A DIRECT MEASUREMENT OF THE INTERGALACTIC MEDIUM OPACITY TO H I IONIZING PHOTONS

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ABSTRACT

We present a new method to directly measure the opacity from H I Lyman limit (LL) absorption $\kappa_{\rm LL}$ along quasar sight lines by the intergalactic medium (IGM). The approach analyzes the average ("stacked") spectrum of an ensemble of quasars at a common redshift to infer the mean free path $\lambda_{\rm mfp}^{912}$ to ionizing radiation. We apply this technique to 1800 quasars at $z = 3.50$–4.34 drawn from the Sloan Digital Sky Survey (SDSS), giving the most precise measurements on $\kappa_{\rm LL}$ at any redshift. From $z = 3.6$ to 4.3, the opacity increases steadily as expected and is well parameterized by $\lambda_{\rm mfp}^{912} = \lambda_0 - b_\lambda(z - 3.6)$ with $\lambda_0 = (48.4 \pm 2.1)\, h_{72}^{-1}$ Mpc and $b_\lambda = (38.0 \pm 5.3)\, h_{72}^{-1}$ Mpc (proper distance). The relatively high $\lambda_{\rm mfp}^{912}$ values indicate that the incidence of systems which dominate $\kappa_{\rm LL}$ evolves less strongly at $z > 3$ than that of the Lyα forest. We infer a mean free path three times higher than some previous estimates, a result which has important implications for the photoionization rate derived from the emissivity of star-forming galaxies and quasars. Finally, our analysis reveals a previously unreported, systematic bias in the SDSS quasar sample related to the survey's color targeting criteria. This bias potentially affects all $z \sim 3$ IGM studies using the SDSS database.

Key words: intergalactic medium – large-scale structure of universe – quasars: absorption lines

1. INTRODUCTION

The observed high transmission of $z \sim 3$ quasars at rest wavelengths λ_r blueward of H I Lyα reveals that the intergalactic medium (IGM) is highly ionized (Gunn & Peterson 1965). The presence of the Lyα forest demands an intense, extragalactic ultraviolet background (EUVB) radiation field. The quasars themselves provide a significant fraction of the required ionizing flux, buoyed by the emission from more numerous yet fainter star-forming galaxies. Several recent studies have argued that the latter population dominates the EUVB at $z \gtrsim 3$ (Faucher-Giguère et al. 2008a; Cowie et al. 2009; Dall'Aglio et al. 2009), where the quasar population likely declines (e.g., Fan et al. 2004). These assertions, however, hinge on the opacity of the IGM to ionizing radiation via H I Lyman limit absorption ($\kappa_{\rm LL}$), which directly impacts estimates of the EUVB measured from the integrated quasar and stellar ionizing emissivity.

Traditionally, $\kappa_{\rm LL}$ has been estimated from the incidence of so-called Lyman limit systems (LLSs) via surveys of quasar spectroscopy (e.g., Lanzetta 1991; Storrie-Lombardi et al. 1994; Péroux et al. 2003). Observationally, one can rather easily identify systems with large optical depths $\tau_{912} \gtrsim 2$ and the majority of these surveys have probed to this limit. The integrated opacity of the IGM, however, includes and is likely dominated by gas with $\tau_{912} < 1$, the so-called partial LLSs and Lyα forest clouds. Systems with these H I column densities ($N_{\rm H I} \approx 10^{14}$–$10^{17}\,{\rm cm}^{-2}$) are especially difficult to survey because the strong lines of the Lyman series (e.g., Lyα, Lyβ) lie on the flat portion of the curve of growth, and they exhibit only weak absorption at the Lyman limit. Therefore, current estimates of $\kappa_{\rm LL}$ are based on an extrapolation/interpolation of the frequency of systems with $N_{\rm H I} < 10^{14}\,{\rm cm}^{-2}$ and $N_{\rm H I} > 10^{17.5}\,{\rm cm}^{-2}$ (Madau et al. 1999; Schirber & Bullock 2003; Faucher-Giguere et al. 2009). Current constraints on $\kappa_{\rm LL}$ span over a magnitude of uncertainty, especially at $z \gtrsim 4$.

In this Letter, we introduce a new technique to estimate $\kappa_{\rm LL}$ that avoids the traditional line-counting statistics of the IGM. We analyze the average rest-frame spectra of 1800 $z > 3.5$ quasars drawn from the Sloan Digital Sky Survey Data Release 7 (SDSS-DR7; Abazajian et al. 2009). As an ensemble at a common redshift, these "stacked" spectra show the exponential drop in flux at $\lambda < \lambda_{912} = c/\nu_{912} = 911.76$ Å from the integrated opacity of the IGM. We precisely evaluate $\kappa_{\rm LL}$ in a series of small redshift intervals covering $z \approx 3.6$–4.3 to explore redshift evolution. We adopt a cosmology with $H_0 = 72\, h_{72}$ km s^{-1} Mpc^{-1}, $\Omega_{\rm m} = 0.3$, and $\Omega_\Lambda = 0.7$, and report proper lengths unless specified.

2. METHODOLOGY

The H I Lyman limit opacity of the IGM has traditionally been expressed as an effective optical depth $\tau_{\rm eff,LL}$ estimated from the observationally constrained $N_{\rm H I}$ frequency distribution $f(N_{\rm H I}, z)$. An ionizing photon ($\nu \geqslant \nu_{912}$) emitted from a quasar with redshift $z = z_q$ will redshift to 1 Ryd at $z = z_{912} \equiv (\nu_{912}/\nu)(1 + z_q) - 1$. The effective optical depth that this photon experiences by the IGM Lyman limit opacity is then (compared to Meiksin & Madau 1993)

$$\tau_{\rm eff,LL}(z_{912}, z_q) = \int_{z_{912}}^{z_q} \int_0^\infty f(N_{\rm H I}, z') \\ \times \{1 - \exp[-N_{\rm H I}\sigma_{\rm ph}(z')]\}dN_{\rm H I}dz', \quad (1)$$

where $\sigma_{\rm ph}$ is the photoionization cross section evaluated at the photon frequency. In practice, this approach is subject to large uncertainties because (1) $f(N_{\rm H I}, z)$ is poorly constrained for systems with $\tau_{912} \lesssim 1$, (2) observational surveys rarely measure $f(N_{\rm H I}, z)$ at the same epoch, forcing interpolation and extrapolation, and (3) estimates of $f(N_{\rm H I}, z)$ for the LLS may always be subject to large systematic uncertainty (Prochaska et al. 2009, in preparation). This traditional approach may never yield a precise and robust estimate of $\tau_{\rm eff,LL}$.

Our new approach is to directly measure $\tau_{\rm eff,LL}$ through analysis of averaged ensembles of quasar spectra. In Figure 1, we present the stacked spectrum of 150 mock quasar spectra at $z \approx 3.6$. Each spectrum was given a unique emission redshift z_q
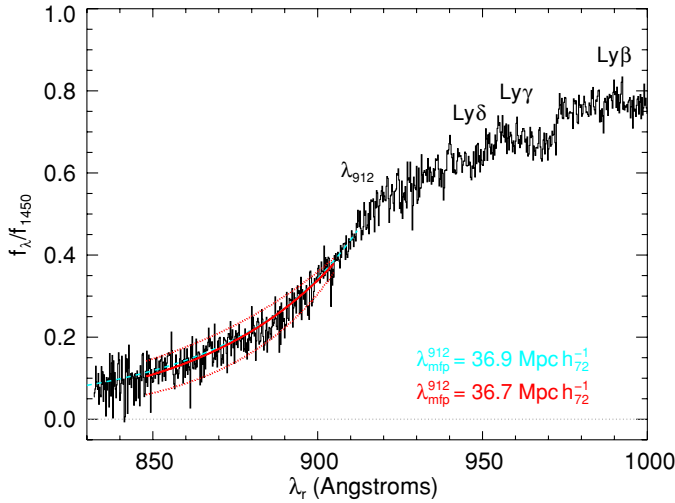
Figure 1. Stacked spectrum (normalized at $\lambda_r = 1450$ Å) from 150 mock quasar spectra with IGM absorption derived from an assumed $f(N_{H_I}, z)$ distribution. The adopted emission redshifts and S/N of the spectra were taken to match our SDSS quasar sample at $z = [3.59, 3.63]$. The sharp decline in the flux at $\lambda_r \approx 912$ Å is associated first with the higher order Lyman series opacity (for $\lambda_r \gtrsim 912$ Å) and then Lyman limit continuum opacity (for $\lambda_r < 912$ Å). The dashed (cyan) curve overplotted on the spectrum is the predicted relative flux corresponding to $\tau_{\mathrm{eff,LL}}$ (Equation (1)), normalized at $\lambda_r = 912$ Å. On top of this curve is the best-fit model of the flux evaluated over the interval $\lambda_r = 848$–905 Å using our new methodology (Equation (6)). The resulting estimates for the mean free path $\lambda_{\mathrm{mfp}}^{912}$ are indicated in the figure and are in excellent agreement. Finally, the dotted curves show the predicted models after offsetting $\lambda_{\mathrm{mfp}}^{912}$ by ± 10 Mpc and indicate the sensitivity of our technique to measuring $\lambda_{\mathrm{mfp}}^{912}$.

and spectral energy distribution[3] (SED; normalized at 1450 Å), and then was blanketed with Lyman series and Lyman limit absorption from an assumed $f(N_{H_I}, z)$ distribution (Dall'Aglio et al. 2008; Worseck & Prochaska 2009, in preparation). The spectra were degraded to the nominal spectral resolution of the SDSS spectrometer (FWHM = 150 km s^{-1}) and Gaussian noise was added to give a distribution[4] of signal-to-noise ratio (S/N) values at $\lambda_r = 1450$ Å. The data were then averaged without weighting.

Inspecting Figure 1, one identifies the effective opacity of the Lyβ forest at $\lambda_r \approx 1000$ Å and corresponding decrements in the spectrum at Lyγ and Lyδ. One then observes a steep drop in the flux starting at $\lambda_r \approx 920$ Å due to the opacity of higher order Lyman series lines of optically thick absorbers (e.g., damped Lyα systems). The continued decline at $\lambda_r < \lambda_{912}$, however, is dominated by the continuum opacity of H I. At all wavelengths, the scatter in the stacked spectrum is due to small-scale variance in IGM absorption, not the noise in the individual spectra.

Overplotted in Figure 1 is the flux model $f = f_{912} \exp[-\tau_{\mathrm{eff,LL}}(z)]$ with $\tau_{\mathrm{eff,LL}}$ evaluated from the input $f(N_{H_I}, z)$ distribution (Equation (1)) and f_{912}, the flux at $\lambda = \lambda_{912}$, estimated from the data. This is a good model of the stacked spectrum; even though the underlying average SED evolves as $f_\lambda \propto \lambda^{2.4}$, the analysis is performed over too small a wavelength interval to note its evolution. The evaluation gives $\tau_{\mathrm{eff,LL}} = 1$ at $(z_q - z_{912}) = 0.22$ corresponding to a proper mean free path $\lambda_{\mathrm{mfp}}^{912} = 36.9\, h_{72}^{-1}$ Mpc at $z = 3.6$.

Now consider an alternate evaluation of $\tau_{\mathrm{eff,LL}}$ which follows the standard definition of optical depth,

$$\tau_{\mathrm{eff,LL}}(r, \nu) = \int_0^r \kappa_{\mathrm{LL}}(r', \nu)\, dr', \qquad (2)$$

where the integral is evaluated to an arbitrary proper distance from the quasar. In an expanding universe, an ionizing photon emitted by the quasar will be attenuated by the Lyman limit opacity κ_{LL} until it is redshifted to $h\nu = 1$ Ryd at $z = z_{912}$. If the photon is not absorbed by IGM line opacity from gas at $z < z_{912}$, it may be observed today at a wavelength $\lambda_{\mathrm{obs}} = (1 + z_{912})c/\nu_{912}$. During the photon's travel from z_q to z_{912}, the opacity κ_{LL} evolves because of the decreasing frequency (redshift) and also from changes to the physical conditions of the universe (e.g., the expanding proper distance). We separate the frequency and radial dependencies in the opacity as follows,

$$\kappa_{\mathrm{LL}}(r, \nu) \equiv \tilde{\kappa}_{912}(r) \left(\frac{\nu}{\nu_{912}} \right)^{-3}, \qquad (3)$$

where the frequency dependence related to σ_{ph} is approximate. This treatment also ignores stimulated emission, i.e., it assumes $\tau_{\mathrm{eff,LL}}$ is dominated by "clouds" with $\tau_{912} \lesssim 1$.

In principle, one can adopt any radial dependence for κ_{LL}. Expressing Equation (3) in redshift space, we have

$$\kappa_{\mathrm{LL}}(z) = \tilde{\kappa}_{912}(z) \left(\frac{1 + z}{1 + z_{912}} \right)^{-3}. \qquad (4)$$

With this functional form for the Lyman limit opacity, it is straightforward to integrate Equation (2) by adopting a Friedman–Walker cosmology where

$$\frac{dr}{dz} \equiv \frac{c}{H(z)(1 + z)} = \frac{c/H_0}{(1 + z)\sqrt{\Omega_m (1 + z)^3 + \Omega_\Lambda}}. \qquad (5)$$

At $z > 3$, the universe is matter dominated and we can express $dr/dz \approx c/(H_0 \Omega_m^{1/2})(1 + z)^{-5/2}$. Altogether, we find

$$\tau_{\mathrm{eff,LL}}(z_{912}, z_q) = \frac{c}{H_0 \Omega_m^{1/2}} (1 + z_{912})^3$$
$$\times \int_{z_{912}}^{z_q} \tilde{\kappa}_{912}(z')\, (1 + z')^{-11/2}\, dz'. \qquad (6)$$

In practice, we find that the analysis of a single stacked spectrum does not constrain the redshift evolution in κ_{LL}. Therefore, we have simply parameterized κ_{LL} by its value at $z = z_q$, i.e., $\tilde{\kappa}_{912}(z') = \kappa_{z_q}$. The thin solid curve in Figure 1 shows the resulting flux model for $\tau_{\mathrm{eff,LL}}$ for a best-fit value $\kappa_{z_q} = 0.028\, h_{72}$ Mpc^{-1}. The corresponding mean free path $\lambda_{\mathrm{mfp}}^{912} = 1/\kappa_{z_q} = 35.2\, h_{72}^{-1}$ Mpc is in excellent agreement with the traditional evaluation. We stress that the analysis was performed without any consideration of the quasar SEDs nor any consideration of evolution in the Lyman series line opacity. Although these contribute to the observed flux in the stacked spectrum, the exponential drop due to $\tau_{\mathrm{eff,LL}}$ dominates over this and any other astrophysical aspect.

3. RESULTS

Our new approach provides a tight constraint on the effective mean free path near the quasar emission redshift (at $z \approx z_q$).

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[3] Quasar SEDs do not show strong emission features at $\lambda_r < 912$ Å (Telfer et al. 2002).
[4] The S/N and emission redshift distributions correspond to the third bin in the observational analysis that follows.
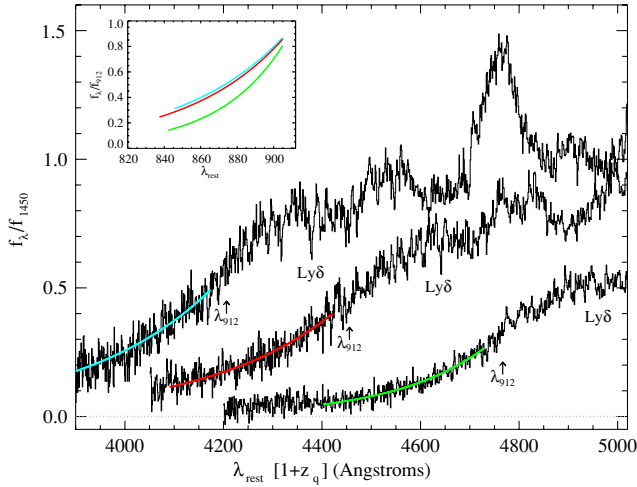
Figure 2. Stacked spectrum for three of our redshift bins (cyan: $z = [3.59, 3.63]$; red: $z = [3.86, 3.92]$; green: $z = [4.13, 4.34]$) plotted against rest-frame wavelengths redshifted to the mean quasar redshift for the bin. Overplotted on each spectrum is our best-fit model for the absorbed flux below λ_{912} due to the Lyman limit opacity. These same curves are shown in the sub-panel against rest wavelengths. The emission lines are from Lyman series and metal-line transitions.

Because the stacked spectrum covers only several tenths in redshift below the Lyman limit, it imposes a very weak constraint on the redshift evolution of κ_{LL}. Instead, one must evaluate the stacked spectrum of quasars at a range of emission redshifts.

We apply the methodology to 1800 quasar spectra drawn from the SDSS-DR7. We began with the vetted quasar list from our survey of LLSs (Prochaska et al. 2009, in preparation) which avoids all purported quasars in the SDSS-DR7 that have erroneous redshift estimates or are not bona-fide quasars. We have also ignored all quasars with strong associated absorption in the C IV, N V, and/or O VI doublets (e.g., broad absorption line systems). We have not, however, removed quasars with evident Lyman limit absorption at $z \approx z_q$. It is our goal to estimate the entire Lyman limit opacity that quasars experience, except for the influence of gas on parsec scales. Therefore, our estimates of κ_{LL} include opacity from the quasar's local galactic environment, i.e., its proximity region.

The sample was further limited to the following criteria: (1) $z_q \geqslant 3.5$ to insure significant coverage of the Lyman limit in the SDSS spectra and minimize the likelihood that LLSs bias the quasar target probability (but see below); (2) S/N $\geqslant 4$ at $\lambda_r = 1450$ Å; (3) $z_q < 4.35$ to insure that a stack of 150 quasars covers a redshift interval $\Delta z < 0.4$. Starting at $z = 3.5$, we constructed a series of bins of 150 quasars each to produce a stacked spectrum. Each quasar spectrum was normalized by the observed flux at $\lambda_r = 1450$ Å (in a 20 Å window) and shifted to its rest frame (nearest pixel). The full ensemble was then averaged (without weighting), ignoring bad pixels. A sample of three of the stacked spectra is given in Figure 2.

As noted above, our analysis of $\tau_{eff,LL}$ includes contributions from the quasar's proximity region. In general, this corresponds to ≈ 10 Mpc or $\Delta\lambda_r \approx 9$ Å (e.g., Dall'Aglio et al. 2008). From the stacked spectrum, Figure 2 shows that the flux at the "edge" of the proximity region (i.e., at $\lambda_r \approx 900$ Å) is ≈ 0.75 times the flux at 912 Å, giving $\tau_{eff,LL} \approx 0.3$. In the absence of the IGM beyond the proximity region, the flux would begin to recover shortward of 900 Å as ν^{-3}. It is evident, therefore, that opacity from the "true" IGM dominates our analysis. Furthermore, we find that our analysis yields models that extrapolate well to 912 Å, suggesting that the opacity of the proximity region
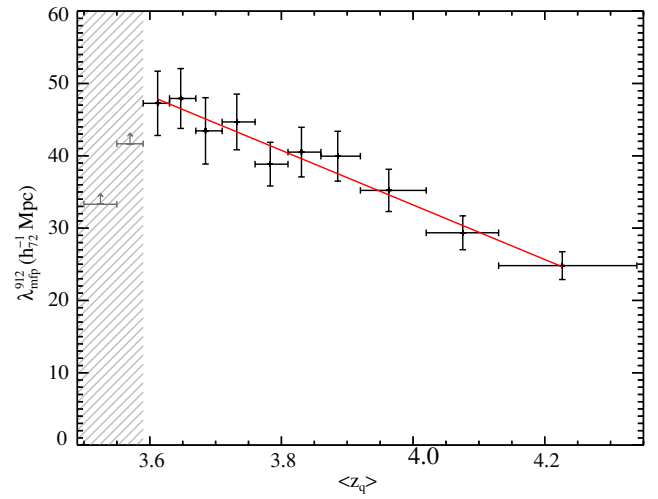

Figure 3. Estimates of the mean free path $\lambda_{mfp}^{912} \equiv (1/\kappa_{z_q})$ vs. redshift. Error bars were estimated from standard bootstrap analysis. The red curve is a simple linear regression of the binned data: $\lambda_{mfp}^{912} = \lambda_0 - b_\lambda(z - 3.6)$ with $\lambda_0 = (48.4 \pm 2.1) h_{72}^{-1}$ Mpc and $b_\lambda = (38.0 \pm 5.3) h_{72}^{-1}$ Mpc. These are the most precise estimates of λ_{mfp}^{912} to date and provide the first robust description of its redshift evolution. The values at $z < 3.6$ (shaded region) are severely affected by a systematic bias in the SDSS database (see the text) and are presented here only as lower limits.

follows the behavior of the general IGM. This is consistent with our finding that there are no strong differences in the incidence of LLSs, near quasars (Prochaska et al. 2009, in preparation).

Using the model for $\tau_{eff,LL}$ (Equation (6)), we have fitted the data to evaluate the opacity κ_{z_q} at $z \approx z_q$ in a series of redshift intervals, each containing 150 quasars. We minimized χ^2 over rest-wavelength intervals starting at 905 Å (to minimize bias from strong Lyman series absorption in the proximity zone of the quasar) and extending down in wavelength corresponding to the larger of 3900 Å and $z_q - z_{912} = 0.4$. The scatter in the stacked spectrum was estimated locally in 21 pixel bins centered at each data point and presumed to be Gaussian. The flux at λ_{912} was estimated from the data but allowed to vary by 10% when minimizing χ^2. Because the observed scatter in the stacked spectrum is systematic (related to stochasticity in the IGM) and is not included in our model, one cannot estimate $\sigma(\kappa_{z_q})$ the uncertainty in κ_{z_q} from standard χ^2 techniques. Instead, we performed a bootstrap analysis of 100 realizations of each stacked spectrum and estimated $\sigma(\kappa_{z_q})$ from the resulting distribution of κ_{z_q} values.

Figure 3 presents the evaluations of κ_{z_q} in terms of the mean free path ($\lambda_{mfp}^{912} = 1/\kappa_{z_q}$). The mean free path exhibits a peak (minimum in opacity) of nearly $50 h_{72}^{-1}$ Mpc at $z = 3.6$ and declines with increasing redshift. We have parameterized the redshift evolution in λ_{mfp}^{912} with a simple linear regression: $\lambda_{mfp}^{912} = \lambda_0 - b_\lambda(z_q - 3.6)$. Restricting the analysis to $z \geqslant 3.59$, a χ^2 minimization of this model to the binned evaluations of λ_{mfp}^{912} gives $\lambda_0 = (48.4 \pm 2.1) h_{72}^{-1}$ Mpc and $b_\lambda = (38.0 \pm 5.3) h_{72}^{-1}$ Mpc.

Formally, our analysis also indicated a rise in the opacity at $z < 3.6$, i.e., the lower limits in Figure 3. This runs contrary to all expectation and current understanding of the IGM. Initially, we suspected that this measurement indicated a systematic error in the SDSS spectra at the bluest wavelengths (e.g., Bernardi et al. 2003). To test this hypothesis, we examined the $(u - g)$ colors of the quasars in the first ($z \approx 3.5$) and third ($z \approx 3.6$) quasar bins. Figure 4 histograms the two $(u - g)$ distributions. The colors
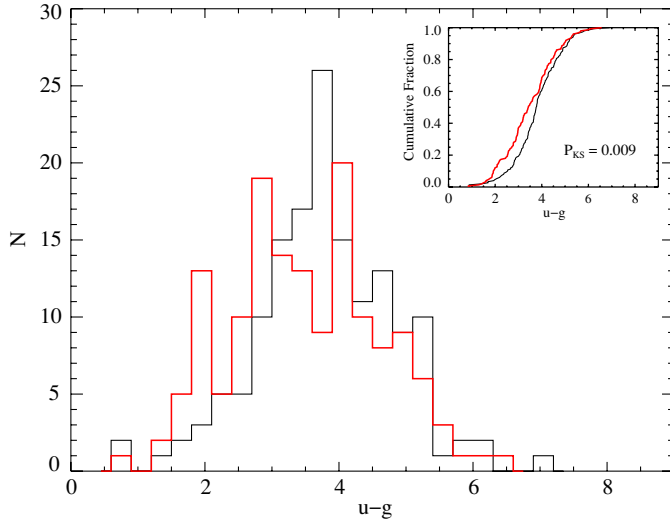
Figure 4. Histograms of the $(u - g)$ colors of quasars in the redshift intervals $z = [3.50, 3.55]$ (black) and $z = [3.59, 3.63]$ (red). The lower redshift cohort shows systematically redder $(u - g)$ colors and a two-sided Kolmogorov–Smirnov test rules out the null hypothesis at greater than 99% c.l. This offset in $(u - g)$ color contradicts standard expectation of a universe where H I absorption monotonically increases with redshift. It occurs because of a bias in the SDSS quasar sample related to the survey's targeting criteria (Richards et al. 2002). Note that objects with $(u - g) \gtrsim 4.5$ are generally not detected in the u band.

of the $z = 3.5$ quasars are systematically *redder* than those at $z = 3.6$; this is the exact opposite of what one predicts if the IGM were monotonically increasing in opacity with redshift. We conclude that the quasars at $z = 3.5$ drawn from the SDSS are redder than the cohort at $z = 3.6$ because of an elevated incidence of the Lyman limit opacity, confirmed by analysis on the incidence of LLS (Prochaska et al. 2009, in preparation).

We then explored whether this elevated opacity is related to observational bias in the SDSS quasar sample. We simulated the SDSS experiment by constructing mock quasar spectra at $z \approx 3.5$ and $z \approx 3.6$ with intrinsic SEDs having mean $(u - g)$ color of 0.57 mag and standard deviation of 0.19 mag. These spectra were blanketed with IGM absorption assuming a monotonically increasing opacity with redshift. After restricting the quasar sample according to the SDSS color-selection criteria (Richards et al. 2002), we found the $z \approx 3.5$ cohort has systematically redder colors than the higher redshift sample. Similarly, we find a correspondingly higher opacity inferred from the stacked spectrum. Because of the targeting criteria, the cohort of $z \approx 3$–3.6 quasars in the SDSS spectroscopic database are systematically biased against having $(u - g) < 1.5$ which biases against sight lines without strong Lyman limit absorption. Our analysis indicates that the bias extends to $z_q = 3.6$, beyond which very few quasars are predicted to ever have such blue color (see Worseck & Prochaska 2009, in preparation, for further details). The results for $z > 3.6$ are presented in Table 1.

4. DISCUSSION

Our new technique provides the most precise measurements on the IGM opacity to H I ionizing radiation at any redshift. Previous estimates of κ_{LL} were limited by large uncertainties in the $N_{H I}$ frequency distribution, especially the incidence of systems with $N_{H I} \approx 10^{17}$ cm^{-2} (Madau et al. 1999; Schirber & Bullock 2003; Miralda-Escudé 2003; Meiksin & White 2004; Faucher-Giguère et al. 2008a). Our results indicate that systems dominating κ_{LL} at $z > 3.6$ must evolve more slowly than the $N_{H I} \leqslant 10^{14}$ cm^{-2} Lyα forest. In turn, we infer a flattening in the

Table 1
Summary Table

z	$\langle z_q \rangle$	$\lambda_{\text{analysis}}$ (Å)	κ_{z_q} (Mpc^{-1})	$\sigma(\kappa_{z_q})$ (Mpc^{-1})
[3.59,3.63]	3.61	[846, 905]	0.0209	0.0023
[3.63,3.67]	3.65	[839, 905]	0.0218	0.0021
[3.67,3.71]	3.68	[834, 905]	0.0224	0.0026
[3.71,3.76]	3.73	[835, 905]	0.0212	0.0020
[3.76,3.81]	3.78	[836, 905]	0.0257	0.0020
[3.81,3.86]	3.83	[837, 905]	0.0254	0.0024
[3.86,3.92]	3.89	[837, 905]	0.0243	0.0024
[3.92,4.02]	3.96	[839, 905]	0.0292	0.0026
[4.02,4.13]	4.08	[840, 905]	0.0341	0.0033
[4.13,4.34]	4.23	[842, 905]	0.0403	0.0036

Notes. Because of the color-criteria bias discussed in this Letter, we caution that the values for the first few bins may systematically underestimate κ_{z_q} by 10%–30%. All wavelengths are in the quasar rest frame and all distances are proper. The assumed cosmology has $\Omega_m = 0.7$, $\Omega_\Lambda = 0.7$, and $H_0 = 72$ km s^{-1} Mpc^{-1}.

H I frequency distribution between $N_{H I} \approx 10^{15}$ and 10^{17} cm^{-2} (Petitjean et al. 1993; O'Meara et al. 2007). We provide a full discussion on the implications for $f(N_{H I}, z)$ in Prochaska et al. (2009, in preparation).

Our analysis also gives the first direct description of the evolution in $\lambda_{\text{mfp}}^{912}$, albeit over a small redshift interval. Our results are well parameterized by a linear decrease in $\lambda_{\text{mfp}}^{912}$ but can also be described by a $(1 + z)^{-\gamma}$ power law with $\gamma = 3.5$–5.5. These values are consistent with the observed evolution in the incidence of LLSs (Prochaska et al. 2009, in preparation).

Our results refine recent inferences (e.g., Faucher-Giguère et al. 2008a; Dall'Aglio et al. 2009) that galaxies contribute significantly to the EUVB at $z > 3$. We assume the photoionization rate at $z = 4$ inferred from the effective Lyα opacity of the IGM ($\log \Gamma_{\text{IGM}} = -12.3$; Faucher-Giguère et al. 2008b). Comparing this value against the photoionization rate inferred from quasars Γ_q using an emissivity $\epsilon_{912}^q = 2 \times 10^{24}$ erg s^{-1} Hz^{-1} Mpc^{-3} (comoving; Hopkins et al. 2007; Bongiorno et al. 2007) and adopting our estimate of $\lambda_{\text{mfp}}^{912}$ at $z = 4$, we find $\Gamma_q = 0.5\,\Gamma_{\text{IGM}}$. This suggests a modest but non-negligible contribution from galaxies to the EUVB at this redshift. The systematic uncertainties in ϵ_{912}^q and Γ_{IGM} are sufficiently large that one could recover $\Gamma_{\text{IGM}} = \Gamma_q$, but Γ_q would overpredict Γ_{IGM} at $z = 3.5$ given the observed rise in $\lambda_{\text{mfp}}^{912}$ and ϵ_{912}^q with decreasing redshift.

The results also revise at least some previous estimates of the EUVB. Haardt & Madau (1996) and their subsequent analyses (CUBA), for example, have adopted an approximately three times shorter mean free path at $z \sim 4$ than our analysis reveals. This implies that (1) the normalization of their EUVB spectrum is several times too low and (2) the EUVB spectrum is softer at energies of ≈ 1 Ryd. The latter point may help to reconcile apparent contradictions in the metal-line analysis of the IGM (Aguirre et al. 2008) without resorting to a large input from galaxies. Other analyses, however, have used estimates for $\lambda_{\text{mfp}}^{912}$ that are in much better agreement with our results (e.g., Faucher-Giguere et al. 2009). An understanding of the full implications of our new constraints on $\lambda_{\text{mfp}}^{912}$ awaits new calculations of the EUVB.

We identified a previously unreported systematic bias in the SDSS quasar spectroscopic sample against sight lines at $z < 3.6$ that are "clear" of optically thick absorbers. This bias has implications for a range of IGM analysis at $z \sim 3$

including (1) the paucity of sight lines for studying He II reionization (Worseck & Prochaska 2009, in preparation), (2) an overestimate of the incidence of LLSs (Prochaska et al. 2009, in preparation) and damped Lyα systems (Prochaska & Wolfe 2009), and (3) analysis of the Lyα forest. We caution that all existing studies of the IGM at $z \sim 3$ using the SDSS database should be reviewed in light of this systematic bias.

In future work, we analyze these same spectra to constrain the SEDs of $z > 3.5$ quasars, infer the $N_{\mathrm{H\,I}}$ frequency distribution and Doppler parameters of gas with $N_{\mathrm{H\,I}} \approx 10^{16}\,\mathrm{cm}^{-2}$, and isolate the opacity of the IGM far from the quasar's proximity region. The technique introduced here is easily extended to higher and lower redshifts by obtaining modest S/N, low-resolution spectroscopy of several hundred quasars. Future ground- and space-based programs will precisely estimate κ_{LL} from $z_q \approx 0.5$ to 5.

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