



Disentangling the Components of the Milky Way

Inferring the Structure of the Milky Way in Phase-Space Using Gaussian
Mixture Modelling with Extreme Deconvolution

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Abstract

The Milky Way Galaxy is a complex system composed of multiple stellar components, each with distinct chemical and kinematic properties. The presence of a very-metal-poor disk structure has been speculated, which would challenge the conventional view that the disk formed later from already enriched gas. By reproducing and extending the Extreme-Deconvolution Gaussian Mixture Model (XD-GMM) study of [Zhang et al. \[2024\]](#), we investigate the presence of a very-metal-poor disk in the Milky Way using data catalogues from [Andrae et al. \[2023\]](#), [Li et al. \[2023\]](#) and [Bailer-Jones et al. \[2021\]](#).

In general, our results are in agreement with that of [Zhang et al. \[2024\]](#): below $[M/H] \simeq -1.6$ we do not find evidence for a coherent, rotationally supported disc component. The prograde rotation signal stems from a prograde halo with a net prograde rotation of $v_\phi \simeq 68 \text{ km s}^{-1}$. In this regime, no component satisfies $v_\phi/\sigma_\phi > 1$; any thick-disc contribution in the very-metal-poor interval $-3.0 < [M/H] < -1.6$ must therefore be negligible.

Splitting the sample by alpha abundance adds further detail to the kinematic picture. This part of the analysis builds on the framework of [Viswanathan et al. \[2024\]](#), extending it using the methodology in [Zhang et al. \[2024\]](#). Our findings are consistent with expectations: a thick-disc component emerges earlier in the high- α sequence, becoming apparent from $[M/H] \gtrsim -1.6$, whereas low- α stars at the same metallicity remain more strongly halo-dominated. An unexpected feature is the appearance of Gaia–Sausage/Enceladus-like components within the high- α regime. This likely reflects misclassified low signal-to-noise abundance measurements, rather than the presence of a truly α -rich GS/E population.

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1 Introduction

The Milky Way Galaxy, host to our solar system, is a spiral galaxy with a centre located approximately 150 000 trillion miles from Earth. Its formation history is complex and remains an active area of research. Being embedded within the Milky Way means we can study it in greater detail than any external galaxy, testing models of galaxy formation with high-precision observational data. One of the central aims of Galactic Archaeology is to reconstruct the Milky Way’s assembly by examining the chemical compositions and dynamical properties of its stars.

In this project, we reproduce and extend the analysis of [Zhang et al. \[2024\]](#), who investigated a very metal-poor disc component in the Milky Way. Very metal poor stars, formed from an interstellar medium unpolluted by earlier generations of supernovae, are among the oldest relics in the Galaxy. Discovering them on disc-like orbits would challenge the conventional view that the disc formed later from already enriched gas [[Bland-Hawthorn and Gerhard, 2016](#)], implying instead an earlier onset of disc assembly. Using Gaia DR3, the study by [Zhang et al. \[2024\]](#) applied a Gaussian Mixture Model with Extreme Deconvolution to the velocity distributions of stars across metallicity bins, probing whether a coherent disc signal persists down to the lowest metallicities.

1.1 Components of the Milky Way

The Milky Way is commonly decomposed into four stellar components: a *thin disc*, a *thick disc*, a central *bulge/bar*, and a roughly spherical *halo* [[Bland-Hawthorn and Gerhard, 2016](#), [Helmi, 2020](#)]. The thin disc describes the majority of stars and most of the interstellar gas in the galaxy. Ongoing star formation is concentrated in the “molecular–gas ring” where young, metal-rich stars trace nearly circular, co-rotating orbits with low velocity dispersion ($\sigma_\phi \simeq 20 \text{ km s}^{-1}$). Above the mid-plane lies the thick disc: an older ($\gtrsim 8\text{--}10$ Gyr), moderately metal-poor population , a scale height of $z_{\text{scale}} \approx 1 \text{ kpc}$, and hotter kinematics ($\sigma_z \simeq 40 \text{ km s}^{-1}$) while still retaining net prograde rotation. Inside $R \lesssim 2 \text{ kpc}$, the central bulge—partly bar-shaped—hosts both old, metal-rich stars and a younger, actively forming component; stellar motions there combine bar-driven streaming with high random velocities ($\sigma \sim 100 \text{ km s}^{-1}$). Encasing all of these is the stellar halo, which contributes only a few per cent of the total stellar mass yet includes the Galaxy’s oldest, most metal-poor stars ($[\text{Fe}/\text{H}] \lesssim -1.5$) on highly eccentric or even retrograde orbits. Its low density, rich substructure, and kinematics indicate an origin in the hierarchical accretion and tidal disruption of dwarf galaxies and globular clusters. The spatial distribution, chemistry, and dynamics of these four components include clues to the Milky Way’s formation history and its sequence of merger events.

1.2 Metallicity as a Cosmic Clock

Precise ages for individual old stars are difficult to measure, so their chemical composition, most commonly the iron-to-hydrogen ratio $[\text{Fe}/\text{H}]$, is used as a surrogate clock. Very metal-poor (VMP) stars must have formed before successive generations of Type II and Type Ia supernovae had substantially enriched the interstellar medium, and therefore have low $[\text{Fe}/\text{H}]$ values. Metallicity is inferred spectroscopically by measuring the equivalent widths of metal absorption lines, particularly Fe I and the Ca II H and K lines at 396.8 nm and 393.4 nm, respectively. After correcting for stellar effective temperature and surface gravity, the relative strengths of these lines yield elemental abundances [[Gray, 2005](#)]. Large surveys (for example APOGEE, GALAH, LAMOST, and the Gaia XP spectra) now provide such measurements

for millions of stars, enabling empirical age–metallicity relations that link chemistry to stellar chronometry [e.g. Nordström et al., 2004, Haywood et al., 2013, Leung and Bovy, 2019, Anders et al., 2023]. These studies consistently show that stars with $[\text{Fe}/\text{H}] \lesssim -1$ are typically older than ~ 10 Gyr, making low-metallicity populations valuable probes of the Milky Way’s earliest epochs.

1.3 Λ CDM Model

In the concordance Λ CDM model, galaxy-sized haloes assemble hierarchically: small dark-matter clumps form first and then merge to build larger structures. Cosmological N -body simulations show that the number of subhaloes of mass M follows $dN/dM \propto M^{-1.9}$ [Springel et al., 2008]. For a Milky-Way–sized halo, this implies that over the ~ 14 Gyr lifetime of the Universe, the halo experiences:

- $\sim 10^2$ **minor** accretion events with $M_{\text{sub}} \lesssim 10^9 M_\odot$, and
- a few **major** mergers with $M_{\text{sub}} \gtrsim 10^{10} M_\odot$.

Only a small fraction of dark-matter haloes form significant numbers of stars. Below a critical virial mass of $M_{\text{vir}} \sim 10^{11} M_\odot$, reionisation and stellar feedback strongly suppress the efficiency of star formation. As a result, the stellar-mass–halo-mass (SMHM) relation becomes steep at the low-mass end [Purcell et al., 2007, Bullock and Boylan-Kolchin, 2017]: most low-mass subhaloes remain effectively “dark”, while only a few relatively massive dwarfs produce observable stellar populations.

Thus, although the Milky Way has accreted hundreds of subhaloes over its history, just one or two of the most massive dwarf galaxies contribute the bulk of the stellar halo’s mass; the rest contribute primarily dark matter and dynamical substructure.

After accretion, dynamical friction drives massive satellites onto radial orbits, dispersing their debris into the inner halo. This debris retains coherent signatures—radial anisotropy, distinctive angular momenta, and narrow chemical abundance patterns [e.g. Helmi and Tim de Zeeuw, 2000]. Therefore, the stellar halo is not smooth but a fossil record of mergers: the inner halo reflects a few massive events (e.g. Gaia-Sausage/Enceladus), while the outer halo bears imprints of numerous minor ones.

1.4 Accreted halo and metal-poor disc origins

Chemical and kinematic evidence confirms that the metal-poor halo is primarily accreted. The debris of the Gaia–Sausage/Enceladus (GSE) event, for instance, is traced by stars with $-2 < [\text{Fe}/\text{H}] < -1$ and extreme orbital anisotropy ($\beta \gtrsim 0.8$; Belokurov et al. 2018, Helmi et al. 2018). At $[\text{Fe}/\text{H}] \lesssim -2$, an even broader mix of minor mergers emerges, removing any global rotation signal [Lancaster et al., 2019, Bird et al., 2021].

Several studies have uncovered stars in the range $-2 < [\text{Fe}/\text{H}] < -1$ with modest eccentricities and net prograde rotation [Norris et al., 1985, Chiba and Beers, 2000, Carollo et al., 2019, An and Beers, 2020], reminiscent of a disc. Gaia data have extended such findings to $[\text{Fe}/\text{H}] < -2$ [Sestito et al., 2019, Venn et al., 2020, Cordoni et al., 2020, Mardini et al., 2022]. Whether these stars reflect (i) a genuine *in-situ* metal-poor disc or (ii) spun-up debris from earlier accretion events remains debated.

Three broad formation scenarios have been proposed:

1. **Early *in-situ* disc.** Stars form in a gas-rich disc before $z \sim 4$ and are later heated or migrate outward, preserving the proto-Galaxy’s chemical signature.
2. **Proto-galactic building blocks.** Very metal-poor (VMP) stars originate in several massive, gas-rich satellites accreted at high redshift, with debris dragged into the disc plane as the gaseous disc settles [e.g. [Sestito et al., 2020](#)].
3. **Late, minor prograde mergers.** Low-mass satellites on aligned orbits are accreted after the disc forms, depositing a kinematically cold, metal-poor stellar component [[Santistevan et al., 2021](#)].

Cosmological simulations generally favour scenario 2: early mergers dominate the VMP star budget, while a coherent disc structure does not emerge until $z \lesssim 2$ [[Gurvich et al., 2023](#)].

Observationally, [Belokurov and Kravtsov \[2022\]](#) identified *Aurora*, a kinematically hot, weakly rotating population with $-2 \lesssim [\text{Fe}/\text{H}] \lesssim -1.3$, arguing against an extremely early disc. Follow-up work shows Aurora to be centrally concentrated [[Rix et al., 2022](#), [Arentsen et al., 2020a,b](#)], consistent with heated debris rather than a long-lived thin disc. Furthermore, secular processes such as bar–halo resonances can impart a modest prograde bias to halo stars, mimicking a disc signal [[Dillamore et al., 2023](#)].

1.5 This Work

By reproducing the work by [Zhang et al. \[2024\]](#), this study assesses the claim that the Milky Way hosts a very-metal-poor (VMP; $[\text{Fe}/\text{H}] < -2$) stellar disc. We use the catalogue from [Andrae et al. \[2023\]](#), which supplies six-dimensional phase–space coordinates—sky position, parallax–based distance, proper motions, and radial velocity—for each star, together with homogeneous metallicity and α -element abundances from the Gaia XP pipeline. To disentangle kinematic sub-populations we model the full three-dimensional velocity distribution (v_R, v_ϕ, v_z) with a Gaussian Mixture Model, in successive narrow metallicity bins, whose parameters are inferred via Extreme Deconvolution (XD). The XD formalism explicitly folds the individual distance and proper-motion uncertainties into the likelihood, ensuring that measurement noise does not bias the recovered velocity values.

Astrophysically, a genuine disc should manifest itself as a high-weight Gaussian centred around $v_\phi \simeq 200 \text{ km s}^{-1}$ with small tangential and vertical dispersions and negligible mean radial motion, whereas the halo or any heated component should appear as a broad, almost isotropic Gaussian with little net rotation and large velocity dispersions. By tracking how the weight of the cold, rotating component varies with metallicity we can determine when ordered rotation first emerged and test whether VMP stars were formed *in situ* or accreted from a satellite.

2 Data

2.1 Sample construction

The catalogue used in this work was the bright red–giant–branch sample of [Andrae et al. \[2023\]](#). Stellar metallicities were predicted with an Extreme Gradient Boosting model trained on high-resolution APOGEE DR17 spectra and a supplementary set of very metal-poor stars, ensuring reliable performance down to $[\text{M}/\text{H}] \simeq -3.5$ [[Andrae et al., 2023](#)]. For each of the 17.6 million giants the catalogue contains values of $[\text{M}/\text{H}]$, T_{eff} , and $\log g$ with a random uncertainty of

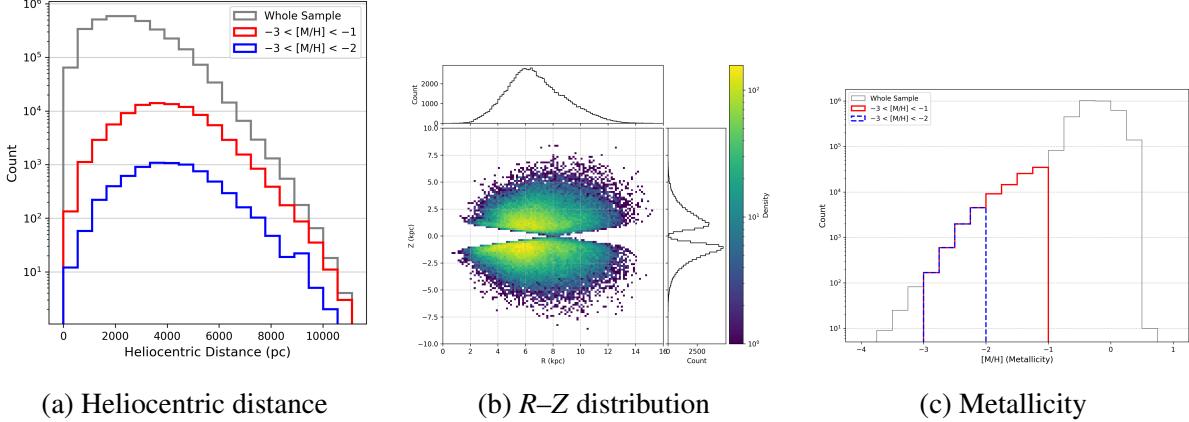


Figure 1: Properties of the final RGB sample after all quality and footprint cuts. *Left:* heliocentric-distance histogram for the whole sample (grey); the subsets with $-3 < [M/H] < -1$ and $-3 < [M/H] < -2$ are shown in solid red and dashed blue, respectively. *Middle:* density map in Galactocentric cylindrical coordinates. The empty band at low $|Z|$ is a selection artefact of our latitude/extinction cuts, which deliberately remove the thin-disc mid-plane; the concentration around $R \simeq 5\text{--}8$ kpc reflects the volume accessible to bright RGB stars interior to the Solar circle and coincides with the molecular ring region where the stellar surface density peaks. *Right:* Metallicity distribution of our data sample. Line colours are the same as in the left panel. (A reproduction of Figure 1 from [Zhang et al. \[2024\]](#))

$\simeq 0.1$ dex in $[M/H]$ at $G \lesssim 15$. The catalogue only retains entries flagged as “high-confidence” and lying in $-3.5 < [M/H] < +0.5$.

Astrometric positions, proper motions, and radial velocities come from the main *Gaia* DR3 tables [[Gaia Collaboration et al., 2023](#)], while heliocentric distances are adopted from the Bayesian photo-geometric catalogue of [Bailer-Jones et al. \[2021\]](#). Because accurate velocities scale with distance precision, following [Zhang et al. \[2024\]](#), a fractional-parallax-uncertainty cut $\sigma_{\varpi}/\varpi < 0.10$ was applied.

XP spectra are susceptible to reddening: heavy extinction dims the sources, lowers the XP signal-to-noise ratio, and biases the machine-learning metallicities. As described [Zhang et al. \[2024\]](#), to minimise such systematics, stars with $E(B-V)_{\text{SFD}} > 0.5$ or Galactic latitude $|b| < 10^\circ$ were excluded. Colour-excess values from the SFD map accessed via dustmaps [[Green, 2018](#)] was used for this. These criteria remove regions where dust corrections are large and spatially variable, at the expense of a loss of sky coverage losing a large number of stars at low galactic latitudes.

Field-star kinematics can also be skewed by dense sub-structures. Accordingly, following [Zhang et al. \[2024\]](#), we removed all objects lying within 1° of any known globular cluster or dwarf-galaxy satellite using the list and data of these compiled by [Pace \[2024\]](#). This step removes obvious non-field populations (e.g. cluster members and recent accretion debris) without significantly reducing the statistical power of the sample.

By following the data cuts of [Zhang et al. \[2024\]](#) of distance-quality, reddening, latitude, and sub-structure cuts, our working data set comprised of $\sim 3.4 \times 10^6$ red-giant stars possessing homogeneous metallicities and six-dimensional phase-space information.

2.2 Galactocentric Positions and Velocities

Six-dimensional phase-space coordinates were obtained with `astropy.coordinates`. We adopted a Galactocentric frame with $R_0 = 8.1$ kpc and $Z_0 = 25$ pc [McMillan, 2016], and a solar velocity¹ $(U_\odot, V_\odot, W_\odot) = (11.1, 245, 7.25)$ km s⁻¹ [Schönrich et al., 2010]. The cylindrical velocity components (v_R, v_ϕ, v_Z) were extracted using the transformed SkyCoord module.

To propagate measurement errors we generated $N_{MC} = 100$ Monte-Carlo realisations per star, drawing parallax, proper motions, radial velocity, and distance from their reported uncertainties. Each realisation was transformed to the Galactocentric frame, yielding distributions of v_R , v_ϕ , and v_Z ; the 1σ widths of those distributions were stored as per-star velocity uncertainties.

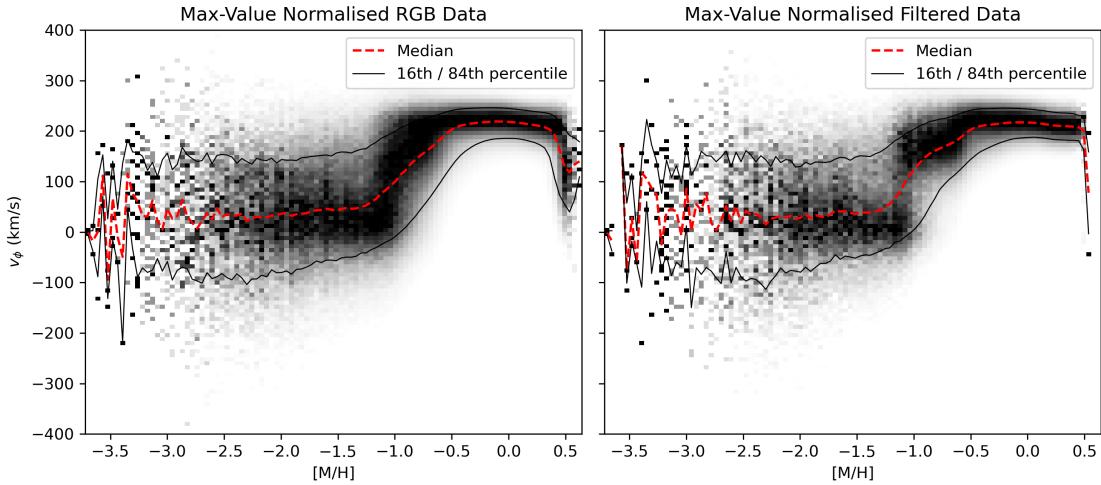


Figure 2: Column-normalised density in the v_ϕ –[M/H] plane. *Left:* the bright-RGB catalogue of Andrae et al. [2023]. *Right:* the same sample after all distance, dust, and quality cuts. Greyscale pixels show the normalised counts in each metallicity bin; the red dashed curve is the median v_ϕ , and the black curves trace the 16th and 84th percentiles. (A reproduction of Figure 2 from Zhang et al. [2024])

As shown in Figure 2, stellar azimuthal velocities evolve increase with metallicity. Halo-like kinematics dominate at $[M/H] \lesssim -1.5$ with statistics indicative of a pressure-supported component. A rapid spin-up appears over the metallicity range $[M/H] \simeq -1.3$ to -0.9 , consistent with Belokurov and Kravtsov [2022]. By $[M/H] \gtrsim -0.5$ the stellar azimuthal velocities reaches the Local Standard-of-Rest value (≈ 220 km s⁻¹) and the velocity dispersion dispersion falls, marking the transition to a disc. Visually, using 2, we can observe that the onset of ordered rotation in the Milky Way occurred when the inter-stellar medium reached roughly one-tenth solar metallicity.

As shown in Fig. 3, stars with $[M/H] \geq -1.0$ have a strong prograde bias in azimuthal velocities, with a peak at $v_\phi \gtrsim 180$ km s⁻¹ and a relatively narrow distribution in v_R . As shown in Fig. 4, this is consistent with the thin- and thick-disc ellipses. Below $[M/H] \simeq -1.0$ the distribution broadens and drops toward $v_\phi \approx 0$, indicating pressure supported kinematics, characteristic of the stellar halo and the radial Gaia–Sausage/Enceladus debris. At the lowest metallicities ($[M/H] \lesssim -1.5$) the contours are nearly isotropic with only a mild prograde bias. Hence,

¹Cartesian components (U, V, W) , where U is radially outwards, V is aligned with Galactic rotation, and W points to the North Galactic Pole.



Figure 3: Galactocentric velocity distributions as a function of metallicity. Each panel shows the column-normalised density of stars in the v_R – v_ϕ plane for the metallicity interval printed at the top.(A reproduction of Figure 3 from [Zhang et al. \[2024\]](#))

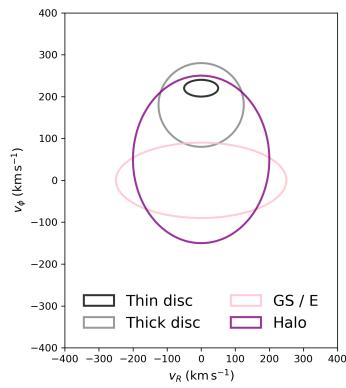


Figure 4: Ellipses mark the approximate extent of the thin disc (black), thick disc (gray), Gaia–Sausage/Enceladus debris (pink) and the pressure-supported stellar halo (purple). The figure should be used as a visual key for interpreting the data panels in Fig. 3.

any rotation-supported very-metal-poor disc must contribute at most a small fraction of the population. In subsequent analysis we quantitatively assess these observations.

Given we are testing whether a very-metal-poor stellar disc exists, following [Zhang et al. \[2024\]](#), we restrict the sample to stars within $|Z| < 2.5$ kpc of the galactic mid-plane. This keeps the focus on stars close to the plane, where any disc-like population (whether formed in situ or deposited by mergers) would be found [[Tkachenko et al., 2025](#)].

Upon comparison with [Zhang et al. \[2024\]](#), we note that our sample is approximately the same size - with a difference of less than 0.002% of the sample size in each bin.

3 Methodology

3.1 Gaussian Mixture Model

As in [Zhang et al. \[2024\]](#), to quantitatively investigate the kinematic structure of metal-poor stars and assess the presence of a potential very-metal-poor disc, we use a Gaussian Mixture Model (GMM) framework. GMMs are a class of unsupervised machine learning algorithms commonly used in data science for clustering and density estimation. They model a dataset as a weighted sum of multivariate Gaussian distributions, each corresponding to a latent sub-population within the data.

From a probabilistic perspective, the GMM assumes that each observed data point is generated from a hidden (latent) variable indicating membership in one of the Gaussian components. This latent space formulation allows the model to assign probabilistic classifications to data points, providing a soft clustering where each star has a fractional likelihood of belonging to each component. In our case, the data is the three-dimensional velocity vectors (v_R, v_ϕ, v_Z) , and the latent space captures distinct kinematic substructures in this space.

We implement the GMM fitting using the pyGMMIs package [[Melchior and Goulding, 2016](#)] in the v_R , v_ϕ , and v_Z phase space. This algorithm extends the standard Expectation-Maximisation (EM) algorithm with the “Extreme Deconvolution” technique developed by [Bovy et al. \[2011\]](#). This method is particularly well-suited to astronomical datasets, as it incorporates measurement uncertainties into the GMM fitting by modifying the EM updates to account for known errors on each data point. In our case, these uncertainties are derived from the Gaia astrometric and spectroscopic data and are represented by diagonal covariance matrices encoding the squared uncertainties in v_R , v_ϕ , and v_Z for each star.

3.2 Number of Gaussian Components

We apply the GMM separately to each metallicity bin with $[M/H] < -1$, as we are specifically interested in detecting any rotationally supported structure among the very metal-poor stars. An important decision must be made when implementing the appropriate number of gaussian components. Too few components may underfit the data, missing real substructures, while too many will overfit, leading to physically uninterpretable results.

We select the optimal number of components using the Bayesian Information Criterion (BIC) [Schwarz 1978](#). The BIC is defined as

$$\text{BIC} = k \ln(n) - 2 \ln \mathcal{L}, \quad (1)$$

where $k = (1 + 3 + 6) \times N - 1$ is the total number of free parameters in a model with N Gaussian components (accounting for the weights, means, and covariances), n is the number of stars in the sample, and \mathcal{L} is the maximum likelihood of the fit. The BIC penalizes model complexity, such that adding extra components without a significant gain in likelihood will result in a higher BIC value.

The Expectation-Maximisation algorithm can become trapped in local minima, so we performed 50 initialisations for each N and recorded the BIC value for each trial. The minimum BIC value across all trials indicated the statistically preferred model and suggests that the global optimum was likely reached.

To improve convergence and model stability, the GMM components in each trial were initialised using the kmeans algorithm. This is due to the sensitivity of GMMs to their starting conditions: poor initialisation can lead to convergence on undesirable solutions, particularly in high-dimensional or overlapping data. KMeans clustering provides an initialisation point by partitioning the velocity space into compact, roughly spherical clusters. This works well with the assumptions of Gaussian components and often gives faster, more stable convergence and more physically meaningful results. In our application, where stellar substructures are partially overlapping in velocity space, this method applies well.

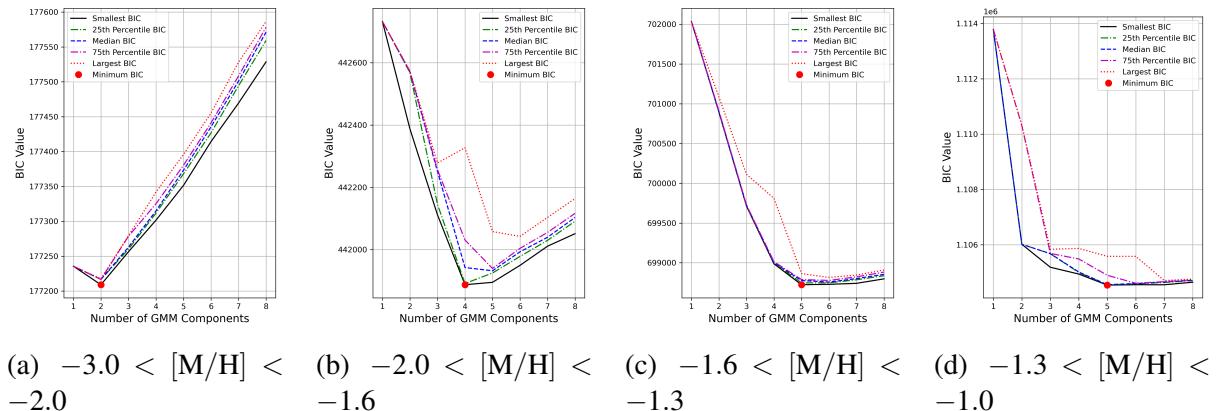


Figure 5: BIC distributions as a function of the number of GMM components in each metallicity bin. The optimum number of components is indicated by the minimum BIC value (highlighted in red). (A reproduction of Figure 4 from [Zhang et al. \[2024\]](#))

In Figure 5, we show the distribution of BIC values across four metallicity bins as a function of N . The resulting preferred number of components are 2, 4, 5, and 5 for the VMP, IMP, MP1, and MP2 bins. This is in agreement with the findings of [Zhang et al. \[2024\]](#). Our analysis, however, selects the lowest BIC for each bin, removing the need for a subjective choice of the number of components when BIC values are very similar. [Zhang et al. \[2024\]](#), when observing the BIC values for the MP1 bin chooses to use 5 components instead of 6 despite the BIC score for the 6-component model being lower than the 5-component model.

4 Results

4.1 Gaussian Mixture Model Fit

Using the number of components selected by the BIC criteria, we fit the GMM to the data in each metallicity bin as shown in Figure 6. We increased the number of initialisations to 100 for

the final fitting to ensure convergence to a stable solution.

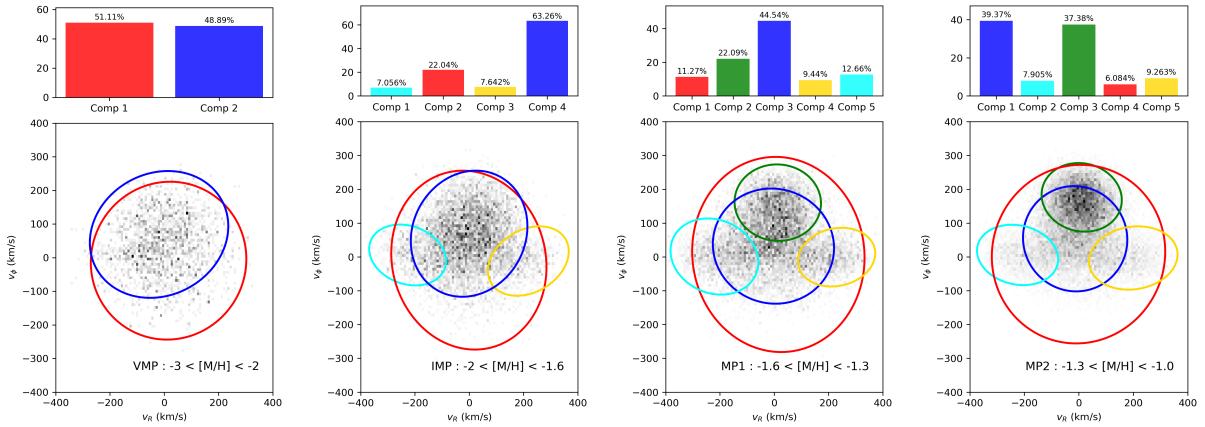


Figure 6: Gaussian Mixture Model decompositions of the stellar velocity distribution in the $v_R - v_\phi$ plane for each metallicity bin. The bottom panel of each subfigure shows the 2D velocity distribution with GMM components overplotted as ellipses representing the 1σ contours in $v_R - v_\phi$. The top panel shows the fractional contribution. (A reproduction of Figure 5 from [Zhang et al. \[2024\]](#))

As shown in Figure 6, the GMM is able to capture kinematic substructures in the velocity space of metal-poor stars. As metallicity increases, a gaussian component with a clear prograde rotation signal emerges (in green), indicative of a disc-like structure. The statistics of the GMM fit are summarised in Table 1, which lists the weights, means, and dispersions of each component in the four metallicity bins.

See the fully interactive 3-D plots for each bin via RawGitHack:

VMP ($-3 < [\text{M}/\text{H}] < -2$): [VMP link](#),

IMP ($-2 < [\text{M}/\text{H}] < -1.6$): [IMP link](#),

MP1 ($-1.6 < [\text{M}/\text{H}] < -1.3$): [MP1 link](#),

MP2 ($-1.3 < [\text{M}/\text{H}] < -1.0$): [MP2 link](#).

At $[\text{M}/\text{H}] < -1.6$ (the VMP and IMP regimes), a disc component is not observed. Instead, the kinematic structure is dominated by a prograde halo and a stationary halo. In the VMP bin, the sample is almost evenly split between these two components, with the stationary halo contributing 51.1% and the prograde halo 48.9%. Both show broad velocity dispersions, with only modest rotational support ($\overline{v_\phi} \sim 69 \text{ km s}^{-1}$ for the prograde halo) and no evidence for a kinematically cold disc-like structure.

In the IMP bin, the picture remains similar, although the prograde halo dominates more strongly (63.3%) and additional substructure becomes apparent. Two minor components, interpreted as fragments of the Gaia–Sausage/Enceladus (GS/E) merger debris, are also identified, each contributing $\sim 7\%$ of the population and exhibiting highly radial orbits ($|v_R| > 200 \text{ km s}^{-1}$). These GS/E components are characterised by strong anisotropy and contribute to the broadening of the halo distribution, but again show no signature of disc-like kinematics. The absence of a cold, rotating disc in these bins suggests that if a very-metal-poor disc exists, it must be a minor contributor to the local stellar population.

Components	Weights (%)	v_R	σ_R	v_ϕ	σ_ϕ	v_Z	σ_Z
VMP: $-3.0 < [\text{M}/\text{H}] < -2.0$ (4768 stars)							
Stationary halo	51.1	15.86	143.17	-9.13	117.36	-0.08	122.64
Prograde halo	48.9	-19.14	127.43	68.86	94.28	-0.90	82.89
IMP: $-2.0 < [\text{M}/\text{H}] < -1.6$ (12052 stars)							
Stationary halo	22.0	-1.45	142.88	-9.59	132.32	-0.91	124.92
Prograde halo	63.3	-0.58	107.32	68.52	93.27	-1.16	72.56
GS/E(1)	7.1	-226.92	70.78	5.60	44.98	8.43	88.40
GS/E(2)	7.6	217.52	74.41	-12.35	51.35	-4.22	89.78
MP1: $-1.6 < [\text{M}/\text{H}] < -1.3$ (19142 stars)							
Stationary halo	11.3	15.09	158.68	7.11	144.33	-2.55	132.08
Prograde halo	44.5	-3.02	111.54	31.79	85.15	-0.54	70.36
GS/E(1)	12.7	-220.40	80.76	1.01	56.29	0.47	91.00
GS/E(2)	9.4	229.32	71.22	-0.82	43.42	1.42	92.69
Thick disc	22.1	12.93	79.56	160.32	56.58	-2.31	68.83
MP2: $-1.3 < [\text{M}/\text{H}] < -1.0$ (30892 stars)							
Stationary halo	6.1	-1.17	159.06	8.07	132.05	-3.37	120.46
Prograde halo	39.4	-13.47	95.65	53.64	77.86	-3.04	71.00
GS/E(1)	7.9	-224.68	74.44	5.73	44.63	2.41	87.27
GS/E(2)	9.3	199.71	81.28	-3.30	46.84	-2.07	88.99
Thick disc	37.4	10.56	73.55	176.07	50.79	0.85	62.03

Table 1: Parameters of the Gaussian mixture model fittings in different metallicity bins. The unit for all velocity columns is km s^{-1} .

Our GMM decomposition recovers the same substructures identified by [Zhang et al. \[2024\]](#) across all metallicity bins. We find similar velocity dispersions and centroid trends for the stationary halo, prograde halo, and GS/E components. Notably, our model assigns a more comparable weight to the prograde halo component in the VMP bin, in contrast to the dominant stationary halo reported by [Zhang et al. \[2024\]](#). We also observe some variation in the centroid v_ϕ values of GS/E substructures, although their bimodal spatial structure is preserved.

4.2 Rotational Support and the Onset of Disc Formation

To assess the degree of rotational support in each component, we compute the ratio $V_{\text{rot}}/\sigma_\phi$, where V_{rot} is the mean azimuthal velocity and σ_ϕ its dispersion. This ratio serves as a conventional diagnostic of disc-like kinematics, with values $\gtrsim 3$ indicating rotation-supported populations. As shown in Figure 7, all components in the VMP and IMP bins fall below this threshold, confirming that the GMM identifies only dynamically hot, dispersion-supported structures in these regimes.

At higher metallicities, a dynamically colder disc population emerges. In the MP1 bin, the thick disc component (green) rotates at $v_\phi \sim 160 \text{ km s}^{-1}$ with $\sigma_\phi \sim 57 \text{ km s}^{-1}$, yielding $V_{\text{rot}}/\sigma_\phi \approx 2.8$. While just below the canonical boundary, this component's low vertical dispersion and substantial weight (22.1%) mark it as a distinct disc-like structure. By the MP2 bin, the disc becomes the dominant component (37.4%), with $V_{\text{rot}}/\sigma_\phi \approx 3.5$, indicating clear rotational support and a well-established thick disc. These trends suggest a transition in the stellar kinematics around $[\text{M}/\text{H}] \sim -1.5$.

Meanwhile, the GS/E components (identifiable in the IMP, MP1, and MP2 bins) remain dynamically distinct, with high $|v_R|$, relatively small v_ϕ , and negligible rotational support, reinforcing their origin as debris from a major radial merger [Helmi et al., 2018]. Notably, given the estimated time of the GS/E accretion event (of order 8–11 Gyr ago [Gallart et al., 2019, Belokurov et al., 2020, Di Matteo et al., 2019]), its stellar debris is expected to be phase-mixed. This implies that the positive and negative v_R components should contribute approximately equally. This symmetry is observed in the IMP and two MP bins.

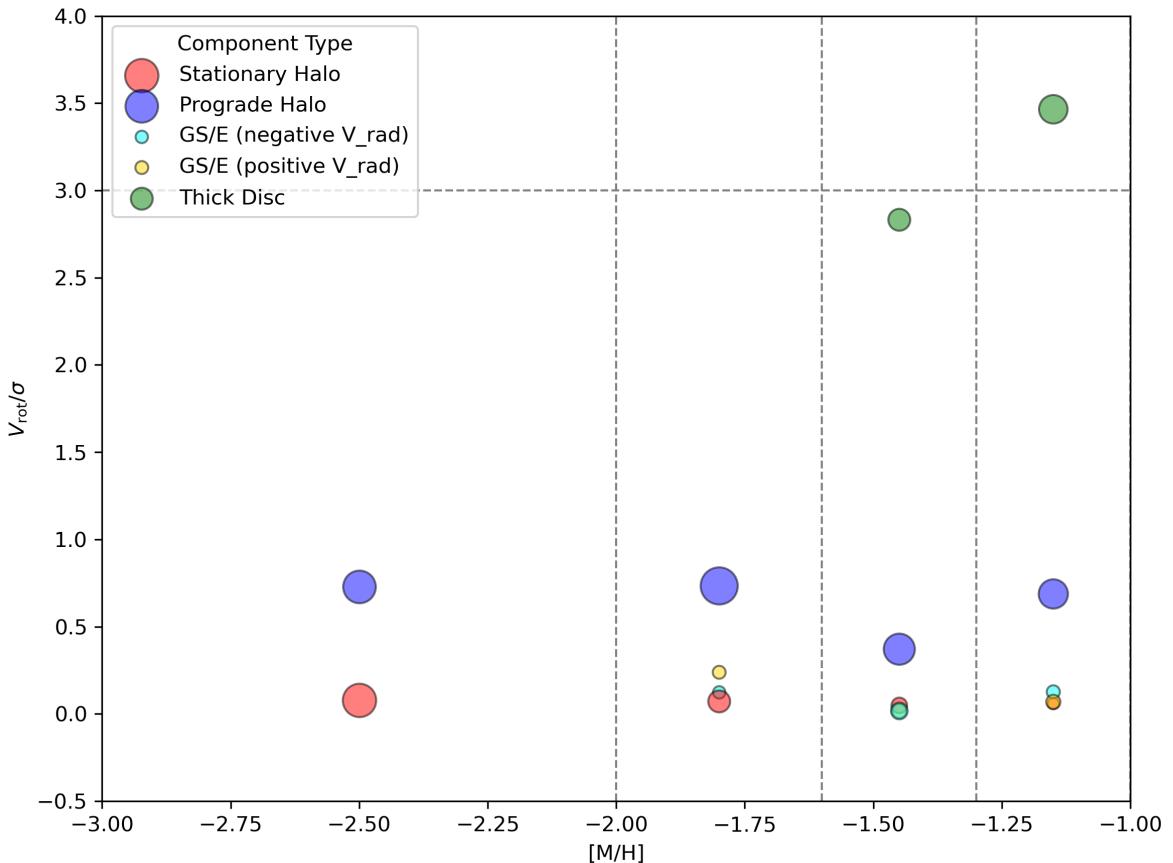


Figure 7: Ratio of mean rotational velocity to azimuthal velocity dispersion, V_{rot}/σ_ϕ , for individual GMM components in each metallicity bin. Larger circles correspond to more dominant components. Colours match the GMM components in Figure 6.(A reproduction of Figure 6 from Zhang et al. [2024])

4.3 Residual Analysis for Hidden Disc Populations

One of the limitations of GMMs is that they can fail to detect substructures that contribute only weakly to the overall distribution. To investigate this, we performed a residual analysis to test whether a disc-like population, too weak to be picked up by the GMM, might be present in the VMP and IMP metallicity regimes.

We generated a synthetic dataset by drawing the same number of mock stars as observed stars from the best-fit GMM in each bin. Since our GMM uses the Extreme Deconvolution algorithm, which accounts for observational uncertainties during fitting, the raw samples from the model are noise-free. However, comparing these directly to real data would be inappropriate

due to the absence of measurement error. To resolve this, we assign each mock star the velocity uncertainties of its nearest neighbour in the observed data (in (v_R, v_ϕ, v_Z) space), and then add Gaussian noise according to these uncertainties.

We then bin both the observed and mock data in the (v_R, v_ϕ) plane and compute the normalized residual map, defined as

$$H_{\text{residual}} = \frac{H_{\text{obs}} - H_{\text{mock}}}{H_{\text{obs}} + H_{\text{mock}} + \epsilon},$$

where H_{obs} and H_{mock} are the 2D histograms and $\epsilon = 10^{-5}$ is added to avoid division by zero. The residual maps are shown in the top row of Figure 8 for the VMP and IMP bins. The grey dashed ellipses highlight a 2σ region of a hypothetical thick disc with mean velocity $(v_R, v_\phi) = (0, 180) \text{ km s}^{-1}$ and dispersions $(\sigma_R, \sigma_\phi) = (70, 50) \text{ km s}^{-1}$, consistent with the thick disc component seen in the MP2 metallicity bin.

To quantify the possible presence of a hidden disc, we count the excess number of observed stars inside this thick disc ellipse relative to the GMM-generated sample. We repeat this calculation for 1000 Monte Carlo realizations to estimate the residual and its uncertainty. The results are shown as horizontal dashed lines in the bottom row of Figure 8, representing a disc-like residual of 38.3 ± 27.1 stars in the VMP bin and 78.1 ± 45.1 stars in the IMP bin.

To interpret these residuals, we repeat the same procedure but now inject a mock disc population—generated from the same thick disc Gaussian—into the GMM baseline. By varying the injected disc fraction from 0% to 5% of the sample and recalculating the residual each time, we construct a relationship between disc fraction and expected residual. This is shown as the solid black line in the bottom panels of Figure 8, with the red shaded region indicating the 1σ uncertainty across Monte Carlo trials.

In the bottom panels of Figure 8, the horizontal blue band shows the observed excess within the thick-disc ellipse ($\pm 1\sigma$), and the solid black curve (with red shading) gives the expected residual when injecting a known disc fraction into the GMM. Their intersection could imply a hidden disc fraction of up to $\sim 1\%$ in the VMP bin and IMP bin. These estimates are indicative only: measurement uncertainties, GMM modelling choices, and Gaia-selection cuts mean that the true disc contribution in the parent VMP/IMP populations may differ [Zhang et al., 2024].

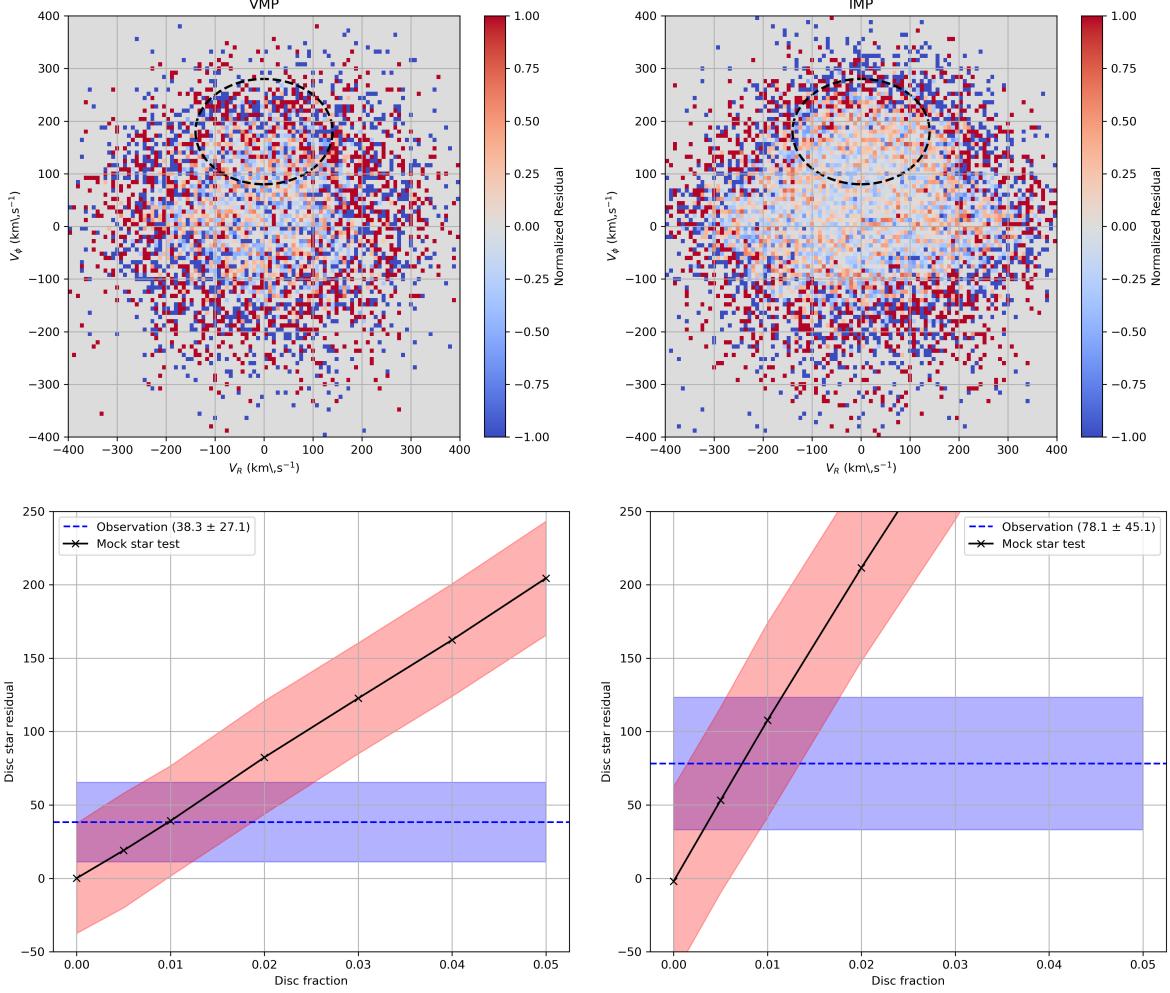


Figure 8: Top: Normalised residual maps between observed and GMM-predicted velocity distributions in (v_R, v_ϕ) for the VMP (left) and IMP (right) metallicity bins. Grey ellipses show the 2σ contour of the fiducial thick disc model. Bottom: Disc residuals as a function of injected disc fraction. The dashed line and blue band indicate the observed residual and its uncertainty. The solid black line and red region show mock test results and their uncertainties. (A reproduction of Figure 7 from [Zhang et al. \[2024\]](#))

5 Extension Analysis

We extend the work of [Zhang et al. \[2024\]](#) by adding the α -element abundance, $[\alpha/\text{Fe}]$, as a fourth dimension in our XD-GMM fits. Following [Viswanathan et al. \[2024\]](#), each metallicity bin is split into “high- α ” and “low- α ” subsets (using XP abundances from [Li et al. 2023](#)), and the GMM is refit independently for each subset. This chemo-kinematic dissection allows us to test whether the kinematic substructures in the very-metal-poor regime vary with α -element enrichment, potentially distinguishing in situ from accreted populations.

5.1 Alpha-Element Abundance

The ratio $[\alpha/\text{Fe}]$ traces the relative contributions of two supernova types: core-collapse (Type II) and thermonuclear (Type Ia). Type II supernovae result from the deaths of massive stars

($\gtrsim 8 M_{\odot}$) and explode on short timescales (~ 10 – 30 Myr) after star formation. They produce large amounts of α -elements (O, Ne, Mg, Si, S, Ca, Ti), which are rapidly injected into the interstellar medium. In contrast, Type Ia supernovae arise from white dwarfs in binary systems, requiring hundreds of millions to billions of years to explode ($\gtrsim 0.1$ – 1 Gyr). These events primarily enrich the interstellar medium with iron and iron-peak elements (e.g. Fe, Ni, Co).

Because the two channels operate on different timescales, the $[\alpha/\text{Fe}]$ ratio encodes a population’s formation history:

- **High- $[\alpha/\text{Fe}]$:** Rapid, intense star formation enriches the gas with α -elements before Type Ia supernovae contribute significant iron, leading to elevated $[\alpha/\text{Fe}]$. These stars typically belong to the in-situ thick disc, stellar halo, and ancient accreted debris. At fixed $[\text{M}/\text{H}]$, accreted populations (e.g. dwarf galaxies) generally show lower $[\alpha/\text{Fe}]$ than in-situ stars [Helmi et al., 2018], reflecting slower chemical evolution.
- **Low- $[\alpha/\text{Fe}]$:** A more extended star-formation history allows iron from Type Ia supernovae to accumulate, diluting the α enrichment. This regime includes the thin disc, metal-rich bulge, and remnants of dwarf satellites.

Hence, in the $[\alpha/\text{Fe}]$ – $[\text{M}/\text{H}]$ plane, populations that overlap in metallicity but formed on different timescales separate cleanly in α enrichment. By fitting separate XD-GMMs to the high- and low- α sequences within each metallicity bin, we can isolate how disc-like rotation and halo substructures depend on chemical evolution history.

5.2 Data

We obtain the α -to-metal abundance ratios from the XP-based catalogue of Li et al. [2023]. Following Viswanathan et al. [2024], we cross-match the XP-based α catalogue of Li et al. [2023] with the metallicity-selected red-giant branch sample. This base sample already incorporates the red-giant selection and photometric quality cuts from Andrae et al. [2023], including limits on surface gravity, colour–magnitude position, and XP calibration reliability.

To further ensure robust α abundances and kinematics, we apply several additional cuts:

- we reject stars within 1° of known dwarf galaxies and those associated with globular clusters;
- a reddening cut $E(B-V) < 0.5$ removes stars in regions of high extinction, where XP abundances are less reliable;
- a Galactic latitude cut $|b| > 10^{\circ}$ similarly avoids the dust-dominated plane;
- we relax the fractional parallax uncertainty cut relative to Section 2.1, requiring only $fpu < 0.2$ to retain sufficient stars in the metal-poor regime, which is the focus of this analysis.

We obtain approximately 3.4 million stars in our final sample with alpha-element abundances.

Figure 9 shows the logarithmic density of our RGB sample in the $[\alpha/\text{M}]$ – $[\text{M}/\text{H}]$ plane. Following Chandra et al. [2024], but with a deliberately tighter boundary as in Viswanathan et al. [2024], we separate the sample into high- α and low- α sequences using the following piecewise definition:

$$\text{high-}\alpha : \begin{cases} [\text{M}/\text{H}] < -0.60 & [\alpha/\text{M}] > 0.28 \\ -0.60 \leq [\text{M}/\text{H}] \leq 0.125 & [\alpha/\text{M}] > -0.25[\text{M}/\text{H}] + 0.13 \\ [\text{M}/\text{H}] > 0.125 & [\alpha/\text{M}] > 0.10 \end{cases}$$

$$\text{low-}\alpha : \begin{cases} [\text{M}/\text{H}] < -0.80 & [\alpha/\text{M}] < 0.21 \\ -0.80 \leq [\text{M}/\text{H}] \leq 0.07 & [\alpha/\text{M}] < -0.21[\text{M}/\text{H}] + 0.04 \end{cases}$$

This high- α threshold is raised so that the chemistry of the Gaia–Sausage/Enceladus debris [Belokurov et al., 2018, Helmi et al., 2018] does not leak into the in-situ sequence; the exclusion band between the branches is widened to maximise purity in both sub-samples [Viswanathan et al., 2024].

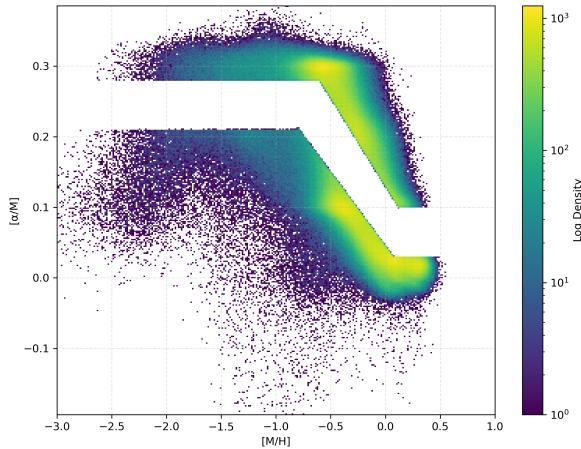
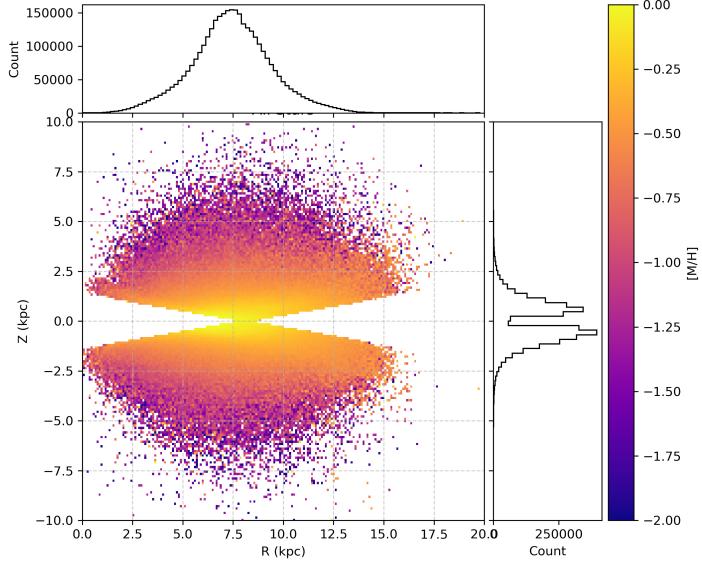
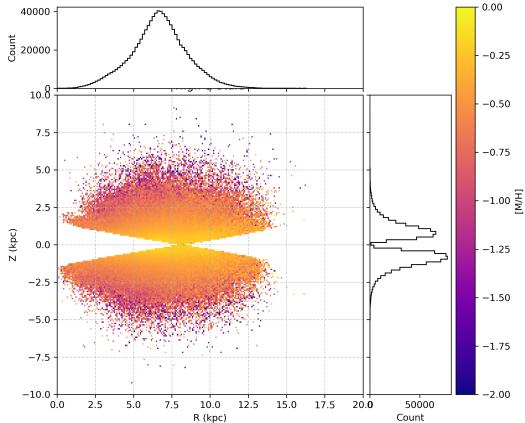


Figure 9: Density of stars in the $[\alpha/\text{M}]$ – $[\text{M}/\text{H}]$ plane. The high- α and low- α sequences are separated by the piecewise linear boundary described in the text. The transition zone between the two sequences is clearly removed from this plot and discarded from subsequent analysis. (A reproduction of Figure 1 in Viswanathan et al. [2024])

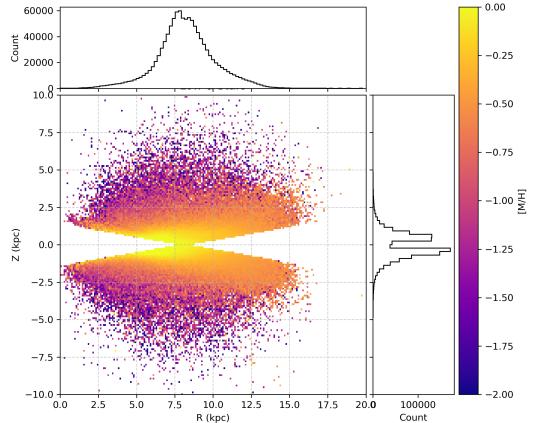
Figure 10 shows the distribution of our sample in cylindrical Galactocentric radius R versus height above the midplane z , colour-coded by the mean metallicity in each pixel. The top panel presents this distribution for the full RGB sample, while the bottom panels split the data into high- α (left) and low- α (right) subsamples.



(a) All RGB stars



(b) High- α subsample



(c) Low- α subsample

Figure 10: Mean $[M/H]$ colour-coded in the Galactocentric $R-z$ plane. *Top:* the complete $G < 16$ RGB sample. *Bottom left:* high- α stars; *bottom right:* low- α stars. The high- α population is more centrally concentrated and shows a shallower vertical metallicity gradient. (A reproduction of Figure 2 in [Viswanathan et al. \[2024\]](#))

The high- α population is more centrally concentrated and spans a smaller range in R and z than the low- α stars. The broader vertical extent of the low- α sequence likely reflects contributions from accreted material at large scale heights. A steep negative metallicity gradient with increasing z is evident in the low- α sample and also visible in the full sample, likely due to the dominance of low- α stars in the overall dataset. In contrast, the high- α population exhibits a shallower vertical metallicity gradient and a thicker disc structure. We also observe a metallicity gradient towards larger R in the low- α disc, indicative of radial migration and consistent with flaring commonly seen in the Galactic thin disc [e.g. [Haywood et al., 2013](#), [Ratcliffe et al., 2023](#)].

To trace how rotational support evolves with metallicity and α -element enrichment, following [Viswanathan et al. \[2024\]](#) we examine the distribution of azimuthal velocity (v_ϕ) as a function

of $[M/H]$. Figure 11 presents this relation for (i) the full RGB sample, and (ii) the chemically separated high- α and low- α subsamples derived in Section 5.1. Each panel shows the column-normalised density in the $[M/H]-v_\phi$ plane together with running medians (red dashed) and 16th/84th percentiles (black).

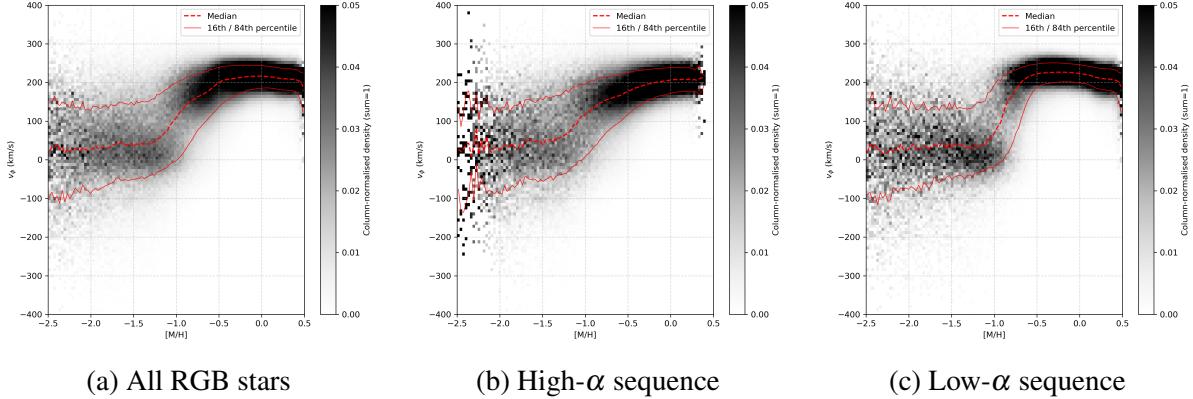


Figure 11: Azimuthal velocity versus metallicity. Left: full RGB sample. Centre: high- α stars. Right: low- α stars. Contours show running medians (red) and the 16th/84th percentiles (black). (Adapted from Fig. 3 of [Viswanathan et al. 2024](#).)

The high- α giants exhibit a gradual spin-up with increasing metallicity. At low metallicities ($[M/H] \lesssim -2$), the population shows mild net prograde rotation with large dispersion, consistent with a dynamically hot stellar halo. From $[M/H] \gtrsim -1.3$, there is a steady increase in rotation velocity and a corresponding decrease in dispersion, indicative of a transition toward the thick disc. By $[M/H] \gtrsim -0.5$, most stars occupy a kinematic regime characterised by elevated v_ϕ and moderate dispersion, consistent with a dynamically hot but rotationally supported thick disc. This continuous trend supports a scenario in which the Galaxy’s in-situ population gradually acquired angular momentum as star formation progressed.

In the low- α panel, stars with $[M/H] \lesssim -1.2$ exhibit kinematics similar to those in the high- α sequence, showing mild prograde rotation and high velocity dispersion. However, above $[M/H] \gtrsim -1.2$, the low- α sequence undergoes a sharper transition to a population with high net prograde rotation and lower velocity dispersion. By $[M/H] \gtrsim -0.5$, the low- α population exhibits kinematics consistent with the thin disc: a strongly rotating, dynamically colder component with lower dispersion than the high- α sequence.

In summary, the high- α population shows a gradual evolution from halo to thick disc, whereas the low- α population exhibits a more abrupt shift to a dynamically cold disc at higher metallicities.

Following the method described in Section 2.2, we obtain full six-dimensional phase-space coordinates for the chemically separated high- and low- α subsamples. Figure 12 shows the distribution of stars in the v_R-v_ϕ plane across metallicity bins for the high- α (top) and low- α (bottom) populations.

The low- α sample exhibits a strong rotationally supported component, particularly at higher metallicities ($[M/H] \gtrsim -1.0$), with a narrow v_ϕ distribution centred around ~ 200 km/s and low radial velocity dispersion—clear hallmarks of thin-disc kinematics. As metallicity decreases, this rotational component becomes less pronounced, with broader distributions in both v_R and

v_ϕ appearing at $[M/H] \lesssim -1.5$, consistent with increased contamination from halo or accreted populations.

The high- α stars, by contrast, show a continuous transition from a kinematically hot, low-rotation regime at low metallicities to a more rotation-supported disc at higher metallicities. This reflects the gradual build-up of the thick disc over time, as also noted in [Chandra et al. \[2024\]](#). The velocity dispersion is broader throughout, indicating that high- α stars trace older, more dynamically heated populations.

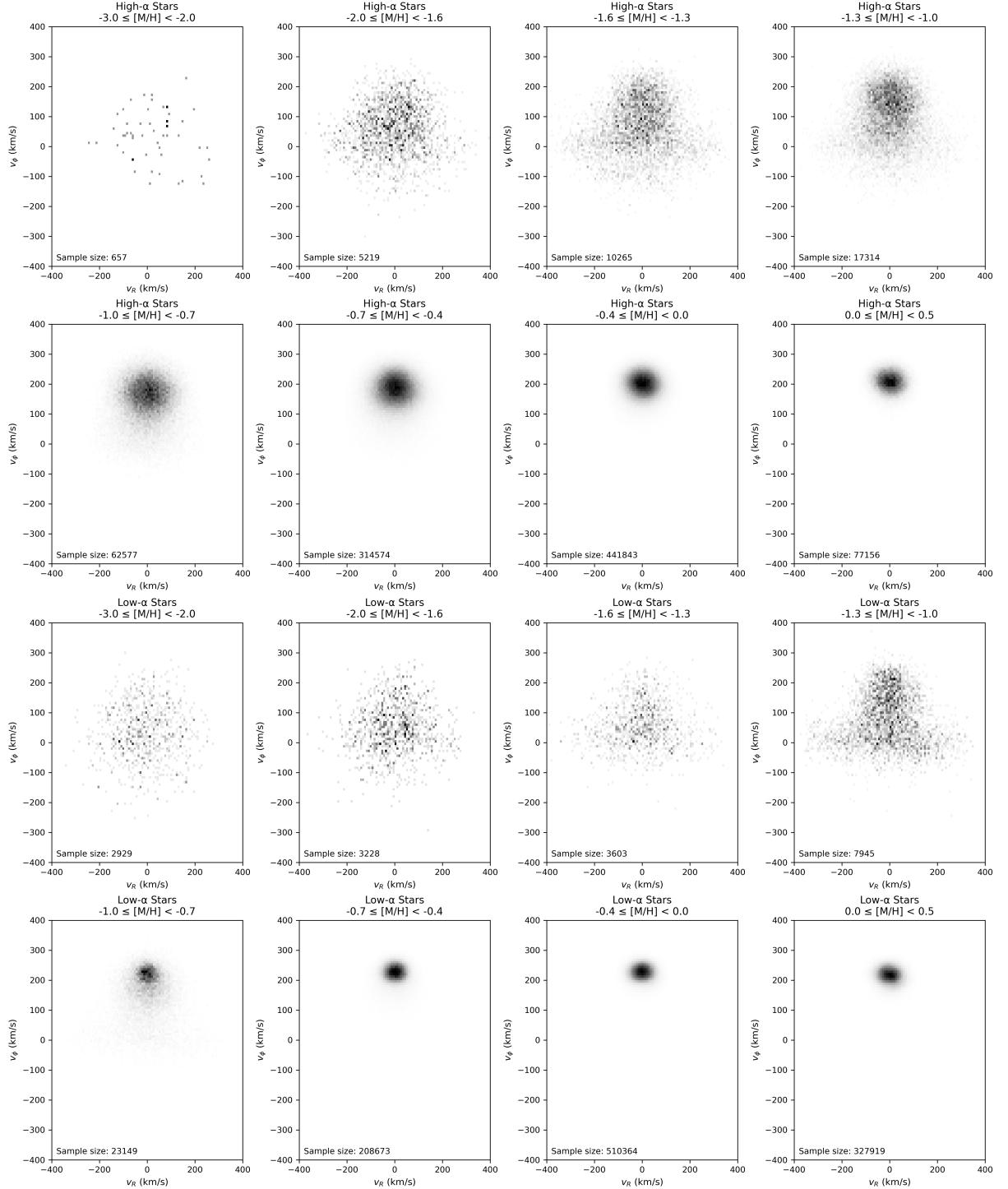


Figure 12: Distribution of stars in $v_R - v_\phi$ space across eight metallicity bins (increasing left to right, top to bottom) for the high- α population (top) and low- α population (bottom). The low- α stars are clearly more rotation-supported at high metallicities, while the high- α sample shows a broader velocity structure across all bins, consistent with thick-disc and halo populations.

5.3 Gaussian Mixture Model Fit

We follow the same Gaussian mixture model fitting procedure as in Section 4.1, but now separately for the high- and low- α sequences.

α -Sequence	VMP ($[M/H] < -2.0$)	IMP ($-2.0 < [M/H] < -1.3$)	MP1 ($-1.3 < [M/H] < -1.0$)	MP2 ($-1.0 < [M/H] < -0.7$)
High- α	1	4	5	3
Low- α	1	2	3	5

Table 2: Number of Gaussian Mixture components selected by the BIC for each metallicity bin, split by α -sequence.

Following the methodology described in Section 3.2, Table 2 summarises the number of components selected by the Bayesian Information Criterion for each metallicity bin. Detailed BIC scores for each metallicity bin are shown in Figure 13.

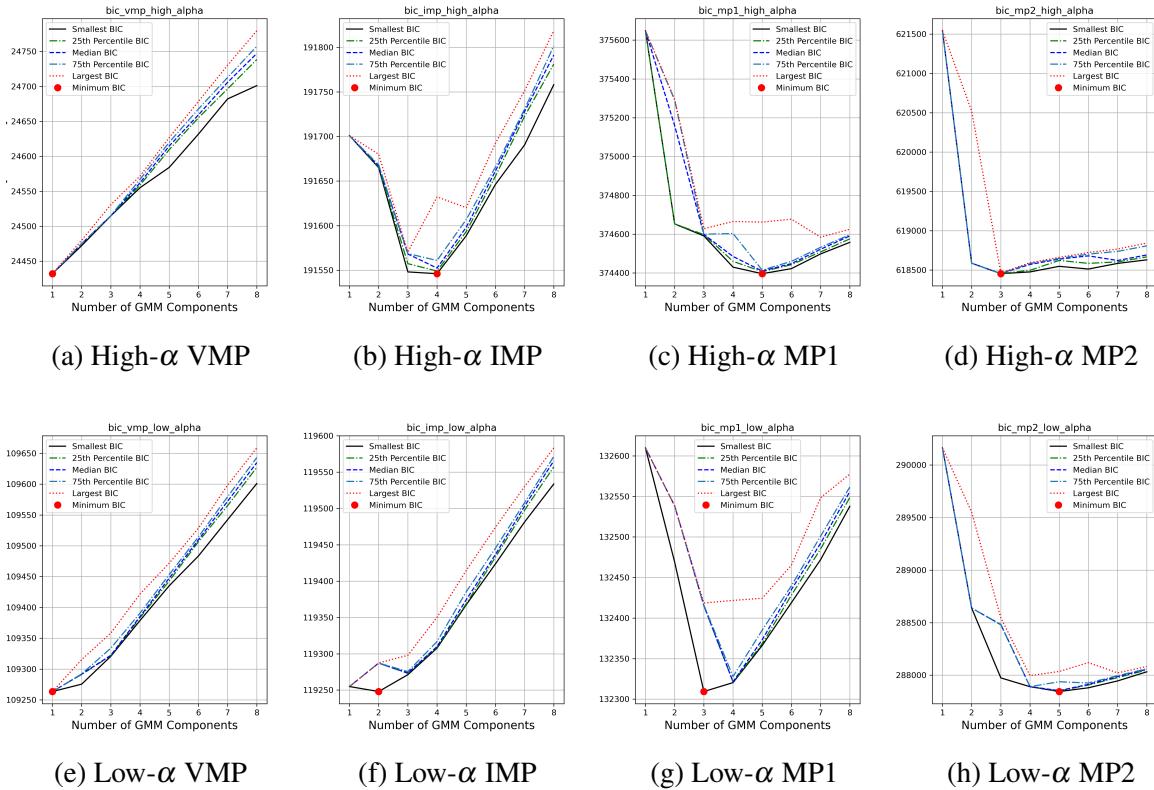


Figure 13: BIC scores for different metallicity bins, grouped by high- and low- α populations.

5.4 High alpha Gaussian Mixture Model Fit

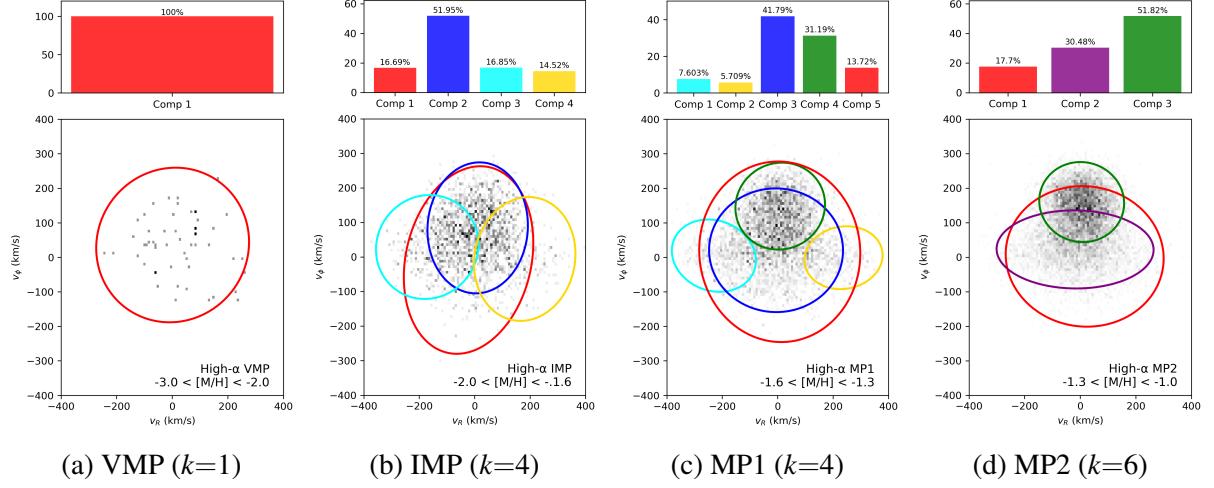


Figure 14: XD–GMM decomposition of high- α giants in four metallicity bins. Grey ellipses mark the 1σ contours of each Gaussian component.

Components	Weights (%)	v_R	σ_R	v_ϕ	σ_ϕ	v_Z	σ_Z
VMP: $-3.0 < [M/H] < -2.0$ (657 stars)							
Stationary halo	100.0	1.40	137.36	35.59	111.99	7.60	102.48
IMP: $-2.0 < [M/H] < -1.6$ (5219 stars)							
Stationary halo	16.7	-22.06	116.74	-8.55	135.74	-2.64	125.40
Prograde halo	51.9	10.84	90.56	84.51	94.79	-0.76	69.42
GS/E (1)	16.8	-171.35	92.33	29.12	75.23	0.50	87.09
GS/E (2)	14.5	180.48	91.48	-5.47	89.66	-2.87	94.37
MP1: $-1.6 < [M/H] < -1.3$ (10265 stars)							
Stationary halo	13.7	7.44	145.24	15.90	130.80	-8.97	127.11
Prograde halo	41.8	-5.85	120.77	20.31	89.59	-1.29	72.85
GS/E (1)	7.6	-230.09	75.81	4.65	52.41	8.69	94.68
GS/E (2)	5.7	237.61	70.43	-1.23	45.51	-3.73	92.81
Thick disc	31.2	9.18	80.79	148.23	62.97	2.99	68.92
MP2: $-1.3 < [M/H] < -1.0$ (17314 stars)							
Stationary halo	17.7	15.91	141.96	2.13	101.75	-1.68	98.90
GS/E	30.5	-18.50	141.08	22.42	56.55	-2.40	77.98
Thick Disk	51.8	5.22	76.27	159.54	57.94	-0.83	67.18

Table 3: GMM component weights and kinematics for the high- α sequence. Velocities are in km s^{-1} . Weights indicate the fraction of stars per metallicity bin.

5.5 Low alpha Gaussian Mixture Model Fit

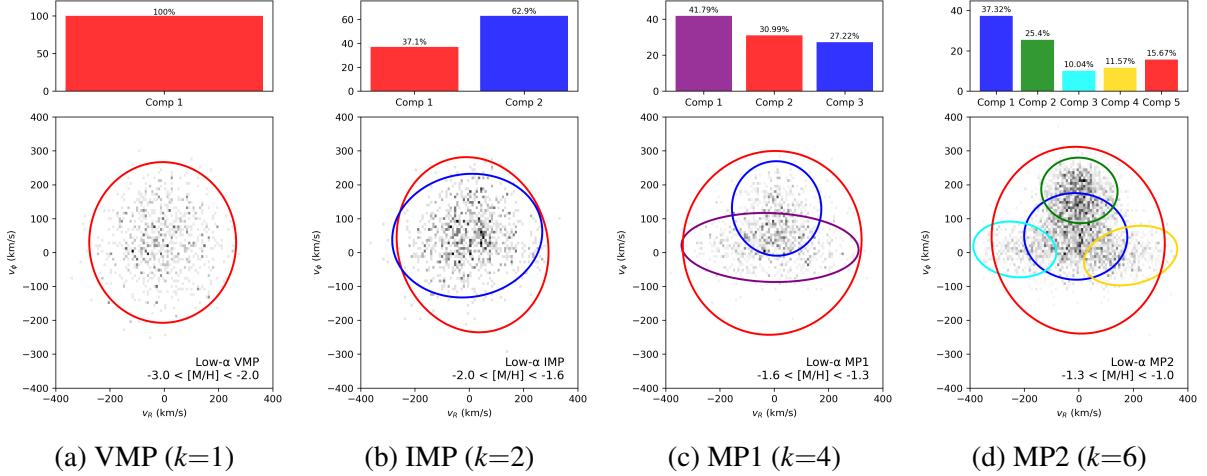


Figure 15: XD–GMM decomposition of low- α giants in four metallicity bins. Grey ellipses show the 1σ contours of every Gaussian component.

Components	Weights (%)	v_R	σ_R	v_ϕ	σ_ϕ	v_Z	σ_Z
VMP: $-3.0 < [M/H] < -2.0$ (2929 stars)							
Stationary halo	100.0	-5.82	135.18	29.55	118.56	1.34	107.10
IMP: $-2.0 < [M/H] < -1.6$ (3228 stars)							
Stationary halo	37.1	12.33	139.85	22.87	129.23	-4.43	120.51
Prograde halo	62.9	-7.34	138.11	49.68	91.25	4.08	75.99
MP1: $-1.6 < [M/H] < -1.3$ (3603 stars)							
Stationary halo	31.0	-7.58	164.57	28.31	135.62	-4.92	121.66
GS/E	41.8	-15.50	162.97	14.71	50.94	-1.93	73.60
Thick Disk/prograde halo	27.2	8.72	82.25	129.87	69.71	3.13	60.68
MP2: $-1.3 < [M/H] < -1.0$ (7945 stars)							
Stationary halo	15.7	-2.30	158.84	36.24	137.81	2.81	131.10
Prograde halo	37.3	-11.31	94.62	47.39	63.89	-1.18	63.12
GS/E (1)	10.0	-234.89	76.09	8.99	41.09	10.70	93.79
GS/E (2)	11.6	189.95	85.58	-9.15	43.99	-8.49	89.69
Thick disc	25.4	2.13	69.98	183.15	48.36	1.75	54.21

Table 4: XD–GMM component weights and kinematics for the low- α sequence. Velocities are in km s^{-1} . Weights are the fractional contribution of each component.

5.6 Contrasting the high and low α GMMS

The two chemical sequences trace very different roads from a dispersion–supported halo to rotationally supported discs.

Very-metal-poor regime (VMP; $[M/H] < -2$). Both data sets are best described by a *single*, non-rotating halo Gaussian. The velocity ellipsoids are close to isotropic ($\sigma_R : \sigma_\phi : \sigma_Z \approx 1 : 1 : 1$), with the high- α stars showing a slightly smaller σ_ϕ , hinting at an early prograde bias. Compared to the results in Section 4.1, where two components were identified in this metallicity

regime, the emergence of only a single Gaussian here may reflect the smaller sample size, limiting the ability of the model to resolve multiple distinct halo structures.

Intermediate regime (IMP; $-2.0 < [\text{M}/\text{H}] < -1.6$). The low- α sample already splits into stationary and prograde halo Gaussians (37 % vs. 63 %), whereas the high- α mixture needs *four* hot components, two of which have the highly radial kinematics characteristic of Gaia–Sausage/Enceladus (GSE). The GSE-like signal in the high- α branch is unexpected, as the GSE merger originated from a dwarf galaxy with a slower star formation history and is therefore expected to contribute primarily α -poor stars. Its appearance here likely reflects mild contamination from $[\alpha/\text{M}]$ systematics at low signal-to-noise rather than a truly α -rich GSE population.

Disc onset (MP1; $-1.6 < [\text{M}/\text{H}] < -1.3$). In both sequences a prograde component with a low dispersion compatible with a thick disk emerges, but with different prominence and kinematics. In the high- α mixture, the thick disc accounts for 31% of the stars, rotating at $v_\phi \sim 148 \text{ km s}^{-1}$, indicating the emergence of a warm, *in-situ* rotating structure. The low- α sample is dominated by non-disc components: the stationary halo and GSE account for 72.8% of the stars, while the remaining 27.2% are captured by a thick-disc/prograde halo Gaussian with $v_\phi \sim 130 \text{ km s}^{-1}$ and higher dispersions. This suggests that α -rich stars settled into a disc structure earlier, while α -poor stars at this metallicity remain primarily associated with accreted debris (GS/E).

Transition to disc dominance (MP2; $-1.3 < [\text{M}/\text{H}] < -1.0$). The high- α fit simplifies to three Gaussians—thick disc (52 %), GSE-like halo (31 %), and a small stationary halo (17 %). The low- α solution remains complex (five Gaussians) but now contains a substantial thick disc (25 %) whose lower dispersion ($\sigma_\phi \sim 48 \text{ km s}^{-1}$) and high v_ϕ are consistent with an early thin-disc precursor. The ultra-cold ($\sigma \lesssim 30 \text{ km s}^{-1}$) thin disc is *absent* because our latitude/extinction cuts remove most mid-plane stars.

6 Discussion

The discussion of the GMM process mirrors that of the discussion in [Zhang et al. \[2024\]](#). Since the primary goal of this report was to reproduce the results of [Zhang et al. \[2024\]](#), the discussion will focus on the differences between the results of this report and those of [Zhang et al. \[2024\]](#). We do however mention the limitations of the GMM approach and selection effects - which are largely the same as those discussed in [Zhang et al. \[2024\]](#).

6.1 Reproducability of [Zhang et al. \[2024\]](#)

The GMM decomposition presented in this report successfully reproduces the same structural features reported by [Zhang et al. \[2024\]](#) by selecting the minimum BIC score for each bin. However, there are several differences in our results worth highlighting.

At the very metal-poor end, our model recovers a nearly even split between stationary and prograde halo components, in contrast to the more dominant non-rotating halo found by [Zhang et al. \[2024\]](#). Although the underlying stellar samples are largely the same, this discrepancy likely reflects the sensitivity of the GMM to initialisation and the precise kinematic distribution of stars in this low-metallicity regime, where substructure is subtle and overlapping. Notably,

the kinematics of the prograde component in our fit are consistent with those reported for Aurora [Belokurov and Kravtsov, 2022], suggesting it may correspond to an early, slightly rotating in-situ population.

In the intermediate metallicity regime ($-2.0 < [\text{M}/\text{H}] < -1.6$), both studies show GSE-like signatures. The agreement in component structure is remarkably consistent, though our fit assigns slightly more weight to the GSE substructures. Interestingly, our GSE components appear with clearer symmetry in radial velocity, which may indicate differences in how the GMM partitions stars with highly eccentric orbits. The recovery of these features in both analyses suggests an ancient radial merger consistent with GS/E [Belokurov et al., 2020].

In the disc onset bin ($-1.6 < [\text{M}/\text{H}] < -1.3$), we find a somewhat higher fraction of thick-disc stars than [Zhang et al. \[2024\]](#), along with slightly lower rotational velocities. This could be due to small variations in how the GMM distinguishes between disc-like and halo-like stars in the transition regime. Despite this, both studies detect a clear emergence of a warm, rotating in-situ component, consistent with the kinematics of a thick disk.

By $[\text{M}/\text{H}] > -1.3$, the thick disc becomes a dominant structure in our analysis and that performed by [Zhang et al. \[2024\]](#), almost surpassing the weight of the prograde halo.

In summary, while small discrepancies exist, particularly in the balance between halo and prograde components at low metallicity, there is an overall agreement in results.

We note that we do not attempt to reproduce the action-space analysis of [Zhang et al. \[2024\]](#) in this report. This decision reflects the fact that our analysis focuses on the use of the XD-GMM algorithm, and because of our understanding that orbital actions do not generally follow Gaussian distributions. Unlike velocity space, which can often be approximated as locally Gaussian for distinct stellar populations, action space retains strong imprints of past dynamical processes (e.g. mergers, phase-mixing, resonances) that result in highly non-Gaussian, often multimodal structures.

6.2 Expansion of [Viswanathan et al. \[2024\]](#)

We extended the XD-GMM analysis to the high and low α sequences identified by [Viswanathan et al. \[2024\]](#). Rather than assuming two Gaussian components as in [Viswanathan et al. \[2024\]](#), we fit a Gaussian mixture model with varying numbers of components in each metallicity bin for both sequences. For each case, the optimal number of components was selected based on the minimum BIC score. This allows us to capture the full complexity of the kinematic structure in both sequences, including the presence of multiple halo components and the emergence of thick disc structures at different metallicities.

We were successfully able to reproduce the figures and data preparation described in [Viswanathan et al. \[2024\]](#), as shown in the plots of Section 5.2.

As outlined in Section 5.1, we expected structural differences between the high- and low- α sequences to emerge in the GMM fits. These expectations were largely met: the high- α stars exhibited a clear transition towards a rotationally supported disc at earlier metallicities, while the low- α stars remained halo-dominated until higher metallicities.

One of the more unexpected outcomes of this analysis is the identification of GS/E components within the high- α sequence, despite the GS/E progenitor being associated with a slow star formation history that should result in an α -poor population. This is likely due to contamination

from low signal-to-noise abundance measurements, or challenges in separating overlapping kinematic substructures when using GMMs. In future work, it would be helpful to consider how abundance uncertainties, particularly at low metallicities, may affect the classification of stars. Additionally, as seen in Figure 9, there appears to be a number of low- α stars at $[M/H] \sim -2.5$ with $[\alpha/M] \sim 0.1$. It remains unclear whether this population is real or the result of measurement uncertainties, and it would benefit from further investigation.

6.3 Selection Effects and

This part of the discussion closely follows the treatment of selection effects in [Zhang et al. \[2024\]](#). As in their study, our application of quality cuts, such as removing stars with high extinction or low Galactic latitude, and enforcing a strict parallax quality threshold, introduces a non-trivial selection function that affects the relative detectability of different Galactic components.

In particular, the Galactic latitude cut removes many stars near the disc midplane, where the low- α thin disc is most prominent. While relaxing this cut does lead to the identification of a colder, low-dispersion disc component in the low- α sequence at higher metallicities, this component is not considered in our analysis. Our focus is on characterising the onset of disc like kinematics across metallicity in both chemical sequences, rather than describing the well established thin disc. Including low latitude stars would bias the sample toward the disc population, introduce uncertainties in the metallicity estimates, since Gaia XP abundances are known to be less reliable in regions of high extinction, and obscure the halo to disc transition that is central to this investigation.

It is important to note that the component weights derived in this work reflect the properties of our selected sample and should not be generalised to the full Galaxy without modelling the underlying selection function.

6.4 Limitations of GMM

Although this part of the discussion mirrors that of [Zhang et al. \[2024\]](#), it is worth reiterating that the Gaussian mixture model approach has several limitations. Gaussian mixtures impose symmetric, ellipsoidal structures in velocity space, which is an oversimplification of the true distribution function of Galactic populations. For instance, the velocity distribution of the thin disc is considered to be non Gaussian due to features such as asymmetric drift, while the GS/E debris has been shown to display bimodal structure in the radial velocity component [[Lancaster et al., 2019](#), [Necib et al., 2019](#)]. In addition, resonant interactions with the Galactic bar may introduce asymmetric and non Gaussian features in the velocity distributions [e.g. [Dillamore et al., 2023](#)].

Following [Zhang et al. \[2024\]](#), we therefore restricted our GMM analysis to the range $-3 < [M/H] < -1$, where the thin disc contribution is minimal. At higher metallicities, the GMM attempts to describe the sharp kinematic features of the thin disc using multiple Gaussians, which are not physically meaningful. While more flexible models based on physically motivated distribution functions could better capture these complexities [e.g. [Binney, 2010](#)], they require significantly more computational effort and are beyond the scope of the aims of both this report and that of [Zhang et al. \[2024\]](#).

Another known limitation of GMMs is their reduced sensitivity to small subpopulations. Like

[Zhang et al. \[2024\]](#), we find little evidence for a significant disc like component at low metallicities, but this absence should be interpreted with caution.

7 Conclusion

In conclusion, we successfully reproduced the XD-GMM analysis of [Zhang et al. \[2024\]](#) which was the primary goal of this research project. This analysis used phase-space information to identify the transition from a dynamically hot halo to a rotationally supported disc in the Milky Way. Analysis found the onset disk-like kinematics at $[M/H] \sim -1.6$ (MP1 bin) - consistent with the findings of [Zhang et al. \[2024\]](#).

We also extended the XD-GMM analysis to the high- and low- α sequences identified by [Viswanathan et al. \[2024\]](#). This extension revealed that the high- α sequence transitions to a more rotationally supported disc at earlier metallicities than the low- α sequence, with the latter remaining dominated by halo-like kinematics until higher metallicities.

Further work should explore the presence of a GS/E-like component in the high- α sequence, which is unexpected given the slow star formation history of the GS/E progenitor. This is likely due to contamination from low signal-to-noise abundance measurements, or challenges in separating overlapping kinematic substructures when using GMMs.

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