

The first ultraviolet quasar-stacked spectrum at $z \simeq 2.4$ from WFC3

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ABSTRACT

The ionizing continuum from active galactic nuclei is fundamental for interpreting their broad emission lines and understanding their impact on the surrounding gas. Furthermore, it provides hints on how matter accretes on to supermassive black holes. Using *Hubble Space Telescope*’s Wide Field Camera 3, we have constructed the first stacked ultraviolet (rest-frame wavelengths 600–2500 Å) spectrum of 53 luminous quasars at $z \simeq 2.4$, with a state-of-the-art correction for the intervening Lyman forest and Lyman continuum absorption. The continuum slope ($f_\nu \propto \nu^{\alpha_\nu}$) of the full sample shows a break at ~ 912 Å with spectral index $\alpha_\nu = -0.61 \pm 0.01$ at $\lambda > 912$ Å and a softening at shorter wavelengths ($\alpha_\nu = -1.70 \pm 0.61$ at $\lambda \leq 912$ Å). Our analysis proves that a proper intergalactic medium absorption correction is required to establish the intrinsic continuum emission of quasars. We interpret our average ultraviolet spectrum in the context of photoionization, accretion disc models, and quasar contribution to the ultraviolet background. We find that observed broad line ratios are consistent with those predicted assuming an ionizing slope of $\alpha_{\text{ion}} = -2.0$, similar to the observed ionizing spectrum in the same wavelength range. The continuum break and softening are consistent with accretion disc plus X-ray corona models when black hole spin is taken into account. Our spectral energy distribution yields a 30 per cent increase to previous estimates of the specific quasar emissivity, such that quasars may contribute significantly to the total specific Lyman limit emissivity estimated from the Ly α forest at $z < 3.2$.

Key words: accretion, accretion discs – galaxies: active – quasars: general.

1 INTRODUCTION

Considerable effort has been devoted to characterizing the shape of the quasar (QSO) ionizing continuum in the ultraviolet (UV) over the past years. The spectral energy distribution (SED) of active galactic nuclei (AGN) shows a prominent bump, the so-called Big Blue Bump (BBB), which appears to peak in the UV and decline at higher energies (Sanders et al. 1989; Elvis et al. 1994). However, the *intrinsic* position and possible luminosity dependence of the BBB has not been properly estimated due to the lack of UV observations corrected for the intergalactic medium (IGM) absorption by neutral hydrogen along the line of sight (Richards et al. 2006a; Trammell et al. 2007; Shang et al. 2011; Elvis et al. 2012).

One of the main predictions of accretion disc models is that the thermal disc should peak at bluer wavelengths for smaller black hole (BH) masses (i.e. $T_{\text{max}} \propto (L/L_{\text{Edd}})^{1/4} M_{\text{BH}}^{-1/4}$; Shakura & Sunyaev

1973). However, both the presence and the position of the break are highly dependent on the IGM correction considered.

From a theoretical perspective, the AGN-ionizing continuum is crucial for interpreting broad and narrow emission lines observed in AGN spectra, and their relative ratios (see Stern et al. 2014 and Baskin et al. 2014 for recent models of these emission line regions). There is a general consensus that AGN emission lines are produced by a photoionizing continuum (extending from optical–UV to X-ray) which emerges from an accretion disc around the BH and by a hot ($T \sim 10^8 - 10^9$ K) plasma of relativistic electrons that Compton up-scatter the photons coming from the disc (Haardt & Maraschi 1991, 1993).

Quasars are relevant (maybe even the dominant) sources of ionizing photons that determine the ionization state and the temperature of the $z \sim 3$ IGM (e.g. Haardt & Madau 1996, 2012; Meiksin & White 2003; Faucher-Giguère et al. 2009). While not numerous enough at $z \gtrsim 6$ to have contributed significantly to H I reionization (e.g. Meiksin 2005; Shankar & Mathur 2007; Jiang et al. 2008; Willott et al. 2010; Fontanot, Cristiani & Vanzella 2012; Fontanot

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et al. 2014), they are likely the only sources responsible for the reionization of He II at $z \sim 3$ (Miralda-Escudé, Haehnelt & Rees 2000; Faucher-Giguère et al. 2008; Furlanetto 2009; McQuinn et al. 2009; Haardt & Madau 2012; Compostella, Cantalupo & Porciani 2013). All these studies rely on parameterizations of the quasar SED in the BBB region motivated by existing observations, the uncertainties of which are rarely discussed.

From an observational point of view, composite spectra of AGN were constructed by taking advantage of several surveys (LBQS, Francis et al. 1991; FIRST, Brotherton et al. 2001; SDSS, Vanden Berk et al. 2001). In all these studies, the rest-frame optical composites indicate that the continuum can be described by a power law of the form $f_\nu \propto \nu^{\alpha_\nu}$, with $-0.5 \lesssim \alpha_\nu \lesssim -0.3$ in the wavelength range 1200–4000 Å. The first composite that suggests a softening in the far-UV (FUV) (blueward of Ly α) was reported by Zheng et al. (1997, hereafter Z97) who analysed 101 quasars from the *Hubble Space Telescope* (HST) in the redshift range $0.33 < z < 3.6$, covering the wavelengths between 350 and 3000 Å. This softening was interpreted as Comptonization of the thermal disc emission in a hot corona above the disc (Czerny & Elvis 1987), due to the similarity between the slope of -1.8 found by Z97 at $\lambda < 1216$ Å and the one measured by Laor et al. (1997) for a sample of radio-quiet X-ray-selected quasars ($\alpha_\nu \sim -1.7$), for which simple accretion disc+X-ray corona models were utilized. The break would thus indicate the peak of the BBB. The work by Z97 has been extended by Telfer et al. (2002, hereafter T02) with more than 80 quasars from HST over a similar redshift range. This analysis confirmed the findings of Z97 that the UV spectral continuum can be parametrized by a broken power law with the break in the vicinity of the Ly α , with a continuum slope of $\alpha_\nu = -1.57$ for the radio-quiet T02 sample. These results are at variance with those presented by Scott et al. (2004, hereafter S04), who considered more than 100 AGN at $z < 0.1$ observed with the *Far Ultraviolet Spectroscopic Explorer* (FUSE), covering the rest-frame wavelength range 630–1100 Å. The spectral slope of the FUSE composite spectrum is $\alpha_\nu = -0.56$, significantly harder than previous estimates in the far-infrared from HST studies. Shull, Stevans & Danforth (2012, hereafter S12) have measured the AGN ionizing continua in 22 AGN at $0.026 < z < 1.44$ using the Cosmic Origins Spectrograph (COS) on HST, covering the rest-frame wavelength range 500–2000 Å. The COS composite shows a break at $\lambda \sim 1000$ Å, in line with the previous estimates by Z97 and T02, but with slightly harder spectral index ($\alpha_\nu = -0.68$ at $\lambda = 1200$ –2000 Å and $\alpha_\nu = -1.41$ at $\lambda = 500$ –1000 Å). Recently, Stevans et al. (2014, hereafter S14) considerably improved the S12 analysis by adding 137 AGN to the original sample for a total of 159 sources, selected from the COS archive at redshifts $0.001 < z < 1.476$. This new COS composite is fully consistent with the one in S12 and shows similar spectral indexes. We note that previous estimates of the break were performed with very few spectra (~ 10 in T02, 3–12 in S12, ~ 20 in S14) contributing at short wavelengths (e.g. < 700 Å).

The differences among various surveys may arise from several factors, such as the small number of observations covering $\lambda < 1216$ Å (e.g. less than 20 AGN at $z > 2$ in the T02 sample), and the crucial placement of the ‘continuum windows’ in order to construct the quasar-ionizing continua. Although the latter may not bias significantly the results, a possible uncertainty arises from the selection criteria involved in defining the various samples. The traditional strategy is to extract any source available from the HST/FUSE archives and apply several cuts such as redshift, signal to noise, and wavelength coverage. However, the brightest UV sources were usually targeted for UV spectroscopy, primarily for

absorption line studies. Furthermore, any archival research has the tendency to include more peculiar objects that were selected for special investigations, and thereby the same objects were re-observed with every generation of UV spectrographs. These issues likely bias any UV archival sample to be very bright in the UV with a highly heterogeneous selection function.

An additional source of variance comes from the correction for the IGM absorption employed by different authors. Although the technique adopted by T02 and S04 was similar, i.e. a statistical correction for the unidentified absorbers by considering an empirical parametrization (see Petitjean et al. 1993; Davé & Tripp 2001), the overall correction applied was very different as pointed out by S04. In fact, S04 noted that the correction considered by T02 was less than 1 per cent over the whole rest-frame wavelength range, which underestimated the number of Ly α absorbers by a factor of ~ 50 at $z = 0.1$. Another simplification usually adopted in previous studies is the use of a single correction for quasar samples spanning a large range of redshift, which, because of strong evolution in the IGM absorption, likely results in significant errors.

In this paper, we present the first *average* quasar spectrum in the rest-frame wavelength range 600–2500 Å corrected for intervening H I Lyman forest and continuum absorption with state-of-the-art IGM transmission functions calibrated to the most recent observations (Prochaska et al. 2014). The sample consists of 53 $z_{\text{em}} \simeq 2.4$ quasars from the HST survey for Lyman limit absorption systems (LLSs) using the Wide Field Camera 3 (WFC3) presented in O’Meara et al. (2011, hereafter O11). The structure of this paper is as follows. In Section 2, we discuss the sample and the selection criteria. In Section 3, we describe the technique to construct the stacked spectrum, whilst the IGM transmission curves adopted to correct the observed average spectrum are presented in Section 4, where we also describe our IGM-corrected stack with uncertainties. The results and implications of our analysis are discussed in Section 5, while the conclusions are presented in Section 6.

We adopt a concordance flat Λ -cosmology with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.3$, and $\Omega_\Lambda = 0.7$ (Komatsu et al. 2009). Unless noted otherwise, we will distinguish between the following wavelength ranges in the UV: (i) the near-UV (NUV; 2000–3000 Å), (ii) the far-UV (FUV; 912–2000 Å), and (iii) the extreme UV (EUV; $\lambda < 912$ Å).

2 THE DATA SET

The data sample employed in the present analysis comes from a survey performed with HST using the WFC3 instrument. Quasar sample, HST observations, and reduction procedures are described in detail in O11. In this section, we provide a brief summary of this data set.

The O11 survey consists of 53 quasars selected from SDSS Data Release 5 (Schneider et al. 2007) with $g < 18.5$ mag, $2.3 < z_{\text{em}} < 2.6$, observed with the WFC3/UVIS-G280 grism in Cycle 17. These data were taken specifically for the scientific goal of surveying the abundance of strong H I Lyman limit absorption features at $1.2 < z < 2.5$. Flux- and wavelength-calibrated 1D spectra for each quasar in this sample are extracted using customized software (see section 3.1 in O11). WFC3/UVIS-G280 spectra span roughly $\lambda = 2000$ –6000 Å and they have relatively high signal-to-noise ratio ($S/N \sim 20$) per pixel down to $\lambda \sim 2000$ Å¹ (FWHM ~ 60 Å at

¹ The data generally have S/N exceeding 10 pixel⁻¹ at all wavelengths $\lambda > 2000$ Å.

$\lambda = 2500 \text{ \AA}$). Uncertainties in the wavelength calibration are of the order of 2 pixels, in the form of a rigid shift in the pixel space (see table 1 in O11).

As already pointed out by O’Meara et al. (2013, hereafter O13), since this sample is based on SDSS optical colour selection (Richards et al. 2002), it might be biased towards bluer colours than a complete sample. However, Worseck & Prochaska (2011) have demonstrated that at the redshifts of this sample this bias is relatively small (see their fig. 16).

Note however that quasar samples constructed from UV spectroscopy archives are based both on optical colour selection and additional UV brightness criteria. Typically, the brightest objects observed in previous *HST* cycles were re-observed when a new UV instrument came online. Objects observed by the *HST* COS were observed in the post *GALEX* era, and are thus explicitly biased to have bright NUV and FUV magnitudes. The selection function of quasars in the UV archives is thus unknown and extremely difficult to quantify, but the expectation is that such samples are biased very blue. In contrast, the selection criteria for our quasar sample is much cleaner, as it is simply an optical apparent magnitude-limited sample $g < 18.5$ of quasars at $z_{\text{em}} \simeq 2.4$.

2.1 X-rays

The X-ray data have been gathered from the *ROSAT*, *XMM-Newton*, and *Chandra* archives. We found two objects (J1253+0516 and J1335+4542) detected in the *ROSAT* All-Sky Survey Faint Source Catalog (RASS-FSC), six quasars have at least one *ROSAT*/PSPC (White, Giommi & Angelini 1994a, 1994b) observation, while two sources have *ROSAT*/HRI images (J1454+0325 and J1119+1302 with also a PSPC observations).

The X-ray fluxes at 0.5–2 keV were calculated from the count rates in the *ROSAT* band (0.1–2.4 keV) and in the PSPC band (0.24–2 keV) by utilizing a power-law spectrum with a photon index $\Gamma = 2.0$ modified by Galactic absorption (Kalberla et al. 2005) only. The X-ray properties of the objects in our sample detected by *ROSAT* are listed in Table A1.

Sources without a *ROSAT* detection might be either intrinsically faint, X-ray obscured, and/or highly variable. We expect the undetected AGN to have X-ray luminosities of the order of $7 \times 10^{45} \text{ erg s}^{-1}$ or lower in the rest-frame 0.5–2 keV band.²

Among these 10 quasars, two have additional spectral information from *XMM-Newton* (J0755+2204 and J1119+1302) and one from *Chandra* (J1220+4608). The soft ($S = 0.5$ –2 keV) and hard ($H = 2$ –10 keV) X-ray luminosities for J0755+2204 and J1119+1302 are tabulated in the SDSS (DR5)/*XMM-Newton* quasar survey catalogue, and the values are $L_S = 4.60 \times 10^{44}$ and $L_H = 1.14 \times 10^{45} \text{ erg s}^{-1}$ for J0755+2204, while J1119+1302 has $L_S = 1.96 \times 10^{45}$ and $L_H = 5.00 \times 10^{45} \text{ erg s}^{-1}$.

J1220+4608 was targeted by *Chandra* with ACIS-S in 2008 March with a nominal exposure of 50 ks. The spectrum has very low number counts, with a $\leq 3\sigma$ detection in the soft band. In this case, assuming a photon index of 1.8 and Galactic N_H , we found that the extrapolated flux in the 2–10 keV band is $1.36 \times 10^{-14} \text{ erg s}^{-1} \text{ cm}^{-2}$. This flux corresponds to a luminosity of $6.4 \times 10^{44} \text{ erg s}^{-1}$.

² The *ROSAT* RASS flux limit is $5 \times 10^{-13} \text{ erg s}^{-1} \text{ cm}^{-2}$ for a mean effective exposure time of 400 s and $\Gamma = 2.0$, but the range of *ROSAT* exposure times is large, and the sensitivity limit is different from field to field (Voges et al. 1999).

Summarizing, the fraction of detected sources is 21 per cent (i.e. 11/53, 10 objects detected by *ROSAT/XMM-Newton* and one additional source with a *Chandra* observation) with a mean soft X-ray luminosity of about $5.6 \times 10^{45} \text{ erg s}^{-1}$.

2.2 Radio

To estimate the fraction of radio emitters in our sample, we matched it with the DR7 quasar property catalogue by Shen et al. (2011). Radio properties are collected from FIRST (Becker, White & Helfand 1995) and from NRAO/VLA Sky Survey (Condon et al. 1998) with a matching radius of 30 arcsec (our matches are all within 1.3 arcsec). We have found 11 quasars in the FIRST survey, while one object (J2338+1504) has a detection from NRAO, for a total of 12 detections. The radio loudness is estimated following Jiang et al. (2007), $R = f_{6\text{cm}}/f_{2500}$ where $f_{6\text{cm}}$ and f_{2500} are the flux density at rest-frame 6 cm and 2500 Å, respectively. According to Jiang et al., 11 quasars are core-dominant and one is lobe-dominant (J1119+1302, radio morphology classification following Jiang et al. 2007). Five objects are not in the FIRST footprint (flag = −1 in the DR7 catalogue, J1259+6720, J1325+6634, J1400+6430, J2111+0024, J2136+1029) and they do not have any detection from NRAO. A summary of the radio properties is given in Table A2. Following Kellermann et al. (1989), quasars with $R > 10$ are defined as ‘radio-loud’. In the WFC3 sample, the radio-loud fraction (RLF) is of the order of 19 per cent. This is consistent with the RLF of quasar at similar optical magnitudes.³

In the following, we will present the WFC3 average spectrum with and without the 10 objects with $R > 10$ (a stand-alone radio-loud quasar stack is not constructed given the poor statistics).

2.3 GALEX

To extend our wavelength coverage to shorter wavelengths, we first considered the *GALEX* photometry in the DR7 quasar property catalogue. We found that the detection rate for the near-UV (NUV at $\lambda_{\text{eff}} = 2316 \text{ \AA}$) and far-UV (FUV at $\lambda_{\text{eff}} = 1539 \text{ \AA}$) bands was 64 per cent (34/53) and 14 per cent (14/53), respectively. Given the low detection rate, we have cross-matched our sample to the *GALEX*-forced photometry catalogue (Schiminovich, private communication), which allows us to obtain a detection for almost all sources in the WFC3 sample. The *GALEX*-forced photometry was not available for only four objects (J0751+4245, J1354+5421, J1540+4138, J2111+0024), for the rest of the sample we have both NUV and FUV bands. The observed *GALEX*-forced photometry for the WFC3 sample is listed in Table A3. The average *GALEX* bands are then corrected for IGM absorption and we will outline how these data have been corrected for IGM absorption in Section 5.4.

3 AVERAGE SPECTRUM CONSTRUCTION

We follow a similar procedure as O13 for the construction of the WFC3 average spectrum. Specifically:

- (i) We correct the quasar flux density⁴ (f_{λ}) for Galactic reddening by adopting the $E(B - V)$ estimates from Schlegel, Finkbeiner &

³ Jiang et al. (2007) have found that the RLF is about 15 per cent for an absolute magnitude at the rest-frame 2500 Å (M_{2500}) of ~ -28 , which is consistent with the average M_{2500} of the WFC3 sample ($M_{2500} \simeq -29$).

⁴ In the following, we will use the word ‘flux’ to mean the flux density (i.e. flux per unit wavelength).

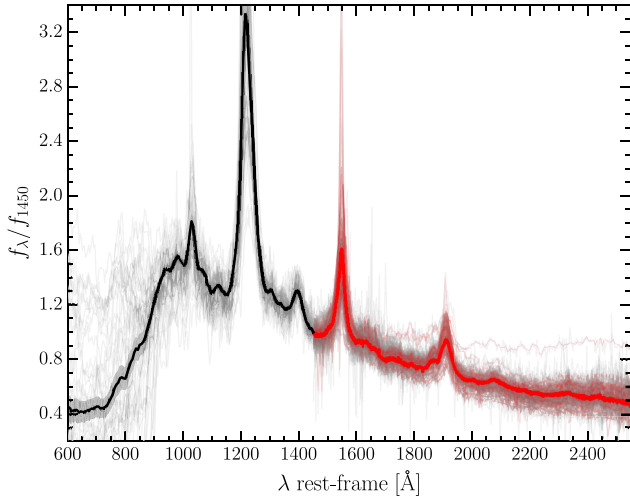


Figure 1. Mean observed QSO spectrum as estimated from a stack of the WFC3 spectra, each normalized to unit flux at 1450 Å (see Section 3 for details). The black line represents the WFC3 QSO average spectrum with uncertainties from bootstrap (black shaded area). The red line represents the SDSS QSO average spectrum (not smoothed to the WFC3 resolution) of our WFC3 sample with uncertainties. Grey lines represent the single QSO spectra, while thin red lines represent SDSS spectra smoothed to WFC3 resolution.

Davis (1998, hereafter SFD, median reddening value is $E(B - V) = 0.02$ mag) and the Galactic extinction curve from Fitzpatrick (1999) with $R_V = 3.0$. The same reddening law has been considered to correct the *GALEX* fluxes.

(ii) We generate a rest-frame wavelength array with fixed dispersion $\Delta\lambda$. The dispersion value was set to be large enough to include at least one entire pixel from the WFC3/UVIS-G280 spectra at rest wavelengths $\lambda < 1215$ Å (i.e. $\Delta\lambda \simeq 6.2$ Å).

(iii) Each quasar spectrum was shifted to the rest frame and linearly interpolated over the rest-frame wavelength array with fixed dispersion $\Delta\lambda$.⁵

(iv) We normalized single spectra by their flux at rest $\lambda = 1450$ Å.

(v) All the flux values were then averaged to produce the stacked spectrum normalized to unity at $\lambda = 1450$ Å.

Recently, Peek & Schiminovich (2013) have found that UV colours at high latitudes are best fitted with a Fitzpatrick (1999) extinction curve with $R_V \simeq 2.2$, while Schlafly & Finkbeiner (2011) found that SFD overestimates reddening by 14 per cent (reddening values should be recalibrated as $E(B - V) = 0.86 \times E(B - V)_{\text{SFD}}$). The average WFC3 spectrum re-estimated following the results outlined above is fully consistent with the one constructed with $R_V = 3.0$ and $E(B - V)$ from SFD.

Uncertainties on the observed stack are estimated through the bootstrap resampling technique. We created 5000 random samplings of the 53 spectra with replacement, and we applied the same procedure as described above. The resulting stack is shown as the solid black line in Fig. 1 for the full sample, while the resulting uncertainties on the stacked spectrum are plotted with a shaded area. The same technique is applied to the SDSS quasar spectra and the resulting stacked spectrum is presented in Fig. 1 with the red solid line. For plotting purpose, we show the SDSS-stacked spectrum down to

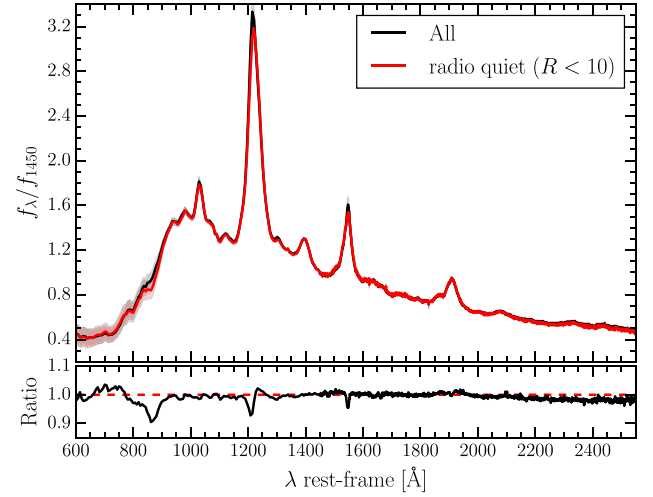


Figure 2. Upper panel: mean observed QSO spectrum for the WFC3 sample (black line) and for the subsample with $R < 10$ (red line) each normalized to unit flux at 1450 Å. Lower panel: ratio of the $R < 10$ to the full mean observed QSO spectrum.

$\lambda = 1450$ Å given that the WFC3 stack already covers wavelengths below this region.

To bring the WFC3 and the SDSS-stacked spectra over a common luminosity scale in the νL_ν plane, we have multiplied the final stacks for the mean flux (estimated from the WFC3 and SDSS) of the sample at 1450 Å. For the sake of matching, the flux scale we have convolved the SDSS spectra to WFC3 resolution. The mismatch between the WFC3 and SDSS spectra at 1450 Å is 11 per cent and it is due to variability and flux calibration errors in the WFC3 data (see section 3.2 in O11 for details). We have thus re-scaled the WFC3-stacked spectrum to match the (convolved) SDSS at 1450 Å. We have also constructed the quasar average spectrum by excluding the 11 objects with $R > 10$ for completeness. The resulting spectrum is shown in Fig. 2 where we have over plotted the WFC3 average spectrum as a comparison. The radio quiet quasar stack is fully consistent with the WFC3 stack within the uncertainties.

4 IGM TRANSMISSION CORRECTION

Blueward of Lyman α emission in the quasar rest frame, absorption from intergalactic H I attenuates the quasar flux, both in the Lyman series (creating the so-called H I forest), and in the Lyman continuum at rest $\lambda < 912$ Å (e.g. Møller & Jakobsen 1990). The large abundance of neutral gas at $z_{\text{em}} \lesssim 2.4$ is very apparent in our average quasar spectrum shown in Fig. 1. With only mild assumptions on the average quasar SED, the exponential flux decline at $\lambda < 912$ Å yields the mean free path of Lyman limit photons in the IGM (O13). Here, we reverse the question and constrain the quasar SED for a range of IGM transmission curves T_λ for $\lambda < 1215.67$ Å.

For a source at emission redshift z_{em} , the effective optical depth to H I Lyman series and Lyman continuum photons at redshift $z < z_{\text{em}}$ is determined from the H I absorber distribution function in redshift and column density $f(N_{\text{H I}}, z) = \partial^2 n / (\partial N_{\text{H I}} \partial z)$. The resulting average IGM transmission T_λ critically depends on the parametrization of $f(N_{\text{H I}}, z)$ (Madau 1995; Meiksin 2006; Inoue et al. 2014) and is statistical in nature due to the stochasticity of Lyman limit systems (Bershady, Charlton & Geoffroy 1999; Inoue & Iwata 2008; Worseck & Prochaska 2011). Quasar composites based on low-resolution spectra need to be corrected for Lyman series and Lyman

⁵ Wavelengths are divided by $(1 + z)$ to shift the spectra into the source rest frame, while fluxes in f_λ are multiplied by $(1 + z)$.

continuum absorption of low-column density absorbers that cannot be identified and corrected by eye. The statistical IGM correction strongly depends on the assumed absorber distribution parameters and their dependence on redshift, which may have resulted in large systematic errors in existing quasar composite spectra if the incorrect parameters were used or if redshift evolution was not properly taken into account. Moreover, the ability to identify weak partial Lyman limit systems ($\tau_{912} \lesssim 0.3$) depends on redshift, the employed spectra (S/N, spectral resolution, wavelength coverage, flux calibration), and the intrinsic quasar continuum, such that heterogeneous samples are prone to ambiguities in the continuum definition and the treatment of Lyman limit systems. In particular, we stress that applying an average correction for partial Lyman limit systems to de-attenuate individual $z_{\text{em}} \lesssim 3$ sightlines is incorrect because of the stochasticity of Lyman limit absorption (Worseck & Prochaska 2011).

Our large sample at similar $z_{\text{em}} \sim 2.4$ ensures a sufficient sampling of partial Lyman limit systems, justifying a redshift-specific IGM correction function T_λ to be applied to our stacked spectrum (Fig. 1). However, the low spectral resolution prevents an unambiguous identification of weak partial Lyman limit systems in individual spectra without knowledge of the underlying quasar continuum. We therefore have to correct for the average Lyman series and continuum absorption of the whole H I absorber population statistically. While this is the simplest and most liberal approach, the significant flux drop in the Lyman continuum implies a large correction to our stacked spectrum, with additional uncertainties related to the parametrization of the IGM.

In this analysis, we consider the recent constraints on the $z \sim 2.4$ IGM presented by Prochaska et al. (2014, and references therein) within their range of uncertainty. Prochaska et al. (2014) introduced a cubic Hermite spline model to describe $f(N_{\text{H I}}, z)$. They performed a Monte Carlo Markov Chain (MCMC) analysis of existing constraints on $f(N_{\text{H I}}, z)$ to derive the posterior probability distribution functions of seven spline points spaced at irregular logarithmic intervals in the range $N_{\text{H I}} = 10^{12} - 10^{22} \text{ cm}^{-2}$. Using the output from their MCMC chains,⁶ we generated 10 000 realizations of $f(N_{\text{H I}}, z)$ at $z = 2.4$ and calculated T_λ in the observed wavelength range with a semi-analytic technique. This modelling assumes that the H I forest is composed of discrete ‘lines’ with Doppler parameter $b = 24 \text{ km s}^{-1}$ and that the normalization of $f(N_{\text{H I}}, z)$ evolves as $(1 + z)^{0.5}$ which implies increasing transmission below $\lambda = 912 \text{ \AA}$ for the lower redshift Lyman series. Opacity due to metal line transitions was ignored since they contribute negligibly to the total absorption in the Lyman continuum. In Fig. 3, we plot our suite of 10 000 IGM transmission functions with grey lines. We then constructed the mean of all of these different realizations and smoothed this to the WFC3 grism resolution (5 pixels). Finally, we shifted the observed wavelengths to rest frame at $z = 2.4$ and resampled the transmission functions on to the rest-frame wavelength grid of our stacked quasar spectrum. This defines our mean IGM transmission function which is shown as the black curve in Fig. 3. Given the narrow redshift range of our quasar sample, we did not account for redshift evolution in T_λ . To summarize, by treating $f(N_{\text{H I}}, z)$ within reasonable uncertainties our approach yields a range of plausible IGM correction functions, a clear advance over previous analyses that assumed a fixed correction. However, since we account for the total H I absorber population, our IGM correction functions in Fig. 3

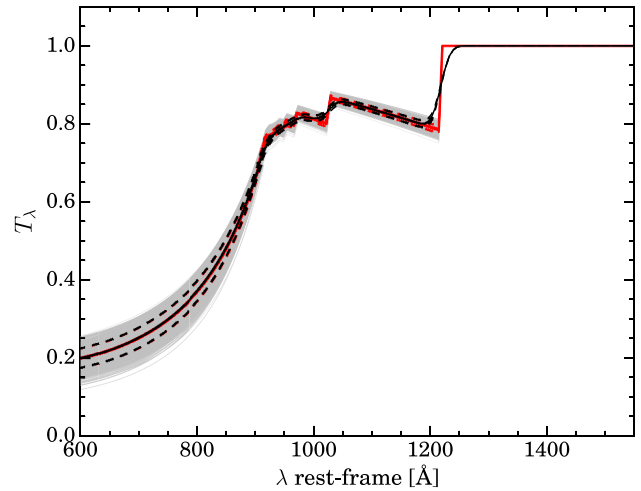


Figure 3. IGM transmission curves as a function of rest-frame wavelength. Grey curves represent 10 000 different realizations of T_λ , while the solid and dashed red curves are the stack and 1σ dispersion, respectively. The black solid curve is the smoothed stack to the WFC3 resolution with the corresponding 1σ dispersion (black dashed lines).

assume that the column density distribution is well sampled at high column densities.

The procedure to correct the observed WFC3 spectra for IGM absorption is outlined as follows.

(i) We generate a set of 10 000 mock quasar stacks, following the same procedure as in Section 3, by drawing randomly from the 53 quasar spectra to assess sample variance allowing for duplications.

(ii) We then randomly draw one IGM transmission function from our suite of 10 000. We smoothed this to the WFC3 grism resolution (5 pixels), and we resampled the transmission function on to the rest-frame wavelength grid of our stacked quasar spectrum. This is repeated for each mock quasar stack.

(iii) We divide the observed spectral flux ($f_{\lambda, \text{obs}}$) by the IGM transmission curve

$$f_{\lambda, \text{corr}} = f_{\lambda, \text{obs}} / T_\lambda. \quad (1)$$

(iv) The 10 000 mock stacks corrected from IGM absorption are then averaged to produce the stacked spectrum (normalized to unity at $\lambda = 1450 \text{ \AA}$).

(v) The uncertainties on the corrected WFC3 stacked spectrum are estimated from the dispersion of these 10 000 mock stacks.

The resulting stacks for the full WFC3 and the radio quiet samples are shown in Fig. 4 with the blue and red lines, respectively, and are tabulated in Table 1. The stacked spectra show a softening at wavelengths $\lambda < 912 \text{ \AA}$ and several emission lines. We will discuss these results further in Section 5. The 1σ uncertainties on the constructed stack are displayed as a light blue shaded area.

One possible concern about the measure of the uncertainties on the stack is that the bootstrap technique outlined above may not have converged far in the blue. In the limit where only a handful of the same transmitting spectra actually contribute to the stack at $\lambda < 912 \text{ \AA}$, the bootstrap may underestimate the true noise in our measurement. In general, the overall uncertainty on the stack depends on a variety of effects such as the intrinsic fluctuations of the underlying quasar continuum (sample variance), our imperfect knowledge of the IGM transmission function (due to uncertainties in the $f(N_{\text{H I}}, z)$; see Fig. 4), spectral noise (i.e. read noise and photon

⁶ <http://www.arcetri.astro.it/~lusso/Site/Research.html>

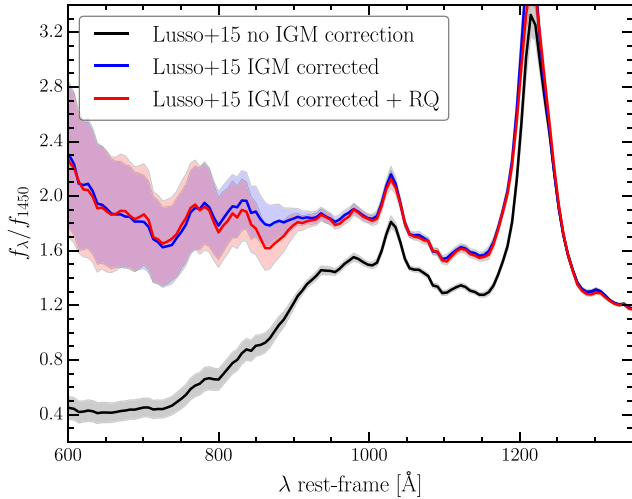


Figure 4. Mean IGM corrected QSO spectrum with uncertainties from bootstrap (shaded area) for the full WFC3 (blue) and for the radio-quiet (red) sample. The black line represents the observed WFC3 QSO stack with uncertainties from bootstrap (grey shaded area).

Table 1. WFC3/UVIS stacked spectrum corrected for IGM absorption.

λ^a	All $f_{\lambda,f}^b$	$\sigma(f_{\lambda,f})^c$	RQ $f_{\lambda,f}$	$\sigma(f_{\lambda,f})$
583.351	2.382	0.577	2.405	0.602
589.540	2.247	0.548	2.302	0.580
595.729	2.326	0.542	2.345	0.576
601.919	2.297	0.533	2.248	0.551
608.108	2.234	0.518	2.205	0.548
614.297	2.092	0.492	2.107	0.522
620.486	2.081	0.472	2.049	0.492
626.675	2.090	0.474	2.044	0.495
632.865	2.037	0.462	1.996	0.477
639.054	1.945	0.444	1.911	0.467
645.243	1.937	0.428	1.916	0.448
651.432	1.908	0.418	1.888	0.442
657.621	1.872	0.402	1.864	0.423
663.811	1.866	0.388	1.871	0.413
670.000	1.867	0.388	1.893	0.413
676.189	1.850	0.379	1.890	0.405
682.378	1.848	0.373	1.857	0.394
688.567	1.814	0.363	1.831	0.384
694.757	1.814	0.354	1.847	0.375
700.946	1.811	0.340	1.866	0.370

Notes. ^aRest-frame wavelength in Angstrom.

^bMean IGM corrected flux per Å normalized to the flux at 1450 Å.

^cFlux uncertainties from our bootstrap analysis (see text).

(This table is available in its entirety in a machine-readable form in the online journal. A portion is shown here for guidance regarding its form and content.)

counting noise), and most importantly, Poisson shot-noise (sample variance) due to the number of LLS and partial LLS intercepted by the quasar sightline.

As we have discussed above, our suite of 10 000 T_λ are created from an integral over the column density distribution, which includes uncertainties in the parameters governing the $f(N_{\text{H I}}, z)$, but it does not include the variance due to the stochasticity of LLS absorption.

To this end, we conducted a Monte Carlo simulation with a different set of empirical IGM transmission curves following the IGM parametrization described in O13. We assumed the $f(N_{\text{H I}}, z)$ defined in table 10 in O13 with a $(1+z)^{1.5}$ redshift evolution in the normalization (see their equation 6). For each quasar in our sample, we drew 100 sightlines from a parent sample of 5000 sightlines populated with $0 < z < 3$ absorbers and computed mock WFC3 IGM transmission spectra in the covered wavelength range, i.e. at $2000 \text{ Å} < \lambda < 1215.67(1+z_{\text{em}}) \text{ Å}$. This allows for the small redshift evolution in the IGM transmission in the narrow redshift range covered by our sample. The result is a set of 100 mock spectra per quasar that accounts for the Poisson statistics of Lyman limit systems. This method does not consider uncertainties in the shape of $f(N_{\text{H I}}, z)$, but samples $f(N_{\text{H I}}, z)$ properly at high column densities (albeit without clustering). The Monte Carlo simulation is then carried out as below.

(i) We start by assuming the true underlying mean QSO spectrum to be the one corrected by the T_λ functions.

(ii) We then create a large number (i.e. 8000) of mock samples of 53 quasars by drawing randomly from the 100 IGM transmission realisations for each quasar in the mock sample.

(iii) We finally calculate the variance of this stack from these many ensembles of 53.

This procedure fully encapsulates the stochastic nature of IGM absorption, and in particular shot noise due to the presence or absence of LLSs. Fig. 5 presents the comparison between the uncertainties from bootstrap and those estimated following the above approach. The amplitude of the uncertainty on the Monte Carlo stack matches the one estimated by applying the bootstrap technique on the data. Our Monte Carlo simulations suggest an ~ 16 per cent fluctuation in samples of 53 due solely to shot-noise in the IGM absorption. Since this fluctuation is comparable to the bootstrap error in our observations, this strongly suggests that the IGM fluctuations dominate the error budget in the stack, and that the other sources of variance (i.e. intrinsic variations in the quasar continuum and spectra noise) are subdominant. Most importantly, this comparison also shows that the bootstrapping is converged far in the blue. In fact, given the narrow

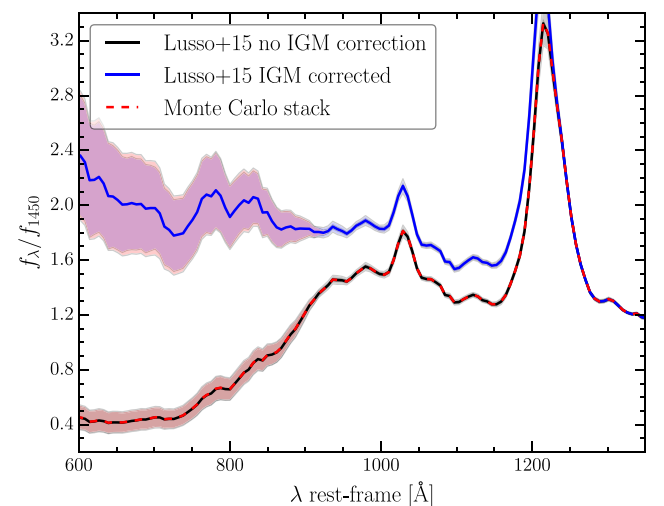


Figure 5. Comparison between the Monte Carlo stack and 1σ uncertainties described in Section 4 (red dashed line) with the one estimated by applying the bootstrap technique on the data. The red shaded area matches the uncertainties due data bootstrap, highlighting the fact that the bootstrapping is converged far in the blue.

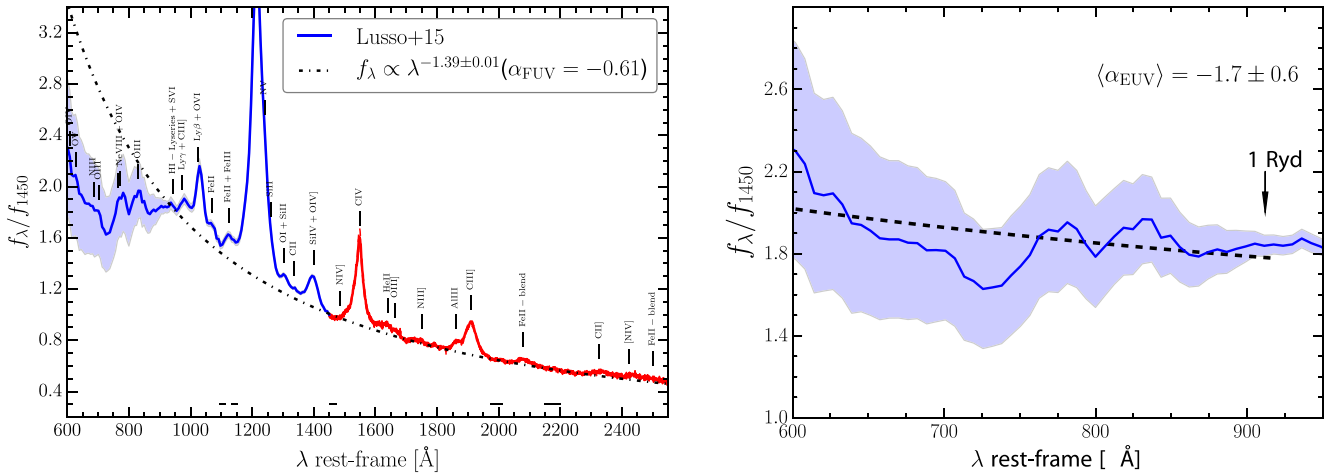


Figure 6. Left-hand panel: average WFC3 and SDSS spectra normalized at 1450 Å with fits to power-law continuum (dot-dashed line). Five continuum windows (from T02) are shown as black horizontal lines below the spectrum. The power-law continuum fit exhibits a break at $\lambda \simeq 920$ Å, with a flatter (softer) spectrum at shorter wavelength. Right-hand panel: zoom-in of the EUV region (600–912 Å). The dashed line is the power-law continuum by employing the average EUV spectral slope.

AGN redshift range, which goes from $z_{\min} = 2.282$ to $z_{\min} = 2.599$ (mean redshift $\langle z_{\text{em}} \rangle \sim 2.44$), almost all spectra contribute appreciably to the total flux at 600 Å. We note that the stack estimated with this second set of IGM transmission curves is in agreement, within the uncertainties, with the one constructed from the 10 000 T_{λ} created from an integral over the column density distribution. We have thus decided to consider the stack constructed from the 10 000 IGM transmission curves for the rest of our analysis.

5 RESULTS

5.1 Spectral fit and emission lines

The AGN-ionizing continuum and its shape in the optical–UV is crucial for several reasons, such as interpreting broad and narrow emission lines observed in AGN spectra, and their relative ratios. It also defines the BBB (i.e. the bulk of the QSO emission), which can be important in interpreting the observed relation between optical and soft X-ray fluxes.

In order to infer the continuum slope of the WFC3 stack, we have to avoid contamination of broad emission lines, broad absorption features, and extended wings of emission lines⁷ especially at $\lambda > 1216$ Å where the emission lines are more prominent.

We have thus fitted the UV stack in each wavelength window free of strong features by means of a single power law. We adopted the same intervals as T02, but given that the uncertainties dramatically increase at $\lambda < 912$ Å we decided to restrict the fit of the continuum redwards of 912 Å (i.e. 1095–1110, 1135–1150, 1450–1470, 1975–2010, and 2150–2200 Å). We note that the continuum window blueward of Ly α (i.e. 1095–1110 Å) may be contaminated by the Fe II + Fe III emission line. Given our limited resolution, we decided to keep it for consistency with T02. Results are unaffected if we neglect this window. We have considered as the error for the χ^2 minimization procedure the uncertainties in the stack, which are rather uniform at $\lambda > 912$ Å. Fig. 6 shows the rest-frame-stacked

spectrum extending from 600 to 2500 Å and the power-law fit to the continuum of the form $f_{\lambda} \propto \lambda^{\alpha_{\lambda}}$, where the best-fitting power law index is $\alpha_{\nu} = -0.61 \pm 0.01^8$ (dashed line).

Our best-fitting slope is consistent with the one estimated by T02 in the FUV (-0.69 ± 0.06 at $\lambda > 1216$ Å, see their table 1), and by S12 (-0.68 ± 0.14) in the same region (see Fig. 8). We note that the continuum regions adopted by S12 are the same as T02. Vanden Berk et al. (2001) also measured the continuum slope of a quasar composite utilizing 2200 spectra from SDSS, which span a redshift range of $0.044 \leq z \leq 4.789$. They find a slope of $\alpha_{\nu} = -0.44 \pm 0.10$ (no IGM absorption correction is applied), which is significantly different than the one we measured at 2.4σ . The difference can be attributed to the different continuum regions adopted by Vanden Berk et al. (2001, i.e. 1350–1365 and 4200–4230 Å).

We have also applied different continuum regions to gauge the dependence of the fit on the windows adopted. We considered the intervals listed by Decarli et al. (2010) for a sample of 96 quasars at redshift lower than three with spectra collected from several instrument (e.g. SDSS, ESO/NTT, Nordic Optical Telescope, and HST). Specifically, the alternative windows adopted are 1351–1362, 1452–1480, 1680–1710, 1796–1834, 1970–2010, and 2188–2243 Å. To also fit the UV part of the spectrum, we add to these the five intervals at $\lambda < 1200$ Å by T02. We find a slope of -0.63 ± 0.01 .

From Fig. 6, it is apparent that a single power-law provides an excellent fit up to ~ 900 Å where the continuum exhibits a break. A flatter (softer) spectrum is present at shorter wavelength, which is not modelled by the simple single power law. Solely for the purposes of a quantitative estimate of the position of the break, we have modelled the stacked spectrum by adding an exponential attenuation to the power law as

$$f_{\lambda} \propto \begin{cases} \lambda^{\alpha_{\lambda}}, & \text{if } \lambda \geq \lambda_b \\ \lambda^{\alpha_{\lambda}} \times e^{-(\lambda_b - \lambda)/\lambda_f}, & \text{if } 600 \text{ Å} < \lambda < \lambda_b \end{cases} \quad (2)$$

where λ_b and λ_f represent the break and the attenuation wavelengths, respectively. The best-fitting power-law index is $\alpha_{\nu} = -0.65 \pm 0.01$,

⁷ Emission from host-galaxy star light is not considered in the fit since our stack covers the rest-frame UV range of very bright AGN, where the expected contribution of the galaxy should be negligible.

⁸ In the following, we will refer to α_{ν} only. The relation between the fluxes in wavelength, $f_{\lambda} \propto \lambda^{\alpha_{\lambda}}$, and frequencies, $f_{\nu} \propto \nu^{\alpha_{\nu}}$, is $\alpha_{\nu} = -(2 + \alpha_{\lambda})$.

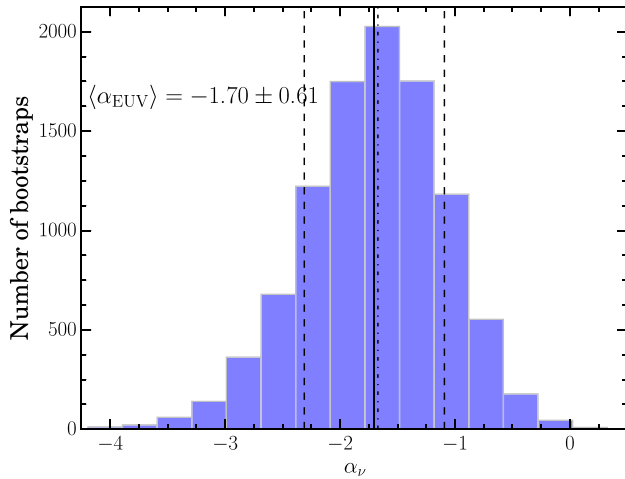


Figure 7. Histogram of the EUV slopes estimated from the procedure outlined in Section 5.1. The solid and the dashed lines represent the mean and 1σ dispersion, respectively, while the dot-dashed line denotes the median.

the break occurs at $\lambda_b = 922 \pm 31$ with an attenuation factor $\lambda_f \simeq 450$ not well constrained given the large uncertainties on the stacked spectrum in the far blue.

The break measured from our stacked spectrum is at bluer wavelengths with respect to the ones estimated from composites at much lower redshift. S12 found the break at $\lambda \sim 1000 \text{ \AA}$, whilst T02 located the wavelength break around $1200\text{--}1300 \text{ \AA}$ (a comparison among the various composites is given in Fig. 8). The difference from the latter is due to the poor IGM correction at those wavelengths, which artificially produces a break around the Ly α . We will analyse this issue further in Section 5.2. Additionally, the composite shown by S04 is much steeper than what we observed in the same wavelength range (i.e. $\alpha_v = 0.56^{+0.38}_{-0.28}$ at $\lambda = 650\text{--}1150 \text{ \AA}$).

Given that emission lines are much fainter in the EUV region than in the FUV/NUV and that uncertainties far in the blue are significant, we have considered the full stacked spectrum (continuum + lines) at $\lambda < 912 \text{ \AA}$ and found that the EUV slope is -1.65 ± 0.07 . The error on the EUV spectral index from the fit of the average spectrum is quite small given the uncertainties on the WFC3 stack itself at these wavelengths. Additionally, the flux in the EUV systematically depends on the IGM correction applied, such that fluxes and errors are strongly correlated. The simple χ^2 fitting technique is not appropriate in the EUV, and the bootstrap procedure is required. We have thus estimated the EUV slope from each realization in the bootstrap described in Section 4. Fig. 7 shows the distribution of the EUV slopes of mock stacks. We found $\alpha_{\text{EUV}} = -1.70 \pm 0.61$ (continuum + lines), which is a more reasonable value of the actual uncertainty. We will quote this as the final EUV slope. Lastly, we comment that it is evident from the right-hand panel of Fig. 6 that a single power law does not seem to be a satisfactory description of the region below 912 \AA . For example, there is no good reason why the feature at $\sim 730 \text{ \AA}$ should be considered an absorption feature. However, given the large uncertainties of our stack and to compare α_{EUV} with previous evaluations from the literature, we refrain either to employ more complicated functions or to define continuum windows to fit this region. We further discuss this issue in Section 5.6.

We also identify most of the emission lines in our stacked spectrum usually seen in optical spectra of high-redshift QSOs such as Ly β , Ly α , Si IV + O IV], C IV, and the semiforbidden line of C III] 1909. A number of weak lines show up in the average spectrum at

Table 2. Emission lines properties.

Line	λ_{obs} (\AA)	Flux ^a ($\text{erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$)	EW (\AA)
H I-Ly α + S IV	938.0	$8.2 \pm 2.7 \times 10^{-16}$	0.5
Ly γ + C III]	987.6	$4.4 \pm 1.1 \times 10^{-15}$	3.0
Ly β + O IV	1029.7	$9.8 \pm 0.4 \times 10^{-15}$	8.5
Fe II + Fe III	1122.0	$1.8 \pm 1.0 \times 10^{-15}$	1.3
Ly α	1216.1	$9.0 \pm 0.9 \times 10^{-14}$	74.0
Si IV + O IV]	1397.5	$6.7 \pm 1.4 \times 10^{-15}$	7.2
C IV	1544.6	$1.7 \pm 0.2 \times 10^{-14}$	17.8
He II	1635.7	$1.2 \pm 0.3 \times 10^{-15}$	2.0
O III]	1663.7	$7.3 \pm 1.0 \times 10^{-16}$	1.2
Al III	1861.0	$1.8 \pm 0.5 \times 10^{-15}$	2.6
C III]	1907.3	$9.7 \pm 0.9 \times 10^{-15}$	15.5

Notes. ^aNormalized fluxes are multiply by the average flux of the WFC3 sample at 1450 \AA .

$\lambda < 1216 \text{ \AA}$, including Ne VIII + O IV 772, O III 831, Ly γ + C III] 873, and Fe II + Fe III 1123. Blended lines from high-ionization states such as O IV 608, O V 630, N III 685, and O III 702, which are important diagnostics for investigating the physical conditions of broad emission line regions, may also be present, although it is impossible to reliably measure their strengths given the noise in our stacked spectrum at blue wavelengths. Weaker lines in the FUV include He II 1640, and O III] 1663. The properties of the emission lines for which the uncertainties on the stacked spectrum are small (i.e. $\lambda > 912 \text{ \AA}$) are listed in Table 2.

Summarizing, we estimated the continuum slope in the NUV/FUV ($\alpha_v = -0.61 \pm 0.01$) and in the EUV ($\alpha_v = -1.70 \pm 0.61$). We confirm the presence of a break in the quasar-stacked spectrum by employing a completely independent sample at high redshift. We also emphasize that previous estimates of the break were performed with very few spectra (~ 10 in T02 and 3–12 in S12) contributing at short wavelengths, whilst all 53 spectra in our high-redshift sample are contributing at 600 \AA .

5.2 Comparison with previous quasar composites

In this section, we will perform a detailed comparison between our $z_{\text{em}} \sim 2.4$ quasar average spectrum with previous works in literature. We caution the reader that a comprehensive and consistent explanation for all the differences hinges on multiple factors as, for example, sample selection biases.

In Fig. 8, we compared our WFC3 average spectrum with the AGN composites found from SDSS by Vanden Berk et al. (2001), from HST by T02, from COS by S12, and from FUSE by S04. The Vanden Berk et al. (2001), T02, and S12 composites are normalized at 1450 \AA , whilst the S04 composite is, instead, normalized to our IGM-corrected mean stack at $\lambda = 1114 \text{ \AA}$ ($\log v = 15.43$). We caution that the latter normalization is more problematic, because it is subject to uncertainties in the Ly α forest IGM correction employed by S04.

Several important points emerge from the comparison in Fig. 8. First, it is apparent that the IGM correction employed by T02 is significantly underestimated at $\lambda < 1216 \text{ \AA}$. In general, T02 attempted to identify individual LLSs by eye and used estimates of the Lyman limit optical depth to correct spectra with strong absorption. Whereas for lower column density systems, a single statistical correction for unidentified Lyman limit absorbers in the Ly α valley

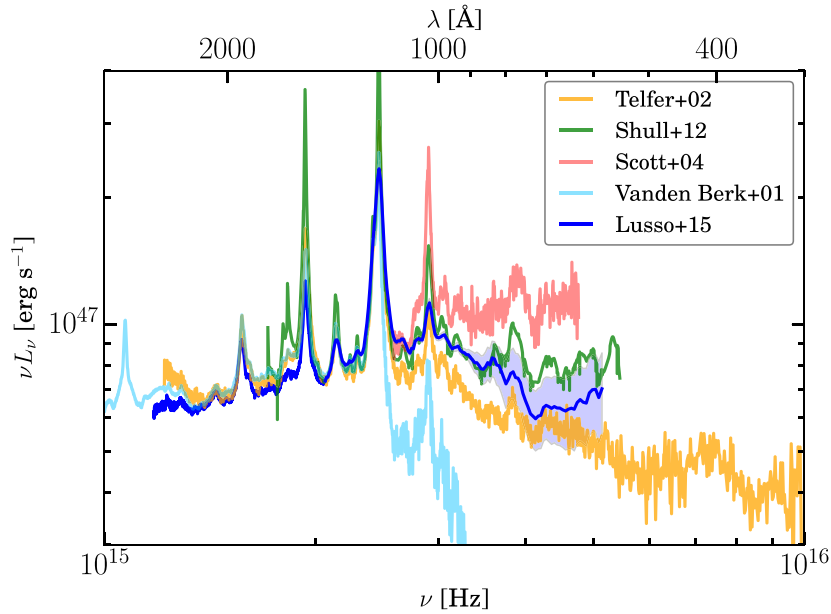


Figure 8. Mean IGM corrected QSO spectra with uncertainties (keys as in Figs 1 and 4) in the rest-frame $\log \nu$ – νL_ν plane. Rest-frame wavelengths (in Angstrom) are plotted on the top x-axis. The T02, S12, and Vanden Berk et al. (2001) composites are normalized to our average spectrum at 1450 Å, and are shown for comparison with the orange, green, and cyan solid lines, respectively. The S04 quasar composite (light red solid line) is normalized to our stack at $\lambda = 1114$ Å ($\log \nu = 15.43$).

(Moller & Jakobsen 1990) was applied per spectra using the column density distribution

$$\frac{\partial^2 n}{\partial z \partial N} \propto (1+z)^\gamma N^{-\beta}, \quad (3)$$

where n is the number of lines, z is the redshift, and N is the column density of neutral hydrogen. Values for the β factor considered by T02 are $\beta = 1.83$ for $3 \times 10^{14} < N_{\text{H I}} < 10^{16} \text{ cm}^{-2}$ and $\beta = 1.32$ for $N_{\text{H I}} > 10^{16} \text{ cm}^{-2}$ (Petitjean et al. 1993). For $3 \times 10^{12} < N_{\text{H I}} < 3 \times 10^{14} \text{ cm}^{-2}$, T02 used $\beta = 1.46$ (Hu et al. 1995). The application of this formula leads to a correction characterized by a stepwise behaviour, in which the opacity decreases toward shorter wavelengths (as already discussed in Binette et al. 2003). We computed the absolute i -band magnitude ($M_i(z=2)$)⁹ for the T02 sample and these values are plotted as a function of redshift in Fig. 9. There are about 20 quasars at $z > 2$ (contributing to the flux at $\lambda \sim 300$ Å in the SED analysis) in the T02 sample and for these objects their correction is lower by a factor of ~ 2 in flux at $\lambda = 600$ Å with respect to ours (see fig. 3 in Binette et al. 2003).

The spectral slope measured by S04 in the EUV is $\alpha_\nu = -0.56 \pm_{0.28}^{0.38}$, which is at variance with the spectral shape shown by our WFC3-stacked spectrum ($\alpha \simeq -1.7$ and with the other spectra from different samples) in the same wavelength range. Like T02, S04 adopted a similar procedure to correct for IGM absorption. LLSs were individually identified by eye and a single statistical correction for the line-of-sight absorption due to the Ly α forest and the Lyman valley was applied on single spectra as in equation (3). At variance with T02, the parameters to correct for the Ly α forest absorption are $\beta = 2$ for $1.6 \times 10^{12} < N_{\text{H I}} < 2.5 \times 10^{14} \text{ cm}^{-2}$, and $\beta = 1.35$ for $2.5 \times 10^{14} < N_{\text{H I}} < 5 \times 10^{16} \text{ cm}^{-2}$. For the red-

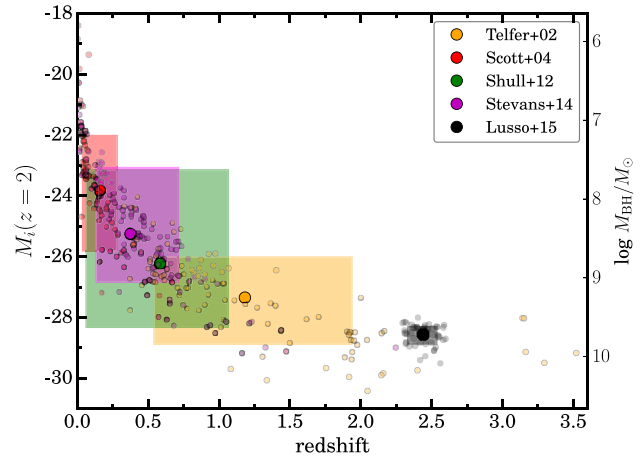


Figure 9. Absolute i -band magnitude (normalized at $z = 2$, K -corrected following Richards et al. 2006b) as a function of redshift. Shaded areas indicate the redshift and magnitude ranges for the different samples, estimated from the 16th and 84th percentiles. Large filled circles represent the median for the different samples: our WFC3 sample (black), T02 (orange), S04 (red), S12 (green), and S14 (magenta). Indicative values of the BH masses (in units of M_\odot) are plotted on the y-axis on the right. M_{BH} is estimated via $\lambda_{\text{Edd}} = L_{\text{bol}}/L_{\text{Edd}}$ assuming the average $\lambda_{\text{Edd}} = 0.35$ for the WFC3 sample. We have estimated the relation between L_{bol} and $M_i(z=2)$ to be $\log L_{\text{bol}} = -10.03 M_i(z=2)/26 + 36.32$ by fitting the sources in the DR7 quasar catalogue.

shift distribution parameter, S04 used $\gamma = 0.15$. This different parametrization leads to an IGM correction that goes from 5 to 10 per cent in the wavelength range 1200–600 Å at $z \sim 0.16$ (see their fig. 4). Any judgment about this IGM correction being underestimated is not trivial given the lower redshifts of the S04 AGN sample, compared with those of the WFC3 sample as shown in Fig. 9. However, biases in the S04 sample, and subtleties in

⁹ The absolute i -band magnitudes normalized at $z = 2$, K -corrected following Richards et al. (2006b), for the WFC3 sample have been taken from the DR7 quasar catalogue by Shen et al. (2011).

correcting for LLSs may contribute in the differences between the *FUSE* and the WFC3 spectra.

We have also compared our WFC3 average spectrum with the one by S12. The COS composite does not significantly differ from our spectrum within the uncertainties, but this similarity cannot be easily explained given the different average redshift, the sample selection,¹⁰ and the IGM correction employed by S12, which is solely performed by visual spectral inspection. The advantage of the S12 sample is the high S/N and the high spectral resolution, which allow them to fit the local continua of their objects (correcting for identified LLSs). Although S12 missed some partial LLSs with small Lyman continuum optical depths, the composite constructed by S14 does not significantly differ from the one in S12. This is due to the fact that the partial LLSs that S14 found have very low columns, $\log N_{\text{H I}} < 16$ (mostly in the 15.0–15.5 range) and their Lyman continuum opacity was negligible. Additionally, S14 used the pattern of higher-Lyman series lines rather than the Lyman edge. The fitted EUV composites of S12 and S14 are in fact quite similar: $\langle \alpha_{\text{v}} \rangle = -1.41 \pm 0.21$ in S12, while S14 found -1.41 ± 0.15 . We also note that, in the S14 analysis only ~ 20 spectra contribute at 700 Å, while few of them do not cover 912 Å such that the continuum normalization may be problematic. Given the bias in the UV samples, and the different approach in the IGM absorption correction, it is unclear what significance to attach to this agreement.

5.3 Comparison with photoionization models

In this section, we compare the line ratios observed in the WFC3 average spectrum and the line ratios predicted by photoionization models. Baskin et al. (2014) presented a radiation-pressure-dominated hydrostatic solution for gas in the broad line region (BLR), using the photoionization code *CLOUDY* (Ferland et al. 1998). We defer the reader to that paper for details of the physical model. Here, we compare their predicted line ratios with our estimates, for several prominent BLR emission lines. The predicted line ratios are based on three types of SEDs, differing from each other in the ionizing slope α_{ion} at energies between 1 Ryd and 1 keV (see their Section 2.2). These three SEDs are plotted in Fig. 10. We compare the observed and predicted line ratios, rather than the equivalent widths (EW) of single lines, since the average covering factor of the BLR in the WFC3 sample could be different than the covering factor of 0.3 assumed by Baskin et al. In order to calculate the predicted line ratios, we use the Baskin et al. model with solar metallicity, and a BLR distance where the line emissivity peaks (see their fig. 5), which is roughly the expected EW if one assumes a BLR which spans a range of distances. The predicted C III]/Ly α , He II/Ly α , and C IV/Ly α for $\alpha_{\text{ion}} = -1.2$, -1.6 , and -2.0 are listed in Table 3, together with the observed line ratios. The observed He II/Ly α and C IV/Ly α are consistent with the $\alpha_{\text{ion}} = -2.0$ model, which is softer than our derived slope of -1.7 , though within the uncertainties. The observed C III]/Ly α is a factor of 2 larger than expected for all assumed values of α_{ion} .

5.4 Comparison with GALEX

We have also increased our coverage at shorter wavelengths by considering the *GALEX* photometry of the WFC3 sample (see Section 2.3). These data have been corrected from IGM absorption as follows.

¹⁰ The COS sample was selected from the archive as the best available high-S/N spectra of AGNs in January 2011.

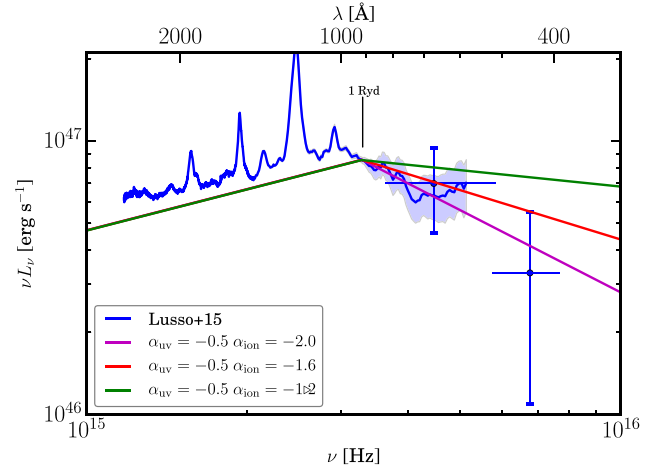


Figure 10. Comparison between the three types of SEDs adopted by Baskin, Laor & Stern (2014) with our WFC3 average spectrum. The green, red, and magenta solid lines correspond to $\alpha_{\text{ion}} = -1.2$, -1.6 and -2.0 , respectively. The blue points represent the IGM-corrected mean *GALEX*-forced photometry, where horizontal bars indicate the *GALEX* bandpasses.

Table 3. EW comparison.

Line	Ionization energy (eV)	EW(line)/EW(Ly α)			
		$\alpha_{\text{ion}} = -1.2$	$\alpha_{\text{ion}} = -1.6$	$\alpha_{\text{ion}} = -2.0$	-1.7 (Lusso + 15)
Ly α	13.6	1.0	1.0	1.0	1.0
C III]	24.0	0.1	0.10	0.13	0.21
He II	54.4	0.09	0.06	0.03	0.03
C IV	48.0	0.78	0.44	0.25	0.24

(i) We generate a set of 6000 mock quasar samples, following a similar procedure as in Section 3. Each sample contains 53 NUV and FUV fluxes (we again allow for duplications), and we compute the mean of these NUV and FUV fluxes.

(ii) We normalize the average *GALEX* fluxes to rest-frame 1450 Å (as for the WFC3 spectra). The 1450 Å flux is estimated from each average WFC3 spectrum of the mock quasar samples.

(iii) We then randomly draw one IGM transmission function from our suite of 10 000 and we integrate this function over the *GALEX* filter curves. This is repeated for each mock quasar sample.

(iv) The 6000 mock stacks are thus corrected from IGM absorption and averaged to produce the final mean NUV and FUV fluxes.

(v) The *GALEX* fluxes are then multiply by the average flux at 1450 Å of the WFC3 sample.

(vi) The uncertainties on the corrected *GALEX* fluxes are estimated from the dispersion of these 6000 mock stacks.

Both IGM absorption-corrected NUV and FUV bands are plotted in Fig. 10. We ignore quasar variability when stacking and plotting the results together with the stacked spectra. The *GALEX* NUV band almost covers the WFC3 spectra in the Lyman continuum, hence the fact that the NUV flux is consistent with the WFC3 stack within the uncertainties further check our results.

None the less, the error bars on the NUV and FUV fluxes are significant and do not allow us to further constrain any of the three SEDs, even though the FUV flux seems to exclude the SED with $\alpha_{\text{ion}} = -1.2$ consistently with the comparison with the predicted EWs.

5.5 AGN accretion disc models

A basic prediction of simple accretion disc models is that the disc temperature decreases as the BH mass increases (Shakura & Sunyaev 1973), thus

$$T = \left(\frac{GM\dot{M}}{4\pi\sigma r^3} \right)^{1/4} \sim 6.3 \times 10^5 \left(\frac{\dot{M}}{\dot{M}_E} \right)^{1/4} M_8^{-1/4} \left(\frac{r}{2R_g} \right)^{-3/4}, \quad (4)$$

where \dot{M}/\dot{M}_E is the Eddington ratio (λ_{Edd} , the accretion rate normalized to the Eddington accretion rate), $M_8 = M_{\text{BH}}/10^8 M_\odot$, and R_g is the gravitational radius ($R_g = GM/c^2$). For a scale r of $6R_g$ and an Eddington ratio $\lambda_{\text{Edd}} = 0.1$, the disc temperature goes from $\sim 5 \times 10^5$ to $\sim 8.7 \times 10^4$ K for an M_{BH} of 10^6 and $10^9 M_\odot$, respectively.¹¹ The disc temperature sets the peak of the BBB and, thus, we expect to see the location of the break changing as a function of M_{BH} .

The UV composites by T02 and S12 show a break between Ly α and 1000 Å, whilst the quasar composite by S04 does not show any. This might be consistent with the fact that the S04 is the lowest luminosity/redshift sample and presumably it has a lower M_{BH} on average than the S12 and T02 samples. Unfortunately, BH mass values are not available for most sources in the T02, S04, and S12 samples. S04 have compiled BH mass values from the literature for 22 objects in their sample (with a median M_{BH} of $1.6 \times 10^8 M_\odot$) and they found a significant correlation between the spectral index and BH mass, with the spectral slope being softer for higher M_{BH} , which may run in the expected direction. However, as we have discussed in the previous sections, both the presence and the position of the break depend strongly on the IGM correction considered.

Furthermore, the standard blackbody disc model depicted above, does not reproduce the observed soft X-ray emission seen in AGN, consequently other physical parameters are required such as BH spin and/or substantial Compton up-scattering (Laor et al. 1997, but see also Capellupo et al. 2015).

We further analysed our WFC3 average spectrum in the context of accretion disc models. In particular, we have considered the publicly available energetically self-consistent model, OPTXAGNF, developed by Done et al. (2012) within the XSPEC spectral fitting package. Their model contains three distinct spectral components, all powered by the energy released by a single accretion flow of constant mass accretion rate, \dot{M} , on to M_{BH} : an outer and inner disc, and an X-ray corona. The outer disc emits as a (colour temperature corrected) blackbody, the inner disc is where a fraction of the disc emission is Compton up-scattered, while the X-ray corona is where also a fraction of the emission is Compton up-scattered producing the power-law tail at high energies (see Section 4 in Done et al. 2012 for further details). This model is set by nine parameters: M_{BH} , λ_{Edd} , BH spin (a), radius of the X-ray corona (r_{corona}), outer accretion disc radius (r_{out}), electron temperature (kT_e), optical depth (τ) of the Comptonization region, photon index (Γ), and the fraction of energy dissipated as a hard X-ray power law from r_{corona} to the innermost stable circular orbit r_{isco} (f_{pl}).

Our data deliver constraints on M_{BH} and λ_{Edd} , which have been collected from the DR7 quasar catalogue (Shen et al. 2011) for all quasars in the WFC3 sample. The average M_{BH} estimated from the C IV line is $6 \times 10^9 M_\odot$, whilst the Eddington ratio is 0.35. The

Table 4. Accretion disc model parameters.

Ref. model	M_{BH} (M_\odot)	λ_{Edd}	r_{corona} (R_g)	a
AD1 (blue)	6×10^9	0.35	8	0.8
AD2 (red)	3×10^9	0.70	20	0.3
AD3 (green)	1.2×10^{10}	0.17	8	1.0
AD4 (black)	6×10^9	0.35	20	0.0

Notes. All models have the following parameters fixed: $\Gamma = 1.9$, $kT_e = 0.15$ keV, $f_{\text{pl}} = 0.2$, $\tau = 18$, and $r_{\text{out}} = 10^5 R_g$.

reliability of utilizing the C IV line is controversial, since C IV can be severely affected by non-virial motions such as outflows, winds, and strong absorption (e.g. Baskin & Laor 2005; Shen et al. 2008). We will refer to this model as AD1. We have thus investigated two additional models with half (AD2) and double (AD3) the average M_{BH} , where the average λ_{Edd} has been re-scaled consequently. The photon index has been fixed to 1.9, which is usually observed in quasars (e.g. Piconcelli et al. 2005). For the other parameters, we do not have any constraint from the data, and hence we have to assume values in order to reproduce both NUV/FUV/EUV and X-ray in a reasonable way. We have thus fixed the electron temperature kT_e at 0.15 keV (Gierliński & Done 2004), $f_{\text{pl}} = 0.2$, $\tau = 18$, and the default outer accretion disc radius, $r_{\text{out}} = 10^5 R_g$. The parameter values for the reference models are given in Table 4, and plotted in Fig. 11. All models are normalized to 1450 Å. We note that all the models are only meant to provide a reasonable description of the observables, since we did not attempt any simultaneous fit of the optical–UV/X-ray data. X-ray data for the 10 quasars detected by *ROSAT* and *XMM–Newton* are shown in Fig. 11 for completeness.

A key aspect of these models is that the peak of the BBB is set by the BH mass, spin and mass accretion rate through the outer disc. Interestingly, the AD4 model (i.e. average M_{BH} and λ_{Edd} and spin zero) do not provide a good representation of the FUV/EUV region of our average spectrum,¹² with the BBB being colder than observed (see also Lawrence 2012). This problem can be solved if we instead consider the AD2 but with the BH spin set to zero. If the BH spin increases, the distance of the innermost stable circular orbit reduces, and this affects the disc temperature as described by equation (4).

Summarizing, if we consider the BH masses from C IV to be correct, we need to have non-zero values of the BH spin to match the data, with higher values of the BH spin as the M_{BH} increases (see Trakhtenbrot et al. 2014 for similar results, i.e. black hole masses above $\sim 3 \times 10^9 M_\odot$ at $z \sim 2–3$ are consistent with very high spins). We do not need BH spin if we instead consider a factor of 2 mass lower, on average. The latter is in agreement with previous results that have found BH mass values from C IV systematically higher, once calibrated with virial mass estimators based on the Mg II (e.g. Shen et al. 2008; Park et al. 2013, but see also Trakhtenbrot & Netzer 2012).

In principle, the combination of reliable M_{BH} and FUV/EUV spectra (properly corrected for IGM absorption) may provide hints on the quasar BH spin.

It is not possible to test if the location of the break changes as a function of M_{BH} within the WFC3 since it has a narrow luminosity range. We thus need to have a comparison sample at low redshift/luminosity, hence lower M_{BH} . The *FUSE* composite by S04

¹² The radius of the corona has been fixed to $20 R_g$, instead of 8 as in AD1.

¹¹ Equation (4) is valid in the Newtonian limit, and therefore may be not accurate to estimate the temperature in the disc, which is highly relativistic. See equations 3–5 and table 1 in Laor & Davis (2011) for a more accurate approach.

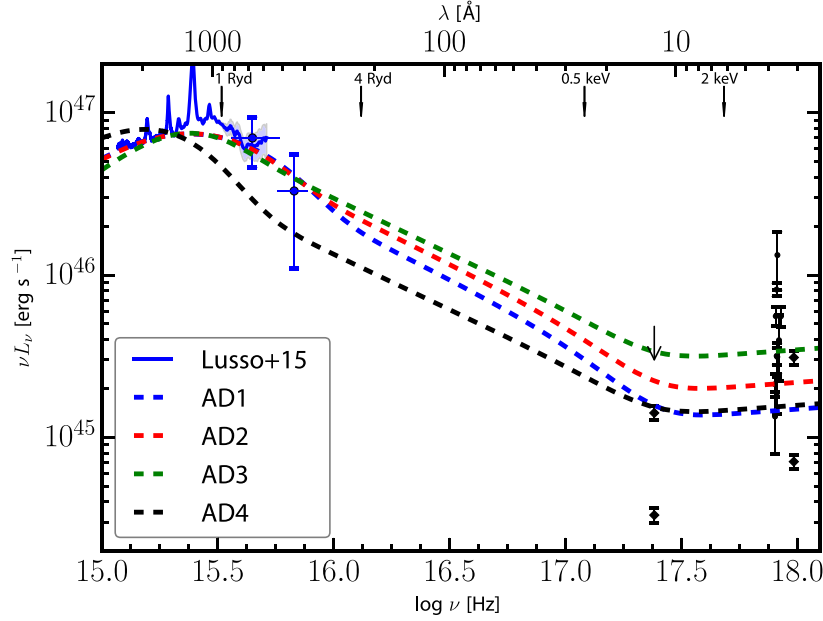


Figure 11. Accretion disc models. Keys as in Fig. 10. The blue (AD1), red (AD2), and green (AD3) dashed lines represent the three reference accretion disc models for $M_{\text{BH}} = 6 \times 10^9$ with $\lambda_{\text{Edd}} = 0.35$ ($a = 0.8$), $3 \times 10^9 M_{\odot}$ with $\lambda_{\text{Edd}} = 0.70$ ($a = 0.3$), and 1.2×10^{10} with $\lambda_{\text{Edd}} = 0.17$ ($a = 1.0$), respectively. The black dashed line is the same as AD1 but with $a = 0$ and $r_{\text{corona}} = 20$ (AD4). All models are normalized to 1450 Å. X-ray data for the 10 quasars detected by *ROSAT* are plotted with filled black points. The *XMM-Newton* soft and hard luminosities for J0755 + 2204 and for J1119 + 1302 are plotted as black diamonds. The black arrow represent the *ROSAT* flux limit at 1 keV.

is constructed with low redshifts/luminosities AGN (see Fig. 9) and interestingly it shows a hard spectral slope in the EUV with no break consistent with the BBB being produced by small M_{BH} . This might point towards the direction of the predicted trend, but further studies are needed. In particular, a natural extension of the present analysis is to cover bluer wavelengths with high-resolution quasar spectra.

5.6 Quasar emissivity and photoionization rate

The UV background is governed by the ionizing emissivities of the relevant source populations (typically quasars and star-forming galaxies) and cosmological radiative transfer in the IGM (e.g. Faucher-Giguère et al. 2009; Haardt & Madau 2012). The specific comoving emissivity of the quasar population at redshift z and frequency ν ,

$$\epsilon_{\nu}(\nu, z) = \int_{L_{\min}}^{\infty} \phi(L, z) L_{\nu}(L, \nu) dL, \quad (5)$$

critically depends on the adopted quasar luminosity function $\phi(L, z)$, its faint end (L_{\min}), and the quasar SED ($L_{\nu}(L, \nu)$). In particular, since the quasar luminosity function is typically determined in the NUV–optical, the specific emissivity at the Lyman limit $\epsilon_{\nu, 912}$ depends on the chosen FUV and EUV SED parametrization, which is commonly taken as a luminosity-independent broken power law based on quasar composite spectra (Z97; T02). Resulting uncertainties in the specific emissivity are rarely discussed (see Faucher-Giguère et al. 2008 for an exception), or limited to variations in the EUV spectral index (Faucher-Giguère et al. 2009). Consequently, current estimates of the specific Lyman limit emissivity of quasars still vary by a factor ~ 2 at $z = 3\text{--}4$ (Hopkins, Richards & Hernquist

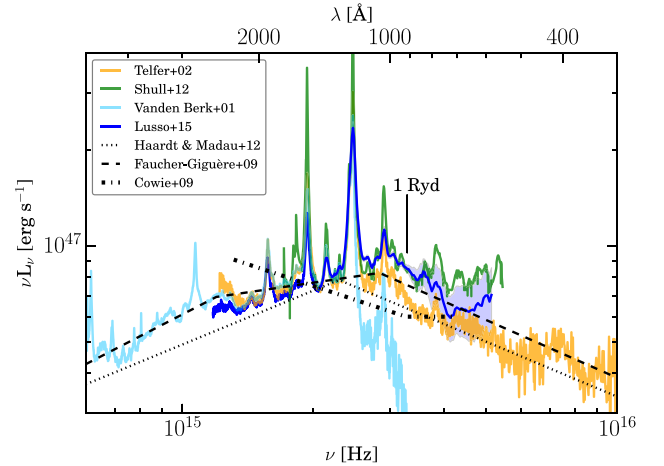


Figure 12. As Fig. 8 but including recent broken power-law parameterizations of the quasar continuum (Cowie et al. 2009; Faucher-Giguère et al. 2009; Haardt & Madau 2012). The SEDs have been normalized at the adopted rest-frame wavelengths (Faucher-Giguère et al. 2009; 4400 Å) or assumed at 1450 Å if not given by the authors (Cowie et al. 2009; Haardt & Madau 2012). All parameterizations underestimate the quasar Lyman limit flux of our stacked spectrum and recent quasar composite spectra (S12, S14).

2007; Siana et al. 2008; Cowie, Barger & Trouille 2009; Masters et al. 2012).

Fig. 12 compares our stacked spectrum and various quasar composites to broken power-law parameterizations used to estimate the specific quasar Lyman limit emissivity and/or the quasar contribution to the UV background (Cowie et al. 2009;

Faucher-Giguère et al. 2009; Haardt & Madau 2012). Most previous SED parameterizations have underestimated the quasar Lyman continuum flux because they have been based on the T02 composite spectrum without a proper IGM correction (Section 5.2). Adopting our power-law fit $f_\nu \propto \nu^{-0.61}$ instead, we obtain Lyman limit flux ratios of $f_{\nu,912}^{\text{L14}}/f_{\nu,912}^{\text{FG09}} = 1.12$ (Faucher-Giguère et al. 2009) and $f_{\nu,912}^{\text{L14}}/f_{\nu,912}^{\text{HM12}} = 1.38$ (Haardt & Madau 2012) as correction factors to the quasar Lyman limit emissivity $\epsilon_{\nu,912}(z)$ for a fixed luminosity function.¹³ We point out that Cowie et al. (2009) computed a power-law slope of -1.35 at $700 < \lambda < 2300$ Å obtained from *GALEX* photometry of a $z_{\text{em}} \sim 1$ AGN sample. This implies a 700 to 2300 Å flux ratio of 0.20. Cowie et al. then assumed the actual SED below the break to be $f_\nu \propto \nu^{-1}$, which returns an extrapolated flux ratio $f_{\nu,912}/f_{\nu,2300} = 0.26$. In Fig. 12, we have taken the break at exactly 912 Å.¹⁴ We also note here that the mean 2300 Å luminosity for the objects in the Cowie et al. sample is much lower ($\nu L_\nu \simeq 5 \times 10^{44}$ erg s⁻¹) than the one of our WFC3 sample. Adopting a standard concave SED shape instead, their $\epsilon_{\nu,912}$ values increase by a factor of $\simeq 2$. While this correction makes their values consistent with higher estimates of $\epsilon_{\nu,912}$ (Hopkins et al. 2007; Siana et al. 2008; Masters et al. 2012), we note that Siana et al. (2008) and Masters et al. (2012) assume a low flux ratio $f_{\nu,912}/f_{\nu,1450} = 0.58$ based on the T02 composite to convert their $\epsilon_{\nu,1450}$ to $\epsilon_{\nu,912}$. Our SED yields a 30 per cent increase to their estimates, such that $\epsilon_{\nu,912} \simeq 5 \times 10^{24}$ erg s⁻¹ Hz⁻¹ Mpc⁻³ at $z \sim 3.2$. With this correction, quasars may account for the total specific Lyman limit emissivity estimated from the Ly α forest at $z \lesssim 3.2$ (Becker & Bolton 2013), although a comparable contribution from star-forming galaxies is allowed within the current large uncertainties (Becker & Bolton 2013; Mostardi et al. 2013; Nestor et al. 2013). The quasar contribution seems to decrease substantially to $z \sim 4$, as the low (corrected) value of $\epsilon_{\nu,912} \simeq 1.8 \times 10^{24}$ erg s⁻¹ Hz⁻¹ Mpc⁻³ from Masters et al. (2012) is just ~ 20 per cent of the total value inferred from the Ly α forest (Becker & Bolton 2013, although see Glikman et al. 2011 for a ~ 15 times higher estimate of $\epsilon_{\nu,1450}$ at $z \sim 4$).

From the Lyman limit emissivity, we can estimate the quasar contribution to the UV background photoionization rate

$$\Gamma_{\text{HI}}(z) = \int_{\nu_{912}}^{\infty} \frac{4\pi J_\nu(\nu, z)}{h\nu} \sigma_{\text{HI}}(\nu) d\nu, \quad (6)$$

where $J_\nu(\nu, z)$ is the mean intensity of the ionizing background at redshift z and $\sigma_{\text{HI}}(\nu) \simeq \sigma_{912}(\nu/\nu_{912})^{-3}$ is the H I photoionization cross-section with $\sigma_{912} = 6.33 \times 10^{-18}$ cm² at the Lyman limit frequency ν_{912} . At high redshifts $z \gtrsim 4$, the mean free path of Lyman continuum photons in the IGM, $\lambda_{\text{mfp}}(\nu, z)$, is much shorter than the Horizon, such that the UV radiation field at any given point is dominated by local sources within a sphere of radius λ_{mfp} , and cosmological redshift effects in $J_\nu(\nu, z)$ can be neglected. In this ‘local-source approximation’, the intensity of the UV background is

$$J_\nu(\nu, z) \simeq \frac{1}{4\pi} \lambda_{\text{mfp}}(\nu, z) (1+z)^3 \epsilon_\nu(\nu, z) \quad (7)$$

(Madau, Haardt & Rees 1999; Schirber & Bullock 2003; Faucher-Giguère et al. 2008; Becker & Bolton 2013), with $\lambda_{\text{mfp}}(\epsilon_\nu)$ given

in proper (comoving) units. The proper mean free path to Lyman limit photons is well described by a power law

$$\lambda_{\text{mfp},912} = 37 \left(\frac{1+z}{5} \right)^{-5.4} \text{ Mpc} \quad (8)$$

at $2.3 < z < 5.5$ (Worseck et al. 2014), whereas its frequency dependence is

$$\lambda_{\text{mfp}}(\nu, z) = \lambda_{\text{mfp},912} \left(\frac{\nu}{\nu_{912}} \right)^{1.5} \quad (9)$$

for a column density distribution $f(N_{\text{HI}}, z) \propto N_{\text{HI}}^{-1.5}$ (Faucher-Giguère et al. 2008).¹⁵ With our power-law parametrization of the quasar SED¹⁶ $f_\nu \propto \nu^{\alpha_{\text{EUV}}}$ at $\lambda > 912$ Å, we can write the comoving specific emissivity of quasars as

$$\epsilon_\nu(\nu, z) = \epsilon_{\nu,912}(z) \left(\frac{\nu}{\nu_{912}} \right)^{\alpha_{\text{EUV}}}. \quad (10)$$

Combining equations (7)–(10) and integrating equation (6), the quasar contribution to the UV background photoionization rate is

$$\begin{aligned} \Gamma_{\text{HI}} &\simeq 4.6 \times 10^{-13} \text{ s}^{-1} \left(\frac{\epsilon_{\nu,912}}{10^{24} \text{ erg s}^{-1} \text{ Hz}^{-1} \text{ Mpc}^{-3}} \right) \\ &\times \left(\frac{1+z}{5} \right)^{-2.4} \frac{1}{1.5 - \alpha_{\text{EUV}}}. \end{aligned} \quad (11)$$

As pointed out by Becker & Bolton (2013), the local-source approximation becomes increasingly inaccurate towards lower redshifts due to redshifting of Lyman limit photons. Consequently, equation (11) overestimates Γ_{HI} for a given $\epsilon_{\nu,912}$. This may be significant already at $z \sim 4$, since the quasar H I photoionization rate estimated from equation (11) with our $\alpha_{\text{EUV}} = -1.7$ and the rescaled emissivity from Masters et al. (2012) is somewhat higher than the value obtained by Haardt & Madau (2012) with approximately the same emissivity and mean free path but including redshift and radiative transfer effects.

Nevertheless, equation (11) provides an estimate on the uncertainty of Γ_{HI} induced by the significant uncertainty in the mean EUV spectral slope ($\alpha_{\text{EUV}} = -1.70 \pm 0.61$, Section 5.1). For any fixed $\epsilon_{\nu,912}$ and z , the ratio $\Gamma_{\text{HI}}/\Gamma_{\text{HI,ref}} = 3.2/(1.5 - \alpha_{\text{EUV}})$ quantifies the variation of the bootstrap realizations of Γ_{HI} with respect to a reference value $\Gamma_{\text{HI,ref}}$ that uses the slope $\alpha_{\text{EUV}} = -1.7$ determined from the stack. As shown in Fig. 13, the uncertain EUV slope results in a ~ 20 per cent uncertainty in Γ_{HI} . We conclude that detailed photoionization models of the IGM need to take into account the uncertainty in the mean EUV spectral index that is dominated by the large IGM correction in our study, but dominated by sample variance in previous work (T02; S14).

Similarly, the uncertainty in α_{EUV} affects estimates of the He II Lyman limit emissivity of quasars in models of He II reionization and the post-reionization UV background. Extrapolating our power-law fit to the He II Lyman limit, we obtain flux ratios

¹⁵ Note that this is an approximation due to departures of $f(N_{\text{HI}}, z)$ from a single power law (e.g. Prochaska, O’Meara & Worseck 2010; Haardt & Madau 2012).

¹⁶ We stress that due to the low resolution we cannot fit the EUV continuum beneath existing EUV emission lines (S12; S14). For computing the photoionization rate, this turns into an advantage, as we approximately account for the EUV lines that also contribute ionizing photons. With sufficient wavelength coverage and spectral resolution, the SED should be integrated instead of the power-law continuum.

¹³ Note that Haardt & Madau (2012) simplistically assume $\epsilon_{\nu,912}(z)$ and $L_\nu(\nu)$ to be independent.

¹⁴ The broken power-law quasar SED is then $f_\nu \propto \nu^{-1.45}$ ($f_\nu \propto \nu^{-1}$) at $\lambda > 912$ Å ($\lambda < 912$ Å).

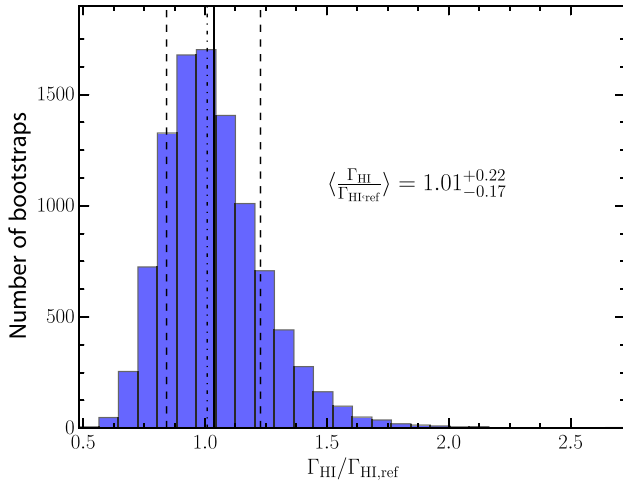


Figure 13. Histogram of the ratio of H I photoionization rates $\Gamma_{\text{H I}} / \Gamma_{\text{H I,ref}} = 3.2 / (1.5 - \alpha_{\text{EUV}})$. The mean, the median, the 16th and the 84th percentile are plotted as solid, dot-dashed, and dashed lines, respectively.

$f_{\nu,1450}/f_{\nu,228} = 13.8^{+18.2}_{-7.5}$ and $f_{\nu,912}/f_{\nu,228} = 10.0^{+14.2}_{-5.4}$, respectively. Currently, the mean quasar SED is poorly constrained at $\lambda < 500 \text{ \AA}$ due to small sample size (< 10 quasars) and incomplete IGM correction in the T02 sample. As the He II photoionization rate depends on the unconstrained spectral slope at $\lambda < 228 \text{ \AA}$, models of the He II-ionizing background are highly uncertain.

Lastly, we emphasize that any broken power law does not have any physical grounds, and thus it is a very poor description of the quasar SED. In principle, one should consider the full spectral information, although spectra do not usually cover energies well beyond 1 Ryd.

6 SUMMARY AND CONCLUSIONS

We presented the first UV average spectrum at high redshift ($\langle z_{\text{em}} \rangle \sim 2.44$) with proper correction for the intervening Ly α forest and continuum absorption by employing the state-of-the-art IGM transmission function which have been calibrated from multiple quasar absorption line observables. The sample consists of 53 quasars selected from the *HST* survey for LLS using the WFC3 presented in O11.

The rest-frame continuum slope of the full sample shows a break at around 912 \AA with an FUV spectral index $\alpha_{\nu} = -0.61 \pm 0.01$ and a softening at shorter wavelengths ($\alpha_{\nu} = -1.70 \pm 0.61$).

Our analysis highlights the fact that slope and spectral break in the *HST* composite reported by T02 are incorrect due to an underestimated intergalactic absorption correction, especially at short wavelength where high- z quasars are contributing. The *FUSE* composite by S04 might be less affected by this problem given the low average redshift of their sample ($z \sim 0.16$). The COS composite by S12 (and S14) does not significantly differ from our spectrum within the uncertainties.

Our observed broad line ratios are in good agreement with those predicted by the photoionization models discussed in Baskin et al. (2014), where the input quasar continuum has an EUV slope consistent with the one we observed.

The accretion disc + X-ray corona models constructed by Done et al. (2012) show that we need to have non-zero values of the BH

spin to match the data, with higher values of the BH spin as the M_{BH} increases (assuming that the BH masses derived from C IV are correct). We do not need BH spin if we instead consider a factor of 2 lower masses (on average). The latter is in agreement with previous results that have found BH mass values from C IV systematically higher once calibrated with virial mass estimators based on the Mg II (e.g. Shen et al. 2008, but see also Runnoe et al. 2013).

Finally, we found that the lack of knowledge of the spectral slope results in an uncertainty on the quasar photoionization rate of the order of 20 per cent. This effect should be taken into account by any photoionization model.

As a final comment, we want to stress that all previous analyses were based on quasar samples selected with an unknown selection function (i.e. the best from the archive), and were characterized by small samples sizes, especially at short wavelengths (less than 10 AGN observations cover the wavelength range between 450 and 600 \AA in the S04 sample, and between 10 and 20 AGN in the T02 sample, and only six contributing at 600 \AA in the S12 sample). Furthermore, errors on the composite were not published, so an assessment of the sample variance is impossible. A single statistical correction for the IGM absorption should not be applied for a broad range of redshifts and to single spectra, since LLSs cannot all be identified or corrected for, especially not the partial LLSs.

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REFERENCES

- Baskin A., Laor A., 2005, MNRAS, 356, 1029
- Baskin A., Laor A., Stern J., 2014, MNRAS, 438, 604
- Becker G. D., Bolton J. S., 2013, MNRAS, 436, 1023

¹⁷ <http://www.mpia-hd.mpg.de/ENIGMA/>

- Binette L., Rodríguez-Martínez M., Haro-Corzo S., Ballinas I., 2003, *ApJ*, 590, 58
- Becker R. H., White R. L., Helfand D. J., 1995, *ApJ*, 450, 559
- Bershady M. A., Charlton J. C., Geoffroy J. M., 1999, *ApJ*, 518, 103
- Brotherton M. S., Tran H. D., Becker R. H., Gregg M. D., Laurent-Muehleisen S. A., White R. L., 2001, *ApJ*, 546, 775
- Capellupo D. M., Netzer H., Lira P., Trakhtenbrot B., Mejía-Restrepo J., 2015, *MNRAS*, 446, 3427
- Compastella M., Cantalupo S., Porciani C., 2013, *MNRAS*, 435, 3169
- Condon J. J., Cotton W. D., Greisen E. W., Yin Q. F., Perley R. A., Taylor G. B., Broderick J. J., 1998, *AJ*, 115, 1693
- Cowie L. L., Barger A. J., Trouille L., 2009, *ApJ*, 692, 1476
- Czerny B., Elvis M., 1987, *ApJ*, 321, 305
- Davé R., Tripp T. M., 2001, *ApJ*, 553, 528
- Decarli R., Falomo R., Treves A., Kotilainen J. K., Labita M., Scarpa R., 2010, *MNRAS*, 402, 2441
- Done C., Davis S. W., Jin C., Blaes O., Ward M., 2012, *MNRAS*, 420, 1848
- Elvis M. et al., 1994, *ApJS*, 95, 1
- Elvis M. et al., 2012, *ApJ*, 759, 6
- Faucher-Giguère C.-A., Lidz A., Hernquist L., Zaldarriaga M., 2008, *ApJ*, 688, 85
- Faucher-Giguère C.-A., Lidz A., Zaldarriaga M., Hernquist L., 2009, *ApJ*, 703, 1416
- Ferland G. J., Korista K. T., Verner D. A., Ferguson J. W., Kingdon J. B., Verner E. M., 1998, *PASP*, 110, 761
- Fitzpatrick E. L., 1999, *PASP*, 111, 63
- Fontanot F., Cristiani S., Vanzella E., 2012, *MNRAS*, 425, 1413
- Fontanot F., Cristiani S., Pfrommer C., Cupani G., Vanzella E., 2014, *MNRAS*, 438, 2097
- Francis P. J., Hewett P. C., Foltz C. B., Chaffee F. H., Weymann R. J., Morris S. L., 1991, *ApJ*, 373, 465
- Furlanetto S. R., 2009, *ApJ*, 703, 702
- Gierliński M., Done C., 2004, *MNRAS*, 349, L7
- Glikman E., Djorgovski S. G., Stern D., Dey A., Jannuzi B. T., Lee K.-S., 2011, *ApJ*, 728, L26
- Haardt F., Madau P., 1996, *ApJ*, 461, 20
- Haardt F., Madau P., 2012, *ApJ*, 746, 125
- Haardt F., Maraschi L., 1991, *ApJ*, 380, L51
- Haardt F., Maraschi L., 1993, *ApJ*, 413, 507
- Hopkins P. F., Richards G. T., Hernquist L., 2007, *ApJ*, 654, 731
- Hu E. M., Kim T.-S., Cowie L. L., Songaila A., Rauch M., 1995, *AJ*, 110, 1526
- Inoue A. K., Iwata I., 2008, *MNRAS*, 387, 1681
- Inoue A. K., Shimizu I., Iwata I., Tanaka M., 2014, *MNRAS*, 442, 1805
- Jiang L., Fan X., Ivezić Ž., Richards G. T., Schneider D. P., Strauss M. A., Kelly B. C., 2007, *ApJ*, 656, 680
- Jiang L. et al., 2008, *AJ*, 135, 1057
- Kalberla P. M. W., Burton W. B., Hartmann D., Arnal E. M., Bajaja E., Morras R., Pöppel W. G. L., 2005, *A&A*, 440, 775
- Kellermann K. I., Sramek R., Schmidt M., Shaffer D. B., Green R., 1989, *AJ*, 98, 1195
- Komatsu E. et al., 2009, *ApJS*, 180, 330
- Laor A., Davis S. W., 2011, *MNRAS*, 417, 681
- Laor A., Fiore F., Elvis M., Wilkes B. J., McDowell J. C., 1997, *ApJ*, 477, 93
- Lawrence A., 2012, *MNRAS*, 423, 451
- McQuinn M., Lidz A., Zaldarriaga M., Hernquist L., Hopkins P. F., Dutta S., Faucher-Giguère C.-A., 2009, *ApJ*, 694, 842
- Madau P., 1995, *ApJ*, 441, 18
- Madau P., Haardt F., Rees M. J., 1999, *ApJ*, 514, 648
- Masters D. et al., 2012, *ApJ*, 755, 169
- Meiksin A., 2005, *MNRAS*, 356, 596
- Meiksin A., 2006, *MNRAS*, 365, 807
- Meiksin A., White M., 2003, *MNRAS*, 342, 1205
- Miralda-Escudé J., Haehnelt M., Rees M. J., 2000, *ApJ*, 530, 1
- Moller P., Jakobsen P., 1990, *A&A*, 228, 299
- Mostardi R. E., Shapley A. E., Nestor D. B., Steidel C. C., Reddy N. A., Trainor R. F., 2013, *ApJ*, 779, 65
- Nestor D. B., Shapley A. E., Kornei K. A., Steidel C. C., Siana B., 2013, *ApJ*, 765, 47
- O'Meara J. M., Prochaska J. X., Chen H.-W., Madau P., 2011, *ApJS*, 195, 16 (O11)
- O'Meara J. M., Prochaska J. X., Worseck G., Chen H.-W., Madau P., 2013, *ApJ*, 765, 137 (O13)
- Park D., Woo J.-H., Denney K. D., Shin J., 2013, *ApJ*, 770, 87
- Peek J. E. G., Schiminovich D., 2013, *ApJ*, 771, 68
- Petitjean P., Webb J. K., Rauch M., Carswell R. F., Lanzetta K., 1993, *MNRAS*, 262, 499
- Piconcelli E., Jimenez-Bailón E., Guainazzi M., Schartel N., Rodríguez-Pascual P. M., Santos-Lleó M., 2005, *A&A*, 432, 15
- Prochaska J. X., O'Meara J. M., Worseck G., 2010, *ApJ*, 718, 392
- Prochaska J. X., Madau P., O'Meara J. M., Fumagalli M., 2014, *MNRAS*, 438, 476
- Richards G. T. et al., 2002, *AJ*, 123, 2945
- Richards G. T. et al., 2006a, *ApJS*, 166, 470
- Richards G. T. et al., 2006b, *AJ*, 131, 2766
- Runnoe J. C., Brotherton M. S., Shang Z., DiPompeo M. A., 2013, *MNRAS*, 434, 848
- Sanders D. B., Phinney E. S., Neugebauer G., Soifer B. T., Matthews K., 1989, *ApJ*, 347, 29
- Schirber M., Bullock J. S., 2003, *ApJ*, 584, 110
- Schlaflly E. F., Finkbeiner D. P., 2011, *ApJ*, 737, 103
- Schlegel D. J., Finkbeiner D. P., Davis M., 1998, *ApJ*, 500, 525 (SFD)
- Schneider D. P. et al., 2007, *AJ*, 134, 102
- Scott J. E., Kriss G. A., Brotherton M., Green R. F., Hutchings J., Shull J. M., Zheng W., 2004, *ApJ*, 615, 135 (S04)
- Shakura N. I., Sunyaev R. A., 1973, *A&A*, 24, 337
- Shang Z. et al., 2011, *ApJS*, 196, 2
- Shankar F., Mathur S., 2007, *ApJ*, 660, 1051
- Shen Y., Greene J. E., Strauss M. A., Richards G. T., Schneider D. P., 2008, *ApJ*, 680, 169
- Shen Y. et al., 2011, *ApJS*, 194, 45
- Shull J. M., Stevans M., Danforth C. W., 2012, *ApJ*, 752, 162 (S12)
- Siana B. et al., 2008, *ApJ*, 675, 49
- Stern J., Behar E., Laor A., Baskin A., Holczer T., 2014, *MNRAS*, 445, 3011
- Stevens M. L., Shull J. M., Danforth C. W., Tilton E. M., 2014, *ApJ*, 794, 75 (S14)
- Telfer R. C., Zheng W., Kriss G. A., Davidsen A. F., 2002, *ApJ*, 565, 773 (T02)
- Trakhtenbrot B., 2014, *ApJL*, 789, L9
- Trakhtenbrot B., Netzer H., 2012, *MNRAS*, 427, 3081
- Trammell G. B., Vanden Berk D. E., Schneider D. P., Richards G. T., Hall P. B., Anderson S. F., Brinkmann J., 2007, *AJ*, 133, 1780
- Vanden Berk D. E. et al., 2001, *AJ*, 122, 549
- Voges W. et al., 1999, *A&A*, 349, 389
- White N. E., Giommi P., Angelini L., 1994a, *BAAS*, 26, 1372
- White N. E., Giommi P., Angelini L., 1994b, *IAU Circ.*, 6100, 1
- Willott C. J. et al., 2010, *AJ*, 139, 906
- Worseck G., Prochaska J. X., 2011, *ApJ*, 728, 23
- Worseck G. et al., 2014, *MNRAS*, 445, 1745
- Zheng W., Kriss G. A., Telfer R. C., Grimes J. P., Davidsen A. F., 1997, *ApJ*, 475, 469 (Z97)

APPENDIX A: MULTIWAVELENGTH PROPERTIES OF THE WFC3 QUASAR SAMPLE

We listed the multiwavelength properties of the WFC3 quasar sample in Tables A1, A2, and A3.

Table A1. RASS properties of the WFC3 quasar sample.

Name	z	Exp. ^a (s)	$N_{\mathrm{H}, \mathrm{I}}^b$ (cm^{-2})	$F_{[0.5-2]\mathrm{keV}}^c$ ($\mathrm{erg\ s^{-1}\ cm^{-2}}$)	$L_{[0.5-2]\mathrm{keV}}^d$ ($\mathrm{erg\ s^{-1}}$)
J0755 + 2204	2.319	23 963	5.73	4.25 ± 1.76	1.87 ± 0.77
J1107 + 6420	2.316	7728	1.00	6.74 ± 1.11	2.95 ± 0.48
J1119 + 1302	2.394	20 105	1.96	9.16 ± 1.93	4.33 ± 0.98
J1235 + 6301	2.383	3311	1.21	5.56 ± 1.45	2.61 ± 0.68
J1253 + 0516	2.398	201	2.17	38.66 ± 15.21	18.41 ± 7.24
J1335 + 4542	2.452	616	1.88	10.94 ± 4.80	5.50 ± 2.41
J1415 + 3706	2.374	18 476	0.91	9.47 ± 1.11	4.40 ± 0.52
J1454 + 0325	2.368	28 496	3.62	24.44 ± 2.70	11.24 ± 0.96
J1625 + 2646	2.518	5006	3.37	14.40 ± 2.04	7.73 ± 1.09
J1651 + 4002	2.343	7649	1.45	17.21 ± 2.29	7.74 ± 1.03

^aExposure time.^bGalactic column density ($\times 10^{20}$) (Kalberla et al. 2005).^cX-ray fluxes ($\times 10^{-14}$) at 0.5–2 keV calculated from the count rates in the *ROSAT* band using a power-law spectrum with a photon index $\Gamma = 2.0$ corrected for Galactic absorption (Kalberla et al. 2005).^dRest-frame luminosity ($\times 10^{45}$) in the 0.5–2 keV band (estimated from the unabsorbed X-ray flux).**Table A2.** Radio properties of the WFC3 quasar sample.

Name	F_{int}^a	Dist. ^b	$f_{6\mathrm{cm}}^c$	$\log f_{2500}^d$	M_{2500}^e	R^f
J0806 + 5041	19.66	0.3	20.02	−26.45	−28.96	57.08
J0854 + 5327	22.07	0.2	22.35	−26.04	−29.96	24.52
J1119 + 1302	12.97	0.4	23.59	−26.58	−28.57	90.48
J1335 + 4542	269.05	0.1	273.80	−26.62	−28.53	1150.32
J1342 + 6015	1.82	1.3	1.83	−26.51	−28.76	5.90
J1415 + 3706	409.69	0.3	412.16	−26.38	−29.05	995.75
J1454 + 0324	4.92	0.2	4.95	−26.48	−28.79	15.09
J1535 + 4836	107.81	0.1	111.14	−26.46	−29.04	321.50
J1540 + 4138	18.28	0.3	18.77	−26.28	−29.45	36.22
J1625 + 2646	9.69	0.1	9.95	−26.37	−29.08	23.11
J1651 + 4002	43.90	0.5	43.96	−26.28	−29.28	83.07
J2338 + 1504	45.70	0.5	43.75	−26.36	−29.15	101.30

^aIntegrated flux density for source in mJy.^bDistance from search position to source in arcsec.^cObserved radio flux density at rest-frame 6 cm in mJy.^dObserved optical flux density at the rest frame 2500 Å in $\log \mathrm{erg\ s^{-1}\ cm^{-2}\ Hz^{-1}}$.^eAbsolute magnitude at rest-frame 2500 Å calculated from f_{2500} .^fRadio loudness defined as $R = f_{6\mathrm{cm}}/f_{2500}$.

APPENDIX B: ABSOLUTE MAGNITUDE RELATIONS

Below, we provide some useful relations between absolute magnitudes at various wavelengths. $M_i(z = 2)$ is the mean absolute i -band magnitude normalized at $z = 2$ (Shen et al. 2011), M_{1450} and M_{912} are the absolute magnitudes at 1450 and 912 Å, respectively, estimated from our WFC3 stack.

$$M_i(z = 2) - M_{1450} = -1.28 \quad (\mathrm{B1})$$

$$M_{912} - M_{1450} = 0.33. \quad (\mathrm{B2})$$

All these absolute magnitudes are estimated from our average WFC3 spectrum rather accurately. We finally note that, typically, the ionizing luminosity is modelled as a power law, thus our esti-

Table A3. GALEX-forced photometry of the WFC3 quasar sample.

Name	z	NUV ^a	NUV ^b _{IVAR}	FUV ^a	FUV ^b _{IVAR}
J0751 + 4245	2.453	0	0	0	0
J0755 + 2204	2.319	10.702	19.899	−0.463	29.254
J0806 + 5041	2.457	2.085	0.001	0.052	0.679
J0833 + 0815	2.581	0.067	0.649	0.291	10.986
J0845 + 0722	2.307	9.277	0.155	−0.103	41.101
J0850 + 5636	2.464	0.365	19.704	0.673	9.939
J0853 + 4456	2.54	0.248	16.721	−0.094	27.238
J0854 + 5327	2.418	0.734	15.757	−0.444	2.087
J0909 + 0415	2.444	−0.151	23.423	0.151	9.292
J0949 + 0522	2.316	0.351	18.34	−0.342	112.288
J1005 + 5705	2.308	0.25	0.713	−0.313	61.293
J1011 + 0312	2.458	0.252	1.919	0.028	21.285
J1053 + 4007	2.482	5.892	0.441	0.178	10.585
J1104 + 0246	2.532	2.159	0.032	−0.011	2453.64
J1107 + 6420	2.316	81.95	4.905×10^{-5}	50.28	3.828×10^{-5}
J1119 + 1302	2.394	0.475	35.985	0.044	51.255
J1135 + 4607	2.496	0.355	4.004	0.486	12.893
J1143 + 0524	2.561	9.468	644.381	0.008	5527.879
J1215 + 4248	2.31	3.616	4.019	0.015	3909.82
J1220 + 4608	2.446	12.665	1.711	14.278	0.894
J1228 + 5107	2.45	23.419	1.31	−0.002	33.744
J1235 + 6301	2.384	106.753	0.107	0.133	35.543
J1248 + 5809	2.599	1.221	0.001	−0.023	287.568
J1253 + 0516	2.398	−0.158	67.674	0.146	43.164
J1259 + 6720	2.443	−0.274	4.891	0.427	7.997
J1300 + 0556	2.446	0.764	1461.025	0.031	5611.715
J1302 + 0254	2.414	1.006	10.751	0.018	85.795
J1311 + 4531	2.403	3.621	1.849	−0.174	7.709
J1318 + 5312	2.321	0.372	29.747	0.126	31.153
J1323 + 4149	2.44	9.324	0.009	5.766	0.022
J1325 + 6634	2.511	0.118	0.121	−0.089	499.471
J1334 + 0355	2.583	1.219	0.031	−0.248	112.13
J1335 + 4542	2.452	1.309	13.318	0.459	11.939
J1335 + 4637	2.474	2.82	4.541	0.583	6.58
J1336 + 0157	2.379	0.378	7.184	0.111	15.703
J1342 + 6015	2.399	−0.158	35.309	0.077	18.418
J1354 + 5421	2.294	0	0	0	0
J1354 + 0020	2.504	0.134	9.851	−0.055	43.665
J1358 + 0505	2.455	0.125	4328.555	−0.013	5909.892
J1400 + 6430	2.359	7.501	0.169	2.517	0.112
J1415 + 3706	2.374	3.522	1.408	−0.067	497.769
J1454 + 0324	2.368	2.602	11.792	−0.168	126.83
J1533 + 3843	2.529	0.108	30.193	0.051	22.415
J1535 + 4836	2.542	1.847	5.627	0.239	1.424
J1540 + 4138	2.516	0	0	0	0
J1610 + 4423	2.588	1.151	6.765	−0.243	13.434
J1625 + 2943	2.357	−0.054	12.542	0.233	1.771
J1625 + 2646	2.518	57.894	0.382	0.267	15.956
J1651 + 4002	2.343	0.184	11.125	0.189	108.886
J1724 + 5314	2.547	0.608	63.628	−0.206	230.608
J2111 + 0024	2.325	0	0	0	0
J2136 + 1029	2.555	0.276	0.643	0.137	6.969
J2338 + 1504	2.419	4.765	0.101	0.435	0.234

^aFluxes in AB nanomaggies (monochromatic AB fluxes are defined as $F_{\mathrm{AB}} = F_{\mathrm{AB}, \mathrm{nano}} \times 10^{-71.1 \times 0.4}$).^bPipeline inverse variance (IVAR) in AB nanomaggies.

mated $\alpha_{\mathrm{EUV}} = -1.7$ can be used to evaluate the ionizing luminosity as

$$L_{\nu} = L_{912} \left(\frac{\nu}{\nu_{912}} \right)^{\alpha_{\mathrm{EUV}}}, \quad (\mathrm{B3})$$

where the normalization, L_{912} , is well determined.

SUPPORTING INFORMATION

Additional Supporting Information may be found in the online version of this article:

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