Search for WIMP Inelastic Scattering Off Xenon Nuclei With Xenon100 Data

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Some nice abstract here.....

I. INTRODUCTION

Astrophysical and cosmological evidence indicates that the dominant mass fraction of our Universe consists of some yet unknown form of dark, or invisible matter. The dark matter could be made of new, yet undiscovered particles, and well-motivated theoretical models going beyond the Standard Model of particle physics predict the existence of Weakly Interacting Massive Particles (WIMPs), which are natural candidates for dark matter. This hypothesis is currently being tested by several direct and indirect detection experiments, as well as at the LHC [**? ?**].

Most direct detection searches focus on elastic scattering of galactic dark matter particles off nuclei, where the keV-scale nuclear recoil energy is to be detected [**? ? ?**]. In this work, the alternative process of inelastic scattering is explored, where a WIMP-nucleus scattering induces a transition to a low-lying excited nuclear state. The experimental signature is a nuclear recoil detected together with the prompt de-excitation photon [**?**].

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We consider the ^{129}Xe isotope, which has an abundance of 26.4% in natural xenon, and a lowest-lying $3/2^+$ state at 36.6 keV above the $1/2^+$ gound state. The electromnagnetic nuclear decay has a half-life of 0.97 s. The sigantures of inelastic scattering in xenon have been studied in detail in [**?**]. It was found that this channel is complementary to spin-dependent, elastic scattering, dominating the integrated rates above $\simeq 10$ keV energy depositions. In addition, in case of a positive signal, the observation of inelastic scattering would provide a clear indication of the spin-dependent nature of the fundamental interaction.

II. THE XENON100 DETECTOR

The XENON100 experiment operated a dual-phase (liquid and gas) xenon time projection chamber (TPC) at the Laboratori Nazionali del Gran Sasso (LNGS) in Italy. It contained 161 kg of xenon in total, with 62 kg in the active region of the TPC. These were monitored by 242 low-radioactivity, UV-sensitive photomultiplier tubes (PMTs) arranged in two arrays, one in the liquid and one in the gas. The PMTs detected the prompt scintillation (S1) and the delayed, proportional scintillation signal (S2) created by a particle interacting in the active TPC region. The S2-signal is generated due to ionisation electrons, drifted and extracted into the gas phase, where the proportional scintillation, or electroluminiscence, is produced. These photons carry the $x-y$ information of the interaction site, while the $z-$information came from the drift time measurement. The TPC thus yielded a three-dimensional event localisation, enabling to reject the majority of background events via fiducial volume cuts, as well a multiple interactions, which are unlikely due to dark matter particles [**?**]. The ratio of S2/S1 provided the basis for distinguishing between nuclear recoils (NRs), as induced by fast neutrons and expected from elastic WIMP-nucleus scatters, and electronic recoils (ERs) produced by β and γ-rays. As we shall see, the inelastic scattering signature will be a superposition of an ER and a NR, due to the low lifetime of the excited state and the short mean free path (∼0.15 mm) of the 39.6 keV photon.

XENON100 has acquired science data between 2008-2015, and has set competitive constraints on spin- independent [**? ?**] and spin-dependent [**? ?**] elastic WIMP-nucleus scatters, as well as on axions/ALPs [**?**] and leptophilic dark matter models [**? ? ?**].

III. DATA ANALYSIS

This analysis is performed using XENON100 Run-II science data, which corresponds to an exposure of 224.6 live days. The detector response to electronic recoil (ER) has been characterized using ^{60}Co and ^{232}Th radioactive sources, while response to inelastic nuclear recoil (NR) scattering was calibrated using an ^{241}AmBe source.

The inelastic scattering of a WIMP with the nucleus of ^{129}Xe produces an energy deposit via nuclear recoil with subsequent emission of a 39.6 keV de-excitation photon. The largest fraction of the energy released in the event is via electronic recoil due to the emitted photon, this represents an unusual signature for this kind of detector and brings the possible signal to overlay a phase space region with large backgrounds. The chosen region of interest (ROI) for this analysis surrounds the 39.6 keV xenon line in the cS1-cS2 plane and is further divided into subregions, as shown in Figure 1 and 2.

Events are asked, other than falling in the defined region of interest, to fulfill several selection criteria which can be summarized as: selection aimed to reduce noise impact and including energy (S1) and S2 thresholds, veto of events with energy release in the detector's outer volume, and additionally, events must be of single scatter nature and fall into a predefined fiducial volume. This analysis follows the selection criteria described in detail in [2] for Run-II, with only few exceptions reported in what follows. The selection on S2 width as a function of drift time has been optimized on a sample of events selected from the 39.6 keV line and set to a 95% acceptance on these. Events are required to be single scatter by applying a threshold on the second largest S2 peak size, for this analysis this threshold has been optimized to 160 PE and set constant as function of S2 signal size. Finally the chosen fiducial volume corresponds to 34 Kg of liquid xenon.

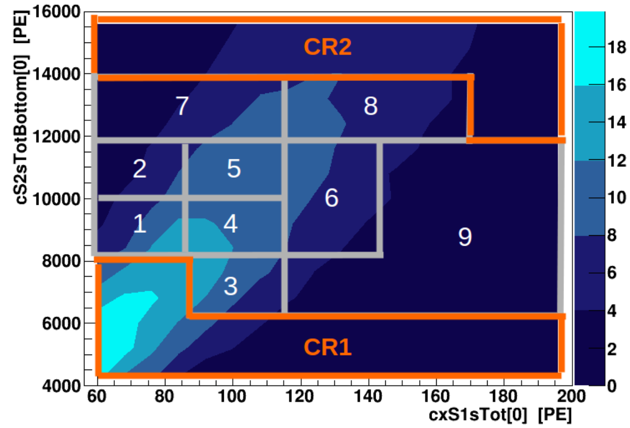

FIG. 1. Signal region and control region.

A. Signal Simulation

The detector response to inelastic scattering of WIMPs off ^{129}Xe nucleus was simulated using an empirical model. The total deposited energy is divided into two independent contributions: one related to the 39.6 keV de-excitation photon and the other relative to the simultaneous nuclear recoil of the xenon atom. The number of photons and charge yield detected is simulated separately for each contribution and then added together.

The distribution in the cS1-cS2 plane of an ER induced by the de-excitation photon is simulated assuming a two dimensional normal pdf, $f(cS1, cS2)$, described except a normalization factor in equation 1:

$$f(cS1, cS2) = exp\Big(- \frac{1}{2(1-\rho^2)} \Big[\frac{(cS1-\mu_{cS1})^2}{\sigma^2_{cs1}} + \frac{(cS2-\mu_{cS2})^2}{\sigma^2_{cs2}} - \frac{2\rho(cS1-\mu_{cs1})(cS2-\mu_{cs2})}{\sigma_{cs1}\sigma_{cs2}} \Big] \Big) \quad (1)$$

where μ_{cS1} and μ_{cS2} represents the average observed cS1 and cS2 given a 39.6 keVee , σ_{cs1} and σ_{cs2} are the standard deviation in cS1 and cS2 respectively, while ρ stands for the correlation between cS1 and cS2. The detector related light yield L_y at 39.6 keV, necessary to evaluate the average number of photon detected (μ_{cS1}), is obtained as the result of a NEST model [3–5] fit to data collected with several lines. The same model is used to predict the charge yield at 39.6 keV which is then scaled according to the detector's secondary scintillation gain Y, determined from detector response to single electrons [6]. The detector resolution at 39.6 keV in cS1 and cS2 has been measured to be respectively 15.8% and 14.7% and used to extract the standard deviations σ_{cs1}, σ_{cs2}. The correlation parameter is assumed to be independent of energy (at least in the considered range) and measured using the 164 keV xenon activated line by ^{124}AmBe calibration data, this line is chosen since allows to disentangle efficiently contribution from nuclear recoil. The measured correlation is $\rho = -0.45 \pm 0.10$.

The cS1 and cS2 distributions from NR contribution are predicted starting from the nuclear recoil energy spectrum of WIMPs inelastic interaction [1]. The average cS1 and cS2 are given by equations 2 and 3 respectively, where $\mathcal{L}_{eff}$ is the liquid xenon NR relative scintillation efficiency, while $S_{ee} = 0.58$ and $S_{nr} = 0.95$ describe the scintillation quenching due to the electric field [7]. The parameterization and uncertainties of $\mathcal{L}_{eff}$ as a function of E_{nr} are based on existing direct measurements [8]. The light yield at 122 keVee originate from the same NEST model fit as described above. For the cS2 the parameterization of $Q_Y(E_{nr})$ is taken from [9]. Finally all detector related resolution effects are introduced following the prescriptions described in [2].

$$cS1_{nr} = E_{nr}\,\mathcal{L}_{eff}(E_{nr})\,L_y\,\frac{S_{nr}}{S_{ee}} \tag{2}$$

$$cS2_{nr} = E_{nr}\,Q_Y(E_{nr})\,Y \tag{3}$$

The pdf of the ER and NR contributions are then convoluted together to obtain the overall pdf of the signal. A 2D (cS1 versus cS2) acceptance map is applied to the signal pdf to reproduce data selection effects. Acceptances are computed separately for each selection criteria using ^{124}AmBe calibration sample, selections as the outer volume veto and the single scatter interaction represent an exception and a dedicated computation has been performed in these cases. The selections acceptance average in the region of interest to about 0.80 ± 0.05. Figure 1 shows an example of full simulated signal model for a WIMP of 100 GeV mass.

The signal simulation procedure has been validated reproducing the 39.6 keV xenon line from interaction due to ^{124}AmBe source and has been compared to data. For this comparison the proper ^{124}AmBe nuclear recoil and acceptances were simulated. The simulated events were found in agreement with calibration data within statistical uncertainties.

B. Background Model

The main expected background contribution in the region of interest is due to eviromental and material radioactivity, its composition is mainly represented by Compton scattering photons. Background contribution due to the activation of the xenon 39.6 keV line from radiogenic neutrons is expected to be negligible.

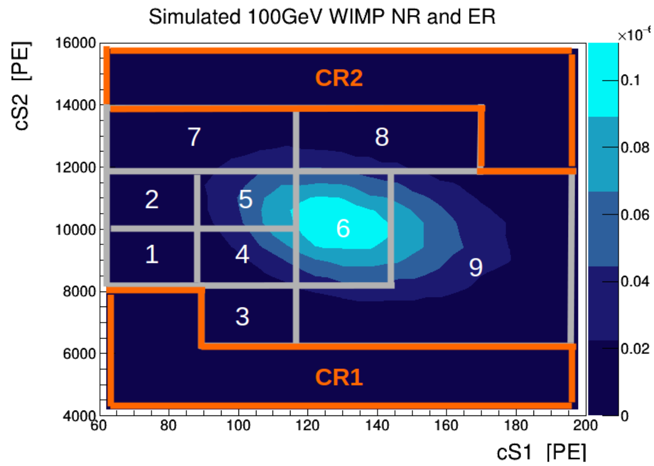

FIG. 2. Signal region and control region, for WIMP of mass 100 GeV.

The background is modeled using data from the ^{60}Co calibration campaign, which are assumed to well represent the background density distribution in the cS1-cS2 plane. The calibration sample yields about 22'000 events in the ROI, these are then scaled to data according to a measured scale factor τ_{bkg}. This scale factor, which is merely the ratio between the data and calibration sample yields, is measured in the two control regions shown in Figure 1 and labelled CR1 and CR2. The two control regions give compatible results and the computed average is $\tau_{bkg} = 0.034 \pm 0.002$, where the reported uncertainty is of statistical nature only.

The distribution of the calibration sample has been compared to the data of the science run in the two control regions, agreement is found within statistical uncertainties. Furthermore, ^{60}Co calibration data have been compared in the region of interest to data from ^{232}Th calibration campaign, the largest deviation between the two shapes is within 4%. An additional systematic uncertainty of 4% has been applied to the expected background yield of each sub-region of the ROI.

C. Systematic Uncertainties

Uncertainties on the total prediction of background events arise from the uncertainty on the measure of the normalization factor, τ_{bkg}, and amount to 6%, contribution of radiogenic neutrons are neglected. Systematic uncertainty on the shape of the predicted background distribution are assessed by the maximal discrepancy in the ROI between the ^{60}Co and ^{232}Th calibration samples, a 4% systematic additional to statistical uncertainty is assigned to the expected yield of each sub-region. Note that uncertainties belonging to different sub-regions are considered independent from each other.

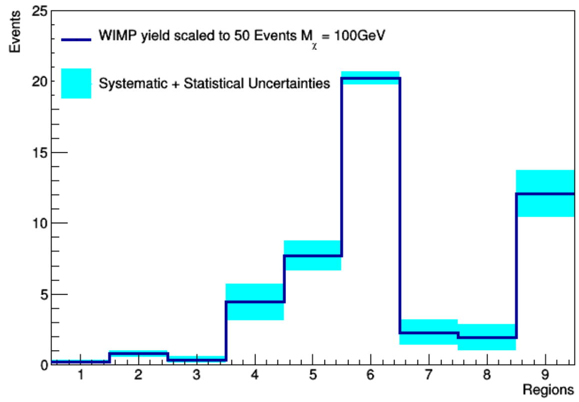

FIG. 3. Signal region, uncertainties for WIMP of mass 100 GeV.

Uncertainty on the total yield of signal arising from selections acceptance uncertainties are found to be very weakly dependent on the WIMP mass, an overall 6% acceptance uncertainty is then applied to all WIMPs hypothesis.

Uncertainties on the energy scale and, more generally, related to detector response are parameterized using the respective uncertainties on the measure of L_y, $\mathcal{L}_{eff}$, Y, Q_Y and ρ. The simulation shows that these type of uncertainties mainly affect the pdf of the signal model in the ROI, and very weakly the total signal yield. They are taken into account by simulating several signal pseudo-sample for each WIMP mass, the pseudo-samples are produced varying the model parameters respectively of ±1 standard deviation. For each sub-region is then computed an overall uncertainty by adding in quadrature the variations of each pseudo-sample with respect to nominal. Figure 3 is an example of such a systematic uncertainty computation for a WIMP of 100 GeV mass.

All the uncertainties discussed here are parameterized within a binned profiled likelihood function using the framework [10, 11]. All parameters related to systematic uncertainties are assumed to be normally distributed.

IV. RESULTS

Using an exposure of 34 kg of liquid xenon and 224.6 live days of data a yield of 764 events is observed in the region of interest, this is compatible with the expectation of $756 \pm 5^{(stat.)} \pm 55^{(syst.)}$ events from the background only hypothesis. Figure 4 shows how these events are distributed in the region of interest, the bottom panel shows the ratio between data and expected background, where the gray and orange shaded areas represent respectively statistical and systematic uncertainty on the background expectation.

This result is interpreted via a binned profiled likelihood approach by means of the test statistic $\tilde{q}$ and its asymptotic distributions described in [12]. Assuming an isothermal WIMP halo with a local density of $\rho_\chi = 0.3\, GeV/cm^3$, a local circular velocity of $v_0 = 220$ km/s, and a galactic escape velocity of $v_{esc} = 544$ km/s, other assumptions..., a 90% CL$_s$ [13] confidence level limit is computed on the spin dependent inelastic WIMP-nucleon cross section, σ_{inel}, as a function of the WIMP mass, m_χ, and shown in Figure 5. The expected median sensitivity is reported with its relative one (green area) and two (yellow area) standard deviation uncertainty. A limit is set on σ_{inel} to $3.3 \times 10^{-38}\ cm^2$ at 90% CL$_s$ confidence level for a WIMP of mass 100 GeV. This limit is compared with decide which other experiment to plot.

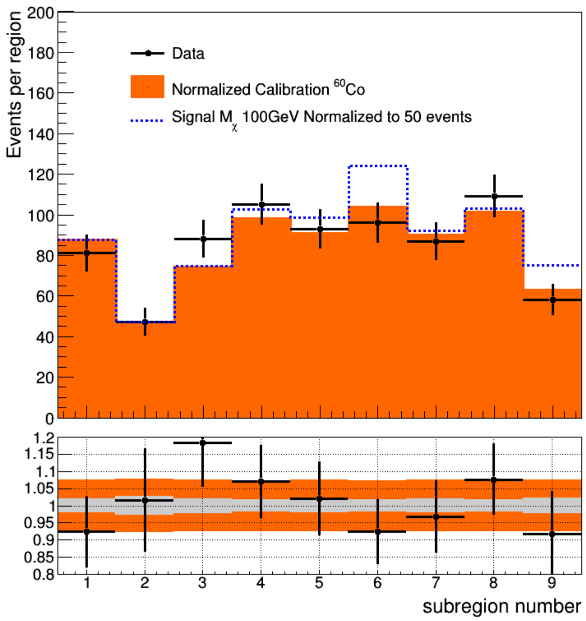

FIG. 4. Results, comparison between data and expected background.

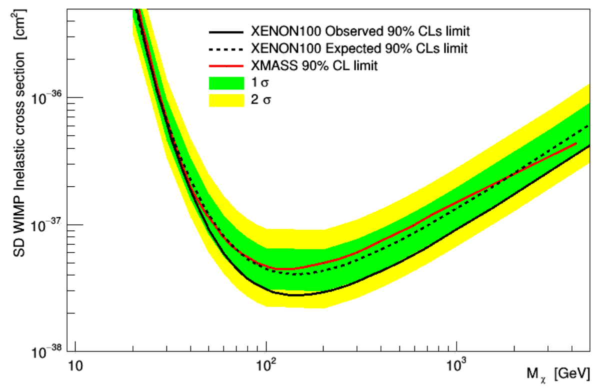

FIG. 5. Observed and expected limits.

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