



Disentangling the Components of the Milky Way

Inferring the Structure of the Milky Way in Phase-Space Using Gaussian Mixture
Modelling with Extreme Deconvolution

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Abstract

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1 Introduction

The Milky Way Galaxy, host to our solar system, is a spiral galaxy with a centre located approximately 150 000 trillion miles (or 25 000 lightyears) from Earth. Its formation history is complex and remains an active area of research. Being embedded within the Milky Way means we can study it in greater detail than any external galaxy, testing models of galaxy formation with high-precision observational data. One of the central aims of Galactic Archaeology is to reconstruct the Milky Way’s assembly by examining the chemical compositions and dynamical properties of its stars.

In this project, we replicate and extend the analysis of [Zhang et al. \[2024\]](#), who investigated a very metal-poor disc component in the Milky Way. Very metal poor stars, formed from an interstellar medium unpolluted by earlier generations of supernovae, are among the oldest relics in the Galaxy. Discovering them on disc-like orbits would challenge the conventional view that the disc formed later from already enriched gas [[Bland-Hawthorn and Gerhard, 2016](#)], implying instead an earlier onset of disc assembly. Using Gaia DR3, the original study applied a Gaussian Mixture Model with Extreme Deconvolution to the velocity distributions of stars across metallicity bins, probing whether a coherent disc signal persists down to the lowest metallicities.

1.1 Components of the Milky Way

The Milky Way is commonly decomposed into four stellar components: a *thin disc*, a *thick disc*, a central *bulge/bar*, and a roughly spherical *halo* [[Bland-Hawthorn and Gerhard, 2016](#), [Helmi, 2020](#)]. The thin disc dominates the majority of all stars and most of the interstellar gas in the galaxy. Ongoing star formation is concentrated in the “molecular-gas ring” at Galactocentric radii $R \simeq 4\text{--}8$ kpc, where young ($\lesssim 1$ Gyr), metal-rich stars trace nearly circular, co-rotating orbits with low velocity dispersion ($\sigma_\phi \simeq 20 \text{ km s}^{-1}$). Above the mid-plane lies the thick disc: an older ($\gtrsim 8\text{--}10$ Gyr), moderately metal-poor population with $[\text{Fe}/\text{H}] \sim -0.6$ to -1.0 , a scale height of $z_{\text{scale}} \approx 1$ kpc, and hotter kinematics ($\sigma_z \simeq 40 \text{ km s}^{-1}$) while still retaining net prograde rotation. Inside $R \lesssim 2$ kpc, the central bulge—partly bar-shaped—hosts both old, metal-rich stars and a younger, actively forming component; stellar motions there combine bar-driven streaming with high random velocities ($\sigma \sim 100 \text{ km s}^{-1}$). Encasing all of these is the stellar halo, which contributes only a few per cent of the total stellar mass yet harbours the Galaxy’s oldest, most metal-poor stars ($[\text{Fe}/\text{H}] \lesssim -1.5$) on highly eccentric or even retrograde orbits. Its low density, rich substructure, and extreme kinematics reveal an origin in the hierarchical accretion and tidal disruption of dwarf galaxies and globular clusters. Together, the spatial distribution, chemistry, and dynamics of these four components encode the Milky Way’s star-formation history and its sequence of merger events.

1.2 Metallicity as a Cosmic Clock

Precise ages for individual old stars are notoriously difficult to measure, so their chemical composition - most commonly the iron-to-hydrogen ratio, $[\text{Fe}/\text{H}]$ - is often used as a surrogate clock. Very metal-poor (VMP) stars must have formed before successive generations of Type II and Type Ia supernovae had substantially enriched the interstellar medium, and therefore exhibit low $[\text{Fe}/\text{H}]$ values. Metallicity is inferred spectroscopically from the equivalent widths of metal absorption lines such as Fe I and the Ca II K line; after

correcting for effective temperature and surface gravity, their relative strengths give elemental abundances. Large surveys (for example APOGEE, GALAH, LAMOST, and the Gaia XP spectra) now provide such measurements for millions of stars, enabling empirical age–metallicity relations that link chemistry to stellar chronometry [e.g. Nordström et al., 2004, Haywood et al., 2013, Leung and Bovy, 2019, Anders et al., 2023]. These studies consistently show that stars with $[\text{Fe}/\text{H}] \lesssim -1$ are typically older than ~ 10 Gyr, making low-metallicity populations valuable probes of the Milky Way’s earliest disc-building epochs.

1.3 Λ CDM: hierarchical growth and a lopsided halo

In the concordance Λ CDM model, galaxy-sized haloes assemble hierarchically: small dark-matter clumps form first and then merge to build larger structures. Cosmological N -body simulations demonstrate that the number of subhaloes of mass M obeys $dN/dM \propto M^{-1.9}$, a near power-law over many decades in mass [Cooper et al., 2010, Fall and Chandar, 2012]. For a Milky-Way-sized halo this translates to

- $\sim 10^2$ **minor** accretions with $M_{\text{sub}} \lesssim 10^9 M_\odot$, and
- a few **major** events with $M_{\text{sub}} \gtrsim 10^{10} M_\odot$

over a Hubble time.

Only a small subset of these haloes ever form appreciable numbers of stars. Below a critical virial mass $M_{\text{vir}} \sim 10^{11} M_\odot$, re-ionisation and stellar feedback drastically reduce the efficiency of turning gas into stars. Consequently, the stellar-mass–halo-mass (SMHM) relation becomes very steep at the low-mass end [Purcell et al., 2007, Bullock and Boylan-Kolchin, 2017]: most low-mass subhaloes are effectively “dark”, whereas a few relatively massive dwarfs are luminous.

Hence, while the Milky Way has absorbed hundreds of subhaloes, **one or two** of the most massive dwarfs contribute the majority of the halo’s stellar mass; the rest add little more than dark matter and dynamical substructure.

Once accreted, dynamical friction drags the most massive satellites deep into the Galactic potential, their orbits radialise, and their debris is dispersed throughout the inner halo. The disrupted stars inherit coherent signatures—high radial anisotropy, distinctive angular momenta, and chemically narrow sequences—that survive to the present [e.g. Helmi and Tim de Zeeuw, 2000]. Consequently, the stellar halo is not a smooth spheroid but a map of the Galaxy’s merger history, with the inner halo overwhelmingly shaped by a few dominant progenitors (e.g. Gaia–Sausage/Enceladus), and the outer halo supplied by many low-mass accretions.

1.4 Accretion versus *in-situ* disc formation

Chemical and kinematic evidence confirms that the metal-poor halo is primarily accreted. The debris of the Gaia–Sausage/Enceladus (GSE) event, for instance, is traced by stars with $-2 < [\text{Fe}/\text{H}] < -1$ and extreme orbital anisotropy ($\beta \gtrsim 0.8$; Belokurov et al. 2018, Helmi et al. 2018). At $[\text{Fe}/\text{H}] \lesssim -2$ an even broader mix of minor mergers emerges, erasing any global rotation signal [Lancaster et al., 2019, Bird et al., 2021].

Against this backdrop, a number of studies have uncovered stars in the range $-2 < [\text{Fe}/\text{H}] < -1$ whose velocities resemble a disc: modest eccentricities and net prograde rotation [Norris et al., 1985, Chiba and Beers, 2000, Carollo et al., 2019, An and Beers, 2020]. Gaia has pushed this frontier to $[\text{Fe}/\text{H}] < -2$ [Sestito et al., 2019, Venn et al., 2020, Cordini et al., 2020, Mardini et al., 2022]. Whether these objects represent (i) an *in-situ* metal-poor disc or (ii) the spun-up debris of earlier mergers remains hotly debated.

1.5 Origins of very-metal-poor disc candidates

Three broad formation scenarios have been proposed:

1. **Early *in-situ* disc.** Stars form in a gas-rich disc before $z \sim 4$, and later migrate outward or are dynamically heated; such stars would share the chemistry of the proto-Galaxy.
2. **Proto-galactic building blocks.** VMP stars originate in several massive, gas-rich satellites accreted at high redshift; their debris is dragged into the disc plane as the gaseous disc settles [e.g. Sestito et al., 2020].
3. **Late, minor prograde mergers.** Low-mass satellites on aligned orbits are assimilated after the disc forms, depositing a thin layer of metal-poor stars that retain disc-like kinematics [Santistevan et al., 2021].

Cosmological simulations generally reproduce scenario 2, finding that early mergers dominate the VMP budget while a coherent disc does not appear until $z \lesssim 2$ [Gurvich et al., 2023].

Observationally, Belokurov and Kravtsov [2022] identified *Aurora*, a kinematically hot, weakly rotating population with $-2 \lesssim [\text{Fe}/\text{H}] \lesssim -1.3$, arguing against an extremely early disc. Follow-up work shows Aurora to be centrally concentrated [Rix et al., 2022, Arentsen et al., 2020a,b], consistent with heated debris rather than a long-lived thin disc. Furthermore, secular processes such as bar–halo resonances can impart a modest prograde bias to halo stars, mimicking a disc signal [Dillamore et al., 2023].

Unravelling these possibilities demands six-dimensional phase-space information and precision abundances—the focus of the present study.

1.6 This Work

In this study we assess the claim that the Milky Way hosts a very-metal-poor (VMP; $[\text{Fe}/\text{H}] < -2$) stellar disc. Our data set is drawn from *Gaia* DR3, which supplies six-dimensional phase-space coordinates—sky position, parallax-based distance, proper motions, and radial velocity—for each star, together with homogeneous metallicity and α -element abundances from the *Gaia* XP pipeline. To disentangle kinematic sub-populations we model, in successive narrow metallicity bins, the full three-dimensional velocity distribution (v_R, v_ϕ, v_z) with a Gaussian Mixture Model whose parameters are inferred via *Extreme Deconvolution*; the XD formalism explicitly folds the individual distance and proper-motion uncertainties into the likelihood, ensuring that measurement noise does not bias the recovered velocity moments.

Astrophysically, a genuine disc should manifest itself as a high-weight Gaussian centred near the Local Standard of Rest ($v_\phi \simeq 220 \text{ km s}^{-1}$) with small tangential and vertical

dispersions ($\sigma_\phi, \sigma_z \lesssim 30 \text{ km s}^{-1}$) and negligible mean radial motion, whereas the halo or any heated component should appear as a broad, almost isotropic Gaussian with little net rotation and dispersions of order $120\text{--}150 \text{ km s}^{-1}$. By tracking how the weight of the cold, rotating component varies with metallicity we can determine when ordered rotation first emerged and test whether VMP stars were formed in situ or accreted from a satellite.

2 Data

2.1 Sample construction

Our parent catalogue is the bright ($G < 16$) red-giant-branch sample of [Andrae et al. \[2023\]](#). Stellar metallicities are predicted with an eXtreme-Gradient-Boosting model trained on high-resolution APOGEE DR17 spectra and a supplementary set of very metal-poor stars, ensuring reliable performance down to $[\text{M}/\text{H}] \simeq -3.5$. For each of the 17.6 million giants the catalogue delivers homogeneous values of $[\text{M}/\text{H}]$, T_{eff} , and $\log g$ with a quoted random uncertainty of $\simeq 0.1 \text{ dex}$ in $[\text{M}/\text{H}]$ at $G \lesssim 15$. The catalogue only retains entries flagged as “high-confidence” and lying in $-3.5 < [\text{M}/\text{H}] < +0.5$.

Astrometric positions, proper motions, and radial velocities come from the main *Gaia* DR3 tables [[Gaia Collaboration et al., 2023](#)], while heliocentric distances are adopted from the Bayesian photo-geometric catalogue of [Bailer-Jones et al. \[2021\]](#). Because accurate velocities scale with distance precision, we impose a fractional-parallax-uncertainty cut $\sigma_\varpi/\varpi < 0.10$ (FPU); stars that fail this threshold are discarded.

XP spectra are susceptible to reddening: heavy extinction dims the sources, lowers the XP signal-to-noise ratio, and biases the machine-learning metallicities. To minimise such systematics we exclude stars with $E(B-V)_{\text{SFD}} > 0.5$ or Galactic latitude $|b| < 10^\circ$, using colour-excess values from the SFD map accessed via DUSTMAPS [[Green, 2018](#)]. These criteria remove regions where dust corrections are large and spatially variable, at the expense of a modest loss of sky coverage.

Field-star kinematics can also be skewed by dense sub-structures. Accordingly we mask all objects lying within 1° of any known globular cluster or dwarf-galaxy satellite, following the list compiled by [Pace \[2024\]](#). This step eliminates obvious non-field populations (e.g. cluster members and recent accretion debris) without significantly reducing the statistical power of the sample.

After the metallicity flag, distance-quality, reddening, latitude, and sub-structure cuts, our working data set comprises $\sim 3.4 \times 10^6$ red-giant stars possessing homogeneous metallicities and full six-dimensional phase-space information. This curated sample underpins the chemo-kinematic analysis presented in the remainder of this report.

2.2 Galactocentric Positions and Velocities

Six-dimensional phase-space coordinates are obtained with `astropy.coordinates`. We adopt a Galactocentric frame with $R_0 = 8.1 \text{ kpc}$ and $Z_0 = 25 \text{ pc}$ [[McMillan, 2016](#)], and a solar velocity¹ $(U_\odot, V_\odot, W_\odot) = (11.1, 245, 7.25) \text{ km s}^{-1}$ [[Schönrich et al., 2010](#)]. The

¹Cartesian components (U, V, W) , where U is radially outwards, V is aligned with Galactic rotation, and W points to the North Galactic Pole.

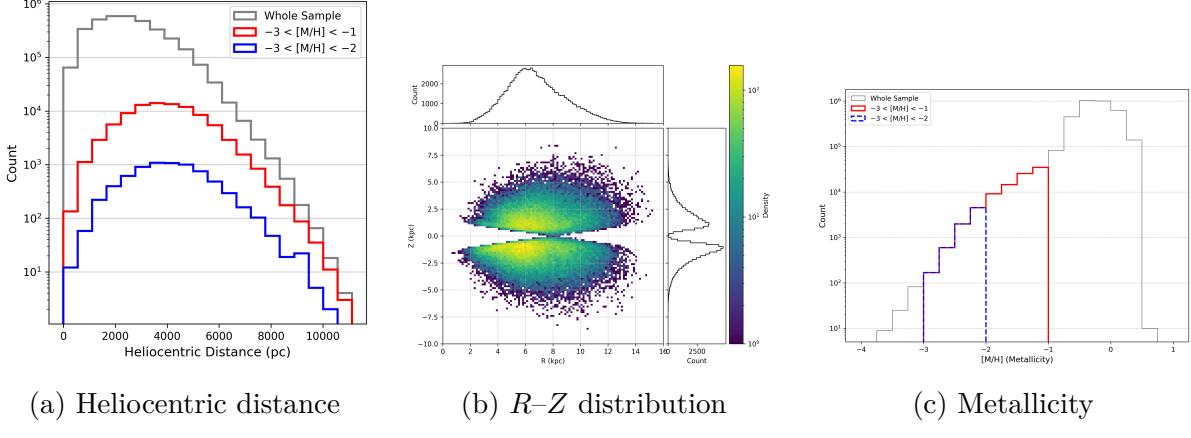


Figure 1: Properties of the final RGB sample after all quality and footprint cuts. *Left:* heliocentric-distance histogram for the whole sample (grey); the subsets with $-3 < [\text{M}/\text{H}] < -1$ and $-3 < [\text{M}/\text{H}] < -2$ are shown in solid red and dashed blue, respectively. *Middle:* density map in Galactocentric cylindrical coordinates. The empty band at low $|Z|$ is a selection artefact of our latitude/extinction cuts, which deliberately remove the thin-disc mid-plane; the concentration around $R \simeq 5\text{--}8\text{ kpc}$ reflects the volume accessible to bright RGB stars interior to the Solar circle and coincides with the molecular ring region where the stellar surface density peaks. *Right:* Metallicity distribution of our data sample. Line colours are the same as in the left panel. (A reproduction of Figure 1 from [Zhang et al. \[2024\]](#))

cylindrical velocity components (v_R, v_ϕ, v_Z) are extracted from the transformed SkyCoord module.

To propagate measurement errors we generate $N_{\text{MC}} = 100$ Monte–Carlo realisations per star, drawing parallax, proper motions, radial velocity, and distance from their reported uncertainties (the proper–motion covariance is honoured through a bivariate normal). Each realisation is transformed to the Galactocentric frame, yielding distributions of v_R , v_ϕ , and v_Z ; the 1σ widths of those distributions are stored as per-star velocity uncertainties.

As shown in Figure 2, stellar azimuthal velocities evolve increase with metallicity. Halo-like kinematics dominate at $[\text{M}/\text{H}] \lesssim -1.5$ with statistics indicative of a pressure-supported component. A rapid spin-up appears over the metallicity range $[\text{M}/\text{H}] \simeq -1.3$ to -0.9 , consistent with [Bekokurov and Kravtsov \[2022\]](#). By $[\text{M}/\text{H}] \gtrsim -0.5$ the stellar azimuthal velocities reaches the Local Standard-of-Rest value ($\approx 220 \text{ km s}^{-1}$) and the velocity dispersion dispersion falls, marking the transition to a disc. Visually, using 2, we can observe that the onset of ordered rotation in the Milky Way occurred when the inter-stellar medium reached roughly one-tenth solar metallicity.

As shown in Fig. 3, stars with $[\text{M}/\text{H}] \geq -1.0$ have a strong prograde bias in azimuthal velocities, with a peak at $v_\phi \gtrsim 180 \text{ km s}^{-1}$ and a relatively narrow distribution in v_R . As shown in Fig. 4, this is consistent with the thin– and thick-disc ellipses. Below $[\text{M}/\text{H}] \simeq -1.0$ the distribution broadens and drops toward $v_\phi \approx 0$, indicating pressure supported kinematics, characteristic of the stellar halo and the radial Gaia–Sausage/Enceladus debris. At the lowest metallicities ($[\text{M}/\text{H}] \lesssim -1.5$) the contours are nearly isotropic with only a mild prograde bias. Hence, any rotation-supported very-metal-poor disc must

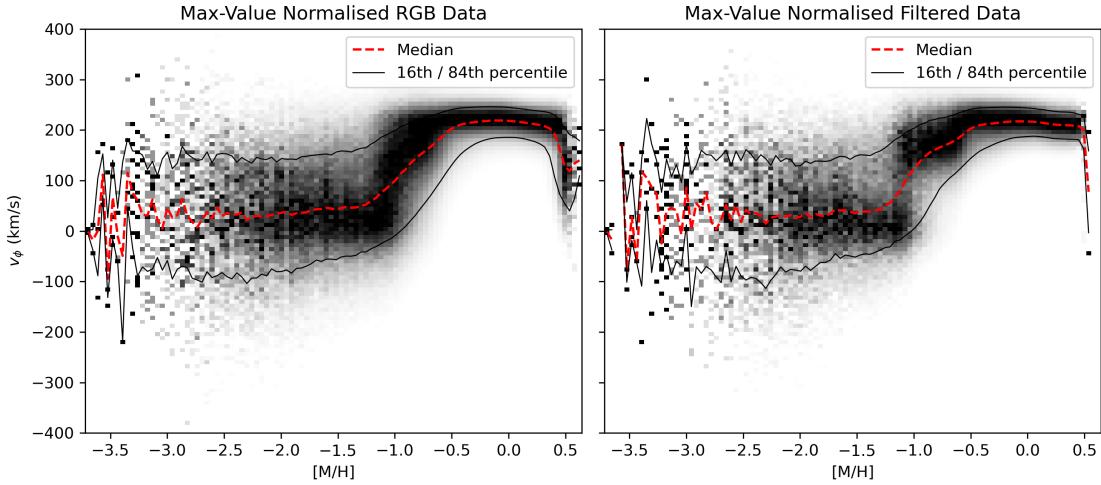


Figure 2: Column-normalised density in the v_ϕ -[M/H] plane. *Left:* the bright-RGB catalogue of [Andrae et al. \[2023\]](#). *Right:* the same sample after all distance, dust, and quality cuts. Greyscale pixels show the normalised counts in each metallicity bin; the red dashed curve is the median v_ϕ , and the black curves trace the 16th and 84th percentiles. (A reproduction of Figure 2 from [Zhang et al. \[2024\]](#))

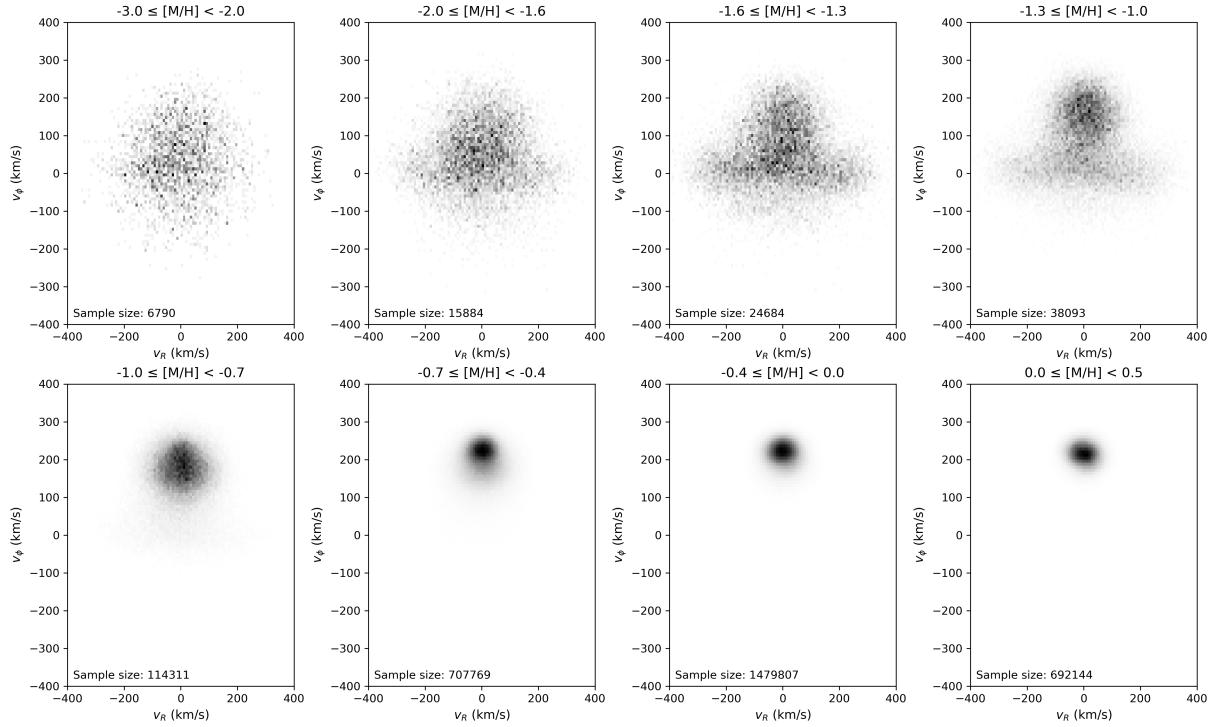


Figure 3: Galactocentric velocity distributions as a function of metallicity. Each panel shows the column-normalised density of stars in the v_R - v_ϕ plane for the metallicity interval printed at the top. With increasing metallicity the distribution contracts in both directions—signalling lower velocity dispersion—while the bulk of stars moves upward to larger prograde azimuthal velocity. (A reproduction of Figure 3 from [Zhang et al. \[2024\]](#))

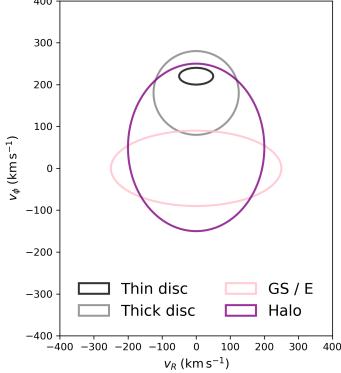


Figure 4: Ellipses mark the approximate extent of the thin disc (black), thick disc (grey), Gaia–Sausage/Enceladus debris (pink) and the pressure-supported stellar halo (purple). The figure serves as a visual key for interpreting the data panels in Fig. 3.

contribute at most a small fraction of the population. In subsequent analysis we quantitatively assess these observations.

Given we are testing whether a very-metal-poor stellar disc exists, we naturally restrict the sample to stars within $|Z| < 2.5$ kpc of the galactic mid-plane. This keeps the focus on stars close to the plane, where any disc-like population (whether formed in situ or deposited by mergers) would be found [Tkachenko et al., 2025].

Upon comparison with Zhang et al. [2024], we note that our sample is approximately the same size - with a difference of less than 0.002% of the sample size in each bin.

3 Methodology

3.1 Gaussian Mixture Model

To quantitatively investigate the kinematic structure of metal-poor stars and assess the presence of a potential very-metal-poor disc, we use a Gaussian Mixture Model (GMM) framework as done by Zhang et al. [2024]. GMMs are a class of unsupervised machine learning algorithms commonly used in data science for clustering and density estimation. They model a dataset as a weighted sum of multivariate Gaussian distributions, each corresponding to a latent sub-population within the data.

From a probabilistic perspective, the GMM assumes that each observed data point is generated from a hidden (latent) variable indicating membership in one of the Gaussian components. This latent space formulation allows the model to assign probabilistic classifications to data points, providing a soft clustering where each star has a fractional likelihood of belonging to each component. In our case, the data are three-dimensional velocity vectors (v_R, v_ϕ, v_Z) , and the latent space captures distinct kinematic substructures in this space.

We implement the GMM fitting using the pyGMMis package [Melchior and Goulding, 2016] in the v_R , v_ϕ , and v_Z phase space. This algorithm extends the standard Expectation-Maximisation (EM) algorithm with the “Extreme Deconvolution” technique developed by Bovy et al. [2011]. This method is particularly well-suited to astronomical datasets,

as it incorporates measurement uncertainties into the GMM fitting by modifying the EM updates to account for known errors on each data point. In our case, these uncertainties are derived from the Gaia astrometric and spectroscopic data and are represented by diagonal covariance matrices encoding the squared uncertainties in v_R , v_ϕ , and v_Z for each star.

3.2 Number of Gaussian Components

We apply the GMM separately to each metallicity bin with $[M/H] < -1$, as we are specifically interested in detecting any rotationally supported structure among the very metal-poor stars. An important decision must be made when implementing the appropriate number of components in applying Gaussian Mixture Models. Too few components may underfit the data, missing real substructures, while too many will overfit, leading to spurious and physically uninterpretable results.

We select the optimal number of components using the Bayesian Information Criterion (BIC) [Schwarz 1978](#). The BIC is defined as

$$BIC = k \ln(n) - 2 \ln \mathcal{L}, \quad (1)$$

where $k = (1 + 3 + 6) \times N - 1$ is the total number of free parameters in a model with N Gaussian components (accounting for the weights, means, and covariances), n is the number of stars in the sample, and \mathcal{L} is the maximum likelihood of the fit. The BIC penalizes model complexity, such that adding extra components without a significant gain in likelihood will result in a higher BIC value.

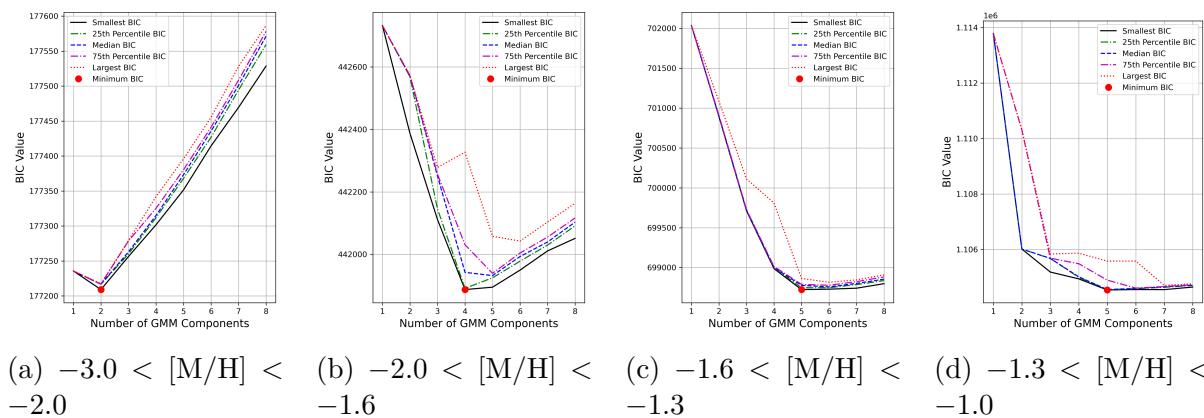


Figure 5: BIC distributions as a function of the number of GMM components in each metallicity bin. The optimum number of components is indicated by the minimum BIC value (highlighted in red). (A reproduction of Figure 4 from [Zhang et al. \[2024\]](#))

The Expectation-Maximisation algorithm can become trapped in local minima, so we performed 50 initialisations for each N and recorded the BIC value for each trial. The minimum BIC value across all trials indicated the statistically preferred model and suggests that the global optimum was likely reached.

To improve convergence and model stability, the GMM components in each trial were initialised using the kmeans algorithm. This is due to the sensitivity of GMMs to

their starting conditions: poor initialisation can lead to convergence on undesirable solutions, particularly in high-dimensional or overlapping data. KMeans clustering provides an initialisation point by partitioning the velocity space into compact, roughly spherical clusters. This works well with the assumptions of Gaussian components and often gives faster, more stable convergence and more physically meaningful results. In our application, where stellar substructures are partially overlapping in velocity space, this method applies well.

In Figure 5, we show the distribution of BIC values across four metallicity bins as a function of N . The resulting preferred number of components are 2, 4, 5, and 5 for the VMP, IMP, MP1, and MP2 bins. This is in agreement with the findings of [Zhang et al. \[2024\]](#). Our analysis, however, selects the lowest BIC each time, removing the need for a subjective choice of the number of components when BIC values are very similar. [Zhang et al. \[2024\]](#), when observing the BIC values for the MP1 bin chooses to use 5 components instead of 6 despite the BIC score for the 6-component model being lower than the 5-component model.

4 Results

4.1 Gaussian Mixture Model Fit

Using the number of components selected by the BIC criteria, we fit the GMM to the data in each metallicity bin as shown in Figure 6. We increased the number of initialisations to 100 for the final fitting to ensure convergence to a stable solution.

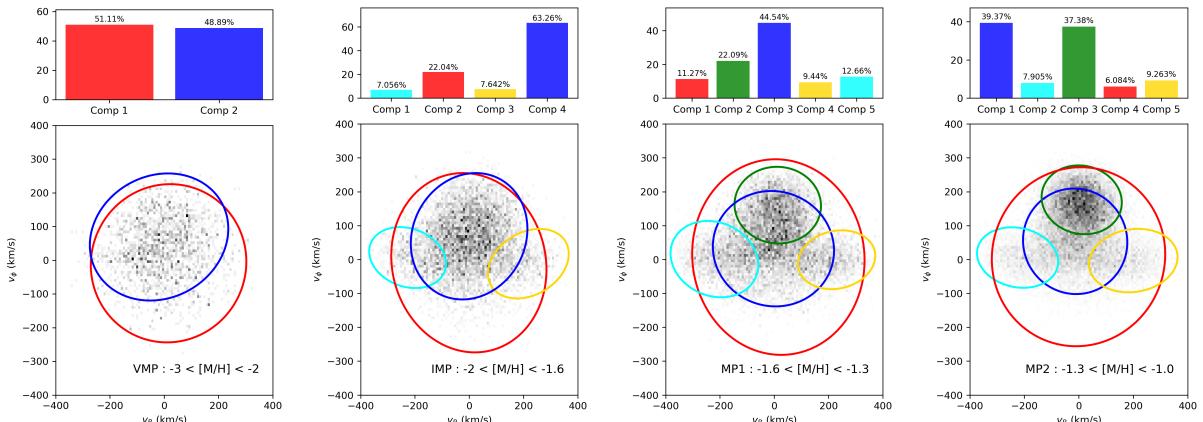


Figure 6: Gaussian Mixture Model decompositions of the stellar velocity distribution in the v_R - v_ϕ plane for each metallicity bin. The bottom panel of each subfigure shows the 2D velocity distribution with GMM components overplotted as ellipses representing the 1σ contours in v_R - v_ϕ . The top panel shows the fractional contribution. (A reproduction of Figure 5 from [Zhang et al. \[2024\]](#))

As shown in Figure 6, the GMM is able to capture kinematic substructures in the velocity space of metal-poor stars. As metallicity increases, a gaussian component with a clear prograde rotation signal emerges (in green), indicative of a disc-like structure. The statistics of the GMM fit are summarised in Table 1, which lists the weights, means, and dispersions of each component in the four metallicity bins.

Components	Weights (%)	v_R	σ_R	v_ϕ	σ_ϕ	v_Z	σ_Z
VMP: $-3.0 < [\text{M}/\text{H}] < -2.0$ (4768 stars)							
Stationary halo	51.1	15.86	143.17	-9.13	117.36	-0.08	122.64
Prograde halo	48.9	-19.14	127.43	68.86	94.28	-0.90	82.89
IMP: $-2.0 < [\text{M}/\text{H}] < -1.6$ (12052 stars)							
Stationary halo	22.0	-1.45	142.88	-9.59	132.32	-0.91	124.92
Prograde halo	63.3	-0.58	107.32	68.52	93.27	-1.16	72.56
GS/E(1)	7.1	-226.92	70.78	5.60	44.98	8.43	88.40
GS/E(2)	7.6	217.52	74.41	-12.35	51.35	-4.22	89.78
MP1: $-1.6 < [\text{M}/\text{H}] < -1.3$ (19142 stars)							
Stationary halo	11.3	15.09	158.68	7.11	144.33	-2.55	132.08
Prograde halo	44.5	-3.02	111.54	31.79	85.15	-0.54	70.36
GS/E(1)	12.7	-220.40	80.76	1.01	56.29	0.47	91.00
GS/E(2)	9.4	229.32	71.22	-0.82	43.42	1.42	92.69
Thick disc	22.1	12.93	79.56	160.32	56.58	-2.31	68.83
MP2: $-1.3 < [\text{M}/\text{H}] < -1.0$ (30892 stars)							
Stationary halo	6.1	-1.17	159.06	8.07	132.05	-3.37	120.46
Prograde halo	39.4	-13.47	95.65	53.64	77.86	-3.04	71.00
GS/E(1)	7.9	-224.68	74.44	5.73	44.63	2.41	87.27
GS/E(2)	9.3	199.71	81.28	-3.30	46.84	-2.07	88.99
Thick disc	37.4	10.56	73.55	176.07	50.79	0.85	62.03

Table 1: Parameters of the Gaussian mixture model fittings in different metallicity bins. The unit for all velocity columns is km s^{-1} .

At $[\text{M}/\text{H}] < -1.6$ (the VMP and IMP regimes), a disc component is not observed. Instead, the kinematic structure is dominated by a prograde halo and a stationary halo. In the VMP bin, the sample is almost evenly split between these two components, with the stationary halo contributing 51.1% and the prograde halo 48.9%. Both show broad velocity dispersions, with only modest rotational support ($\overline{v_\phi} \sim 69 \text{ km s}^{-1}$ for the prograde halo) and no evidence for a kinematically cold disc-like structure.

In the IMP bin, the picture remains qualitatively similar, although the prograde halo dominates more strongly (63.3%) and additional substructure becomes apparent. Two minor components, interpreted as fragments of the Gaia–Sausage/Enceladus (GS/E) merger debris, are also identified, each contributing $\sim 7\%$ of the population and exhibiting highly radial orbits ($|v_R| > 200 \text{ km s}^{-1}$). These GS/E components are characterised by strong anisotropy and contribute to the broadening of the halo distribution, but again show no signature of disc-like kinematics. The absence of a cold, rotating disc in these bins suggests that if a very-metal-poor disc exists, it must be a minor contributor to the local stellar population.

Our GMM decomposition recovers the same substructures identified by [Zhang et al. \[2024\]](#) across all metallicity bins. We find similar velocity dispersions and centroid trends for the stationary halo, prograde halo, and GS/E components. Notably, our model assigns a more comparable weight to the prograde halo component in the VMP bin, in contrast to the dominant stationary halo reported by [Zhang et al. \[2024\]](#). We also observe some variation in the centroid v_ϕ values of GS/E substructures, although their bimodal spatial structure is preserved.

4.2 Rotational Support and the Onset of Disc Formation

To assess the degree of rotational support in each component, we compute the ratio $V_{\text{rot}}/\sigma_\phi$, where V_{rot} is the mean azimuthal velocity and σ_ϕ its dispersion. This ratio serves as a conventional diagnostic of disc-like kinematics, with values $\gtrsim 3$ indicating rotation-supported populations. As shown in Figure 7, all components in the VMP and IMP bins fall below this threshold, confirming that the GMM identifies only dynamically hot, dispersion-supported structures in these regimes. This is consistent with the interpretation of the dominant components as stationary and prograde halo populations, with $V_{\text{rot}}/\sigma_\phi \approx 0.7$ in both cases (see also Table 1).

At higher metallicities, a dynamically colder disc population emerges. In the MP1 bin, the thick disc component (green) rotates at $v_\phi \sim 160 \text{ km s}^{-1}$ with $\sigma_\phi \sim 57 \text{ km s}^{-1}$, yielding $V_{\text{rot}}/\sigma_\phi \approx 2.8$. While just below the canonical boundary, this component's low vertical dispersion and substantial weight (22.1%) mark it as a distinct disc-like structure. By the MP2 bin, the disc becomes the dominant component (37.4%), with $V_{\text{rot}}/\sigma_\phi \approx 3.5$, indicating clear rotational support and a well-established thick disc. These trends point to a sharp transition in the stellar kinematics around $[\text{M}/\text{H}] \sim -1.5$, consistent with the early onset of disc formation.

Meanwhile, the GS/E components — identifiable in the IMP, MP1, and MP2 bins — remain dynamically distinct, with high $|v_R|$, relatively small v_ϕ , and negligible rotational support, reinforcing their origin as debris from a major radial merger [Helmi et al., 2018]. Notably, given the estimated time of the GS/E accretion event (of order 8–11 Gyr ago [Gallart et al., 2019, Belokurov et al., 2020, Di Matteo et al., 2019]), its stellar debris is expected to be phase-mixed. This implies that the positive and negative v_R components should contribute approximately equally. This symmetry is indeed observed in the two MP bins. However, in the IMP bin the distribution is clearly asymmetric, with the negative v_R GS/E component (aqua) containing nearly twice as many stars as the positive one (orange). This may suggest incomplete phase mixing or contamination by other structures at this intermediate metallicity.

4.3 Residual Analysis for Hidden Disc Populations

One of the limitations of GMMs is that they can fail to detect substructures that contribute only weakly to the overall distribution. To address this, we perform a residual analysis to test whether a disc-like population—too weak to be picked up by the GMM—might nonetheless be present in the VMP and IMP metallicity regimes.

We generated a synthetic dataset by drawing the same number of mock stars as observed stars from the best-fit GMM in each bin. Since our GMM uses the Extreme Deconvolution algorithm, which accounts for observational uncertainties during fitting, the raw samples from the model are noise-free. However, comparing these directly to real data would be inappropriate due to the absence of measurement error. To resolve this, we assign each mock star the velocity uncertainties of its nearest neighbour in the observed data (in (v_R, v_ϕ, v_Z) space), and then add Gaussian noise according to these uncertainties.

We then bin both the observed and mock data in the (v_R, v_ϕ) plane and compute the normalized residual map, defined as

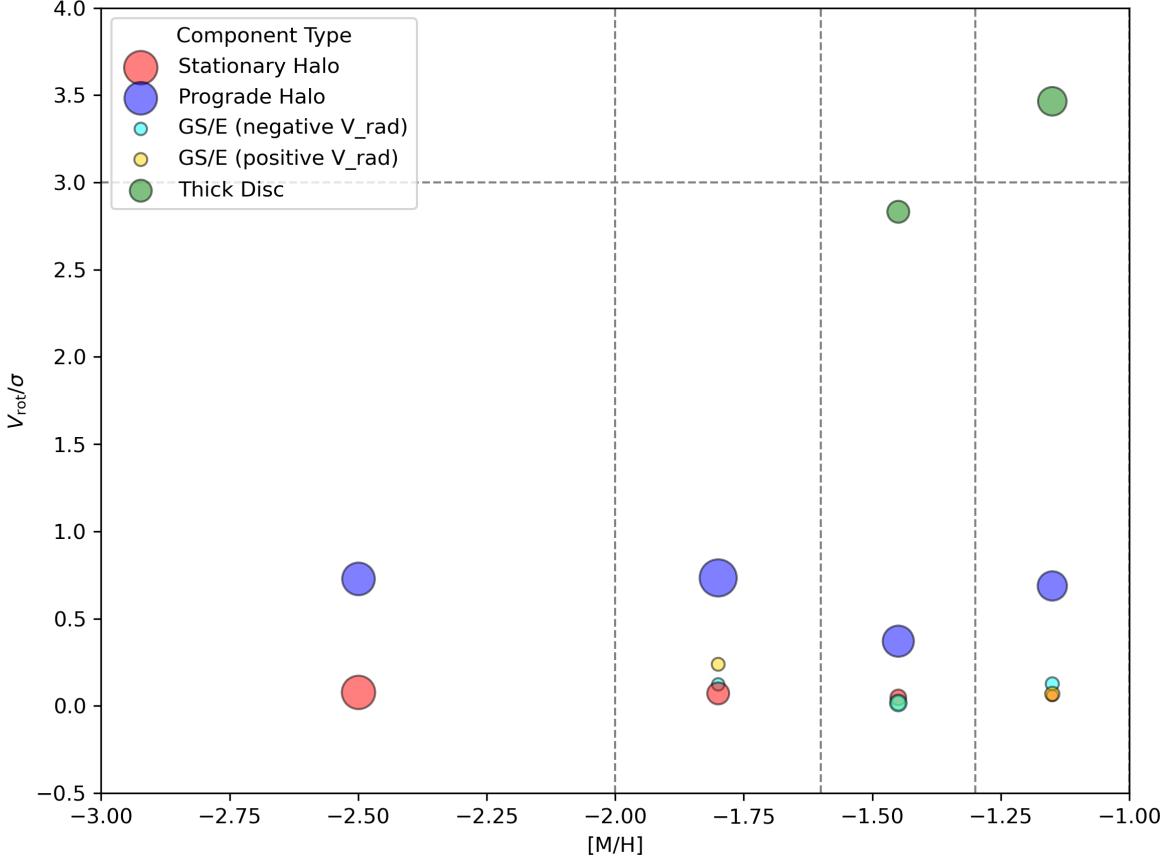


Figure 7: Ratio of mean rotational velocity to azimuthal velocity dispersion, $V_{\text{rot}}/\sigma_\phi$, for individual GMM components in each metallicity bin. Larger squares correspond to more dominant components. Colours match the GMM components in Figure 6.(A reproduction of Figure 6 from [Zhang et al. \[2024\]](#))

$$H_{\text{residual}} = \frac{H_{\text{obs}} - H_{\text{mock}}}{H_{\text{obs}} + H_{\text{mock}} + \epsilon},$$

where H_{obs} and H_{mock} are the 2D histograms and $\epsilon = 10^{-5}$ is added to avoid division by zero. The residual maps are shown in the top row of Figure 8 for the VMP and IMP bins. The grey dashed ellipses highlight a 2σ region of a hypothetical thick disc with mean velocity $(v_R, v_\phi) = (0, 180)$ km s $^{-1}$ and dispersions $(\sigma_R, \sigma_\phi) = (70, 50)$ km s $^{-1}$, consistent with the thick disc component seen in the MP2 metallicity bin.

To quantify the possible presence of a hidden disc, we count the excess number of observed stars inside this thick disc ellipse relative to the GMM-generated sample. We repeat this calculation for 200 Monte Carlo realizations to estimate the residual and its uncertainty. The results are shown as horizontal dashed lines in the bottom row of Figure 8, representing a disc-like residual of 40.4 ± 24.8 stars in the VMP bin and 76.9 ± 42.8 stars in the IMP bin.

To interpret these residuals, we repeat the same procedure but now inject a mock disc population—generated from the same thick disc Gaussian—into the GMM baseline. By varying the injected disc fraction from 0% to 5% of the sample and recalculating

the residual each time, we construct a relationship between disc fraction and expected residual. This is shown as the solid black line in the bottom panels of Figure 8, with the red shaded region indicating the 1σ uncertainty across Monte Carlo trials.

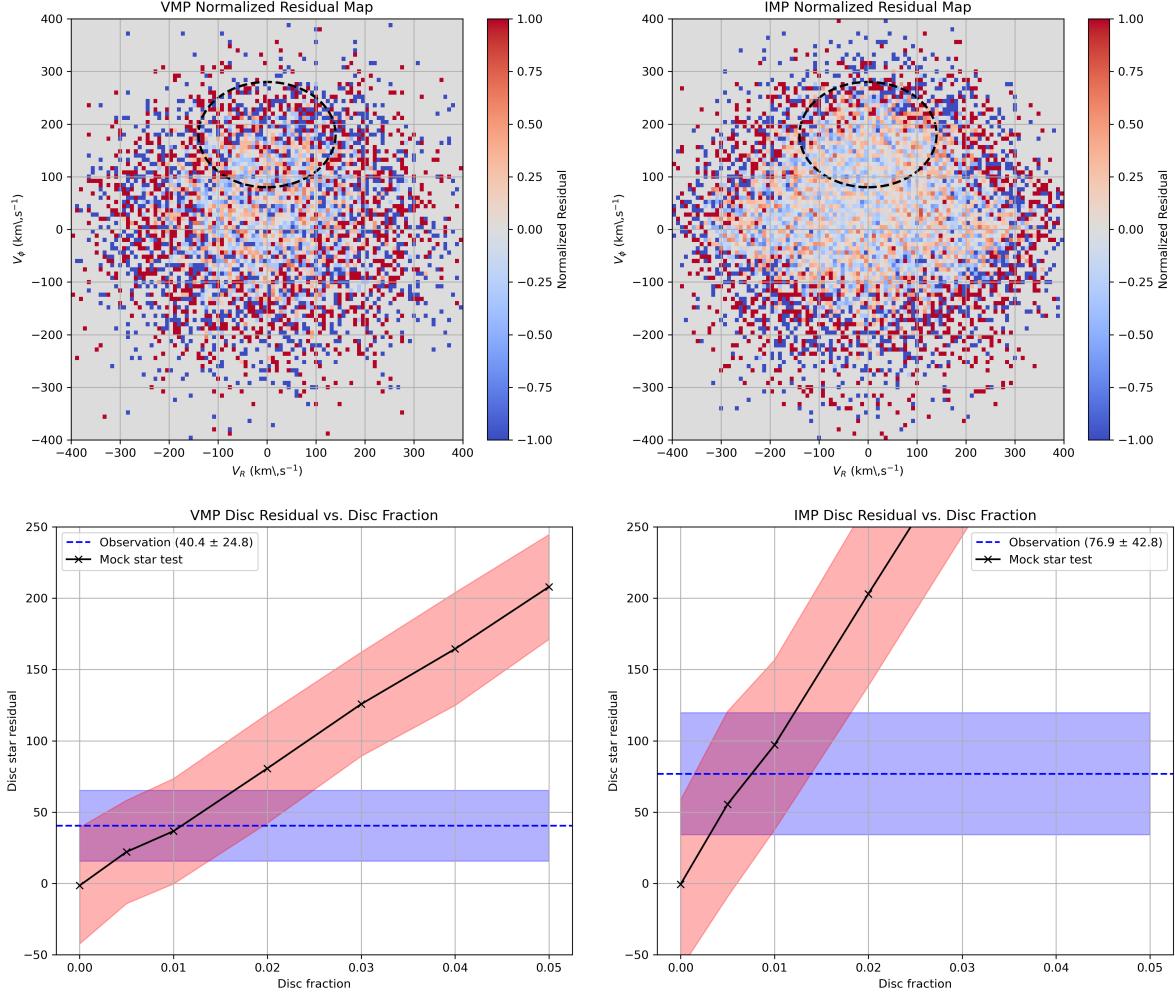


Figure 8: Top: Normalized residual maps between observed and GMM-predicted velocity distributions in (v_R, v_ϕ) for the VMP (left) and IMP (right) metallicity bins. Grey ellipses show the 2σ contour of the fiducial thick disc model. Bottom: Disc residuals as a function of injected disc fraction. The dashed line and blue band indicate the observed residual and its uncertainty. The solid black line and red region show mock test results and their uncertainties. (A reproduction of Figure 7 from [Zhang et al. \[2024\]](#))

5 Extension analysis

5.1 Alpha-element abundance

We extend the chemo-kinematic analysis by adding the alpha-to-iron ratio $[\alpha/\text{Fe}]$, where the “alpha” elements (O, Ne, Mg, Si, S, Ca, Ti) are produced almost entirely in core-collapse (Type II) supernovae, whereas most iron is released later by thermonuclear (Type Ia) explosions. Because the two channels operate on very different time-scales— $\sim 10\text{--}30$ Myr for Type II versus $\gtrsim 0.1\text{--}1$ Gyr for Type Ia—the quantity $[\alpha/\text{Fe}]$ encodes the star-formation history of a stellar population:

- **High** $[\alpha/\text{Fe}]$. Rapid, intense star formation enriches gas with α -elements before significant iron production from Type Ia supernovae occurs, leading to elevated $[\alpha/\text{Fe}]$. This chemical signature is typical of the in-situ thick disc, stellar halo, and some ancient accreted material. However, at fixed metallicity, accreted populations, especially dwarf galaxies, tend to show lower $[\alpha/\text{Fe}]$ compared to in-situ stars, reflecting their slower chemical evolution.
- **Low** $[\alpha/\text{Fe}]$. A more prolonged formation history allowed Type Ia ejecta to accumulate, diluting the alpha enhancement. The present-day thin disc and metal-rich bulge occupy this regime.

Hence the $[\alpha/\text{Fe}]-[\text{M}/\text{H}]$ plane acts as a chemical clock: populations that overlap in $[\text{M}/\text{H}]$ but differ in formation time-scale separate cleanly in $[\alpha/\text{Fe}]$.

Following Viswanathan et al. [2024], we separate the RGB catalogue into high- and low- $[\alpha/\text{M}]$ sequences using the XP abundances of Li et al. [2023] and model each subset with an error-convolved Gaussian-mixture as described earlier. Before fitting the GMM, we can expect to see some differences between the high- and low- $[\alpha/\text{M}]$ sequences as summarised below.

In the high- $[\alpha/\text{M}]$ sequence, the oldest, in-situ population formed rapidly from well-mixed, α -rich gas. Below $[\text{M}/\text{H}] \approx -1.6$ we therefore expect a single non-rotating halo Gaussian. As enrichment proceeds, gas spins up and produces the thick disc: a prograde, kinematically warm Gaussian whose weight rises steeply between $[\text{M}/\text{H}] \simeq -1.6$ and -1.0 and dominates for $[\text{M}/\text{H}] > -0.7$. Accreted debris from the last major merger (Gaia–Sausage/Enceladus, GS/E) should not be prominent here because GS/E stars are typically *less* α -enhanced than the in-situ thick disc at the same metallicity [e.g. Helmi et al., 2018]; our conservative high- $[\alpha/\text{M}]$ cut should remove most of this material.

In the low- $[\alpha/\text{M}]$ sequence, lower $[\alpha/\text{M}]$ signals slower enrichment, typical of dwarf satellites or the Milky Way disc after Type-Ia supernovae began adding iron. A single halo Gaussian should suffice below $[\text{M}/\text{H}] \approx -1.3$, with a cold, high-weight thin-disc Gaussian—centred near the Local Standard of Rest and $\sigma \lesssim 30 \text{ km s}^{-1}$ —emerging above that threshold. Our footprint excludes $|b| < 10^\circ$, however, and thus removes most of the geometrically thin disc; the ultra-cold Gaussian may therefore be absent at the low metallicities we are inspecting. The aim is instead to pinpoint the onset of disc-like rotation in the low- $[\alpha/\text{M}]$ sequence despite this selection bias.

5.2 Data

We obtain the α -to-metal abundance ratios from the XP-based catalogue of Li et al. [2023], whose neural-network model, trained on high-resolution APOGEE DR17 spectra, achieves a typical precision of $\simeq 0.05$ dex for bright sources ($G < 16$). Following Viswanathan et al. [2024], we cross-match the XP-based α catalogue of Li et al. [2023] with the metallicity-selected red-giant branch sample. This base sample already incorporates the red-giant selection and photometric quality cuts from Andrae et al. [2023], including limits on surface gravity, colour-magnitude position, and XP calibration reliability.

To further ensure robust α abundances and kinematics, we apply several additional cuts:

- we reject stars within 1° of known dwarf galaxies and those associated with globular clusters;
- a reddening cut $E(B - V) < 0.5$ removes stars in regions of high extinction, where XP abundances are less reliable;
- a Galactic latitude cut $|b| > 10^\circ$ similarly avoids the dust-dominated plane;
- we relax the fractional parallax uncertainty cut relative to Section 2.1, requiring only $f_{\text{pu}} < 0.2$ to retain sufficient stars in the metal-poor regime, which is the focus of this analysis.

We obtain approximately 3.4 million stars in our final sample with alpha element abundances.

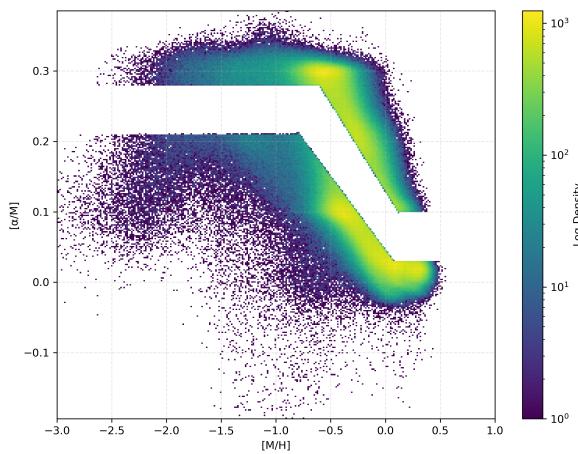


Figure 9: Density of stars in the $[\alpha/\text{M}]$ – $[\text{M}/\text{H}]$ plane. The high- α and low- α sequences are separated by the piece-wise linear boundary . The transition zone between the two sequences is clearly removed from this plot and discarded from subsequent analysis. (A reproduction of Figure 1 in [Viswanathan et al. \[2024\]](#))

Figure 9 displays the logarithmic density of our RGB sample in the $[\alpha/\text{M}]$ – $[\text{M}/\text{H}]$ plane. Following [Chandra et al. \[2024\]](#), but with a deliberately tighter boundary as in [Viswanathan et al. \[2024\]](#):

$$\begin{aligned} \text{high-}\alpha : & \begin{cases} [\text{M}/\text{H}] < -0.60 & [\alpha/\text{M}] > 0.28 \\ -0.60 \leq [\text{M}/\text{H}] \leq 0.125 & [\alpha/\text{M}] > -0.25[\text{M}/\text{H}] + 0.13 \\ [\text{M}/\text{H}] > 0.125 & [\alpha/\text{M}] > 0.10 \end{cases} \\ \text{low-}\alpha : & \begin{cases} [\text{M}/\text{H}] < -0.80 & [\alpha/\text{M}] < 0.21 \\ -0.80 \leq [\text{M}/\text{H}] \leq 0.07 & [\alpha/\text{M}] < -0.21[\text{M}/\text{H}] + 0.04 \end{cases} \end{aligned}$$

and discard the intermediate “transition” band.

This high- α threshold is raised so that the chemistry of the Gaia–Sausage/Enceladus debris [[Belokurov et al., 2018](#), [Helmi et al., 2018](#)] does not leak into the in-situ sequence; and the exclusion band is widened to maximise purity in both sub-samples [[Viswanathan et al., 2024](#)].

Figure 10 shows how metallicity varies with Galactocentric radius R and height z . The all-star map already hints at a global decline in [M/H] away from the plane and the Solar circle. This decline is more pronounced in the low- α sequence.

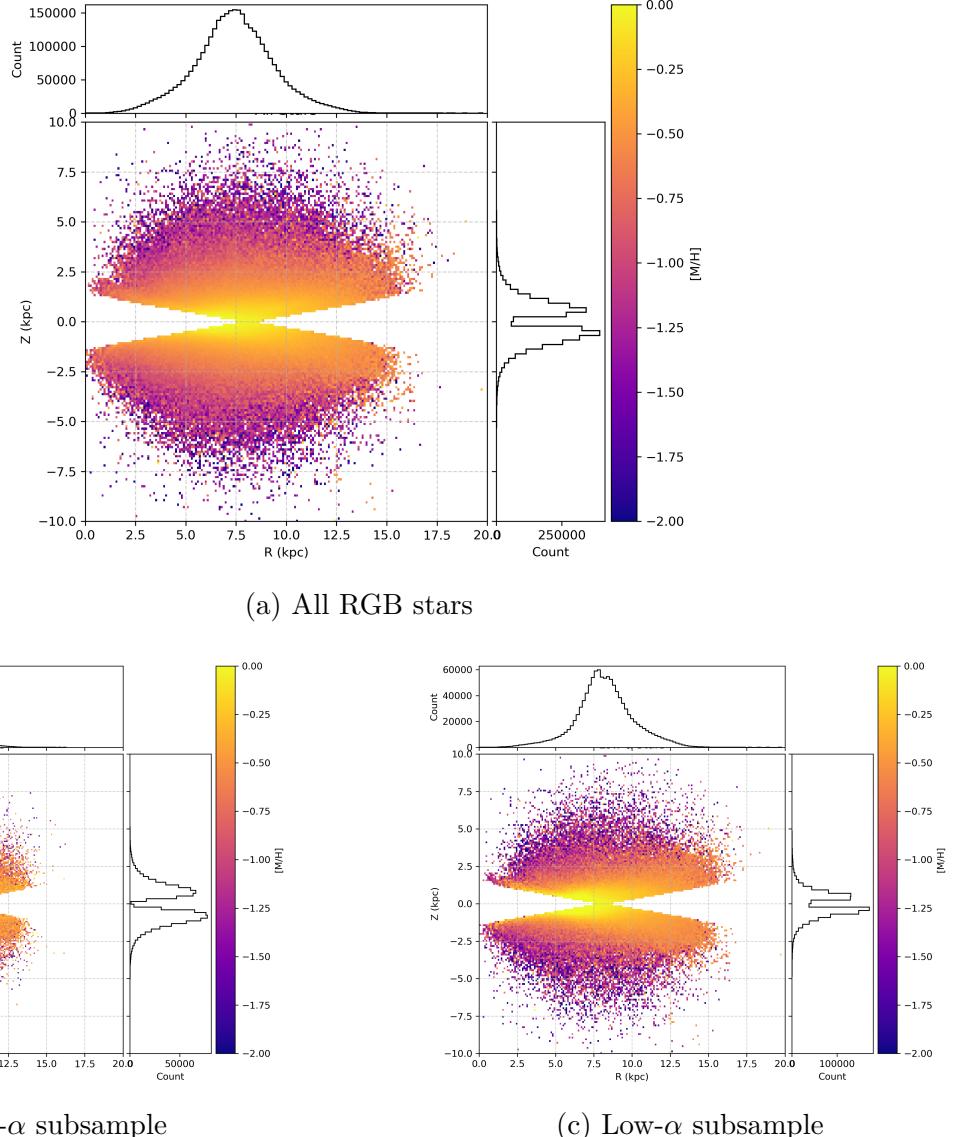


Figure 10: Mean [M/H] colour-coded in the Galactocentric R - z plane. *Top*: the complete $G < 16$ RGB sample. *Bottom left*: high- α stars; *bottom right*: low- α stars. The high- α population is more centrally concentrated and shows a shallower vertical metallicity gradient. (A reproduction of Figure 2 in [Viswanathan et al. \[2024\]](#))

As shown in Figure 11, high- α giants exhibit a gradual spin-up with increasing metallicity. At low metallicities ($[M/H] \lesssim -2$), the population shows mild net prograde rotation with large dispersion, consistent with a dynamically hot stellar halo. From $[M/H] \gtrsim -1.3$, there is a steady increase in rotation velocity and a corresponding decrease in dispersion, indicative of a transition toward the thick disc. By $[M/H] \gtrsim -0.5$, most stars occupy a kinematic regime characterised by elevated v_ϕ and moderate dispersion, consistent with a dynamically hot but rotationally supported thick disc. This continuous trend supports a scenario in which the Galaxy's in-situ population gradually acquired angular momentum



Figure 11: Azimuthal velocity as a function of metallicity. Each panel shows the column-normalised density in the $[M/H]$ – v_ϕ plane together with running medians (red dashed) and 16th/84th percentiles (black). Left: full RGB sample. Centre / right: chemically separated high- α and low- α subsamples. (A reproduction of Figure 3 in Viswanathan et al. [2024])

as star formation progressed.

In the low- α panel, stars with $[M/H] \lesssim -1.2$ exhibit kinematics similar to those in the high- α sequence, showing mild prograde rotation and high velocity dispersion. However, above $[M/H] \gtrsim -1.2$, the low- α sequence undergoes a sharper transition to a population with high net prograde rotation and lower velocity dispersion. By $[M/H] \gtrsim -0.5$, the low- α population exhibits kinematics consistent with the thin disc: a strongly rotating, dynamically cold component with lower dispersion than the high- α sequence.

In summary, the high- α sequence shows a gradual transition from a dynamically hot halo to a thick disc, while the low- α sequence shows a more abrupt transition to a dynamically cold disc at higher metallicities.

Following the method described in Section 2.2, we obtain full six-dimensional phase-space coordinates for the chemically separated high- and low- α subsamples. Figure 12 shows the distribution of stars in the v_R – v_ϕ plane across metallicity bins for the high- α (top) and low- α (bottom) populations.

The low- α sample exhibits a strong rotationally supported component, particularly at higher metallicities ($[M/H] \gtrsim -1.0$), with a narrow v_ϕ distribution centred around ~ 200 km/s and low radial velocity dispersion—clear hallmarks of thin-disc kinematics. As metallicity decreases, this rotational component becomes less pronounced, with broader distributions in both v_R and v_ϕ appearing at $[M/H] \lesssim -1.5$, consistent with increased contamination from halo or accreted populations.

The high- α stars, by contrast, show a continuous transition from a kinematically hot, low-rotation regime at low metallicities to a more rotation-supported disc at higher metallicities. This reflects the gradual build-up of the thick disc over time, as also noted in Chandra et al. [2024]. The velocity dispersion is broader throughout, indicating that high- α stars trace older, more dynamically heated populations.

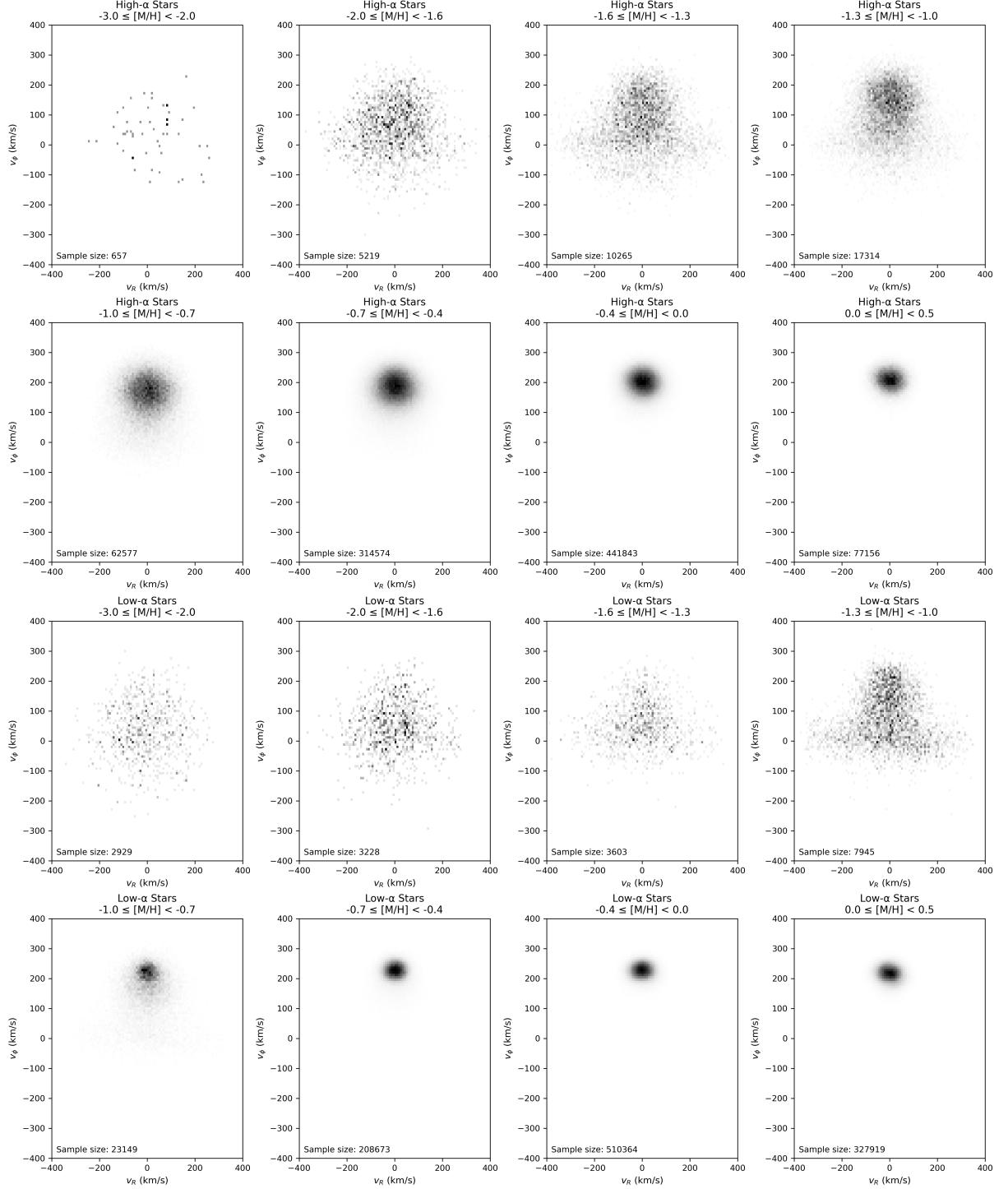


Figure 12: Distribution of stars in $v_R - v_\phi$ space across eight metallicity bins (increasing left to right, top to bottom) for the high- α population (top) and low- α population (bottom). The low- α stars are clearly more rotation-supported at high metallicities, while the high- α sample shows a broader velocity structure across all bins, consistent with thick-disc and halo populations.

5.3 Gaussian Mixture Model Fit

We follow the same Gaussian mixture model fitting procedure as in Section 4.1, but now separately for the high- and low- α sequences.

α -Sequence	VMP ($[M/H] < -2.0$)	IMP ($-2.0 < [M/H] < -1.3$)	MP1 ($-1.3 < [M/H] < -1.0$)	MP2 ($-1.0 < [M/H] < -0.7$)
High- α	1	4	5	3
Low- α	1	2	3	5

Table 2: Number of Gaussian Mixture components selected by the BIC for each metallicity bin, split by α -sequence.

Following the methodology described in Section 3.2, Table 2 summarises the number of components selected by the Bayesian Information Criterion for each metallicity bin. Detailed BIC scores for each metallicity bin are shown in Figure 13.

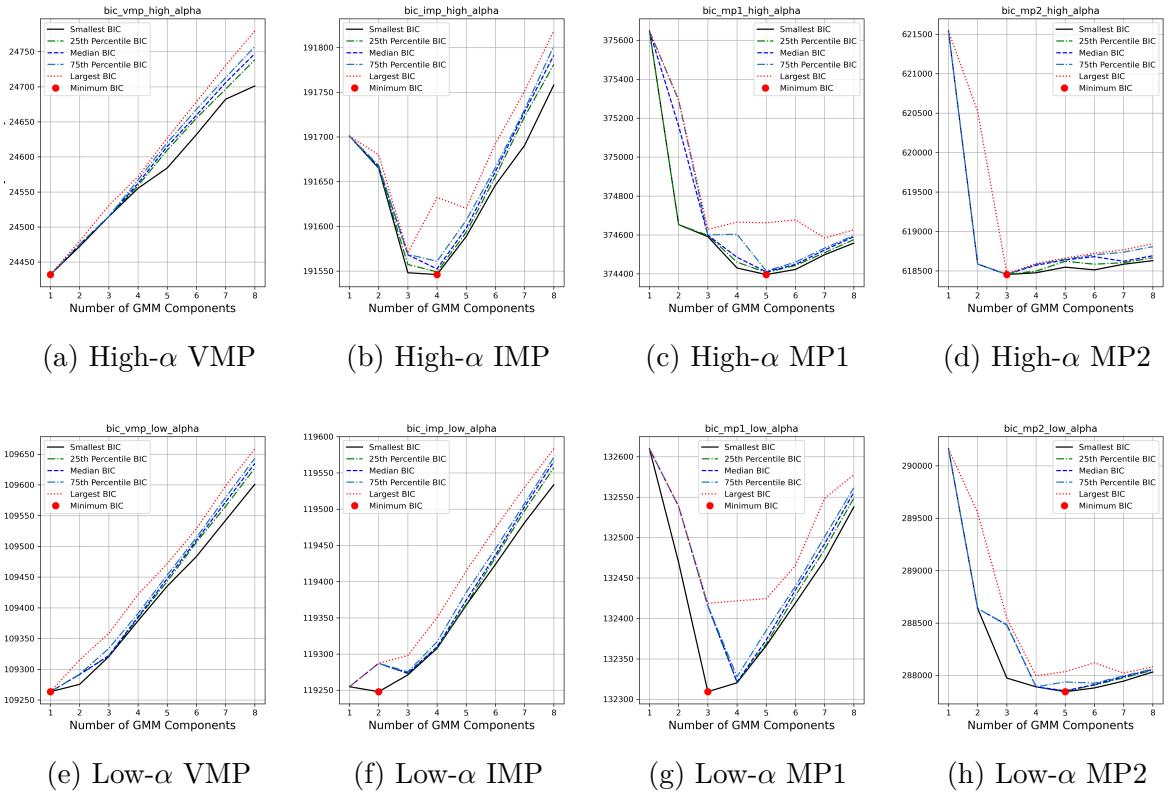


Figure 13: BIC scores for different metallicity bins, grouped by high- and low- α populations.

5.4 High alpha Gaussian Mixture Model Fit

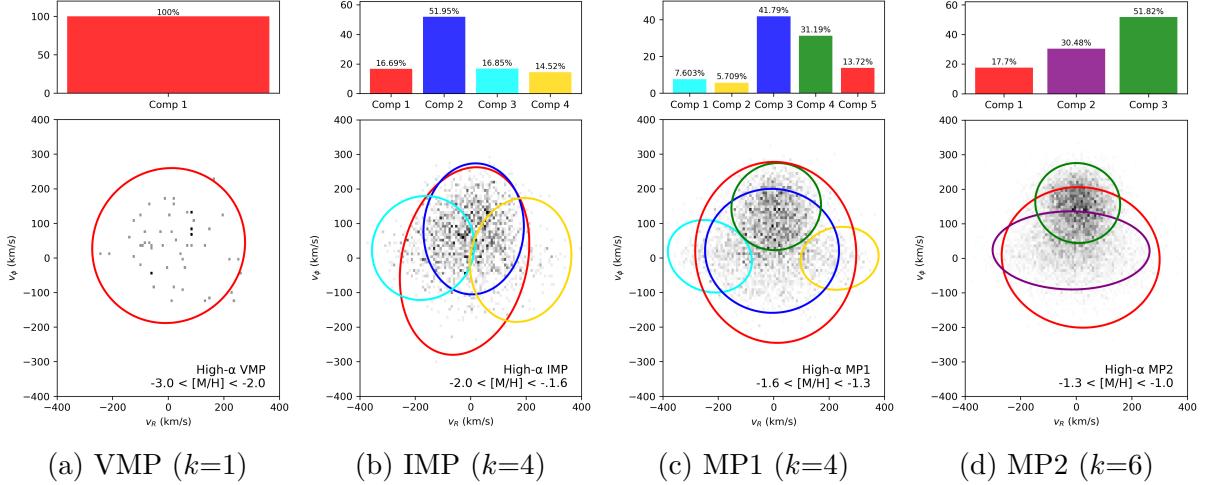


Figure 14: XD-GMM decomposition of high- α giants in four metallicity bins. Grey ellipses mark the 1σ contours of each Gaussian component.

Components	Weights (%)	v_R	σ_R	v_ϕ	σ_ϕ	v_Z	σ_Z
VMP: $-3.0 < [M/H] < -2.0$ (657 stars)							
Stationary halo	100.0	1.40	137.36	35.59	111.99	7.60	102.48
IMP: $-2.0 < [M/H] < -1.6$ (5219 stars)							
Stationary halo	16.7	-22.06	116.74	-8.55	135.74	-2.64	125.40
Prograde halo	51.9	10.84	90.56	84.51	94.79	-0.76	69.42
GS/E (1)	16.8	-171.35	92.33	29.12	75.23	0.50	87.09
GS/E (2)	14.5	180.48	91.48	-5.47	89.66	-2.87	94.37
MP1: $-1.6 < [M/H] < -1.3$ (10265 stars)							
Stationary halo	13.7	7.44	145.24	15.90	130.80	-8.97	127.11
Prograde halo	41.8	-5.85	120.77	20.31	89.59	-1.29	72.85
GS/E (1)	7.6	-230.09	75.81	4.65	52.41	8.69	94.68
GS/E (2)	5.7	237.61	70.43	-1.23	45.51	-3.73	92.81
Thick disc	31.2	9.18	80.79	148.23	62.97	2.99	68.92
MP2: $-1.3 < [M/H] < -1.0$ (17314 stars)							
Stationary halo	17.7	15.91	141.96	2.13	101.75	-1.68	98.90
GS/E	30.5	-18.50	141.08	22.42	56.55	-2.40	77.98
Thick Disk	51.8	5.22	76.27	159.54	57.94	-0.83	67.18

Table 3: GMM component weights and kinematics for the high- α sequence. Velocities are in km s^{-1} . Weights indicate the fraction of stars per metallicity bin.

5.5 Low alpha Gaussian Mixture Model Fit

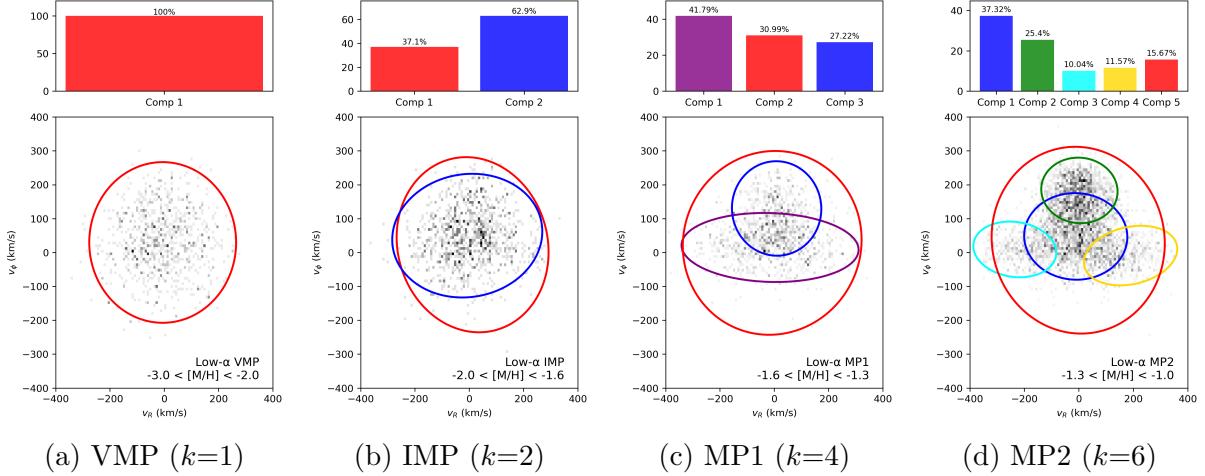


Figure 15: XD–GMM decomposition of low- α giants in four metallicity bins. Grey ellipses show the 1σ contours of every Gaussian component.

Components	Weights (%)	v_R	σ_R	v_ϕ	σ_ϕ	v_Z	σ_Z
VMP: $-3.0 < [M/H] < -2.0$ (2929 stars)							
Stationary halo	100.0	-5.82	135.18	29.55	118.56	1.34	107.10
IMP: $-2.0 < [M/H] < -1.6$ (3228 stars)							
Stationary halo	37.1	12.33	139.85	22.87	129.23	-4.43	120.51
Prograde halo	62.9	-7.34	138.11	49.68	91.25	4.08	75.99
MP1: $-1.6 < [M/H] < -1.3$ (3603 stars)							
Stationary halo	31.0	-7.58	164.57	28.31	135.62	-4.92	121.66
GS/E	41.8	-15.50	162.97	14.71	50.94	-1.93	73.60
Thick Disk/prograde halo	27.2	8.72	82.25	129.87	69.71	3.13	60.68
MP2: $-1.3 < [M/H] < -1.0$ (7945 stars)							
Stationary halo	15.7	-2.30	158.84	36.24	137.81	2.81	131.10
Prograde halo	37.3	-11.31	94.62	47.39	63.89	-1.18	63.12
GS/E (1)	10.0	-234.89	76.09	8.99	41.09	10.70	93.79
GS/E (2)	11.6	189.95	85.58	-9.15	43.99	-8.49	89.69
Thick disc	25.4	2.13	69.98	183.15	48.36	1.75	54.21

Table 4: XD–GMM component weights and kinematics for the low- α sequence. Velocities are in km s^{-1} . Weights are the fractional contribution of each component.

5.6 Contrasting the high and low α GMMs

The two chemical sequences trace very different roads from a dispersion-supported halo to rotationally supported discs.

Very-metal-poor regime (VMP; $[M/H] < -2$). Both data sets are best described by a *single*, non-rotating halo Gaussian. The velocity ellipsoids are close to isotropic ($\sigma_R : \sigma_\phi : \sigma_Z \approx 1 : 1 : 1$), with the high- α stars showing a slightly smaller σ_ϕ , hinting at an early prograde bias. Compared to the results in Section 4.1, where two components

were identified in this metallicity regime, the emergence of only a single Gaussian here may reflect the smaller sample size, limiting the ability of the model to resolve multiple distinct halo structures.

Intermediate regime (IMP; $-2.0 < [\text{M}/\text{H}] < -1.6$). The low- α sample already splits into stationary and prograde halo Gaussians (37% vs. 63%), whereas the high- α mixture needs *four* hot components, two of which have the highly radial kinematics characteristic of Gaia–Sausage/Enceladus (GSE). The GSE-like signal in the high- α branch is unexpected, as the GSE merger originated from a dwarf galaxy with a slower star formation history and is therefore expected to contribute primarily α -poor stars. Its appearance here likely reflects mild contamination from $[\alpha/\text{M}]$ systematics at low signal-to-noise rather than a truly α -rich GSE population.

Disc onset (MP1; $-1.6 < [\text{M}/\text{H}] < -1.3$). In both sequences a prograde component with a low dispersion compatible with a thick disk emerges, but with different prominence and kinematics. In the high- α mixture, the thick disc accounts for 31% of the stars, rotating at $v_\phi \sim 148 \text{ km s}^{-1}$, indicating the emergence of a warm, *in-situ* rotating structure. The low- α sample is dominated by non-disc components: the stationary halo and GSE account for 72.8% of the stars, while the remaining 27.2% are captured by a thick-disc/prograde halo Gaussian with $v_\phi \sim 130 \text{ km s}^{-1}$ and higher dispersions. This suggests that α -rich stars settled into a disc structure earlier, while α -poor stars at this metallicity remain primarily associated with accreted debris (GS/E).

Transition to disc dominance (MP2; $-1.3 < [\text{M}/\text{H}] < -1.0$). The high- α fit simplifies to three Gaussians—thick disc (52 %), GSE-like halo (31 %), and a small stationary halo (17 %). The low- α solution remains complex (five Gaussians) but now contains a substantial thick disc (25 %) whose lower dispersion ($\sigma_\phi \sim 48 \text{ km s}^{-1}$) and high v_ϕ are consistent with an early thin-disc precursor. The ultra-cold ($\sigma \lesssim 30 \text{ km s}^{-1}$) thin disc is *absent* because our latitude/extinction cuts remove most mid-plane stars.

Overall, the high/low alpha GMM analysis shows that the high- α stars transition smoothly from a dispersion-supported halo to a rotationally supported thick disc, while the low- α stars remain halo-dominated to higher metallicity before eventually forming a colder disc component. One surprising outcome is the emergence of GSE-like components in the high- α sequence, particularly at intermediate metallicities. Given that the Gaia–Sausage/Enceladus progenitor is thought to have formed stars slowly, enriching predominantly in iron via Type Ia supernovae, its stellar debris should be α -poor. The appearance of GSE-like kinematics among high- α stars is therefore unexpected and may reflect contamination due to measurement uncertainties or classification errors in $[\alpha/\text{M}]$, especially low signal to noise data. Alternatively, some *in-situ* stars may share similar orbits, making it challenging for the GMM to disentangle overlapping kinematic substructures.

6 Discussion

7 problems with Zhang et al. [2024]

i do not believe random initial coditions are used for bic and gmm fitting in order to get results with astrophysical meaning, k means inistialisation is required this is because gmms are very senstitive to initial conditions, and can easily get stuck in local minima.

8 Conclusion

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