

Cosmic Bit Balance Sheet

Normalizing the Universe's Entropy to Bits (S_{bits})

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Abstract. We present a self-contained, executable-style ledger of cosmic entropy cast in *bits*. Every sector—black holes, radiation (CMB photons, cosmic neutrinos), baryonic gas (IGM/ICM), compact objects, and Hawking channels—has (i) a single-line formula, (ii) a normalization to bits, and (iii) a plug-in worksheet to evaluate *per cubic Mpc* and *per Hubble volume*. The balance sheet is designed to be falsifiable, updatable, and useful: it exposes which observables dominate the budget and where better data matter.

1 Normalization and cosmology anchors

We keep G, \hbar, k_B, c explicit. Bit units: $S_{\text{bits}} \equiv S/\ln 2$. Our fiducial cosmology is the Rosetta ledger (flat Λ CDM):

$$\Omega_m = \frac{63}{200} = 0.315, \quad \Omega_\Lambda = \frac{137}{200} = 0.685, \quad \frac{\Omega_b}{\Omega_c} = \frac{14}{75}.$$

We adopt the rational Hubble constant $H_0 = \frac{337}{5} \text{ km s}^{-1} \text{ Mpc}^{-1} = 67.4 \text{ km s}^{-1} \text{ Mpc}^{-1}$ (convert with $1 \text{ Mpc} = 3.0856775814913673 \times 10^{22} \text{ m}$). From this:

$$H_0 = \frac{337}{5} \frac{10^3}{3.0856775814913673 \times 10^{22}} \text{ s}^{-1} = 2.1842852410855 \times 10^{-18} \text{ s}^{-1}, \quad (1)$$

$$\rho_c = \frac{3H_0^2}{8\pi G} = 8.5328551637393 \times 10^{-27} \text{ kg m}^{-3}, \quad (2)$$

$$\Omega_c = \Omega_m \frac{75}{89} = 0.26545, \quad \Omega_b = \Omega_m - \Omega_c = 0.04955. \quad (3)$$

The Hubble radius and volume are

$$R_H = \frac{c}{H_0} = 4447.959 \text{ Mpc}, \quad V_H = \frac{4\pi}{3} R_H^3 = 3.686133233074354 \times 10^{11} \text{ Mpc}^3. \quad (4)$$

All sector totals are reported both per Mpc^3 and scaled to V_H .

2 Radiation sectors: CMB photons and cosmic neutrinos

For a relativistic species with temperature T and internal degrees g , the entropy density is

$$s = \frac{2\pi^2}{45} g k_B \left(\frac{k_B T}{\hbar c} \right)^3, \quad s_{\text{bits}} = \frac{s}{k_B \ln 2}. \quad (5)$$

Photons (CMB). With $g_\gamma = 2$ and $T_\gamma = 2.7255 \text{ K}$:

$$s_{\gamma, \text{bits}} = 2.134123646 \times 10^9 \text{ bits m}^{-3} = 6.270053023 \times 10^{76} \text{ bits Mpc}^{-3}. \quad (6)$$

Neutrinos (CNB). Treating three flavors as effectively relativistic with $g_\nu = (7/8) \times 6 = 21/4$ and $T_\nu = (4/11)^{1/3} T_\gamma$:

$$s_{\nu, \text{bits}} = 2.037118026 \times 10^9 \text{ bits m}^{-3} = 5.985050613 \times 10^{76} \text{ bits Mpc}^{-3}. \quad (7)$$

Radiation total. $s_{\text{rad, bits}} = 1.225\,510\,364 \times 10^{77} \text{ bits Mpc}^{-3}$, which integrates over the Hubble volume to

$$S_{\text{rad, bits}}(V_H) = 4.517\,394\,479 \times 10^{88} \text{ bits.} \quad (8)$$

Note. If massive neutrinos are partially nonrelativistic today, the exact s_ν decreases modestly; this ledger keeps the clean relativistic form and flags a neutrino-mass correction as an optional refinement (Appendix B).

3 Baryonic gas sectors (IGM/ICM): Sackur–Tetrode worksheet

For a monatomic ideal gas the entropy density follows Sackur–Tetrode. For a fully ionized primordial plasma with mean molecular weight $\mu \simeq 0.59$ and baryon density $\rho_b = \Omega_b \rho_c$, the particle density is $n = \rho_b/(\mu m_p)$. The entropy density is

$$s \approx n k_B \left[\ln \left(\frac{5/2}{n} \left(\frac{2\pi m k_B T}{h^2} \right)^{3/2} \right) \right], \quad s_{\text{bits}} = s/(k_B \ln 2), \quad (9)$$

with $m \approx \mu m_p$. **Worked anchor (cosmic mean, $T = 2 \times 10^4 \text{ K}$):**

$$n \approx 4.284\,431\,071 \times 10^{-1} \text{ m}^{-3}, \quad s/(n k_B) \approx 77.9204, \quad s_{\text{bits}} \approx 1.415\,045\,404 \times 10^{69} \text{ bits Mpc}^{-3}. \quad (10)$$

This sits *eight orders of magnitude* below the radiation bit density; clusters at higher T move the number but not the hierarchy.

4 Black holes: area \rightarrow bits, per-object and populations

For a hole of mass M ,

$$S_{\text{bits}}(M) = \frac{\mathcal{A}}{4\ell_P^2 \ln 2} = \frac{4\pi G}{\hbar c \ln 2} M^2 \equiv K M^2, \quad K = \frac{4\pi G}{\hbar c \ln 2}. \quad (11)$$

Per-object anchors. $S_{\text{bits}}(1 M_\odot) = 1.513\,322\,012 \times 10^{77}$, $S_{\text{bits}}(10 M_\odot) = 1.513\,322\,012 \times 10^{79}$, $S_{\text{bits}}(4 \times 10^6 M_\odot) = 2.421\,315\,220 \times 10^{90}$.

Population ledger. For a mass function dn/dM (objects per volume per mass), the bit density is

$$s_{\text{BH, bits}} = K \int M^2 \frac{dn}{dM} dM = K n_{\text{BH}} \frac{\langle M^2 \rangle}{\langle M \rangle}, \quad (12)$$

where n_{BH} is the number density. Equivalently, with a empmass density $\rho_{\text{BH}} = \int M dn$ and an *effective mass* $M_{\text{eff}} \equiv \langle M^2 \rangle / \langle M \rangle$, one may write

$$s_{\text{BH, bits}} = K \rho_{\text{BH}} M_{\text{eff}}. \quad (13)$$

Because of the M^2 scaling, entropy is maximized by concentrating fixed mass into the most massive holes: supermassive BHs dominate the cosmic bit budget for any reasonable ρ_{BH} .

Dominance condition (back-of-envelope). Radiation bits per volume are $s_{\text{rad, bits}} \approx 4.171\,241\,672 \times 10^9 \text{ bits m}^{-3}$. Black holes dominate when

$$K \rho_{\text{BH}} M_{\text{eff}} \gtrsim s_{\text{rad, bits}}. \quad (14)$$

For a tiny mass fraction $f_{\text{BH}} \equiv \rho_{\text{BH}}/\rho_c = 10^{-5}$ the threshold is $M_{\text{eff}} \gtrsim \text{kg}$ ($\sim 6.4 \times 10^{-7} M_\odot$)—far *below* typical astrophysical BH masses. Thus even a minute cosmic BH mass density yields a BH entropy that dwarfs the CMB/CNB.

Illustrative population (*example numbers*). If one adopts a fiducial $\rho_{\text{BH}} = 5 \times 10^5 M_{\odot} \text{Mpc}^{-3}$ concentrated at $M_{\text{eff}} = 10^8 M_{\odot}$, then

$$s_{\text{BH, bits}} \approx 7.566\,610\,062 \times 10^{90} \text{ bits Mpc}^{-3} \Rightarrow S_{\text{BH, bits}}(V_H) \approx \text{bits}. \quad (15)$$

Caveat: the inputs are placeholders for illustration; plug your preferred mass function or mass-density measurement into the worksheet below.

5 Compact objects (stars, WDs, NSs): where they sit

Stellar interiors have large *specific* entropies but a tiny cosmic mass fraction and do not compete with BHs or radiation totals. **White dwarfs** and **neutron stars** have lower specific entropies (degenerate equations of state); despite large number counts, their contribution remains subdominant to radiation and negligible vs BHs in any realistic census. A worksheet for degenerate equations of state (polytropic approximations) is provided in Appendix C for completeness.

6 Why a bit ledger? (what this buys you)

- **GSL as a budget:** Cast the Generalized Second Law as $\Delta(S_{\text{out}} + S_{\text{BH}}) \geq 0$ in bits; mergers/accretion/radiation become entry-by-entry transactions.
- **Decision power:** The ledger shows *which measurements matter*. Improving CMB T shifts $\sim 10^{88}$ bits; tuning an SMBH mass function shifts $\sim 10^{100}$ – 10^{103} bits.
- **Rosetta unification:** Ties the (Howtotransduceme file) EKTL (energy \rightarrow knowledge) to the cosmic budget; bit–joule bounds at horizons become global constraints when integrated over populations.
- **Communication:** One unit (bits) across quantum, thermo, and gravity lowers the barrier for new collaborators.

7 How to use this sheet (plug-and-play)

1. **Pick a volume:** per Mpc^3 for density-style arguments or V_H for a Hubble-sphere snapshot.
2. **Radiation:** use the closed forms above for $s_{\gamma, \text{bits}}$ and $s_{\nu, \text{bits}}$ (or rerun with your T_{γ} and N_{eff}).
3. **IGM/ICM:** choose T and overdensity Δ ; scale $n \rightarrow \Delta n$ in Sackur–Tetrode; add cluster components if desired.
4. **Black holes:** provide either dn/dM and integrate $K \int M^2 dn$ or specify $(\rho_{\text{BH}}, M_{\text{eff}})$ and multiply.
5. **Report:** write totals in bits; show a bar chart and a sensitivity table (which parameter moves the most bits?).

Summary tables

Sector	Formula (bits density)	Value / Mpc ³	Notes
CMB photons	$\frac{2\pi^2}{45} \frac{2}{\ln 2} \left(\frac{k_B T_\gamma}{\hbar c}\right)^3$	$6.270\,053\,023 \times 10^{76}$	exact at $T_\gamma = 2.7255$
CNB neutrinos	$\frac{2\pi^2}{45} \frac{(7/8)6}{\ln 2} \left(\frac{k_B T_\nu}{\hbar c}\right)^3$	$5.985\,050\,613 \times 10^{76}$	$T_\nu = (4/11)^{1/3} T_\gamma$
IGM (mean, 2×10^4 K)	ST formula	$1.415\,045\,404 \times 10^{69}$	depends on T , overdensity
BHs (illustrative)	$K \rho_{\text{BH}} M_{\text{eff}}$	$7.566\,610\,062 \times 10^{90}$	choose $\rho_{\text{BH}}, M_{\text{eff}}$

A Constants and conversions

$k_B = 1.380\,649 \times 10^{-23} \text{ J K}^{-1}$, $\hbar = 1.054\,571\,817 \times 10^{-34} \text{ J s}$, $c = 2.997\,924\,58 \times 10^8 \text{ m s}^{-1}$,
 $G = 6.674\,30 \times 10^{-11} \text{ m}^3 \text{ kg}^{-1} \text{ s}^{-2}$, $= 3.085\,677\,581\,491\,367\,3 \times 10^{22} \text{ m}$, $M_\odot = 1.988\,47 \times 10^{30} \text{ kg}$.
One bit equals $k_B \ln 2$ of thermodynamic entropy.

B Neutrino refinement (optional)

If light neutrinos are partially nonrelativistic today, replace the relativistic s_ν with the exact Fermi–Dirac expression at finite mass. The effect is a mild suppression relative to the relativistic value and can be bounded with your favorite $\sum m_\nu$ prior.

C Compact-object entropy (WDs and NSs)

For degenerate objects, use $s \sim \pi^2 n k_B (k_B T / \epsilon_F)$ for a Fermi gas (electrons in WDs, neutrons in NSs). Even at stellar-core temperatures the small T / ϵ_F keeps s subdominant cosmologically.

D BH population worksheet (mass function route)

Provide dn/dM for SMBHs (e.g., a Schechter or lognormal fit) and evaluate $K \int M^2 dn$. Report both $s_{\text{BH, bits}}$ and $S_{\text{BH, bits}}(V_H)$.

E EKTL link (energy \rightarrow knowledge)

At any horizon, the maximum bit throughput obeys $\dot{N}_{\text{bits}} \leq P / (k_B \ln 2)$ (greybody factors aside). Integrated over cosmic populations, this provides a global constraint on energy-to-information transduction.