

A repeating fast radio burst

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Fast radio bursts are millisecond-duration astronomical radio pulses of unknown physical origin that appear to come from extragalactic distances^{1–8}. Previous follow-up observations have failed to find additional bursts at the same dispersion measure (that is, the integrated column density of free electrons between source and telescope) and sky position as the original detections⁹. The apparent non-repeating nature of these bursts has led to the suggestion that they originate in cataclysmic events¹⁰. Here we report observations of ten additional bursts from the direction of the fast radio burst FRB 121102. These bursts have dispersion measures and sky positions consistent with the original burst⁴. This unambiguously identifies FRB 121102 as repeating and demonstrates that its source survives the energetic events that cause the bursts. Additionally, the bursts from FRB 121102 show a wide range of spectral shapes that appear to be predominantly intrinsic to the source and which vary on timescales of minutes or less. Although there may be multiple physical origins for the population of fast radio bursts, these repeat bursts with high dispersion measure and variable spectra specifically seen from the direction of FRB 121102 support an origin in a young, highly magnetized, extragalactic neutron star^{11,12}.

FRB 121102 was discovered⁴ in the PALFA survey, a deep search of the Galactic plane at 1.4 GHz for radio pulsars and fast radio bursts (FRBs) using the 305-m William E. Gordon Telescope at the Arecibo Observatory and the 7-beam Arecibo L-band Feed Array (ALFA)^{13,14}. The observed dispersion measure (DM) of the burst is roughly three times the maximum value expected along this line of sight in the NE2001 model¹⁵ of Galactic electron density, that is, $\beta_{\text{DM}} \equiv \text{DM}_{\text{FRB}} / \text{DM}_{\text{max}}^{\text{Gal}} \approx 3$, suggesting an extragalactic origin.

Initial Arecibo follow-up observations were limited in both dwell time and sky coverage and resulted in no detection of additional bursts⁴. In 2015 May and June we carried out more extensive follow-up using the Arecibo telescope, covering an $\sim 9'$ radius with a grid of six ALFA pointings around the then-best sky position of FRB 121102 (Fig. 1 and Extended Data Tables 1 and 2). As described in Methods, high-time-resolution, total intensity spectra were recorded, and the data were processed using standard radio-frequency interference (RFI) excision, dispersion removal, and single-pulse-search algorithms implemented in the PRESTO¹⁶ software suite and associated data reduction pipelines^{14,17}.

We detected ten additional bursts from FRB 121102 in these observations. The burst properties, and those of the initial FRB 121102 burst, are listed in Table 1. The burst intensities are shown in Fig. 2. No other periodic or single-pulse signals of a plausible astrophysical origin were detected at any other DM. Until the source's physical nature is clear, we continue to refer to it as FRB 121102

and label each burst chronologically starting with the original detection.

The ten newly detected bursts were observed exclusively in two adjacent sky positions of the telescope pointing grid located $\sim 1.3'$ apart (Fig. 1 and Extended Data Table 1). The unweighted average J2000 position from the centres of these two beams is right ascension $\alpha = 05^{\text{h}} 31^{\text{m}} 58^{\text{s}}$ and declination $\delta = +33^{\circ} 08' 04''$, with an uncertainty radius of about $3'$. The corresponding Galactic longitude and latitude are $l = 174.89^{\circ}$, $b = -0.23^{\circ}$. This more accurate position is $3.7'$ from the beam centre of the discovery burst⁴, meaning that FRB 121102 burst 1 was detected well off-axis, as originally concluded⁴.

The measured DMs of all 11 bursts are consistent to within the uncertainties, and the dispersion indices (dispersive delay $\Delta t \propto \nu^{-\xi}$, where ν is the radio frequency) match the $\xi = 2.0$ value expected for radio waves travelling through a cold, ionized medium. This is strong evidence that a single astronomical source is responsible for the events. In addition, the ~ 0.002 DM index uncertainty we calculate for burst 11 (see Methods) is slightly less than that reported for FRB 110523 (ref. 8), making this the most precise determination of dispersion index for any FRB thus far. The upper bound on the dispersion index is identical to that of FRB 110523 (ref. 8) and, hence, following the same arguments used there, burst 11 provides a similar lower limit of ten astronomical units (1 AU is the Earth–Sun distance) for the size of the dispersive region.

The 11 bursts have peak flux densities $gS_{1400} \approx 0.02\text{--}0.3$ Jy at 1.4 GHz, where g is the antenna gain at the source's unknown location in the beam normalized to unit amplitude on the beam axis. The other known FRBs typically have peak flux densities an order of magnitude higher, $gS_{1400} \approx 0.2\text{--}2$ Jy. The wide range of flux densities seen at Arecibo, some near the detection threshold, suggests that weaker bursts are also produced, probably at a higher rate. The rate of burst detections is $\sim 3\text{ h}^{-1}$ for bursts with $gS_{1400} \geq 20$ mJy over all observations in which an ALFA beam was within $3.5'$ of the improved position. We note, however, that the bursts appear to cluster in time with some observing sessions showing multiple bright bursts and others showing none.

The observed burst full-widths at half-maximum are $w_{50} = 2.8\text{--}8.7$ ms, which are consistent with the $w_{50} = 1.3\text{--}23.4$ ms widths seen from other FRBs. No clear evidence for scatter broadening was seen in any of the bursts. Bursts 8 and 10 show double-peaked profiles, which has also been seen in FRB 121002 (ref. 7). Furthermore, the morphologies of bursts 8 and 10 evolve smoothly with frequency.

Within our observing band (1.214–1.537 GHz) the burst spectra are remarkably variable. Some are brighter towards higher frequencies, as in the initial discovery, burst 1, while others are brighter towards lower frequencies. The spectra of bursts 8 and 10 are not monotonic. The detections of bursts 6–11 exclusively in beam 0 of the ALFA receiver

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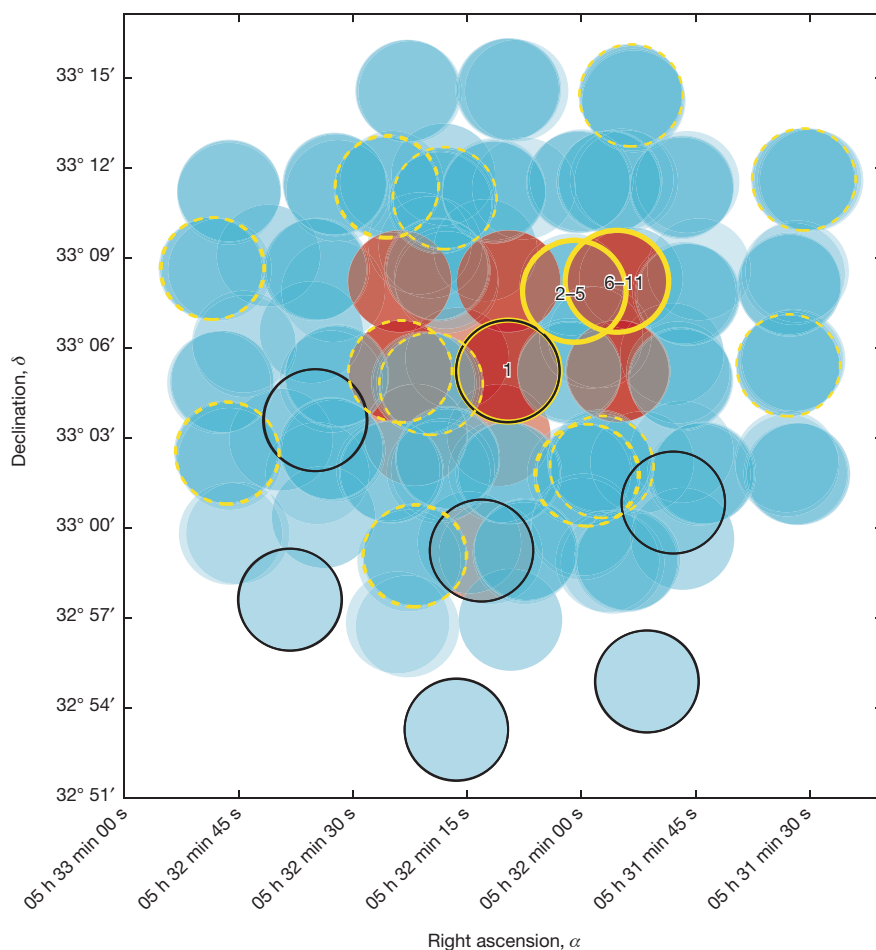


Figure 1 | Discovery and follow-up detections of FRB 121102. For each seven-beam ALFA pointing, the central and outer six beams are shown schematically, in red and blue, respectively (see Extended Data Tables 1 and 2). The circles indicate the $\sim 3.5'$ half-power widths of the beams at 1.4 GHz. Darker shading indicates sky positions with multiple grid observations at roughly the same position. The initial discovery pointing⁴ and second survey observation are outlined in black (these overlap). Beam positions in which bursts were detected are outlined in solid yellow (dashed yellow outlines for the other six beams from the same pointing) and the corresponding burst identifier numbers (Table 1) are given.

(see Extended Data Table 1) means that the bursts must have been detected in the main beam and not in a side-lobe. Although the frequency-dependent shape of the main beam attenuates the bursts' intrinsic spectra at higher frequencies if the source is off-axis⁴, this bias is either not large enough or in the wrong direction to cause the observed spectral variability of bursts 6–11. Given our improved position, burst 1 is consistent with its detection in a side-lobe, which, unlike in the main beam, could have caused attenuation of the spectrum at lower frequencies. This spectral volatility is reflected by the wide range of spectral indices $\alpha \approx -10$ to $+14$ obtained from fitting a power-law model ($S_\nu \propto \nu^\alpha$, where S_ν is the flux density at radio frequency ν) to burst spectra (Table 1).

There is no evidence for fine-scale diffractive interstellar scintillation, most probably because it is unresolved by our limited spectral resolution. In principle, the spectra could be strongly modulated if the source is multiply imaged by refraction in the interstellar medium¹⁸ or by gravitational lensing. However, the splitting angle between sub-images required to produce spectral structure across our band ($\ll 1$ milliarcsecond) is much smaller than the expected diffraction angle from interstellar plasma scattering. The fine-scaled diffraction structure in the spectrum will therefore wash out the oscillation. Lastly, positive spectral indices could also be explained by free-free absorption at the source¹⁹, but this is ruled out by the large spectral differences

Table 1 | Properties of detected bursts

Burst number	Barycentric peak time (MJD)	Peak flux density (Jy)	Fluence (Jy ms)	Gaussian width (ms)	Spectral index	DM (pc cm ⁻³)
1	56233.282837008	0.04	0.1	3.3 ± 0.3	8.8 ± 1.9	$553 \pm 5 \pm 2$
2	57159.737600835	0.03	0.1	3.8 ± 0.4	2.5 ± 1.7	$560 \pm 2 \pm 2$
3	57159.744223619	0.03	0.1	3.3 ± 0.4	0.9 ± 2.0	$566 \pm 5 \pm 2$
4	57175.693143232	0.04	0.2	4.6 ± 0.3	5.8 ± 1.4	$555 \pm 1 \pm 2$
5	57175.699727826	0.02	0.09	8.7 ± 1.5	1.6 ± 2.5	$558 \pm 6 \pm 4$
6	57175.742576706	0.02	0.06	2.8 ± 0.4		$559 \pm 9 \pm 1$
7	57175.742839344	0.02	0.06	6.1 ± 1.4	-3.7 ± 1.8	
8	57175.743510388	0.14	0.9	6.6 ± 0.1		$556.5 \pm 0.7 \pm 3$
9	57175.745665832	0.05	0.3	6.0 ± 0.3	-10.4 ± 1.1	$557.4 \pm 0.7 \pm 3$
10	57175.747624851	0.05	0.2	8.0 ± 0.5		$558.7 \pm 0.9 \pm 4$
11	57175.748287265	0.31	1.0	3.06 ± 0.04	13.6 ± 0.4	$556.5 \pm 0.1 \pm 1$

Uncertainties are the 68% confidence interval, unless otherwise stated. MJD, modified Julian day.

The barycentric peak time is the arrival time corrected to the Solar System barycentre and referenced to infinite frequency (that is, the time delay due to dispersion is removed).

The peak flux density and the fluence are lower limits because it is assumed that the burst is detected at the centre of the beam (that is, with an assumed gain of 10 KJy^{-1} yielding a system equivalent flux density of 3 Jy). Gaussian widths are the full-width at half-maximum. For the spectral index, bursts 8 and 10 are not well fitted by a power-law model and burst 6 is too corrupted by RFI to include. Quoted errors on DM are, in order, statistical and systematic (see Methods). The DM for burst 7 was too weak and corrupted by RFI to include.

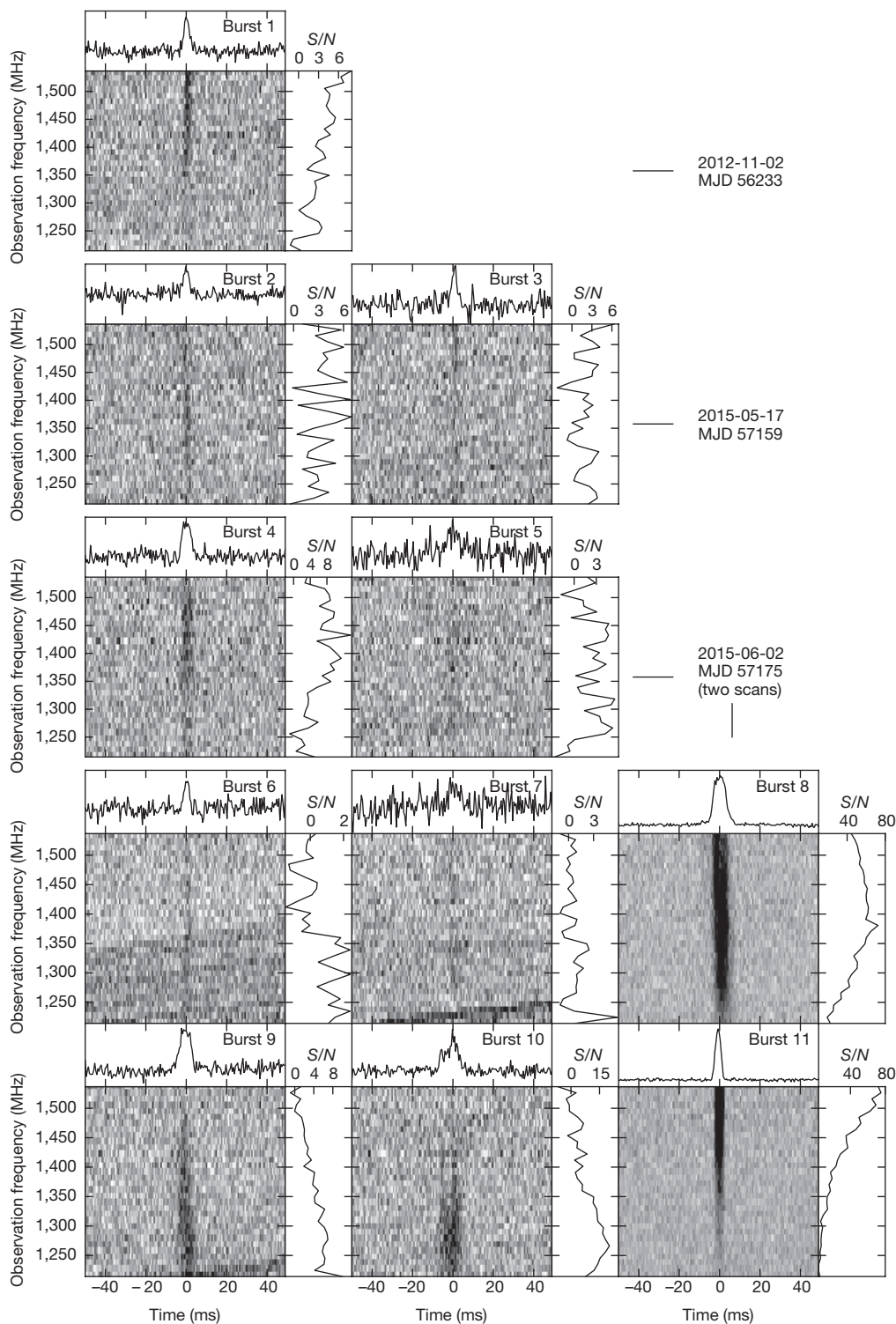


Figure 2 | FRB 121102 burst morphologies and spectra. The central greyscale (linearly scaled) panels show the total intensity versus observing frequency and time, after correcting for dispersion to a value of $DM = 559 \text{ pc cm}^{-3}$. The data are shown with frequency resolution 10 MHz and time resolution 0.524 ms. The diagonal striping at low radio frequencies for bursts 6, 7 and 9 is due to RFI that is unrelated to FRB 121102. The upper sub-panels are burst profiles summed over all frequencies. The band-corrected burst spectra are shown in the right sub-panels. The signal-to-noise ratio (S/N) scales for the spectra are shown on each sub-panel. All panels are arbitrarily and independently scaled.

among bursts. We therefore conclude that the spectral shapes and variations are likely to be predominantly intrinsic to the source.

An analysis of the arrival times of the bursts did not reveal any statistically significant periodicity (see Methods). If the source has a long period ($\geq 1 \text{ s}$), then it is probably emitting at a wide range of rotational phases, which is not uncommon for magnetars²⁰, making a convincing period determination difficult. Owing to the small number of detected bursts, we are not sensitive to periodicities much shorter than $\sim 100 \text{ ms}$.

Repeat bursts rule out models involving cataclysmic events—such as merging neutron stars²¹ or collapsing super-massive neutron stars¹⁰. Bursts from Galactic flare stars have been proposed as a model for FRBs

with the DM excess originating in the stellar corona²². However, temporal density variations in the corona should produce bursts with varying DMs, which we do not observe. Planets orbiting in a magnetized pulsar wind may produce a millisecond-duration burst once per orbital period²³; however, the observed intra-session separations of our bursts (23–572 s) are too short to correspond to orbital periods. Repeated powerful radiative bursts are associated with magnetars, and indeed giant flares from the latter have been suggested as a FRB source^{12,19,24}. However, no Galactic magnetar has been seen to emit more than a single giant flare in over four decades of monitoring, arguing against a magnetar giant flare origin for FRB 121102. Magnetars have

been observed to exhibit repeating bright radio pulses²⁰, but not yet at the energy scale implied if FRB 121102 is more than several hundred kiloparsecs away.

Giant pulse emission from an extragalactic pulsar remains a plausible model¹¹. The most prominent giant pulses are from the Crab pulsar, which has a large spin-down energy loss rate. Spectral indices calculated from wideband measurements of giant pulses from the Crab pulsar²⁵ have a broad distribution ranging from $\alpha \approx -15$ to $+10$, as well as frequency ‘fringes’—that is, a banded structure to the emission brightness as a function of frequency²⁶. These fringes have characteristic widths of a few hundred megahertz, and we speculate that—given our 322-MHz observing bandwidth—a similar phenomenon could create the spectral variability we have seen in FRB 121102. The double-peaked nature of some FRB 121102 bursts is also possible in the giant pulse model^{7,25}, and the evolution of these burst morphologies with frequency could imply rapid spectral variation between consecutive (sub-)pulses only milliseconds apart.

The low Galactic latitude and relatively small β_{DM} of FRB 121102 compared with other FRBs raises the question of whether it is genuinely extragalactic in origin (see also Methods). However, no H α or H II regions are seen in archival data along the line of sight to FRB 121102, as might be expected for an intervening ionized nebula⁴ that can give $\beta_{\text{DM}} \gg 1$. Furthermore, a detailed multi-wavelength investigation, which searched for a compact nebula in a sky region that includes the refined position presented here, concluded that FRB 121102’s high DM cannot be explained by unmodelled Galactic structure along the line of sight and that FRB 121102 must therefore be extragalactic¹⁹. Conclusively establishing that FRB 121102 is extragalactic will require arcsecond localization and association with a host galaxy. The repeating nature of the bursts facilitates such localization with a radio interferometer.

Although the FRB 121102 bursts share many similarities to the FRBs detected using the Parkes^{1–3,5–7} and Green Bank⁸ telescopes, it is unclear whether FRB 121102 is representative of all FRBs. The ten bursts from FRB 121102 in 2015 were detected near the best-known position in 3 h of observations. In contrast, follow-up observations of the Parkes FRBs, again using the Parkes telescope, range in total time per direction from a few hours⁹ to almost 100 hours¹ and have found no additional bursts. Arecibo’s sensitivity is at least ten times higher, possibly allowing detection of a broader range of the burst-energy distribution of FRBs, and thus increasing the chances of detecting repeated bursts; for example, of the 11 bursts from FRB 121102, Parkes may have been capable of detecting only bursts 8 or 11. More sensitive observations of the Parkes FRBs may therefore show that they also sporadically repeat.

Alternatively, FRB 121102 may be fundamentally different from the FRBs detected at Parkes and Green Bank. As was the case for supernovae and gamma-ray bursts, multiple astrophysical processes may be required to explain the diversity of observational properties of FRBs.

Online Content Methods, along with any additional Extended Data display items and Source Data, are available in the online version of the paper; references unique to these sections appear only in the online paper.

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Author Contributions L.G.S. and J.W.T.H. led the design and execution of the observing campaign described here. P.S. performed the analysis that discovered the radio bursts. More detailed analysis of the signal properties was done by L.G.S., P.S., S.M.R., M.A.M., J.W.T.H., S.C. and J.M.C. L.G.S., J.W.T.H., P.S. and V.M.K. led the writing of the manuscript. All authors contributed substantially to the interpretation of the analysis results and to the final version of the manuscript.

Author Information Reprints and permissions information is available at www.nature.com/reprints. The authors declare no competing financial interests. Readers are welcome to comment on the online version of the paper. Correspondence and requests for materials should be addressed to J.W.T.H. (j.w.t.hessels@uva.nl).

METHODS

Positional gridding strategy and burst localization. Extended Data Table 1 summarizes all the observations made in both the initial 2013 follow-up and 2015 May/June observations (project code p2886). The ALFA beam 0 pointing positions in J2000 equatorial coordinates are summarized in Extended Data Table 2. In the p2030/p2886 observations, the major axis of the ALFA receiver was rotated $19^\circ/90^\circ$ with respect to North⁴.

In 2015 May/June, we searched for additional bursts from FRB 121102 using a grid of six pointings using the seven-beam ALFA receiver to cover a generous $\sim 9'$ radius around the discovery beam position and side-lobe. The ALFA receiver was aligned East–West to optimize the sky coverage for this specific purpose. The centre beams of the six grid pointings are shown in red in Fig. 1, and the six outer ALFA beams are shown in blue. Each grid pointing position was observed at least four times for $\sim 1,000$ s. The beam positions of the discovery observation and 2013 follow-up gridding⁴ with ALFA (in that case rotated 19° with respect to North) are also indicated using the same colour scheme. The outer six ALFA beams in the multiple grid observations are only at roughly the same position because the projection of the ALFA beams on the sky depends on the position of the telescope feed with respect to the primary reflecting dish, and these do not overlap perfectly between independent observations.

Two bursts on May 17 (bursts 2 and 3) and two on June 2 (bursts 4 and 5) were detected at a single grid position: FRBGRID2b in ALFA beam 6, which had positions of $\alpha = 05\text{ h }32\text{ min }01\text{ s}$, $\delta = +33^\circ 07' 56''$ and $\alpha = 05\text{ h }32\text{ min }01\text{ s}$, $\delta = +33^\circ 07' 53''$ (J2000) at the two epochs—that is, only a few arcseconds apart. Six more bursts (bursts 6–11) were detected on June 2 at a neighbouring grid position, FRBGRID6b in ALFA beam 0, $\sim 1.3'$ away at $\alpha = 05\text{ h }31\text{ min }55\text{ s}$, $\delta = +33^\circ 08' 13''$. In all cases bursts were detected in only one beam of the seven-beam ALFA receiver at any given time. This shows that the bursts must originate beyond Arecibo's Fresnel length of ~ 100 km (ref. 19).

The intermittency of FRB 121102 makes accurate localization more challenging. Nonetheless, the detection in adjacent grid positions is informative, and to refine the position of FRB 121102, we simply take the average position between FRBGRID2b ALFA beam 6 and FRBGRID6b ALFA beam 0, which gives: $\alpha = 05\text{ h }31\text{ min }58\text{ s}$, $\delta = +33^\circ 08' 04''$ (J2000) and, equivalently, Galactic longitude and latitude $l = 174.89^\circ$, $b = -0.23^\circ$. The approximate uncertainty radius of $\sim 3'$ is based on the amount of overlap between the two detection beam positions and the ALFA beam width at half power, which is $\sim 3.5'$. The distance from the initially reported burst 1 position is $3.7'$, consistent with the interpretation that this burst was detected in a side-lobe. Although FRB 121102 bursts have been detected in beams with different central sky positions, all detections are consistent with a well defined sky position when one considers the imprint of the ALFA gain pattern on the sky during each observation⁴.

Galactic versus extragalactic interpretation. Noteworthy is the fact that FRB 121102 lies directly in the Galactic plane, whereas the other claimed FRBs lie predominantly at high Galactic latitudes. The PALFA survey is only searching in the Galactic plane, however, and no comparable FRB survey at 1.4 GHz with Arecibo has been done at high Galactic latitudes. Therefore, this difference may simply be a consequence of where Arecibo has most deeply searched for FRBs and does not necessarily suggest that FRB 121102 is of Galactic origin. Furthermore, FRB 121102 was found in the Galactic anti-centre region of the PALFA survey, whereas searches in the inner-Galaxy region have thus far found no FRBs¹⁴. This may be because the Galactic foregrounds in the anti-centre region are comparatively low, so the deleterious effects of DM smearing and scattering, which may reduce our sensitivity to FRBs, are less important in the outer Galaxy than the inner Galaxy.

The low Galactic latitude of FRB 121102 also contributes to its low DM excess factor $\beta_{\text{DM}} \approx 3$ compared to the $\beta_{\text{DM}} \approx 1.2$ – 40 range seen for the other 15 FRBs in the literature. Only FRB 010621 (ref. 27), with $\beta_{\text{DM}} \approx 1.2$, has a lower β_{DM} than FRB 121102, and it has been proposed to be Galactic²⁸. We note, however, that six of 16 FRBs have DMs comparable to or lower than FRB 121102. Furthermore, its total Galactic DM excess $\text{DM}_{\text{FRB}} - \text{DM}_{\text{max}}^{\text{Gal}} \approx 370\text{ pc cm}^{-3}$ is larger than that of the first-discovered FRB¹. Lastly, within a generous 20-degree radius of FRB 121102, the highest-DM pulsar known is the millisecond pulsar PSR J0557+1550 (ref. 29; also a PALFA discovery), which has $\text{DM} = 103\text{ pc cm}^{-3}$ and $\beta_{\text{DM}} = 0.6$, as well as the highest DM-inferred distance¹⁵ of any pulsar in this region, $d = 5.7\text{ kpc}$. FRB 121102's DM is clearly anomalous, even when compared to this distant Galactic anti-centre pulsar. At an angular offset of 38° , we note the existence of PSR J0248+6021, with $\text{DM} = 370\text{ pc cm}^{-3}$ and $\beta_{\text{DM}} = 1.8$. Although the DM of this young, 217-ms pulsar is in excess of the maximum Galactic contribution in the NE2001 model¹⁵, this can be explained by its location within the dense, giant H II region W5 in the Perseus arm³⁰ at a distance of 2 kpc. A similar association for FRB 121102 has been sought to explain its $\beta_{\text{DM}} \approx 3$, but multi-wavelength investigations have as yet found no unmodelled Galactic structure^{4,19}. In summary,

FRB 121102's comparatively low β_{DM} does not strongly distinguish it from other FRBs, or necessarily suggest it is more likely to be Galactic.

Observations and search processing. Here we provide a brief description of the Arecibo Mock spectrometer data and search pipeline¹⁴ used for our follow-up observations of FRB 121102. The 1.4-GHz data were recorded with the Mock spectrometers, which cover the full ALFA receiver bandwidth in two subbands. Each 172-MHz subband was sampled with 16 bits, a time resolution of $65.5\text{ }\mu\text{s}$, and frequency resolution of 0.34 MHz in 512 channels. The data were later converted to 4-bit samples to reduce the data storage requirements. Before processing, the two subbands were combined into a single band of 322 MHz (accounting for frequency overlap between the two subbands), which was centred at 1,375 MHz and spans 1,214.3–1,536.7 MHz.

We used the PALFA PRESTO-based¹⁶ search pipeline¹⁴ to search for astrophysical signals in the frequency and time domains. These data were processed using the McGill University High Performance Computing Centre operated by Compute Canada and Calcul Québec. The presence of RFI can have a detrimental effect on our ability to detect bursts. We therefore applied PRESTO's *rfifind* software tool to identify contaminated frequency channels and time blocks. Flagged channels and time blocks were masked in subsequent analyses. Time blocks contaminated by RFI are identified using data that are not corrected for dispersive delay (that is, $\text{DM} = 0\text{ pc cm}^{-3}$), in order to avoid removing astrophysical signals. The data were corrected for dispersion using 7,292 trial DMs in the range 0 – $9,866.4\text{ pc cm}^{-3}$, generating a time series at each trial. We performed Fourier analyses of all the time series to look for periodic signals using PRESTO's *accelsearch* software tool and detected no significant signal of a plausible astrophysical origin.

We searched for single pulses in each dispersion-corrected time series by convolving a template bank of boxcar functions with widths ranging from 0.13 ms to 100 ms . This optimizes the detection of pulses with durations longer than the native sample time of the data. Single-pulse events at each DM were identified by applying a signal-to-noise ratio (S/N) threshold of 5.

These single-pulse events were grouped and ranked using the RRATrap sifting algorithm¹⁷. An astrophysical pulse is detected with maximum S/N at the signal's true DM and is detected with decreasing S/N at nearby trial DMs. This is not generally the case for RFI, whose S/N does not typically peak at a non-zero trial DM. The RRATrap algorithm ranks candidates based on this DM behaviour and candidate plots are produced for highly ranked single-pulse groups. These plots display the S/N of the pulse as a function of DM and time as well as an image of the signal as a function of time and observing frequency (for example, Fig. 2). The resulting plots were inspected for astrophysical signals, and pulses were found at a DM of $\sim 559\text{ pc cm}^{-3}$ at a sky position consistent with the discovery position of FRB 121102 (ref. 4). It is possible that the analysed data contain weaker bursts, which cannot be reliably identified because their S/N is too low to distinguish them from RFI or statistical noise. If, in the future, the bursts are shown to have an underlying periodicity, then this would enable a deeper search for weak bursts.

Timing analysis of burst arrival times. Using several approaches, we searched for an underlying periodicity matching the arrival times of the eight bursts detected in the 2015 June 2 observing session. There are no significant periodicities detected through a standard fast Fourier transform of the time series. We then carried out a similar analysis to that routinely used to detect periodicities in sporadically emitting radio pulsars³¹. In this analysis, we calculate differences between all of the burst arrival times and search for the greatest common denominator of these differences. We found several periods, not harmonically related, that fitted different subsets of bursts within a tolerance of 1% of the trial period, but none that fitted all of the bursts. We subsequently calculated residuals for the times-of-arrival for the eight bursts detected on 2015 June 2 for a range of trial periods using the pulsar timing packages TEMPO and PINT (see 'Code availability' section). We found that some of the periods returned by the differencing algorithm also resulted in residuals with root-mean-square value of less than 1% of the trial period. However, there were many non-harmonically related candidate periods resulting in residuals of a comparable root-mean-square value. Furthermore, given the number of trials necessary for this search, none of these trial periods was statistically significant. In addition, owing to the small number of detected bursts, and the widths of the pulses, we were not sensitive to periodicities much shorter than $\sim 100\text{ ms}$ because our tolerance for a period match (or acceptable root-mean-square value) becomes a large fraction of the period and there are many possible fits. The 16-day gap between the 2015 May and June detections precluded us from including the May bursts in any search for periodicity in the single pulses.

Spectral fitting. To produce the spectra shown in the right panels of Fig. 2, we corrected each spectrum for the bandpass of the receiver. We estimated the bandpass by taking the average of the raw data samples for each frequency channel. We then median-filtered that average bandpass with a width of 20 channels to remove the effects of narrow-band RFI and divided the observed spectrum of each burst by this median-filtered bandpass. The band-corrected burst spectra shown in the

right sub-panels of Fig. 2 are still somewhat contaminated by RFI, however. The bottom and top ten channels (3.4 MHz) of the band were ignored owing to roll-off in the receiver response.

To characterize the bandpass-corrected spectrum of each burst, we applied a power-law model using least-squares fitting. The power-law model is described by $S_\nu \propto \nu^\alpha$, where S_ν is the flux density in a frequency channel, ν is the observing frequency, and α is the spectral index. These measured spectral indices and their uncertainties are shown in Table 1. We do not include a spectral index value for burst 6 because of the RFI in the lower half of the band. For bursts 7 and 9, we exclude data below 1,250 MHz, because of RFI contamination. For bursts 8 and 10, the power-law model was not a good descriptor, and therefore no value is reported in Table 1.

We verified this technique by applying the bandpass correction to PALFA data of pulsar B1900+01. The measured spectral index was calculated for ten bright single pulses, and the values are consistent with the published value.

Measurement of DM and dispersion index. We measured the DM for 10 of the 11 bursts and additionally the dispersion index ξ (from the dispersive delay $\Delta t \propto \nu^{-5}$) for the brightest two. The DM and the dispersion index were calculated with a least-squares routine using the SIMPLEX and MIGRAD functions from the CERN MINUIT package (<http://www.cern.ch/minuit>). The user specifies the assumed form of the intrinsic pulse shape, which is then convolved with the appropriate DM smearing factor. For these fits a boxcar pulse template was used.

Subbanded pulse profiles for each burst were generated by averaging blocks of frequency channels. The number of subbands generated depended on the S/N of the burst to ensure that there was sufficient S/N in each subband for the fit to converge. Subbands with no signal were excluded from the fit. Furthermore, the data were binned in time to further increase the S/N and reduce the effects of frequency-dependent flux evolution. As the true intrinsic pulse width is not known, each burst was fitted with a range of boxcar widths. The parameters corresponding to the input template yielding the cleanest residuals are reported.

The DM value was fitted keeping the DM index fixed at 2.0. We note that burst 7 was too weak and corrupted by RFI to obtain reasonable fits. Additionally, for the brightest two bursts (8 and 11), we also did a joint fit of DM and dispersion index. The resulting dispersion index fits were 2.00 ± 0.02 and 1.999 ± 0.002 for bursts 8 and 11, respectively. These values are as expected for radio waves travelling through a cold, ionized medium.

Frequency-dependent pulse profile evolution introduces systematic biases into the times of arrival in each subband. These biases in turn bias the DM determination. These systematics cannot be mitigated without an accurate model for the underlying burst shape versus frequency, which is not available in this case, and is further complicated by the fact that the burst morphology also changes randomly from burst to burst. We estimated the systematic uncertainty by considering what DM value would produce a delay across our observing band that is comparable to half the burst width in each case.

Table 1 presents the results of the fits with the statistical and systematic uncertainties both quoted. The DM estimates do not include barycentric corrections (of the order of 0.01 – 0.1 pc cm^{-3}). Although FRB 121102 is close to the ecliptic, the angular separation from the Sun was always much larger than 10° , and any annual contribution to the DM from the solar wind was small ($< 10^{-3} \text{ pc cm}^{-3}$)^{32,33}. These effects are, therefore, much smaller than the aforementioned systematics in modelling the DMs of the bursts.

The $\pm 1\sigma$ range of DMs for the ten new bursts is $558.1 \pm 3.3 \text{ pc cm}^{-3}$, consistent with the discovery value⁴, $557.4 \pm 2.0 \text{ pc cm}^{-3}$. The DM values and dispersion indices reported here and previously⁴ were calculated using different methods. These two approaches fitted for different free parameters, so different co-variances between parameters may result in slightly different values. Also, different time and frequency resolutions were used. Nonetheless, the burst 1 parameters quoted here and previously⁴ are consistent within the uncertainties. The consistency of the DM values is conclusive evidence that a single source is responsible for the events.

The role of interstellar scattering. Some FRBs have shown clear evidence for multi-path propagation from scattering by the intervening interstellar or extragalactic material along the line of sight^{2,6–8}. However, the burst profiles from FRB 121102 show no obvious evidence for asymmetry from multi-path propagation. An upper bound⁴ on the pulse broadening time from burst 1 is 1.5 ms at 1.5 GHz. Using the NE2001 model for a source far outside the Galaxy, the expected pulse broadening is $\sim 20 \mu\text{s} \times \nu^{-4.4}$ with ν in gigahertz, an order of magnitude smaller than the ~ 2 -ms pulse widths and ~ 0.7 -ms intra-channel dispersion smearing. The features of the spectra cannot be explained by diffractive interstellar scintillations; the predicted scintillation bandwidth for FRB 121102 is $\sim 50 \text{ kHz}$ at 1.5 GHz, which is unresolved by the 0.34-MHz frequency channels of our data. We would, therefore, also not expect to observe diffractive interstellar scintillation in our bursts. Additional scattering occurring in a host galaxy and the intergalactic medium is at a level below our ability to detect. However, observations at frequencies below 1.5 GHz may reveal pulse broadening that is not substantially smaller than the upper bound if we use as a guide the observed pulse broadening from other FRBs^{2,6–8}. Future observations that quantify diffractive interstellar scintillations can provide constraints on the location of extragalactic scattering plasma relative to the source, as demonstrated for FRB 110523 (ref. 8).

The upper bound on pulse broadening for FRB 121102 implies that the apparent, scattered source size for radio waves incident on the Milky Way's interstellar medium is small enough that refractive interstellar scintillation (RISS) from the interstellar medium is expected. For the line of sight to FRB 121102, we use the NE2001 model to estimate an effective scattering-screen distance of $\sim 2 \text{ kpc}$ from Earth and a scattering diameter of 6 milliarcseconds. The implied length scale for phase-front curvature is then $l_{\text{RISS}} \approx 2 \text{ kpc} \times 6 \text{ milliarcseconds} = 12 \text{ AU}$. For an effective, nominal velocity, $V_{\text{eff}} = 100 \times V_{100} \text{ km s}^{-1}$, the expected RISS timescale is $\Delta t_{\text{RISS}} = l_{\text{RISS}} / V_{\text{eff}} = 215 V_{100}^{-1} \nu_{1\text{GHz}}^{-2.2} \text{ days}$. At 1.5 GHz and with an effective velocity due to Galactic rotation of about 200 km s^{-1} in the direction of FRB 121102, RISS timescales of 20–40 days are expected. Modulation from RISS can be several tens of per cent³⁴. This level of modulation could play a part in the detections of bursts in 2015 mid-May and 2015 June and their absence in 2015 early-May and at other epochs. However, the Solar System and the ionized medium have the same Galactic rotation, so the effective velocity could be smaller than 100 km s^{-1} , leading to longer RISS timescales.

Data availability. The beam positions used in Fig. 1 and the data of the bursts used to generate Fig. 2 are provided as Source Data files (available online with the figures).

Code availability. The code used to analyse the data are available at the following sites: PRESTO (<https://github.com/scotttransom/presto>), RRATrap (<https://github.com/ckarako/RRATrap>), TEMPO (<http://tempo.sourceforge.net/>), and PINT (<http://github.com/nanograv/PINT>).

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Extended Data Table 1 | Arecibo FRB 121102 discovery and follow-up observations

UTC Date	Project	Position	Receiver	Frequency (GHz)	Backend	Dwell time (s)	# Bursts
Survey discovery observations presented in Spitler et al. (2014)							
2012-11-02	p2030	FRBDISC	ALFA	1.4	Mocks	200	1 (Beam4)
2012-11-04	p2030	FRBDISC	ALFA	1.4	Mocks	200	0
Follow-up observations presented in Spitler et al. (2014)							
2013-12-09	p2886	FRBGRID1a	ALFA	1.4	Mocks	2700	0
2013-12-09	p2886	FRBGRID2a	ALFA	1.4	Mocks	970	0
2013-12-09	p2886	FRBGRID2a	ALFA	1.4	Mocks	1830	0
2013-12-10	p2886	FRBGRID3a	ALFA	1.4	Mocks	2700	0
2013-12-10	p2886	FRBDISC	327-MHz	0.327	PUPPI	2385	0
Follow-up observations presented here for the first time							
2015-05-02	p2886	FRBDISC	L-wide	1.4	PUPPI	7200	0
2015-05-03	p2886	FRBGRID1b	ALFA	1.4	Mocks	1502	0
2015-05-03	p2886	FRBGRID2b	ALFA	1.4	Mocks	1502	0
2015-05-03	p2886	FRBGRID3b	ALFA	1.4	Mocks	343	0
2015-05-03	p2886	FRBGRID3b	ALFA	1.4	Mocks	1502	0
2015-05-03	p2886	FRBGRID1b	ALFA	1.4	Mocks	921	0
2015-05-05	p2886	FRBGRID1b	ALFA	1.4	Mocks	1002	0
2015-05-05	p2886	FRBGRID2b	ALFA	1.4	Mocks	1002	0
2015-05-05	p2886	FRBGRID3b	ALFA	1.4	Mocks	1002	0
2015-05-05	p2886	FRBGRID4b	ALFA	1.4	Mocks	1002	0
2015-05-05	p2886	FRBGRID5b	ALFA	1.4	Mocks	1002	0
2015-05-05	p2886	FRBGRID6b	ALFA	1.4	Mocks	1002	0
2015-05-09	p2886	FRBGRID1b	ALFA	1.4	Mocks	1002	0
2015-05-09	p2886	FRBGRID2b	ALFA	1.4	Mocks	1002	0
2015-05-09	p2886	FRBGRID3b	ALFA	1.4	Mocks	1002	0
2015-05-09	p2886	FRBGRID4b	ALFA	1.4	Mocks	1002	0
2015-05-09	p2886	FRBGRID5b	ALFA	1.4	Mocks	1002	0
2015-05-09	p2886	FRBGRID6b	ALFA	1.4	Mocks	1002	0
2015-05-09	p2886	FRBGRID6b	ALFA	1.4	Mocks	425	0
2015-05-17	p2886	FRBGRID1b	ALFA	1.4	Mocks	1002	0
2015-05-17	p2886	FRBGRID2b	ALFA	1.4	Mocks	1002	2 (Beam6)
2015-05-17	p2886	FRBGRID3b	ALFA	1.4	Mocks	1002	0
2015-05-17	p2886	FRBGRID4b	ALFA	1.4	Mocks	707	0
2015-05-17	p2886	FRBGRID4b	ALFA	1.4	Mocks	391	0
2015-05-17	p2886	FRBGRID5b	ALFA	1.4	Mocks	1002	0
2015-05-17	p2886	FRBGRID6b	ALFA	1.4	Mocks	1002	0
2015-06-02	p2886	FRBGRID1b	ALFA	1.4	Mocks	1002	0
2015-06-02	p2886	FRBGRID2b	ALFA	1.4	Mocks	1002	2 (Beam6)
2015-06-02	p2886	FRBGRID3b	ALFA	1.4	Mocks	1002	0
2015-06-02	p2886	FRBGRID4b	ALFA	1.4	Mocks	1002	0
2015-06-02	p2886	FRBGRID5b	ALFA	1.4	Mocks	1002	0
2015-06-02	p2886	FRBGRID6b	ALFA	1.4	Mocks	1002	6 (Beam0)
2015-06-02	p2886	FRBGRID6b	ALFA	1.4	Mocks	300	0

The observing setup of these observations is described in Methods. Data from ref. 4 and this study.

Extended Data Table 2 | FRB 121102 gridding positions

Grid ID	Right Ascension	Declination
FRBDISC	05 ^{<i>h</i>} 32 ^{<i>m</i>} 09 ^{<i>s</i>}	+33 ^{<i>d</i>} 05 ^{<i>m</i>} 13 ^{<i>s</i>}
FRBGRID1a	05 ^{<i>h</i>} 32 ^{<i>m</i>} 16 ^{<i>s</i>}	+33 ^{<i>d</i>} 05 ^{<i>m</i>} 39 ^{<i>s</i>}
FRBGRID2a	05 ^{<i>h</i>} 32 ^{<i>m</i>} 22 ^{<i>s</i>}	+33 ^{<i>d</i>} 03 ^{<i>m</i>} 06 ^{<i>s</i>}
FRBGRID3a	05 ^{<i>h</i>} 32 ^{<i>m</i>} 11 ^{<i>s</i>}	+33 ^{<i>d</i>} 03 ^{<i>m</i>} 06 ^{<i>s</i>}
FRBGRID1b	05 ^{<i>h</i>} 32 ^{<i>m</i>} 10 ^{<i>s</i>}	+33 ^{<i>d</i>} 05 ^{<i>m</i>} 13 ^{<i>s</i>}
FRBGRID2b	05 ^{<i>h</i>} 32 ^{<i>m</i>} 24 ^{<i>s</i>}	+33 ^{<i>d</i>} 05 ^{<i>m</i>} 13 ^{<i>s</i>}
FRBGRID3b	05 ^{<i>h</i>} 31 ^{<i>m</i>} 55 ^{<i>s</i>}	+33 ^{<i>d</i>} 05 ^{<i>m</i>} 13 ^{<i>s</i>}
FRBGRID4b	05 ^{<i>h</i>} 32 ^{<i>m</i>} 10 ^{<i>s</i>}	+33 ^{<i>d</i>} 08 ^{<i>m</i>} 13 ^{<i>s</i>}
FRBGRID5b	05 ^{<i>h</i>} 32 ^{<i>m</i>} 24 ^{<i>s</i>}	+33 ^{<i>d</i>} 08 ^{<i>m</i>} 13 ^{<i>s</i>}
FRBGRID6b	05 ^{<i>h</i>} 31 ^{<i>m</i>} 55 ^{<i>s</i>}	+33 ^{<i>d</i>} 08 ^{<i>m</i>} 13 ^{<i>s</i>}

The J2000 ALFA beam 0 positions are listed.