

# ASTROPHYSICS AND GRAVITY DIVISION

## Overview

Astrophysics and Gravity Division consists of Gravitational Wave Group, The Observational Cosmology Group, Primary Cosmic Ray Group and Theory Group.

The Gravitational Wave Group conducts experimental research of gravitational wave with researchers of gravitational wave experiment and theory in Japan. The main items are the construction of the large scale cryogenic interferometer(KAGRA) at Kamioka underground and the operation of CLIO. For this purpose, KAGRA observatory was established at the beginning of the fiscal year of 2016 to assist the construction of KAGRA gravitational wave telescope.

The Observational Cosmology Group studies the cosmic history based on deep multi-wavelength observations in collaboration with worldwide researchers. This group has started a new optical deep survey project with the wide-field imager of Hyper Suprime-Cam mounted on the Subaru telescope.

Theory Group conducts both theoretical study of the Universe and astroparticle physics.

## Gravitational Wave Group

### KAGRA Project Status

[Spokesperson : Takashi UCHIYAMA]

ICRR, The Univ. of Tokyo, Hida, Gifu 506-1205

### Overview

KAGRA, Large-scale Cryogenic Gravitational wave Telescope, aims at detecting gravitational waves and developing gravitational wave astronomy, which was established by the first detection of gravitational waves by LIGO. KAGRA employs a 3 km L-shaped laser interferometer with a cryogenic mirror system placed underground at Kamioka[1]. The KAGRA development is divided into two stages: the initial KAGRA (iKAGRA) and baseline KAGRA (bKAGRA). The iKAGRA interferometer is a simple Michelson interferometer with a 2-Watt laser, room-temperature mirrors, and a simple vibration isolation system. We completed the iKAGRA interferometer with a test run in April 2016[2]. Then we proceeded to bKAGRA.

Figure 1 and 2 show a schematic view of optical layout of the bKAGRA interferometer and the KAGRA vibration isolation systems. Table 1 shows design parameter of the bKAGRA interferometer[4]. The bKAGRA interferometer will employ a Resonant Sideband Extraction (RSE) interferometer with 180-Watt laser, cryogenic Sapphire mirrors, and several kinds of vibration isolation systems. The bKAGRA interferometer should attain the sensitivity high enough for the detection of gravitational waves with the help of the high power

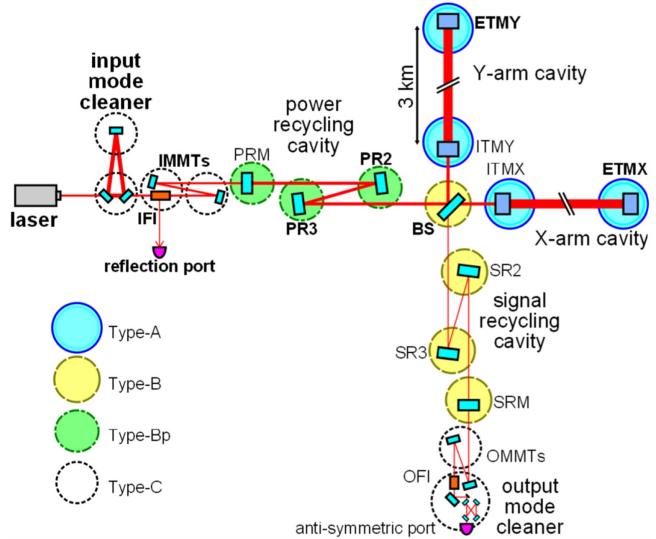


Fig. 1. Schematic view of the bKAGRA interferometer[4]. Type-A, Type-B, Type-Bp, and Type-C are the names of vibration isolation system for each mirror.

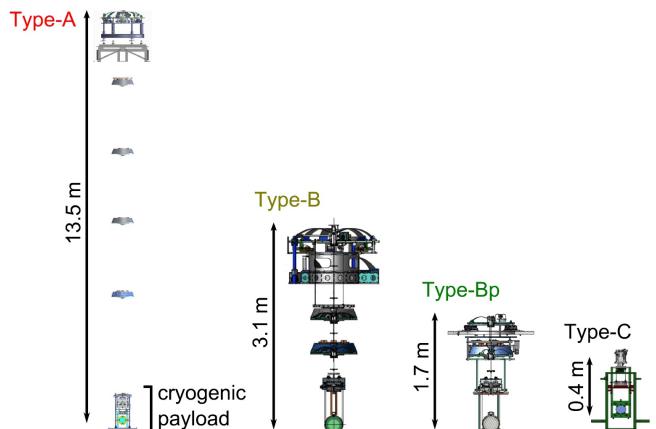


Fig. 2. KAGRA vibration isolation systems[4]. KAGRA equips four kinds of vibration isolation systems such as Type-A, Type-B, Type-Bp, and Type-C.

laser and RSE interferometer to reduce the quantum noise, the cryogenic Sapphire mirrors to reduce the thermal noise, and the vibration isolation systems to reduce the seismic noise. Figure 3 shows designed sensitivities of bKAGRA in case of Broadband RSE (BRSE) and of Detuned RSE (DRSE), where incoherent sum of the fundamental noise sources is assumed. Observation range for an in-spiral and merger of neutron-star binary reaches 135 Mpc in BRSE and 153 Mpc in DRSE with the same definition of the observation range as LIGO and Virgo.

Figure 4 shows the international collaborative observation scenario[3]. LIGO conducted Observation 1 (O1) from

Table 1. The design parameters of the bKAGRA interferometer[4].

Arm cavity length	3000 m	Test mass size	$\phi 22\text{ cm} \times 15\text{ cm}$
Laser wave length	1064 nm	Mass of test mass	22.8 kg
Input power at PRM	67W	Temperature of test mass	22 K
Arm intra-cavity power	340 kW	Beam radius at test mass	3.5 cm
ITM transmittance	0.4 %	PRC/SRC lengths	66.6 m
PRM transmittance	10 %	Detuning angle	3.5 deg
SRM transmittance	15 %	Homodyne angle	135.1 deg

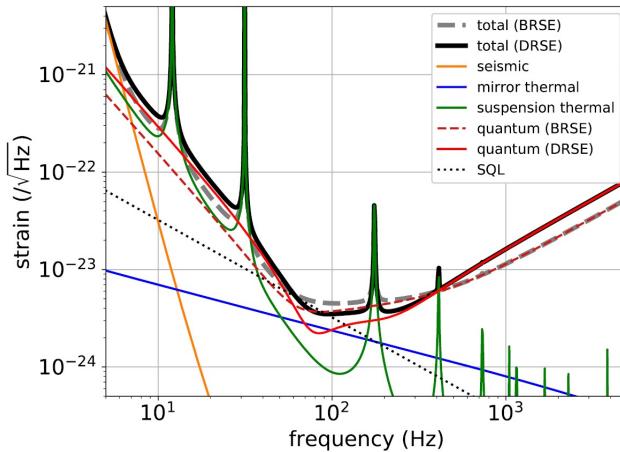


Fig. 3. The designed sensitivity of the bKAGRA interferometer[4]. "total", "seismic", "mirror thermal", "suspension thermal", "quantum", and "SQL" mean total sum of fundamental noise sources shown in this figure, seismic noise including gravity gradient noise, mirror thermal noise, suspension thermal noise, quantum noise, and standard quantum limit, respectively. The figure shows "total" and "quantum noise" in both Broadband RSE (BRSE) and Detuned RSE (DRSE) case. Observation range for an in-spiral and merger of neutron-star binary reaches 135 Mpc in BRSE and 153 Mpc in DRSE with the same definition of the observation range as LIGO and Virgo.

September 12th, 2015 to January 19th, 2016 and Observation 2 (O2) from November 30th, 2016 to August 25th, 2017. Virgo joined O2 from August 1st, 2017. LIGO and Virgo started Observation 3 (O3) from April 1st, 2019 and O3 will continue by the end of April in 2020. KAGRA is aiming to join O3 in 2019.

In FY2018 we started with an operation of KAGRA interferometer as bKAGRA phase 1 which is 3 km Michelson interferometer with two sapphire mirrors suspended by the Type-A vibration isolation systems. One sapphire mirror was cooled at 18 K. The operation was done from April 28 to May 6 in 2018 and it was the first demonstration of operating km-class interferometer at cryogenic temperature. Figure 5 and Figure 6 shows a summary of daily status of the operation and a strain sensitivity comparing with noise sources[4], respectively. Duty factor in the first half of the operation reached 88.6%. The observation range for an in-spiral and merger of neutron-star binary and BH binary reached 17 pc and 100 pc, respectively. The longest continuous operation time was 11.1 hour.

After the bKAGRA phase 1 operation, we started con-

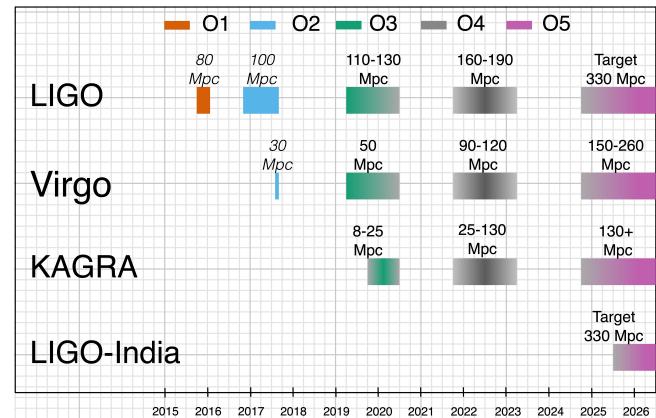


Fig. 4. International observation scenario[3]. Virgo was joined in Observation 2 (O2) from August 1st in 2017. LIGO and Virgo started Observation 3 (O3) from April 1st in 2019. O3 will continue by the end of April in 2020. KAGRA is aiming to join O3 in 2019.

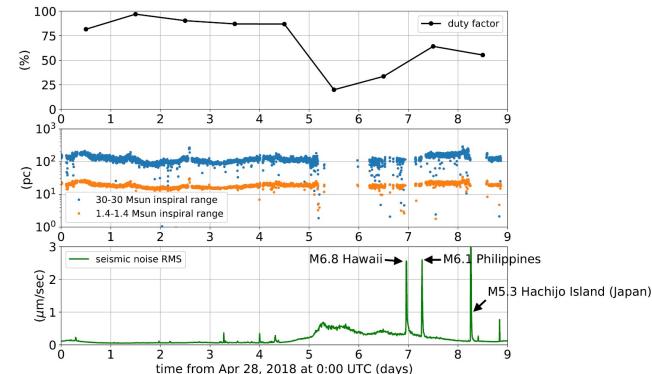


Fig. 5. Daily status of bKAGRA phase 1 operation[4]. The figure shows daily duty factor (Top panel), inspiral range (Middle panel), and seismic noise level (Bottom panel) during the operation. The operation was done from April 28 to May 6 in 2018.

struction of the bKAGRA interferometer with 40 W laser power. What we have installed were an infrared laser system with the maximum power of 40 W, two sets of arm length stabilization system using a green laser, calibration systems using photon radiation pressure, large beam baffles, transmission monitor systems, some optics consisting a signal recycling cavity and output optics, two input test masses called ITMX and ITMY in Figure 1, and so on. Physical Environmental Monitor (PEM) is a sensor network consisting several kinds of environmental sensors such as accelerometers,

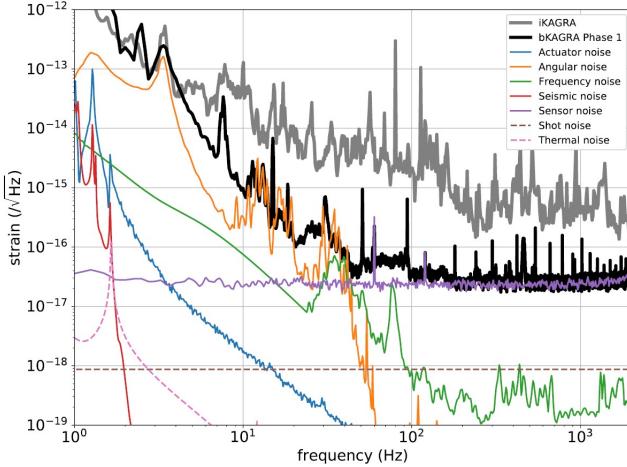


Fig. 6. Strain sensitivity of KAGRA in phase 1 operation[4].

seismometers, magnetometers, thermometers, acoustic sound monitors, power monitors and so on. Purpose of PEM is to check the detector health, noise sources, and data quality in cooperation with the detector characterization group. We placed many sensors in the KAGRA site and monitoring has already started.

We have tried lock acquisition of the X-arm cavity for the first time in parallel with the installation works mentioned above. The lock acquisition of the X-arm cavity was successfully achieved with helps of the arm length stabilization system. Then we carried out characterization of the X-arm cavity. Table2 shows a summary of optical parameters of the X-arm cavity comparing with designed and measured values.

We also enhanced the international collaborations with the Einstein Telescope (ET) project, LIGO, Virgo, Korean and other Asian groups mainly based on the JSPS core-to-core program.

The rapidly progressing status of KAGRA were presented in many international conferences. Many papers about the progress of KAGRA were also published [1], [4], [5]. We also presented activities on our web-page.[6]

## Bibliography

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## Vacuum system for KAGRA

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We firstly pumped six vacuum chambers those are IFI, IMM, PRM, PR3, PR2, and BS and vacuum pressure achieved order of  $10^{-5}$  Pa. We made control boxes for seven large gate valves those are GVmc between MCF and IFI, GVbsx between BS and IXC, GVitm between IXA and the X-arm tube, GVitm between the X-arm and EXA, GVbsy between BS and IYC, GVitm between IYA and the Y-arm tube, and GVitm between the Y-arm and EYA. The control boxes has open/close control buttons and indicators. The GVs will be controlled by remotely via KAGRA machine control system which has Programable Logic Circuits through the control boxes.

## Input and Output Optics

[Spokesperson : Keiko KOKEYAMA]

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The input and output optics of KAGRA consists of the pre-stabilization system for the laser, auxiliary locking system, input optics chain, output optics chain, and detectors. The pre-stabilization system includes the frequency stabilization system, intensity stabilization system, pre-mode cleaner, and modulation system for the main interferometer. The auxiliary locking system includes the phase locking system for the green beam, the fiber system, and the locking system for the arm cavity. The input optics chain includes the input mode cleaner, input Faraday isolator, and input mode matching telescope. The output optics chain includes the output mode matching telescope, output Faraday isolator, and output mode cleaner. The detectors are for the symmetric, antisymmetric, and pick-off ports.

In the fiscal year 2017, major developments towards the bKAGRA phase-1 operation, planned in the spring of 2018, were conducted. The reference cavity was installed on the optical table where the laser source is placed. The reference cavity consists of a cylindrical spacer made of a ultra-low expansion glass, and two mirrors attached on the both side of the cylinder. The length of the cavity is extremely stable, therefore works as a frequency reference for the laser. The frequency of the laser was successfully stabilized, with an acousto-optic modulator installed as the frequency actuator. The input mode cleaner, which had been operating since the iKAGRA operation, was re-commissioned as a part of the frequency stabilization system. At the low frequency, the cavity length of the input mode cleaner is controlled by actuating one of the three mode-cleaner mirrors, whereas the high frequency feedback is applied to the acousto-optic modulator for the frequency stabilization. The cross over frequency of the

Table 2. Optical parameters of the X-arm cavity[5].

Parameter name	Designed	Measured
Cavity length	3000 m	29999.990(2) m
Finesse for 1064 nm	1530	1410(30)
Roundtrip loss for 1064 nm ± 100 ppm	86(3) ppm	
Finesse for 532 nm	49.2	41.0(3)

control was 3 Hz. In addition, the operation procedure was automated. For the intensity stabilization system, the main optical components were installed in the laser room, and the loop was closed. The performance of the intensity stabilization did not reach the goal stability yet, and to be improved in the next fiscal year.

There are two input mode-matching telescopes downstream of the input mode cleaner, where is upstream of the main interferometer. The input mode-matching telescopes are the curved mirrors, hung by the double pendulum, for matching the spacial mode of the beam to the main interferometer mode. The installation of the suspensions had been finished in the previous fiscal year, and in 2017, the fused silica mirrors (diameter 100 mm) were hung from the suspensions. The beam profile after the second input mode-matching telescope was measured, and the beam shape was confirmed to be as designed.

In November, KAGRA hosted the commissioning workshop in Kamioka. Researchers working at the LIGO and VIRGO sites visited the KAGRA site to accelerate the KAGRA experiments. The input and output optics group provided the topics to study in the workshop. Particularly, the noise couplings on the input mode cleaner and laser system were investigated, and useful findings and suggestions were obtained.

The interferometric tilt sensors, developed by Sogang university and tested at the KAGRA site in 2016, was upgraded. The upgraded tilt sensor can sense the orthogonal two degrees of freedom, pitch and yaw motions of a target. The performance was compared with the currently-used sensors (optical levers), and found to be as good in the lower frequencies, and to be better at the frequencies higher than 10 Hz. The tilt sensor was installed permanently at one of the mirrors of the input mode clear.

### Cryogenic system

[Spokesperson : Takafumi USHIBA]

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One of unique features of KAGRA is cooling sapphire mirrors, which are installed at 3-km arm cavity. Members working for KAGRA cryogenic system, which plays an important role for achieving this unique characteristics, are mainly constituted by that of ICRR, KEK, and the University of Toyama. Here, we summarize the activity of members of ICRR in FY 2019.

**Cryogenic payload** A KAGRA sapphire mirror is suspended by 9-stage suspension and its bottom 4 stages that include sapphire mirror is called cryogenic payload. Angular and

translational motion of these sapphire mirrors must be controlled well for operating them as an interferometer. Therefore, damping and global control of suspension were implemented and improved in this fiscal year.

One large improvement of the suspension control is implementation of hierarchical control of suspension. After the implementation of this hierarchical control, angular motion of sapphire mirror became very small: less than 100 nrad in RMS, which is small enough to operate the interferometer. Thanks to this improvement, we succeeded to operate a power-recycling Fabry-Perot Michelson interferometer with the cryogenic payloads.

**Heatlink vibration isolation system** KAGRA cryogenic payload has heatlinks, which are directly connected to cryocoolers, for cooling sapphire mirror to 20 K. Since cryocooler makes large vibration compared to the ground motion at Kamioka site, the vibration through these heatlinks can be contaminate the detector sensitivity of KAGRA. So, we developed heatlink vibration isolation system (HLVIS) to mitigate the vibration through heatlinks.

HLVIS consists of three-stage pendulum. Each stage is suspended by four tension springs in order to mitigate not only horizontal vibration but also vertical vibration. Vibration isolation ratio was measured below 10 Hz and measurement result was almost matched with designed value.

**New design of mirror inclination control system** A cryogenic payload has an inclination adjustment system called moving mass system. the moving mass system consists of three components: cryogenic compatible stepper motor, oil-free ball screw, and copper block. We can drive a copper block by rotating stepper motor and change the mass balance of the cryogenic payload to tilt the mirror.

Basically, this moving mass system works well but in terms of long-term stability, there is an issue to be solved. So, we started to re-design the moving mass system. In the new moving mass system, we adopted pulleys instead of ball screw. In FY 2019, an initial design was finished and prototype was fabricated.

**Sapphire fiber thermal conductivity** In KAGRA cryogenic payload, high purity aluminum heatlink is installed except for the sapphire mirror to cool the suspension effectively. Instead of this, sapphire mirror is cooled only by sapphire fibers in order not to induce additional vibration and mechanical losses to the mirror. So, thermal conductivity of sapphire fibers directly affects the cooling performance of the cryogenic payload.

In FY2019, we measured thermal conductivity of several sapphire samples and obtained  $4000 \text{ W/m}\cdot\text{K}$  at 20 K. This value is slightly lower than KAGRA requirement but we started the collaboration with foreign institute for making better sapphire fibers.

**Molecular adosorption at cooled mirror surface** Gas molecules that hit to cryogenic objects are trapped at these surface. This effect is called cryopumping effect and its pumping power is very large. So, cryogenic mirror of KAGRA can trap residual gas molecules inside vacuum chamber and make thin layers on the mirror surface. Since the thin layers will change the reflectivity of the mirror, this effect can affect the detector sensitivity of KAGRA. First, we measure the speed of molecular adosorption in the KAGRA cryostat. Then, we calculated the impact for the KAGRA detector sensitivity based on the adosorption speed we measured. In addition to this, we considered the effective way to remove the adsorbed molecules and demonstrated the disorption by using  $\text{CO}_2$  laser.

**Doctoral theses** One doctoral thesis, *Optical and thermal study of molecular thin layers on cryogenic mirrors in next-generation gravitational wave telescopes* was accepted in FY 2019.

**Acknowledgement** Mechanical workshops of Institute for Solid State Physics (The University of Tokyo, Kashiwa campus) and Mechanical Engineering Center of KEK make a large contribution through providing many products for our research.

### Vibration Isolation System Type-A

[Spokesperson : Lucia TROZZO]

ICRR, The Univ. of Tokyo, Hida, Gifu 506-1205 lucia,s contribution

### Introduction

Seismic noise gives the most important contribution and it is dominant in the range [0.1,10] Hz. This region is critical for the ground based GW detector, because the seismic noise influence the detector sensitivity and the possibility to operate with high duty cycle: The selection of an experimental site has a key role to limit the impact of seismic noise on interferometric detectors for Gravitational Waves observations. In underground locations the perturbations caused by the atmospheric fluctuation due to the weather conditions are minimal, the environmental conditions are more stable and, as a consequence, the measurable noise on the surface level is attenuated: seismic noise, in the underground facilities, is about 100 time smaller than those ones on surface: for all these reasons KAGRA has been designed to be underground. Anyway to achieve the target sensitivity of  $10^{-18} \text{ m}/\sqrt{\text{Hz}}$  at 10 Hz, suspending optical components is a crucial task in the construction of GW interferometers. In fact the gravitational waves observation depends on the capability to include into the experimental apparatus a free falling mass (test

mass), well isolated from all the noise sources which are relevant in the observational frequency band. For these reasons the ground based and broad-band interferometers have been equipped with an appropriate suspension system of the test masses including complex mechanical structures studied to inhibit noise transmission and consider test masses as free-falling from few Hz.

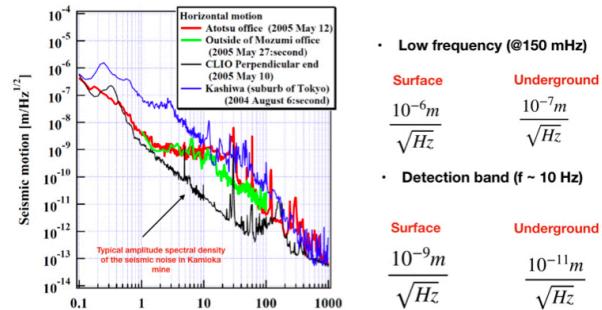


Fig. 7. Seismic noise spectrum on surface in the Tokyo area and the measurements performed at different levels inside the Kamioka mine

### Vibration Isolation System for the Cryogenic Mirrors: Type A

A good approximation of a free falling mass is represented by a simple pendulum composed by a thin wire, as elastic element, connected to a body of mass  $m$ . Thanks to this simple solution it is possible to filter seismic noise in a real experiment to the optical level by using a harmonic oscillator. In GW detectors seismic isolation with a capability attenuation 10 orders of magnitude is needed: It is evident that a multiple stage pendulum represents a good approach to improve the total attenuation factor of seismic noise at the level of the test masses. Solution adopted in KAGRA is based on the idea to replicate harmonic oscillators of length 2 m to obtain a sophisticated mechanical structure: Type-A suspensions. The Type-A suspension is a nine-stage pendulum with the height of 13.5 m based on the second floor of the underground mine. The top of the suspension has a pre-isolation stage supported by three inverted pendulum legs. The suspension chain consists of cascaded geometric anti-spring filters that show low frequency mechanical resonances. The bottom four stages including the sapphire mirror are called cryogenic payload and cooled down to about 20 K in order to reduce the thermal noises. The whole suspension provide a seismic attenuation of 15 orders of magnitude at 10 Hz.

The suspension system has a negative aspect of amplifying the mirror fluctuation at its resonant frequencies: the excess of the mirror displacement even at an out-of-band frequency disturbs stable operation of the interferometer. We need to implement damping controls on suspension stage and on the payload, in order reduce the free swinging of the test masses and to make the system robust against external disturbances.

To monitor and control the displacement of the suspension, the system is equipped with several sensors and actuators: position sensors (LVDT), inertial sensors (accelerometer or geophone), photosensors, optical levers, actuators (coil-

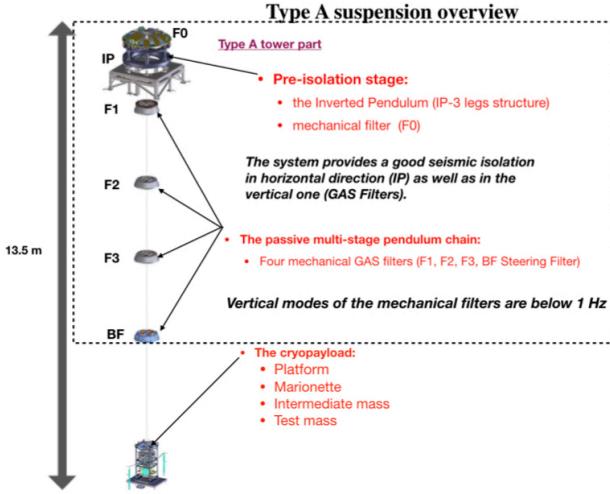


Fig. 8. Drawing of the Type A suspension system

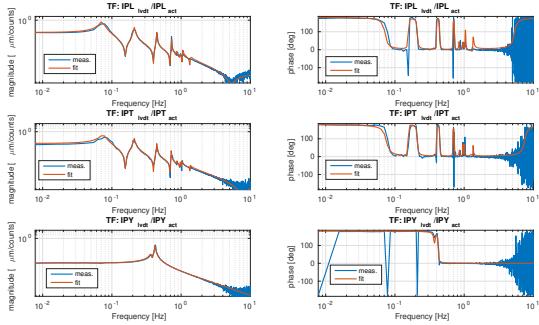


Fig. 9. Normal modes of Type A suspension system

magnet type). Thanks to the presence of these sensors and actuators a feedback control could be implemented in different points: Inverted Pendulum, vertical GAS filters, Bottom Filter, Marionette stage and the Mirror. The intent of the feedback action is to reduce the swinging of the free falling masses along: -longitudinal, transverse direction by acting on the Top stage (IP); - vertical direction by acting on GAS filters; - angular directions (Pitch and Yaw) by acting on the Bottom Filter, the Marionette and the Intermediate mass stages. At the IP stage the "sensor correction technique" and the "tidal compensation" are also implemented. Thanks to them it is possible to achieve a suppression of microseismic noise, in the range [0.1 0.5] Hz, by about a factor 3 and to compensate for long-term drifts in Fabry-Perot cavities due to the tidal effect, respectively. All these control loops play a crucial role in the locking acquisition process of the whole interferometer: in fact they are engaged to achieve and to keep its standard working conditions even after a violent unlock event.

## Vibration Isolation System Type-B

[Spokesperson : Fabian Peña Arellano ]

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Besides the four Type-A suspensions KAGRA relies on smaller suspensions for other mirrors which are always at room temperature. The Type-B suspension is used for the beam splitter and the three signal recycling mirrors, whereas the Type-Bp is used for the three power recycling mirrors.



Fig. 10. The IP of the beam splitter in the vacuum chamber and the installation team.

As in Type-A suspension, in Type-B suspension the first vibration isolation stage is the Inverted Pendulum (IP), whose main goal is to passively attenuate the persistent horizontal microseismic motion produced by ocean activity. Typical resonant frequencies for these IPs are between 60 mHz and 80 mHz in order to attenuate the microseismic peak at around 200 mHz. The following three stages, intended for vertical isolation, are three geometric anti-spring filters. The first one lies directly on top of the IP table while the other two hang from it as the masses of a multi-stage pendulum. As in the case of the IP, GAS filters are devices capable of supporting loads of hundreds of kilograms and at the same time achieving resonant frequencies of a few hundreds of millihertz. The horizontal position of the IP and the vertical positions of the GAS filters are measured with Linear Variable Differential Transformers (LVDTs) and are adjusted with coil-magnet actuators which also damp the mechanical resonant motion. The typical resolution of the LVDTs are below 0.2 μm. From the lowermost GAS filter the payload hangs. The payload, which is common to Type-B and Type-Bp suspensions, comprises the optic, its marionette and their respective recoil masses. The recoil mass of the marionette holds local displacement sensors to monitor six degrees of freedom and coil-magnet actuators for damping the resonant modes of oscillation of the suspension itself. The sensors have typical resolutions of about 20 nm. The recoil mass of the mirror holds coil-magnet actuators only and the system relies on an optical lever to monitor the tilt and position of the optic from the ground.

The Type-Bp suspension is shorter and comprises two GAS filters and the payload. The uppermost GAS filter lies directly on the ground while the lowermost one has its own recoil mass in order to attenuate motion which may be excited by the microseismic motion of the ground.

## Integrated DAQ/control system using real time computers

[Spokesperson : Shoichi OSHINO]

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The 2019 fiscal year, we started the observation with a power-recycled Fabry-Perot-Michelson interferometer from February 2020. During this observation, we continued to maintain stable DAQ/control system.

**Stable operation with the real time control system** The first part of 2019 was the process of replacing computers with a faster ones. This is a countermeasure to the glitches that were identified last year, which occur under high loads. Basically the control computers use a real time operating system. Some delay due to the heavy task causes a serious problem for control loops and it emerges as jumps or glitches on many signals. Eventually, by replacing 17 computers, these glitches no longer occur and we are able to control the interferometer with stable digital system.

In the 2018 fiscal year, we used a simple Michelson interferometer to perform operations, but we have to build a more complicated interferometer for the actual observations. Therefore, we have installed various types of hardware to enable more complicated control. In particular, we have installed and enhanced the hardware for length sensing control and output mode cleaner. Finally, we used the digital system of 25 RT-PCs, 50 ADCs, 39 DACs and 73 BIOs to control the interferometer.

In late 2019, commissioning work of interferometer began and the sensitivity of gravitational wave is visible. Since the DAQ/control system was already stable, we began reducing the noise to improve sensitivity. A number of circuits have been installed in the KAGRA mine. A number of AC-DC converters are also installed to supply power to these circuits. These DC power supplies generate a lot of magnetic noise, which is expected to affect the sensitivity of the interferometer. We solved this problem by replacing the power supplies farther away from the circuit and wiring the power cables. Most of the power supplies have already been replaced. We will continue to replace the remaining power supplies in fiscal year 2020.

## Detector Charcterization

[Spokesperson : Takahiro Yamamoto]

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**Detector commissioning support** The goal of detector characterization is providing the reliable data for gravitational waves searches. In order to achieve this goal, we developed some data monitor tools such as SummaryPages shown in figure 1. SummaryPage is the web-based data monitor system for supporting detector commissioning activities. Commissioning activities are performed to achieve the stable operation of the interferometer and the target sensitivity. On the gravitational wave observation, more than 100,000 signals are acquired as the observational data. SummaryPage helps us to monitor such a lot of signals efficiently. In addition, since



KAGRA joined the international observing network of gravitational waves, SummaryPage also served as an interface for LIGO, Virgo, and KAGRA to know each other's situation.

**Observing data validation** During the observing run in the end of March 2020, we also provided indicators of the data quality as another major role. These indicators are used to judge whether the data can be used for gravitational wave searches. For the joint gravitational wave search with LIGO and Virgo, we also established to share these data quality indicators.

## Physical Environment Monitors

[Spokesperson : Takaaki YOKOZAWA ]

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Because the amplitude of gravitational waves(GWs) are very small, we must take care everything which can be a noise source. One of the major noise sources is environmental disturbance through earthquakes, effects from magnetic and acoustic fields, temperature and so on. To investigate the effect from the environmental noise, we installed various types of the monitors, which are called as physical environmental monitors (PEMs). Also, because KAGRA interferometer was developed in the underground environment and operated with cryogenic temperature (20K). Those environment and technique is essential to the next generation GW interferometer. So the understanding of the KAGRA interferometer environment is key technology.

Six seismometers were installed to evaluate the ground motion in the underground environment and those seismometer were also used for the suspension control to reduce the mirror motion. the accelerometers and microphones were also installed to monitor the optical tables. Such auxiliary optics on the optical tables were used for many purpose, such as laser source stabilization, mode matching, sensing for the interferometer control and controls of the photon calibrator. Those PEMs were used for evaluating not only the stationary interferometer noise but also narrow band frequency noise and transient noise identification.

The temperature monitor is also important for keep the suspensions healthy. If the temperature changed 1 degree, the height of the mirror motion was changed and it became difficult to keep the interferometer condition. For monitoring the temperature, we installed total 77 thermometer, which can monitor both temperature and humidity and made the web site to monitor the temperature drift.

Finally, to evaluate the effect from the environments, we established the system of the environmental injection, such as acoustic injection and magnetic injection system. By performing the acoustic injection, we turned out that the peaks of the 160, 280 and 360 Hz came from the acoustic environmental noise.

## Commissioning

[Spokesperson : Osamu MIYAKAWA]

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For KAGRA, it is not an exaggeration to say that the fiscal year of 2019 has been a year of commissioning. By FY2018, almost all of the subsystems had been installed and in FY2019 we targeted to operate as a whole interferometer. We had several engineering runs with the Michelson interferometer, operating in a single arm cavity, and finally carried it through to observation. It is a remarkable result that we were able to achieve the minimum target sensitivity of 1 Mpc. This led to a joint observation with GEO.

In this process, one major problem was exposed: the presence of birefringence in the Sapphire mirror, which was thought to reduce the returning light from the interferometer. This can be a direct problem for sensitivity since the target power-recycling gain cannot be achieved. It was then found that the modal healing effect of the arm cavity relieved this problem and that sufficient power recycling gain could be obtained. However, the sidebands would still be lost due to the birefringence, so we should consider making the mirror again in order to achieve the final sensitivity.

The Fig.11 shows the improvement in sensitivity over a period of about six months. We successfully achieved the first FPMI operation in late August, with both 3 km arms storing light into the arm cavities, and the first sensitivity was measured. At this point, it was found that 4-5 orders of magnitude were necessary to reach the minimum target sensitivity, and we realized that noise-hunting was the most important task for commissioning. We improved the sensitivity initially by resolving a few minor problems. When the sensitivity improvement was a bit slow, the introduction of the ISS (Intensity Stabilization Servo) improved by one order, and the removal of the FSS (Frequency Stabilization Servo) improved by another order of magnitude, total improvement in sensitivity by two orders of magnitude was not insignificant at that time. We continued to try to improve the sensitivity, but since we were using only 10% of the transmitted light of the power recycling mirror, it was difficult to increase the laser power in the interferometer further, which limited the sensitivity at high frequencies.

In January, the PRFPMI which aimed to increase laser power in the interferometer was successfully operated. The laser power could be increased by a factor of 100 compared to the FPMI state. Although there were some concerns about the noise in the control systems, it was found that they did not limit the final sensitivity of the PRFPMI for the observations in 2019, and the operation of PRFPMI was fine. However, we believe that we need to move to DRFPMI to realize the final sensitivity as designed.

The following February we successfully established the OMC (output mode cleaner) in operation. This OMC allowed us to measure the sensitivity of the interferometer using DC power. This is called as a DC readout method that was designed as an final method without using RF sideband signals. The first sensitivity of the DC readout was 40 kpc, which was almost the same as the highest sensitivity of FPMI, and we aimed to improve the sensitivity further. From this point, we improved the sensitivity by one more order of magnitude, and when the sensitivity was about 400 kpc, we started the first observation for two weeks. During the observation, the sensitivity fluctuated several times due to uncertainties of calibration. This issue is still under investigation, and further verification is necessary. Finally, we achieved the minimum target sensitivity of 1 Mpc, and this led to the second round of observations April.

Through this commissioning, it was found that both sensitivity and stability are highly dependent on the alignment of mirrors. The angular control of the mirror has been successfully introduced in some degrees of freedom. In order to achieve the final sensitivity, it is necessary to introduce the angular control for all the degrees of freedom. It was also exposed that weakness for climate change, such as micro seismic motion in the 0.1 Hz band existed. This needs to be resolved by improving the control system of the vibration isolation system.

There were various problems and issues during the commissioning process, and most of them were solved and the stable operation was achieved. Although, there are still some problems to be solved, but it is remarkable that we have established the 4-5 orders of magnitude improvement of the sensitivity in only 6 months. In the future, we will update the vibration isolation system and try to achieve the final target sensitivity through further commissioning.

## Observation

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## Data Analysis

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There are variety of data related activities in KAGRA. The main data server of KAGRA is located at ICRR Kashiwa. It has a 2.5PiB data storage. All KAGRA data taken at Kamioka are packed into one file for every 32 seconds, and are transferred continuously to the main data server at Kashiwa. Beside this, low latency data transfer is also done by packing only main interferometer data into one file for every 1 seconds. For low latency data transfer, the latency of about 3 seconds is achieved from Kamioka to Kashiwa (this time include the time necessary for calibration).

KAGRA detector is producing several hundreds thousands of channels of data which record signals from various sensors,

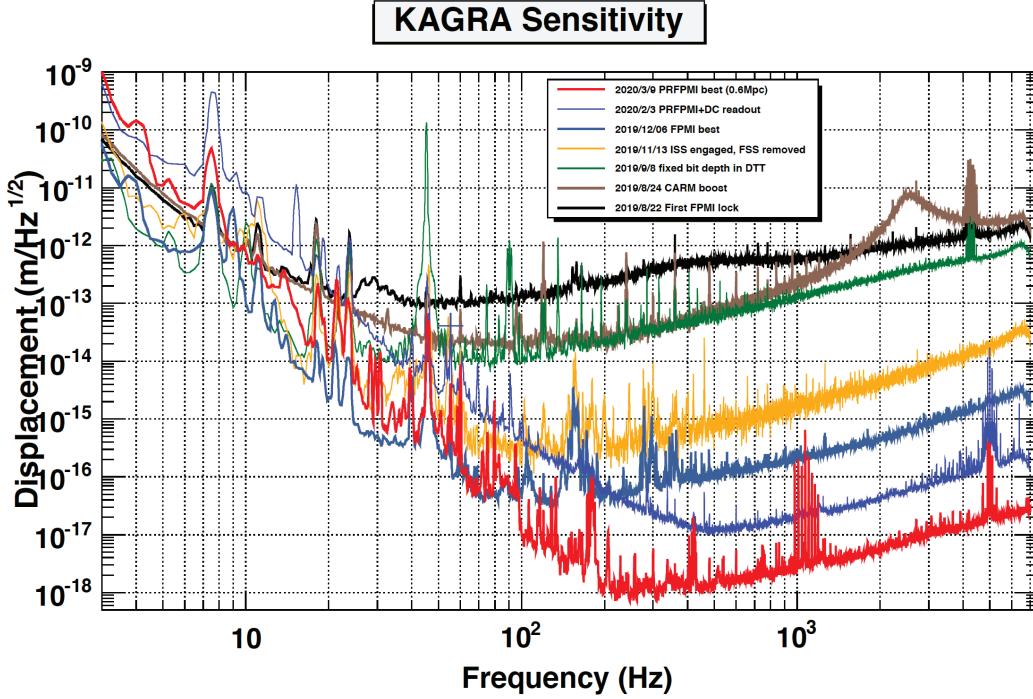


Fig. 11. Improvement of the KAGRA sensitivity in 6 months.

signals to control instruments, signals to monitor environment of the detector. Those data are used to check the status of detector and to improve the sensitivity. It is important to introduce convenient tools to visualize the data in order to accelerate the installation and commissioning works. Web based visualization tools are now being developed. Some of tools developed by LIGO group are also installed. These tools are also useful when gravitational wave signals are detected. In order to have a confidence of detection of gravitational wave signals, it is important to investigate environmental channels whether there are any noise sources which might produce data which are similar to real gravitational wave signals. These visualization tools can be used to check various environmental channel data.

In order to detect gravitational wave signals, several pipelines have been developed in KAGRA. Among them, a pipeline to search for gravitational waves from compact binary coalescences (CBC) are developed in KAGRA Algorithmic Library (KAGALI). KAGALI is a common data analysis library written mainly in C. The CBC pipeline have been used to analyze KAGRA data during iKAGRA operation. Improvement of the CBC pipeline are now ongoing in order to treat multiple detectors and to introduce the spin parameters in the waveform. These tasks will be continued in 2019. The improvement of the parameter estimation pipeline for CBC signals based on the Markov Chain Monte Carlo method was continued from the last year. This work is lead by Hyung Won Lee (Inje Univ).

There are several efforts to introduce new data analysis methods in the analysis of gravitational wave data. Among them, the performance of Non-Harmonic Analysis (NHA) in visualizing the time-frequency behavior of the data was

evaluated. NHA is a method to evaluate the spectrum of data by evaluating multiple instantaneous frequencies and amplitudes of data in a way which is different from discrete Fourier transform. We find that there are various advantage in NHA in visualizing CBC signals compared with the method of short time Fourier transform. We apply NHA to public data of LIGO-Virgo events, like GW150914, GW170817, and demonstrated the visualization of the signal on time-frequency plane. This work has been done in collaboration with the group of Shigeki Hirobayashi (Univ. Toyama).

Ref. Kenta Yanagisawa , Dongbao Jia, Shigeki Hirobayashi, Nami Uchikata , Tatsuya Narikawa, Koh Ueno, Hirotaka Takahashi, Hideyuki Tagoshi, PTEP 2019 (2019) no.6, 063F01.

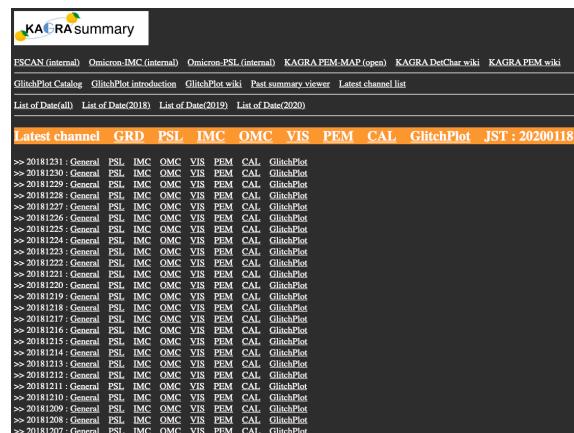


Fig. 12. Example of summary page

## KAGRA OBSERVATORY

KAGRA observatory is located in the Ikenoyama-mountain on the border between Gifu and Toyama prefecture, about 35 km south of Toyama city in Japan. The observatory was established in 2016 in order to operate Large -scale Cryogenic Gravitational Wave Telescope (nicknamed “KAGRA”). KAGRA itself has a L-shape tunnel facility, and it is located more than 200m under Mt.Ikeno-yama. The corner station of the L-shape tunnel is accessible through a 500-m horizontal access tunnel from Atotsu area. The observatory has its own surface research buildings and rental space in the community center of Hida city located about 5km away from the Atotsu entrance of KAGRA.

KAGRA aims to observe several gravitational waves (GWs) per a year with its designed sensitivity as one of observatories of the world GW detection network including Advanced-LIGO, Advanced-Virgo and planned LIGO-India. KAGRA project (formerly named LCGT) was partially approved in 2010 as one of Leading-edge Research Infrastructure Program, and also supported by Program for Promoting Large-scale Science Projects, Subsidy for Facilities Expense

and Grants-in-Aid for Scientific Research from Ministry of Education, Culture, Sports, Science and Technology (MEXT).

In KAGRA project, Institute for Cosmic Ray Research plays a role of a host promoting institute, and National Astronomical Observatory in Japan (NAOJ) and High Energy Accelerator Research Organization (KEK) are the main support organizations, then more than 297 researchers in 85 institutes and universities in the world are collaborating for construction and data analysis of KAGRA.

The tunnel excavation started in May 2012, and finished in March 2014. After that, the basic laboratory environment was prepared until September 2015. A Michelson interferometer with 3km arm (iKAGRA) was demonstrated in March 2016, and the first engineering run was performed until May 2016. At present (April 2019), all the interferometer components had been installed to complete the KAGRA Observatory that adopts a power recycled Fabry-Perot Michelson type interferometer with the resonant sideband extraction and the interferometer is under commissioning. We plan to start the joint observation with LIGO and Virgo within 2019.



Fig. 1. Surface Research Building.



Fig. 2. Atotsu Entrance of KAGRA.