACY and ‘Kages’ (left) and [31] (right). In these cases the asteroseismic $v \sin(i)$ has been calculated using our measure of asteroseismic $\nu_s \sin(i)$ and the known asteroseismic radii. Three stars (KICs 6603624, 8760414 and 8938364) have been excluded from this figure due to strong disagreements of measured rotation rates with the literature (see above).

For the LEGACY and ‘Kages’ sample, we find no strong deviation from the 1:1 line except at very low velocities, which is likely due to biases inherent to spectroscopic line broadening measurements [64, 65]. For the [31] sample stars lie a lot closer to the 1:1 line.

Overall, there appears to be a global offset where spectroscopic measurements of projected rotation appear faster than asteroseismic measures. Based on the LEGACY and ‘Kages’ $v \sin(i)$ values, this offset is roughly 18% (i.e. spectroscopic projected rotation rates are faster than asteroseismic rates). This offset is much smaller ($\sim 5\%$) for the [31] sample, albeit for far fewer stars. These offsets are within the typical disagreement between spectroscopic methods, based on comparisons of projected rotation measurements for red giant stars [65, see Figure 2], especially at $< 5 \text{ km s}^{-1}$.

It is worth addressing the impact this comparison would have on our conclusions for gyrochronology, were we to take the spectroscopic rotation rates as if they were the truth, even at velocities of $< 5 \text{ km s}^{-1}$. For the majority of cases, the seismic velocities are slower than the spectroscopic velocities, which would bias our ensemble towards favouring a standard evolution over a weakened magnetic braking scenario (as stars would overall appear to be rotating slower at late ages). Based on the comparison to spot rotation presented in the main body of this paper, and given known issues comparing seismic and spectroscopic rotation at low velocities, we conclude that the divergence seen in **Extended Data** Figure 9 does not undermine our results.

3.4 Asteroseismic detection biases

Stars that are rotating very slowly will show small rotational splittings, which may be indistinguishable from a non-rotating case, whereas stars that are spinning very fast may find that their splitting is so wide that they cross over with split modes at different radial orders (n). Asteroseismic selection techniques may also introduce additional biases on a population level. We consider the limitations this places on our asteroseismic method.

- *Fast rotating stars:* In principle, a star could rotate fast enough that the highest-frequency peak in a split dipole ($\ell = 2$) mode would overlap with the radial ($\ell = 0$)

mode. This would occur if $2\nu_s \lesssim \delta\nu_{02}$, the small separation. For the star in our sample with the smallest $\delta\nu_{02}$ (so the highest risk of this happening), this would occur at a rotation of 8.8 days or faster. For the star with the largest $\delta\nu_{02}$, this limit was roughly 2 days.

In practice, the model would still properly fit a spectrum where this kind of mode-overlapping occurred, as the $\ell = 0$ modes don't split, meaning that there are no degenerate solutions in this regime. A more serious issue may be when the $\ell = 1$ mode at lower frequencies overlapped with the $\ell = 2$ mode as well, which would only occur at a rotation rate of 2 days or faster for the star with the smallest $\delta\nu_{01}$.

- *Slow rotating stars:* Except for the most fast rotating cases, asteroseismic mode splitting is detected by a broadening of the mode frequencies as the split components begin to separate from the central mode. If there is to be a hard limit at which this measurement is not possible, it should be in the case that $\nu_s < 2/T$, where T is the observation time, and $1/T$ is the frequency bin width. For the star in our sample with the shortest observing time, this limit would be ~ 89 days, which quickly rises for longer observing baselines. While this constitutes a hard limit, slower rotation rates will become harder and harder to constrain as the rate of splitting becomes smaller. A benefit of the Bayesian approach used in this paper is that it should accurately reflect the increased uncertainty associated with the small rotational splitting.

While these two limits cause concern when studying the most extreme rotators, the results presented in this work are driven by stars with rotation rates in the regime of 10 to 40 days, and so we do not expect these limits to bias our conclusions.

Finally, we also consider bias in asteroseismic detections. The probability of an asteroseismic detection with *Kepler* scales with temperature and radius (i.e. scales inversely with $\log(g)$, and therefore main sequence age) [58, 66]. More active, magnetic stars may also suppress their oscillation modes, making it more likely to detect seismic oscillations in older stars [?]. Through these two effects, asteroseismology is biased towards detecting older over younger stars. As it is older stars, undergoing magnetic braking, that are required to drive the distinction between stellar models, we do not expect this observing bias to effect our conclusions.

4 Verifying consequences for gyrochronology

4.1 Limits of our stellar models

The rotational stellar models used in this work [18] are constructed for metallicities of $-0.4 < [\text{Fe}/\text{H}] < 0.4$ dex, in steps of 0.1 dex. Our sample of 91 stars contained 4 stars with metallicities below -0.4 dex, which are shown as shaded symbols in Figure 3. Of these, 3 were included in our final sample of 73 stars used to evaluate our stellar models: KICs 7970740, 8684723 and KIC 9965715. All three are classed as MS stars, with metallicities of -0.54 ± 0.10 , -0.42 ± 0.10 and -0.44 ± 0.18 dex respectively, placing them within 3σ of the metallicity limits of our stellar models. KICs 8684723 and 9965715 strongly agree with the WMB model,

whereas KIC 7970740 weakly prefers the standard model. Excluding these stars was not found to significantly alter the joint posterior distribution shown in Figure 4.

A recent study [67] compared different rotational evolution models [1, 17, of which we use the former in this work] while studying the effect metallicity has on rotation. They found that metal-rich stars spin down significantly more effectively than metal poor stars. The population of 73 stars used in our stellar model comparisons is roughly centered on a $[\text{Fe}/\text{H}]$ of 0 dex, with a spread of 0.16 dex, with no stars significantly above or below ± 0.4 dex (as discussed above). While differences in stellar rotational evolution as a function of metallicity in this region are still somewhat pronounced, they are much less so than for more metal-rich or poor stars [67, see Figure 2]. However, for stars with $[\text{Fe}/\text{H}] < 0$, the alternative model prescription [17] sees stars spin down more slowly than the models used in this work (i.e. they will rotate faster at later ages). This work only explores the presence of weakened magnetic braking for a single braking law, and a comparison to the [17] braking model will be explored in a future paper.

When constructing KDEs from our stellar model samples (see Section 2), we selected fixed resolutions (or band-widths) for the KDEs. In mass, the band-width of $0.02 M_{\odot}$, was larger than the uncertainties of 24 of 73 stars used to construct our joint posterior. For the stars with the smallest uncertainties, this significantly limits the size of the KDE being evaluated (as subsections of the full stellar models are used to evaluate individual stars, for computational efficiency). In order to confirm that these stars do not significantly affect the ensemble’s preference towards the WMB model, we recalculated the joint posterior distribution for P_s , excluding stars with an uncertainty on mass smaller than the KDE band-width. While the 24 stars with small uncertainties do favour the WMB model, they do so very weakly, whereas the remaining 49 stars with larger uncertainties strongly favour the WMB model. Their removal from the total joint posterior probability does not significantly alter it from the distribution shown in Figure 4.

4.2 Systematic uncertainties from asteroseismology

In our model analysis, we used asteroseismic mass and age obtained using BASTA, as reported in LEGACY and Kages. Asteroseismic properties obtained through stellar models can be subject to systematic errors arising from differences in input physics and choice of stellar models not included in the reported statistical uncertainties. A quantification of these different systematic effects can be found in the ‘Kages’ catalogue paper [8]. Combining their reported median systematic uncertainties due to input physics results in median uncertainties of 20% (up from 14%) on age and 5% (up from 3%) on mass for the ‘Kages’ sample. For LEGACY, the median uncertainties are 18% (up from 10%) and 5.6% (up from 4%) for age and mass respectively.

We re-ran our model analysis after inflating uncertainties on mass and age. We increased uncertainties by the fractional difference between the BASTA statistical uncertainties and the median full statistical and systematic uncertainties described above. For example, a LEGACY star with a mass of $2.0 \pm 0.5 M_{\odot}$ would have its uncertainty inflated by 1.6% of its mass, to $0.53 M_{\odot}$.

To further test the limits of this analysis, we also reran our mixture model fit, this time only shifting the asteroseismic ages younger by the systematic uncertainty, and retaining the

statistical uncertainty (i.e. the ages of LEGACY and ‘Kages’ stars were reduced by 8% and 6% respectively). In this scenario where all asteroseismic ages are overestimated, ‘true’ fast rotators at young ages would have been mistaken for fast rotators at old ages, suggesting the presence of weakened magnetic braking where none existed.

The results of these tests are discussed further in the main body of the paper.

4.3 Binaries and Planet Hosts

For stars to be good probes of existing gyrochronology relations, their rotational evolution must occur in isolation. If a star interacts with a close binary companion (through tides, or a merger) the natural angular momentum loss can be disturbed, causing gyrochronology to mispredict ages [5, 68]. Between LEGACY and Kages, we have 8 known binaries.

First, KIC 8379927, KIC 7510397, KIC 10454113 and KIC 9025370 are spectroscopic binaries. This does not affect the asteroseismic analysis, but may affect their rotational evolution. Of these, KICs 8379927, 7510397 and 9025370 were included in the gyrochronology analysis. None of them preferred one model strongly over the other, with all three finding flat posteriors for P_s .

Second, the binary pairs of KIC 9139151 & 9139163 and 16 Cyg A & B are individually observed binary components with wide orbital separations, so we do not expect their binarity to have affected their rotational evolution.

While we chose not to account for star-planet tidal interactions in this work, we note that this may also disturb the natural stellar rotational evolution when tidal forces are at play [69, 70, 71], although this has been disputed by observations of asteroseismic planet hosts [72].

5 Supplementary Information

5.1 Asteroseismic Model

In order to extract signatures of stellar rotation from the asteroseismic mode frequencies, we built a model that simultaneously treats the convective background ($B(\nu)$), the oscillations ($O(\nu)$), and the white noise (W) in the style of e.g. [11]. Our data, observed with *Kepler*, are also subject to the apodization (attenuation) of signals in the frequency-domain, where we fit our model [58]. The apodization in power is given by

$$\eta^2(\nu) = \text{sinc}^2\left(\frac{\pi}{2} \frac{\nu}{\nu_{\text{nyq}}}\right), \quad (4)$$

where ν_{nyq} is the Nyquist frequency for the *Kepler* short cadence, which was treated as a free parameter in our model to account for gaps in the data. Apodization only affects signals with characteristic timescales, meaning that it does not affect the white noise level, only the oscillations and convective background components. Given the above, our comprehensive model for the power spectrum is

$$M(\nu) = W + \eta^2(\nu)[O(\nu) + B(\nu)]. \quad (5)$$

5.1.1 Convective Background

To model the convective background we used three Harvey components [73], which express the background in power as Lorentzian-like functions centered on zero frequency. The Harvey components take the form

$$H(\nu, a, b, x) = \frac{4a^2/b}{1 + (2\pi b\nu)^x}, \quad (6)$$

where a and b are the free parameters in our model, and x is fixed. The three Harvey components together form our background function as

$$B(\nu) = H(\nu, a, b, x = 4) + H(\nu, c, d, x = 4) + H(\nu, j, k, x = 2), \quad (7)$$

where we have labeled parameters for the separate components. The $x = 2$ term here contributes to the background at high frequencies, whereas the $x = 4$ terms contribute the background at low frequencies.

5.1.2 Modes of Oscillation

Modes of oscillation appear in the power spectrum as Lorentzian peaks [74]. Due to stellar rotation, each mode with an angular degree of $\ell > 0$ is split into its $(2\ell + 1)$ Lorentzian components, labeled by m , the azimuthal order. For all $\ell = (0, 1, 2)$ modes identified in [20, 21] we add a (set of) Lorentzian(s) to our model, building a composite model representing all visible modes. The construction of our oscillation model takes the form

$$O(\nu) = \sum_n \sum_\ell \sum_{m=-\ell}^{\ell} \frac{H_{n,\ell,m}}{1 + \frac{4}{\Gamma_{n,\ell}^2}(\nu - \nu_{n,\ell,m})^2}, \quad (8)$$

where n is the radial order of a mode (i.e. the overtone number of the oscillation), $H_{n,\ell,m}$ is the height of the mode, $\Gamma_{n,\ell}$ is the linewidth of the mode (approximated to be equal for all split components at a single n and ℓ) and $\nu_{n,\ell,m}$ is the frequency of the mode. The range of n differs per star depending on how many radial orders were reported in LEGACY or Kages, and the range of ℓ depends on how many angular degrees were reported for the corresponding radial order.

5.1.3 Mode Frequencies and Rotational Splitting

The mode frequencies of main sequence stars are described by the asymptotic expression [75, 76]. The asymptotic expression defines the locations of the modes as regularly spaced, with structured deviation around ν_{\max} , the frequency of maximum oscillation amplitude. The expression takes the form

$$\nu_{n,\ell,m} = \Delta\nu \left(n + \epsilon + \delta\nu_{0,\ell} + \frac{\alpha}{2} \left(n - \frac{\nu_{\max}}{\Delta\nu} + \epsilon \right)^2 \right) + m\nu_s, \quad (9)$$

where $\Delta\nu$ is the large frequency separation between two consecutive radial orders, ϵ is a phase offset, $\delta\nu_{0,\ell}$ is the small frequency separation between two oscillation modes of different ℓ at the same radial order, α describes the curvature of the spacing around ν_{\max} , and ν_s is the rotational splitting. Note that here we have expressed the small separation $\delta\nu_{0,\ell}$ as a fraction of $\Delta\nu$. In order to improve the computational efficiency of this analysis, we fixed $\Delta\nu$ to the values reported in LEGACY and Kages.

Instead of calculating mode frequencies directly from Equation 9 for the model, we treated the individual mode frequencies as parameters as well, drawn from Equation 9. This is called a ‘latent parameter’ implementation [77, 78], as it forms a step between the parameters we want to draw inference on (also called hyperparameters) and our data. The parameters $\nu_{n,\ell,m}$ were allowed to vary within an uncertainty σ_ℓ , which has a single value for each angular degree and also varied as a free parameter. This allowed us to account for small shifts in frequency due to sudden changes in the stellar structure [79, called acoustic glitches]. The mode frequency latent parameters were drawn from a normal distribution using Equation 9 as its mean function, as

$$\nu_{n,\ell,m} \sim \mathcal{N}(\nu_{n,\ell,m}, \sigma_\ell), \quad (10)$$

where the expression $\nu_{n,\ell,m}$ on the right hand side represents the contents of Equation 9, and \mathcal{N} represents a normal distribution with a mean equal to $\nu_{n,\ell,m}$ and a standard deviation equal to σ_ℓ . The symbol ‘ \sim ’ indicates that the parameters on the left hand side of the equation are drawn from the probability distribution on the right hand side. This notation will be used throughout this paper.

5.1.4 Mode Linewidth

The linewidths of asteroseismic p modes vary roughly as a function of mode frequency, and do so slowly relative to $\Delta\nu$. This can be expressed as an empirical relation [80, 81, 21]. However, this relation has six free parameters, none of which are directly relevant to this work. Instead of fitting this relation, we chose to employ a more flexible Gaussian Process [82, GP] to act as a prior on the linewidths. This can be considered as us modelling the linewidths as correlated measurements. As our GP covariance kernel we used a Squared Exponential Kernel to capture the slight periodicity of linewidth with frequency, as

$$K_{i,j} = \rho^2 \exp \left[-\frac{(n_{f,i} - n_{f,j})^2}{2L^2} \right], \quad (11)$$

where n_f is the fractional radial order of a given mode. For each star, the overtone numbers n were rescaled to be between 0 (for the lowest n) and 1 (for the highest n)^{††}. This approximation was used to describe the change in linewidth as a function of frequency without depending on the exact frequencies of the modes. $K_{i,j}$ represents an element of the covariance matrix \underline{K} , describing the covariance between two values of linewidth at different fractional radial orders. The GP kernel has two hyperparameters: ρ , which determines the spread of the kernel, and L , which determines the length scale in terms of n_f . The length scale was significantly larger than the large frequency separation ($\Delta\nu$) in all cases, and so we considered the use of fractional radial orders a valid approximation in this model.

A linear function was used for the mean of the GP, as

$$\mu = m \times n_f + c, \quad (12)$$

where m and c are the slope and intercept of the line. The linewidth latent parameters were then drawn from the multivariate probability distribution

$$\Gamma \sim \mathcal{N}(\mu, \underline{K}), \quad (13)$$

where Γ represents the linewidths of all the modes in the model. The parameters m , c and ρ were marginalised over, whereas L was fixed to a pre-determined value.

5.1.5 Mode Heights and Angle of Inclination

The height in power of each mode, $H_{n,\ell,m}$, varies not only as a function of distance in frequency from ν_{\max} , but also due to observation conditions, such as inclination angle and passband. In our model, we treated $H_{n,\ell,m}$ as a deterministic parameter, as

$$H_{n,\ell,m} = \varepsilon_{\ell,m}(i) \frac{2(A_{n,\ell})^2}{\pi\Gamma_{n,\ell}}, \quad (14)$$

where $\varepsilon_{\ell,m}(i)$ modulates the height as a function of inclination angle (i) (see below), and $A_{n,\ell}$ and $\Gamma_{n,\ell}$ are the mode amplitude and linewidth respectively for a given radial order and angular degree. Instead of modeling and modulating height directly, we instead sampled

^{††}Overtone of $\ell = 2$ modes were increased by 1 to ensure this approximation applied to all modes.

in amplitude and linewidth. This approach mitigates the correlations between height and linewidth in the sampling process [83].

As done above for the mode frequencies and linewidths, the mode amplitudes $A_{n,\ell}$ were also treated as latent parameters drawn from a probability distribution governed by hyperparameters. We used a Gaussian function $G(\nu)$, centered on ν_{\max} , so that

$$G(\nu) = A \times \exp \left[-\frac{(\nu - \nu_{\max})^2}{2w^2} \right], \quad (15)$$

where A is the amplitude at ν_{\max} , and w is the width of the Gaussian, both free parameters in our model. The mode amplitude latent parameters were then drawn from the probability distribution

$$A_{n,\ell} \sim \mathcal{N}(G(\nu_{n,\ell}) \times V_\ell, \sigma_A), \quad (16)$$

where V_ℓ is a free parameter for the mode visibility of different angular degrees, which should be consistent for all *Kepler* observations. The mode visibility for V_0 is fixed at 1, and $V_{1,2}$ are treated as free parameters. The parameter σ_A , the uncertainty on the distribution, is also a free parameter, and takes the same value for all amplitudes regardless of angular degree.

The angle of inclination of a star with respect to Earth changes the net perturbation by a given mode when integrated across the stellar disc, changing the amplitudes of modes of different azimuthal orders. This is a geometric problem, and is expressed by $\varepsilon_{\ell,m}(i)$, which takes the form [84]

$$\varepsilon_{\ell,m}(i) = \frac{(\ell - |m|)!}{(\ell + |m|)!} \left[P_\ell^{|m|}(\cos(i)) \right]^2, \quad (17)$$

where $P_\ell^{|m|}$ are associated Legendre functions. For the first three angular degrees, Equation 17 takes the form [85]

$$\begin{aligned} \varepsilon_{0,0}(i) &= 1, \\ \varepsilon_{1,0}(i) &= \cos^2(i), \\ \varepsilon_{1,\pm 1}(i) &= \frac{1}{2} \sin^2(i), \\ \varepsilon_{2,0}(i) &= \frac{1}{4} (3 \cos^2(i) - 1)^2, \\ \varepsilon_{2,\pm 1}(i) &= \frac{3}{8} \sin^2(2i), \\ \varepsilon_{2,\pm 2}(i) &= \frac{3}{8} \sin^4(i), \end{aligned} \quad (18)$$

where the sum of available components for a single ℓ are normalized to one.

5.1.6 Likelihood Function

If data have Gaussian noise in the time domain, they will appear in the frequency domain with noise following a χ^2 distribution with two degrees of freedom [?, χ_2^2 hereafter]]woodard1984.

The noise properties of χ^2_2 distributed data are multiplicative, and require a specific treatment when fitting a model. As our frequency bins can be approximated to be independent, we used the likelihood function [86],

$$\ln p(\mathcal{P}|M) = \sum_{j=0}^{N-1} \left[\ln[M_j(\nu)] + \frac{\mathcal{P}_j}{M_j(\nu)} \right], \quad (19)$$

where \mathcal{P} is the power spectral density (and thus our data), and $M(\nu)$ represents our model. The subscript j denotes an individual datum, for a total of N data. We have omitted the dependence of the model $M(\nu)$ on its parameters, for clarity. This equation is functionally equivalent to the evaluation of a gamma distribution of the form $\gamma(\mathcal{P}|1, \beta)$, where $\beta = 1/M(\nu)$, which is the implementation we used in the sampling process.

5.1.7 Fitting the convective background

Fitting the convective background, apodization and white noise component must be done for the full range of the power spectrum in order to be accurately constrained. However, fitting a single model to the full range of frequencies is computationally inefficient when we are interested in the modes of oscillation, as these occupy a relatively small range of frequencies.

In order to speed up this process, the background was first fit independently to a subset of our data for each star. This subset was created by removing all frequencies within a range $0.1 \times \Delta\nu$ below and above the minimum and maximum mode frequencies reported in LEGACY and Kages. For KIC 3427720 we also removed frequencies in the range $90 \mu\text{Hz} < \nu < 400 \mu\text{Hz}$, where there were large peaks not of asteroseismic origin, skewing the background fit.

For each star the model function (see Eq. 5) was fit, as

$$M_B(\nu) = W + \eta^2(\nu)B(\nu), \quad (20)$$

where $B(\nu)$ is the background model described in Equation 7. The parameter components of our background fit are then $\phi_B = \{\log(a), \log(b), \log(c), \log(d), \log(j), \log(k), W, \nu_{\text{nyq}}\}$, where the parameters of the Harvey components were sampled in log space. The model was fit to the background data using PyStan [87], run for 10,000 iterations on each star^{‡‡}.

5.1.8 Obtaining First Guesses and Prior Values

In order to utilise some of the prior measurements of our targets without using them as hard constraints on our parameters, some of our model equations were fit to LEGACY and Kages data to obtain first guesses and mean values on hyperparameter priors.

For first guesses for parameters in the asymptotic expression, we fit Equation 9, *not* including the rotational component $m\nu_s$, to the $\ell = (0, 1, 2)$ mode frequencies reported in LEGACY and Kages for each star, using their reported uncertainties. This yielded estimates of $\hat{\epsilon}$, $\hat{\delta\nu}_{01}$, $\hat{\delta\nu}_{02}$ and $\hat{\alpha}$, where the hat symbol ‘ $\hat{}$ ’ indicates a prior value (e.g. $\hat{\nu}_{\text{max}}$ is taken from LEGACY or Kages). While not precise, as we did not mitigate any perturbations due to acoustic glitches (sudden changes in interior sound speed, which the latent parameter

^{‡‡}These PyStan runs were initiated with a random seed of 11, as were all other random processes in this work.

treatment takes care of in the full model), these rough results act as functional first guesses and prior mean values. The relation was fit to each star using PyMC3 [40] using 5000 iterations on 4 chains.

To obtain first guesses for the parameters used to set the GP prior on linewidth, we fit a GP constructed as in Equation 13 to the linewidths of the $\ell = 0$ modes reported in LEGACY. Linewidths were not reported for the other angular degrees in LEGACY, but the estimates may be generalised to other ℓ , as linewidth is a strong function of frequency. The relation was fit to each star using PyMC3 using 2500 iterations on 4 chains.

Fitting the LEGACY linewidths yielded rough estimates of \hat{m} , \hat{c} , $\hat{\rho}$ and L for each star. As is noted in Equation 13, L was fixed to this fit value when fitting our full model to our data. For stars in Kages, for which no linewidths were reported, we instead fixed these prior values to $\hat{m} = 1$, $\hat{c} = 0.5$, $\hat{\rho} = 0.1$, and the length scale to $L = 0.3$. These values were chosen to reflect those found for the LEGACY stars.

Finally, we obtained prior values for the Gaussian function describing the distribution of mode amplitudes around ν_{\max} . The mode amplitude of the highest peak in the spectrum was used for the amplitude, which was typically at or near ν_{\max} . For the width of the Gaussian we used the empirical function [21]

$$\hat{w} = 0.25 \times \nu_{\max}. \quad (21)$$

For the mode visibilities, we used $\hat{V}_1 = 1.2$ and $\hat{V}_2 = 0.7$, which reflect the results for these parameters reported in the LEGACY catalogue.

5.1.9 Priors on our Hyperparameters

Given our first guesses and measured prior values, we can define the prior probabilities of the hyperparameters on which our model depends. For the mode frequency hyperparameters, these are

$$\begin{aligned} \nu_{\max} &\sim \mathcal{N}(\hat{\nu}_{\max}, 10), \\ \epsilon &\sim \mathcal{N}(\hat{\epsilon}, 1), \\ \alpha &\sim \ln \mathcal{N}(\ln(\hat{\alpha}), 0.01), \\ \delta\nu_{01} &\sim \ln \mathcal{N}(\ln(\widehat{\delta\nu_{01}}), 0.1), \\ \delta\nu_{02} &\sim \ln \mathcal{N}(\ln(\widehat{\delta\nu_{02}}), 0.1), \\ \sigma_{0,1,2} &\sim \mathcal{C}_{1/2}(\beta = 2), \end{aligned} \quad (22)$$

where $\ln \mathcal{N}$ represents a log-Normal distribution and $\mathcal{C}_{1/2}$ represents a half-Cauchy distribution, and other symbols are as described above. All three hyperparameters $\sigma_{0,1,2}$ describing the uncertainty on the latent parameters of different angular degree were subject to the same prior.

For the mode linewidths, our hyperparameter priors took the form

$$\begin{aligned} m &\sim \mathcal{N}(\hat{m}, 1), \\ c &\sim \mathcal{N}(\hat{c}, 1), \\ \rho &\sim \ln \mathcal{N}(\ln(\hat{\rho}), 0.1), \end{aligned} \quad (23)$$

where the conventions are the same as above. For our mode amplitudes, they took the form

$$\begin{aligned}
w &\sim \ln \mathcal{N}(\ln(\hat{w}), 10), \\
A &\sim \ln \mathcal{N}(\ln(\hat{A}), 1), \\
V_1 &\sim \ln \mathcal{N}(\ln(\hat{V}_1), 0.1), \\
V_2 &\sim \ln \mathcal{N}(\ln(\hat{V}_2), 0.1), \\
\sigma_A &\sim \mathcal{C}_{1/2}(\beta = 1).
\end{aligned} \tag{24}$$

As the convective background had already been fit to our data excluding the region where the modes are present, the results from that fit could be used as extremely informative priors on our fit to the region containing the modes, where there is little information present to constrain the background. To do so, we modeled the background parameters ϕ_B in our full model as being drawn from a multivariate normal distribution as

$$\phi_B \sim \mathcal{N}(\hat{\phi}_B, \underline{\Sigma}_{\hat{\phi}_B}), \tag{25}$$

where $\hat{\phi}_B$ are the median values of our posterior distributions from our prior background fit, and $\underline{\Sigma}_{\hat{\phi}_B}$ is the full covariance matrix of all the posterior distributions from our prior background fit, taking into account the correlations between the different Harvey components.

Finally, we defined the priors on the rotational parameters: the mode splitting (ν_s), and the inclination angle (i). In order to give these an appropriate treatment, we made two reparameterizations. First, we sampled the projected rotational splitting, $\nu_s \sin(i)$, which is more efficiently sampled due to the strong correlations between i and ν_s [88, 89]. A prior was applied over this as

$$\nu_s \sin(i) \sim \ln \mathcal{N}(\ln(0.75), 0.75), \tag{26}$$

where conventions are as above. This subjective prior was chosen to reflect that most stars will have a solar-like rotation, with a long tail to allow for the fastest rotators. Second, we sampled in $\cos(i)$, and gave this a prior of

$$\cos(i) \sim \mathcal{U}(0, 1), \tag{27}$$

which is equivalent to stating that probability to observe an inclination angle i is equal to $\sin(i)$. Here, the $\mathcal{U}(0, 1)$ indicates a uniform prior between 0 and 1. Using a uniform prior on $\cos(i)$ allowed us to account for the geometric effect that stars with a large inclination angle with respect to us are more common [74].

5.1.10 Fitting procedure

Using our prior information and model described above, we fit Equation 5 to our data \mathcal{P} , using the likelihood function described in Section 5.1.6.

In order to speed up the fitting process, we only applied our model to the region of the power spectrum that contains visible modes of oscillation. We created this subset by removing

all frequencies outside a range $0.25 \times \Delta\nu$ below and above the minimum and maximum mode frequencies reported in LEGACY and Kages. This region overlaps minimally with the data used to fit for our prior information of the convective background (see Section 5.1.7), so that both model fits are independent.

To improve computational efficiency, we reduced the number of oscillation modes being fit in five targets. For 16 Cyg A & B, KIC 7970740 and KIC 8478994, we excluded any modes with a Bayes Factor ($\ln(K)$) of less than 6, as reported in LEGACY [90, 20, 21]. For KIC 8478994, which is reported without a value for $\ln(K)$ in Kages, we only included modes of an overtone number that contained a detection for all of $\ell = (0, 1, 2)$, retaining 5 sets of higher signal-to-noise overtones. We do not expect this reduced scope to bias our results, although they may reduce the precision on our measured rotation rates.

We fit our model to our power spectrum data with PyMC3, using 2500 iterations each on 4 chains. An example of our model fit to an asteroseismic power spectrum of 16 Cyg A is shown in Figure 5.

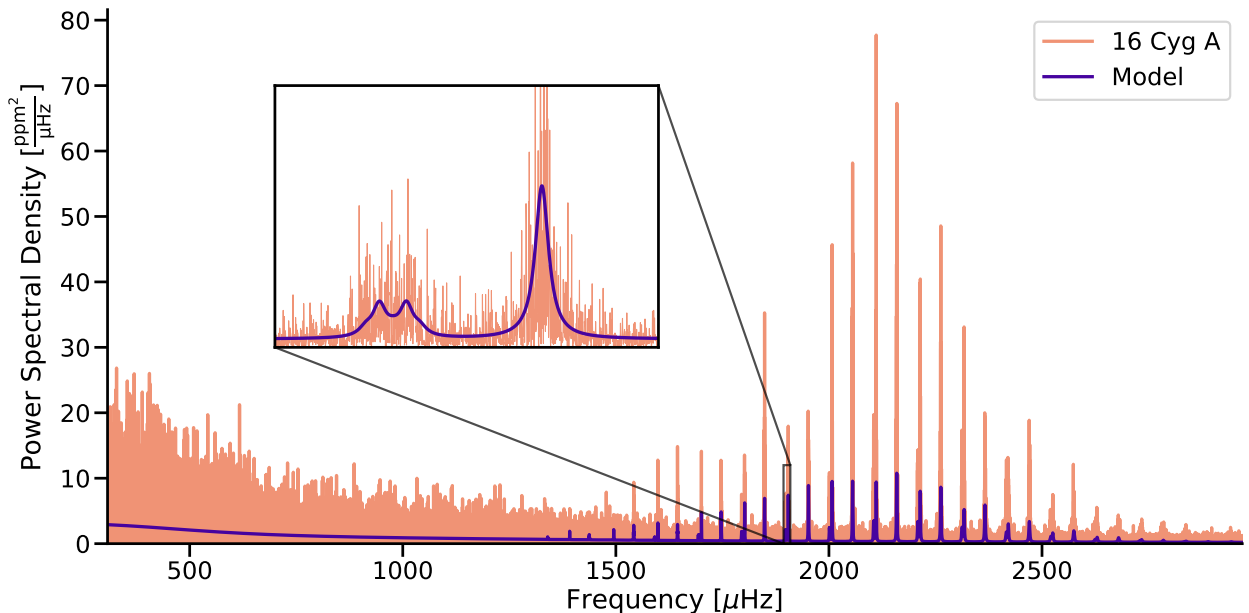


Figure 5: A power spectrum constructed from four years of *Kepler* observations of 16 Cyg A (KIC 12069424). Plotted over the top is the model resulting from the fit to the data described in this work. The model implements both the mode frequencies, seen on the right hand side of the plot, and the convective background, the effects of which are seen on the left. Low frequencies have been cropped out for clarity. *Inset*: A zoom in on a radial (right) and quadrupole (left) ($\ell = 0, 2$) pair of modes. The quadrupole mode is split into five components by the star’s rotation. Due to the star’s inclination angle with respect to us, two out of five peaks are more distinct. The height and spacing of the mode components is a function of the star’s rotational splitting ($0.56 \mu\text{Hz}$, equivalent to $P_{\text{rot}} = 20.5$ days) and angle of inclination (45°).

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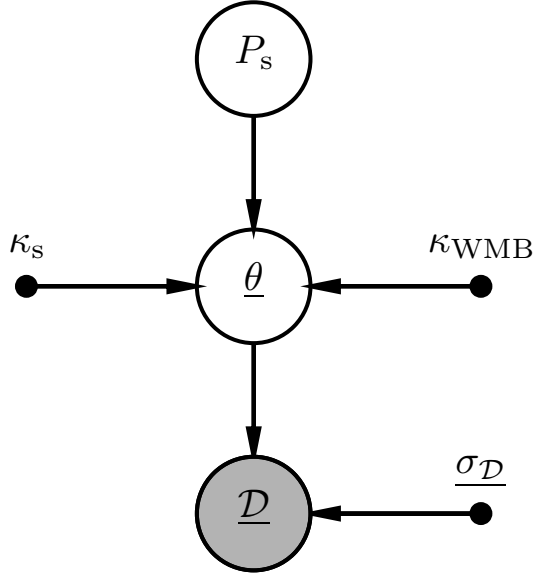


Figure 6: A probabilistic graphical model (PGM) represented algebraically in Equation 1. The shaded circle indicates observed data, and solid black points represent other fixed information, such as the KDEs and observational uncertainties. The remaining circles represent parameters. The underline indicates that the symbol represents a set of parameters or data. Here, κ_s and κ_{WMB} represent the KDEs of standard and WMB model populations respectively. P_s is the mixture model weighting factor. The latent parameters $\underline{\theta}$, our observations $\underline{\mathcal{D}}$ and their uncertainties $\underline{\sigma_{\mathcal{D}}}$ include temperature (T_{eff}), mass (M), log-age ($\ln(t)$), metallicity ($[\text{Fe}/\text{H}]$) and log-rotation ($\ln(P)$). This model is *hierarchical*, as all the latent parameters are drawn from the common probability distribution set by P_s and described in Equation 2.

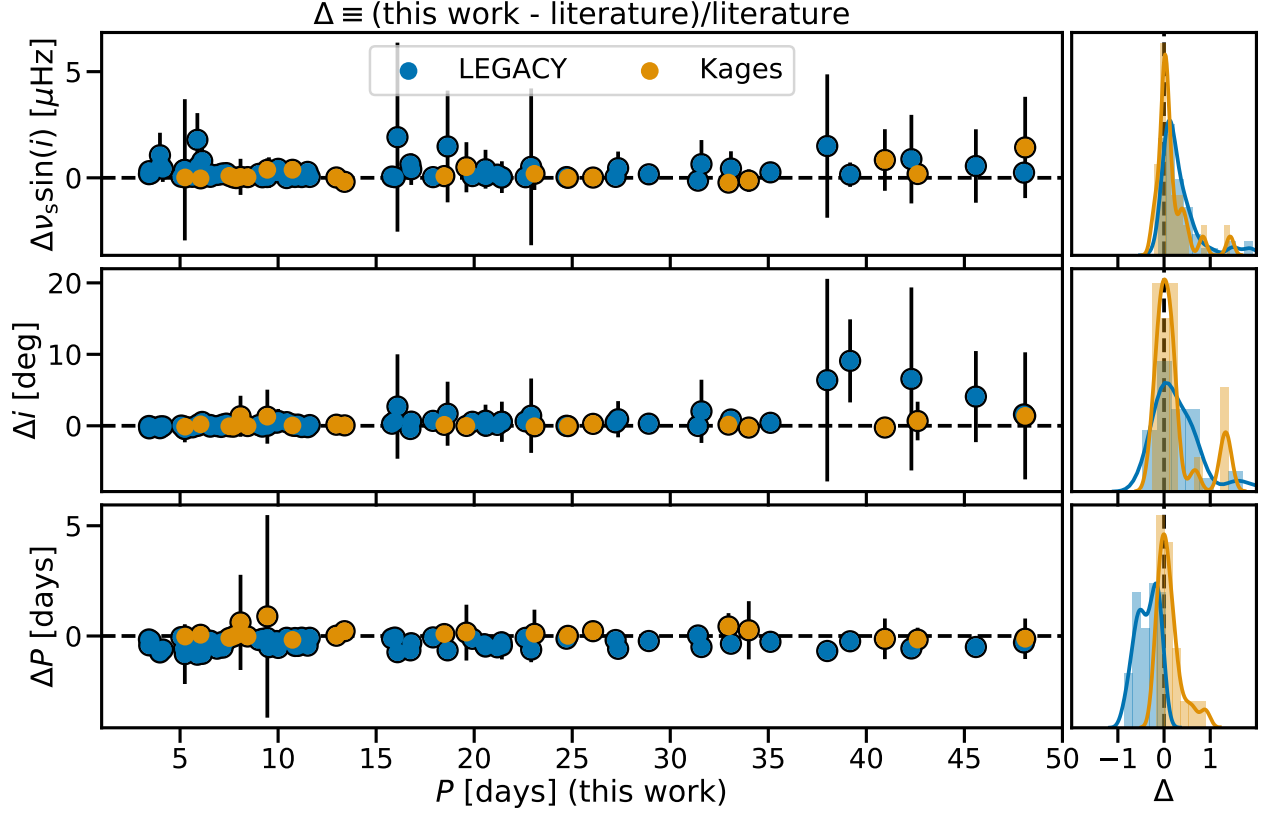


Figure 7: Comparisons between posterior estimates of rotational parameters from this work, LEGACY and Kages [20, 21, private communication]. Fractional differences are plotted against stellar rotation obtained in this work. The Δ indicates the fractional difference between this work and the literature (i.e. stars above the zero-line have higher values in this work). The right hand panels show the distribution of the fractional differences around the zero line. The colour legend is consistent throughout all panels. The x-axis units on the right hand panels are equivalent to the y-axis of the left hand panels. 10 stars have been omitted from this plot and are discussed in more detail in the text: KICs 5094751, 6196457, 8349582, 8494142, 8554498, 105114430 and 11133306 all have extremely low rotation periods in Kages, with high uncertainties. Conversely, KICs 6603624, 8760414 and 8938364 have extremely high rotation periods in LEGACY with low uncertainties. In cases where stars had asymmetric error bars, the larger of the two was used when propagating uncertainty for the purposes of this figure.

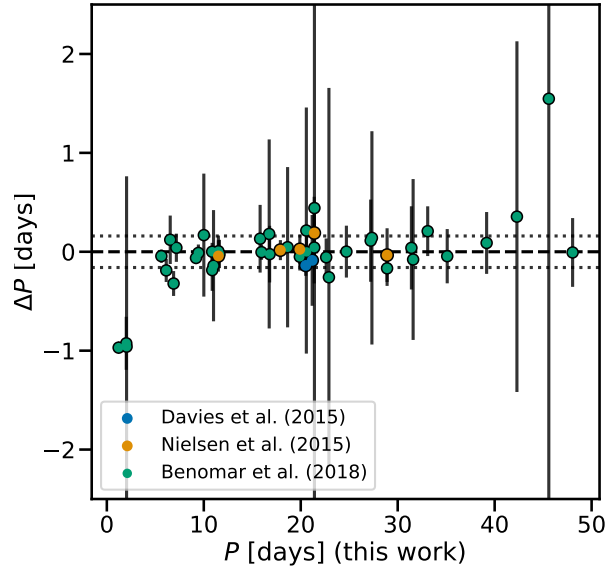


Figure 8: Fractional differences between posterior estimates of asteroseismic rotation period from this work. Literature sources are: [11] (16 Cyg A & B), [10] (5 stars) and [29] (40 stars). We used the reported parameter a_1 from [29], which represents the rotational splitting in the case of uniform latitudinal rotation in their model. The dashed line represents the median of the sample shown, with the dotted lines representing the 15.9th and 84.1st percentiles. In cases where stars had asymmetric error bars, the larger of the two was used when propagating uncertainty for the purposes of this figure.

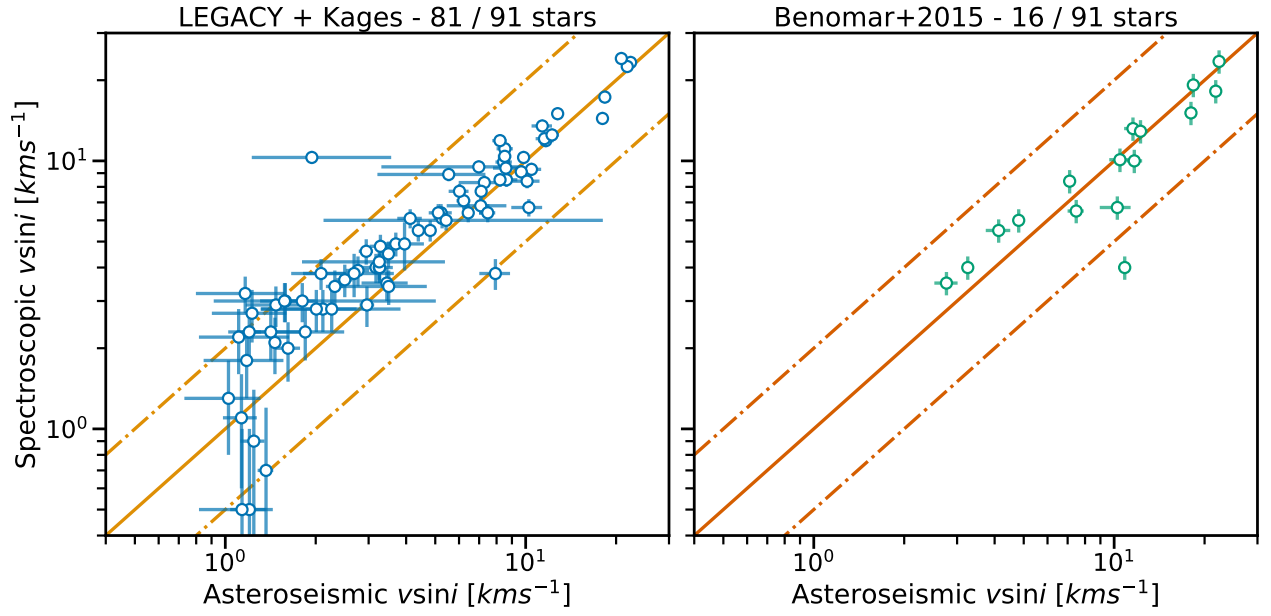


Figure 9: Comparisons between asteroseismic and spectroscopic measures of projected surface rotation, $v \sin(i)$. All asteroseismic (x-axis) values are from this work, all spectroscopic (y-axis) values are from the literature. *Left*: comparisons to 81 stars values reported in LEGACY and Kages. *Right*: comparisons to 16 stars observed by [31]. Asteroseismic values are transformed from projected splitting ($\nu_s \sin(i)$) using the asteroseismic radius measurements presented in LEGACY and Kages. The solid lines indicate the 1:1 line, while the dash-dotted lines represent the 2:1 and 1:2 lines.

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