

A Comprehensive Consistency Check between Synchrotron radiation and the Observed Gamma-ray Burst Spectra

DAO-ZHOU WANG,¹ XIAO-HONG ZHAO,^{2,3} ZHAO ZHANG,⁴ BIN-BIN ZHANG,^{4,5,6} AND ZHAO-YANG PENG¹

¹College of Physics and Electronics, Yunnan Normal University, Kunming 650500, China

²Yunnan Observatories, Chinese Academy of Sciences, Kunming, China

³Center for Astronomical Mega-Science, Chinese Academy of Sciences, Beijing, China

⁴School of Astronomy and Space Science, Nanjing University, Nanjing 210093, China

⁵Key Laboratory of Modern Astronomy and Astrophysics (Nanjing University), Ministry of Education, China

⁶Department of Physics and Astronomy, University of Nevada Las Vegas, NV 89154, USA

ABSTRACT

We performed a time-resolved spectral analysis of 53 bright gamma-ray bursts (GRBs) observed by *Fermi*/GBM. Our sample consists of 908 individual spectra extracted from the finest time slices in each GRB. We fitted them with the synchrotron radiation model by considering the electron distributions in five different cases: mono-energetic, single power-law, Maxwellian, traditional fast cooling, and broken power-law. Our results were further qualified through Bayesian Information Criterion (BIC) by comparing with the fit by empirical models, namely the so-called Band function and cut-off power-law models. Our study showed that the synchrotron models, except for the fast-cooling case, can successfully fit most observed spectra, with the single power-law case being the most preferred. We also found that the electron distribution indices for the single power-law synchrotron fit in more than half of our spectra exhibits flux-tracking behavior, i.e., the index increases/decreases with the flux increasing/decreasing, implying that the distribution of the radiating electrons is increasingly narrower with time before the flux peaks and becomes more spreading afterward. Our results indicate that the synchrotron radiation is still feasible as a radiation mechanism of the GRB prompt emission phase.

Keywords: gamma-ray burst: general — radiation mechanisms: non-thermal — methods: spectral fitting

1. INTRODUCTION

The radiation mechanism of the sub-MeV emission in gamma-ray bursts (GRBs) is still an unsolved issue, despite extensively studied in the past decades. One helpful tool to confront observed data to physical radiation models is the spectral fit. It has been widely accepted that GRB spectra can be modeled using some empirical models, such as the Band function (BAND, Band et al. 1993), cut-off power-law (CPL), or smoothly broken power-law models (SBPL, Preece et al. 2000; Kaneko et al. 2006; Gruber et al. 2014). Those models, however, are purely mathematical functions and possess no direct physical meanings. From the theoretical point

of view, the observed GRB photons can be generated either through thermal or non-thermal radiation, both directions having been investigated with intensive efforts by, e.g., relating the thermal-like spectra to the photospherical radiation of the fireball (Mészáros & Rees 2000; Mészáros et al. 2002; Pe’er et al. 2006) and non-thermal spectra to synchrotron radiation of Poynting flux in a large emission radius or internal shocks (e.g., Zhang et al. 2016). However, the measured low-energy photon indices through the BAND fit to the observed spectra peak at around ~ -1 (Preece et al. 2000; Kaneko et al. 2006; Poolakkil et al. 2021), which is inconsistent with any segment of the theoretical synchrotron spectrum and also deviates from the photosphere model.

In the internal shock model, the expected low-energy photon slope is $\sim -3/2$, different from the observed -1 (Poolakkil et al. 2021). This is so-called fast cooling problem (Ghisellini et al. 2000). On the other hand,

the “death-line” limit of the low-energy index of the synchrotron model is $-2/3$, which has been violated by some GRBs (Crider et al. 1998; Preece et al. 1998, 2002). Mounting efforts have been made in order to solve these problems. For example, additional physical effects such as Inverse Compton (IC) cooling in Klein-Nishina (KN) regime is considered in Derishev et al. (2001); Daigne et al. (2011); Wang et al. (2009); Nakar et al. (2009). A decaying magnetic field in the emission region is introduced in Pe’er & Zhang (2006); Zhao et al. (2014); Uhm & Zhang (2014); Geng et al. (2018). Liu et al. (2020, 2021) modified the injection rate of electrons from a usually-used constant form to a flexible time-dependent function to phenomenally convey the observed spectral shapes. These models all resulted in flatter electron distributions than that in the fast-cooling case and better explained the observed data. Recently, Burgess et al. 2020 claimed that a synchrotron model with a flexible synchrotron cooling break can accommodate the observed spectra of bursts from a large sample and that the model even overcomes the synchrotron “death-line” problem. In all the above models, the electron distribution is a crucial factor that determines the shape of photon spectrum and diagnoses the physics of particle acceleration and radiation.

In this paper, we first performed an extensive study on nearly all possible forms of electron distributions and derived the consequent synchrotron spectra models (§2). Then, by processing a large sample of brightest GRBs from *Fermi* Gamma-ray Burst Monitor (GBM) data (§3), we fit the observed time-resolved spectra with our models, present the model comparisons, analyze the parameter distributions and their correlations of the best models (§4). Finally, we summarize and discuss the implications of our results in §5.

2. THE MODELS

We consider seven different models including five synchrotron models from various electron distributions and two empirical models, as listed below.

1. Synchrotron model with a power-law electron distribution (SYNPL). The electron energy distribution follows:

$$\frac{dn_e(\gamma_e)}{d\gamma_e} \propto \gamma_e^{-p} \quad (1)$$

where p is the power-law index.

The observed flux density of the synchrotron radiation, F_ν [$\text{erg s}^{-1} \text{cm}^{-2} \text{Hz}^{-1}$] can be calculated as

$$F_\nu = C \int_{\gamma_m}^{\gamma_{max}} \frac{dn_e(\gamma_e)}{d\gamma_e} P(\nu, \gamma_e) d\gamma_e, \quad (2)$$

where C is a normalization factor and γ_m and γ_{max} are the minimum and maximum injection energies of the electron population, respectively. γ_{max} is fixed to 10^7 in our analysis. The synchrotron emissivity averaged over an isotropic distribution of pitch angles scales, $P(\nu, \gamma_e)$, is given as (Crusius & Schlickeiser 1986):

$$P(\nu, \gamma_e) \propto t^2 \left\{ K_{4/3}(t) K_{1/3}(t) - \frac{3}{5} t \left[K_{4/3}^2(t) - K_{1/3}^2(t) \right] \right\}, \quad (3)$$

where $t = \nu / (3\gamma_e^2 \nu_L)$, $\nu_L = eB / (2\pi m_e c)$, γ_e is the electron energy, ν is the synchrotron photon frequency, B is the magnetic field and $K_n(t)$ is the modified Bessel function of order n . The magnetic field is fixed to be $B = 5000$ G throughout this paper, since it does not change the shape of synchrotron spectrum.

For any given free parameter set $[C, \gamma_m, p]$, one can calculate the observed flux via Eq. (2), which can be further compared with observed data.

2. Synchrotron model with a broken power-law electron distribution (SYNBPL). The electron energy distribution follows:

$$\frac{dn_e(\gamma_e)}{d\gamma_e} \propto \begin{cases} \gamma_e^{-p_1}, & \gamma_m < \gamma_e < \gamma_b, \\ \gamma_b^{p_2-1} \gamma_e^{-p_2}, & \gamma_b < \gamma_e < \gamma_{max}, \end{cases} \quad (4)$$

where γ_b is the break energy of electrons. The index, p_1 , is set as a free parameter in our model, which can flexibly account for the effects of several cases such as IC cooling in Klein-Nishina (KN) regime (e.g. Wang et al. 2009; Nakar et al. 2009), magnetic field decay (e.g., Zhao et al. 2014; Uhm & Zhang 2014), or a time-dependent injection rate (Liu et al. 2020). Note that $p_1 = 2$ corresponds to the fast cooling case, while $p_1 = p_2 - 1$ corresponds to the slow cooling case. The SYNBPL model actually includes the SYNPL case, provided $p_1 < 1/3$ or $p_2 \gg p_1$. Note that in the traditional fast cooling model of GRBs in the prompt phase, the magnetic field in GRB emission region is as high as the equipartition field so that the $1/3$ section should be out of the gamma-ray bands toward low energies. However, recently, as shown in Burgess et al. (2020), the $1/3$ section can fall into gamma-ray bands, in the case that the magnetic field is much lower than the equipartition field. Our SYNBPL model also includes the case discussed in Burgess et al. (2020) by changing γ_b or γ_m .

Similar to case 1, for any given free parameter set $[C, \gamma_m, \gamma_b, p_1, p_2]$, one can substitute Eq. (4) to Eq. (2) to derive the observed flux.

3. Synchrotron model with the fast cooling electron distribution (SYNFC). The distribution is:

$$\frac{dn_e(\gamma_e)}{d\gamma_e} \propto \begin{cases} \gamma_e^{-2}, & \gamma_e < \gamma_m \\ \gamma_e^{-(p+1)}, & \text{otherwise,} \end{cases} \quad (5)$$

For any given free parameter set $[C, \gamma_m, p]$, one can substitute Eq. (5) to Eq. (2) to derive the observed flux.

4. Synchrotron model with a Maxwellian electron distribution (SYNMAX). The distribution is

$$n_e(\gamma_e) \propto \left(\frac{\gamma_e}{\gamma_{th}}\right)^2 \exp\left(-\frac{\gamma_e}{\gamma_{th}}\right) \quad (6)$$

where γ_{th} is the dimensionless electron temperature. Such a distribution has been found in a long-term particle-in-cell simulation of an unmagnetized shock and even carries most of the shock energy (Spitkovsky 2008; Giannios & Spitkovsky 2009).

For any given free parameter set $[C, \gamma_{th}]$, one can substitute Eq. (6) to Eq. (2) to derive the observed flux.

5. Synchrotron model with mono-energetic electron distributions. In this model, the energy of all the electrons are assumed to be the same, following a distribution of

$$\frac{dn_e(\gamma_e)}{d\gamma_e} \propto \begin{cases} C, & \gamma_e = \gamma_m \\ 0, & \text{otherwise,} \end{cases} \quad (7)$$

where C is a normalization factor and γ_m is the Lorentz factor of all the electrons.

For any given free parameter set $[C, \gamma_m]$, one can substitute Eq. (7) to Eq. (2) to derive the observed flux.

6. Band function model (BAND). The widely-used empirical model introduced by Band et al. 1993 follows

$$f_{\text{Band}}(E) = A \begin{cases} \left(\frac{E}{100\text{keV}}\right)^\alpha \exp\left[-\frac{(\alpha+2)E}{E_p}\right], & E < E_c, \\ \left(\frac{E}{100\text{keV}}\right)^\beta \exp(\beta - \alpha) \left(\frac{E_c}{100\text{keV}}\right)^{\alpha-\beta}, & E \geq E_c \end{cases} \quad (8)$$

where A is the normalization factor at 100 keV in units of $\text{ph s}^{-1} \text{ cm}^{-2} \text{ keV}^{-1}$, α and β are the low- and high-energy power-law photon indices, respectively, E_p is the peak energy in units of keV, and

$E_c(E_c = \left(\frac{\alpha-\beta}{\alpha+2}\right) E_p)$ is the characteristic energy in units of keV.

One can calculate the flux giving any free parameter set of $[A, \alpha, \beta, E_p]$.

7. Cut-off Power Law model (CPL). Another empirical model follows

$$f_{\text{CPL}}(E) = A \left(\frac{E}{100\text{keV}}\right)^\alpha \exp\left[-\frac{(\alpha+2)E}{E_p}\right], \quad (9)$$

where A is the normalization factor at 100 keV in units of $\text{ph s}^{-1} \text{ cm}^{-2} \text{ keV}^{-1}$, α is the low-energy power-law photon index, E_p is the peak energy in units of keV.

One can calculate the flux giving any free parameter set of $[A, \alpha, E_p]$.

Fig 1 depicts the synchrotron spectrum (F_ν) with the models 1-5. In Section 3, we will employ those models as well as the two empirical models (6-7) to fit the observed spectra of a large sample of bright GRBs.

3. FIT AND RESULTS

The *Fermi* GBM has been triggered by more than three thousand GRBs to date (Poolakkil et al. 2021). We used the data observed by the GBM, which consists of twelve sodium iodide (NaI, 8 keV to 1 MeV) detectors and two bismuth germanium oxide (BGO, 200 keV to 40 MeV) detectors and covers an energy range from 8 keV to 40 MeV (Bissaldi et al. 2009; Meegan et al. 2009). In order to avoid the K-edge¹ at 33.17 keV in spectrum analysis, we use energy spectrum analysis range of 8 keV \sim 30 keV and 40 keV \sim 900 keV for NaI detectors, and use 250 keV \sim 40 MeV for BGOs. We use the Time-Tagged Events (TTE) data with the highest temporal and spectral resolutions for spectrum analysis. Following Yu et al. (2016), we select 53 long bright GRBs (listed in Table 1) from Fermi Science Support Center (FSSC) according to the following criteria: the energy fluence f and peak photon flux F_p in 10 keV \sim 1 MeV bands with the time bin size of 64, 256, and 1024 ms are greater than $4 \times 10^{-5} \text{ erg cm}^{-2}$ and $20 \text{ ph s}^{-1} \text{ cm}^{-2}$, respectively. Data from the two brightest NaI detectors and the brightest BGO detector are selected. The three detectors used in each GRB, the spectrum extraction range, and the number of spectra are shown in Table 1. To study the spectral evolution within a GRB, we divide the whole burst into finest slices allowed by statistics using the same criteria in Yang et al. (2020) and

¹ https://fermi.gsfc.nasa.gov/ssc/data/analysis/GBM_caveats.html

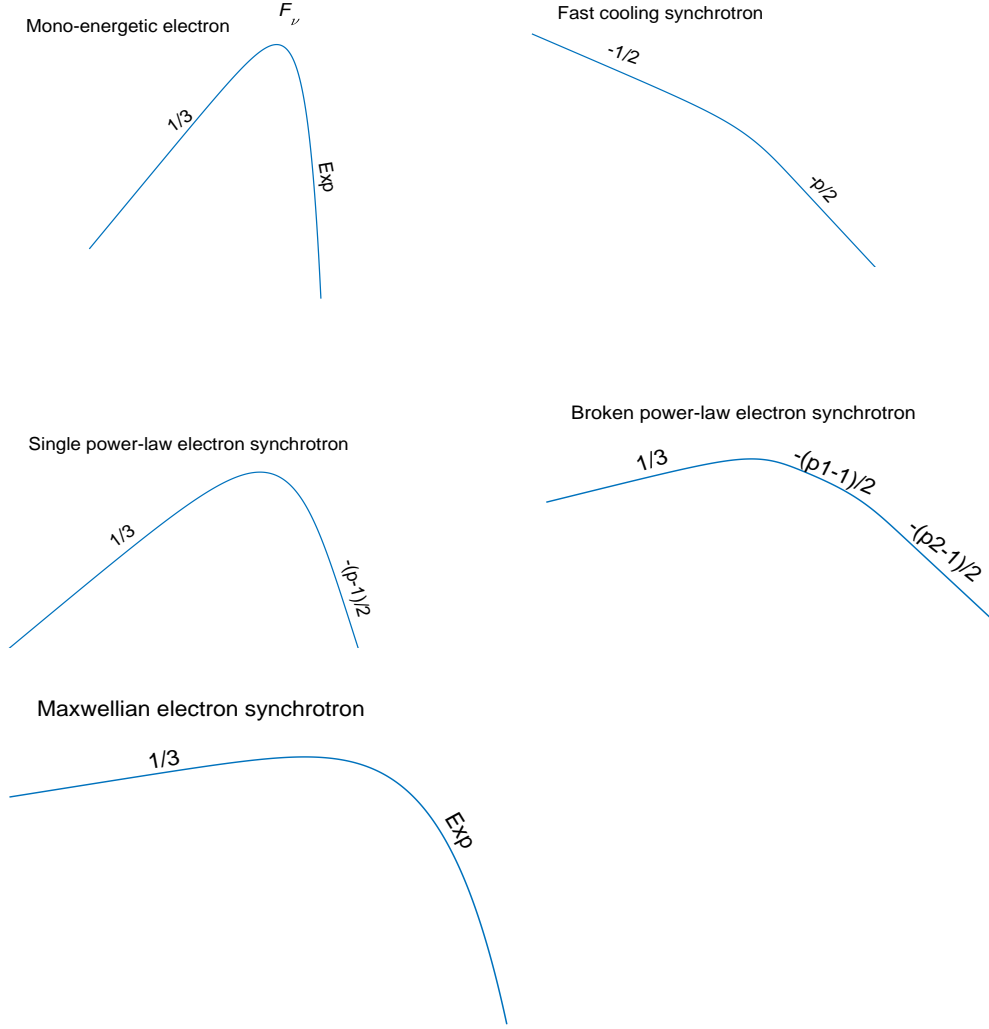


Figure 1. Schematic sketch of the F_ν synchrotron spectra from various electron distributions, including mono-energetic, fast-cooling, single power-law, broken power-law and Maxwellian.

Yang et al. (2021), as listed in Table 1. For each slice, we extract source spectra, background spectra, and corresponding response files following the standard procedure described in Zhang et al. (2016, 2018). Each set of the spectra in each slice are then fitted by the aforementioned seven models using our own spectral fitting package McSpecFit.

An example of our spectral fit of seven models to one slice of data in GRB 130821674 is shown in Fig. 2, where one can see the fast cooling model fit presents a significant deviation in the low energies. The corresponding

parameter constraints are shown in Fig. 3. The best model in each slice is chosen according to the Bayesian Information Criteria (BIC). Table 2 shows the spectral fit for all the slices of spectra of a burst, GRB 130121835. For each of the 908 spectra, we calculate the BIC difference between any two model fits. Finally, we plot the distributions of the difference in Fig. 4. Our results suggest that the SYNPL model is overall the most preferred one in our sample. The distributions of parameters of the models are shown in Fig. 5. The median values and standard deviations of the distribution of these parameters are summarized in Table 3.

Table 1. GRBs, detectors, time slices and the number of spectra selected in the data fitting.

GRB	Detectors	ΔT_{source}	Spectral number	GRB	Detectors	ΔT_{source}	Spectral number
130121835	n4 n8 b1	0.00-13.00	15	160215773	n4 n5 b0	168.00-200.00	15
130305486	n6 n9 b1	2.55-14.46	12	160530667	n1 n2 b0	4.25-9.36	14
130327350	n0 n1 b0	3.75-30.75	10	160720767	n2 n5 b0	78.25-100.41	31
130502327	n6 n7 b1	9.65-30.85	17	160905471	n6 n9 b1	7.15-23.79	22
130518580	n3 n4 b0	20.25-32.97	13	160910722	n1 n5 b0	7.5514.22	13
130606497	n8 nb b1	8.95-52.85	20	170114917	n1 n2 b0	1.05-31.85	5
130609902	n4 n5 b0	5.85-23.15	15	170115743	n0 n1 b0	1.05-31.85	5
130614997	n0 n3 b0	0.23-4.65	5	170210116	n2 na b0	38.05-69.85	6
130815660	n3 n4 b0	30.00-42.00	8	170308221	n9 na b1	7.45-13.85	10
130821674	n6 n9 b1	10.75-45.95	17	170409112	n2 n5 b0	31.85-67.35	27
130925173	n6 n7 b1	25.05-167.85	38	170522657	n1 n2 b0	1.25-5.68	5
131014215	n9 nb b1	1.33-3.97	23	170527480	n3 n6 b0	2.95-27.55	15
131108862	n0 n3 b0	0.85-16.75	14	170826819	na nb b1	2.15-10.22	11
131127592	n2 n5 b0	0.65-7.34	7	170906030	n3 n4 b0	21.65-86.05	14
131231198	n0 n3 b0	18.65-39.45	30	171210493	n0 n1 b0	3.65-18.45	22
140206275	n0 n1 b0	8.00-35.86	39	171227000	n3 n5 b0	15.25-28.35	27
140306146	n3 n4 b0	2.25-32.45	10	180113418	n2 n9 b1	8.35-27.25	25
140723499	n9 na b1	24.65-37.55	18	180210517	n0 n1 b0	6.25-30.90	26
141028455	n6 n9 b1	8.45-22.65	15	180305393	n1 n2 b0	2.65-8.85	12
150105257	n8 nb b1	12.15-69.65	25	180720598	n6 n7 b1	6.65-19.64	32
150213001	n7 n8 b1	0.85-3.69	6	190114873	n3 n4 b0	0.00-6.00	18
150306993	n4 n5 b0	0.00-20.00	12	190530430	n1 n2 b0	8.00-19.82	38
150314205	n1 n9 b1	0.00-12.00	15	190531840	n0 n3 b0	13.55-35.55	30
150510139	n1 n5 b0	0.00-24.00	12	190731943	n6 n9 b1	2.25-13.82	14
150902733	n0 n3 b0	4.95-15.45	16	191227069	n2 n5 b0	12.15-26.81	19
151021791	n9 na b1	0.00-10.00	6	200412381	n7 n8 b1	7.63-10.94	13
160113398	n8 nb b1	29.65-45.90	21				

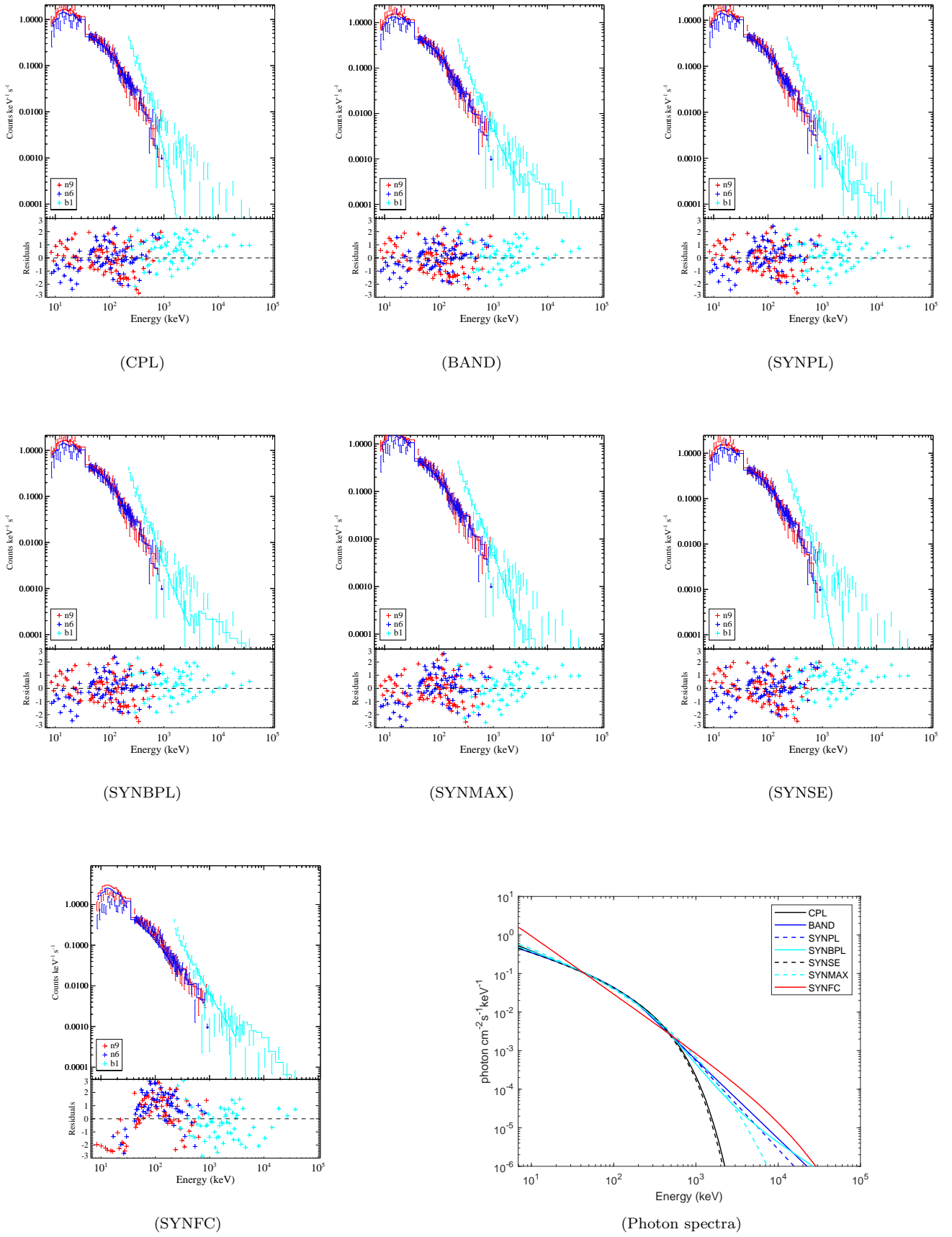


Figure 2. Comparison of observed count spectra with the fitted empirical models (up-left: CPL, up-middle: BAND) and synchrotron models (up-right: SYNPL; middle: SYNBPL, SYNMAX and SYNSE; bottom-left: SYNFC) for the time slice 8 of GRB 130821674. Bottom-right: comparison between the spectra of the seven models

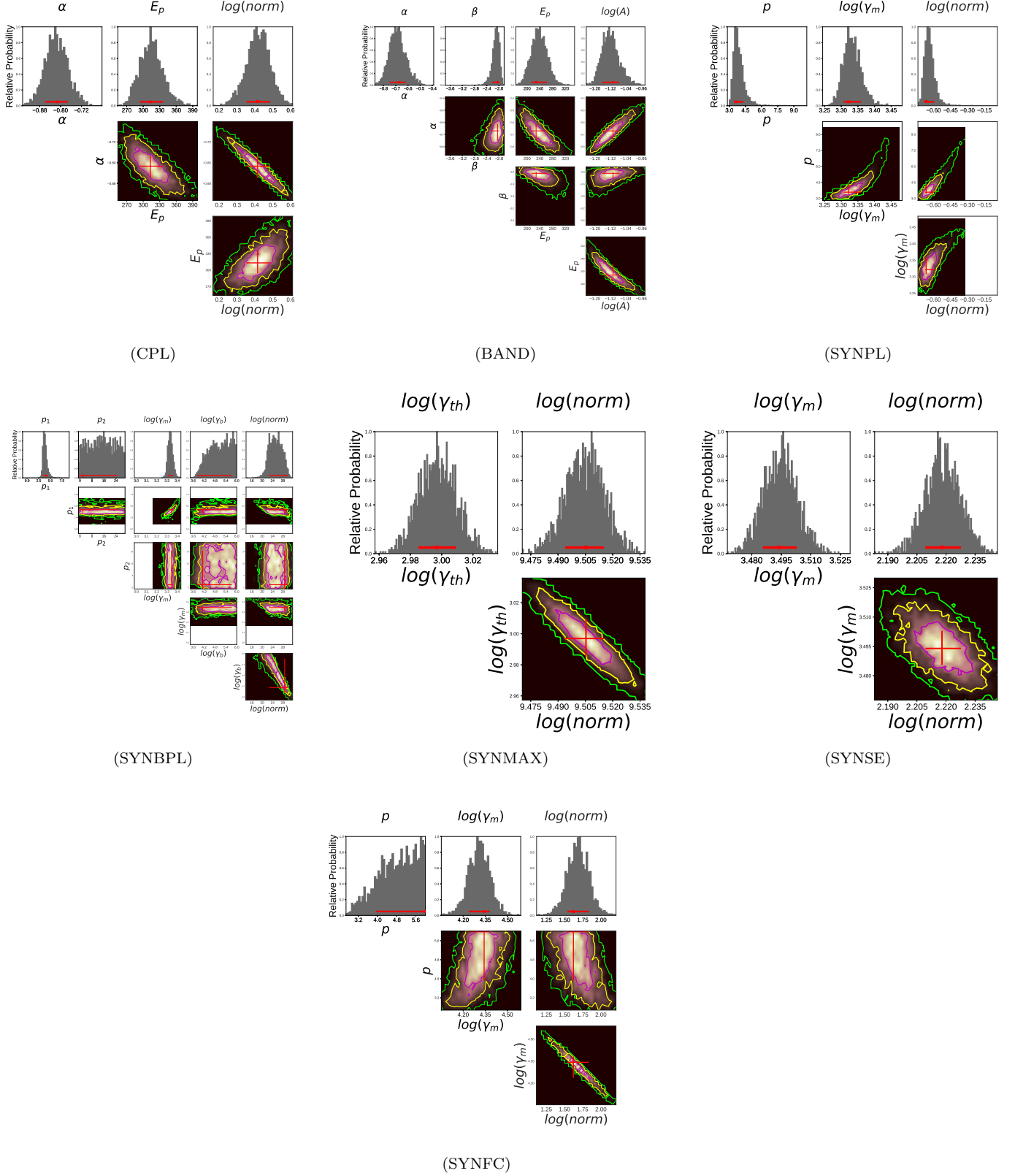


Figure 3. Parameter constraints of the empirical models (up-left: CPL, up-middle: BAND) and synchrotron models (up-right: SYNPL; middle: SYNBPL, SYNMAX and SYNSE; bottom: SYNFC) for the fit to the time slice 8 of GRB 130821674. Histograms and contours show the likelihood map of the parameter-constraint outputs by the McSpecFit package. Red crosses mark the best-fit values and 1σ error bars.

Table 2. Time-resolved spectral analysis results of GRB 130121835

$t_{start}:t_{end}$	Models	α/p	$E_p/\log(\gamma_m)/\log(\gamma_{th})$	PGSTAT/dof	BIC	$F \times 10^{-6}$	Models	α/p_1	β/p_2	$\log(\gamma_m)$	$E_p/\log(\gamma_b)$	PGSTAT/dof	BIC	$F \times 10^{-6}$
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)	(13)	(14)	(15)
0.00:0.87	CPL	$-1.11^{+0.08}_{-0.04}$	$2990.84^{+9.05}_{-410.86}$	282.4/349	300.01	$6.39^{+1.20}_{-0.99}$	BAND	$-1.11^{+0.07}_{-0.05}$	$-6.26^{+0.79}_{-11.01}$	—	$2999.99^{+15.31}_{-18.29}$	282.4/348	305.84	$6.43^{+1.10}_{-1.05}$
	SYNPL	$1.88^{+23.4}_{-0.07}$	$2.97^{+0.98}_{-0.06}$	308.8/349	326.42	$6.23^{+3.04}_{-1.58}$	SYNBPL	$1.13^{+0.17}_{-0.24}$	$28.41^{+1.59}_{-19.22}$	$2.45^{+0.22}_{-1.14}$	$4.28^{+0.07}_{-0.09}$	283.7/347	313.03	$7.09^{+1.73}_{-1.32}$
	SYNSE	—	$3.96^{+0.02}_{-0.04}$	316.7/350	328.42	$8.00^{+1.77}_{-1.63}$	SYNFC	—	$1.80^{+0.06}_{-0.05}$	$1.69^{+1.20}_{-0.39}$	—	314.9/349	332.45	$6.59^{+0.81}_{-0.86}$
	SYNMAX	—	$3.37^{+0.04}_{-0.06}$	302.1/350	313.87	$6.77^{+1.45}_{-1.47}$								
0.87:1.73	CPL	$-1.15^{+0.05}_{-0.03}$	$2985.91^{+13.75}_{-635.28}$	300.8/349	318.37	$8.41^{+1.99}_{-0.5}$	BAND	$-1.15^{+0.05}_{-0.03}$	$-4.86^{+2.79}_{-12.55}$	—	$2981.22^{+15.31}_{-18.29}$	300.8/348	324.22	$8.30^{+1.00}_{-1.38}$
	SYNPL	$2.09^{+0.13}_{-0.11}$	$3.16^{+0.03}_{-0.07}$	328.0/349	345.61	$7.70^{+1.38}_{-1.33}$	SYNBPL	$1.10^{+0.18}_{-0.16}$	$29.04^{+0.59}_{-20.17}$	$1.65^{+0.98}_{-0.32}$	$4.21^{+0.09}_{-0.10}$	303.8/347	333.14	$8.53^{+1.78}_{-1.78}$
	SYNSE	—	$3.77^{+0.04}_{-0.04}$	356.5/350	368.22	$5.05^{+1.55}_{-1.13}$	SYNFC	—	$1.80^{+0.05}_{-0.04}$	$1.67^{+1.04}_{-0.35}$	—	348.6/349	366.21	$9.05^{+0.87}_{-0.90}$
	SYNMAX	—	$3.26^{+0.04}_{-0.04}$	323.4/350	335.11	$6.72^{+1.35}_{-1.15}$								
1.73:2.60	CPL	$-1.07^{+0.04}_{-0.03}$	$2602.41^{+213.82}_{-345.82}$	316.3/349	333.86	$11.50^{+1.50}_{-1.60}$	BAND	$-1.07^{+0.04}_{-0.03}$	$-8.42^{+2.24}_{-8.78}$	—	$2641.91^{+15.31}_{-18.29}$	316.3/348	339.75	$11.60^{+1.40}_{-1.80}$
	SYNPL	$4.66^{+12.92}_{-2.17}$	$3.65^{+0.15}_{-0.01}$	368.7/349	386.33	$8.44^{+2.36}_{-1.71}$	SYNBPL	$0.80^{+0.10}_{-0.27}$	$26.66^{+0.03}_{-15.42}$	$1.97^{+0.63}_{-0.67}$	$4.14^{+0.02}_{-0.08}$	322.9/347	352.20	$11.40^{+1.60}_{-1.90}$
	SYNSE	—	$3.83^{+0.02}_{-0.03}$	376.1/350	387.81	$8.82^{+1.65}_{-1.45}$	SYNFC	—	$1.80^{+0.03}_{-0.03}$	$1.43^{+1.15}_{-0.14}$	—	459.9/349	477.53	$12.10^{+0.90}_{-0.90}$
	SYNMAX	—	$3.29^{+0.03}_{-0.03}$	339.1/350	350.80	$10.00^{+1.30}_{-1.20}$								
2.60:3.47	CPL	$-1.09^{+0.03}_{-0.04}$	$1432.36^{+260.78}_{-175.05}$	346.9/349	364.50	$11.80^{+2.10}_{-1.70}$	BAND	$-1.09^{+0.04}_{-0.03}$	$-9.90^{+3.75}_{-7.49}$	—	$1452.60^{+15.31}_{-18.29}$	346.9/348	370.37	$11.90^{+2.00}_{-1.90}$
	SYNPL	$3.30^{+0.31}_{-0.22}$	$3.43^{+0.04}_{-0.02}$	367.2/349	384.78	$10.70^{+1.80}_{-1.30}$	SYNBPL	$0.90^{+0.07}_{-0.39}$	$28.96^{+1.02}_{-18.38}$	$3.04^{+0.03}_{-0.67}$	$4.02^{+0.01}_{-0.09}$	342.1/347	371.42	$12.10^{+2.20}_{-1.80}$
	SYNSE	—	$3.65^{+0.01}_{-0.01}$	402.4/350	413.75	$7.27^{+0.66}_{-0.54}$	SYNFC	—	$5.83^{+0.17}_{-1.93}$	$4.43^{+0.09}_{-0.06}$	—	632.6/349	650.15	$12.80^{+0.90}_{-1.00}$
	SYNMAX	—	$3.16^{+0.02}_{-0.02}$	347.8/350	359.55	$11.10^{+1.00}_{-0.90}$								
3.47:4.33	CPL	$-1.12^{+0.04}_{-0.03}$	$1139.18^{+218.03}_{-184.47}$	333.3/349	350.93	$9.88^{+1.92}_{-1.63}$	BAND	$-1.12^{+0.04}_{-0.04}$	$-8.22^{+2.12}_{-9.28}$	—	$1138.01^{+15.31}_{-18.29}$	333.4/348	356.81	$9.92^{+1.81}_{-1.78}$
	SYNPL	$3.22^{+0.36}_{-0.15}$	$3.38^{+0.03}_{-0.03}$	344.1/349	361.72	$9.75^{+1.34}_{-1.53}$	SYNBPL	$1.49^{+0.19}_{-0.73}$	$28.82^{+1.18}_{-18.33}$	$3.16^{+0.05}_{-1.03}$	$4.03^{+0.03}_{-0.12}$	320.5/347	349.82	$10.70^{+2.30}_{-2.30}$
	SYNSE	—	$3.59^{+0.01}_{-0.01}$	383.2/350	394.90	$6.08^{+0.44}_{-0.42}$	SYNFC	—	$5.93^{+0.07}_{-1.52}$	$4.31^{+0.04}_{-0.07}$	—	557.9/349	575.49	$12.00^{+1.00}_{-1.20}$
	SYNMAX	—	$3.10^{+0.02}_{-0.02}$	327.5/350	339.18	$9.19^{+0.81}_{-0.70}$								
4.33:5.20	CPL	$-1.23^{+0.03}_{-0.03}$	$1738.74^{+401.92}_{-280.67}$	363.7/349	381.33	$9.50^{+1.84}_{-1.81}$	BAND	$-1.24^{+0.03}_{-0.03}$	$-15.40^{+9.38}_{-1.92}$	—	$1727.71^{+15.31}_{-18.29}$	363.7/348	387.19	$9.35^{+1.91}_{-1.73}$
	SYNPL	$2.51^{+0.28}_{-0.00}$	$3.22^{+0.06}_{-0.00}$	366.1/349	383.68	$9.59^{+0.79}_{-1.93}$	SYNBPL	$1.78^{+0.19}_{-0.43}$	$27.77^{+2.21}_{-18.17}$	$3.09^{+0.06}_{-0.16}$	$4.15^{+0.04}_{-0.14}$	352.5/347	381.81	$9.41^{+1.97}_{-2.11}$
	SYNSE	—	$3.54^{+0.02}_{-0.01}$	430.1/350	441.87	$4.14^{+0.36}_{-0.32}$	SYNFC	—	$5.95^{+0.04}_{-2.33}$	$4.38^{+0.07}_{-0.13}$	—	463.2/349	480.77	$10.20^{+1.00}_{-1.10}$
	SYNMAX	—	$3.06^{+0.02}_{-0.02}$	372.1/350	383.79	$6.45^{+0.72}_{-0.58}$								
5.20:6.07	CPL	$-1.30^{+0.03}_{-0.02}$	$2300.42^{+397.88}_{-414.19}$	382.7/349	400.34	$11.30^{+1.70}_{-2.0}$	BAND	$-1.29^{+0.02}_{-0.03}$	$-17.34^{+11.77}_{-0.08}$	—	$2215.80^{+15.31}_{-18.29}$	382.8/348	406.24	$11.10^{+1.90}_{-1.70}$
	SYNPL	$2.64^{+0.06}_{-0.16}$	$3.21^{+0.02}_{-0.04}$	371.9/349	389.52	$9.33^{+1.72}_{-0.89}$	SYNBPL	$2.07^{+0.17}_{-0.19}$	$26.80^{+3.20}_{-17.44}$	$3.12^{+0.04}_{-0.06}$	$4.29^{+0.07}_{-0.12}$	359.9/347	389.24	$10.60^{+1.70}_{-1.90}$
	SYNSE	—	$3.51^{+0.02}_{-0.01}$	461.6/350	473.35	$4.10^{+0.33}_{-0.29}$	SYNFC	—	$5.99^{+0.01}_{-1.99}$	$4.38^{+0.04}_{-0.11}$	—	460.3/349	477.94	$11.20^{+0.90}_{-1.10}$
	SYNMAX	—	$3.02^{+0.02}_{-0.04}$	395.2/350	406.94	$6.35^{+0.67}_{-0.55}$								
6.07:6.93	CPL	$-1.26^{+0.03}_{-0.03}$	$1627.65^{+359.52}_{-246.78}$	359.2/349	376.77	$10.40^{+1.90}_{-1.70}$	BAND	$-0.83^{+0.08}_{-0.46}$	$-1.77^{+0.07}_{-15.47}$	—	$256.87^{+15.31}_{-18.29}$	349.8/348	373.24	$10.70^{+1.60}_{-2.40}$
	SYNPL	$2.66^{+0.12}_{-0.12}$	$3.22^{+0.03}_{-0.03}$	346.0/349	363.63	$9.95^{+1.55}_{-1.16}$	SYNBPL	$1.95^{+0.25}_{-0.14}$	$28.43^{+1.56}_{-19.20}$	$3.10^{+0.06}_{-0.03}$	$4.17^{+0.06}_{-0.10}$	332.0/347	361.34	$10.40^{+1.60}_{-1.90}$
	SYNSE	—	$3.52^{+0.01}_{-0.01}$	437.4/350	449.15	$4.52^{+0.32}_{-0.30}$	SYNFC	—	$5.88^{+0.12}_{-1.95}$	$4.31^{+0.08}_{-0.06}$	—	463.4/349	481.01	$11.50^{+1.20}_{-1.00}$
	SYNMAX	—	$3.02^{+0.02}_{-0.02}$	362.1/350	373.78	$6.95^{+0.63}_{-0.58}$								
6.93:7.80	CPL	$-1.16^{+0.04}_{-0.03}$	$1093.38^{+206.64}_{-172.04}$	374.9/349	392.48	$11.70^{+2.0}_{-1.90}$	BAND	$-0.98^{+0.09}_{-0.20}$	$-2.11^{+0.14}_{-14.76}$	—	$481.93^{+15.31}_{-18.29}$	371.3/348	394.72	$11.70^{+2.0}_{-1.90}$

Table 2 continued

Table 2 (continued)

$t_{start}:t_{end}$	Models	α/p	$E_p/\log(\gamma_m)/\log(\gamma_{th})$	PGSTAT/dof	BIC	$F \times 10^{-6}$	Models	α/p_1	β/p_2	$\log(\gamma_m)$	$E_p/\log(\gamma_b)$	PGSTAT/dof	BIC	$F \times 10^{-6}$
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)	(13)	(14)	(15)
7.80:8.67	SYNPL	$3.09^{+0.17}_{-0.15}$	$3.33^{+0.02}_{-0.03}$	367.6/349	385.15	$12.10^{+1.20}_{-1.70}$	SYNBPL	$1.85^{+0.29}_{-0.20}$	$19.41^{+8.70}_{-8.59}$	$3.18^{+0.06}_{-0.04}$	$4.07^{+0.06}_{-0.07}$	342.3/347	371.57	$12.80^{+2.00}_{-1.90}$
	SYNSE	—	$3.56^{+0.01}_{-0.01}$	438.6/350	450.35	$6.90^{+0.42}_{-0.38}$	SYNFC	—	$6.00^{+0.00}_{-1.35}$	$4.30^{+0.03}_{-0.07}$	—	624.2/349	641.82	$14.80^{+1.10}_{-1.30}$
	SYNMAX	—	$3.07^{+0.02}_{-0.01}$	358.1/350	369.81	$10.40^{+0.80}_{-0.70}$								
	CPL	$-1.38^{+0.04}_{-0.03}$	$1517.33^{+634.99}_{-426.4}$	324.0/349	341.61	$8.01^{+2.14}_{-2.34}$	BAND	$-1.05^{+0.02}_{-0.33}$	$-1.88^{+0.03}_{-9.79}$	—	$245.76^{+15.31}_{-18.29}$	309.3/348	332.80	$7.51^{+2.61}_{-2.32}$
	SYNPL	$2.76^{+0.12}_{-0.14}$	$3.15^{+0.02}_{-0.03}$	311.3/349	328.94	$7.46^{+1.42}_{-1.01}$	SYNBPL	$2.28^{+0.29}_{-0.15}$	$17.23^{+8.70}_{-8.59}$	$3.06^{+0.07}_{-0.04}$	$4.22^{+0.19}_{-0.12}$	304.0/347	333.31	$7.93^{+1.84}_{-1.69}$
	SYNSE	—	$3.44^{+0.01}_{-0.01}$	382.5/350	394.24	$3.32^{+0.22}_{-0.20}$	SYNFC	—	$5.94^{+0.06}_{-1.77}$	$4.20^{+0.03}_{-0.10}$	—	348.0/349	365.63	$9.26^{+1.18}_{-1.18}$
8.67:9.53	SYNMAX	—	$2.92^{+0.02}_{-0.02}$	327.2/350	338.93	$4.71^{+0.39}_{-0.37}$								
	CPL	$-1.42^{+0.03}_{-0.02}$	$2962.80^{+36.41}_{-821.9}$	377.8/349	395.40	$9.21^{+0.78}_{-1.45}$	BAND	$-1.04^{+0.00}_{-0.39}$	$-1.76^{+0.00}_{-15.39}$	—	$215.02^{+15.31}_{-18.29}$	373.2/348	396.61	$7.80^{+2.18}_{-0.24}$
	SYNPL	$2.48^{+0.10}_{-0.01}$	$3.07^{+0.03}_{-0.04}$	372.3/349	389.86	$7.92^{+1.28}_{-1.16}$	SYNBPL	$2.17^{+0.21}_{-0.13}$	$26.35^{+3.63}_{-17.94}$	$3.00^{+0.03}_{-0.06}$	$4.33^{+0.19}_{-0.11}$	365.4/347	394.70	$8.22^{+1.79}_{-1.69}$
	SYNSE	—	$3.41^{+0.02}_{-0.01}$	478.3/350	490.04	$2.70^{+0.20}_{-0.17}$	SYNFC	—	$5.97^{+0.03}_{-2.07}$	$4.28^{+0.04}_{-0.12}$	—	390.7/349	408.28	$9.03^{+1.10}_{-1.19}$
9.53:10.40	SYNMAX	—	$2.90^{+0.02}_{-0.02}$	414.8/350	426.48	$3.98^{+0.37}_{-0.34}$								
	CPL	$-1.47^{+0.04}_{-0.03}$	$2976.46^{+22.16}_{-1036.1}$	354.7/349	372.27	$6.48^{+0.71}_{-1.37}$	BAND	$-1.47^{+0.04}_{-0.03}$	$-9.57^{+4.32}_{-7.54}$	—	$2880.26^{+15.31}_{-18.29}$	354.7/348	378.18	$6.40^{+0.68}_{-1.26}$
	SYNPL	$2.48^{+0.12}_{-0.10}$	$2.99^{+0.05}_{-0.03}$	351.4/349	368.96	$5.70^{+1.10}_{-1.80}$	SYNBPL	$2.29^{+0.19}_{-0.24}$	$20.82^{+5.07}_{-13.54}$	$2.96^{+0.05}_{-0.12}$	$4.34^{+0.76}_{-0.15}$	347.6/347	376.89	$5.63^{+1.58}_{-1.42}$
	SYNSE	—	$3.37^{+0.02}_{-0.02}$	430.9/350	442.66	$1.85^{+0.15}_{-0.14}$	SYNFC	—	$5.98^{+0.02}_{-2.28}$	$4.20^{+0.06}_{-0.15}$	—	356.8/349	374.35	$6.41^{+1.10}_{-1.08}$
10.40:11.27	SYNMAX	—	$2.85^{+0.03}_{-0.02}$	384.2/350	395.93	$2.63^{+0.28}_{-0.26}$								
	CPL	$-1.41^{+0.04}_{-0.03}$	$2983.48^{+15.89}_{-583.05}$	355.9/349	373.52	$5.87^{+0.79}_{-0.87}$	BAND	$-1.41^{+0.04}_{-0.04}$	$-14.53^{+9.99}_{-2.53}$	—	$2998.47^{+15.31}_{-18.29}$	355.9/348	379.37	$5.96^{+1.71}_{-0.95}$
	SYNPL	$2.09^{+0.01}_{-0.1}$	$2.78^{+1.12}_{-0.51}$	356.0/349	373.58	$6.64^{+1.61}_{-0.64}$	SYNBPL	$1.90^{+0.15}_{-0.07}$	$16.74^{+8.54}_{-12.00}$	$2.63^{+0.27}_{-1.34}$	$4.50^{+0.80}_{-0.18}$	353.2/347	382.48	$7.18^{+1.23}_{-1.66}$
	SYNSE	—	$3.44^{+0.02}_{-0.03}$	476.2/350	487.93	$1.70^{+0.23}_{-0.19}$	SYNFC	—	$3.10^{+2.13}_{-0.73}$	$4.39^{+0.81}_{-0.06}$	—	355.1/349	372.67	$6.81^{+0.87}_{-1.05}$
11.27:12.13	SYNMAX	—	$2.92^{+0.04}_{-0.03}$	420.4/350	432.17	$2.50^{+0.44}_{-0.35}$								
	CPL	$-1.40^{+0.05}_{-0.03}$	$2993.46^{+6.24}_{-761.98}$	405.4/349	422.96	$5.95^{+0.78}_{-1.05}$	BAND	$-1.39^{+0.05}_{-0.03}$	$-8.77^{+3.57}_{-8.27}$	—	$2986.05^{+15.31}_{-18.29}$	405.4/348	428.86	$6.01^{+0.78}_{-1.13}$
	SYNPL	$2.24^{+0.05}_{-0.16}$	$2.96^{+0.03}_{-0.21}$	417.8/349	435.42	$5.87^{+1.36}_{-1.12}$	SYNBPL	$1.80^{+0.09}_{-0.11}$	$25.52^{+0.50}_{-16.89}$	$2.75^{+0.17}_{-1.46}$	$4.27^{+0.10}_{-0.10}$	406.6/347	435.92	$6.43^{+1.10}_{-1.05}$
	SYNSE	—	$3.42^{+0.02}_{-0.03}$	504.6/350	516.33	$1.60^{+0.21}_{-0.18}$	SYNFC	—	$6.00^{+0.00}_{-2.65}$	$4.34^{+0.11}_{-0.17}$	—	413.9/349	431.49	$5.89^{+0.90}_{-0.99}$
12.13:13.00	SYNMAX	—	$2.93^{+0.04}_{-0.03}$	456.2/350	467.93	$2.56^{+0.46}_{-0.36}$								
	CPL	$-1.46^{+0.05}_{-0.04}$	$2600.74^{+197.62}_{-750.94}$	310.8/349	328.40	$4.81^{+0.81}_{-0.98}$	BAND	$-1.46^{+0.04}_{-0.04}$	$-17.56^{+11.93}_{-2.44}$	—	$2537.86^{+15.31}_{-18.29}$	310.8/348	334.27	$4.79^{+0.82}_{-0.99}$
	SYNPL	$2.30^{+0.04}_{-0.16}$	$2.88^{+1.04}_{-0.26}$	324.6/349	342.24	$4.78^{+1.41}_{-1.08}$	SYNBPL	$1.93^{+0.09}_{-0.13}$	$21.11^{+5.10}_{-12.08}$	$2.62^{+0.22}_{-1.34}$	$4.19^{+0.10}_{-0.10}$	312.4/347	341.68	$4.57^{+1.38}_{-1.00}$
	SYNSE	—	$3.39^{+0.03}_{-0.02}$	402.8/350	414.50	$1.42^{+0.17}_{-0.15}$	SYNFC	—	$5.86^{+0.14}_{-2.03}$	$4.14^{+0.11}_{-0.10}$	—	314.4/349	332.03	$4.53^{+1.18}_{-0.79}$
	SYNMAX	—	$2.87^{+0.04}_{-0.03}$	355.7/350	367.47	$2.05^{+0.34}_{-0.26}$								

NOTE—We list the start and stop times of each time slice of each burst (Col. 1), fitting models (Cols. 2 and 8), model fitting parameters (Cols. 3 ~ 4 and Cols. 9 ~ 12), PGSTAT/dof (Cols. 5 and 13), Bayesian Information Criteria (BIC) (Cols. 6 and 14) and energy flux (F) in $1\text{keV} \sim 10^4\text{keV}$ (Cols. 7 and 15). All time parameters have units of s, energies have units of keV, and fluxes have units of $\text{erg s}^{-1}\text{cm}^{-2}$.

Table 3. Values of the median and standard deviation of the parameter distributions.

Models	$\alpha/p/p_1$	$E_p/\log(\gamma_b)$	β	$\log(\gamma_m)/\log(\gamma_{th})$
CPL	-0.87 ± 0.37	332.88 ± 40.98	—	—
BAND	-0.78 ± 0.42	278.08 ± 59.41	-2.86 ± 1.21	—
SYNPL	4.91 ± 1.34	—	—	3.38 ± 0.33
SYNSE	—	—	—	3.49 ± 0.21
SYNBPL	2.70 ± 2.78	3.75 ± 0.51	—	3.27 ± 0.68
SYNMAX	—	—	—	2.98 ± 0.25

The distribution of the low energy index, α , which is in the empirical model (such as CPL and BAND) fits, peaks at slightly less than $-2/3$, which is the synchrotron slope below the minimum injected energy without cooling. The peak energy, E_p , of the BAND model, is lower than that of the CPL model (see Table 3). This result is consistent with the statistical results of previous GBM spectral catalogs (e.g., Yu et al. 2016, 2019). The high-energy index β of the BAND spectrum is smaller than the previous spectral catalogs. A small part of the sample's β is very small (-10 to -4), indicating that this part of the energy spectrum is a high-energy cut-off single power-law spectrum.

The distribution of the spectral index p of the SYNPL model has two peaks (see Fig. 5). About 3/4 of the spectral sample peaks at ~ 3.6 , while the left 1/4 peaks at ~ 29.6 , which suggests these spectra actually can be fitted by the SYNSE model. The distribution of p_1 of the SYNBPL model is very wide, varying from ~ 0 to ~ 10 . This may be attributed to the poor constraint on the parameters of the SYNBPL model. The distributions of the minimum injection energy (γ_m) for the SYNSE, SYNPL, and SYNBPL models peak at $\sim 10^{3.4}$, which corresponds to the synchrotron peak of ~ 100 keV, while the peak of γ_b for the SYNBPL model is $10^{3.8}$ corresponding to the synchrotron peak of ~ 400 keV (see Table 3).

The electron power-law index, p , is a critical parameter in the SYNPL model. Interestingly, a correlation between the flux F and the index p is found in our analysis. Using Spearman's rank correlation analysis method, we calculate the linear correlation coefficients (r) and chance probabilities (P) for each burst and evaluate the correlation of $F - p$. The distribution of r is presented in Fig. 6. One can find that the distribution of r is somewhat diverse, but clearly, there is a peak around $r \sim 0.6$. It is $r > 0.4$ for more than a half bursts in our samples, suggesting strong positive correlations between F and p . We list 6 cases in Fig. 7. The evolution of the

electron power-law index traces the light curve for these bursts.

4. SUMMARY AND IMPLICATIONS

In this paper, using several synchrotron models with different electron distributions and two empirical models, a time-resolved spectroscopy analysis is performed for 53 long bright bursts of *Fermi*/GBM, including a total of 908 time-resolved spectra. According to the Bayesian information standard, we compare the fitting results of each model. We find that: (1) most of the GRB energy spectra are well fitted with the synchrotron models, except the SYNFC model, and the SYNPL model is better than the SYNSE, SYNMAX, and SYNBPL models based on the BIC criterion; (2) a strong correlation between $F - p$ is found for the SYNPL model, i.e., the electron power-law index in more than a half bursts traces their flux; (3) the empirical models are comparably good to synchrotron models when fitting to the observed spectra.

Our results support the previous conclusion that the synchrotron mechanism can be responsible for the observed GRB spectra (e.g., Oganessian et al. 2017; Beniamini et al. 2018; Ravasio et al. 2018 and 2019; Burgess et al. 2014 and 2020). However, the traditional fast cooling synchrotron model is not favored in our fits. Our findings that SYNPL model can fit the data suggest that the heated electrons in the emission region is not significantly cooled or γ_c is right around γ_m . These suggest that GRBs may result from magnetic dissipation at much larger radii with a weak magnetic field instead of the internal shock at smaller radii with a large magnetic field (e.g., Kumar et al. 2007; Zhang & Yan 2011).

We note that our model comparison is subject to the caveat that the forward-folding method, which requires the input models to be first convolved with the instrumental response (RSP) matrix before compared with the observed count spectrum, may be incapable of discriminating between various models due to the nonlinear nature of the RSP matrix.

The $F - p$ tracking behavior in Fig. 7 suggests that the electrons are distributed in a much narrower range as the flux approaches its peak. The slope of p at the peak time of flux for some bursts is larger than 5, which actually corresponds to the mono-energetic electron synchrotron spectrum with the electron energy and number concentrated in the minimal injection energy of electrons. The steep electron spectra is inconsistent with the typical slope expected in theoretical studies of the relativistic shock model (e.g., Kirk et al. 2000), where $p \sim 2.3$. Taking cooling into account, the typical theoretical index is $p \sim 3.3$, which is still much smaller than

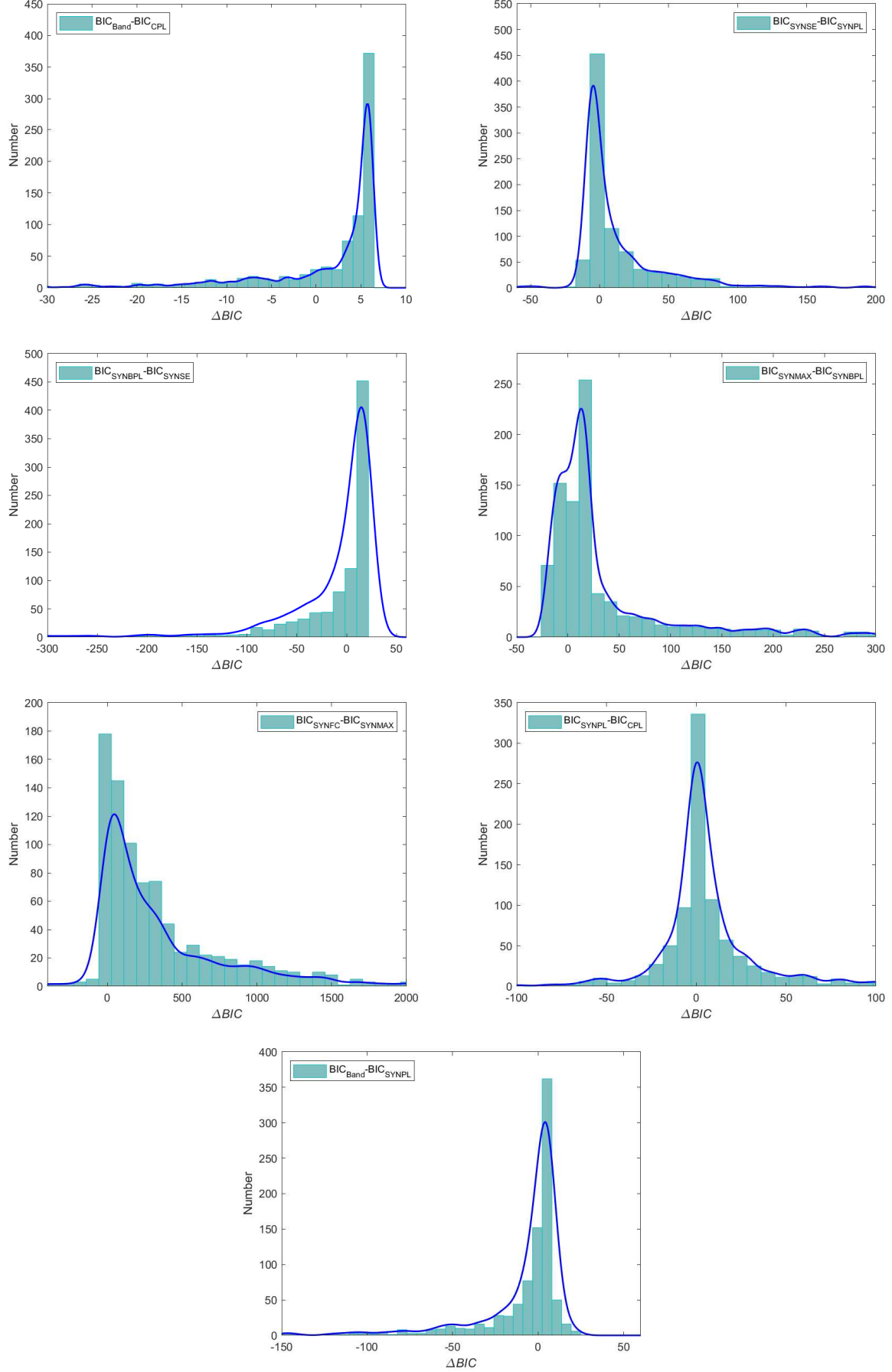


Figure 4. Distributions of the BIC difference between the seven models are shown. The blue curves represent the kernel density estimation (KDE) of the distributions, where Gaussian kernels are used.

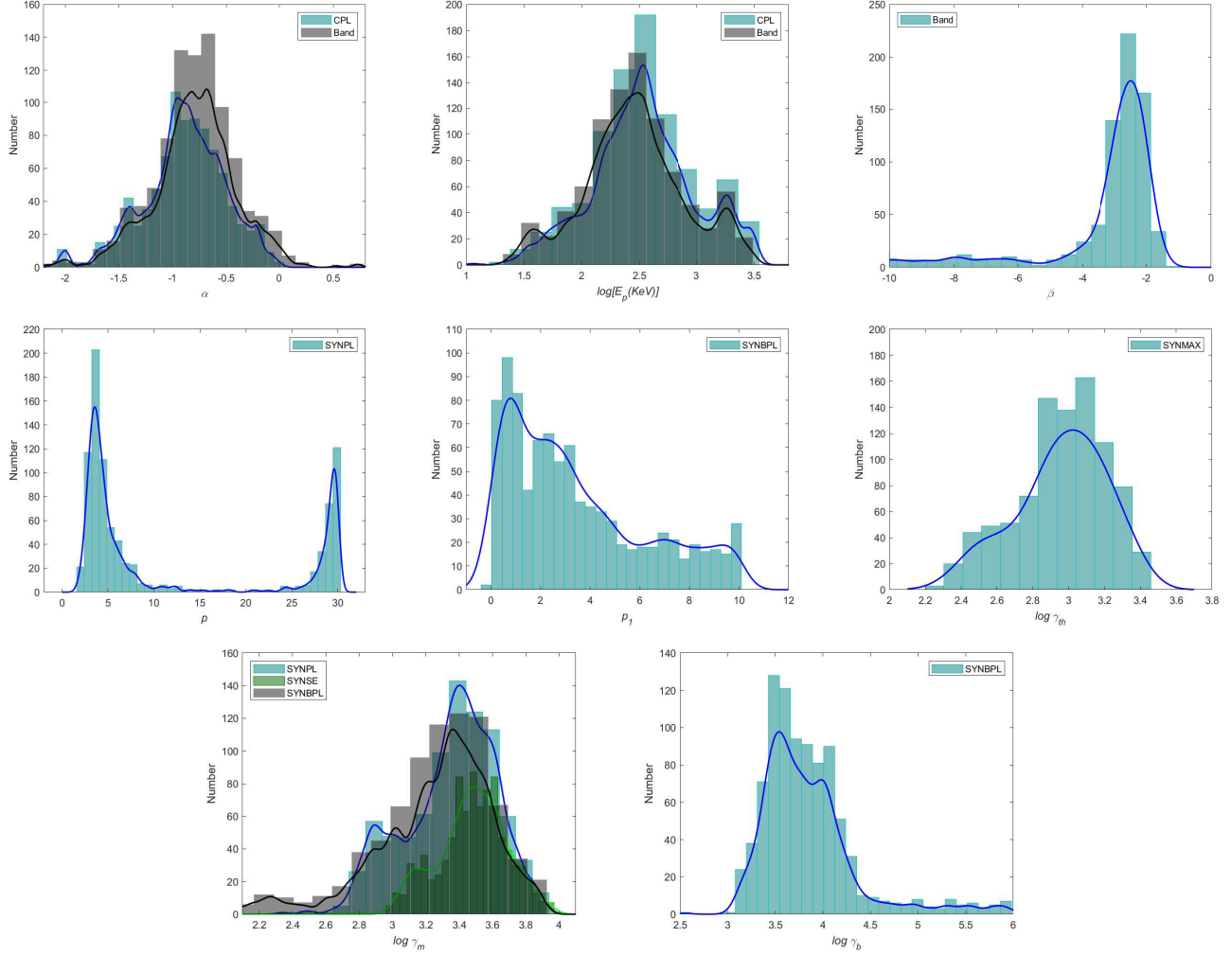


Figure 5. Distributions of each model parameter are shown, including α and $\log(E_p)$ of CPL and BAND; β of BAND; p and $\log(\gamma_m)$ of SYNPL; p_1 , $\log(\gamma_m)$ and $\log(\gamma_b)$ of SYNBPL; $\log(\gamma_m)$ of SYNSE; $\log(\gamma_{th})$ of SYNMAX. The SYNFC is not included because it fits the data poorly. The curves represent the kernel density estimation (KDE) of the distributions, where Gaussian kernels are used.

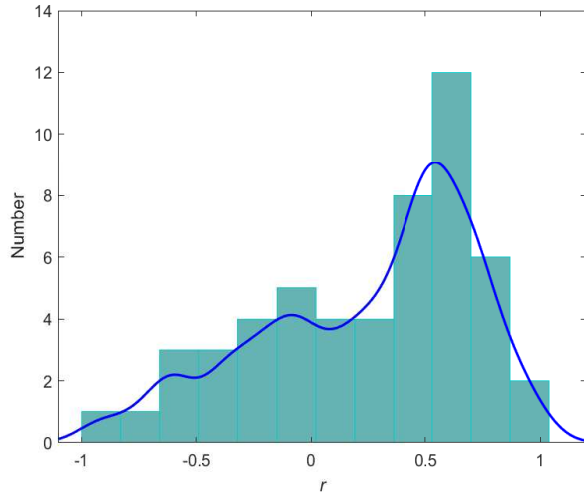


Figure 6. Distribution of the correlation coefficient (r) between the electron distribution index of the SYNPL model and the flux of 53 GRBs.

some bursts. This suggests that the electrons may be accelerated by other mechanisms.

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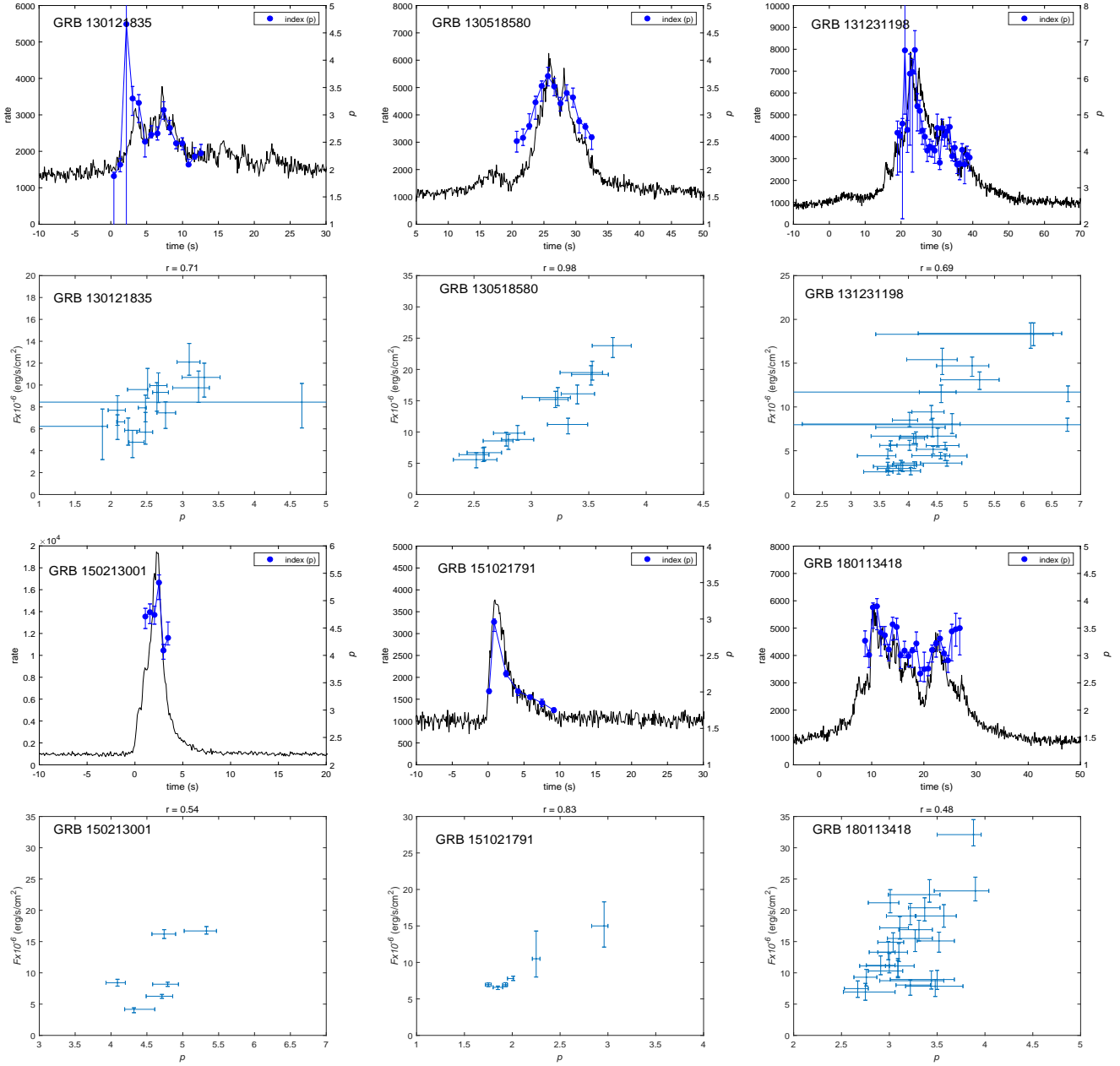


Figure 7. Evolution of p parameter in the SYNPL model (blue data points, right axis) of six sample bursts with the light curve (black curve, left axis) and correlations between p and the flux

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