

Acoustic detection of ultra-high energetic neutrinos - a snap shot -

Rolf Nahnauer

DESY, Platanenallee 6, D-15738 Zeuthen, Germany

Abstract

Already more than 30 years ago the acoustic particle detection method has been considered to be one possibility to measure signals from ultra-high energetic neutrinos. The present status and problems of corresponding model predictions are discussed in comparison with existing experimental measurements. Available acoustic sensors and transmitters are described and new ideas for corresponding applications are mentioned. Different methods for in-situ calibrations are discussed. Results of measurements of acoustic test arrays at different sites are presented in some detail. Future activities for applications of the technology in large size detectors are evaluated.

Keywords: acoustic neutrino detection, cosmogenic neutrinos, acoustic transducers

PACS: 43.58, 43.60Fg, 95.85 Ry

1. Introduction

The detection and study of ultra-high energetic neutrinos above 10^{17} eV created in cosmic sources or by interaction or decay of even higher energetic particles became of increasing interest during the last decade. Several interesting questions of particle physics, astrophysics and cosmology, could be studied measuring interactions of such neutrinos on Earth with reasonable statistics [1].

Corresponding neutrino flux predictions were already small ten years ago [2] but had to be decreased further by new bounds from cosmic ray [3] and high energy gamma ray measurements [4]. Present day experiments could derive therefore until now only flux limits [5, 6, 7, 8]. A first detection of such neutrinos is expected with detectors of $\sim 100 \text{ km}^3$ size. To get reasonable statistics will need probably about an order of magnitude larger detector volumes. It seems impossible today to instrument such experiments with conventional optical detectors within reasonable cost limits. **Acoustic particle detection may be one option among others to overcome this problem [9].**

The possibility to detect charged particles by the sound they produce passing through matter was the first time mentioned in 1957 by G. Askaryan [10]. About 20 years later a corresponding model was formulated

[11, 12], and first ideas about a 100 km^3 acoustic detector were discussed seriously [13]. In the following sections it will be shown, how far the predictions of the Thermo-acoustic Model could be confirmed experimentally and what questions have still to be answered.

During the last 10 years different groups tried to use acoustic test arrays to extract basic information needed to build large scale detectors in different materials and environmental conditions. Their results will be summarized in the second part of this paper. Finally future steps for the improvement of the acoustic technology and its application will be mentioned.

2. The Thermo-Acoustic Model

Ideas about the Thermo-acoustic Model of the creation and propagation of sound in particle interactions were displayed for the first time at the 1976 DUMAND meeting [14]. They were formulated independently by Bowen [12] and Askaryan and Dolgoshein [11]. More detailed descriptions of both concepts were published three years later [15, 16]. Recently a new approach based on [16] has been published including signal attenuation effects [17].

In neutrino interactions a charged or neutral lepton and a hadronic particle cascade is produced. The cascade gives rise to a large energy deposition in a small

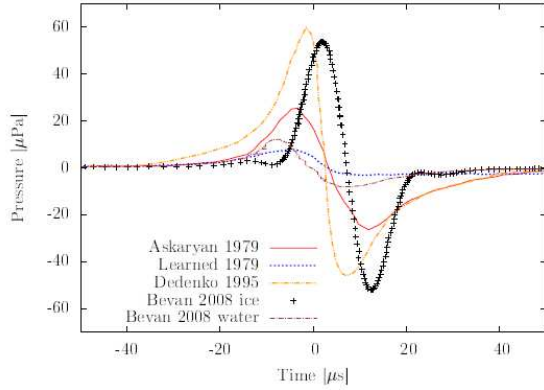


Figure 1: Signal strength from three different acoustic model parameterizations [18] and inclusion of attenuation effects [17] (from [19])

volume in a very short time. The volume is overheated what gives rise to a pressure wave which develops orthogonal to the cascade and therefore the incident neutrino direction.

2.1. model predictions

An illustrative way to describe the dependencies of the important quantities and variables of the process is given in [15].

$$\begin{aligned} p &= (k/c_p)(E/R)M \\ M &= (f^2/2)(\sin x/x) \\ f &= v_s/(2d) \\ x &= (\pi L/2d)(\sin \delta) \end{aligned} \quad (1)$$

with p : pressure amplitude, E : cascade energy, R : distance to receiver, f : frequency, v_s : speed of sound, d : cascade diameter, k : volume expansion coefficient, c_p : specific heat, L : cascade length, δ : angle between normal to cascade direction and receiver.

The model predicts a linear dependence of the pressure amplitude on the particle cascade energy, which is related to the incoming neutrino energy and allows to determine the neutrino direction from measuring the pressure wave propagation through the medium.

Unfortunately absolute signal predictions are uncertain mainly due to the not well known particle interaction and energy loss process at ultra-high energies. Different assumptions about what values should be used for the width and the length of a cascade of a certain energy give rise to signal predictions which differ by about an order of magnitude (see fig. 1).

In the considered energy range two other phenomena have also to be taken into account: the Landau-

Pomeranchuk-Migdal-effect and photo- and electro-nuclear reactions [20]. The LPM-effect predicts decreasing cross section for bremsstrahlung and pair production, elongating hadronic cascades above about 10^{18} eV whereas the second effect act in the opposite direction at energies above 10^{20} eV.

2.2. target material effects

One could rewrite eq. 1 in the form

$$p = \gamma (E/R) M' \quad (2)$$

with $\gamma = v_s^2 (k/c_p)$ and $M' = (1/2)(1/d^2)(\sin x/x)$.

The Grueneisen coefficient γ is a strongly material dependent quantity. In tab. 1 this is displayed for three materials under discussion for acoustic detector applications. At the same incoming energy signals in ice should be therefore nearly an order of magnitude larger than those in water. In salt even larger signals are expected. Also permafrost was recently suggested to give rise to quite large signals [22].

Table 1: Thermo-acoustic model parameters and boundary conditions for three different materials (adapted from [21])

	water	South Pole ice	salt
c (m/s)	1530	3880	4560
$(k/10^{-5})[K^{-1}]$	25.5	12.5	11.6
$c_p[J(K \text{ kg})^{-1}]$	3900	1720	839
γ	0.153	1.12	2.87
$f_{max}[kHz]$	7.7	20	42
refraction	moderate	very small	small?
λ_{att}	>1000 m	~300 m	>100 m
noise	variable	stable, <14 mPa	small?

For water things are even more complicated, because the volume expansion coefficient for water depends strongly on the water temperature. At 4 degree Celsius it is equal to zero having different signs below and above this temperature. The signal strength for acoustic pulses in water depends therefore strongly on the specific location of the neutrino interaction and may vary also by a factor 10 (see [23, 24]).

3. Experimental verifications

Already at the end of the seventies of the last century several experiments were performed to test the predictions of the Thermo-acoustic Model using accelerator or laser beams [24, 25]. The basic experimental arrangement is shown in fig. 2.

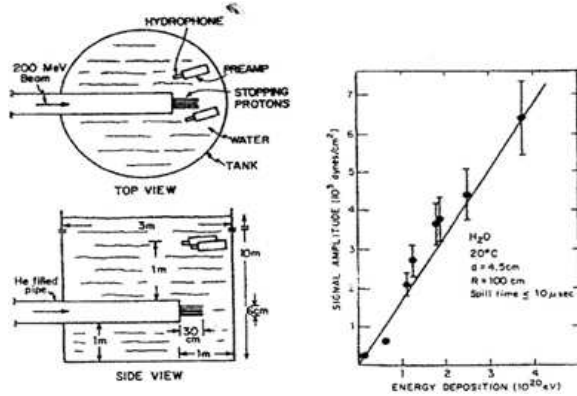


Figure 2: Experimental setup and results for signal dependence on deposited energy for an early Thermo-acoustic model check (from [24])

An intense laser or low energy proton beam is stopped in a water tank. Varying the beam parameters allowed several model checks like the following:

- varying intensity (E -dependence)
- varying beam diameter (f -dependence)
- varying distance (R -dependence)
- varying liquids (γ -dependence)

During the last decade several similar tests have been made by different groups for water [26, 27, 28], ice [29, 30] and permafrost [22]. The general conclusions from all these tests are the following:

- many predictions of the Thermo-acoustic Model could be confirmed
- the dominant mechanism for acoustic signal production is thermal expansion
- other contributions, e.g. from micro-bubble formation could not completely be excluded
- absolute signal values could not be checked, because the energy deposit in the different tests was difficult to calculate precisely and was different to the ultra-high energetic particle case.

An unsolved problem was until recently also the relative signal strength in different target materials. It is therefore very valuable that an experiment is performed this year at the Aachen-Acoustic-Laboratory aiming to answer this question [31]. A well controlled laser beam is stopped in a large tank placed in an even larger freezer. This allows measurements in a temperature range from +20 to -25 degree Celsius in water and ice. In fig. 3 acoustic signals for both configurations are shown. The time difference of the signals is due to the

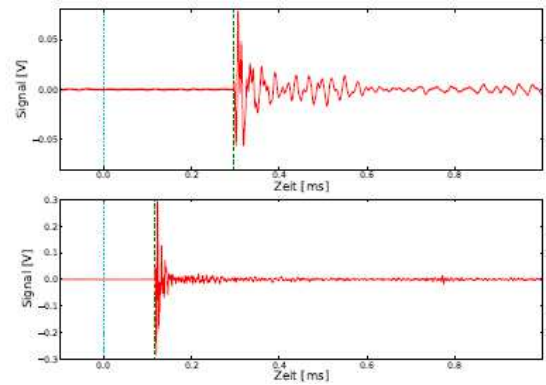


Figure 3: Acoustic signals from a laser beam shot in a tank filled with water (top) and ice (bottom) (from [31])

different velocity of sound in both materials. The signal strength in ice is about a factor six larger than in water consistent with the theoretical expectation. The exact comparison has however to be done for the corresponding pressure amplitudes. More detailed results of these measurements will be available soon.

4. Acoustic transducers

Acoustic transducers for ultra-sound have been developed since long for applications in water. Their quality profited from their use in military projects. Today commercial products are available from several companies. Special requirements like stability at high pressure for use at large water depth lead to considerably high prices per piece. Because large numbers of corresponding devices are necessary to build large detector arrays, several attempts were started to build own sensitive sensors and transmitters. For applications in ice, salt and permafrost this was unavoidable anyway.

4.1. piezo ceramic sensors

Nearly all acoustic sensors in use in today's test arrays are based on piezo-ceramic elements [32]. Within the AMADEUS project a nice comparison of different commercial and self-made devices has been performed [33]. In fig. 4 the response of three of their sensors to the same transmitter signal is shown. An interesting development is their "Acoustic Module", where piezo-ceramic elements are integrated in a glass pressure sphere otherwise used to contain photomultipliers for light detection in the ANTARES experiment [34].

The SPATS group connected to the IceCube experiment started their own developments for ice applications with a similar concept [29, 35] but changed their

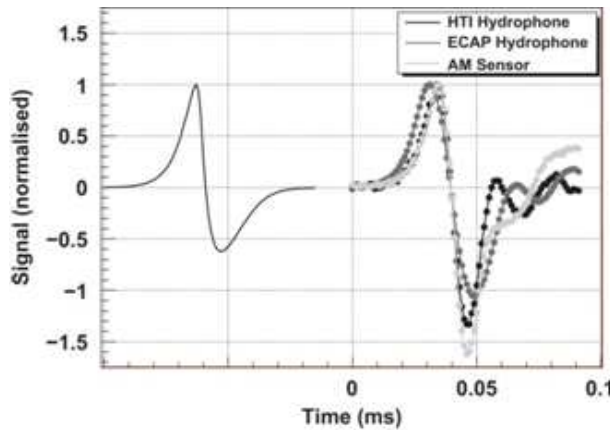


Figure 4: Comparison of the response of three different acoustic sensor types to a bipolar pulse (from [33])

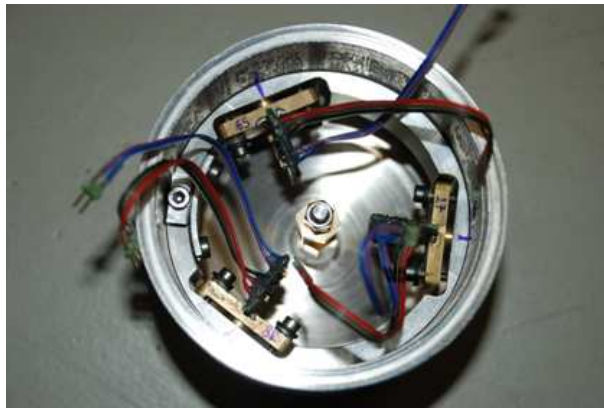


Figure 5: SPATS second generation sensor module (from [36])

design finally to a cylindrical steel pressure housing where three piezo-ceramics are pressed against the inner wall (see fig. 5). Typical sensitivities of sensors in use are in the range -190 to -110 dB re 1V/ μ Pa.

Recently the use of piezo-foils for the design of acoustic sensors has been reported [31]. This would have the advantage that one would be more flexible in sensor shape and size. Results from forthcoming tests will be hopefully published soon.

4.2. other sensor concepts

A “new” concept for acoustic sensors under discussion in particle detection applications since a few years is the use of acousto-optic hydrophones. In fact the corresponding working principles using either frequency or intensity modulations in optical fibers has been discussed already more than 30 years ago [37].

Recently several groups published results for a fiber laser, where the modulation of the cavity size by pres-

sure pulses leads to modulations of the laser wavelength detectable by interferometry [38]. The concept has been proven to work at a static pressure of 35 MPa, i.e. 3500 m depth with higher sensitivity and better resolution than piezo-ceramic based devices. An application in a real open water test array is however still missing.

The application of coupled waveguide intensity modulated hydrophones is under discussion for an upgrade of the BAIKAL acoustic test setup [23].

4.3. transmitters

Acoustic transmitters are normally used in acoustic arrays to get position and calibration information. An exceptional idea was realized for this purpose in the SAUND-1 experiment, where light bulbs of different size were deployed on top of seven hydrophones at the AUTECH military array near the Bahamas [39].

The SPATS group published results from different types of pingers fixed in ice or used in water filled holes to measure the acoustic attenuation length in South Pole ice [40]. A byproduct of these measurements was the observation of shear waves in the ice [41].

Position monitoring is also a problem of all deep water optical neutrino telescopes. It is normally solved by fixing strong acoustic transmitters at the sea bed and by adding some hydrophones to the optical strings. Pulses emitted with short time differences allow then precise position monitoring at the 10 cm scale [42]. For the planned KM3Net detector a corresponding system is under development, which will allow to observe pulses from particle cascades with high sensitivity [43].

To test the sensitivity of acoustic arrays in-situ is still an unsolved problem. Two groups try to develop acoustic pulser systems which mimic pulses from ultra-high energetic neutrino interaction in strength and shape. Successful tests have been made using arrays with more than 5 transmitters taking into account the transfer function between emitted and received signals [44]. In another approach a parametric acoustic source is used. The overlay of two high frequency signals fed to an emitter leads to the production of a bipolar pulse at lower frequency whereas the remaining high frequency components are quickly absorbed in the medium [45].

5. Sensor calibration

For the application of the acoustic technology for particle detection the use of carefully calibrated sensors is mandatory. Several methods are applied for this purpose in the laboratory [46], e.g. comparison with a calibrated reference hydrophone and reciprocity calibration or use

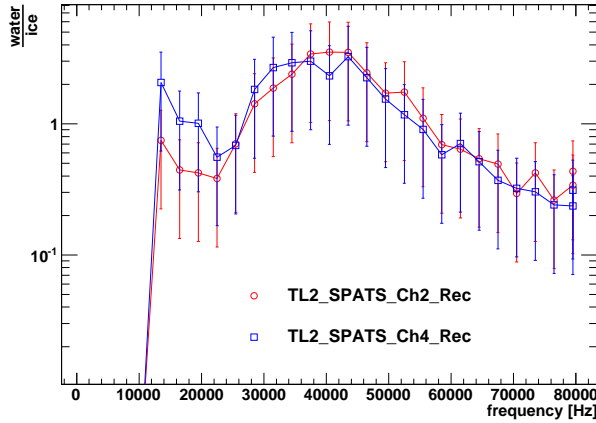


Figure 6: Preliminary result for the ratio of sensitivities of a third generation SPATS sensor deployed in water and ice (from [49])

of calibrated emitters. With the last method all sensors used for the AMADEUS project were calibrated in dependence of the azimuthal and zenith angle in a water tank [33]. One has to keep in mind, however, that the sensor sensitivity depends on the specific environmental conditions at the deployment location as e.g. temperature and pressure. Together with a NATO institute the ONDE-group has developed a standard procedure for under pressure calibration. They reported a sensitivity change of about 2 dB/1000 m water depth [47].

In-situ calibration in ice is even more complicated than in water. Beside pressure effects, deep temperature and impedance changes between ice and water have to be taken into account. The SPATS group has calibrated their sensors in water and studied separately the dependence on pressure (in a water-oil mixture), temperature (in air) and impedance (measuring noise level changes during sensor freeze-in). Assuming that the corresponding sensitivity changes can just be multiplied, they got a final result with about 40 percent error [48].

A direct comparison of calibration results in water and ice at normal pressure was done for a third generation SPATS sensor in the Aachen-Acoustic Laboratory [49]. Preliminary results of a reciprocity calibration, shown in fig. 6 are consistent with the estimation described above. The final result of these measurements is expected to be published soon.

6. Acoustic test arrays

During the last decade results were reported from six test sites either using part of military hydrophone arrays or installing own acoustic sensors. The main purpose of these studies was to clarify how sensitive acoustic

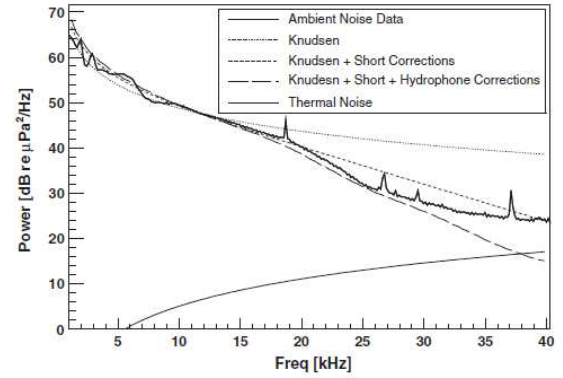


Figure 7: Average ambient noise level measured by SAUND (from [50])

detectors could be to signals from neutrino interactions at given environmental conditions like acoustic noise, signal attenuation and refraction. A few prominent results of the different groups will be discussed below. It is however out of the scope of this presentation to give a detailed overview about presently available data. To get more information the reader is pointed to the references given in this paper and to the proceedings of the ARENA conferences since 2005 [9].

6.1. SAUND

The “Study of Acoustic Ultra-high Neutrino Detection - SAUND” started in 2003. Seven hydrophones of the military AUTECH array near the Bahamas were used to search for acoustic signals. In 196 days 65×10^6 triggers were taken in a sensitive volume of about 15 km^3 . The data were used to calculate the first acoustic neutrino flux limit [39].

In the second phase of this experiment the number of used hydrophones increased to 49 and the volume under study to 1500 km^3 . 130 days of data taking allowed detailed studies of acoustic noise behaviour (see fig. 7) [50]. In a complex data reduction and signal processing procedure two events were found compatible with showers from neutrino interactions above 10^{22} eV . The result was used to derive an improved neutrino flux limit (see fig. 12) [51].

6.2. Acorne

The “Acoustic COsmic Ray Neutrino Experiment - ACORNE” is an activity of different groups in the UK. Eight hydrophones of the RONA military array in North-West Scotland are used for acoustic signal searches. Between 2006 and 2008 nearly 30 Tb of raw data have been collected. Advanced signal processing

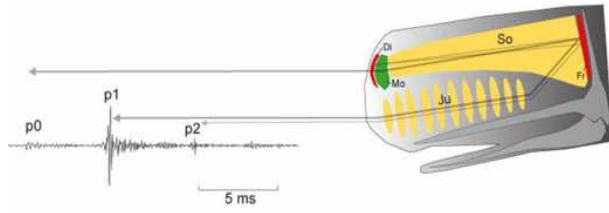


Figure 8: Acoustic signals from a sperm whale (from [47])

and filtering technologies were developed by the group and applied to the data [52, 53]. Corresponding data reduction and efficiency calculations allowed to derive a neutrino flux limit in the same energy range as the SAUND result (see fig. 12). Furthermore information about signal attenuation and localization was collected. The examination of the data for black hole signatures lead to the calculation of a corresponding upper limit [54].

6.3. ONDE

Within the “Ocean Noise Detection Experiment - ONDE” four hydrophones forming a tetrahedron were deployed in 2005 to about 2000 m depth in the Mediterranean Sea 25 km offshore of Catania/Italy. Noise was monitored about five minutes every hour. The average noise level in the 20-43 kHz frequency region was found to be $5.4 \pm 2.2 \pm 0.3$ mPa [55]. The noise level was strongly correlated with the actual environmental conditions at the sea surface.

Signals observed in the noise data more often than expected were produced by a population of sperm whales. Analyzing these signal in detail (for an example see fig. 8) even allowed to derive age and gender of the animal emitting it [47]. Corresponding studies will continue in collaboration with marine biologists.

6.4. Baikal

A subgroup of the collaboration which constructed and operates the first optical neutrino telescope in lake Baikal [56] has deployed a digital hydro-acoustic module with four hydrophones at a regular tetrahedron of 1.5 m edge length in 150 m depth at one of the outer strings of the optical array. With this device and its predecessors extensive environmental studies have been done. The noise level was measured most of the time below 5 mPa. Accepting only signals from the deep lake one interesting neutrino-like event has been observed (see fig. 9).

In March 2011 an acoustic string with three acoustic modules has been deployed [23]. Data taking is ongoing

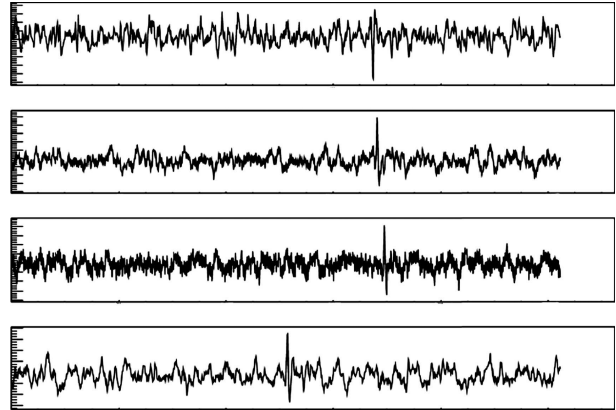


Figure 9: Acoustic event reconstructed as upward going neutrino-like signal (from [23])

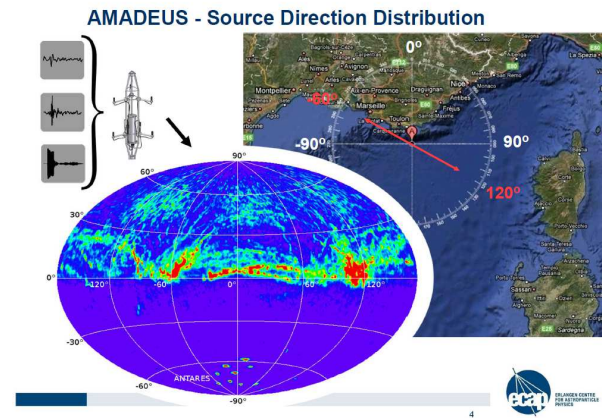


Figure 10: Angular distribution of transient events observed with the AMADEUS detector (from [58])

and first results are expected to be shown at next years conferences.

6.5. ANTARES

A detailed description of the design and performance of the “ANTARES Module for the Acoustic DEtection Under the Sea - AMADEUS” can be found in [33]. 36 sensors are located at three storeys of both string 12 and the instrumentation line of the ANTARES optical neutrino telescope [34] located at a depth of ~ 2500 m in the Mediterranean sea about 40 km distant to Toulon/France. Data taking started in 2007. Extensive noise studies have shown strong correlations to actual weather conditions. A high statistics frequency dependent noise measurement reflects similar results like those found with ONDE in the deep sea near Sicily (see section 6.3).

The arrival direction of transient acoustic signals could be determined (see fig. 10) [57]. It was found that these signals have mostly antropogenic origin being due to heavy ship traffic in this region of the sea.

6.6. SPATS

The “South Pole Acoustic Test Setup - SPATS” has been deployed in the upper empty part of holes drilled for the IceCube neutrino observatory [5]. Seven acoustic stations with transmitters and receivers are positioned between 80 m and 500 m in the ice at the South Pole, with a maximum distance of about 520 m between strings [36]. Data taking started in early 2007. Since then results have been published for the speed of sound of pressure and shear waves and their refraction versus depth [41], the acoustic attenuation length [40] and the ambient noise level [48].

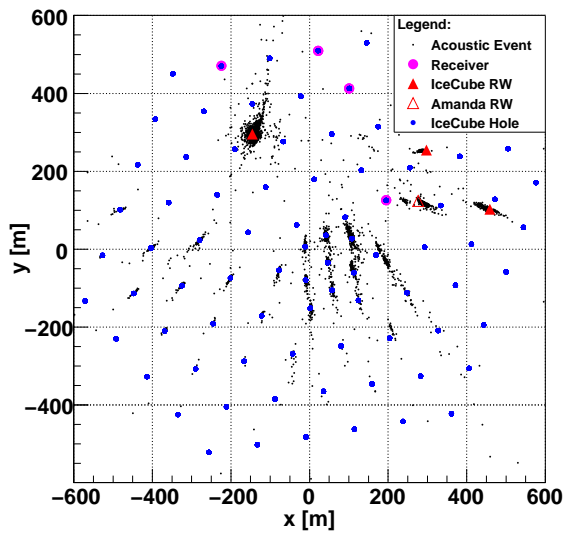


Figure 11: Distribution of transient signal source locations in the x-y plane compared with the position of IceCube holes and Rod-wells (from [48])

Transient events were observed from re-freezing IceCube holes and from water reservoirs used for hole drilling (see fig. 11). From the non-observation of acoustic signals in a region outside the IceCube construction area a neutrino flux limit has been estimated (see fig. 12).

7. Acoustic neutrino-flux limits

Neutrino flux limits have been derived until now in three acoustic test experiments [51, 53, 48], all not optimized or not designed for this purpose. In fig. 12

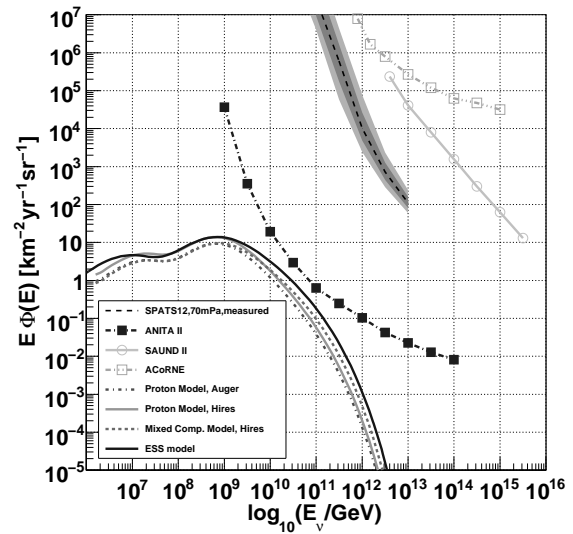


Figure 12: Available neutrino flux limits from acoustic test arrays compared to some model predictions and results from some radio detectors sensitive at highest energies (from [48])

the corresponding results are shown and compared with predictions of one frequently quoted cosmogenic neutrino flux model and its recent modifications [2]. The presently best limit in the considered energy region has been published recently by the ANITA collaboration using radio antennas as payload of a balloon circling in a height of 35 km around Antarctica [8]. The detector was searching for pulses of coherent radio emission from neutrino induced cascades in the ice below in a volume of about $1.6 \times 10^6 \text{ km}^3$.

The acoustic neutrino flux limits are today still more than four orders of magnitude less sensitive than the best radio limits which is partly explained by the huge difference in the corresponding detection volumes used. What makes all limits difficult to interpret are the partly unknown systematic errors of the measurements and the reliability of assumptions made for efficiency and sensitivity calculations. This has been discussed e.g. in [48].

8. Future activities

The experiments mentioned in section 6 have in the meantime either finished or achieved their primary goals. Follow up programs are under discussion or already in a planning stage.

8.1. acoustic detection in water

With the present facilities “fake neutrino sources” (see section 4.3) will be used to get better knowledge

about in-situ detector efficiencies and trigger schemes. On an intermediate time scale design studies towards a km³-sized hybrid acousto-optic detector will hopefully be followed by its construction in the Mediterranean KM3Net project [47]. Most European acoustic activities should converge around this program.

8.2. acoustic detection in ice

Ice provides probably the best conditions to build a hybrid radio-acoustic-optical detector allowing background-free detection of the rare neutrino signals above 10¹⁸ eV. Having evaluated the acoustic ice properties at the South Pole, the boundary conditions for such a detector are now mostly known at that location. Corresponding Monte Carlo design studies are under way. They show, that already a 100 km³ scale detector would need a large number of holes to be drilled. The development of robotic drilling and deployment methods is mandatory for such a project. First attempts in that direction are just starting [59].

The detection of ultra-high energy neutrinos and the detailed study of their interactions remains a top priority topic of astroparticle physics also in the long term future. Corresponding experiments need most probably to employ multiple detection methods with different systematics for successful operation. The acoustic particle detection technology is qualified to be one of them.

Acknowledgements

I would like to thank the organizers of this conference for the invitation to give this presentation. I thank also N. Budnev, G. Gratta, D. Heinen, R. Lahmann and G. Riccobene for providing new information about activities of their groups. I am indebted to J. Berdermann for valuable suggestions and critical reading of the manuscript.

References

- [1] R. Nahnauer, NIM A (2010), doi:10.1016/j.nima.2010.11.010
- [2] R. Engel, D. Seckel, T. Stanev, Phys. Rev. D64 (2001) 093010, update 2008: (ftp://ftp.bartol.udel.edu/seckel/ess-gzk/2008)
- [3] L.A. Anchordoqui et al., Phys. Rev. D76, (2007) 123008
- [4] M. Ahlers et al., Astropart. Phys. 34, (2010) 106
- [5] R. Abbasi et al., Phys. Rev. D83, (2011) 092003
- [6] J. Abraham et al., Phys. Rev. D79, (2009) 102001
- [7] I. Kravchenko et al., Phys. Rev. D73, (2006) 082002, update: arXiv 1106.1164
- [8] P.W. Gorham et al., Phys. Rev. D82, (2010) 022004
- [9] R. Nahnauer, S. Boeser (eds.), Proceedings of ARENA 2005, Int. J. Mod. Phys. A21 (2006); L. Thompson, S. Danaher (eds.) Proceedings of ARENA2006, J. Phys. Conf. ser. 81, (2007); A. Capone et al., (eds.) Proceedings of ARENA2008, NIM A 604

- (2009); P. Latridou et al., (eds.) Proceedings of ARENA2010, in Press
- [10] G.A. Askaryan, Atomnaja Energiya V3 (1957)152
- [11] G.A. Askaryan, B. Dolgoshein, JETP Lett.25 (1977) 213
- [12] T. Bowen, Proceedings of the 15th ICRC, Plowdiv, vol.6 (1977)p.277
- [13] V.S. Berezinsky, G.T. Zatsepin, Sov. Phys. Usp.. 20 (1977) 361
- [14] A. Roberts, Rev. Mod. Phys. 64 (1992) 259
- [15] G.A. Askaryan et al., NIM 164 (1979) 267
- [16] J. Learned, Phys. Rev. D 19 (1979) 3293
- [17] S. Bevan et al., NIM A607 (2009) 398
- [18] A.V. Butkevich et al., Fiz. El. Sch. Atom. Jad. 29 (1998) V3
- [19] F. Descamp, Thesis, Universiteit Gent, (2009)
- [20] L. Gerhard, S. Klein, Phys. Rev. D82 (2010) 074017
- [21] P.B. Price, J. Geophys. Res. 111 (2006) B02201.
- [22] R. Nahnauer et al., NIM A 587 (2008) 29
- [23] V. Aynutdinov et al., NIM A (2010) doi:10.1016/j.nima.2010.11.153
- [24] L. Sulak et al., NIM 161 (1979)203
- [25] R. Golubnichy et al., Proceedings of DUMAND 1979, p.148
- [26] V. I. Abdul et al., Instr. Exp. Techn. 44 (2001) 327
- [27] K. Graf et al., Int. J. Mod. Phys. A21 (2006) 127
- [28] G. De Bonis et al., NIMA 604 (209) 199
- [29] S. Boeser, Thesis, Humboldt Universitaet Berlin, 2006
- [30] J. Stegmeier, Diploma Thesis, Johannes Gutenberg Universitaet Mainz, 2004
- [31] D. Heinen, Diploma Thesis RWTH Aachen, 2011
- [32] G. Anton et al., Astropart. Phys. 26 (2006) 301
- [33] J. Aguilar et al., NIM A 626-627 (2011) 128
- [34] M. Angeron et al., arXiv:1104.1607 (2011), subm. to NIM A
- [35] J. Boeser, Int. J. Mod. Phys. A21 (2006) 107
- [36] Y. Abdou et al., arXiv:1105.4339 (2011) subm. to NIM A
- [37] J. N. Fields, Appl. Optics, V18, (1979) 3533
- [38] P.E. Bagnoli et al., NIMA 567(2006) 515, E. Maccione et al., NIMA 572 (2007) 490; A. Cotrufo et al. NIMA 604 (2009) 219; J. Wang et al., Optoe. Let. V3 (2007) doi:10.1007/s11801-007-6169-1; S. Goodman et al., Proc. SPIE V 7004 (2008), doi:10.1117/12.785937; B. Guan et al. , Opt. Expr. V17 (2009)19544 and references therein
- [39] J. Vandenbroucke et al., Astrophys. J. 621 (2005) 301
- [40] R. Abbasi et al., Astropart. Phys. 34 (2011) 382
- [41] R. Abbasi et al., Astropart. Phys. 33 (2010) 277
- [42] M. Ardid, NIM A 602 (2009) 174
- [43] P. Ameli et al., NIM A 626 (2011) 211
- [44] W. Ooppakaev, S. Danaher, NIM A (2011), doi:10.1016/j.nima.2010.11.142
- [45] M. Ardid et al., NIM A (2010), doi:10.1016/j.nima.2010.11.139
- [46] M. Ardid, NIM A 604 (2009) 203
- [47] G. Riccobene, contribution to this proceedings
- [48] R. Abbasi et al., arXiv:1103.1216
- [49] L. Paul, RWTH Aachen internal report, 2011, unpublished
- [50] N. Kurahashi, G. Gratta Phys. Rev. D78 (2008) 092001
- [51] N. Kurahashi et al., Pys. Rev. D82, (2010) 073006
- [52] S. Danaher, J. Phys. Conf. ser. 81, (2007) 012011
- [53] S. Bevan, NIM A 604 (2009) 143
- [54] S. Danaher et al., NIM A (2010), doi:10.1016/j.nima.2010.11.075
- [55] g. Riccobene et al., NIM A 604 (2009) 149
- [56] V. Aynutdinov et al., Astropart. Phys. 7 (1997) 263
- [57] R. Lahmann, NIM A, doi:10.1016/j.nima.2010.11.157
- [58] R. Lahmann, private communication
- [59] K. Laihem, contribution to this proceedings