

# MACE: a Machine-learning Approach to Chemistry Emulation

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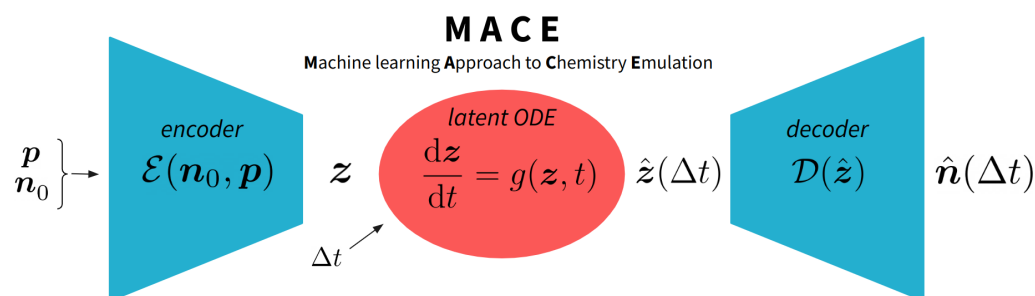
## Summary

Astrochemistry is the study of chemical species in astrophysical objects. By combining observations and theoretical predictions, the physical conditions of the object can be analysed as well as its chemical composition and yields, since chemistry is closely coupled to the dynamics of the environment. Most often, such astrophysical environments are known to be complex and out of thermodynamic equilibrium. For that reason, the chemical evolution is usually modelled using a chemical kinetics approach, in which a set of non-linear coupled ordinary differential equations (ODEs) is solved for a given network of chemical species and reactions. For large chemical networks, however, this method is computationally slow. Therefore, we developed MACE, a *Machine-learning Approach to Chemistry Emulation*, inspired by similar work in the literature, e.g., Holdship et al. (2021); Grassi et al. (2022); Sulzer & Buck (2023). Their code is publicly and can be found in their respective papers. However, not all of their code can be easily adapted or used as such.

MACE is a PyTorch module that offers a trainable surrogate model that is able to emulate chemical kinetics in only 5% of the computation time of its classical analogue, in the current implementation. More speed up is expected when using GPUs instead of only CPUs. MACE provides a machine-learning architecture, consisting of an autoencoder and a trainable ODE, implemented in PyTorch and [torchode](#) (Lienen & Günnemann, 2022). Mathematically, MACE is given by

$$\hat{n}(t) = \mathcal{D}\left(g(\mathcal{E}(\mathbf{n}, \mathbf{p}), t)\right), \quad (1)$$

where  $\hat{n}(t)$  represents the predicted chemical abundances at time  $t$ . The symbols  $\mathbf{n}$  and  $\mathbf{p}$  represent the chemical abundances and physical parameters, respectively. The autoencoder consists of an encoder  $\mathcal{E}$ , able to map  $\mathbf{n}$  and  $\mathbf{p}$  to a latent space representation  $\mathbf{z}$ , and decoder  $\mathcal{D}$ , working the other way around. The function  $g$  represents the trainable ODE, which needs to be solved for a time step  $\Delta t$ . A schematic representation of the architecture and flow of the emulator is given in [Figure 1](#). More details can also be found in Maes et al. (2024). MACE is developed in the context of circumstellar envelopes (CSEs) of evolved stars, but is flexible to be retrained on data of other astrophysical environments.



**Figure 1:** Schematic representation of the architecture of MACE. The autoencoder is shown in blue, the time evolution happens in the red part of the emulator. (Adapted from [Maes et al., 2024](#).)

## Statement of need

Astrophysical objects or environments entail 3D dynamical processes, as well as chemical and radiation processes. Thus, in order to properly model such an environment, these three constituents need to be included in a single simulation. However, modelling each constituents separately already poses a computational challenge for a modest parameter space, let alone running a simulation with the three constituents integrated, and on an extended parameter space. Specifically on the chemistry side, it is generally established that, from a mathematical point of view, chemical kinetics is a stiff problem. When working with an extended chemical network ( $\gtrsim 100$  species, connected by  $\gtrsim 1000$  reactions), the system becomes computationally very slow to solve. Therefore, it is essential to speed up the solver in an alternative way, also because the solver is applied multiple times when calculating the chemical evolution of a certain environment. A fast and accurate chemistry solver would make it computationally feasible to run it alongside a 3D hydrodynamical and/or radiation model.

## Example case

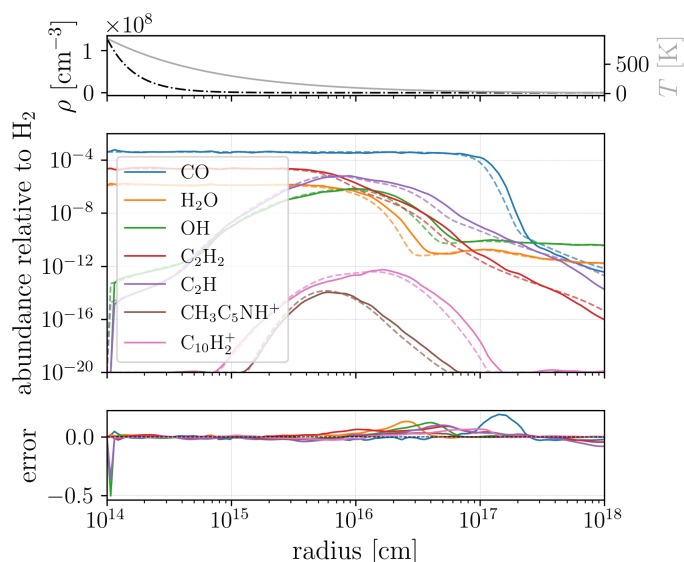
In the context of circumstellar envelopes, we aim to build a comprehensive model, consisting of 3D hydrodynamics, chemistry, and radiation. This is needed since recently, complex structure (such as spiral, arcs, and bipolarity) are observed in CSEs around evolved, low- and intermediate mass stars, currently attributed to the presence of (sub-)stellar companions ([Decin et al., 2020](#); [Gottlieb et al., 2022](#); [Nordhaus & Blackman, 2006](#)). In order to study these stars and their CSEs, taking into account companions, accurate 3-dimensional models are needed. Currently, only hydrodynamical modelling is done for these systems in 3D (e.g., [Maes et al., 2021](#); [Malfait et al., 2021](#)), though the aim is to include chemistry as well as radiation ([Maes et al., 2022](#)). Starting with integrating chemistry and 3D hydrodynamics, this coupling will only be computationally feasible when a faster alternative to chemical kinetics exists.

The development of MACE is a first step towards this coupling. Using the MACE architecture, we are able to reproduce the 1D abundance profiles of 468 chemical species with an accuracy between 10 and 18, averaged over 3000 tests, for MACE models with different sets of hyperparameters, for the following error metric:

$$\text{error} = \frac{\log_{10} \mathbf{n} - \log_{10} \hat{\mathbf{n}}}{\log_{10} \mathbf{n}}, \quad (2)$$

which is executed element-wise and subsequently summed over the different chemical species. More details on the accuracy of the MACE models can be found in [Maes et al. \(2024\)](#). [Figure 2](#) shows the abundance profiles of seven chemical species, where the full curves indicate the

MACE test of model *int4* from Maes et al. (2024), and the dashed curves give the result for the classical model. On average, the MACE routine provides a speed-up of a factor 26, not taking into account the potential extra speed-up factor due to efficient vectorisation when coupling MACE with an SPH hydrodynamical model, or making it compatible with GPU calculations.



**Figure 2:** Chemical abundance profiles from a test of trained MACE model model *int4* (full curves), compared to the classical model (dashed curves). The error on the MACE model is calculated according to Equation 2. More details in Maes et al. (2024).

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