

# Simulating the Ionosphere

Anna Scaife

Cavendish Astrophysics  
University of Cambridge.

MCCT SKADS  
Summer School on Third Generation Calibration

# Outline

- 1 The Basics
- 2 interpolation
- 3 Model fitting

# Let's start with the **Z-Jones**...

## Z-Jones

$$Z = \begin{bmatrix} e^{i\phi_{\text{iono}}} & 0 \\ 0 & e^{i\phi_{\text{iono}}} \end{bmatrix}$$

$$\phi_{\text{iono}} = -25\lambda\text{TEC}$$

$$\text{TEC} = \dots$$

# Let's start with the **Z-Jones**...

## Z-Jones

$$Z = \begin{bmatrix} e^{i\phi_{\text{iono}}} & 0 \\ 0 & e^{i\phi_{\text{iono}}} \end{bmatrix}$$

$$\phi_{\text{iono}} = -25\lambda\text{TEC}$$

$$\text{TEC} = \dots$$

# Let's start with the **Z-Jones**...

## Z-Jones

$$Z = \begin{bmatrix} e^{i\phi_{\text{iono}}} & 0 \\ 0 & e^{i\phi_{\text{iono}}} \end{bmatrix}$$

$$\phi_{\text{iono}} = -25\lambda\text{TEC}$$

$$\text{TEC} = \dots$$

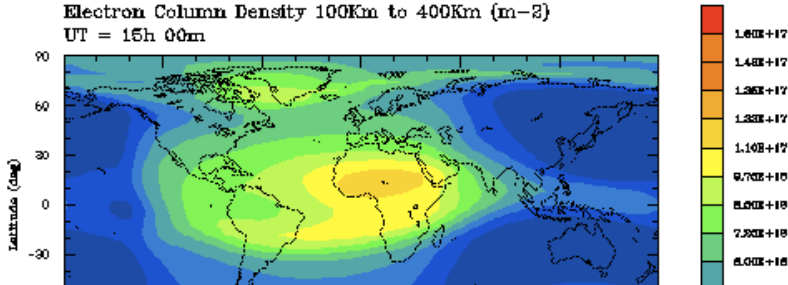
# TEC

$$1 \text{ TECU} = 10^{16} \text{ m}^{-2}$$

$$\text{TEC} = \int_0^{\infty} n_e dl + \text{const}$$

**Quiet Ionosphere UT = 15h 00m**

**Electron Column Density 100Km to 400Km (m<sup>-2</sup>)**  
UT = 15h 00m



- It seems that we should be able to predict the behaviour of the ionosphere...
- ...but, in fact we are limited to interpolation of scattered measurements.
- Counter-intuitively, simulating the ionosphere is in fact an inverse problem.
- Fortunately there are lots of ways of approaching inverse problems.

# Interpolation

Let's look at interpolating between measurements...

**DEFINITION:** Approximating measurements at intermediate scales/positions from scattered measurements.

- We have *sparse* measurements
- i.e. We have *under-sampled* data



# Interpolation

Let's look at interpolating between measurements...

**DEFINITION:** Approximating measurements at intermediate scales/positions from scattered measurements.

- We have *sparse* measurements
- i.e. We have *under-sampled* data

# Interpolation schemes

- 1 Kriging
- 2 Triangulation based
- 3 Natural neighbour
- 4 Splining

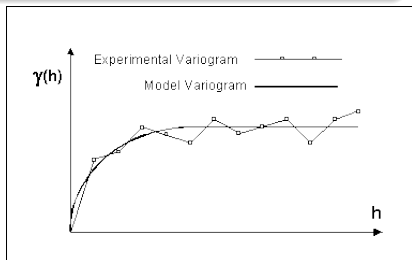
# Kriging

## Semivariance

$$\gamma(\mathbf{h}) = \frac{1}{2} (f(\mathbf{x}) - f(\mathbf{x} + \mathbf{h}))^2$$

A typical model:

$$\gamma = \begin{cases} s\left(\frac{3}{2}\left(\frac{h}{a}\right) - \frac{1}{2}\left(\frac{h}{a}\right)^3\right), & 0 \leq h \leq a \\ s, & h > a \end{cases}$$



# Triangulation

Given three measured points we can interpolate to any point within the triangle using:

## Delaunay triangulation

$$f(x, y) = \sum_{i=1}^3 \phi_i(x, y) f_i$$

- $\phi_i(\mathbf{x})$  is our basis function
- In a simple case we can use linear equations:

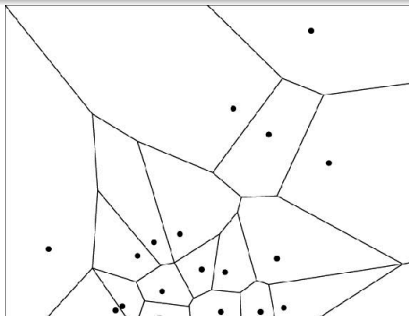
$$f(x, y) = c_1 x + c_2 y + c_3$$

- We can just solve  $\mathbf{A}\mathbf{c} = \mathbf{f}$ , where  $\mathbf{f} = (f_1, f_2, f_3)^T$  and  $\mathbf{A} = \{(x_i, y_i, 1)\}_{1 \leq i \leq 3}$ .

# Natural neighbour

## Voronoi tessellation

- Voronoi tessellation divides the data into cells defined by the positions of the measurements.
- We use the interpolation point to define a new Voronoi cell.
- The value of this cell can be evaluated as the weighted sum of the contributions from its overlapping cells.



# Splining

## Bayes Theorem

$$p(D|M, \Theta) = \frac{L(M, \Theta|D)\pi(\Theta)}{E}$$

$$\textit{posterior} = \frac{\textit{likelihood} \times \textit{prior}}{\textit{Evidence}}$$

## Bayes Theorem

$$p(D|M, \Theta) = \frac{L(M, \Theta|D)\pi(\Theta)}{E}$$

$$\textit{posterior} = \frac{\textit{likelihood} \times \textit{prior}}{\textit{Evidence}}$$



# Maximum likelihood

We maximise the likelihood of the **DATA** w.r.t the **MODEL**.  
When we perform a  $\chi^2$  test we are in fact calculating a Gaussian ML.

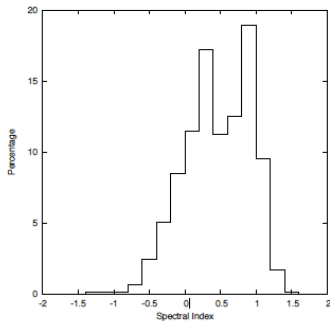
$$N(D_i, \sigma) \propto \exp \frac{-(D_i - M_i)^2}{2\sigma^2}$$

# Maximum A Posteriori

If we know something about our parameters then we can utilise that prior information.

## Example

Say we are fitting a spectral index...  $S = A\nu^{-\alpha}$



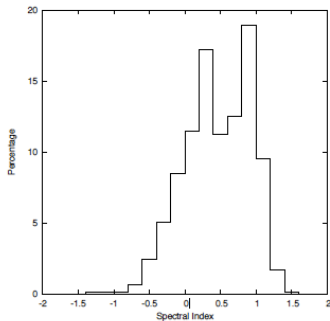
$$\ln L = -0.5 \frac{(d_i - m_i)^2}{\sigma_i^2} + \ln \pi$$

# Maximum A Posteriori

If we know something about our parameters then we can utilise that prior information.

## Example

Say we are fitting a spectral index...  $S = A\nu^{-\alpha}$



$$\ln L = -0.5 \frac{(d_i - m_i)^2}{\sigma_i^2} + \ln \pi$$

# Model comparison

What if we have **more than one** model...?

The Evidence:

$$E = \int L(\Theta)\pi(\Theta)d^D\Theta$$

The model selection ratio:

$$R = \frac{\Pr(H_1|D)}{\Pr(H_2|D)} = \frac{\Pr(D|H_1)\Pr(H_1)}{\Pr(D|H_2)\Pr(H_2)} = \frac{Z_1 \Pr(H_1)}{Z_2 \Pr(H_2)}$$

## Worked Example

Let's look at Kriging again...

### Kriging variance

$$\gamma(\mathbf{h}) = \frac{1}{2} (f(\mathbf{x}) - f(\mathbf{x} + \mathbf{h}))^2$$

$$\text{TEC}(\mathbf{x}_0) = \sum_{i=1}^n \lambda_i \text{TEC}(\mathbf{x}_i)$$

$$\Delta = 2 \sum_{i=1}^n \lambda_i \gamma(\mathbf{x}_i, \mathbf{x}_0) - \sum_{i=1}^n \sum_{j=1}^n \lambda_i \lambda_j \gamma(\mathbf{x}_i, \mathbf{x}_0)$$

$$\Gamma_{ij} \lambda = \Gamma_{i0}$$

## Worked Example

Let's look at Kriging again...

### Kriging variance

$$\gamma(\mathbf{h}) = \frac{1}{2} (f(\mathbf{x}) - f(\mathbf{x} + \mathbf{h}))^2$$

$$\text{TEC}(\mathbf{x}_0) = \sum_{i=1}^n \lambda_i \text{TEC}(\mathbf{x}_i)$$

$$\Delta = 2 \sum_{i=1}^n \lambda_i \gamma(\mathbf{x}_i, \mathbf{x}_0) - \sum_{i=1}^n \sum_{j=1}^n \lambda_i \lambda_j \gamma(\mathbf{x}_i, \mathbf{x}_0)$$

$$\Gamma_{ij} \lambda = \Gamma_{i0}$$

## Worked Example

Let's look at Kriging again...

### Kriging variance

$$\gamma(\mathbf{h}) = \frac{1}{2} (f(\mathbf{x}) - f(\mathbf{x} + \mathbf{h}))^2$$

$$\text{TEC}(\mathbf{x}_0) = \sum_{i=1}^n \lambda_i \text{TEC}(\mathbf{x}_i)$$

$$\Delta = 2 \sum_{i=1}^n \lambda_i \gamma(\mathbf{x}_i, \mathbf{x}_0) - \sum_{i=1}^n \sum_{j=1}^n \lambda_i \lambda_j \gamma(\mathbf{x}_i, \mathbf{x}_0)$$

$$\Gamma_{ij} \lambda = \Gamma_{i0}$$

We need a model for  $\gamma$

- We can calculate  $\gamma$  directly from the data:

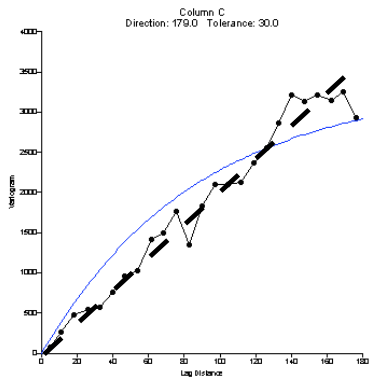
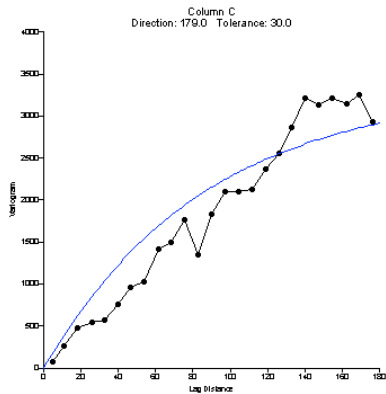
$$\gamma(\mathbf{x}_i - \mathbf{x}_j) = \frac{1}{2}(\text{TEC}(\mathbf{x}_i) - \text{TEC}(\mathbf{x}_j))^2$$

- But we need an **analytic** form for  $\gamma$
- So we have to pick a model...
- Typical models would be:

$$\gamma_1 = a + b * \mathbf{h}$$

$$\gamma_2 = a + b * \mathbf{h}^\alpha$$





# For Further Reading I



S. Dodelson.

*Modern Cosmology.*

Elsevier, 2003.



W. Hu & S. Dodelson

Cosmic Microwave Background Anisotropies

*Ann. Rev. Astron. and Astrophys.*, 2002



R. Subrahmanyan & R. Ekers

CMB observations using the SKA

*SKA Memo Series*, 26