

\*thesis title\*

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2022

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# Chapter 1

## Introduction

### 1.1 Introduction

#### 1.1.1 Neutrino Physics

There are a plethora of physics phenomenon in which neutrinos are involved, including beta decay, cosmic rays, and supernovae. As part of the Standard Model, they are described as being Dirac fermions with no electric charge with three flavours: the electron neutrino, the muon neutrino and the tau neutrino, corresponding with their associated leptons: the electron, muon and tauon. Prior to the discovery of neutrino oscillations, it was believed that neutrinos were massless, but they in fact have small but non zero masses ( $\leq 1\text{eV}$ ). The next subsections of this chapter will discuss a brief history of neutrino physics including the discovery of neutrino oscillations, the manner in which neutrinos interact with nuclei in the Super-Kamiokande detector, and the motivation behind an NCQE neutron tagging analysis.

### 1.1.2 History of Neutrino Physics

To correct a violation of energy and momentum conservation discovered in beta-decay, Wolfgang Pauli put forward the idea of a neutrino as a solution [8]. He originally named this particle the “neutron”, prior to Chandwick’s discovery of the actual neutron two years later. Enrico Fermi then proposed to call it the neutrino, Italian for “little neutral one”. In 1934, Enrico Fermi’s theory of beta decay stated that a neutron could decay to a proton, electron and an anti-electron neutrino and in 1956 [22], Clyde Cowan and Frederick Reines directly confirmed the existence of neutrinos [11], by detecting the electron antineutrino originating from inverse beta decay produced from nuclear reactor or fission products. In inverse beta decay an electron antineutrino interacts with a proton to produce a neutron and positrons. These positrons pair-annihilate with electrons and produce two 0.5 MeV gamma rays which go in opposite directions. Scintillator material was placed in a tank of water, which was used to detect the gamma photons, and the scintillator light produced flashes of visible light which were detected by photomultiplier tubes. Cadmium chloride was used to detect the coincident neutron, which after exciting and de-exciting, emits a gamma ray within 5 microseconds after the pair-annihilation gammas are detected. In 1962, Ledermen, Schwartz and Steinberger detected the muon neutrino [12]. To discover the muon neutrino, the AGS (Alternating Gradient Synchrotron) at Brookhaven was used - the AGS’s beam of protons was used to produce a shower of pions after slamming into a beryllium target. These pions would travel toward a 5,000-ton steel wall (constructed from old battleship plates!) and they would decay into muons and neutrinos, but only neutrinos would be able to pass through the steel wall into a spark chamber filled with neon. As neutrinos coursed through this chamber, one neutrino would occasionally hit a proton in an aluminium nucleus which would produce muon tracks, instead of



Figure 1.1: Proton-proton chain cycle with branching percentages shown

the showers produced from electrons, proving that there were muon neutrinos in the neutrinos that passed through the chamber. In 2000 the DONUT collaboration (Direct Observation of NU Tau) at Fermilab detected the existence of the tau neutrino [13]. The tau neutrino will interact and form a tau which has an 86% chance of decaying which involves only one charged particle (called a kink), which is what DONUT scientists were trying to observe.

### 1.1.3 The Solar Neutrino Problem

The Standard Solar Model predicted that most of the solar neutrino flux (about 90%) come from the proton-proton chain reaction, where the solar neutrinos produced have energies below 0.4 MeV. Figure 1.1 shows the three main cycles which make up the pp chain, along with the branching percentages where it produces a neutrino.



Figure 1.2: Solar neutrino fluxes predicted by Bachall's Standard Solar Model, and which experiments are able to detect them.

Only experiments that use Gallium can detect neutrinos with this energy, whereas experiments whose detector medium is Chlorine can observe neutrinos from  ${}^7\text{Be}$  (production equation of which is shown in Figure 1.1). Experiments whose detector medium is water (such as Super-Kamiokande and SNO) can only see the  ${}^7\text{B}$  neutrinos, as shown in Figure 1.2.

In the 1960s, the Homestake experiment made the first measurement of solar electron neutrino flux [10]. The experiment used a perchloroethylene-based detector, placed 1,478 metres underground in the Homestake Gold Mine in South Dakota. When an electron neutrino interacts with  ${}^{37}\text{Cl}$  in the perchloroethylene, the  ${}^{37}\text{Cl}$  becomes a radioactive isotope of  ${}^{37}\text{Ar}$  which is extracted by bubbling helium through the tank, and then counted in order to determine how many neutrinos had been captured. The theoretical solar neutrino flux calculated by John Bachall was about three times as much as Raymond Davis' results from

the experiment: with Bachall's calculations predicting a solar neutrino flux of  $8.1 \pm 1.2$  SNU, whereas the Homestake experimental results showed a flux of  $2.56 \pm 0.25$  SNU. These results were consistent with those from the subsequent Kamiokande, SAGE and GALLEX experiments. Super-Kamiokande's main detection channel for solar neutrinos is the elastic-scattering channel, which has a minimum energy threshold value of 5 MeV: neutrinos with energies below this value will not generate an electron with enough momentum to be detected. In this channel, Super-Kamiokande observed a solar neutrino capture rate of  $0.45 \pm 0.02$  SNU, while the solar neutrino model prediction was  $1.0 \pm 0.2$  SNU. Turning our attention to the results of the Gallium experiments, namely SAGE and GALLEX, they also showed a discrepancy regarding the solar model predictions. Due to their sensitivity to the bulk of the proton-proton chain neutrino flux, they have a larger solar neutrino capture flux rate, with SAGE observing  $70.8 \pm 5.0$  SNU, and GALLEX observing  $77.5 \pm 8.0$  SNU, while the model prediction was  $129 \pm 9.0$  SNU, a difference of about 40%. This discrepancy of experimental results compared with the Standard Solar Model prediction became known as the Solar Neutrino Problem.

The way to combat this discrepancy and prove that the cause of the Solar Neutrino Problem was neutrino oscillations was finding a way to detect solar neutrino flux which was not dependent on neutrino flavour. Due to these experiments only detecting solar neutrinos through charged current interactions, they wouldn't be able to detect muon or tau neutrinos. However, with the Sudbury Neutrino Observatory (SNO) experiment this would change, as SNO used heavy water  $D_2O$  as its detector medium. Deuteron consists of a proton and a neutron and due to how little energy it takes to break it apart (2 MeV), an electron, muon or tau neutrino can very easily break it apart, so that SNO can detect the final state neutron and muon and tau neutrinos would be visible

to the detector. The three main channels SNO uses to detect neutrinos is the elastic scattering channel, the charged current channel, and the neutral current channel. Elastic scattering allows SNO to see muon and tau neutrino flux given by  $\phi(\nu_e + 0.15(\phi(\nu_\mu) + (\nu_\tau)))$ . The most important neutrino interaction channel regarding SNO being able to measure the muon and tau neutrino flux is the neutral current (NC) channel (shown in Equation 1.1), as it allows SNO to measure the total flux.



Using these two reaction channels, along with the charged current channel (CC), which can only measure electron neutrino flux, SNO was able to measure the individual neutrino flavour fluxes, shown in Equation 1.2.

$$\phi_{CC} = \phi(\nu_e) = 1.76 \pm 0.0110^{-8} \text{ cm}^{-2} \text{ s}^{-1} \quad (1.2)$$

$$\phi_{ES} = \phi(\nu_e) + 0.15(\phi(\nu_\mu) + \phi(\nu_\tau)) = 2.39 \pm 0.2610^{-8} \text{ cm}^{-2} \text{ s}^{-1}$$

$$\phi_{NC} = \phi(\nu_e) + \phi(\nu_\mu) + \phi(\nu_\tau) = 5.09 \pm 0.6310^{-8} \text{ cm}^{-2} \text{ s}^{-1}$$

From Equation 1.2, you can see that the flux of muon and tau neutrinos from the Sun is  $\phi(\nu_\mu) + (\nu_\tau) - \phi(\nu_e) = 3.33 \pm 0.62$ , which is 3 times that of the value for  $\phi(\nu_e)$ , and as we know that the Sun only produces  $\nu_e$ , it means that the neutrinos must be changing flavour to mu and tau neutrinos on their journey to Earth, definitely solving the solar neutrino problem.

#### 1.1.4 Neutrino Oscillation Theory

In 1957, Bruno Pontecorvo postulated that neutrinos could transition from neutrinos to antineutrinos and vice versa (similarly to how two kinds of neutral

kaons  $\bar{K}_0$ , and  $K_0$  were found to oscillate) [24]. Neutrino flavour oscillation theory was then developed by Maki, Nakagawa and Sakata in 1962. The PMNS matrix (Pontecorvo-Maki-Nakagawa-Sakata matrix) is the neutrino analogue of the Cabibbo-Kobayashi-Masakawa quark mixing matrix [18], and is shown in Equation 1.3 as ( $U_{\alpha i}^*$ ) which shows the relationship between the mass and flavour eigenstates for a neutrino with a definite flavour of  $\alpha$  and a definite mass of  $m_i$ .

$$\begin{aligned} |\nu_\alpha\rangle &= \sum_i U_{\alpha i}^* |\nu_i\rangle & (1.3) \\ |\nu_i\rangle &= \sum_\alpha U_{\alpha i} |\nu_\alpha\rangle \end{aligned}$$

In Equation 1.3, the terms  $U_{\alpha i}^*$  and  $U_{\alpha i}$  are the complex conjugate and normal PMNS matrix. Equation 1.4 shows the 3x3 form of the PMNS matrix, where  $c_{ij} = \cos\theta_{ij}$  and  $s_{ij} = \sin\theta_{ij}$ .

$$\begin{aligned} U &= \begin{pmatrix} 1 & 0 & 0 \\ 0 & c_{23} & s_{23} \\ 0 & -s_{23} & c_{23} \end{pmatrix} \begin{pmatrix} c_{13} & 0 & s_{13}e^{-i\delta_{CP}} \\ 0 & 1 & 0 \\ -s_{13}e^{i\delta_{CP}} & 0 & c_{13} \end{pmatrix} \begin{pmatrix} c_{12} & s_{12} & 0 \\ -s_{12} & c_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix} \\ &= \begin{pmatrix} c_{12}s_{13} & s_{12}c_{13} & s_{13}e^{-i\delta_{CP}} \\ -s_{12}c_{23} - c_{12}s_{13}s_{23}e^{i\delta_{CP}} & c_{12}c_{23} - s_{12}s_{13}s_{23}e^{i\delta_{CP}} & c_{13}s_{23} \\ s_{12}s_{23} - c_{12}s_{13}c_{23}e^{i\delta_{CP}} & c_{12}s_{23} - s_{12}s_{13}c_{23}e^{i\delta_{CP}} & c_{13}c_{23} \end{pmatrix} \end{aligned} \quad (1.4)$$

In Equation 1.4, if the  $\sin \delta_{CP}$  terms are not equal to 0, it means that there will be imaginary terms in the matrix, which will contribute to CP (charge-parity) violation. The angles  $\theta_{12}$ ,  $\theta_{23}$  and  $\theta_{13}$  are mixing angles. An additional 3 x 3 matrix term is present if neutrinos are proven to be their own antiparticle

(Majorana particles): it has not been proven whether neutrinos are Majorana or not. Equation 1.5 shows this extra term, where the two Majorana CP-violating phases are given as  $(\alpha_{21}, \alpha_{31})$ .

$$\begin{pmatrix} 1 & 0 & 0 \\ 0 & e^{i\lambda_{21}} & 0 \\ 0 & 0 & e^{i\lambda_{31}} \end{pmatrix} \quad (1.5)$$

### Two-flavour neutrino oscillations in vacuum

If a neutrino is produced in the flavour state  $|\nu_\alpha\rangle$  in a vacuum, using the PMNS matrix for the relationship between the mass eigenstate and the flavour eigenstate, propagating the neutrino according to the time-dependent Schrodinger equation with no potential would yield a solution to this equation as a plane wave, shown in Equation 1.6.

$$i\frac{\partial}{\partial t} |\nu_i(x, t)\rangle = E |\nu_i(x, t)\rangle = -\frac{1}{2m_i} \frac{\partial^2}{\partial x^2} |\nu_i(x, t)\rangle \quad (1.6)$$

The probability of oscillation to flavour state  $|\nu_\beta\rangle$  at time  $t$  is given by Equation 1.7.

$$\begin{aligned} P(\nu_\alpha \rightarrow \nu_\beta) &= |\langle \nu_\beta | \nu_\alpha(t) \rangle|^2 \\ &= \sum_{i,j} (U_{\beta i}^*)(U_{\beta i})(U_{\alpha j})(U_{\beta j}^*) e^{-i(E_i - E_j)t} \end{aligned} \quad (1.7)$$

For a two-flavour oscillation probability calculation, the matrix which transforms a vector in flavour basis into mass basis is simply a 2x2 rotation matrix (where  $\theta$  is the mixing angle) so that:

$$\begin{pmatrix} \nu_\alpha \\ \nu_\beta \end{pmatrix} = \begin{pmatrix} \cos \theta & \sin \theta \\ -\sin \theta & \cos \theta \end{pmatrix} \begin{pmatrix} \nu_1 \\ \nu_2 \end{pmatrix}$$

One could calculate the probability of a neutrino of flavour  $\alpha$  oscillating into a neutrino of flavour  $\beta$  by using the formula in Equation 1.8. Here  $L$  is the distance travelled by the neutrino in kilometres,  $\Delta m^2$  is the difference between the mass eigenstates, and  $E$  is the difference between the energy of the mass eigenstates (where an assumption has been made that the mass eigenstates are the same for both.) The survival probability of a neutrino of flavour  $\alpha$  (i.e. the probability of generating  $\nu_\alpha$  and detecting  $\nu_\alpha$ ) is  $P(\nu_\alpha \rightarrow \nu_\alpha) = 1 - P(\nu_\alpha \rightarrow \nu_\beta)$  for two flavour oscillations.

$$P(\nu_\alpha \rightarrow \nu_\beta) = \sin^2(2\theta) \sin^2 \left( 1.27 \Delta m^2 \frac{L}{E_\nu} \right) \quad (1.8)$$

### Neutrino oscillation matter effects

The presence of dense matter affects the vacuum neutrino oscillation probability, since neutrinos interact via the weak interaction with matter. This is called the Mikheyev-Smirnov-Wolfenstein (MSW) effect: due to ordinary matter containing only electrons, and muons and tauons, the charge current interactions only affects electron neutrinos and creates an additional potential which affects the vacuum oscillation probability for electron neutrinos, shown in Equation 1.9, where  $G_F$  is the Fermi coupling constant, and  $n_e$  is the number density of electrons in the matter, where the positive sign applies to neutrinos, and the negative sign applies to anti-neutrinos.

$$V_{CC} = \pm \sqrt{2} G_F n_e \quad (1.9)$$

Replacing  $\Delta m^2$  with  $\Delta m_M^2$  and  $\sin^2 \theta$  with  $\sin^2 \theta_M$  gives the two flavour



Figure 1.3: Feynman diagrams of charged-current (left) and neutral-current (right) neutrino interactions

oscillation probability modified to take into account the MSW effect (shown in Equation 1.10, where the factor of  $A = 2|V_{CC}|E$ ).

$$\begin{aligned}\sin 2\theta_M &= \frac{\sin 2\theta}{\sqrt{\sin^2 2\theta + (\cos 2\theta - A/\Delta m^2)^2}} \\ \Delta m_M^2 &= \Delta m^2 \sqrt{\sin^2 2\theta + (\cos 2\theta - A/\Delta m^2)^2}\end{aligned}\quad (1.10)$$

### 1.1.5 Neutrino-nucleus interactions in Super-Kamiokande Gd

Understanding neutrino interaction modes, and understanding neutrino nucleus interaction modes, in particular the neutral current quasielastic reaction is key to understanding this analysis.

There are two main types of neutrino interaction: charged-current (CC) and neutral current (NC). The former occurs when a  $W^\pm$  boson is used in a nuclear exchange, and the latter occurs when a  $Z^0$  is used (see Figure 1.3).

There are multiple types of neutrino-nucleus interactions that occur which can be either charged current or neutral current interactions, or both.

One such interaction, and often the simplest interaction, is elastic scattering



Figure 1.4: Feynman diagrams of neutral-current (left) and charged-current (right) neutrino-electron scattering

$(\nu + e^- \rightarrow \nu + e^-)$ , which occurs when a neutrino scatters off an electron with a virtual vector boson being exchanged. This type of scattering is used in the detection of low energy neutrinos, primarily those from the sun. Figure 1.4 shows the Feynman diagrams for this kind of interaction.

Single mesons can also be produced via neutrino-nucleon reactions: these are mostly pions, but some kaons and eta particles can also be produced. Here a neutrino with a high enough energy interacts with and excites a nucleon, producing a resonant baryon which decays to a nucleon and a single pion (shown in Equation 1.11), where  $N$  and  $N'$  are nucleons.

$$\nu_l + N \rightarrow l + N^* \quad (1.11)$$

$$N^* \rightarrow \pi + N'$$

The resonant baryon produced during the reaction is usually a  $\Delta(1232)$  resonance.

Single pion final states can also be produced by a neutrino which scatters off

an entire nucleus (X), shown in Equation 1.12 for the charged current reactions and Equation 1.13 for neutral current reactions.

$$\nu_l + X \rightarrow l^- + X + \pi^+, \quad \bar{\nu}_l + X \rightarrow l^+ + X + \pi^- \quad (1.12)$$

$$\nu_l + X \rightarrow \nu_l + X + \pi^0, \quad \bar{\nu}_l + X \rightarrow \bar{\nu}_l + X + \pi^0 \quad (1.13)$$

At higher energies (above 1 GeV), neutrino interactions can also produce kaons in the final state, due to the higher energies being able to produce strange quarks.

Deep inelastic scattering is a type of neutrino interaction where the neutrino scatters off a quark inside the proton or neutron involved in the exchange, via a W (CC) or Z (NC) boson (Equation 1.14), shown in Figure 1.5.

$$\nu_l + N \rightarrow l^- + X, \quad \bar{\nu}_l + N \rightarrow l^+ + X(\text{CC}) \quad (1.14)$$

$$\nu_l + N \rightarrow \nu_l + X, \quad \bar{\nu}_l + N \rightarrow \bar{\nu}_l + X(\text{NC})$$

The next two neutrino-nucleus interactions explained here are of particular relevance to the analysis in this thesis: inverse beta decay (IBD) and quasi-elastic scattering.

Inverse beta decay is the reaction by which Cowan and Reines first detected electron antineutrinos: it is important at low energies: from the minimum energy for the reaction to take place ( $E_\nu = 1.806$  MeV) to tens of MeV. Diffuse Supernova Neutrino Background and low energy antineutrinos produced from nuclear reactors can be detected via this process. Figure 1.6 shows the Feynman diagram for this reaction. The neutron produced by this reaction is integral to the motivation behind the Gadolinium-doping upgrade to Super-Kamiokande,



Figure 1.5: Feynman diagram for a charged current deep inelastic scattering interaction with an incoming muon neutrino

which will be explained in the next section.

Finally, we get to the type of interaction investigated in this thesis: quasi-elastic scattering. This makes up the majority of the neutrino-nucleus interaction cross-sections at the energy range from 100 MeV to 2 GeV, and therefore vital for the study of neutrinos from long baseline neutrino experiments and also low energy atmospheric neutrinos. Equation 1.15 shows the equations for both the charge current (CCQE) and neutral current (NCQE) version of this interaction where the incoming neutrino scatters off a nucleon.

$$\text{CC : } \nu(k) + n(p) \rightarrow l^-(k') + p(p') \quad (1.15)$$

$$\bar{\nu}(k) + p(p) \rightarrow l^+(k') + n(p')$$

$$\text{NC : } \nu(k) + N(p) \rightarrow \nu(k') + N(p')$$

$$\bar{\nu}(k) + N(p) \rightarrow \bar{\nu}(k') + N(p')$$

As this analysis relates to the various properties of neutrons produced from



Figure 1.6: Feynman diagram for the inverse beta decay reaction

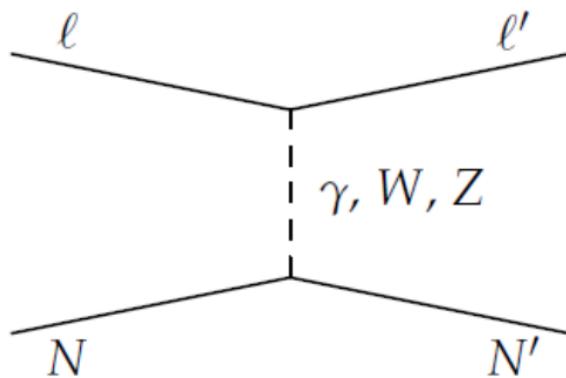


Figure 1.7: Feynman diagram for a quasi elastic scattering interaction off a nucleon

NCQE interactions in Super-Kamiokande Gd, it is worth looking at the interactions where neutrons are produced inside the detector medium. In Water Cherenkov detectors, the incoming neutrino or anti-neutrino interacts with an  $^{16}\text{O}$  nucleus, and at neutrino energies greater than 200 MeV, a nucleon is knocked out, as shown in Equation 1.16.



This interaction happens in stages, where there is the initial reaction between the incoming neutrino and the specific nucleon inside the  $^{16}\text{O}$  nucleus, hadronic final state interactions (FSI) inside the nucleus itself, and subsequent hadronic secondary interactions (SI) inside the detector medium. These are outlined in the schematic shown in Figure 1.8, with the primary interaction, taken as the CCQE reaction  $(\bar{\nu}(k) + p(p) \rightarrow l^+(k') + n(p'))$ . There are stages to neutrino-nucleus reactions, and for the CCQE reaction shown in Figure 1.8, there are three processes: the initial neutrino-nucleon reaction  $(\bar{\nu}(k) + p(p) \rightarrow l^+(k') + n(p'))$ , the final state interaction (FSI) inside the  $^{16}\text{O}$  nucleus  $(p + p \rightarrow p + n + \pi^+)$ , and the secondary (SI) interaction in the detector medium  $(n + ^{16}\text{O} \rightarrow n + n + ^{15}\text{O}^*)$ . All interactions inside the nucleus are referred to as hadronic final state interactions, and when they move through the detector medium and leave the target nucleus they are known as hadronic secondary interactions.

### 1.1.6 Diffuse Supernova Neutrino Background

A key feature of the analysis presented in this thesis is that it is an investigation into the significant background of the signal for SRN. It is therefore important to state and understand the process behind the production of supernova relic

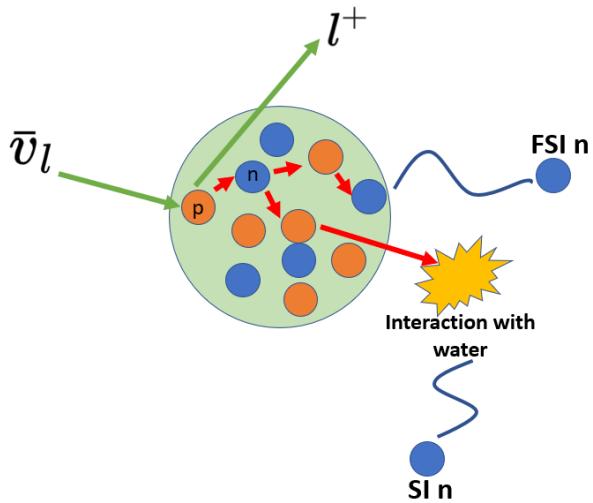


Figure 1.8: Schematic of neutron production following the interaction of neutrino with a nucleon inside an  $^{16}\text{O}$  and the resulting FSI and SI interactions in the detector medium.

neutrinos in order to get a firm handle on the motivation behind this analysis.

### Supernovae Classification

Supernovae occur when a star with a mass around eight times the mass of our sun explodes, and in a galaxy these occur only a few times in a century. Supernovae are classified into different types: Type 1a, Type 1b, Type 1c and Type II. The classification of supernovae are determined by looking at the spectral lines in the light emitted from these supernovae. Table 1.1.6 shows how these supernovae are classed and which spectral line elements are associated with each class.

Supernova Classification	Element lines present in spectra
Type 1a	No hydrogen, silicon
Type 1b	No hydrogen, no silicon, helium
Type 1c	No hydrogen, no silicon, no helium
Type II	Hydrogen

The kinetic energy of a supernova is  $10^{44}$  J, and 99 % of the energy from core-collapse supernovae (CCSN) are released in the form of neutrinos. Unlike Type 1a supernovae which are usually thermonuclear supernovae, Type 1b, Type 1c and Type II are core-collapse supernovae, from which more neutrinos are emitted which is why these types of supernovae are of more interest.

### Core-Collapse Supernovae Mechanism

Using the pressure produced by the process of nuclear fusion, a star is able to support itself against gravitational collapse. During the proton-proton chain reaction, hydrogen will fuse to produce helium and once temperatures and pressures are high enough, helium fusion will occur. After all the helium in the core is used up in the fusion process, the star will contract until the pressure and temperatures get even higher, allowing more massive nuclei to fuse. This reaction will carry on until iron nuclei are produced, this being the element with the highest binding energy, causing the fusion to stop.

As more and more iron accumulates in the core of the star, the density and temperature of the core will increase, and these higher energy electrons will increase the rate of electron capture on protons that will occur in the iron nuclei. This will cause a reduction in the electron degeneracy pressure, which is further enhanced by the breakdown of the iron nuclei which occurs at higher temperatures when gamma rays interact with them. The degeneracy pressure is no longer greater than the gravitational forces acting inwards, and gravitational core collapse occurs.

The point at which in the core the neutrino mean free paths become approximately the same size as the proto-neutron star is called the “neutrinosphere”. Figure 1.9 shows the formation of a proto-neutron star (PNS) and the neutrino flux and neutrinosphere produced by the core-collapse mechanism.

The radius of the neutrinosphere becomes as large as that of the inner core



Figure 1.9: Schematic showing core collapse of a star, and the resultant shockwave and neutrino production

of the star, when the inner core of the star reaches a density of  $10^{11} g cm^{-3}$ , and the electron neutrinos produced from electron capture become unable to escape. Gravitational collapse of the star continues until the inner core reaches nuclear density, at which point a shock wave is produced due to the repulsive force between nuclei.

When this shock wave reaches the neutrinosphere, neutrino emission begins, which lasts less than 10 milliseconds. After this shockwave passes, nucleons and electrons fall back onto the proto-neutron star which heats it up. This causes neutrinos of all flavours to be produced via pair production and electron capture. This is called the “accretion phase”. Due to this expulsion of neutrinos the shock wave loses energy, but it is revived through matter behind the shock wave being heated by neutrino absorption from the proto-neutron star region. The mechanism for this shown in Figure 1.10, where the shockwave produced

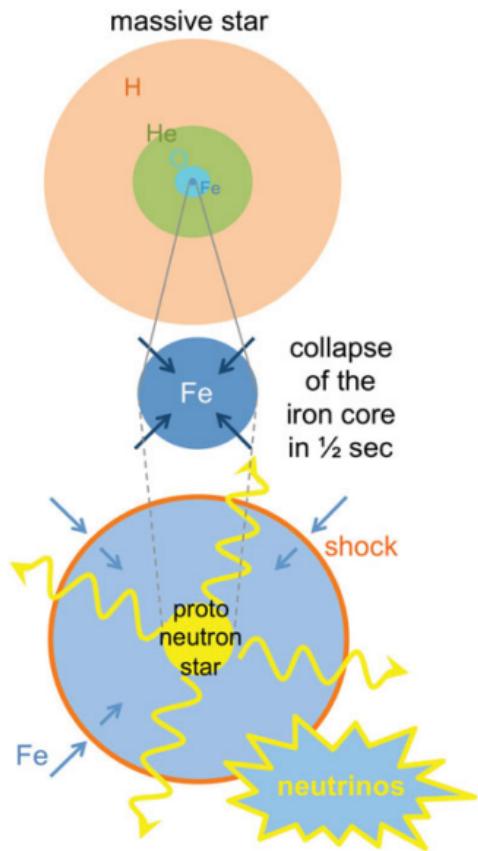


Figure 1.10: Schematic showing core collapse of a star, and the resultant shockwave and neutrino production

depends on antineutrinos being absorbed by the dense post-shock gas.

After the shock wave is revived, if it has enough energy to blow off the outer layer of matter a supernova occurs. Then, depending on the mass of the PNS, it cools to either become a neutron star or a black hole. If the shock wave energy is not high enough to blow off the outer layer of matter, the accretion phase continues until a black hole is formed.

The energy of the emitted neutrinos depend on their flavour - neutrinos emitted from a deeper layer inside the supernova will be higher in temperature and therefore have a higher energy. For electron neutrinos and electron anti-neutrinos, the dominant interactions are charged current interactions with nucleons. Due to the number of neutrons in a proto-neutron star greatly outnumbering the number of protons, the interaction involving  $\nu_e$  will be far more efficient than those involving  $\nu_{\bar{e}}$ , meaning the neutrinosphere for  $\nu_{\bar{e}}$  is smaller than for  $\nu_e$ , so they are emitted from the PNS with greater energies.

Due to there being no charged current interactions involving muon neutrinos in the medium, they are involved instead in neutral current reactions, including Bremsstrahlung, neutrino-pair annihilation, and electron-positron pair annihilation. Due to only undergoing these reactions their neutrinosphere is even smaller than that of their electron neutrino and electron anti-neutrino counterparts, therefore being emitted with even higher energies [21]. Figure 1.11 shows the luminosity (top panels) and average energy (bottom panels) for three different neutrino flavours as a function of time for the neutronisation phase (left), accretion phase (centre) and cooling phase (right). Taken from [9].

Kamiokande-II observed neutrinos produced from a supernova, later named 1987A, in the Large Magellanic Cloud. These neutrinos were also observed by the Baskan Neutrino Observatory and the Irvine-Michigan-Brookhaven detector (IMB) [16]. The neutrinos detected from 1987A were detected via the inverse beta decay (IBD) interaction (shown in Figure 1.6), as opposed to the electron elastic scattering, CC interaction with  $^{16}\text{O}$  or NC interaction with  $^{16}\text{O}$  channels. The positron from this reaction is what is searched for, and these positrons are produced isotropically within the detector. A neutrino burst search from SN1987A was carried out on the data run which ran continuously between 16:09 21 February 1987 to 07:31 24th February 1987 JST [17].



Figure 1.11: Luminosity (top panels) and average energy (bottom panels) for three different neutrino flavours as a function of time for the neutronisation phase (left), accretion phase (centre) and cooling phase (right). Taken from [9].

The event selection criteria for a supernova neutrino event were as follows: firstly, the total number of photoelectrons per event should be less than 170, as this corresponds to an electron with an energy of 50 MeV, and the neutrinos released from a supernovae have a typical energy of the order of 10 MeV. Secondly, the total number of electrons in the outer detector per event needed to be less than 30, in order for the event to be considered fully contained, and finally, the time between previous events has to be greater than 20  $\mu$ sec, in order to avoid electrons produced from Michel electrons. Figure 1.12 shows the time-sequence of all low-energy events. If the neutrinos were to show up in Kamiokande-II, they would arrive in a cluster lasting of order 10s, so intervals of 10 seconds were used to search for clusters of events. In order to classify the event as a supernova neutrino candidate, the event had to be registered by atleast 30 PMTs, and 6 events were found in the first 10s interval which satisfied this, and 2 events in the next 10s interval. A second search was performed on a larger data sample of 42.9 days, from 9th of January to 25th of February 1987, and no other supernova neutrino burst candidates were found. The first two events shown in Figure 1.12 also correlate in angle to the direction of 1987A. As a result, this was considered detection of a genuine supernova neutrino burst, and was the first direct observation of this in neutrino astronomy.

Due to neutrinos carrying away 99% of the energy produced by supernovae, they are the most important signal a supernova can produce, and can be used for finding out when the supernova explosion actually happened - due to neutrinos leaving the supernovae immediately while gamma rays from the the explosion can only leave the supernova when the shock wave reaches the surface of the star [7].

The supernova relic neutrinos emitted from all past CCSN are accumulated and form an everpresent background due to the fact that neutrinos only weakly



Figure 1.12: Events detected centered on 16:35:35 JST, where the solid lines are the number of hit PMTs and dashed lines are muon events. The zoomed in plot in the top-right shows an expanded version of the 0.0 - 2.0 second time scale.

interact with matter. The total flux of this neutrino background has been theoretically predicted and is called the “diffuse supernova neutrino background” (DSNB). When calculating the total DSNB flux, the redshift caused by the expansion of the universe needs to be taken into account.

Equation 1.17 gives the differential of the DSNB flux with respect to the DSNB energy at Earth.

$$dn_\nu(E_\nu) = R_{\text{SN}}(z) \frac{dt}{dz} dz \frac{dN_\nu(E'_\nu)}{dE'_\nu} (1+z)dE_\nu \quad (1.17)$$

Here  $E'_\nu = (1+z)E_\nu$  is the energy of neutrinos at a certain redshift  $z$ ,  $R_{\text{SN}}(z)$  is the supernova rate at a comoving volume at  $z$ ,  $dN_\nu/dE_\nu$  is the number spectrum of neutrinos emitted by one supernova explosion, and  $(1+z)^{-3}$  is the expansion of the universe factor. The relationship between redshift and

time ( $dt/dz$ ), is given by the Friedmann equation (Equation 1.18), where  $H_0$  is the Hubble constant,  $\Omega_m$  is the matter density and  $\Omega_\Lambda$  is the cosmological constant [5].

$$\frac{dz}{dt} = -H_0(1+z)\sqrt{\Omega_m(1+z)^3 + \Omega_\Lambda} \quad (1.18)$$

It is worth reviewing the lastest searches from Super-Kamiokande for DSNB neutrinos, and the limit on the flux prediction from different Super-Kamiokande phases. While Figure 1.13 shows that the majority of DSNB neutrinos have an energy below 10 MeV, due to massive backgrounds these DSNB are hard to detect, so Super-Kamiokande therefore searches for the DSNB signal in the order of 10 MeV. As mentioned previously the main detection channel is inverse-beta decay, with the prompt signal being the emitted positron and the delayed signal constituting of the photon emission from neutron capture. Up to and including now, there has been no DSNB signal evidence but upper-limits on the flux have been set from both KamLAND and from Super-Kamiokande phases I-IV, shown in Figure 1.14.

Regarding important backgrounds to the DSNB signal, below neutrino energies of 15 MeV NCQE interactions from atmospheric neutrinos form the main background to DSNB searches. The importance of this background motivates the analysis in this thesis: by measuring the NCQE interaction using T2K beam neutrinos, the large uncertainty of this interaction in the low energy region can be reduced due to the reason that the atmospheric neutrino energy region is similar to that of the T2K beam energy region. The differential cross section of the neutral-current elastic interaction of a neutrino on a free proton or neutron can be written as shown in Equation 1.19.



Figure 1.13: DSNB flux against neutrino energy for various theoretical models, "NH" and "IH" are the normal and inverted mass hierarchies respectively



Figure 1.14: The 90% C.L. for observed and expected upper limits on the electron antineutrino flux from SK-I - SK-IV and KamLAND, with the theoretical predictions in Figure 1.13 shown in grey.

$$\frac{d\sigma^{\nu(\bar{\nu})}}{dq^2} = \frac{M^2 G_F^2}{8\pi E_\nu^2} \left\{ A(q^2) \pm B(q^2) \frac{s-u}{M^2} + C(q^2) \frac{(s-u)^2}{M^4} \right\} \quad (1.19)$$

where  $E_\nu$  is the initial neutrino or antineutrino energy,  $M$  is the mass of the nucleon,  $G_F$  is the Fermi coupling constant,  $s$  and  $u$  are the Mandelstam variables, and  $q$  is the transfer of four momentum between incoming (anti)neutrinos and outgoing (anti)neutrinos. Depending on whether it is an incoming neutrino or antineutrino, there is either a plus or a minus between the A and B terms. A,B and C are terms which are made up of form factors specific to neutral current interactions. The ideal goal of this analysis is to complement prior analyses which investigate the NCQE cross-section and analyses which produce neutron multiplicity distributions for NCQE interactions, with the added benefit of comparisons to Run 11 T2K data, which brings with it the first beam events with gadolinium in the far detector. The emphasis on the distributions of neutrons in this thesis and the methods of neutron tagging used are of increasing interest due to the fact that the main purpose behind adding Gadolinium to the detector is to make these neutron captures more visible.

## Chapter 2

# The Tokai-to-Kamioka experiment

### 2.1 The Tokai-to-Kamioka experiment

The Tokai-to-Kamioka (T2K) experiment is a long baseline neutrino oscillation experiment based in Japan and its purpose is to study neutrino oscillations: specifically a precision measurement of the neutrino oscillation parameters  $\Delta m_{23}^2$  and  $\sin^2 \theta_{23}$  and to improve the measurement of the leptonic CP violating phase  $\delta_{CP}$ , which is mentioned in Chapter 1. The experiment produces a beam of intense muon neutrinos at J-PARC (Japan Proton Accelerator Research Complex) in Tokai, which is located on the east coast of Japan in Ibaraki Prefecture. The muon neutrino beam travels 295 km west towards the far detector Super-Kamiokande (see Chapter 3.) The neutrino beam is measured by other detectors such as ND280 and INGRID, before they oscillate and reach Super-Kamiokande which is important with regards to measuring the neutrino oscillation parameters. Both ND280 and Super-Kamiokande are off-axis detec-

tors, meaning that they are located  $2.5^\circ$  off axis from the centre of the neutrino beam. This results in the peak of the energy of the muon neutrinos to be 0.6 GeV, maximising the neutrino oscillation on the 295 km baseline. The spread in muon neutrino energy is reduced, this means that these detectors are far less susceptible to potential backgrounds. This chapter explains production of the muon neutrino beam from the JPARC proton beam line and the near detector complex.

### 2.1.1 Neutrino beam production

#### JPARC proton beam production

A sequence of three accelerators results in the production of a proton beam from the JPARC accelerator complex, these are the LINAC (linear accelerator), an RCS (rapid cyclic synchrotron) and the main ring synchrotron (MR). A negative hydrogen ion is accelerated to a kinetic energy of 400 MeV, from which a beam of protons is created by converting the negative hydrogen ion beam using charge-stripping foils. This proton beam is then accelerated to a kinetic energy of 3 GeV by the RCS, and about 5% of the bunches produced from this process are passed to the MR where the proton beam will be accelerated up to 30 GeV.

#### Neutrino beam production

The neutrino beam is produced using the primary and secondary beamline as shown in Figure 2.2. The primary beamline takes the proton beam from the MR and directs it towards Kamioka. The beam is then transferred through a succession of beam monitors which measure facets of the proton beam including the beam profile, intensity and position. The beam monitor which is closest to the graphite target measures the "Protons-On-Target" (POT), a value used to

determine the neutrino beam flux. The secondary beamline takes the proton bunches and passes them through a target station, the decay volume and the beam dump. After interacting with the target station, the proton bunches are collimated through a 1.7 m graphite rod where the collimated hole the proton bunches pass through are 30 mm in diameter. Beam profile reconstruction occurs in the OTR (Optical Transition Radiation) monitor, made up of titanium alloy foil placed at a 45 degree angle in order to intercept the beam. As the proton beam enters the foil, the visible light that is produced escapes through a collection of mirrors and is then captured by a charge injection device camera, which creates the beam profile.

After beam profiling, the proton beam then impacts upon a graphite rod target which is 91.4 cm long and 2.6 cm in diameter - this collision produces secondary hadrons, including pions which are focused by three magnetic horns (shown in Figure 2.1). These magnetic horns can be used to produce either a muon neutrino or muon antineutrino beam depending on the polarity of the 250kA current they are pulsed with. If a +250 kA current is used, the positive pions and kaons produced can go on to make muon neutrino beams whereas if a -250 kA current is used negative pions and kaons can decay to create muon antineutrino beams (both are shown in Equation 2.1). The +250 kA mode is called Forward Horn Current (FHC) mode and the -250 kA mode is called Reversed Horn Current (RHC) mode. The analysis in this thesis will occur in FHC mode only.

$$\begin{aligned} \pi^+ &\longrightarrow \mu^+ + \nu_\mu & \text{FHC} \\ \pi^- &\longrightarrow \mu^- + \bar{\nu}_\mu & \text{RHC} \end{aligned} \tag{2.1}$$

A 75 ton volume beam dump made of graphite and iron stops the particles, specifically the protons, secondary hadrons and mesons which have a momentum below 5 GeV/c, and therefore end up being absorbed by the beam dump. A

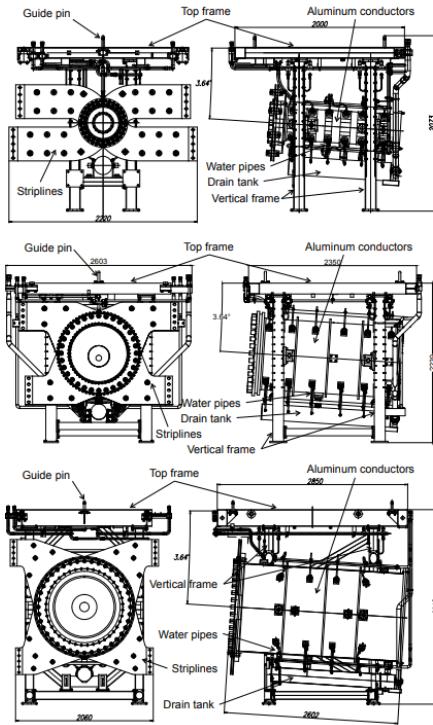


Figure 2.1: Front and side (left and right) view drawings of the three T2K magnetic horns (top, middle bottom).

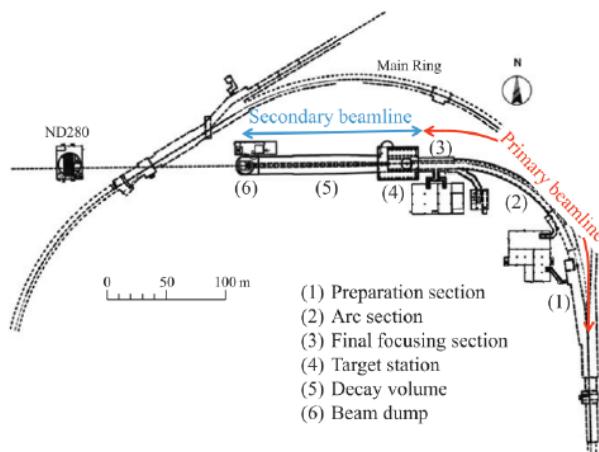


Figure 2.2: Schematic of the neutrino beam line

muon monitor is placed after the beam dump in order to directly measure the beam intensity and beam direction. Muons can be used to monitor the beam properties because along with the neutrinos these are the main particles produced from pion decay. After the muon monitor, a nucleon emulsion plate detector measures the flux and momentum of the muons.

### 2.1.2 Near detectors

#### INGRID detector

The INGRID (Interactive Neutrino GRID) detector is a neutrino detector which unlike ND280 is placed on-axis instead of off-axis. This allows it to directly monitor the direction of the neutrino beam and the intensity of the neutrino beam by measuring the interactions of the neutrinos with the alternating iron plates that make it up. INGRID is also placed 280 m from the graphite target and consists of 14 modules placed in a cross formation, with the centre of the cross placed at the centre of the neutrino beam. The INGRID modules are comprised of nine iron plates alternating with 11 tracking scintillator planes, which are themselves surrounded by scintillator plates the purpose of which is to reject interactions that occur outside the module. A schematic of the INGRID cross is shown in [2.3](#) and a schematic of the modules is shown in [2.4](#).

An additional module, called the proton module was added to measure the muons in combination with protons produced by the neutrino beam in INGRID. This module is used to distinguish the quasi-elastic interaction channel in order to compare it with Monte Carlo simulations of the beamline and neutrino interactions. The Proton Module is made of scintillator planes (no alternating iron plates) and is contained by veto planes. The Proton Module was placed in the centre of the INGRID cross at the intersection of the vertical and horizontal modules. Figure [2.5](#) shows what a standard neutrino event looks like inside the



Figure 2.3: Schematic of the INGRID cross taken from [2].



Figure 2.4: Individual INGRID module schematic taken from [2].



Figure 2.5: INGRID event display showing a typical INGRID event, taken from [1].

INGRID module. A neutrino enters from the left and after interacting with the scintillator cells (shown in green) produces tracks (shown in red), with the relative size of each circle corresponding to the observed signal in that cell. The blue cells show the position of the veto scintillators, while the gray planes show the iron plates.

The hit efficiency in INGRID is calculated using muons: the muon track is reconstructed, but not using the hit information, and then channels in the scintillator plane expected to contain hits from the trajectory of the track are checked. Figure 2.6 shows the hit efficiency as a function of reconstructed track angle measured by cosmic-ray data.

## ND280

ND280 is a near detector which sits 280 m from the target. It is an off-axis detector, meaning that just like Super-Kamiokande, it is placed  $2.5^\circ$  off-axis from the center of the beam. This stems from the relationship between the energy of the neutrinos produced from the decay of the pions, a relationship shown in Equation 2.2.

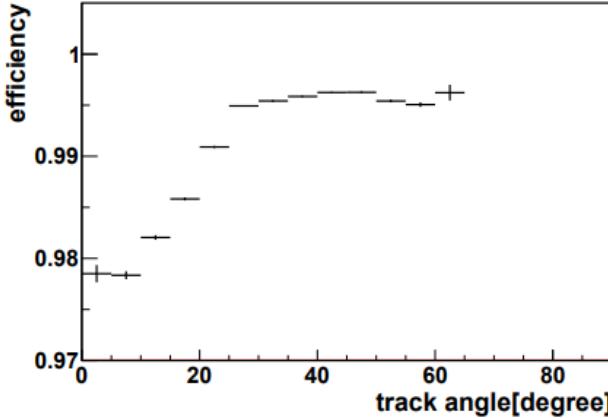


Figure 2.6: Hit efficiency plotted against reconstructed track angle

$$E_\nu = \frac{m_\pi^2 - m_\mu^2}{2(E_\pi - \sqrt{E_\pi^2 - m_\pi^2} \cos \theta)} \quad (2.2)$$

where  $E_\pi$  is the energy of the parent pion and  $\theta$  is the scattering angle between the direction of the outgoing neutrino and the direction of the parent pion's momentum,  $m_{pi}$  is the parent pion mass and  $m_\mu$  is the mass of the outgoing muon. Figure 2.7 shows the neutrino energy from pion decay plotted against the energy of the parent pion for a range of off-axis angles.

As off-axis angle increases, the intensity of the neutrino beam decreases, and therefore picking an off-axis angle of  $2.5^\circ$  at which to place the ND280 detector complex strikes a good balance between keeping a high beam intensity while ensuring a peak energy of 0.6 GeV in order to have the neutrino oscillation be maximised at 295km. The relationship between muon neutrino oscillation probability and muon neutrino energy is shown at the top in Figure 2.8, and at the bottom the muon neutrino flux 295 km away from the graphite target can be seen for three different off-axis angles.

Figure 2.9 shows a schematic of the ND280 detector complex. It has five



Figure 2.7: Energy of the neutrino plotted against the energy of the parent pion for multiple different off-axis angles



Figure 2.8: The probability of survival of muon neutrinos (top plot) and neutrino beam flux at the 295km far detector (bottom) Taken from [2].



Figure 2.9: Near detector ND280 schematic taken from [1].

main goals: firstly, to measure the cross sections of muon neutrino interactions, so neutrino-nucleus interaction models can reduce their systematic uncertainties. Secondly, measuring the component of the neutrino beam which is made of electron neutrinos, hence being able to better constrain the background to electron neutrino appearance at the far detector. Thirdly, to predict the event rate at Super-Kamiokande by measuring the energy spectrum of the muon neutrinos produced. It also aims to study non quasi-elastic processes which produce pions below the Super Kamiokande Cherenkov threshold, and to measure the neutral current (NC) pion-nought production rates.

N2D280 is made of a neutral pion detector (*P0D*), three Time Projection Chambers (TPCs) and two Fine Grained Detectors (FGDs). These are enclosed within Electromagnetic Calorimeters (ECals) and a Side Muon Range Detector (SMRD). These detectors are magnetised using a magnet originally used in the UA1 detector at CERN.

The neutral pion detector is important due to its ability to detect a process

that can imitate the signal event of electron neutrino appearance at Super-Kamiokande. Neutral pions are produced during the neutral current interactions on water ( $\nu_\mu + N \rightarrow \nu_\mu + N' + \pi^0 + X$ ) and the purpose of *P0D* is to measure the cross-section of this interaction. The central part of the *P0D* detector is made of planes of scintillator, brass and water bags which are placed in alternating layers as shown in Figure 2.11. There are also two electromagnetic calorimeters which remove events entering the detector from the outside. The SMRD (Side Muon Range Detector) is used as a way to measure the momenta of muons which escape the detector complex at a large angle relative to the direction of the beam. There are three time projection chambers placed downstream of the neutral pion detector. These are used for reconstruction of charged particle tracks, particle identification and determination of the momentum of the particle. Each individual TPC is comprised of a box filled with an argon based drift gas which is surrounded by an insulating outer box which contains  $CO_2$ . The inward facing walls of the inner box have copper strips with a separation of precisely 11.5 mm, and along with a central cathode produces a uniform electric field. Therefore, charged particles ionise the drift gas as they pass through the TPC which produces electrons which are pushed away from the central cathode and towards one of the readout planes. These electrons are then amplified by an electric field of 27 kV/cm and the signals are read out by MicroMegas detectors. Figure 2.10 shows a schematic drawing of a TPC.

There are two fine grained detectors (FGDs) in ND280, which have two purposes. Firstly to yield a mass for the neutrino interaction, and secondly, to track the charged particles issuing from the neutrino interaction vertex. Each FGD has dimensions of 2300 mm x 2400 mm x 365 mm and consist of 1.1 tons of target material. The FGD contains layers of scintillator bars, with each layer orientated in the x and y directions perpendicular to the neutrino beam which



Figure 2.10: Schematic of a TPC taken from [1].

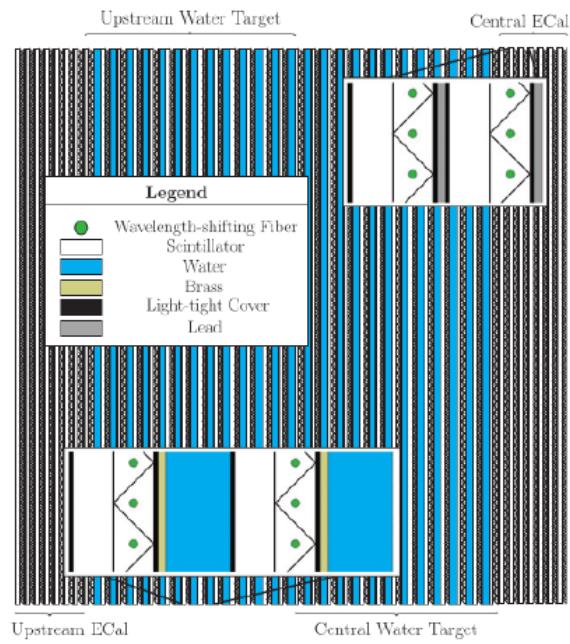


Figure 2.11: P0D schematic taken from [1].

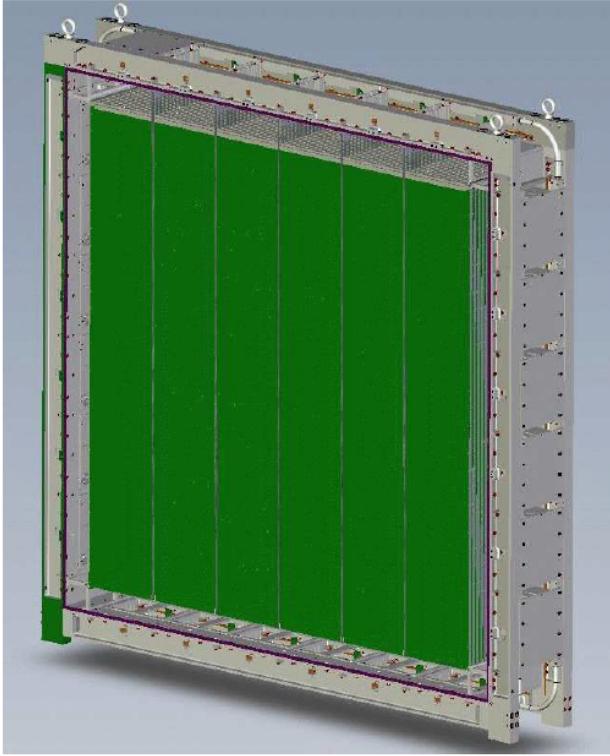


Figure 2.12: FGD with the front cover removed: the X-Y scintillator modules are shown in green.

allows for these charged particles to be tracked. The first FGD is comprised of 15 scintillator modules, while the second FGD has 7 scintillator X-Y modules alternating with 6 2.5cm thick water target modules. Comparing the rate at which neutrinos interact with the first FGD and the second FGD, neutrino interaction cross sections on carbon and on water can be measured. Figure 2.12 shows a view of one of these FGDs.

The ECal is an electromagnetic calorimeter which surrounds the P0D, the TPCs and the FGDs. It consists of plastic scintillator layers sandwiched between lead absorber sheets. Its purpose is to aid in the full event reconstruction through detecting photons and measuring their energy and direction, along with



Figure 2.13: Electromagnetic calorimeter module external view

detecting charged particles and getting information which helps in their identification (separating out electrons, muons and pions.) As well as this, the ECal helps reconstruct pi-zero events produced by neutrino interactions inside the inner detectors. Figure 2.13 shows the external view of one ECal module.

## Chapter 3

# The Super Kamiokande Detector

### 3.1 The Super-Kamiokande detector

#### 3.1.1 Background and detector design

Super-Kamiokande is a neutrino observatory consisting of a cylindrical tank which is 41.4 m in height and 39.3 m in diameter, and filled with 50kton of ultrapure water and gadolinium sulphate. It is used as a neutrino detector for atmospheric, solar and astrophysical neutrinos, as well as being a far detector of the Tokai-to-Kamioka neutrino beam. It is based in the Mozumi mine, located in Gifu Prefecture, Japan. Due to its location being underneath Mount Ikenoyama, 1000m underground, it is shielded as much as possible from the cosmic ray muon detector background. Super-Kamiokande is divided into two concentric cylinder volumes, consisting of the inner detector (ID) and outer detector (OD) using a stainless steel structure which supports the photomultiplier tubes. Tyvek and

black polyethylene terephthalate sheets are mounted on this structure in order to optically separate the inner and outer detector [25].

The inner detector is a cylinder which has a diameter of 33.8 m and a height of 36.2m, and has a fiducial volume of 22.5 ktons of water. The fiducial volume is defined as the region inside a surface drawn 2.00 m from the inner detector wall: using this fiducial volume gives protection against events produced from natural radioactivity in the surrounding rock. It is host to 11,129 photomultiplier tubes which give 40% photocoverage of its inner surface, with the specific photomultiplier model chosen being the hemispherical, 50.8 cm diameter Hamamatsu R3600 model. The 2 m wide outer detector has only 1885 photomultiplier tubes which are mounted on the outside of stainless steel structure, each with a smaller diameter of 20cm, and are either the R1408 or R5912 Hamamatsu model. Each outer detector photomultiplier tube is attached to a 50cm x 50cm wavelength shifting plate, improving the light collection ability in the OD [15].

A detailed schematic of the photomultiplier tubes used in Super-Kamiokande can be seen in Figure 3.1. The photocathodes used in these PMTs are comprised of bialkali antimony-potassium-caesium material (Sb-K-Cs), which give a greater sensitivity to longer wavelengths, giving the greatest spectral response at 360 nm, near the ultraviolet range. This makes these types of photocathodes suitable to use for matching with light sources in the blue region of the visible spectrum, which is the region in which the wavelength of Cherenkov photons lie. The 11 dynodes which are inside the photomultiplier tube are arranged in a “venetian blind” fashion, meaning that the dynodes consist of an assembly of parallel strips. This results in a good collection efficiency of the multiplied electrons and gives decent protection from external magnetic fields. More protection from external magnetic fields is provided by a set of Helmholtz coils which are aligned outside the inner detector, which reduce the ambient



Figure 3.1: Schematic of an ID 50 cm Super-Kamiokande PMT [15].

background geomagnetic field from 450mG to 50mG. This is needed due to the systematic bias that could occur due to the strength and uniform direction of the geomagnetic field affecting the photo-electron trajectories and consequently the photomultiplier tube hit timing.

Radioactivity from radon, uranium and thorium radioisotopes in the Super-Kamiokande tank water and radon in the surrounding air could provide a low energy background to measurements which should be negated. In low-energy analyses, such as the analysis in this thesis, this becomes an even more significant issue. Microbes present in the mine water also present a problem as they cause a reduction in the value of the light attenuation length by a factor of 2.72. Therefore the water used in Super-Kamiokande has to undergo a purification process - the water used is continuously reprocessed at a rate of  $30 \text{ tons } h^{-1}$  in a closed loop system, which is the same method used to purify the water when refilling the tank. The first step of this water purification process is to use a  $1\mu\text{m}$  mesh filter to remove large particulates of impurities from the water. [14].A heat exchanger is then used to cool down the water to a constant

13.0 °C to reduce the dark noise hits of the photomultiplier tubes and the growth of the microbes. A treatment with ultraviolet light is then used to kill any remaining microbes in the water. Radon-free air is then dissolved into the water to aid in the later process of radon removal from the water and a high performance membrane ultra filter (UF) is used to remove organic compounds about 10 nm in diameter from the water. After this step a membrane degasifier removes the dissolved radon, where 30 L of radon reduced air is supplied to the membrane degasifier. The dissolved radon will transfer across the membrane but the water will not, allowing for efficient radon reduction. The inner detector tank water is circulated by injecting the water at the bottom of the tank and extracting it from the top, where the convection currents are able to maintain the temperature in the region within 11m to the bottom of tank, however, outside this region the present temperature gradient causes an asymmetry in the attenuation of light, which is discussed more in Chapter 3.

An air purification system is also installed in Super-Kamiokande in order to reduce the amount of Radon present which could dissolve into the tank and affect measurements. The level of radon activity achieved after this air purification has taken place is  $1 \text{ mBq}/\text{m}^3$ .

The Super-Kamiokande experiment began taking data on 1st April 1996, and due to maintenance was shut down in July 2001, which was phase I of the experiment. A table of the Super-Kamiokande phases is shown in Table 3.1. During the refilling of the tank after maintenance, there was cascade of PMT implosions that occurred on the 12th of November 2001, which were triggered by the implosion of a single photomultiplier tube, due to a microfracture in the neck of the tube. This implosion destroyed about 7,000 of the PMTs and in order to avoid such chain reactions in the future, from 2002 onwards all of the inner detector PMTs were fitted with acrylic covers and fiber-glass reinforced plastic

(FRP) cases. Detector shutdown and rebuild took nine months and Super-Kamiokande resumed data taking with the full number of photomultiplier tubes in July 2006 which marked the beginning of Super-Kamiokande phase III. Super-Kamiokande phase IV began in September 2008 where a new data acquisition system and charge to time (QTC) based electronics with Ethernet (QBEE) was deployed in order to measure arrival times and integrated charge for inner detector and outer detector photomultiplier tube signals. This replaced the ATM (Analogue and Timing Module) which was used in Phase I, II and III of Super-Kamiokande. The improvements in electronics and calibration methods meant that when Super-Kamiokande phase IV started running in September of 2008, electrons with energies as low as 3.5 MeV were able to be detected.

Phase	Phase Period	Total PMT number		FRP case?	Electronic type
		ID (Coverage)	OD		
SK-I	Apr. 1996 - Jul. 2001	11146(40%)	1884	no	ATM
SK-II	Oct. 2002 - Oct. 2005	5182(19%)	1884	yes	ATM
SK-III	Jul. 2006 - Sep. 2008	11129(40%)	1884	yes	ATM
SK-IV	Sep. 2008 - May 2018	11129(40%)	1884	yes	QBEE
SK-V	Jan. 2019 - Jul. 2020	11129(40%)	1884	yes	QBEE
SK-VI	Aug. 2020 - May 2022	11129(40%)	1884	yes	QBEE
SK-VII	Jun. 2022 - Today	11129(40%)	1884	yes	QBEE

Table 3.1: Phases of Super-Kamiokande and main properties of each phase

### 3.1.2 Data aquisition system

As shown in Table 3.1, Phase IV of the experiment marked the beginning of Super-Kamiokande using QTC-based Electronics with Ethernet (QBEEs). Each QBEE board used in Super-Kamiokande contains 24 photomultiplier tube input channel, where each channel uses a charge-to-time (QTC) converter and a time-to-digital converter placed in series as shown in Figure 3.2. After a single

photoelectron is produced by the photomultiplier tubes after incident photons have been received upon it, the dynodes inside the photomultiplier tube amplify this photoelectron so that each photoelectron that strikes the surface of the dynode produces several more photoelectrons, which are produced from the dynodes as an analogue signal. This signal then enters a charge-to-time converter (QTC), an application specific integrated circuit (ASIC) which was specifically designed for Super-Kamiokande in order to detect photomultiplier tube signals using built-in discriminators and to produce output timing signals whose widths represent the integrated charge of the PMT signal. The QTC used has three input channels per chip, which has three gain ranges (Small, Medium, Large as shown in Figure 3.2). There is a built in discriminator inside the QTC which determines whether the the signal from a PMT is a “hit”. If the PMT signal exceeds the threshold value of this discriminator, the QTC integrates the charge of the signal over the next 400ns, and a square wave pulse is generated whose pulse width is proportional to the intergrated charge of the input signal from the PMT. The next 400ns is used to discharge the integrated charge from the QTC, leading to a total channel dead time of 900 ns due to summing the charge integration and discharge of the QTC.

A time-to-digital converter (TDC) then digitises the output time signal, so the PMT charge information is retained. These digitised outputs are collated by 20 front-end computers which collects all the information from the inner and outer detectors, with each computer taking the PMT hit information from 30 inner detector and 20 outer detector QBEE boards and sorting the PMT hit time information in order of the raw hit time. This information is then sent to “merger” computers, who then produce a full time-ordered list of all PMT hits. These merger computers then apply software triggers to select event



Figure 3.2: Schematic of the charge-to-time converter circuit [23].

candidates, using  $N_{200}$ , a quantity which determines the number of PMT hits in a 200 ns timing window. When the value of  $N_{200}$  surpasses the threshold value of a certain trigger type (whether it is a SLE (Super Low Energy), LE (Low Energy), HE (High Energy), SHE (Super High Energy), OD (Outer Detector) or AFT (After Window) trigger), the trigger is used to select an event candidate. The SLE, LE, HE and SHE triggers roughly define the energy of a certain event, based on the number of hits detected. The OD trigger is used to veto events, and the AFT trigger is of special importance to the analysis in this thesis: it is used to discern when a neutron is produced after a neutrino interaction. The AFT Trigger is issued to take 500  $\mu$ sec of data after SHE trigger in order to identify the positron from the prompt event and the delayed 2.2 MeV gamma released from the neutron capture on the proton during the IBD reaction. Another set of computers acts as an “organiser”: it takes all the information regarding the event candidates from the “merger” computers and writes them onto disks [15].

## 3.2 Event Reconstruction

### 3.2.1 Vertex Reconstruction

For low energy events (events up to 100 MeV), Super-Kamiokande currently uses BONSAI (Branch Optimisation Navigating Successive Annealing Interactions) for event reconstruction. Vertex reconstruction for Super-Kamiokande has undergone changes and improvements depending on the phase of the experiment.

For Phase I of Super-Kamiokande, vertex reconstruction depended on a lattice of test vertices with 4 m spacing throughout the detector, with a specific measure of goodness for each test vertex: the test vertex with the highest measure of goodness would have around it a more finely spaced grid, and the process would be repeated. For Phase II of Super-Kamiokande due to the reduced number of PMTs, this approach was no longer as successful as it was in Phase I and as a result the reconstruction performance declined, and BONSAI was created as a replacement. Instead of using a fixed grid which was the case with SK-I and SK-II, BONSAI creates test vertices by selecting groups of four PMT hits and seeing where the timing residuals of the PMT hits would be most reduced. After these test vertices have been identified, a maximum likelihood fit over all the PMT hits in the event is performed, shown in Equation 3.1.

$$\mathcal{L}(\vec{x}, t_0) = \sum_{i=1}^{N_{\text{hit}}} \log(P(t - t_{\text{tof}} - t_0)) \quad (3.1)$$

where  $(\vec{x}, t_0)$  is the test vertex, and  $(P(t - t_{\text{tof}} - t_0))$  is the probability density function of the timing residual, which for each PMT hit is defined as  $(t - t_{\text{tof}} - t_0)$ , where  $t_0$  is the time of the interaction,  $t_{\text{tof}}$  is the time of flight from the interaction vertex position to the position of the hit PMT,  $t$  is the PMT hit time. The  $1\sigma$  difference between the true vertex position and the reconstructed



Figure 3.3: Probability density of the timing residual  $P(t - t_{\text{tof}} - t_0)$ , where  $t_0$  is used for the vertex reconstruction maximum likelihood fit. The peaks at 30ns and 100ns are caused by PMT after-pulsing.

vertex position plotted as a function of true electron energy is shown in Figure 3.3.

### 3.2.2 Direction Reconstruction

Cherenkov light is emitted in a conical formation as electrons and positrons travel through water, with a Cherenkov angle of  $\approx 42^\circ$ . BONSAI can reconstruct the direction of these particles by using this information along with the reconstructed vertex. This reconstruction occurs using a maximum likelihood function defined in Equation 3.2.

$$\mathcal{L}(\vec{d}) = \sum_i^{N_{20}} \log(f(\cos \theta_i, E)) \times \frac{\cos \theta_i}{a(\theta_i)} \quad (3.2)$$



Figure 3.4: The vertex resolution (the point at which 68% of the events in the distance distribution between the actual and reconstructed vertex are contained) for the different SK phases. SK-I (Blue), SK-III (Red), SK-IV (Black).

$f(\cos \theta_i, E)$  is the expected distribution of the angle between the vector of the direction  $\vec{d}$  of the particle, and the observed Cherenkov photon from the position of the reconstructed vertex. The reason there is a spread in this energy distribution is because while the highest value of this distribution occurs at the cosine of the opening Cherenkov angle of  $42^\circ$ , due to the particle travelling through the water being Coulomb scattered multiple times, there is a variation in the angle because of the varying particle energy.  $N_{20}$  is the number of hits whose residual hit time is within 20 ns of the time of the reconstructed event, which is used in order to reduce the amount dark noise and scattered photons contribute to the direction reconstruction calculation. The variable  $a(\theta_i)$  is used in the second term in Equation 3.2, and it is linked to the angle of incidence of the photon on the PMT, and is a correction factor stemming from the acceptance of PMTs and therefore linked to the shape of the PMT and its acrylic case.

### 3.2.3 Energy reconstruction

The kinetic energy of a particle is proportional to the amount of Cherenkov photons emitted from it, and if we assume that the Cherenkov photons in a single event come from a single electron, we can reconstruct the total energy of the electron. Instead of using the number of photoelectrons of all hit photomultiplier tubes to reconstruct the energy of low energy events, the number of hit photomultiplier tubes is used instead. The reasons for this are threefold - firstly, low energy events emit a small number of Cherenkov photons, and therefore average about one photon per hit PMT. Secondly, at single photoelectron level, the resolution of the PMTs is poor, and third, the number of photoelectrons produced is related to the gain of the photomultiplier tubes, which is given by equation:

$$G(i) \propto \frac{Q_{obs}(i)}{N_{obs}(i)} \quad (3.3)$$

where  $G(i)$  is the gain of each PMT and  $Q_{obs}(i)$  is the average charge for each inner detector photomultiplier tube, and  $N_{obs}(i)$  is the number of times that photomultiplier tube  $i$  registers a charge which is greater than the threshold charge value. Due to the variation in gain value not affecting the number of hit photomultiplier tubes as much as it does for the number of photoelectrons, number of hit PMTs is used instead. Energy reconstruction uses  $N_{50}$ , which is the number of photomultiplier tube hits in a 50 ns window, which allows for the rejection of dark noise hits for the photomultiplier tubes. The number of effective photomultiplier tubes which are hit in this timing window of 50ns are summed up, while being weighted with correction factors, shown in Equation 3.4 [26].

$$N_{eff} = \sum_{i=1}^{N_{50}} \left[ (X_i - \epsilon_{dark} + \epsilon_{tail}) \times \frac{N_{all}}{N_{alive}} \times \frac{1}{S(\theta_i, \phi_i)} \times \exp\left(\frac{r_i}{\lambda}\right) \times G(i) \right] \quad (3.4)$$

where  $X_i$  is the correction factor hits with many photoelectrons. This correction factor is important because if some photomultiplier tubes are hit by multiple photons (for example, if the edge of the fiducial volume is where the event vertex took place). The number of photoelectrons produced by each hit photomultiplier tube is estimated using the occupancy of the eight photomultiplier tubes which surround it. Using the number of hit photomultiplier tubes ( $n_i$ ) and the number of functional photomultiplier tubes that surround the  $i$ -th photomultiplier tube ( $N_i$ ), the formula for  $X_i$  is shown in Equation 3.5.

$$X_i = \begin{cases} \log(1 - n_i/N_i)^{-N_i/n_i} & (n_i < N_i) \\ 3 & (n_i = N_i) \end{cases} \quad (3.5)$$

$\epsilon_{dark}$  in Equation 3.5 is a correction factor for dark noise hits, shown in Equation 3.6, where  $R_{dark}$  is the average value for the dark rate during the run period that the event is in and  $N_{PMTalive}$  is the number of active photomultiplier tubes in the inner detector.

$$\epsilon_{dark} = \frac{N_{PMTalive} \times R_{dark} \times 50 \text{ ns}}{N_{50}} \quad (3.6)$$

$\epsilon_{tail}$  is the correction factor for photomultiplier tube hits which are in the tail end of the 50ns timing window, and is defined in Equation 3.7.

$$\epsilon_{tail} = \frac{N_{100} - N_{50} - N_{alive} \times R_{dark} \times (100 - 50)\text{ns}}{N_{50}} \quad (3.7)$$

$\frac{1}{S(\theta_i, \phi_i)}$  is the inverse of the effective area of the  $i$ th hit photomultiplier tube photocathode, from the direction of the incident photon given by  $(\theta_i, \phi_i)$ .

$G(i)$  is the gain correction for the quantum efficiency of the photomultiplier tubes and  $\exp(\frac{r_s}{\lambda})$  is the correction for water transparency which accounts for the amount of attenuation undergone by the photons in water, where  $\lambda$  is the water transparency measured during the run period which includes the event, and  $r_i$  is the distance between the reconstructed event vertex and the i-th hit PMT.

The average of each  $N_{eff}$  distribution is taken, after producing multiple  $N_{eff}$  distributions with fixed energies using Monte Carlo. These energies are fitted with a polynomial which is a function of the averaged  $N_{eff}$  distribution, so the reconstructed energy is converted from  $N_{eff}$ .

Super-Kamiokande has undergone a lot of transformations over its lifetime, with modifications made to its data aquisition system and methods of event reconstruction with each successive new phase. The biggest change made to the detector is the upgrade to Super-Kamiokande Gd, where in the summer of 2020 0.02% of gadolinium sulphate was added to the detector: a long-awaited next step which would enhance sensitivity to the detection of neutrons emitted in inverse beta decay compared to that in water, something discussed in much more detail in Chapter 6.

## Chapter 4

# Super-Kamiokande Detector Calibration

In order to achieve optimal event reconstruction for physics analyses, calibration of the Super-Kamiokande detector is crucial. For example, when conducting Monte Carlo simulations of certain processes in the detector, facets of the experiment such as properties of the water, photomultiplier tube response and the inner detector and outer detector electronics are all calibrated so that input parameters for the Monte Carlo simulations can be obtained. This chapter will concern itself with the inner and outer detector calibration, including photomultiplier tube and electronics calibration, PMT gain calibration, quantum efficiency determination and hit timing and charge information calibration.

## 4.1 Inner detector calibration

### 4.1.1 PMT High-voltage setting calibration

The high-voltage (HV) setting for all photomultiplier tubes need to be adjusted individually so all the PMTs produce the same amount of charge for a certain light intensity received by them. Placing a light source which distributes light isotropically in the centre of the inner detector to achieve this calibration means that there is no position in the detector from which the inner detector PMTs are equidistant, so each PMT will not receive the same amount of light from the light source. To avoid this problem, a set of 420 pre-calibrated PMTs inside the detector were used, separated into groups relating to their geometrical distance from the HV calibration light source (see Figure 4.1 for their location with respect to the other photomultiplier tubes.)

### 4.1.2 Relative gain calibration

Understanding the timing information from the hit photomultiplier tubes depends on how well the charge from the hit PMT is calculated. To conceive charge calibration, a quantity called photomultiplier tube "gain" must be calculated. "Gain" is the conversion factor from the number of photoelectrons produced by the hit PMT and charge, and calibration of this quantity is what interpretation of very high energy events (TeV scale) rely on. Quantum efficiency is another quantity used for the calibration of low energy physics events (such as detection of solar neutrinos), due to them consisting of single photoelectron (single-pe) hits: it is the ratio of the number of the number of photoelectrons emitted by the cathode to the number of photons that are incident on the photomultiplier tube window. Super-Kamiokande calibration converts this measure of quantum efficiency into "QE" by multiplying the quantum efficiency by the collection effi-

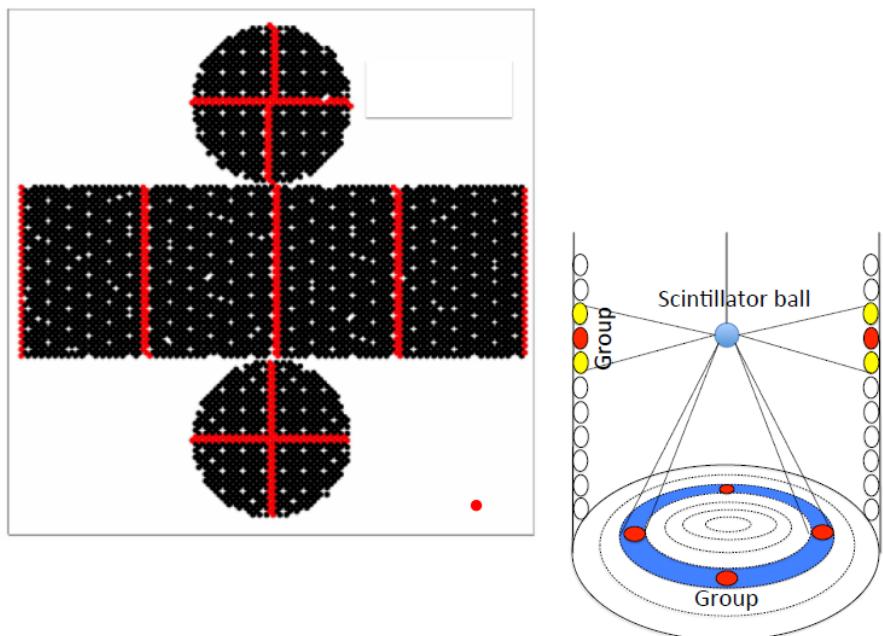


Figure 4.1: Location of 420 reference PMTs used for HV setting calibration. The red lines in show the placement of these PMTs with respect to the others (left). The grouping of these PMTs due to their geometry in relation to the light source is also shown (right)

ciency of the photoelectrons onto the first dynode inside the PMT [3]. Knowing the gain and QE of each PMT in the detector is important in order to accurately measure the output charge from each individual PMT, which is done by first calculating the relative gain difference among all PMTs and then work out the average gain difference over all PMTs in the detector. After this, the variation away from this average gain value can be calculated for each separate inner detector photomultiplier tube, and the gain value for each can be extracted.

The relative gain difference is calculated by two measurements using a light source to produce constant-intensity flashes. The first measurement involves using the light source to produce high-intensity flashes so that all photomultiplier tubes in the detector gets a certain number of photons, and the second measurement has the light source produce low-intensity flashes so that only a few PMTs are hit. The first measurement provides an average charge value ( $Q_{obs}(i)$ ) for each inner detector PMT, while the second measurement gives single photoelectron hits, providing a number of times ( $N_{obs}(i)$ ) that a single PMT gives a charge which is greater than the PMT threshold value. Equation 4.1.2 shows how these two values are calculated from the the high and low intensity flash values ( $I$ ), the acceptance of the PMT(i) ( $a(i)$ ), the QE value of the PMT ( $\varepsilon_{qe}$ ) and the PMT gain  $G$ .

$$Q_{obs}(i) \propto I_{high} \times a(i) \times \varepsilon_{qe}(i) \times G(i) \quad (4.1)$$

$$N_{obs}(i) \propto I_{low} \times a(i) \times \varepsilon_{qe}(i) \quad (4.2)$$

Therefore, by simply dividing these two values of  $Q_{obs}(i)$  and  $(N_{obs}(i))$  the average gain over all PMTs can be calculated. Figure 4.2 shows the spread of the relative gain over all the PMTs.

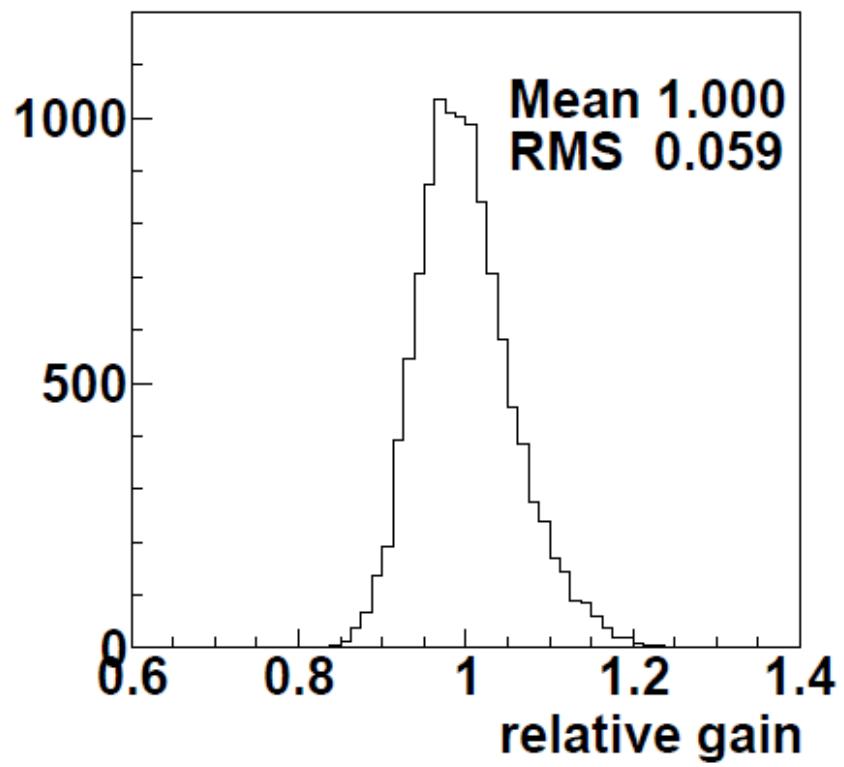
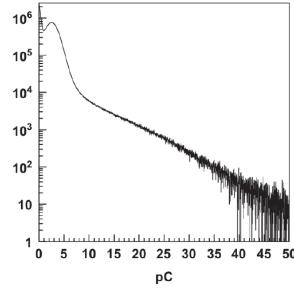


Figure 4.2: Relative gain of PMTs in Super-Kamiokande



(a) Photograph of nickel source



(b) Single p.e distribution of charge in pC

### 4.1.3 Absolute gain calibration

In order to calculate absolute gain, the single photoelectron distribution needs to be measured, this is because absolute gain relates to the observed charge in the photomultiplier tube with the number of photoelectrons produced. A nickel-californium source (shown in Figure 4.3a) is used for this measurement due to it releasing gamma rays isotropically. Thermal neutron capture on nickel is used as the gamma ray source, with the neutrons provided by the spontaneous fission of  $^{252}\text{Cf}$ . This nickel source is placed in the centre of the inner detector and the gamma rays produced are detected by all the inner detector PMTs. On average the observed number of photoelectrons is 0.004 per event per PMT, meaning that single p.e. hits are observed for more than 99% of the hits. The observed charge distribution of all the hits from this nickel source is used to give the average charge, which is used as a conversion factor from a charge measurement in picoColoumbs and single photo-electrons. The factor is 2.658 pC per photoelectron (calculated at the beginning of SK-IV) which is then used to extract the single-p.e. distribution, shown in Figure 4.4.

### 4.1.4 Relative quantum efficiency measurement

To measure the relative quantum efficiency, the gamma rays from neutron capture on the nickel source are simulated. Inside this simulation, a common value

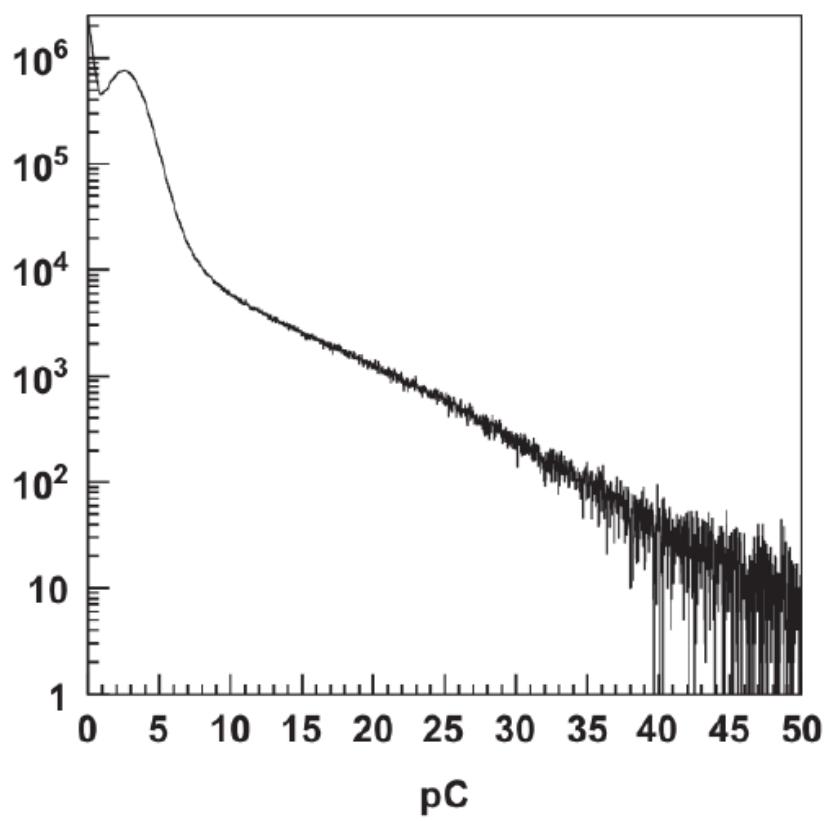


Figure 4.4: Single p.e distribution of charge in pC

of QE is used for all the ID PMTs to predict the number of hits for each PMT. Comparing this number of hits to the actual data obtained for each individual PMT by calculating the ratio between them provides us with a value for relative QE for each inner detector PMT which is then used inside the simulation.

#### 4.1.5 Timing information calibration

Calibrating the hit timing information of the photomultiplier tubes is essential due to its importance in accurately being able to reconstruct events. Event reconstruction makes use of determining exactly where interaction vertices are and the direction in which particles travel, and to do this the time response needs to be very carefully calibrated. The response time of the PMT also relates to the amount of charge observed: in order for a hit to be registered, the PMT signal must pass the discriminator value of the hit threshold, and the time in which this happens is dependent on height of the pulse, which is correlated with observed charge. All these factors need to be considered when calibrating hit timing information.

To aid the timing calibration, a diffuser ball is placed near the centre of the inner detector, into which a nitrogen laser injects pulsed laser light. By varying the intensity of this light, the laser light can be outputted in flashes, and the timing of the laser pulses is monitored using a 2-inch monitor PMT. The schematic of this timing calibration system is shown in Figure 4.5.

The value of these laser pulse timings, and the time-of-flight value from the diffuser ball are subtracted from the PMT hit time. Using the observed charge values and these adjusted hit times, "TQ" (time and charge) distributions can be plotted for each inner detector PMT. An example TQ distribution is shown for an inner detector PMT (cable number 00010) in Figure 4.6, where the vertical axis shows is the TOF corrected and laser pulse time corrected PMT hit time,

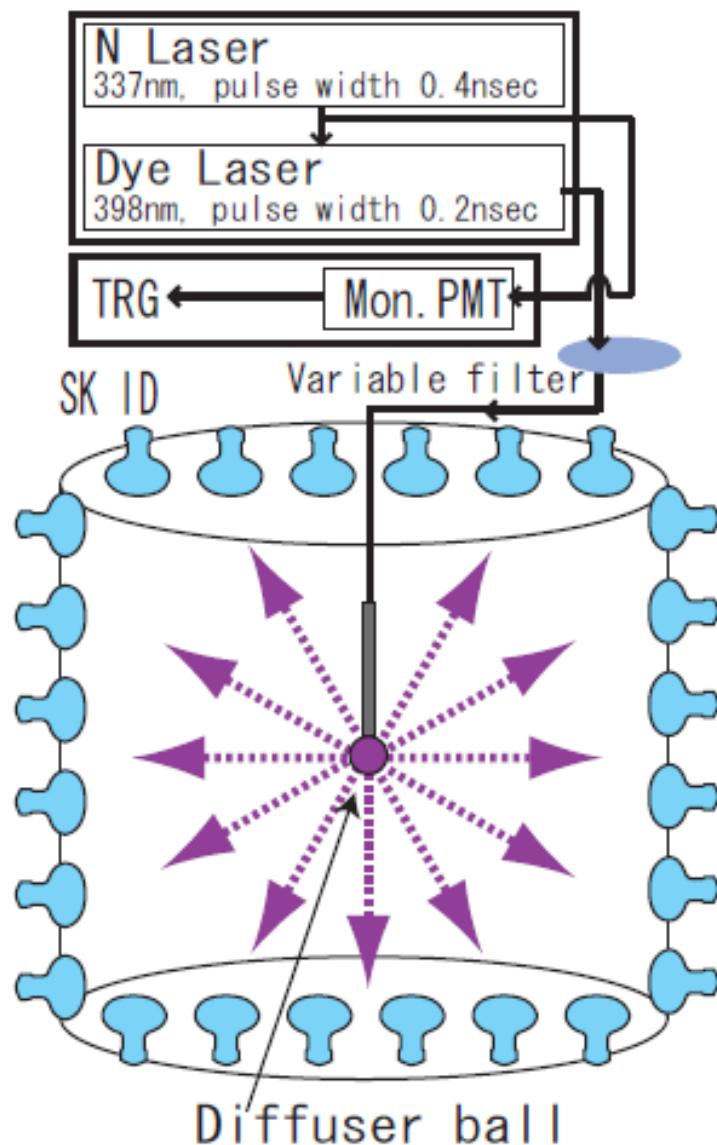


Figure 4.5: Schematic of timing calibration system, with SK ID PMTs in blue and the diffuser ball in purple. The dye laser shifts the wavelength of the laser light to 398 nm to maximise the quantum efficiency of the PMTs and light absorption.

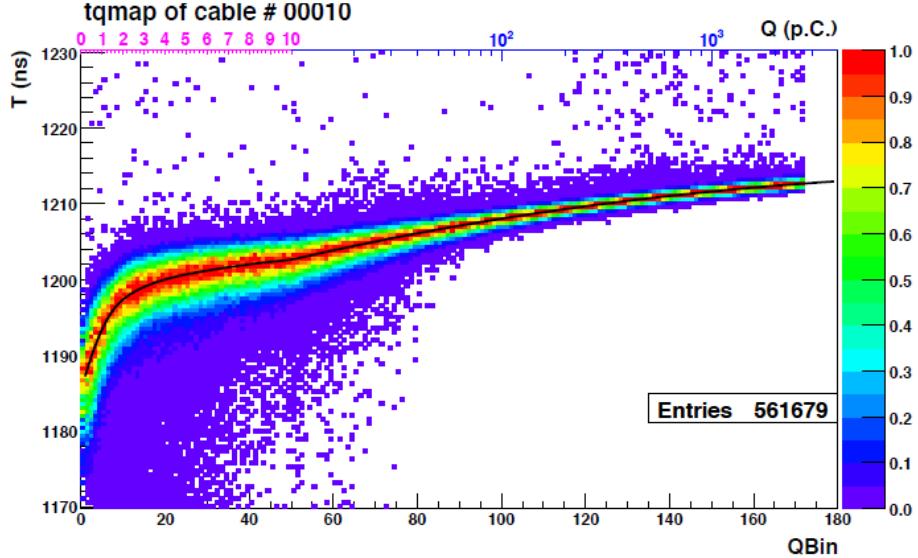


Figure 4.6: Example TQ distribution for an inner detetector PMT

the horizontal axis shows the observed charge of each hit. Figure taken from [3].

The timing resolution of the PMTs (also known as the transit time spread) is also something that must be calibrated. This is calculated by using the same timing and charge data used to produce the TQ distributions. By correcting all the ID PMT hits by their TQ distributions, these residual timing distributions have an asymmetric Gaussian function fitted to them, from which two values of sigma are extracted:  $\sigma_t$  and  $\sigma_{t'}$ . These are the values for the timing resolution for before and after the peak time of this distribution. Equation 4.4 shows how the asymmetric Gaussian is defined.

$$f(t; t > T_{\text{peak}}) \equiv A_1 \cdot \exp\left(-(t - T_{\text{peak}})^2 / \sigma_t^2\right) + B_1 \quad (4.3)$$

$$f(t; t \leq T_{\text{peak}}) \equiv A_2 \cdot \exp\left(-(t - T_{\text{peak}})^2 / \sigma_{t'}^2\right) + B_2 \quad (4.4)$$

Figure 4.7 shows an example of the residual timing distribution fitted with

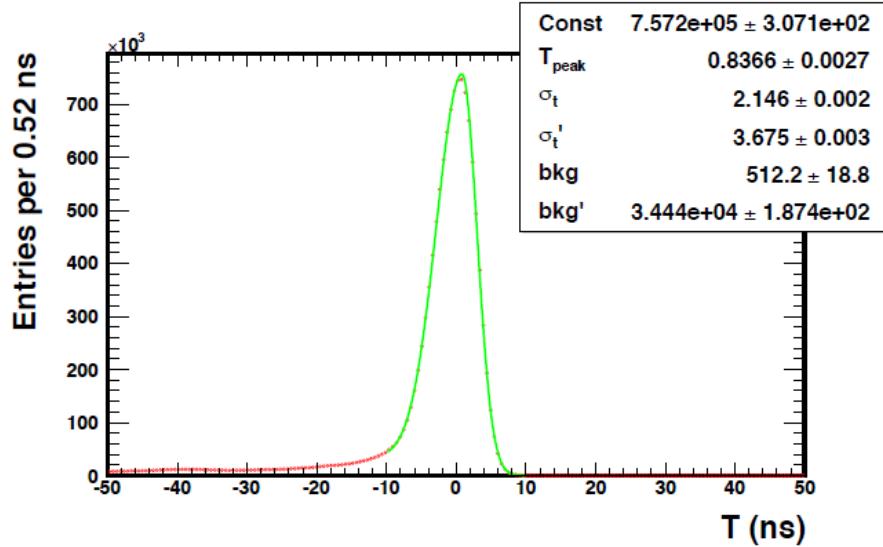


Figure 4.7: Residual timing distribution summed over all the readout channels in charge bin 14 (QBin in 4.6.)

an asymmetric gaussian for a certain value of binned charge.

Figure 4.8 shows timing resolution plotted as a function of charge for both before and after the peak time of the residual time distribution (red and blue points respectively) for SK-IV TQ calibration data.

#### 4.1.6 Measurement of the properties of the water in Super-Kamiokande

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#### 4.1.7 Measurement of light reflection on the black sheet and the PMTs

Using the same method as described in

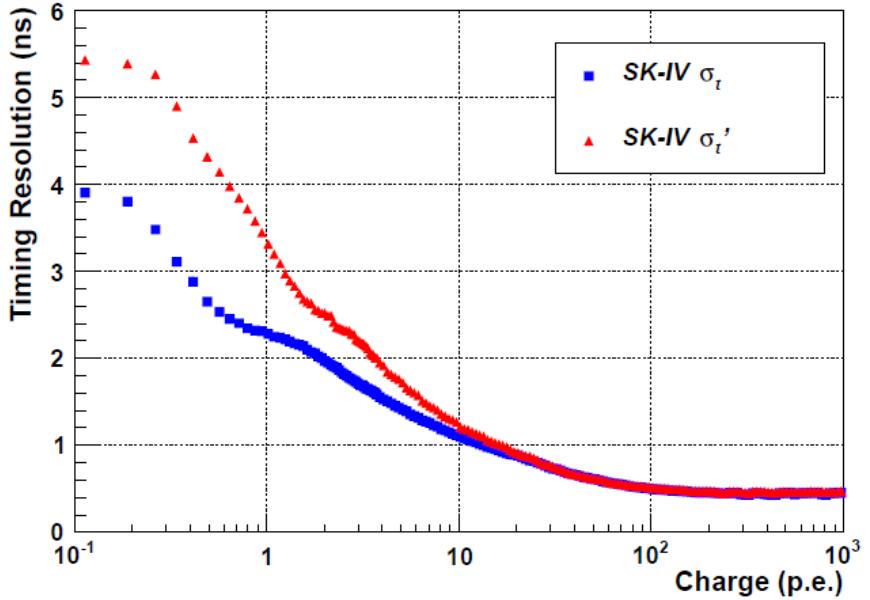


Figure 4.8: Timing resolution as a function of charge for SK-IV

## 4.2 Measurement of the properties of the water in Super-Kamiokande

the amount of light reflected at the PMT surface is determined. For SK-IV, four types of material, corresponding to four different refractive indices are taken into account: water, glass, bialkali, and vacuum. Table 4.1 shows the values for each of these materials,  $\lambda$  is the wavelength of the light in nm and  $n_{real}$  and  $n_{img}$  are the real and imaginary parts of the complex refractive index.

Regarding the black sheet used to line the inside of Super-Kamiokande, it can either reflect or absorb Cherenkov photons. The amount of Cherenkov photons measured is measured by a light injector setup, shown in Figure 4.9. For three different incident angles of laser light ( $30^\circ$ ,  $45^\circ$  and  $60^\circ$ ) and three different laser light wavelengths (337 nm, 400 nm and 420 nm) the charge from

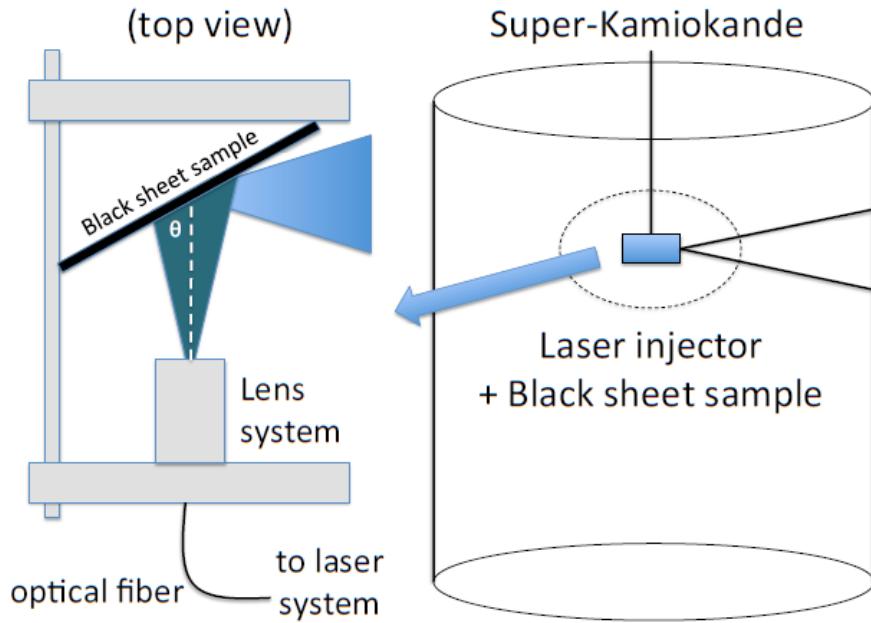


Figure 4.9: Schematic of the laser light reflectivity. Bird's-eye view (left) and setup inside of Super-Kamiokande (right).

the light scattered off the black sheet was measured. The direct charge (i.e. the same setup without the black sheet present) was also measured, with the total black sheet reflectivity being the ratio between the scattered charge and the direct charge.

Material	Refractive index
Water	1.33
Glass	$1.472 + 3670/\lambda^2$
<i>Bi – alkali</i>	$n_{real} + i n_{img}$
Vacuum	1.00
$x_{abs}^N$	Nucleon absorption probability
$x_\pi^N$	Nucleon $\pi$ -production probability

Table 4.1: Refractive indices of materials inside the Super-K detector.

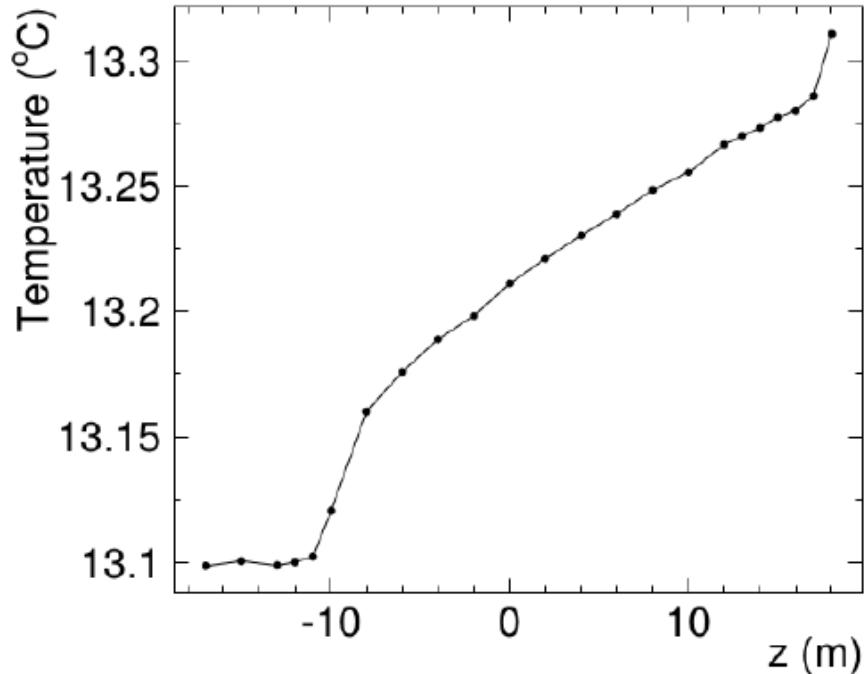


Figure 4.10: Vertical position in the detector ( $z$ ), plotted against the temperature of the water.

#### Temperature gradient and water transparency measurements

There is a vertical temperature gradient in the Super-Kamiokande tank due to the fact that ultrapure water is supplied upwards from the bottom of the detector shown in Figure 4.10. Below  $z = 11\text{m}$ , there is no variation in temperature, whereas above this value there is an increase of temperature as height in the detector increases. This variation in temperature causes a variation in the attenuation length of light in the detector, which means that the probability of PMTs being hit changes depending on the depth in the detector.

Using the nickel source mentioned in the section in absolute gain calibration, the dependence between the amount of light attenuation and vertical position in the tank is calculated, with the equation for this model shown in 4.5, where

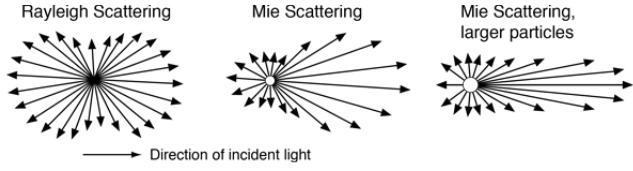


Figure 4.11: Schematic of the differences between Rayleigh and Mie scattering

$\langle N_{top} \rangle$ ,  $\langle N_{bottom} \rangle$  and  $\langle N_{barrel} \rangle$  are the mean rate at which the PMTs at the top, bottom and barrel regions are hit.

$$E_\nu = \frac{m_\pi^2 - m_\mu^2}{2(E_\pi - \sqrt{E_\pi^2 - m_\pi^2} \cos \theta)} \quad (4.5)$$

#### Variation of water transparency over time

In Super-Kamiokande, the detector medium scatters and absorbs the Cherenkov light, attenuating its wavelength. The amount of light attenuation after the light has travelled distance  $l$  is given by  $\exp(-l/L(\lambda))$ . Here  $L(\lambda)$  is the attenuation length of the detector medium and is shown by Equation 4.6,

$$L(\lambda) = \frac{1}{\alpha_{\text{sym}}(\lambda) + \alpha_{\text{asym}}(\lambda) + \alpha_{\text{abs}}(\lambda)} \quad (4.6)$$

In Equation 4.6, the  $\alpha_{\text{sym}}$  parameter consists of two types of scattering: the symmetric component of Mie scattering and Rayleigh scattering, whereas  $\alpha_{\text{asym}}$  only consists of asymmetric Mie scattering. Figure 4.11 shows the difference between these two types of scattering:

Rayleigh scattering occurs when light is scattered from very small particles, the size of these particles being less than 1/10 of the wavelength of the light. For particles larger than the wavelength of the light being scattered, Mie scattering becomes more dominant, with more asymmetric scattering occurring the larger the particle is (i.e. with a greater forward lobe for larger particles.) As a result

of this  $\alpha_{\text{sym}}$  has an angular dependance of  $1 + \cos(\theta)^2$  where  $\theta$  is the angle between the incoming and outgoing scattered photon vectors, but  $\alpha_{\text{asym}}$  has a  $\cos\theta$  dependence, only for forward scattering, and an amplitude of zero for backward scattering. The  $\alpha_{\text{abs}}(\lambda)$  is the amplitude of the light absorption.

## **Chapter 5**

# **The UK Light Injection System**

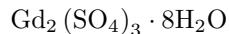
# **Chapter 6**

## **Super-Kamiokande Gadolinium Upgrade**

### **6.1 Physics motivation behind Super-Kamiokande Gadolinium Upgrade**

In order to be able to possibly observe the supernova relic neutrino (SRN) flux, also known as diffuse supernova neutrino background (DSNB) flux it was proposed that Gadolinium (Gd) should be added to the the water in Super-Kamiokande. Natural isotopes of gadolinium have large thermal neutron capture cross sections. As a result of this, when neutrons are captured on them there is a cascade of gamma rays that occurs, with an energy totalling 8 MeV, whereas neutron capture that occurs on hydrogen produces a single 2.2 MeV gamma ray. Two such natural isotopes, Gd-155 and Gd-157 have thermal neutron capture cross sections of 60740 barns and 253700 barns respectively, while the thermal neutron capture cross section of hydrogen is just 0.329 barns [20]. The Super-

Kamiokande with Gadolinium experiment, formerly known as GADZOOKS! (Gadolinium Antineutrino Detector Zealously Outperforming Old Kamiokande, Super!) was proposed in 2003, which stated the intention of enriching Water Cherenkov detectors with water soluble gadolinium salt. The ultimate aim is to load a total amount of gadolinium in the form of gadolinium sulphate octahydrate(



) in Super-Kamiokande which equates to 0.2% of Gd by mass, which would allow for 90% neutron capture efficiency. The ability to tag neutrons efficiently in Super-Kamiokande will benefit multiple physics topics, not only for the aforementioned observation of SRN flux, but also for analyses involving atmospheric neutrinos and proton decay.

A massive amount of energy is released when a core-collapse supernova (SN) occurs, about  $10^{46}$  J. The vast majority of this energy (99%) is released in the form of neutrinos, and due to neutrinos interacting with matter only very weakly, these traverse space with barely any attenuation. The interaction through which neutrino detectors such as Super-Kamiokande detect SRN flux is through inverse beta decay (IBD), shown in Equation 6.1.

$$\bar{\nu}_e + p \rightarrow n + e^+ \quad (6.1)$$

Due to the large cross section of the interaction, these events constitute about 88% of the total number of events in the detector [19]. With efficient neutron tagging in Super-Kamiokande, the backgrounds (charged current interactions and muon spallation) in the search for SRN flux neutrinos would be largely reduced. The neutral current quasielastic (NCQE) background would still remain

due to the way the gamma rays arising from neutron capture mimic the signal of the inverse beta decay (IBD) reactions: a schematic of both NCQE and IBD reactions are shown in Figure 6.1. The measurement of the NCQE interactions using T2K beam flux can aid in understanding this background more due to the T2K flux peak being near the atmospheric neutrino flux peak ( 600 MeV).

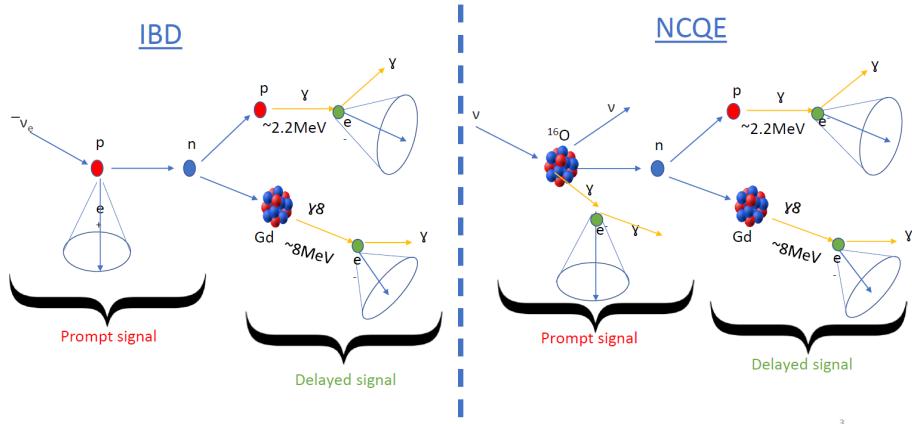


Figure 6.1: Schematic showing the IBD and NCQE interaction mechanisms

Efficient neutron tagging aids information about neutrino energy and neutrino interaction type, and when it comes to studying atmospheric neutrino oscillations, accurate neutrino energy reconstruction is particularly important. Figure 6.2 shows the fraction of non-visible energy as a function of the number of tagged neutrons from simulations of atmospheric neutrino interactions at Super-Kamiokande. Here  $E_\nu$  is the energy of the atmospheric neutrino and  $E_{vis}$  is the energy that is reconstructed from charged particles. Due to these neutrinos interacting with nuclei in the detector, more neutrons are produced, and with efficient neutron tagging on gadolinium the neutrino energy can be properly reconstructed.

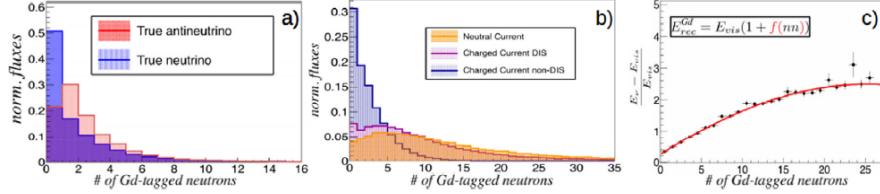


Figure 6.2: MC study of (a) neutron multiplicity production for  $\nu$  and  $\bar{\nu}$ , (b) neutral current, charged current and deep and non-deep inelastic scattering, (c) the correction to the visible energy as a function of neutrino multiplicity. Taken from [19].

Proton decay searches benefit from the addition of gadolinium to Super-Kamiokande because the main background to proton decay analyses come from atmospheric neutrino interactions, due to Figure 6.2 showing that atmospheric neutrinos cause at least one neutron to be produced.

## 6.2 The EGADS project

In 2009, prior to the addition of gadolinium in Super-Kamiokande, the EGADS (Evaluating Gadolinium's Action on Detector Systems) project was used to evaluate how the inclusion of gadolinium sulphate octahydrate would affect water quality and detector components inside Super-Kamiokande and their analyses. It was also used to observe how to reduce the visible neutron background from spallation and neutrons from fission chains from the uranium and thorium impurities in the gadolinium sulphate. EGADS is a cylindrical 200 ton stainless steel tank in a cavern near Super-Kamiokande and has its own water purification system and gadolinium sulphate octahydrate dissolving system.

Observing the impact the addition of gadolinium sulphate octahydrate had on the water quality and components was especially crucial, and after loading 0.2% gadolinium sulphate into the the EGADS tank in 2013, 240 PMTs were installed into the detector. 224 of these are the 50 cm Super-Kamiokande inner

detector PMTs, with 60 of these having the same FRP and acrylic covers as the inner detector PMTs. Much like Super-Kamiokande, the active photocoverage for EGADS is 40%, with the inside walls of the detector also being covered in black polyethylene terephthalate sheets. However, unlike Super-Kamiokande, there is no outer detector in EGADS. Figure 6.3 shows the PMT map for the EGADS detector along with the PMT types. Along with the PMTs which are identical to the ones inside Super-Kamiokande, EGADS also contains PMTs which are used for research and development for use inside Hyper-Kamiokande [19].

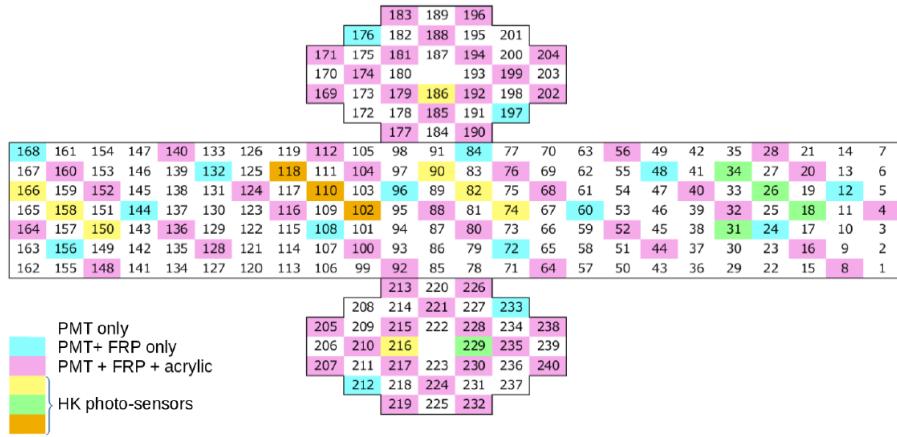
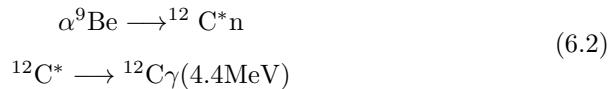


Figure 6.3: Map of unrolled cylindrical EGADS detector with PMT types. Taken from [19].

Measurements regarding neutron tagging were also taken using an Americium/Beryllium (Am/Be) source placed inside EGADS. Using an Am/Be neutron source to observe gadolinium's efficacy with respect to tagging neutrons is feasible because the Am/Be source decays as in Equation 6.2. It produces a prompt 4.4 MeV gamma ray alongside a neutron during its decay process, and as a result the prompt 4.4 MeV signal can serve as a trigger signal, while the following hundreds of microseconds can be used as a timing window within which

to scan for the neutron. Due to its similarity to the neutral current quasielastic events studied in the analysis in this thesis, it can serve as a helpful control sample and is used in the calculation of the detector response uncertainty in Chapter 7.



The delayed neutron capture time from an Am/Be source used in EGADS with the gadolinium sulphate concentration of 0.2% is shown in Figure 6.4. Here, we can see that the neutron capture time from the data is  $29 \pm 0.3 \mu\text{s}$  and for the Monte Carlo it is  $30 \pm 0.8 \mu\text{s}$ .

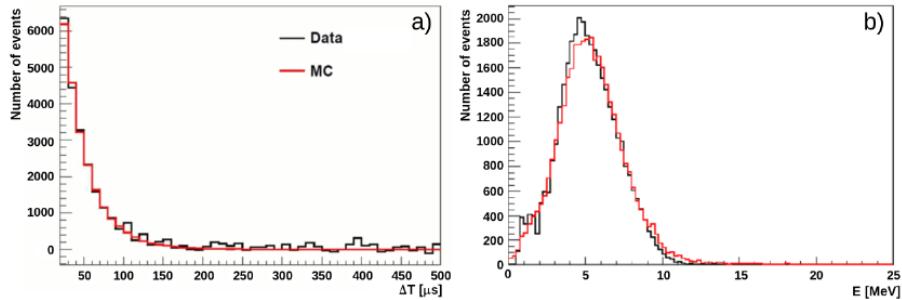


Figure 6.4: a) Delayed neutron capture time from prompt event with Am/Be source.  
b) Reconstructed energy from gamma rays after neutron capture.

EGADS water transparency and and gadolinium sulphate concentration was monitored daily to ensure similarity to ensure no negative effects to the detector components. Figure 6.5 shows the transparency and concentration from bottom, top and centre sampling ports in EGADS, with the blue band being typical water transparency values for Super-Kamiokande. It shows that the water transparency values remain comparable to Super-Kamiokande values and that there is very little variation from the final target gadolinium sulphate concentration,

as there is little variation from the black dashed line.

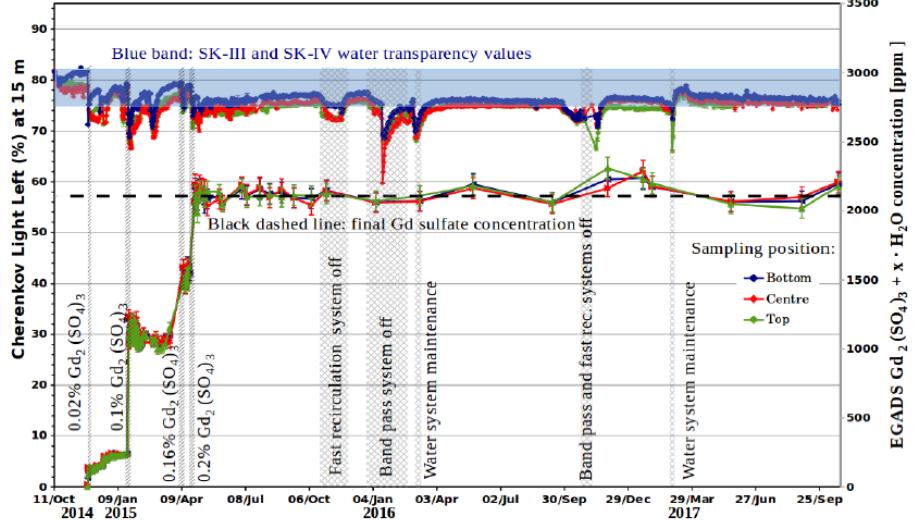


Figure 6.5: The upper lines are Cherenkov light left (%) and the lower three lines represent gadolinium sulphate concentration. The blue, red and green lines represent measurements taken from the bottom, centre and top sampling ports, respectively. Taken from [19]

EGADS also represented the most realistic possible soak test, and after two and a half years of running EGADS was emptied in November 2017 to check the condition of the inner structure and the photomultiplier tubes. There was no deterioration of any of the components, which was an excellent sign for a detector designed to so closely resemble the conditions for the Super-Kamiokande Gadolinium upgrade.

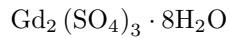
### 6.3 Gadolinium loading into Super-Kamiokande

After the success of EGADS, the Super-Kamiokande Gadolinium project was formed in 2015 when the final goal of adding 0.2% gadolinium sulphate octahydrate by mass to the detector. 10% of the target concentration was then loaded into Super-Kamiokande from July 14th to August 17th 2020. The details of the

various aspects of the gadolinium loading are mentioned in the next subsections: the SK-Gd water system, the flow of the water in the detector, and the loading of the gadolinium sulphate octahydrate powder.

### The SK-Gd water system

The SK-Gd water system was designed to dissolve gadolinium sulphate octahydrate into the detector water, pass it through a "pretreatment" system to remove contaminants from the water, and then to continuously circulate it. Figure 6.6 shows a schematic of the system. The



powder is transported to a weighing hopper which dissolves it into water held in a dissolving tank, which allows for exact amounts of the powder to be dissolved to attain the specific concentration required. The water in Super-Kamiokande is fed into a solvent tank, and then to the dissolving tank, from which it receives the gadolinium sulphate octahydrate from the shear blender. Water from this dissolving tank mixes the powder, and this combination is sent to the pretreatment system. The design for this treatment system was successfully tested in EGADS and involves treating this solution with ultraviolet light. Positively charged impurities dissolved in the solution (such as radium ions) and negatively charged impurities (such as uranium) are removed using a cation and anion exchange resins respectively. The gadolinium loaded water is next transferred to a UV steriliser to remove bacteria introduced during the dissolving process. This pretreatment process only occurs during gadolinium octahydrate sulphate dissolution, and not during the recirculation process. The recirculation process is similar to the pretreatment process, but with ultrafiltration modules installed in the final stage, and heat exchmage units used before the water is returned to

the detector, to help maintain exact control over the water temperature.

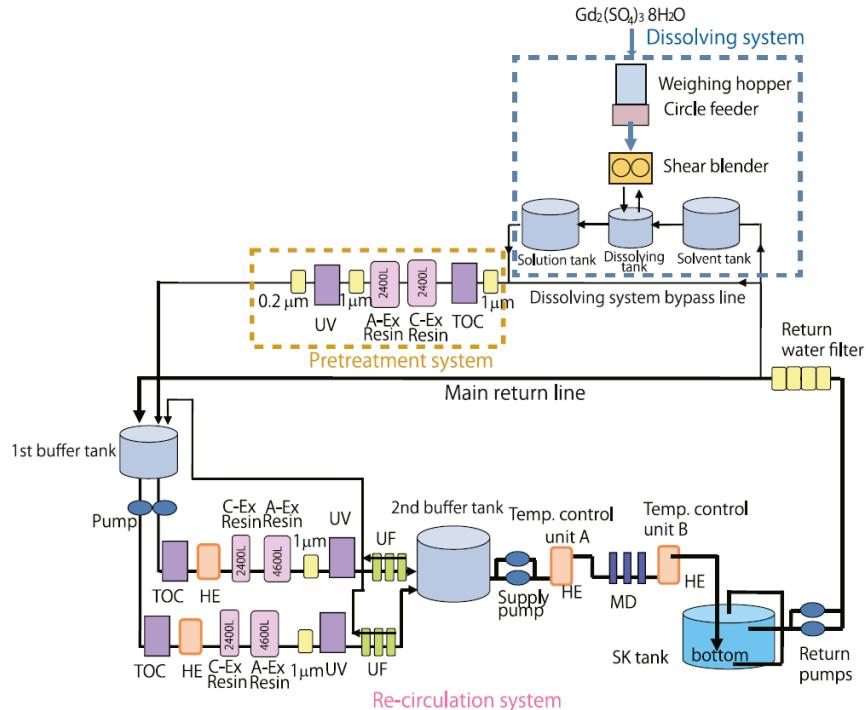


Figure 6.6: Schematic of the SK-Gd water system

### Control of water parameters in SK-Gd

Before gadolinium loading occurred, the water temperature in the tank was set at 13.9 °C exactly and was circulated at this temperature for about 45 days. The temperature of the "supply water" (water sent from the water system to the detector tank) was set to a lower temperature (13.55 °C) as Gd-loading began. This gradient in the density between the tank water and supply water meant that spatial configuration of the Gd-loaded water could be observed, by taking the temperature of the water at different locations in the tank. Figure 6.7 shows a schematic of the water piping during Gd dissolving for both the inner and outer detectors.

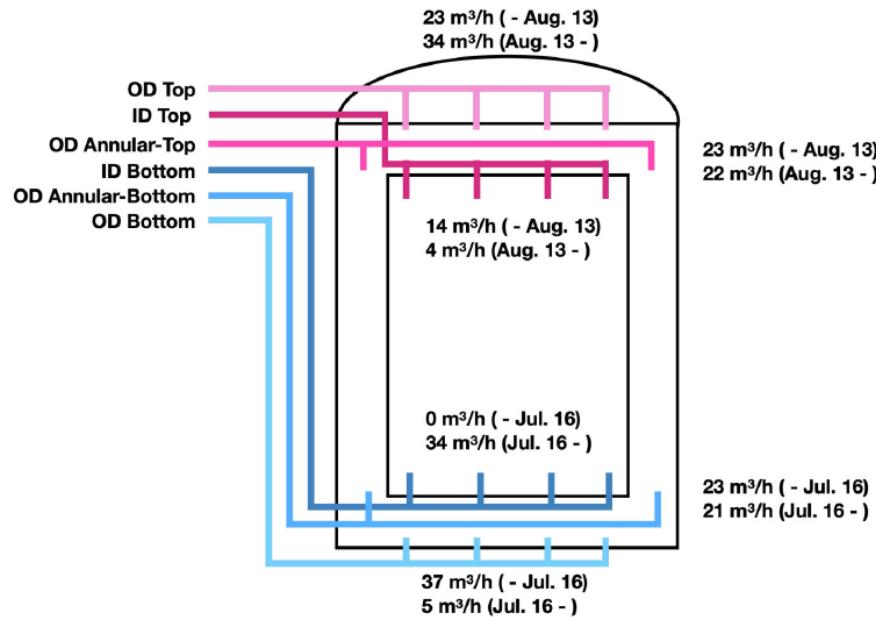
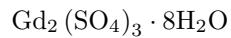


Figure 6.7: Schematic of the SK-Gd water piping for both the ID and OD

#### Gd-powder specification and loading

The acceptable background rate after the final value for Gd loading (set to be a concentration of 0.1% Gadolinium) was set to be less than double that of the background rate for when Super-Kamioknde ran with pure water. As a result, rigorous standards of cleanliness were set for the



powder, which meant setting maximum allowed levels of radioactive impurities which are shown in Table 6.1.

Chain	Isotope	Criterion [mBq/kg]	Physics target
$^{238}\text{U}$	$^{238}\text{U}$	< 5	SRN
	$^{226}\text{Ra}$	< 0.5	Solar
$^{232}\text{Th}$	$^{232}\text{Th}$	< 0.05	Solar
	$^{228}\text{Ra}$	< 0.05	Solar
$^{235}\text{U}$	$^{235}\text{U}$	< 30	Solar
	$^{227}\text{Ac}/^{227}\text{Th}$	< 30	Solar

Table 6.1: Table of impurities in the gadolinium sulphate octahydrate powder

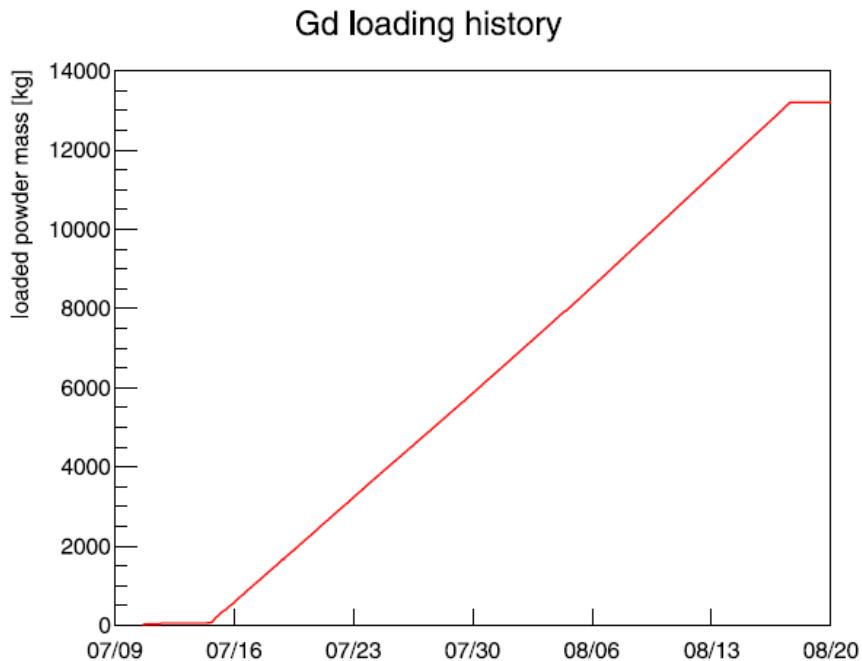


Figure 6.8: Schematic of the SK-Gd water piping for both the ID and OD

To sustain this very small level of impurity, the gadolinium powder batches went under routine screening at three collaboration laboratories at Boulby (UK), Canfranc (Spain), and the Kamioka Observatory.

Figure 6.8 shows the loaded mass of gadolinium powder in Super-Kamiokande at each dissolving cycle. The loaded mass of gadolinium sulphate octahydrate powder increased linearly until the target value of 5426 kg of it dis-

solved in 44,878,429 kg of water in Super-Kamiokande, resulting in a gadolinium concentration of 0.011% and a gadolinium sulphate ( $\text{Gd}_2(\text{SO}_4)_3$ ) concentration of 0.021%. After the loading of gadolinium sulphate powder into Super-Kamiokande, water transparency and attenuation length were continuously monitored. This had also been conducted in EGADS (see previous section), but due to the flow rate of the water and method by which gadolinium is loaded into the detector being different in Super-Kamiokande, it is important to monitor these attributes in the SK-Gd upgrade as well. Figure 6.9 shows how the attenuation length of the Cherenkov light in SK-Gd varies with time: this attenuation length is measured using cosmic ray through going muons (as explained in Chapter 3). There is a clear decrease in the attenuation length after the loading of  $(\text{Gd}_2(\text{SO}_4)_3 \cdot 8\text{H}_2\text{O})$ , reaching a minimum of 75 m in August of 2020, and increasing again to its original pre-Gadolinium loading value of 90 m by the beginning of December 2020.

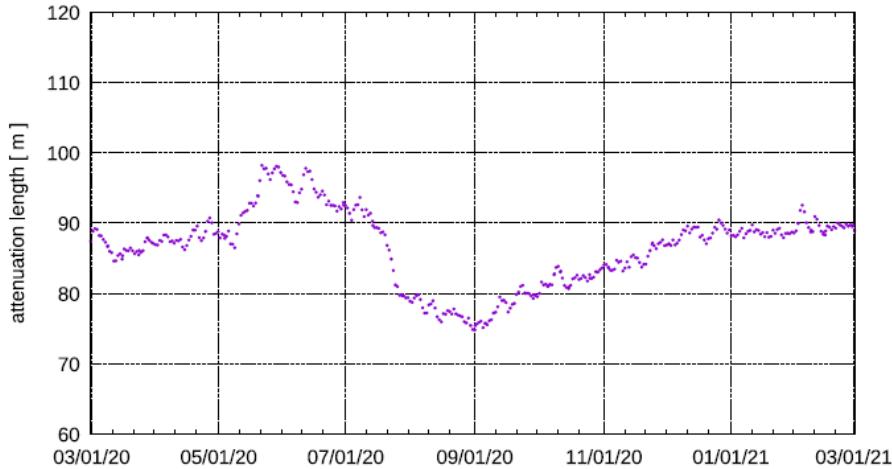


Figure 6.9: The attenuation length of Cherenkov light measured using cosmic ray muons between March of 2020 and March of 2021

There is also great importance in monitoring the concentration of the  $\text{Gd}_2(\text{SO}_4)_3 \cdot 8\text{H}_2\text{O}$  in the detector - this can be done using either direct sampling via calibration ports or using the Americium/Beryllium neutron source. The first method involves using a sampling probe (a 40m long flexible plastic tube) which took samples through two ports in the inner detector and the outer detector. This probe was attached to a flow meter, a pump and a conductivity meter which comprised the sampling system. The Super-Kamiokande water conductivity is a good measure of the  $\text{Gd}_2(\text{SO}_4)_3 \cdot 8\text{H}_2\text{O}$  concentration because due to the thorough water recirculation systems, all contaminants should have been removed, and all that would be measured would be the amount of the gadolinium and sulphate ions. Figure 6.10 shows the variation of the  $\text{Gd}_2(\text{SO}_4)_3 \cdot 8\text{H}_2\text{O}$  concentration in the detector at different depths in the detector during various points in time in the gadolinium loading period. Figure 6.11 shows the latest  $\text{Gd}_2(\text{SO}_4)_3 \cdot 8\text{H}_2\text{O}$  concentration measured using an atomic absorption spectrometer (AAS) in both the inner and outer detector. It shows that the  $\text{Gd}_2(\text{SO}_4)_3 \cdot 8\text{H}_2\text{O}$  is unvarying with an average value of  $271 \pm 4$  ppm.

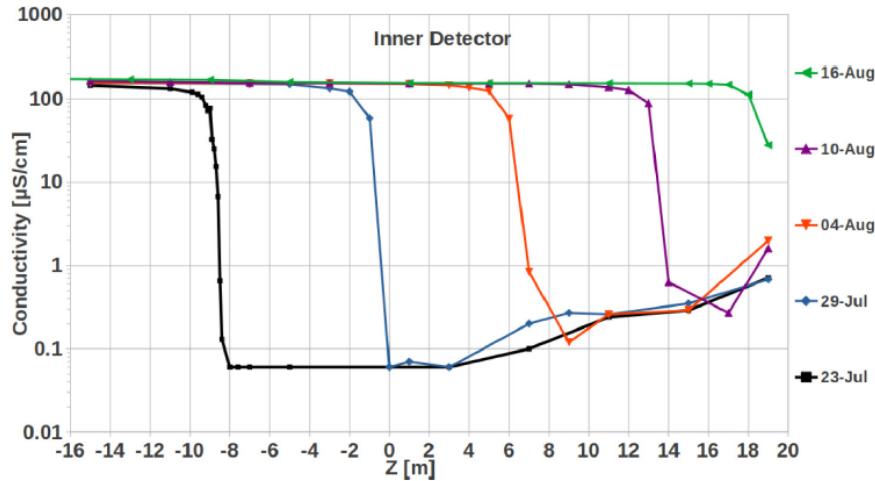


Fig. 8. Conductivity vs. Z position in the ID during Gd loading. Z=+19 m is inside the calibration guide pipe.

Figure 6.10: Conductivity vs. z position during Gd loading inside the inner detector

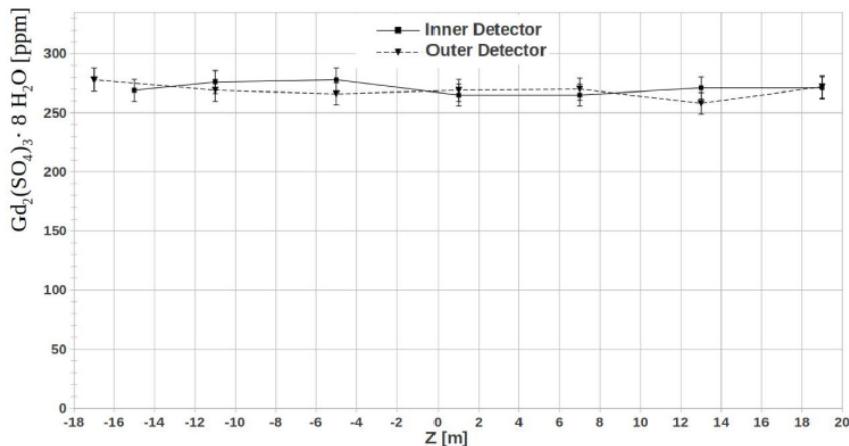


Figure 6.11: Latest gadolinium concentration measurements against z-position in the detector.

The second method of evaluating the  $\text{Gd}_2(\text{SO}_4)_3 \cdot 8\text{H}_2\text{O}$  concentration is using an Am/Be source. This source was placed within a crystal cube made of bismuth germanate (BGO) and the gamma ray released by the decay of the

source (shown in Equation 6.2) can be detected by the scintillation photons which pass through the BGO cube. The data was taken by placing the Am/Be source into the inner detector of Super-Kamiokande using central calibration port and then monitoring the source at three positions along the Z-axis. One in the middle of the detector ( $Z=0$  m) and two at  $Z = +12$  and  $Z = -12$  m respectively. As mentioned in Section 2, there are specialised triggers used when looking at events, and those used specifically for looking at neutron capture events at Super-Kamiokande are the SHE (super-high-energy) and AFT (after) triggers. The SHE trigger is generated whenever there is a minimum of one photon detected within a 200 ns timing window. 35  $\mu$ s after the SHE trigger is when all the photons from these neutron capture events were recorded, whereas for the AFT trigger, photons detected within a following 500  $\mu$ s window were recorded. Neutron captures which happened specifically on Gadolinium were taken from this data by using a 200 ns window within which there were 30 PMTs receiving atleast one photon hit and applying event vertex reconstruction using the method outlined in Section 2. Event reconstruction parameters such as timing and hit pattern goodness (a measure of how symmetrical the Cherenkov cone hit pattern is) were used to select events. The selected event vertex also had to be within 4 m from the Am/Be source position, the events during the SHE trigger had to be include 800 to 1300 active PMTs within 1.3  $\mu$ s, ensuring the selection of a 4.4 MeV gamma ray from the Am/Be source.

Figure 6.12 shows an example of the time distribution of neutron capture event candidates fit with a function. The Am/Be source position here was  $Z=0$  m and the 0  $\mu$ s time value on the x-axis is the time at which the scintillation photons produced by the 4.4 MeV gamma ray from its passing through the BGO cube are detected.

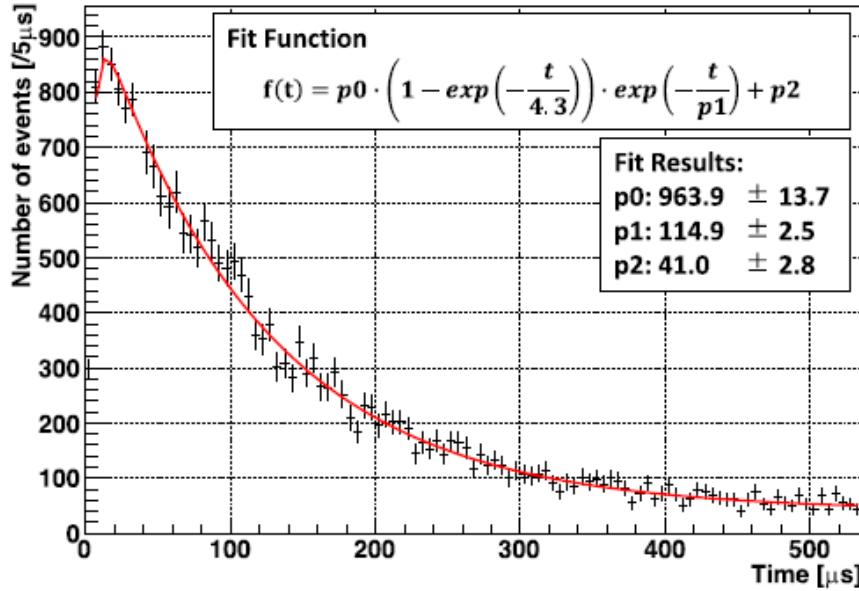


Figure 6.12: Latest gadolinium concentration measurements against z-position in the detector.

To convert the capture time of the neutron to a value for Gd concentration, GEANT4 Monte Carlo simulations were used. Figure 6.13 shows the neutron capture time constant plotted as a function of the concentration of Gd nuclei. The red points are the values from MC simulation and the black line is a polynomial function which is fitted to the red points. The horizontal blue stripe is the mean neutron capture lifetime calculated from Figure 6.12 ( $115 \pm 1$ )  $\mu$  s and the vertical blue stripe is the corresponding Gadolinium concentration value of  $111 \pm 2$  ppm which is consistent with the Gd concentration value of 0.011 % [4].

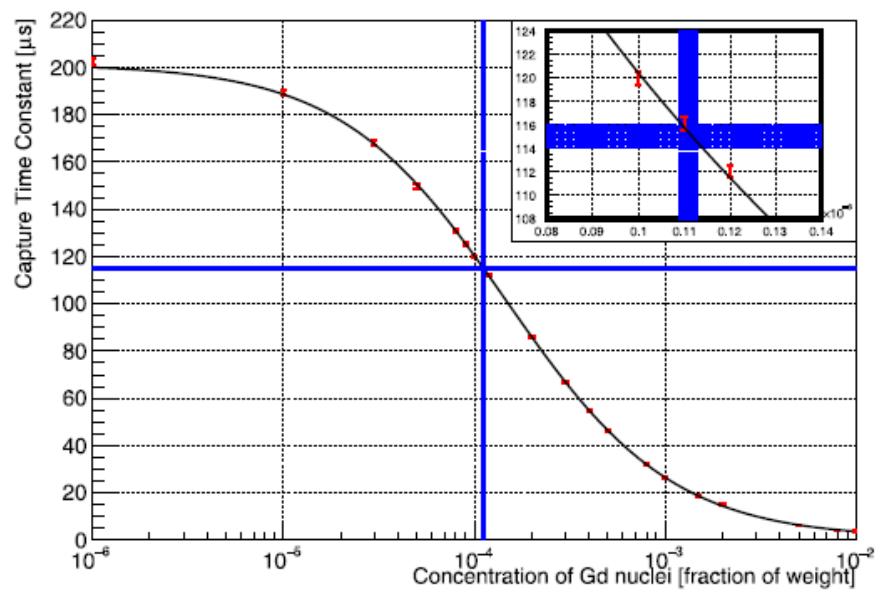


Figure 6.13: Neutron capture time constant ( $\mu s$ ) plotted against gadolinium concentration.

## **Chapter 7**

# **Measurement of Neutral Current Quasielastic Interactions with Super-Kamiokande Gadolinium Upgrade**

### **7.1 Event Simulation**

This chapter discusses details about the event simulation, specifically the software used and the way neutrino interactions are simulated in order to produce Monte Carlo. Table 7.1 summarises all the software used in this analysis.

Name	Version
Flux	13 a tuning v4.0
NEUT	5.3.3
SKDETSIM-SKGd	ANNRI-Gd model
Geant	4.10.05.p01
T2KReWeight	v1r23
GENIE	R2-12-10

Table 7.1: Software versions used in analysis

### 7.1.1 Neutrino flux

The neutrino beam flux used in this analysis is 13a tuning v4.0, taken from the NA61/SHINE fixed target experiment at CERN [27]. The NA61/SHINE experiment provides hadron production measurements for T2K and other long baseline experiments, and measures the yield of charged hadrons from a proton beam fired at a thin graphite target (2 cm long) and a T2K replica graphite target (90 cm long). From this experimental data, Monte Carlo simulations of the neutrino flux are predicted. The oscillation effect on the neutrino flux needs to be considered after choosing which flux to use, and although the neutrino-oxygen NCQE cross section does not depend on flavour (and therefore neutrino oscillation affects would have no impact on a fully pure NCQE sample) there is a small amount of charged current events which seep into the NCQE sample, and oscillation weights need to be applied to the charged current events in the sample. Monte Carlo reweighting takes care of this issue by assuming that they are all muon neutrino or muon anti-neutrino events, due to the fact that the number of electron neutrino events in FHC mode would be negligible.

### 7.1.2 Primary interaction

The interaction between the incoming neutrino event and an Oxygen nucleus, and it's consequent de-excitation is modelled by NEUT, which produces a vector

of primary particles. Because these are primary particles, they do not take into consideration all the particles produced from the detector response, i.e. the secondary particles produced from the interactions within the medium of the detector, in our case, water doped with gadolinium sulphate octahydrate. The way NEUT treats the interaction between a nucleus and a neutrino for the case of a neutrino interaction with a  $^{16}\text{O}$  isotope is shown in Equation 7.1.

$$\sigma(\nu^{16}\text{O}) = \sum_{i=1}^8 \sigma_p(p_i^{(p)}) + \sum_{j=1}^8 \sigma_n(p_j^{(n)}) + \sigma(2p2h) \quad (7.1)$$

Here  $\sigma_p$  and  $\sigma_n$  are the neutrino-proton cross section and the neutrino-neutron cross section respectively, and these cross-sections are dependent on the momenta of the protons and neutrons in the model, and therefore on the choice on nuclear model. The Benhar Spectral Function (SF) is used by NEUT as the nuclear model for NCQE interactions, and for CCQE interactions the Smith-Moniz relativistic Fermi Gas (RFG) model is used [6], with 2p2h interactions included for these interactions.

The lowest rung of nucleons in Figure 7.1 is the  $s_{1/2}$ , with  $p_{3/2}$  above this level and  $p_{1/2}$  above that. The protons in these rungs have removal energies of 12.1 MeV, 18.4 MeV and 42 MeV respectively, and due to neutron levels being more tightly bound, these have an extra removal of 3.24 MeV compared to their proton counterparts. The shell model is imperfect due to how it allocates the probability of  $^{16}\text{O}$  transitioning to the possible nucleon states. In the shell model, probabilities are allocated by counting the number of nucleons in each energy level and assigning probabilities according to how many there are. In Figure 7.1 it can be seen that the  $p_{3/2}$  state has double the number of nucleons compared to the  $s_{1/2}$  model and therefore double the probability of transition is assigned for the  $p_{3/2}$  state, and the probability of transitioning to any other state is assigned to be 0, because in the shell model they don't even exist. The Benhar

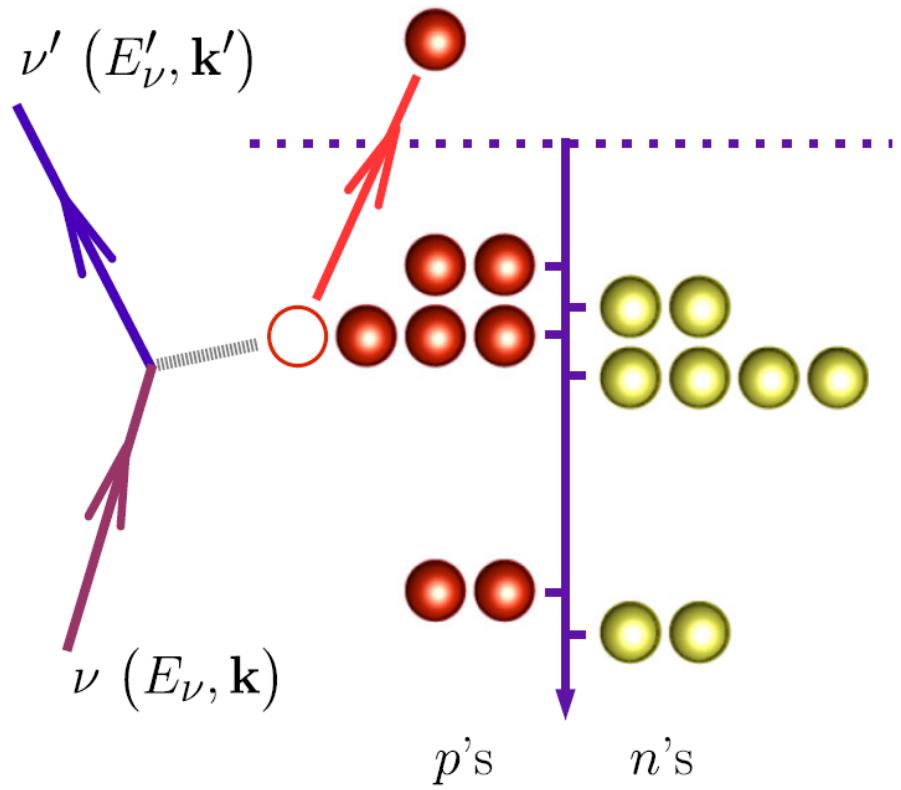


Figure 7.1: Representation of NC neutrino scattering off  $^{16}\text{O}$  with protons on the left hand side and neutrons on the right arranged according to the shell model.

	$(s_{1/2})^{-1}$	$(p_{3/2})^{-1}$	$(p_{1/2})^{-1}$	others
Shell model	0.25	0.5	0.25	0
Spectral Function	0.1055	0.3515	0.158	0.385
This analysis	0.4905	0.3515	0.158	0

Table 7.2: Transition probabilities for different models and states

Spectral Function model however is complex and is tuned using electron-nucleus scattering data. The model used in this analysis is a modified version of the Benhar Spectral Function model, where the case of other transition states is dealt with by merging the "others" state into the  $s_{1/2}$  state. Table 7.2 gives the probabilities of transition to different states for different models.

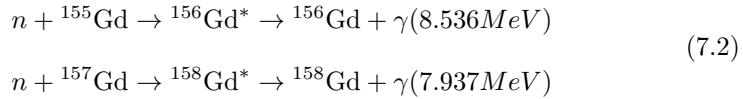
### 7.1.3 Detector response and interactions in the detector medium

Prior analyses to this used SKDETSIM (Super-Kamiokande Detector Simulator) to simulate the trajectories of particles through the water in Super-Kamiokande and output detector response MC. This analysis uses SKDETSIM-SKGd to propagate the particles, due to the requirement of needing gadolinium sulphate present in the simulation. The particular isotopes of gadolinium used in the simulation are  $^{155}Gd$  and  $^{157}Gd$  due to their excellent thermal neutron capture cross sections. Table 7.3 shows the relative abundance of various gadolinium isotopes inside natural gadolinium and their associated thermal neutron capture cross sections.

As can be seen in Table 7.3, not only are  $^{155}Gd$  and  $^{157}Gd$  the most abundant isotopes, they also have extremely high neutron capture cross sections compared to the other isotopes. Equation 7.2 shows the neutron capture on both of these isotopes, and the energy of the subsequent gamma rays.

Isotope	Abundance [%]	Cross-section [b]
<sup>152</sup> Gd	0.200	735
<sup>154</sup> Gd	2.18	85
<sup>155</sup> <b>Gd</b>	<b>14.80</b>	<b>60900</b>
<sup>156</sup> Gd	20.47	1.8
<sup>157</sup> <b>Gdd</b>	<b>15.65</b>	<b>254000</b>
<sup>158</sup> Gd	24.84	2.2
<sup>160</sup> Gd	21.86	1.4

Table 7.3: Abundance and thermal neutron capture cross section of various isotopes of Gadolinium



It is important to therefore model the gamma ray emission spectra from both of these isotopes. The model used by SKDETSIM-SKGd is the ANNRI-Gd model. This uses gamma energy spectrum data from the germanium spectrometer at the ANNRI (Accurate Neutron Nucleus Reaction Measurement Instrument) experiment. This experiment uses the incoming pulsed neutron beam from the Japan Spallation Neutron Source (JSNS) at the Material and Life Science Experimental Facility (MLF) of J-PARC. After a 300 kW beam of protons from the JSNS facility hits a target of mercury and produces neutrons, this neutron beam hits an enriched <sup>155</sup>Gd target or a <sup>nat</sup>Gd film. The ANNRI spectrometer is placed 21.5 m away from the neutron beam source, with two cluster detectors on either side of the neutron capture target material, which is 13.4 cm away from each cluster detector. Surrounding the target, there are also 8 co-axial germanium detectors. The schematic for the ANNRI-Gd experiment is shown in Figure 7.2.

By defining how many Germanium crystals were hit and how many neigh-

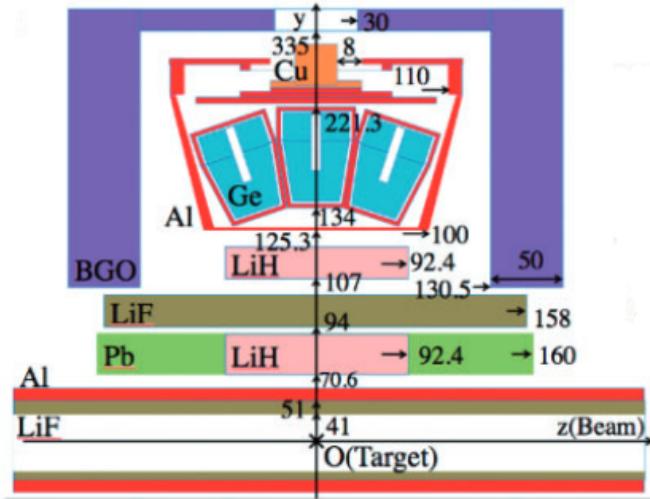


Figure 7.2: Schematic of the ANNRI Ge spectrometer (dimensions in mm). The beam pipe along with one of the Ge cluster detectors (light blue) is shown. The shaded purple area are the anti-coincidence shields made of bismuth-germanium-oxide (BGO) crystals.

bouring crystal received gamma ray hits, values of H (hit number) and M (multiplicity) can be assigned in order to classify the events. After selecting for neutron energy and subtracting the background, the events remaining were separated into samples based on their H and M values. Figure 7.3 shows the energy spectra for different multiplicity values (left) and hit values (right) from one of the germanium crystals in the ANNRI spectrometer.

One of the key differences between the default Geant4 model for thermal neutron capture on Gadolinium and the ANNRI-Gd model is how it conserves energy: the photon evaporation model conserves only the final sum of energy for the captured event, but performs poorly when modelling the gamma-ray energy on an individual event by event basis. One way the ANNRI-Gd model combats this is by separately describing the continuous and discrete peaks in Figure 7.3, where the discrete peaks are shown as spikes below 1500 keV and above

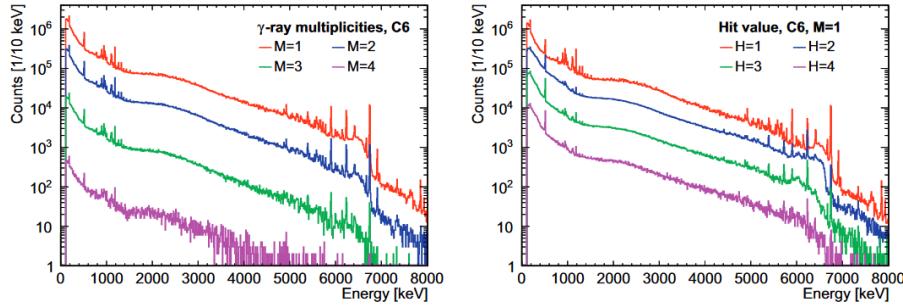


Figure 7.3: Energy spectra for thermal neutron capture on  $^{157}\text{Gd}$  energy spectra for different multiplicity values (left) and hit values (right) from one of the germanium crystals in the ANNRI spectrometer.

4500 keV. These come from the different ways in which gadolinium de-excites after thermal neutron capture. The continuous spectrum in the plots in Figure 7.3 show the de-excitation of  $^{158}\text{Gd}^*$  in multiple steps, producing multiple low energy gamma rays, which accounts for about 93% of the spectrum. The discrete spikes in the spectra come from a two-step cascade, which produces a high energy gamma ray instead, accounting for the remaining 7% of the spectrum. Figure 7.4 shows the continuous and discrete components of the ANNRI-Gd MC, along with data from the experiment. Figure 7.5 shows ratio plot of data and MC for the GLG4sim model, the default Geant4 photon evaporation model and the ANNRI-Gd model: here it is clear the ANNRI-Gd model fits the data better than the other two models, especially at energies above 3500 keV.

## 7.2 Event Reconstruction

### 7.2.1 Bonsai output reconstruction quantities

Due to this analysis looking specifically at the low energy region, a fitter specific to low energies (called LOWFIT) is used to reconstruct events. Both MC and data neutrino events undergo a reconstruction phase, where the low-energy

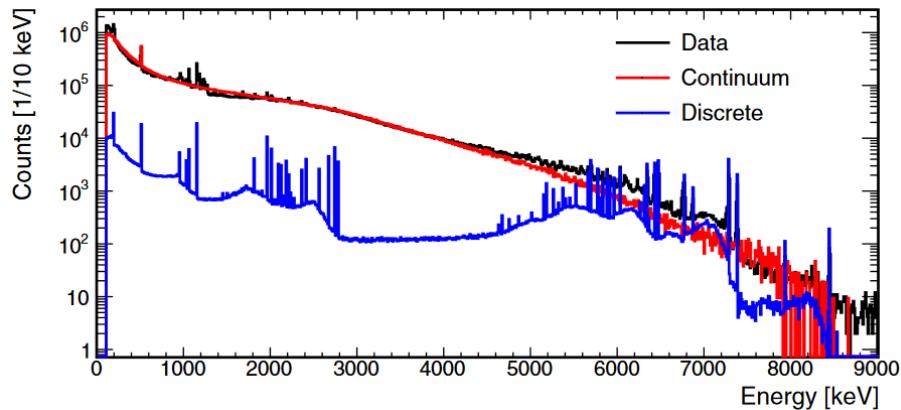


Figure 7.4: Energy spectra for the ANNRI-Gd model broken down into its continuous and discrete components along with data from the ANNRI-Gd experiment.

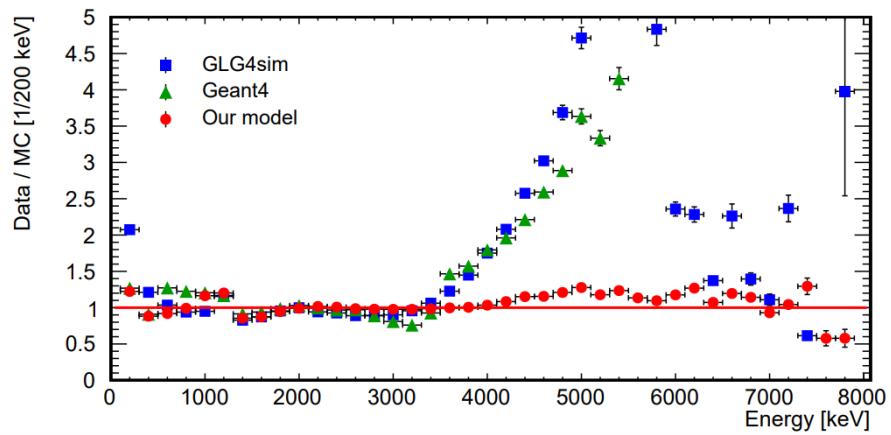


Figure 7.5: Comparison between neutron capture MC models and data.

fitter BONSAI is applied to the event, which is discussed in Chapter 2. This reconstruction is carried out using timing and cable information, however charge information is omitted. The output of BONSAI gives information which will be used in the reduction phase of the data and allow for the selection of the NCQE sample. The following quantities comprise the BONSAI output, the first two being helpful spectator variables and the latter five constituting parameters which are used in the reduction phase of the analysis, from which the neutrino NCQE sample is determined.

#### Neutrino vertex direction

This vector points towards the direction which is an average over all the Cherenkov cone axes which are produced, due to there being multiple leptons induced in the interaction.

#### Neutrino vertex position

The reconstructed location of the neutrino interaction event.

#### Reconstructed energy

In line with the standard SK low energy analysis definition, this energy is simply the reconstructed energy with the 0.511 MeV electron mass omitted. The range for Erec in this variable is 3.49 MeV to 29.49 MeV - the estimated kinetic energy under the hypothesis that the event is a singular electron.

#### Dwall

This variable gives the minimum distance of the neutrino vertex position from the closest wall of the Super-Kamiokande detector.

### Effwall

Thus variable gives the distance between the neutrino vertex positon and the closest wall, but moving back from the vertex position along the neutrino vertex direction vector.

### Vertex direction and goodness coefficient

The coefficient  $ovaQ$  (defined in Equation 7.3) describes the quality of the vertex reconstruction. It consists of two parameters  $g_{vtx}^2$  and  $g_{dir}^2$  where the former describes the goodness of the vertex which is based on PMT hit timings, and increases the sharper an event is in time. The latter is the directional goodness and measures the azimuthal uniformity in the ring pattern produced by the Cherenkov cone, which decreases the more uniform an event is in space. As a result of this,  $ovaQ$  increases the more uniform and sharp in time an event is.

$$ova\ Q = g_{vtx}^2 - g_{dir}^2 \quad (7.3)$$

$g_{vtx}$  is calculated using a fit of the PMT timing distribution and using the hit times of the PMT it is defined as in Equation 7.4.

$$g_{vtx} = \frac{\sum_i w_i e^{-\frac{1}{2}(\frac{\Delta t_i}{\sigma})^2}}{\sum_i w_i} \text{ with } w_i = -\frac{1}{2}(\frac{\Delta t_i}{\omega})^2 \quad (7.4)$$

Here  $\sum_i w_i$  is the weight given to the i-th hit PMT for the reduction of dark noise, where  $\omega$  has a value of 60ns.  $\sigma$  has a value of 5ns which is used to test the goodness, and as a result, a sharp timing distribution produces a large vertex goodness.  $g_{dir}$  is calculated by looking at how spatially uniform the hit PMTs are around the reconstructed neutrino vertex direction. In order to quantify this uniformity, the Kolmogorov-Smirnov (KS) test is used as in Equation 7.5.

$$g_{\text{dir}} = \frac{\max_i\{\angle_{\text{uni}}(i) - \angle_{\text{data}}(i)\} - \min_i\{\angle_{\text{uni}}(i) - \angle_{\text{data}}(i)\}}{2\pi} \quad (7.5)$$

where  $\angle_{\text{data}}(i)$  is the azimuthal angle of i-th hit real PMT included in the number of hits in 50ns.  $\angle_{\text{uni}}(i) = 2\pi i/N_{50}$  is the azimuthal angle of the i-th virtual PMT hit, but only when uniform distribution of the hits is assumed. As the uniformity of the hit pattern increases, the goodness decreases.

#### Cherenkov angle $\theta_{\text{C}}$

For relativistic electrons in water, the value of the Cherenkov opening angle is  $\approx 41^\circ$ , due to the relation:

$$\cos \theta_{\text{Cherenkov}} = \frac{1}{n\beta} \quad (7.6)$$

where  $\beta = v/c \approx 1$  and  $n$  is the refractive index of water, 1.33. However due to other particles in the simulation, such as protons or muons, the Cherenkov cone is expected to be narrower, or if multiple leptons are present, the Cherenkov cones will be less distinct and more spread out, leading to deviations from the  $41^\circ$  value.

### 7.2.2 True neutron tagging information

### 7.2.3 Primary selection criteria

### 7.2.4 Secondary selection criteria

When the neutron vertex is found by this method, 14 variables which describe different aspects of the neutron candidate are calculated. For each of the neutron candidates the vector of these variables are computed and fed into the neural network and this produces an output value which is between 0 and 1. These variables relate to different features regarding categorising hits from neutron

capture on Gd or H, including the number of the hits from neutron capture, the isotropy of these hits, the Cherenkov angles of these hits and the position of the neutron vertex in the detector when capture occurs. A description of these variables are given as follows:

#### N10nvx

This is the number of hits in the 10ns sliding window of the neutron candidate

#### N300S

Excluding the number of hits in the 10ns sliding window (N10nvx), this is the number of hits in the extended window of 300ns.

#### NcS

This variable is defined as:

$$NcS = N10nvx - Nclushit \quad (7.7)$$

Where *Nclushit* is the number of clusterised hits: if hit *i* and *j* are hits on PMTs, then for hit *i* and hit *j* the hit vector  $\hat{r}_i$  can be written as:

$$\hat{r}_i = \frac{\overrightarrow{PMT_i} - \overrightarrow{VTX_n}}{\left\| \overrightarrow{PMT_i} - \overrightarrow{VTX_n} \right\|} \quad (7.8)$$

where the angle at the point of the neutron capture vertex between  $\hat{r}_i$  and  $\hat{r}_j$  of the PMT hits is defined as:

$$\theta_{ij} = \arccos(\hat{p}_i \cdot \hat{p}_j) \quad (7.9)$$

where the hits are defined as clustered if  $\theta_{ij}$  is less than 14

#### llrca

This variable is the log likelihood ratio calculated using triplets of hits from

N10nvx that make up a rudimentary Cherenkov cone, from which the opening angle  $\theta$  is calculated. Two PDFs ( $\theta_{Ci}$ ) and ( $\theta_{Cj}$ ) are calculated from each  $\theta_{Ci}$  where p\_s and p\_b are the probability density functions of  $\theta_C$  depending on whether the hits come from a true neutron capture on Gd or H or a false neutron capture which makes up the background. The log likelihood ratio variable is computed using Equation 7.10.

$$llrca = \sum_{i \in \{triplets\}} \log \left( \frac{f_B(\theta_{Ci})}{f_S(\theta_{Ci})} \right) \quad (7.10)$$

#### beta-n

These variables (where n = 1,2,3,4,5) are defined using Legendre polynomials, shown in Equation 7.11, which gives the isotropy of the Cherenkov hits.

$$\text{beta\_n} = \frac{2}{N10nvx(N10nvx - 1)} \sum_{i \neq j} \text{Legendre}_n(\cos \theta_{ij}) \quad (7.11)$$

where  $\text{Legendre}_n$  gives the Legendre polynomial of order  $n$  and  $\theta_{ij}$  is the angle between hit PMTs relative to the neutron capture vertex.

#### ndwall

This parameter, similar to dwall, gives the shortest distance of the neutron capture vertex from the wall of the Super-Kamiokande tank.

#### ntowall

This variable (similar to effwall), gives the distance of the neutron capture vertex from the wall, however, unlike ndwall it gives the direction of the neutron capture specifically along the direction of the centre of the hits. The direction ( $\vec{R}$ ) is given by:

$$\overrightarrow{\text{dir}} = \sum_{i=1}^{N10nvx} \hat{p}_i \quad (7.12)$$

## Chapter 8

# Systematic and statistical uncertainty calculations

### 8.1 Systematic uncertainty calculation methodology

The systematic uncertainties for this analysis are calculated using the probability distribution functions of each quantity appearing in the formula for the mean neutron multiplicity, which is given by:

$$M = \frac{\#n_{\text{det}} - R \times \#\nu_{\text{det}}}{T} \frac{1}{\#\nu_{\text{det}}} \quad (8.1)$$

By random sampling the probability distribution functions for each of the terms in Equation (8.1) one can calculate the multiplicity probability distribution functions for both the statistical uncertainty and the systematic uncertainty. The statistical uncertainty for the value for the multiplicity is related to the variation in the number of detected neutrons, while the systematic uncer-

tainty is related to the variation on the tagging efficiency and the background rate. The total search time for the tagged neutrons is dependent on the number of "windows" in which the neutron is searched for in, and therefore the term for the number of detected neutrinos. Because any variation on the number of neutrinos which are detected is unrelated to the value for the mean neutron multiplicity, calculating a probability mass function for the number of neutrinos is unnecessary.

A Poissonian distribution is used to model the distribution for the number of detected neutrons, due to its value being approximated by counting the positives in the timing window that the neutron tagging search is carried out in. The mean value of this Poisson distribution is the number of detected neutrons  $\langle \#n_{\text{detected}} \rangle$ .

$$PMF(\#n_{\text{detected}}) = \frac{1}{(\#n_{\text{detected}})!} \langle \#n_{\text{detected}} \rangle^{\#n_{\text{detected}}} e^{-\langle \#n_{\text{detected}} \rangle} \quad (8.2)$$

Regarding the background rate, this is estimated from dummy spill data. The background rate error is associated with the statistical variation of the Monte Carlo size that the background rate is associated with, and secondly the change of the background rate value during the SK-V period. The statistical variation of the MC is modelled using a Poisson, but the statistics are high enough so that it appears Gaussian, while the uncertainty relating to time variation is characterised by its own probability distribution function. In contrast, the tagging efficiency is model dependent and has systematic uncertainties relating to this. The two ways in which the systematic error are estimated are either using MC re-weighting or MC regeneration.

For the MC-reweighting approach, weights are applied to a quantity and the tagging efficiency of the re-weighted MC is extracted. The general methodology

is to have the input of a model (given by a set of parameters) and to vary them one by one and then calculate the reweighted tagging efficiencies - the set of relative discrepancies  $\delta_i$  are computed from this set of reweighted tagging efficiencies  $T_i$  and the nominal tagging efficiency  $T_{nom}$  using Equation (8.3).

$$\delta_i = \frac{T_i - T_{nom}}{T_{nom}} \quad i \in \{ \text{parameters} \} \quad (8.3)$$

These relative discrepancies  $\delta_i$  are used to calculate the one individual discrepancy  $\delta_{reweighted}$  that would describe the final deviation from the nominal tagging efficiency  $T_{nom}$  due to the systematic error. The parameter  $\delta_{reweighted}$  describes the model which has been produced through  $1\sigma$  variations of these parameters, therefore the final probability distribution function which describes the deviation from the nominal MC has a Gaussian distribution with the standard deviation being equal to  $\delta_{reweighted}$ .

The other method to estimate the systematic error on the tagging efficiency is the method of Monte Carlo regeneration. This is carried out by varying a parameter then regenerating the whole Monte Carlo and then extracting the tagging efficiency - therefore unlike with MC re-weighting there is no set of discrepancies  $\delta_i$  but instead two single discrepancies  $\delta_{min}$  and  $\delta_{max}$ . The resulting probability distribution which describes the deviation from the nominal Monte Carlo is a Gaussian which has the mean and standard deviation relating to the discrepancies shown in Equation (8.4).

$$\begin{cases} \mu = \frac{\delta_{max} + \delta_{min}}{2} \\ \sigma = \frac{\delta_{max} - \delta_{min}}{2} \end{cases} \quad (8.4)$$

## 8.2 Neutrino beam flux uncertainty

The uncertainty on neutrino beam fluxes can be evaluated by looking at the dependence of the tagging efficiency on the flux variations. The beam fluxes for the four flavour modes ( $\nu_e \bar{\nu}_e \nu_\mu \bar{\nu}_\mu$ ) have the fractional uncertainties given for each mode, FHC and RHC. The binned uncertainties are shown in Figure 8.1.

Each individual bin for the flux is increased/decreased by its error, the Monte Carlo re-weighting method is then used to extract the tagging efficiency for each flux bin, and Equation (8.5) is used to calculate the fractional uncertainty.

$$\delta_i(\pm\sigma) = \frac{T_i(\pm\sigma) - T_{\text{nom}}}{T_{\text{nom}}} \quad i \in \{ \text{each flux bin} \} \quad (8.5)$$

Figure 8.2 shows the fractional errors calculated from the reweighted Monte Carlo, with the red bars showing the  $-1\sigma$  variation and the blue bars showing the  $+1\sigma$  variation. Table ?? contains the value for the total uncertainty resulting from the neutrino beam flux, which was calculated using Equation (8.6), where the maximum value between the increased and decreased discrepancy is taken and summed over to produce the final neutrino flux beam uncertainty.

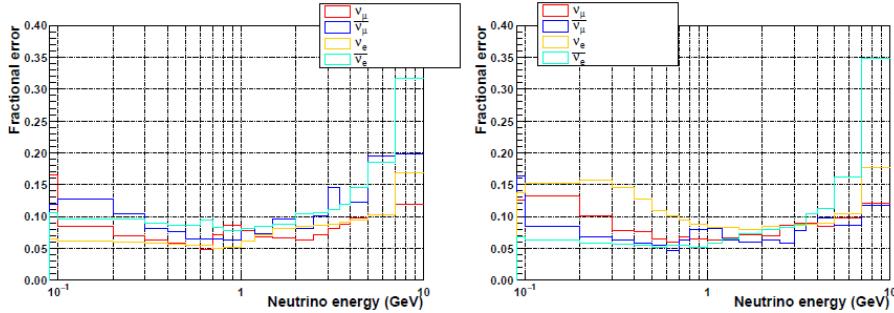


Figure 8.1: Fractional uncertainties of beam fluxes.

$$\delta_{\nu \text{ flux}} = \sum_{i \in \{\text{ bins}\}} \max [|\delta_i(+\sigma)|, |\delta_i(-\sigma)|] \quad (8.6)$$

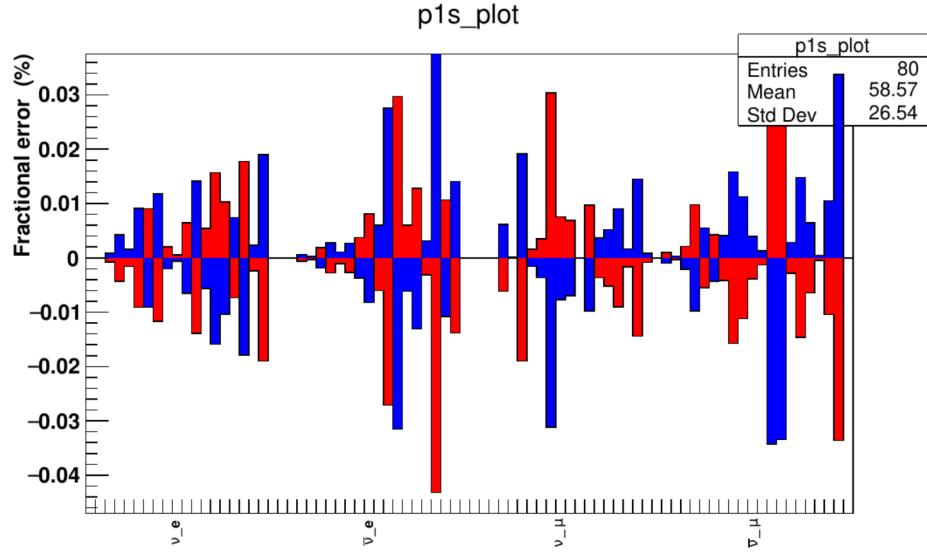


Figure 8.2: Tagging efficiency fractional uncertainties caused by neutrino beam flux discrepancies. From left to right the sections in this plot are comprised of the beam fluxes elements of  $(\nu_e \bar{\nu}_e \nu_\mu \bar{\nu}_\mu)$  respectively.

### 8.3 Neutrino cross section uncertainty

A group of default neutrino cross section values are used to make up the nominal Monte Carlo from which the tagging efficiency is calculated. The values of the parameters that determine the cross sections are shown in Table 8.1. Each of the parameter values relate to a specific interaction type and are either a normalisation parameter or a parameter which shows a kinematic dependence.

For charged current quasi-elastic interactions, the uncertainty is described by the Fermi momentum of the oxygen nucleus, ( $p_F^O$ ), the binding energy of the

oxygen nucleus, ( $E_B^O$ ) and the axial mass  $M_A^{CCQE}$ . The axial mass for CCQE interactions relates to the axial form factor which along with vector form factors is proportional to the cross section of the interactions. For neutrino interactions where two nucleons produce two holes (2p2h), an overall normalisation parameter takes the uncertainty of these interactions into account. For  $CC$  and  $NC1\pi$  interactions, the uncertainty is described by the isospin background, the axial form factor  $C_{A5}^{RES}$  which just like for CCQE interactions relates to the axial mass  $M_A^{RES}$ . For neutral current and charged current interactions (both elastic and inelastic) there are normalisation parameters and energy dependent parameters to take the uncertainty into account. Finally, for charged current interactions with electron neutrinos, the braking radiation from the lepton in the final state is also considered when calculating the uncertainty and is treated using a normalisation parameter.

The Monte Carlo re-weighting method is used to reweight the nominal Monte Carlo on an event by event basis with each parameter value being increased and decreased by its uncertainty, and for each reweighted Monte Carlo the equivalent tagging efficiency value is extracted. Equation (8.7) shows how the fractional discrepancies are extracted from the nominal and reweighted tagging efficiency values.

$$\delta_i(\pm\sigma) = \frac{T_i(\pm\sigma) - T_{\text{nom}}}{T_{\text{nom}}} \quad i \in \{ \text{parameters} \} \quad (8.7)$$

Figure 8.3 shows the reweighted Monte Carlo fractional uncertainty plotted for the FHC sample. Since this sample contains a lot of NCother interactions, the uncertainty for this interaction type is greater than for the others.

Parameter	Interaction	Type	Value
$p_F^O$	CCQE	$^{16}\text{O}$ Fermi momentum	$225 \pm 31\text{MeV}/c$
$E_B^O$	CCQE	$^{16}\text{O}$ binding energy	$27 \pm 9\text{MeV}$
$M_A^{CCQE}$	CCQE	Axial mass	$1.2 \pm 0.41\text{GeV}/c^2$
$2p2h$	$2p2 h$	Normalization par.	$1.0 \pm 1.0$
$C_{A5}^{RES}$	CC and NC1 $\pi$	Axial form factor	$1.01 \pm 0.12$
$M_A^{RES}$	CC and NC1 $\pi$	Axial mass	$0.95 \pm 0.15\text{GeV}/c^2$
$BG_A^{RES}$	CC and NC1 $\pi$	I = 1/2 continuum background	$1.3 \pm 0.2$
CC other	CC other	E-dependent par.	$0.0 \pm 0.4$
CC elastic	CC elastic	Normalization par.	$1.0 \pm 0.3$
NC other	NC other	E-dependent par.	$1.0 \pm 0.3$
NC elastic	NC elastic	Normalization par.	$1.0 \pm 0.3$
FSe $^-$ -Bremsstrahlung	CC $\nu_e$	Normalization par.	$1.00 \pm 0.03$

Table 8.1: Neutrino cross section parameters

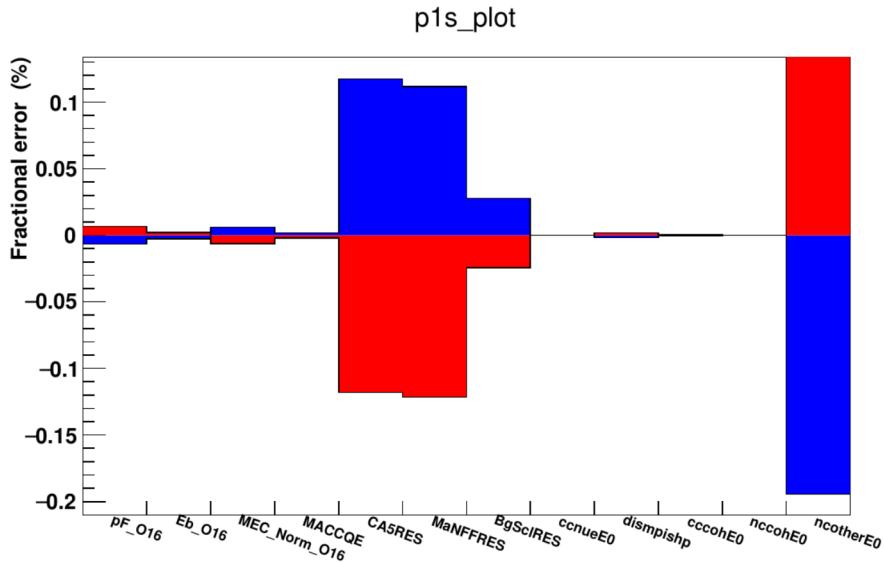


Figure 8.3: Tagging efficiency uncertainty caused by the cross-section parameters variations for the FHC mode

## 8.4 Pion final state interaction (FSI) and secondary interaction (SI) uncertainties

Even though this is an analysis concerned with neutral current quasi elastic interactions, pion events are a significant contribution to the background, and as a result it is important to examine the pion interaction uncertainties both for final state interactions and secondary interactions as their trajectories span the detector.

The neutrino-nucleus interaction simulator used in this analysis (NEUT) handles pion final state interactions and secondary interactions using a cascade model. This cascade model contains parameters which will have uncertainties on them and these will be transferred to a possible change in the tagging efficiency.

Depending on the momentum of the pions, different interaction types occur in the model. For pions with a momentum less than 500 MeV, the interactions in place are absorption (ABS), quasi-elastic scattering (QE) and charge exchange (CX).

Absorption occurs when the incident pion is absorbed by the nucleus and no pions remain in the final state. Quasi-elastic (QE) scattering occurs when there is only one pion observed in the final state and it has the same charge as the incident beam. Charge exchange occurs when the charged pion interacts wth he nucleus and a single  $\pi_0$  can be seen in the final state.

For pions with a momentum of greater than 500 MeV, a different set of interactions are used. Inelastic interactions (INEL) can now produce hadrons and replace absorption processes, but quasi-elastic scattering (QE) and charge exchange (CX) will still occur. The final state interaction parameters and the pion momentum range they are used in can be seen in Table 8.2. Each parameter

scales the relevant very small probability of the charged pion interaction at every stage of the intra-nuclear cascade, aside from the parameter for charge exchange which scales only the fraction of low momentum QE scattering.

Parameter	Description	Momentum Region (MeV/c)
$f_{\text{ABS}}$	Absorption	< 500
$f_{\text{QE}}$	Quasi-elastic scatter	< 500
$f_{\text{CX}}$	Single charge exchange	< 500
$f_{\text{QE}H}$	Quasi-elastic scatter	> 500
$f_{\text{CX}H}$	Single charge exchange	> 500
$f_{\text{INEL}}$	Hadron (N + n $\pi$ ) production	> 500

Table 8.2: Table showing the pion final state interaction parameters in NEUT and the pion momentum range they are used in

A set of paramater variations which determine a surface in paramater space have been estimated by pion scattering experiments, the values for which are shown in Table 8.3. The  $1\sigma$  surface has been explored using the nominal Monte Carlo re-weighting method and the analagous tagging efficiency uncertainty is shown in Equation 8.8, and the uncertainty stemming from the models shown in Table 8.3 is shown in Figure 8.4

$$\delta_i = \frac{T_i - T_{\text{nom}}}{T_{\text{nom}}} \quad i \in \text{ parameter vector} \quad (8.8)$$

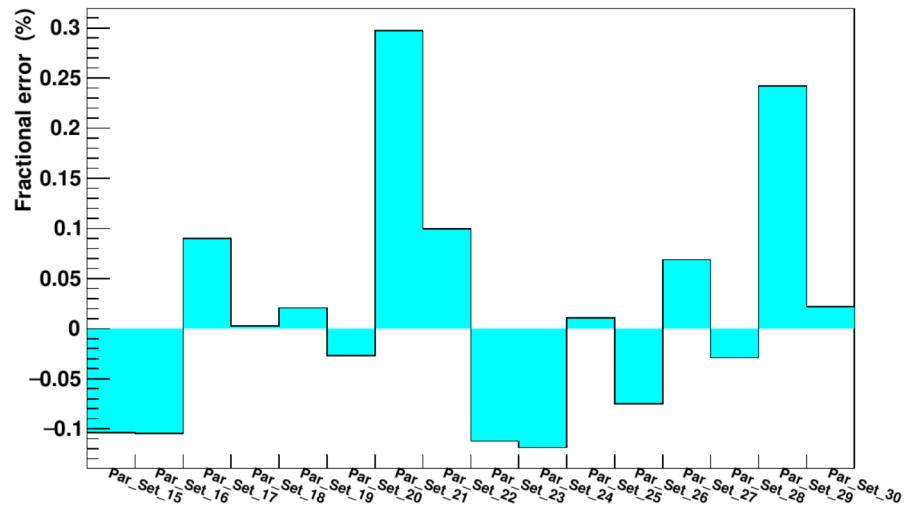


Figure 8.4: Tagging efficiency fractional uncertainty caused by the variation in the FSI/SI model parameters for the FHC mode.

Set	ABS	QE	CX	INEL	QEH	CXH
Nominal	1.1	1.0	1.0	1.0	1.8	1.8
	0.7	0.6	0.5	1.5	1.1	2.3
	0.7	0.6	1.6	1.5	1.1	2.3
Hadron production Up	1.6	0.7	0.4	1.5	1.1	2.3
	1.6	0.7	1.6	1.5	1.1	2.3
	0.6	1.4	0.6	1.5	1.1	2.3
	0.7	1.3	1.6	1.5	1.1	2.3
	1.5	1.5	0.4	1.5	1.1	2.3
	1.6	1.6	1.6	1.5	1.1	2.3
	0.7	0.6	0.5	0.5	2.3	1.3
	0.7	0.6	1.6	0.5	2.3	1.3
Hadron production Down	1.6	0.7	0.4	0.5	2.3	1.3
	1.6	0.7	1.6	0.5	2.3	1.3
	0.6	1.4	0.6	0.5	2.3	1.3
	0.7	1.3	1.6	0.5	2.3	1.3
	1.5	1.5	0.4	0.5	2.3	1.3
	1.6	1.6	1.6	0.5	2.3	1.3

Table 8.3: Pion FSI/SI model parameter nominal value and variations grouped according to inelastic hadron production value

## 8.5 Nucleon final state interactions

Uncertainties regarding the nucleon final state interactions can change the number of nucleons knocked out of  $^{16}\text{O}$ , therefore how the tagging efficiency is changed due to the variation in nucleon final state interactions needs to be investigated. This uncertainty is extracted using GENIE, a Monte Carlo event generator which contains the INTRANUKE (hA) intranuclear transport model. The uncertainties in the total scattering probability for hadrons inside the target nuclei ( $x_{mfp}^N$ ) and the uncertainties in the likelihood of each hadron rescattering method: (elastic ( $x_{el}^N$ ), inelastic ( $x_{inel}^N$ ), charge exchange ( $x_{cex}^N$ ), pion production ( $x_\pi^N$ ) and absorption ( $x_{abs}^N$ )) are taken into account. The frac-

tional uncertainties for these modes for pions is shown in Table 8.4.

Abbreviation	Description of uncertainty	Fractional uncertainty
$x_{mfp}^N$	Nucleon mean free path (total rescattering probability)	±20%
$x_{cex}^N$	Nucleon charge exchange probability	±50%
$x_{el}^N$	Nucleon elastic reaction probability	±30%
$x_{inel}^N$	Nucleon inelastic reaction probability	±40%
$x_{abs}^N$	Nucleon absorption probability	±20%
$x_\pi^N$	Nucleon $\pi$ -production probability	±20%

Table 8.4: Nucleon final state interaction parameters of the hA model executed inside GENIE.

A nominal GENIE Monte Carlo sample is generated (different from the previously used NEUT Monte Carlo) and this shifted using the re-weighting method to a varied GENIE Monte Carlo by individually increasing and decreasing the parameters in Table 8.4 by its error. For each shifted Monte Carlo produced, the fractional uncertainty can be written as in Equation 8.9.

$$\delta_i(\pm\sigma) = \frac{T_i(\pm\sigma) - T_{\text{nom}}}{T_{\text{nom}}} \quad i \in \{ \text{parameters} \} \quad (8.9)$$

The tagging efficiency fractional uncertainties are displayed in Figure 8.5, showing which parameter from Table 8.4 each uncertainty has arisen from.

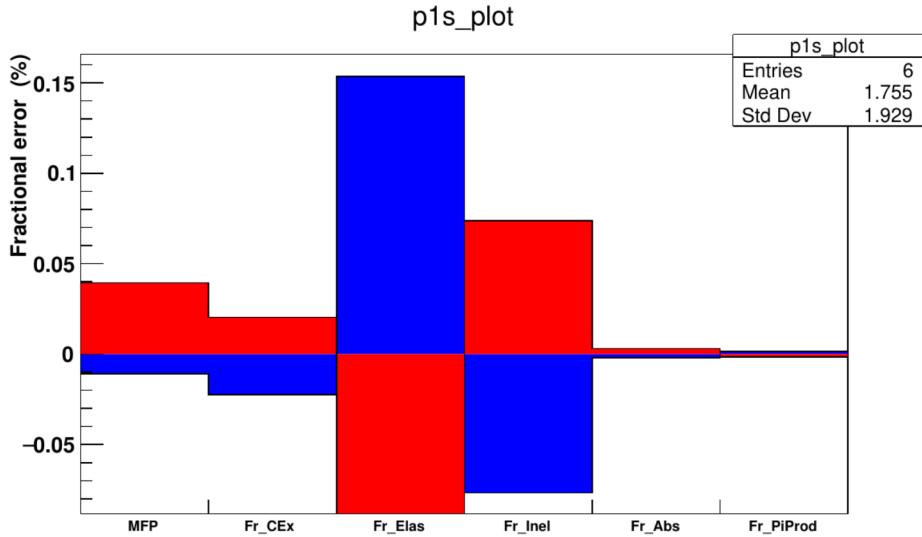


Figure 8.5: Tagging efficiency fractional uncertainties caused by the nucleon final state interaction model parameter variation for the FHC mode

## 8.6 Muon and pion capture on Oxygen-16

Neutrons are produced from negative muon capture on  $^{16}\text{O}$  as shown in Equation 8.10.



Direct neutrons are produced from pion capture on  $^{16}\text{O}$ , but also a number of evaporation neutrons that leave the nucleus. For the capture of muons and pions on  $^{16}\text{O}$ , the energy spectra of the neutrons produced have been measured: for muons the spectra can range up to 15 MeV, while in the case of pions the spectra can reach up to 100 MeV.

Geant4 simulates the capture processes for muons and pions, but there are alternate models that can be used: for example, the Chiral Invariant Phase Space (CHIPS) model for muon captures (based on non perturbative QCD) and

two different routines for pion capture, one which is based on CHIPS and one based on intra-nuclear cascade.

Because any change in the model can alter the energy spectra of the neutrons, these alternative functions can be used to estimate the fractional uncertainties for the tagging efficiency. This is done by using the MC regeneration method, where the nominal Monte Carlo is regenerated by replacing the default Geant4 routines with the alternative models. For the alterantive muon capture model and the two alternative pion capture models, the fractional discrepancies are shown in Equation 8.11.

$$\begin{aligned}\delta_{muonCHIPS} &= \frac{T_{muonCHIPS} - T_{nom}}{T_{nom}} \\ \delta_{pionCHIPS} &= \frac{T_{pionCHIPS} - T_{nom}}{T_{nom}} \\ \delta_{pionBert} &= \frac{T_{pionBert} - T_{nom}}{T_{nom}}\end{aligned}\quad (8.11)$$

Figure 8.6 shows the fractional uncertainties caused by muon and pion capture on oxygen for each model.

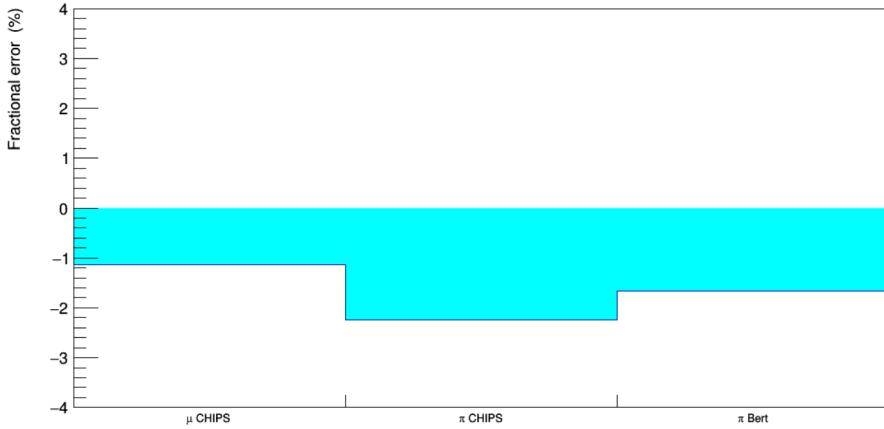


Figure 8.6: Fractional uncertainties in the tagging efficiency caused by muon and pion capture on the muon capture CHIPS model, pion capture CHIPS model and the pion capture Bertini model.

## 8.7 Nucleon SI

Uncertainties in how Nucleon SI interactions are modelled can affect the tagging efficiency, due to how these uncertainties can affect the final number of nucleons present in the simulation and how far they travel in the detector medium. There are two Monte Carlo code systems used in order to determine how nucleons travel in the simulation, MICAP and HETC. These come as part of GCALOR, used to determine the energy and direction of incident hadrons, leptons and photons. MICAP (Monte-carlo-Ionization-Chamber-Analysis-Program), which simulates interactions based on calculated cross section and angular distributions for secondary particles, and is called for neutrons with a kinetic energy below 20 MeV. HETC on the other hand, is the High-Energy-Transport-Code, and is responsible for transporting charged hadrons above 20 MeV (up to an energy of 10 GeV) through the detector medium.

### 8.7.1 MICAP uncertainty calculation

A library of cross-section data called ENDF (Evaluated Nuclear Data Files) are used by MICAP to determine what processes the neutrons go through in the detector medium, and their respective cross sections. There are two versions of libraries which are used in evaluating the uncertainty in the tagging efficiency, version B release V (released in 1978) and version B release VIII, released in 2018. There is very little difference between the total neutron on hydrogen cross sections between the two versions of the code, however, in the energy range of 0.1 MeV to 20 MeV, there are differences of up to 40% in the cross-sections of neutron on oxygen cross sections. Both an elastic and inelastic part comprise the total cross-section of a nucleus, and these can effect the way neutrons travel in the simulated detector medium and also the way secondary particles from these interactions are distributed. The way these inelastic processes are simulated

depend on the nuclei involved: hydrogen nuclei capture the neutrons in the range ( $10^{-11}$ ,  $10^{-7}$ ) MeV, while neutron captures on oxygen mainly happen in the 4 MeV to 20 MeV energy region. To calculate the way uncertainties arise from the way MICAP simulates neutron captures, the ENDF-V library is replaced by ENDF-VIII, and the tagging efficiency using this library ( $T_{VIII}$ ) is evaluated from regenerated MC. Equation ?? shows the fractional uncertainty  $\delta_{MICAP}$  regarding MICAP.

$$\delta_{mcp} = \frac{T_{VIII} - T_{nom}}{T_{nom}} \quad (8.12)$$

### 8.7.2 HETC uncertainty calculation

Due to there being no experimental data for cross-section calculations on nucleon-oxygen scattering in the energy range at which T2K functions, experimental data of proton-carbon scattering is used to assign error on the cross sections. In the proton-carbon scattering analysis, NEUT was used to evaluate the theoretical cross sections of carbon, and this uses cross sections calculated using the Bertini model, which is also used in HETC. The comparison of these calculated cross sections to other theoretical calculations as well as to data, showed that the total cross sections calculated by Bertini need to varied by  $\pm 30\%$  in order to be consistent with them. As a result of this, the Monte Carlo is regenerated twice where the free nucleon-nucleon cross sections are scaled by  $\pm 30\%$ . Equation 8.13 shows the fractional uncertainty  $\delta_{HETC}(\pm)$ , related to the corresponding tagging efficiencies ( $T_{HETC}(\pm)$ ).

$$\delta_{HETC}(\pm) = \frac{T_{HETC}(\pm) - T_{nom}}{T_{nom}} \quad (8.13)$$

## **Chapter 9**

## **Conclusion**

# Bibliography

- [1] K. e. a. Abe. The T2K Experiment. *Nuclear Instruments and Methods in Physics Research Section A: Accelerators, Spectrometers, Detectors and Associated Equipment*, 659(1):106–135, Dec. 2011. arXiv: 1106.1238.
- [2] K. e. a. Abe. The T2K Neutrino Flux Prediction. *Physical Review D*, 87(1):012001, Jan. 2013. arXiv: 1211.0469.
- [3] K. e. a. Abe. Calibration of the Super-Kamiokande Detector. *Nuclear Instruments and Methods in Physics Research Section A: Accelerators, Spectrometers, Detectors and Associated Equipment*, 737:253–272, Feb. 2014. arXiv: 1307.0162.
- [4] K. e. a. Abe. First gadolinium loading to Super-Kamiokande. *Nuclear Instruments and Methods in Physics Research Section A: Accelerators, Spectrometers, Detectors and Associated Equipment*, page 166248, Dec. 2021.
- [5] S. Ando and K. Sato. Relic neutrino background from cosmological supernovae. *New Journal of Physics*, 6:170–170, Nov. 2004. arXiv: astro-ph/0410061.
- [6] O. Benhar, N. Farina, H. Nakamura, M. Sakuda et al. Electron- and neutrino-nucleus scattering in the impulse approximation regime. *Physical Review D*, 72(5):053005, Sept. 2005.

- [7] H. A. Bethe. Supernova mechanisms. *Reviews of Modern Physics*, 62(4):801–866, Oct. 1990.
- [8] L. M. Brown. The idea of the neutrino. *Physics Today*, 31(9):23–28, Sept. 1978. Publisher: American Institute of Physics.
- [9] S. Chakraborty, P. Bhattacharjee et al. Observing supernova neutrino light curve in future dark matter detectors. *Physical Review D*, 89(1):013011, Jan. 2014.
- [10] B. T. Cleveland, T. Daily, J. Raymond Davis, J. R. Distel, K. Lande, C. K. Lee, P. S. Wildenhain et al. Measurement of the solar electron neutrino flux with the homestake chlorine detector. *The Astrophysical Journal*, 496(1):505–526, mar 1998.
- [11] J. Cowan, C. L., F. Reines, F. B. Harrison, H. W. Kruse et al. Detection of the Free Neutrino: A Confirmation. *Science*, 124(3212):103–104, July 1956.
- [12] G. Danby, J.-M. Gaillard, K. Goulian, L. M. Lederman, N. Mistry, M. Schwartz et al. Observation of high-energy neutrino reactions and the existence of two kinds of neutrinos. *Phys. Rev. Lett.*, 9:36–44, Jul 1962.
- [13] G. Danby, J.-M. Gaillard, K. Goulian, L. M. Lederman, N. Mistry, M. Schwartz et al. Observation of high-energy neutrino reactions and the existence of two kinds of neutrinos. *Phys. Rev. Lett.*, 9:36–44, Jul 1962.
- [14] P. Fernández. Status of GADZOOKS!: Neutron Tagging in Super-Kamiokande. *Nuclear and Particle Physics Proceedings*, 273-275:353–360, Apr. 2016.
- [15] S. Fukuda and Y. e. a. Fukuda. The Super-Kamiokande detector. *Nuclear Instruments and Methods in Physics Research Section A: Acceler-*

ators, Spectrometers, Detectors and Associated Equipment, 501(2-3):418–462, Apr. 2003.

- [16] K. Hirata and T. e. a. Kajita. Observation of a neutrino burst from the supernova SN1987A. *Physical Review Letters*, 58(14):1490–1493, Apr. 1987.
- [17] K. S. Hirata and others. Observation in the Kamiokande-II Detector of the Neutrino Burst from Supernova SN 1987a. *Phys. Rev. D*, 38:448–458, 1988.
- [18] Z. Maki, M. Nakagawa et al. Remarks on the Unified Model of Elementary Particles. *Progress of Theoretical Physics*, 28(5):870–880, 11 1962.
- [19] L. Marti and M. e. a. Ikeda. Evaluation of gadolinium’s action on water Cherenkov detector systems with EGADS. *Nuclear Instruments and Methods in Physics Research Section A: Accelerators, Spectrometers, Detectors and Associated Equipment*, 959:163549, Apr. 2020.
- [20] S. L. Meo, C. Massimi et al. Measurement of the neutron capture cross section for  $^{155}\text{Gd}$  and  $^{157}\text{Gd}$  for Nuclear Technology. page 8.
- [21] H. Nagakura and K. Hotokezaka. Non-thermal neutrinos created by shock acceleration in successful and failed core-collapse supernova. *Monthly Notices of the Royal Astronomical Society*, 502(1):89–107, Jan. 2021. arXiv: 2010.15136.
- [22] L. Nanni. Fermi theory of beta decay: A first attempt at electroweak unification. *Advanced Studies in Theoretical Physics*, 13:281, 07 2019.
- [23] H. Nishino, K. Awai, Y. Hayato, S. Nakayama, K. Okumura, M. Shiozawa, A. Takeda, K. Ishikawa, A. Minegishi et al. High-speed charge-to-time converter ASIC for the Super-Kamiokande detector. *Nuclear Instruments and Methods in Physics Research Section A: Accelerators, Spectrometers,*

*Detectors and Associated Equipment*, 610(3):710–717, Nov. 2009. arXiv: 0911.0986.

- [24] B. Pontecorvo. Mesonium and anti-mesonium. *Sov. Phys. JETP*, 6:429, 1957.
- [25] Y. Suzuki. The Super-Kamiokande experiment. *The European Physical Journal C*, 79(4):298, Apr. 2019.
- [26] K. Ueno. Analysis of nuclear de-excitation gamma-rays induced by neutrino-nucleon neutral-current interactions at SK using the T2K RUN-I/II data. page 87.
- [27] T. Vladislavljevic. Constraining the T2K Neutrino Flux Prediction with 2009 NA61/SHINE Replica-Target Data. June 2018. arXiv:1804.00272 [hep-ex, physics:physics].