

Formation of Planetary Systems

Lecture 1 - Observations of planetary systems

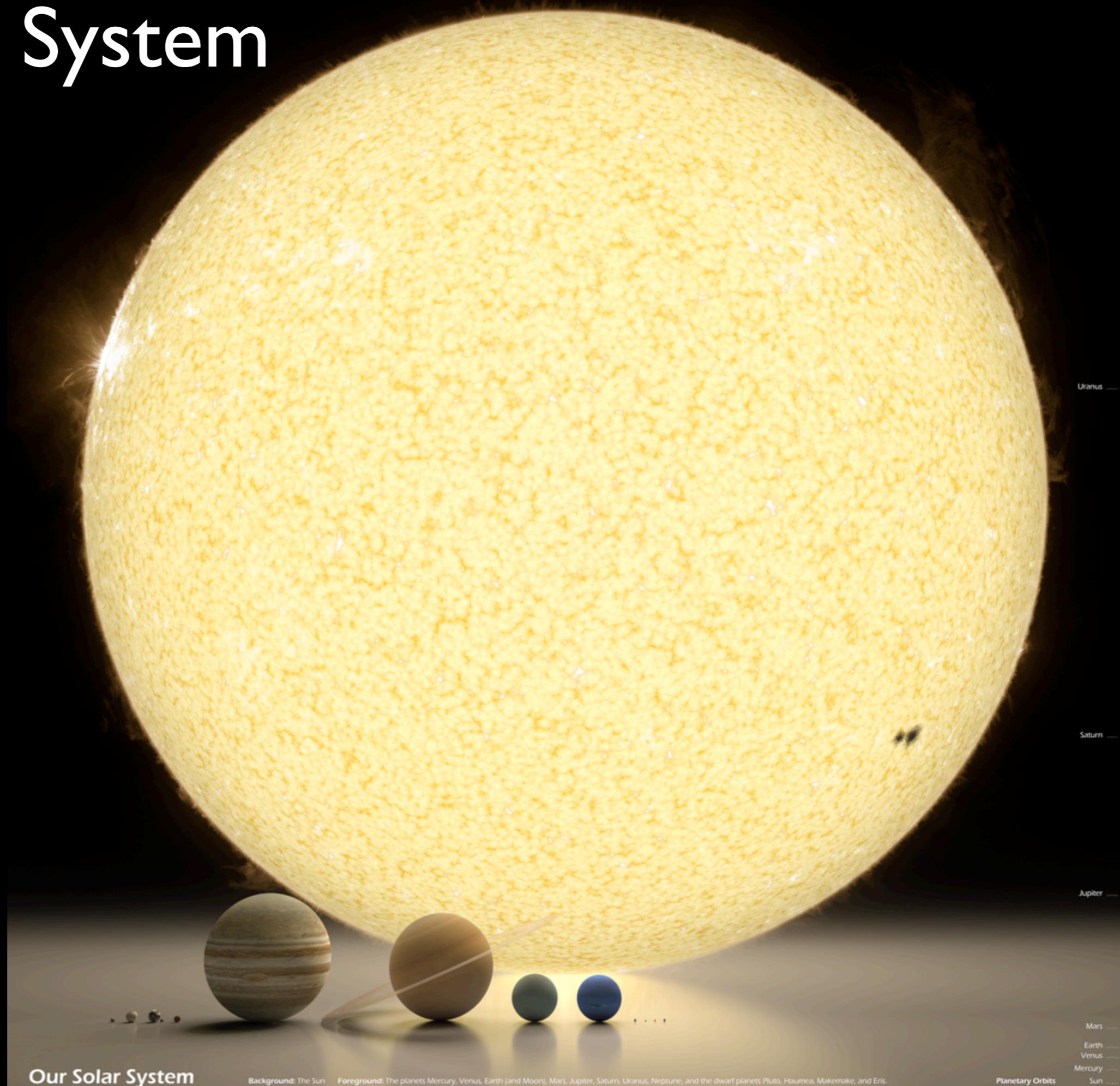


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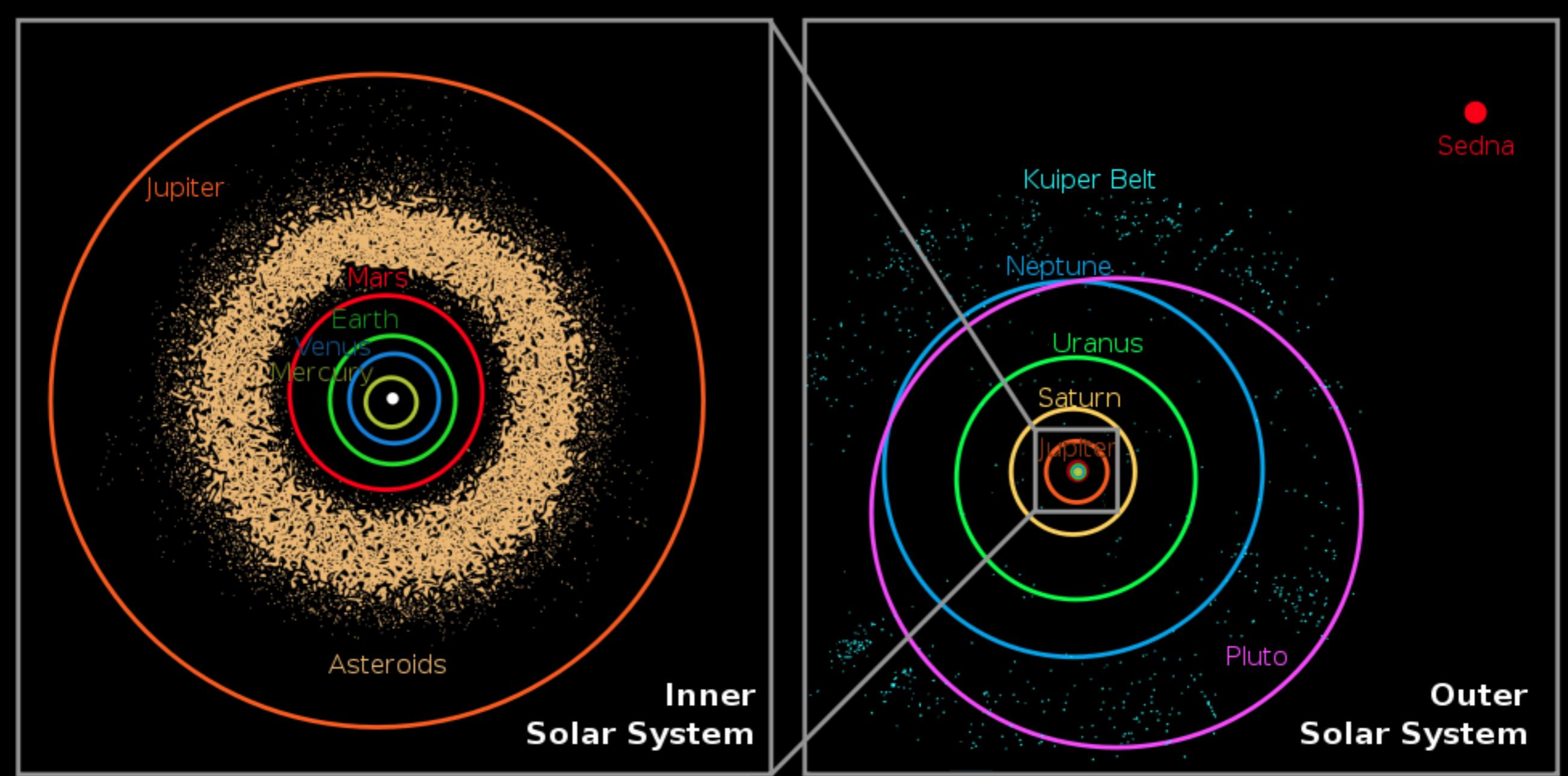
Course Outline

- 5 Lectures, 2 hours each (with a break in the middle!).
 - 1) Observations of planetary systems
 - 2) Protoplanetary discs
 - 3) Dust dynamics & planetesimal formation
 - 4) Planet formation
 - 5) Planetary dynamics
- Notes for each lecture will be placed on the course home page in advance - you may find it useful to annotate these as we go.
- These slides will also be posted online.
- Textbooks: Armitage - *Astrophysics of planet formation* (CUP).
Protostars & Planets series (VI - 2014; VII - 2023)

The Solar System



The Solar System



The Solar System

	a AU	e	M_p M_{Jup}
Mercury	0.387	0.206	1.74×10^{-4}
Venus	0.723	0.007	2.56×10^{-3}
Earth	1.000	0.017	3.15×10^{-3}
Mars	1.524	0.093	3.38×10^{-4}
Jupiter	5.203	0.048	1.00
Saturn	9.537	0.054	0.299
Uranus	19.19	0.047	0.046
Neptune	30.07	0.009	0.054

The Solar System

- Gas giants (Jupiter & Saturn):
 - massive: >90% of total planetary mass.
 - primarily H/He, but metal-rich w.r.t. Sun.
 - $\sim 10M_{\text{Earth}}$ solid cores (probably!).
- Ice giants (Uranus & Neptune):
 - H_2O , NH_3 , CH_4 , etc.
 - $\sim 1 M_{\text{Earth}}$ solid cores.
- Terrestrial planets (Mercury, Venus, Earth, Mars).
- Minor bodies: “dwarf planets”, moons, asteroids, comets, Kuiper belt, Oort cloud.
- All 8 planets are nearly co-planar, with near-circular orbits.

The Solar System

- >99% of total mass resides in the Sun.
- >99% of total angular momentum resides in the planets (mostly in Jupiter).
- Planets very metal-rich w.r.t. Sun (though majority of heavy elements are in the Sun).
- Radioactive dating (e.g. $^{87}\text{Rb} \rightarrow ^{87}\text{Sr}$) finds age of 4.57Gyr.
- **Planet formation processes must:**
 - **grow solid bodies from ISM grains to $>\text{M}_{\text{Earth}}$.**
 - **separate mass from angular momentum.**
 - **separate metals from H/He.**

Methods of detecting extra-solar planets

- Directly:
 - Light emitted/reflected by planet
direct imaging
- Indirectly:
 - Motion of star due to planet
astrometry
radial velocity
timing methods
 - Obscuration of stellar light by planet
transits
 - Obscuration/amplification of background star by planet
gravitational microlensing

Methods of detecting extra-solar planets

• Directly:	- Light emitted/reflected by planet direct imaging	69
• Indirectly:	- Motion of star due to planet astrometry	3
	- Motion of star due to planet radial velocity	1068
	- Motion of star due to planet timing methods (inc TTVs)	43
	- Obscuration of stellar light by planet transits	4128
	- Obscuration/amplification of background star by planet gravitational microlensing	204

Methods of detecting extra-solar planets

- Directly:

- Light emitted/reflected by planet

direct imaging

69

- Indirectly:

- Motion of star due to planet

astrometry

3

radial velocity

1068

timing methods (inc TTVs)

43

- Obscuration of stellar light by planet

transits

4128

- Obscuration/amplification of background star by planet

gravitational microlensing

204

Total: 5535

Direct Imaging

- Planets are very faint. How faint?



- Fraction of star-light reflected by planet is*:

$$f = A \left(\frac{\text{Cross-sect. area of planet}}{\text{Area of sphere radius } a} \right) = A \left(\frac{\pi R_p^2}{4\pi a^2} \right)$$

$$\Rightarrow f_{\oplus} \simeq 2 \times 10^{-10} \quad f_{Jup} \simeq 1 \times 10^{-9}$$

- Two problems for detecting in exo-planetary systems:
brightness and **contrast**. Contrast is usually dominant.

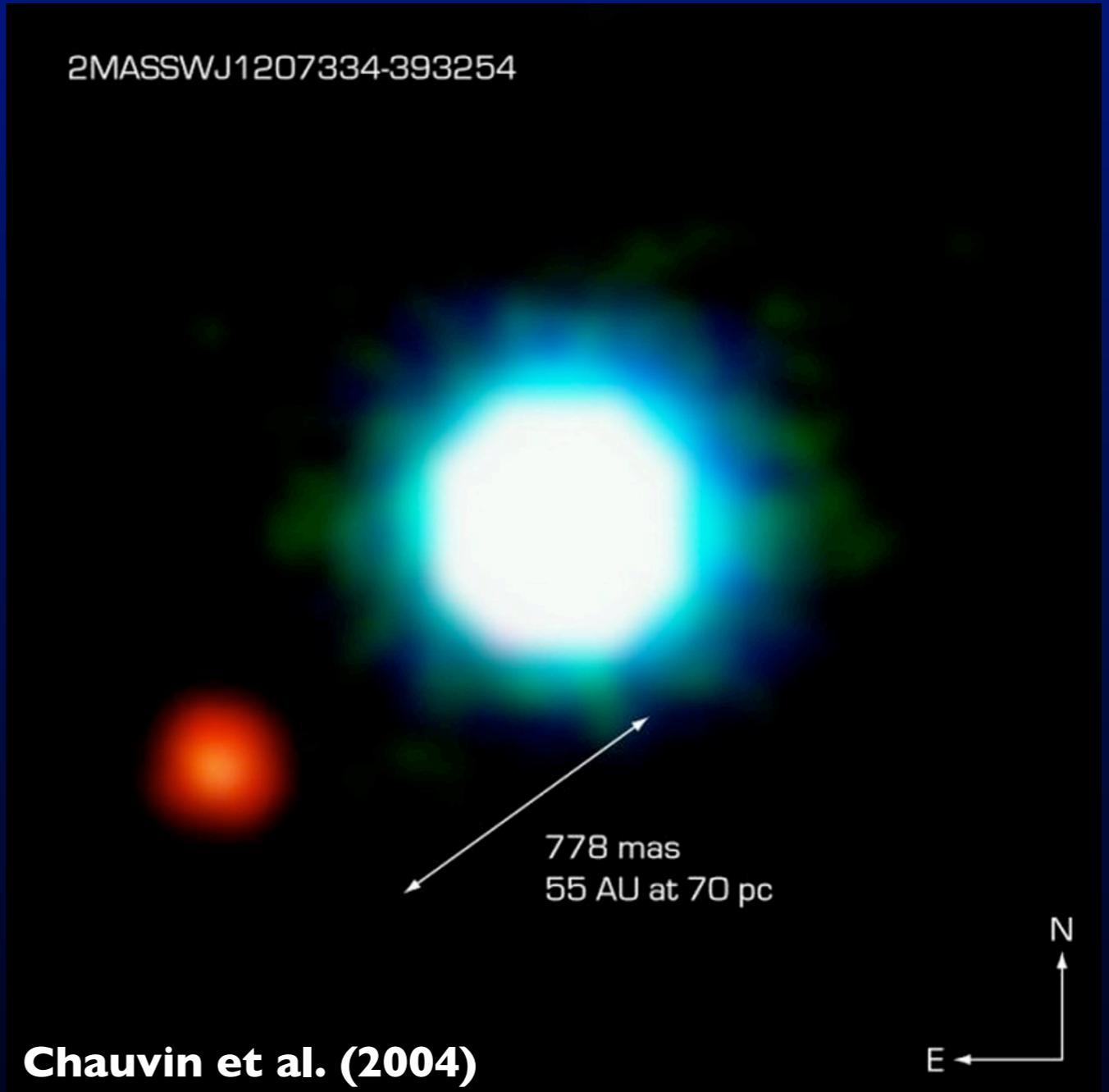
* A is the albedo.

Direct Imaging

- Two ways around the contrast problem:
 - a) Look for planets around faint stars
 - b) Try to mask out star-light

Direct Imaging

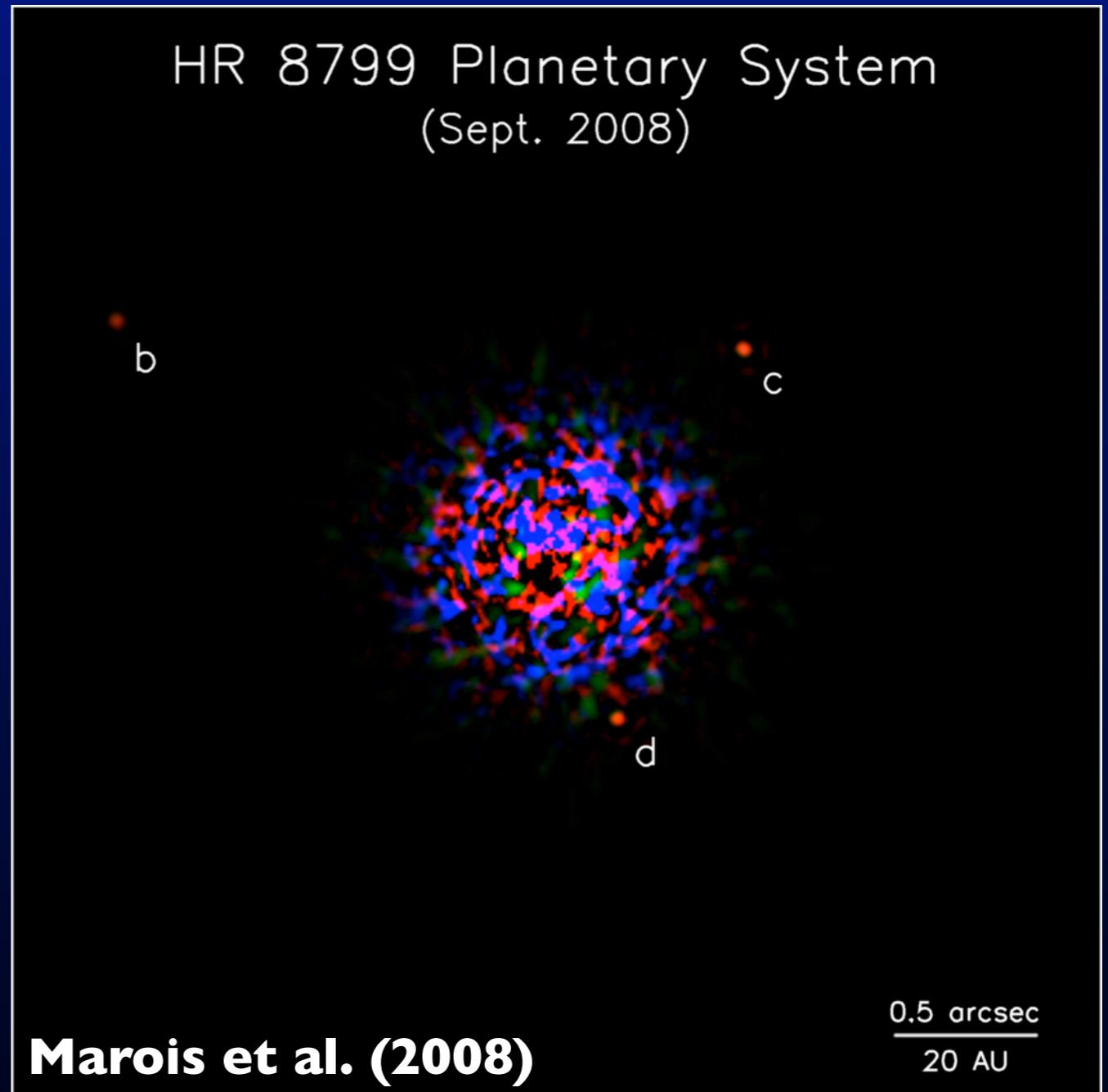
- Two ways around the contrast problem:
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“Planet” around brown dwarf 2M1207 discovered in 2004.
Primary is $\sim 25 M_{Jup}$; secondary is $\sim 5 M_{Jup}$. Wide separation.
More akin to a low-mass binary than a true planetary system.

Direct Imaging

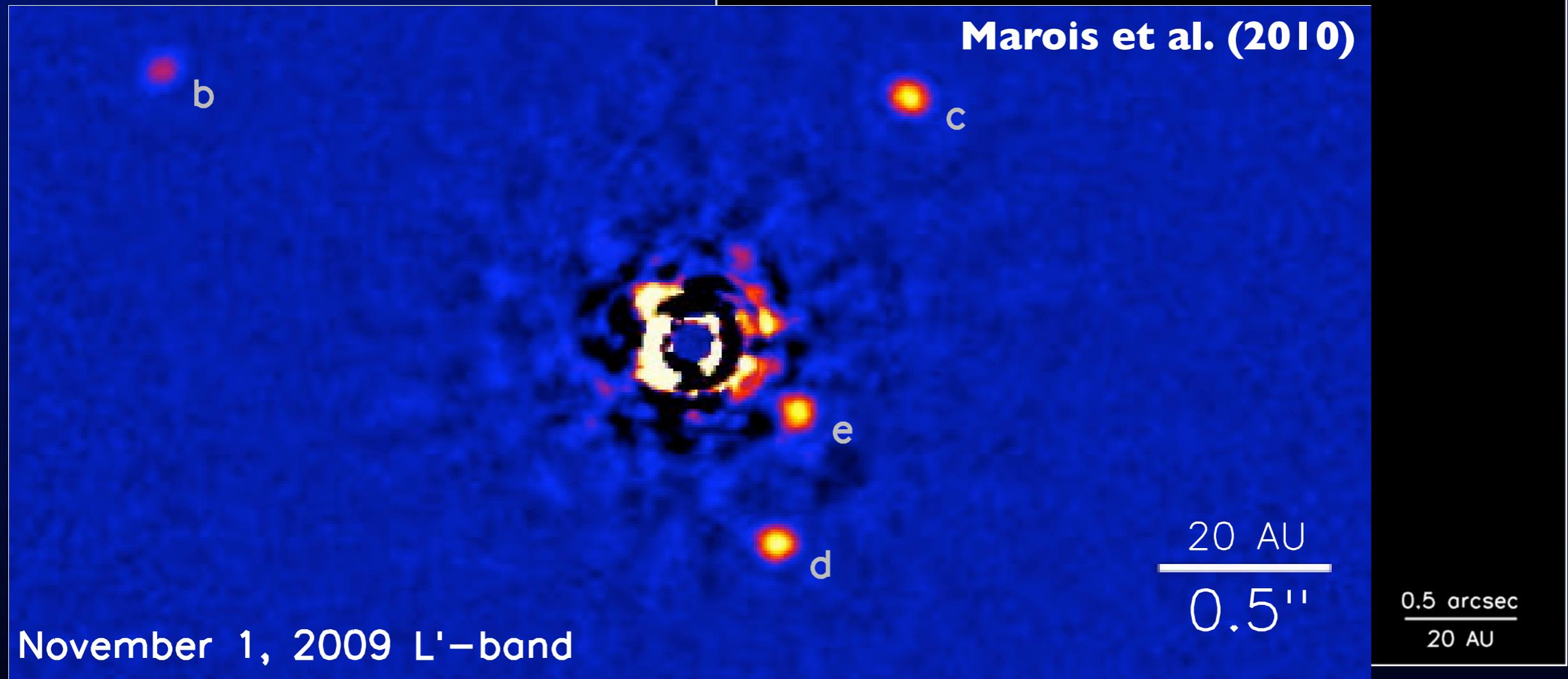
- Two ways around the contrast problem:
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Planets around HR8799 discovered in 2008.
Star is $\sim 1.5M_{\odot}$. Planet masses all estimated to be $\sim 10M_{Jup}$.
Wide orbits - “d” is beyond orbit of Uranus.

Direct Imaging

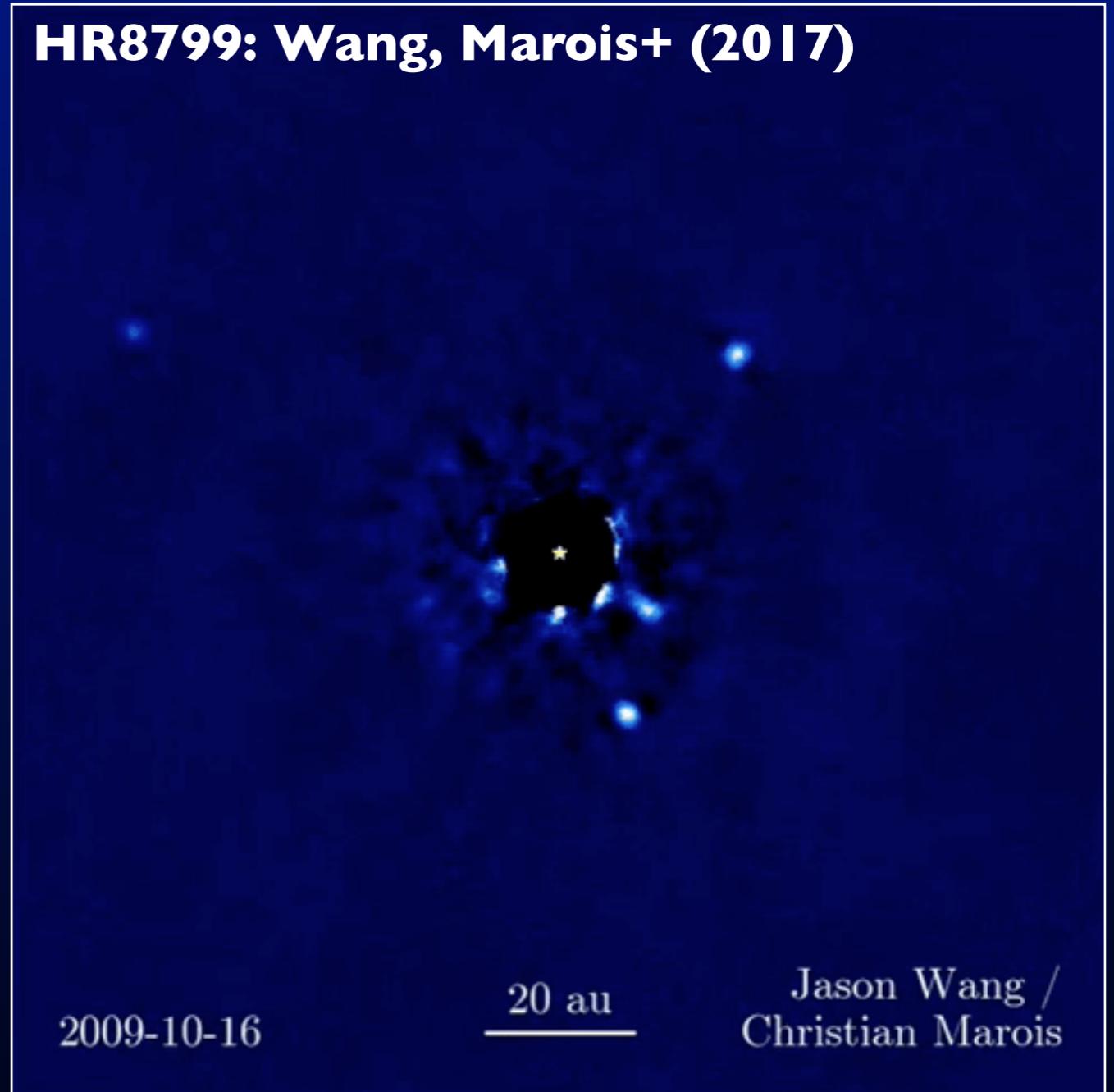
- Two ways around the contrast problem:



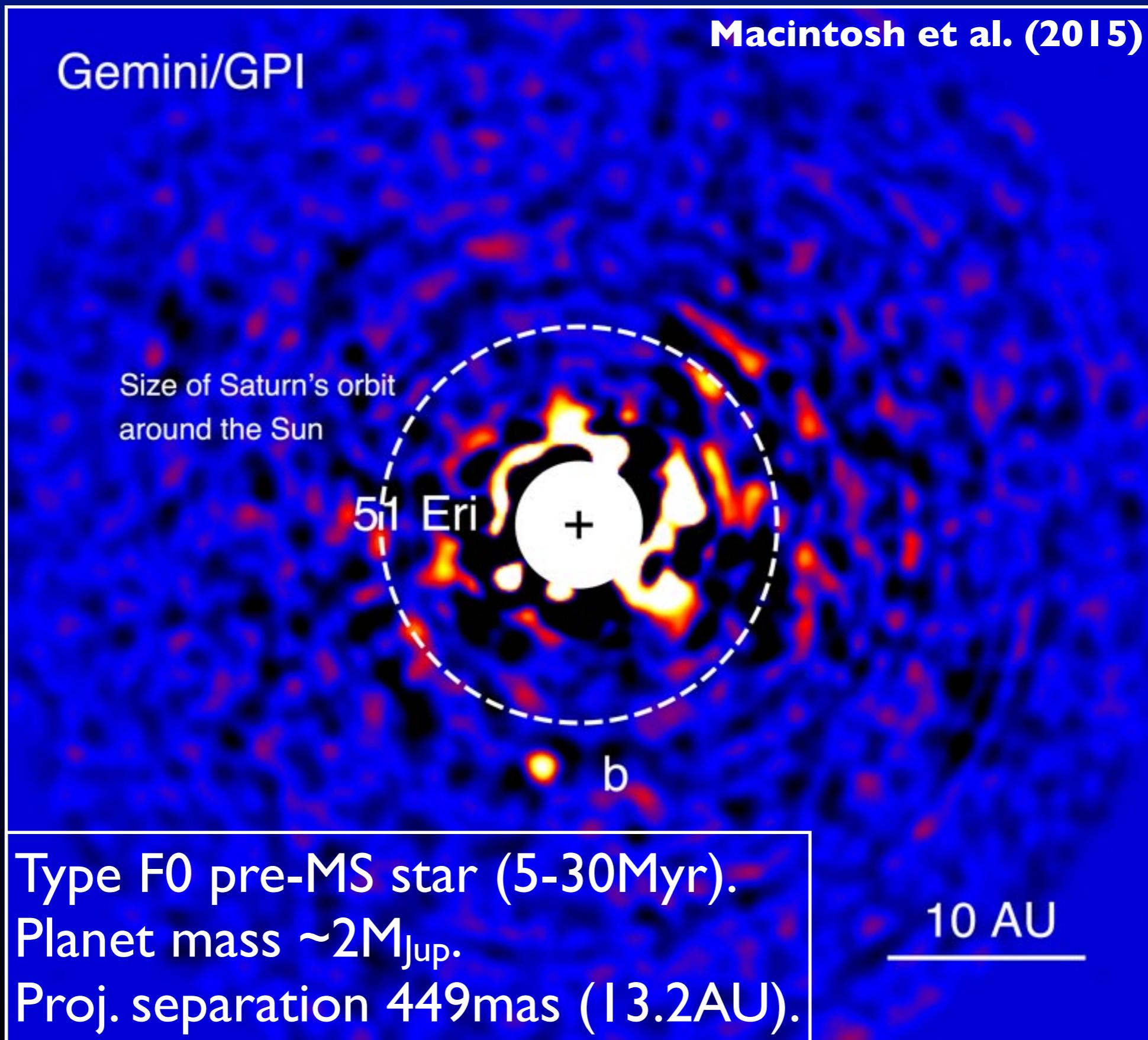
Direct Imaging

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HR8799: Wang, Marois+ (2017)

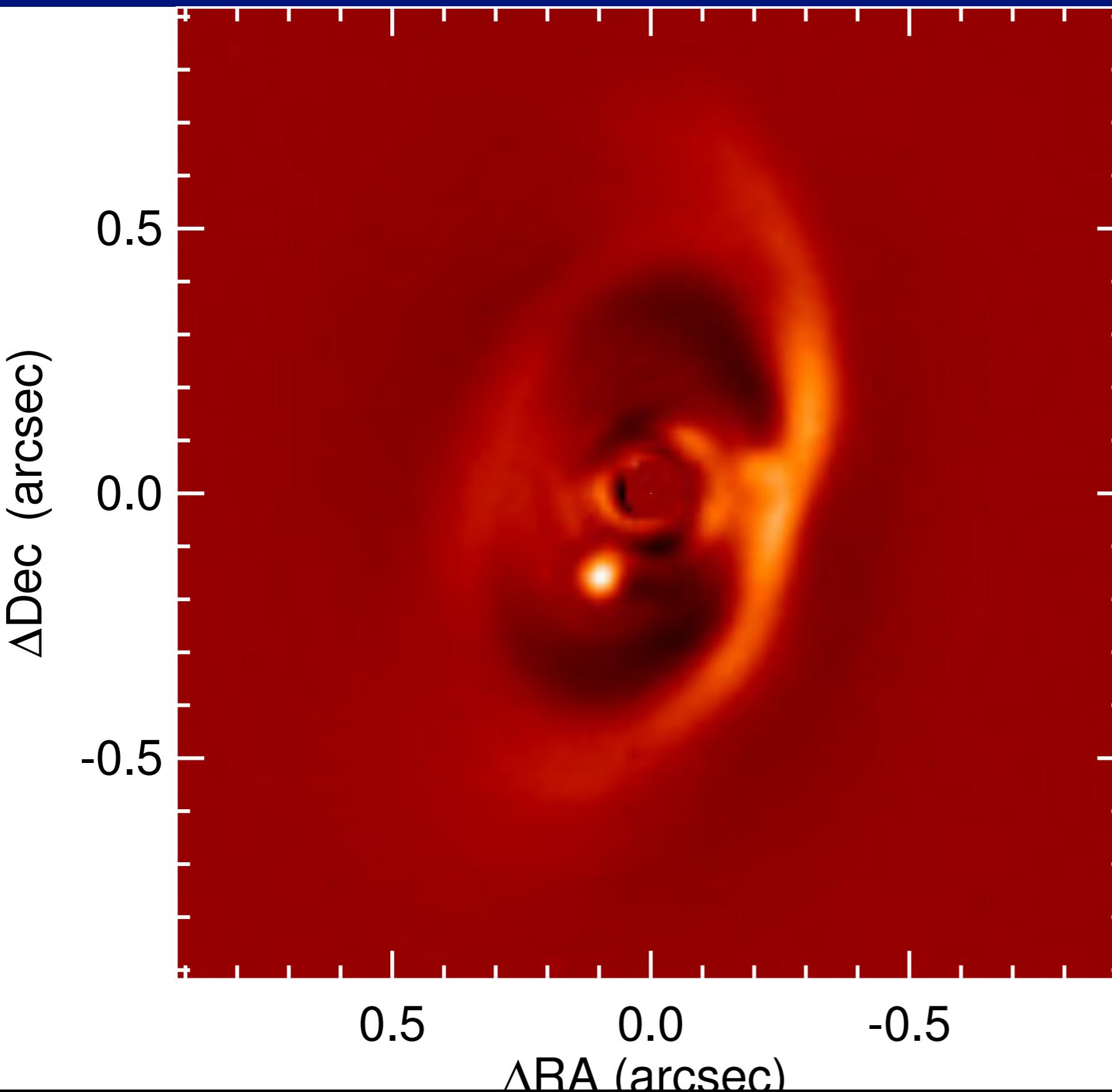


New facilities...



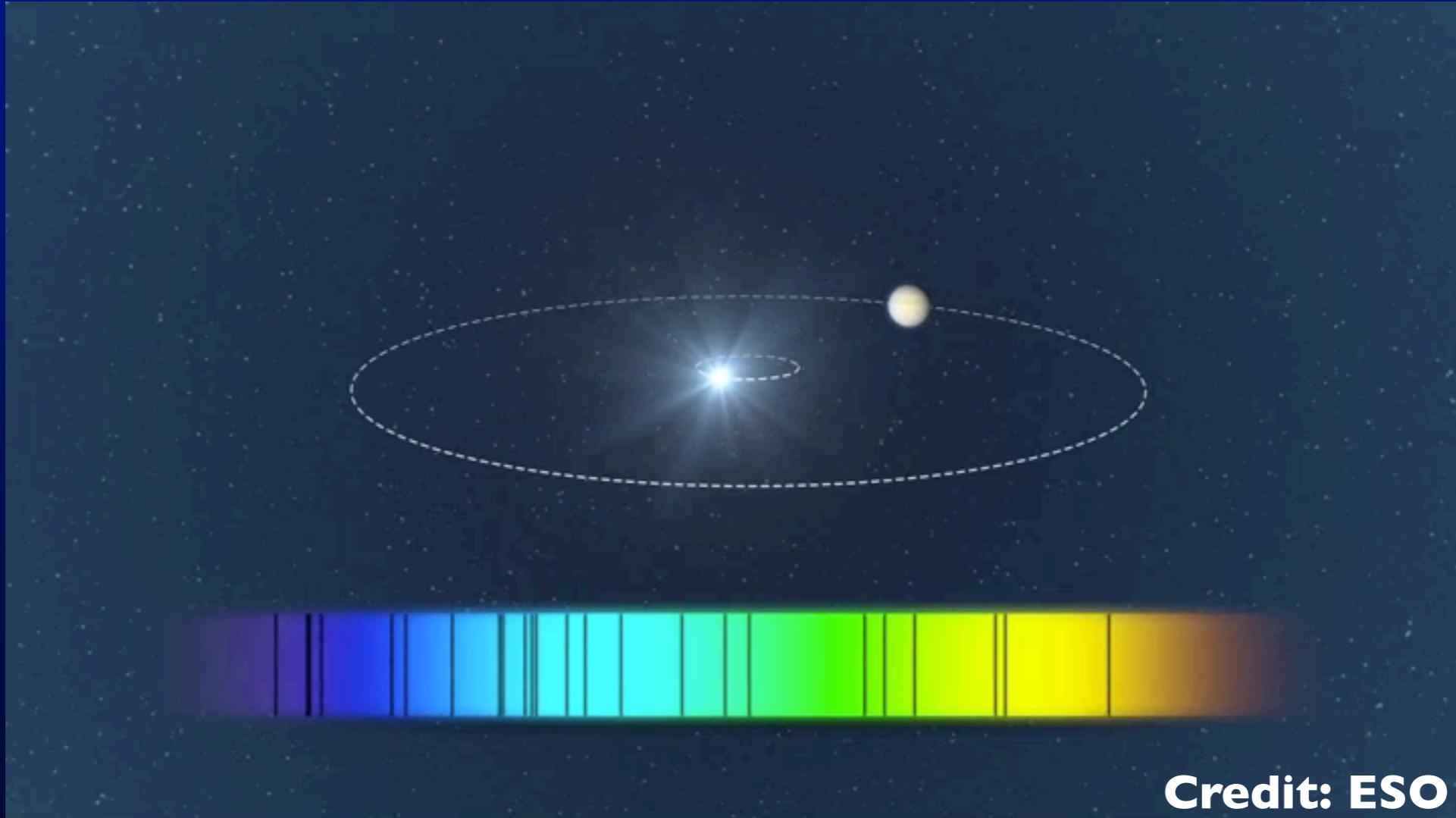
Young planets

PDS 70
Keppler et al. (2018)



T Tauri star (5.4Myr).
Planet mass $\sim 5M_{Jup.}$.
Proj. separation of
195mas (22AU).

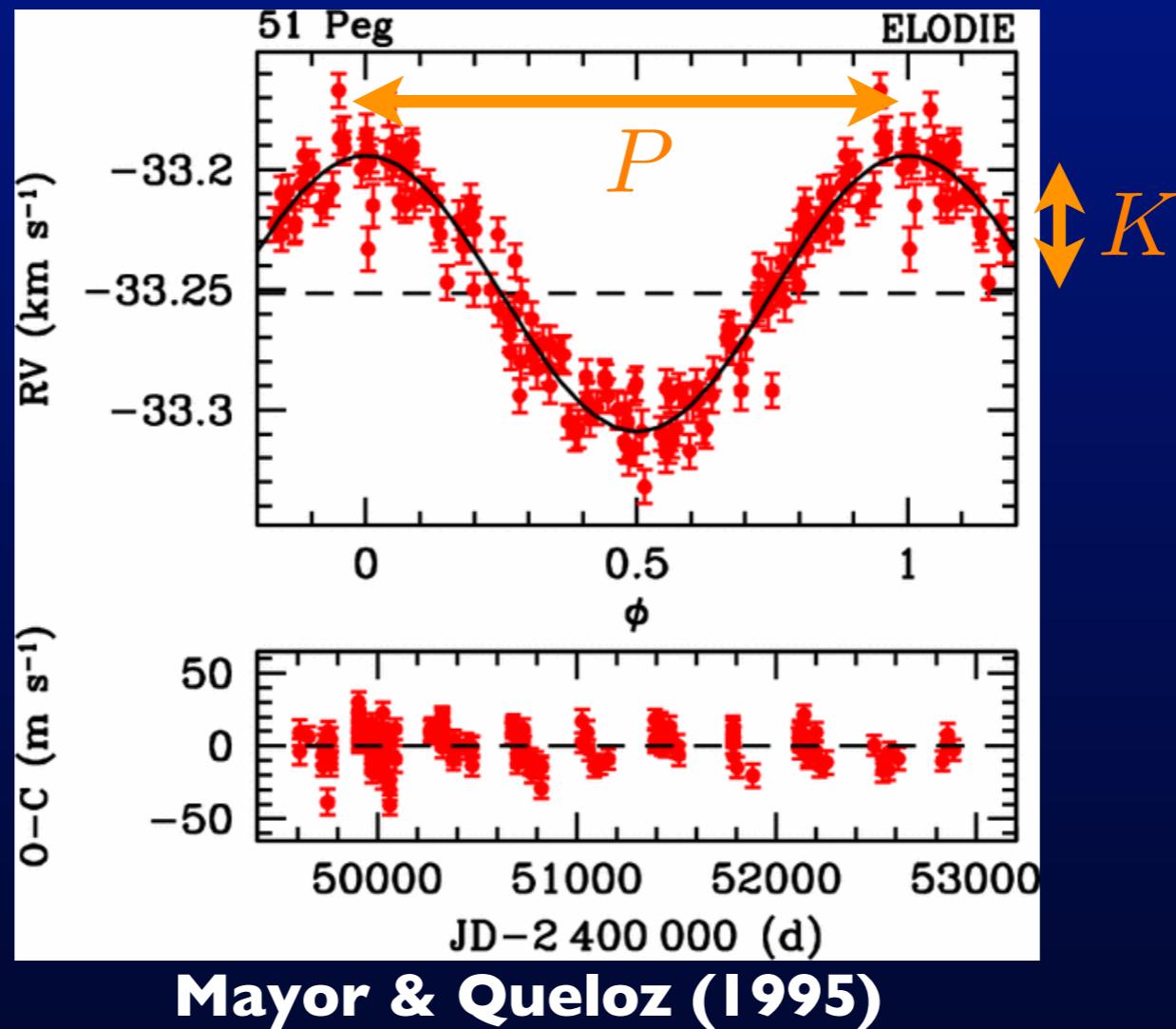
Radial velocity methods



Credit: ESO

- Look for Doppler shifts caused by stellar reflex motion.
- RV surveys on-going since first detection in 1995. Now ~500 detections: until *Kepler*, was most successful detection method.
- Originally pioneered by Latham, Mayor, Griffin and others. Most discoveries have come from two groups: Geneva & Lick/California.

Radial velocity methods

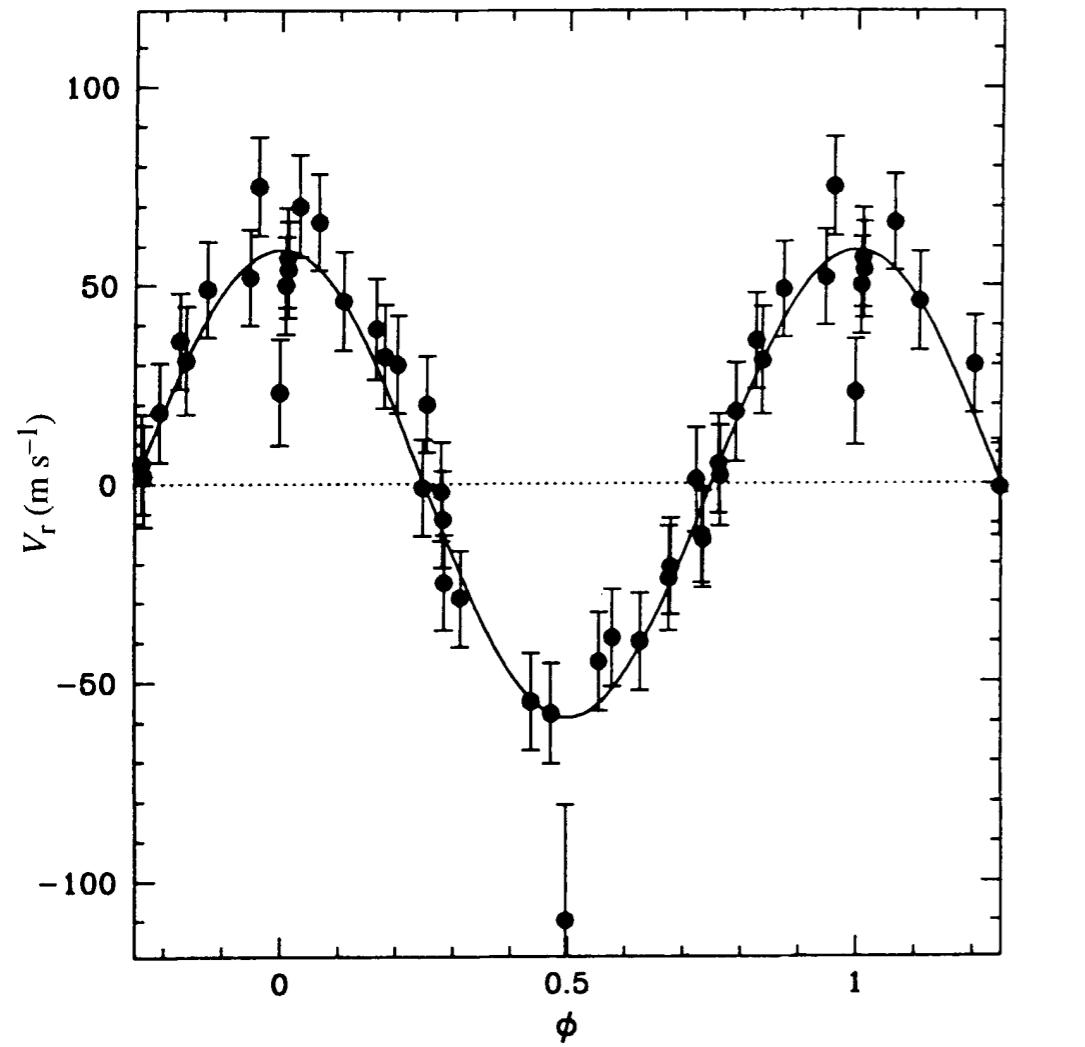


- Fit semi-major axis a , eccentricity e , and stellar mass $M_p \sin i$:

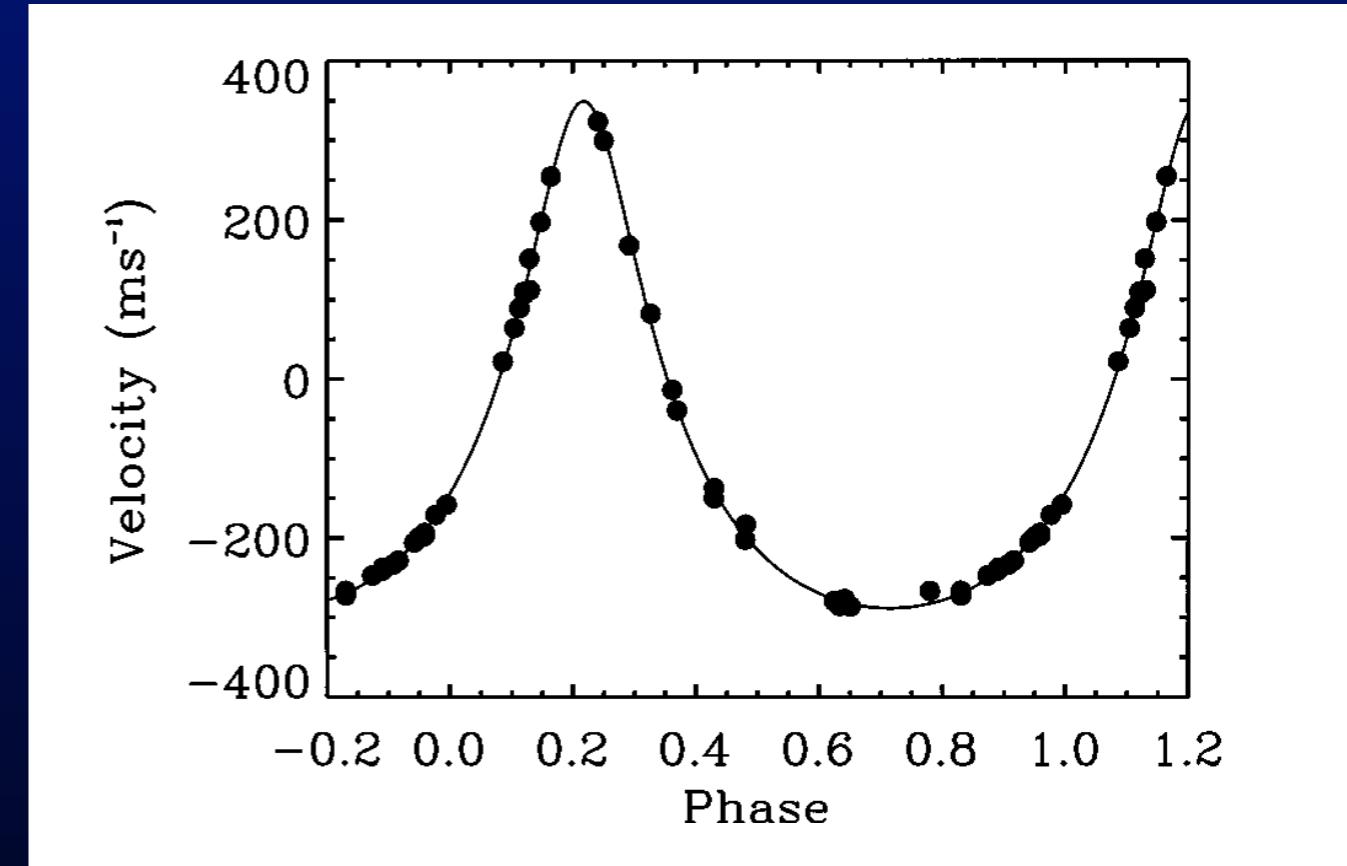
$$K = v_* \sin i = \frac{1}{\sqrt{1 - e^2}} \frac{M_p \sin i}{M_*} \sqrt{\frac{GM_*}{a}}$$

- $K_{\text{Jup}} \sim 12 \text{ m/s}$; $K_{\text{Earth}} \sim 10 \text{ cm/s}$.

First detections...

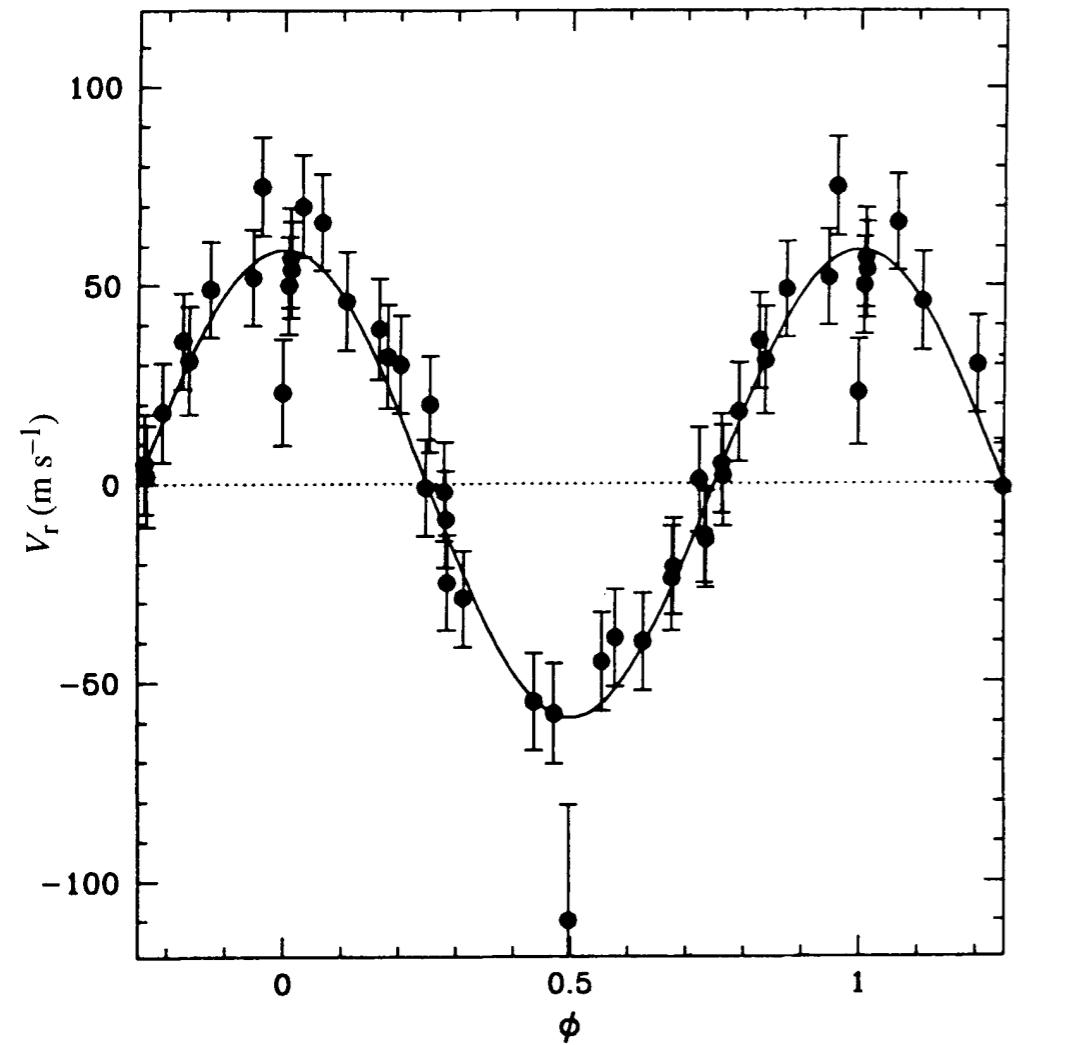


51 Peg b: Mayor & Queloz (1995)
Planet mass $0.47\text{M}_{\text{Jup}}$, Period 4.23d

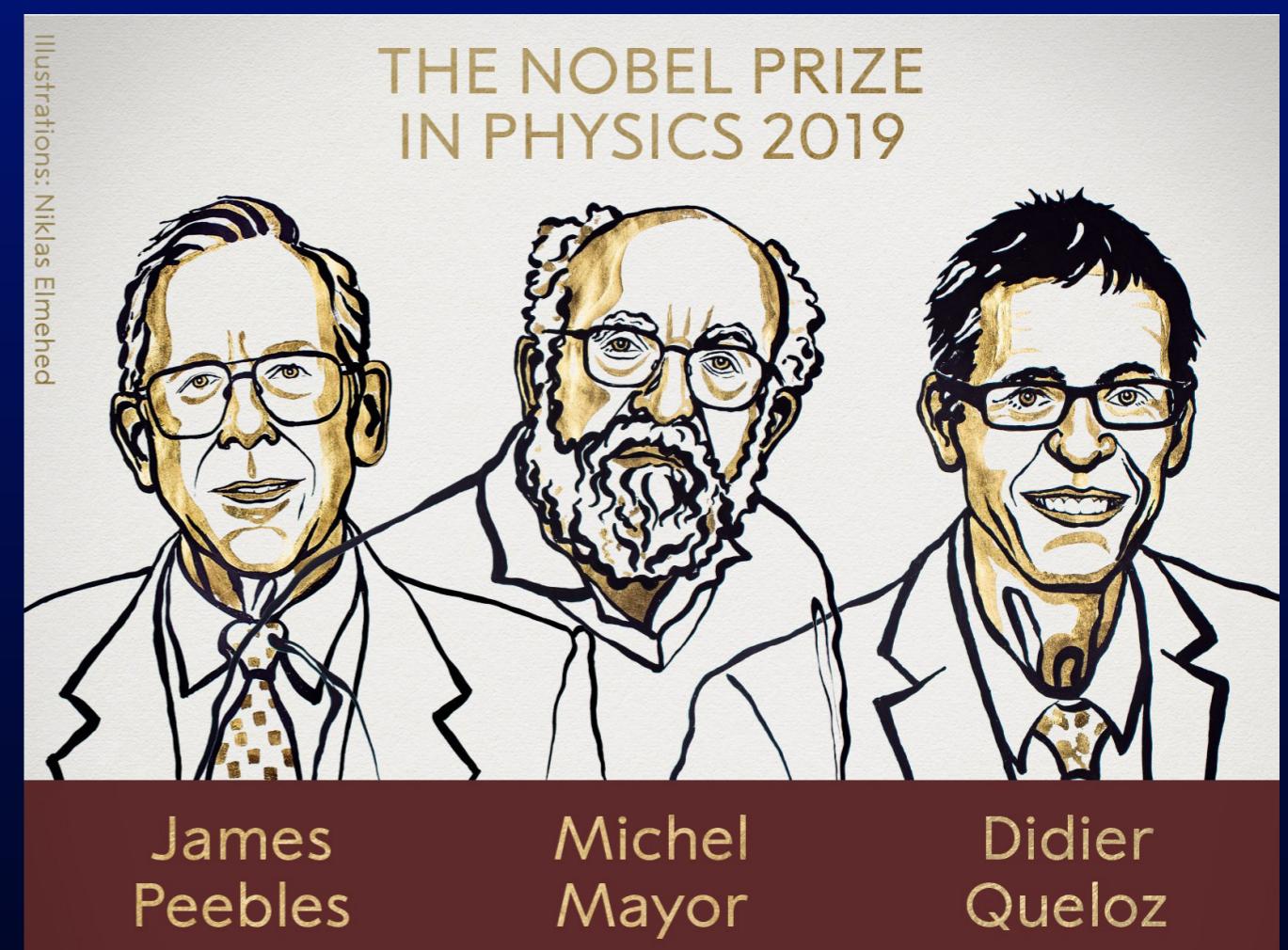


70 Vir b: Marcy & Butler (1996)
Planet mass 7.5M_{Jup} , Period 117d

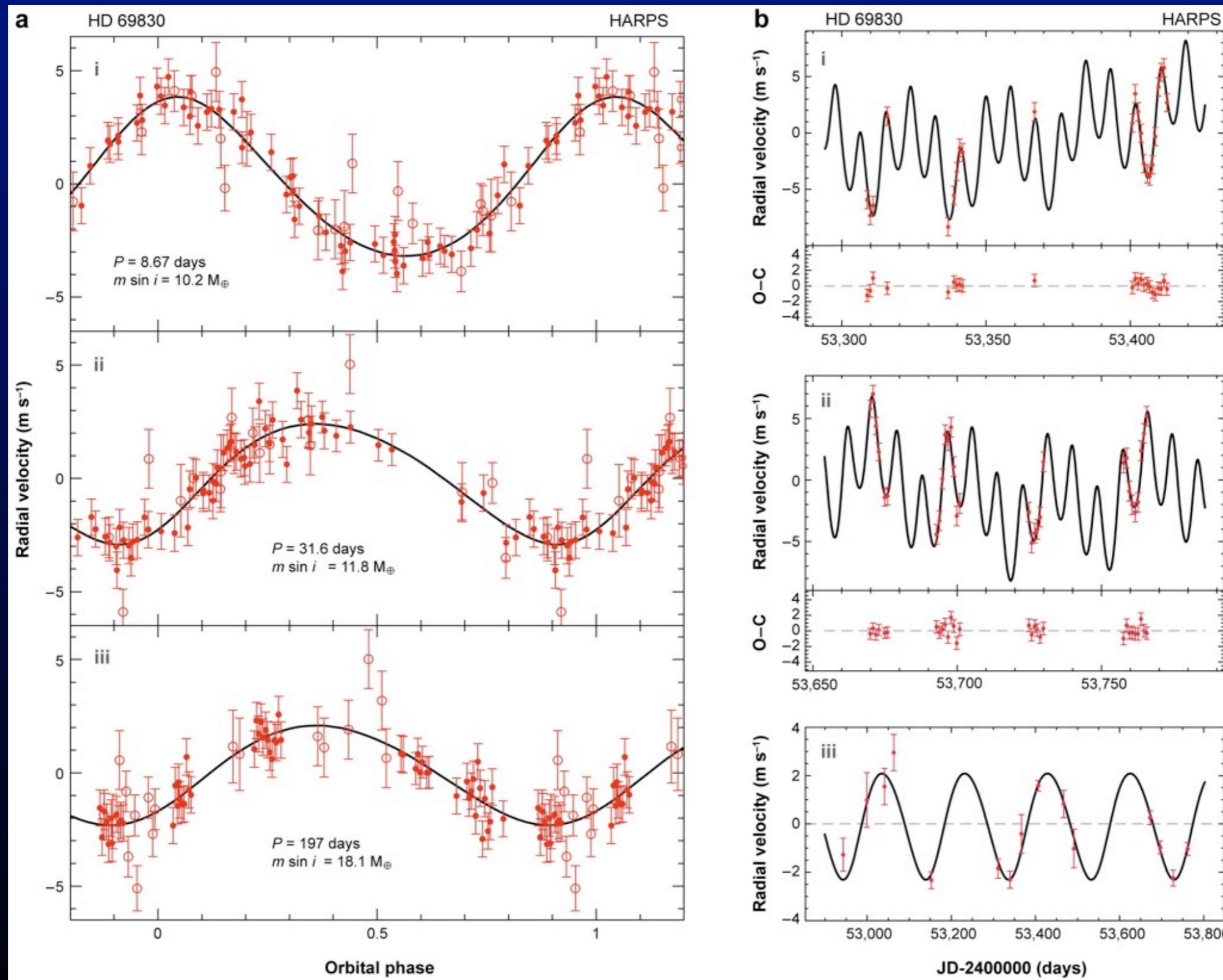
First detections...



51 Peg b: Mayor & Queloz (1995)
Planet mass $0.47 M_{\text{Jup}}$, Period 4.23d

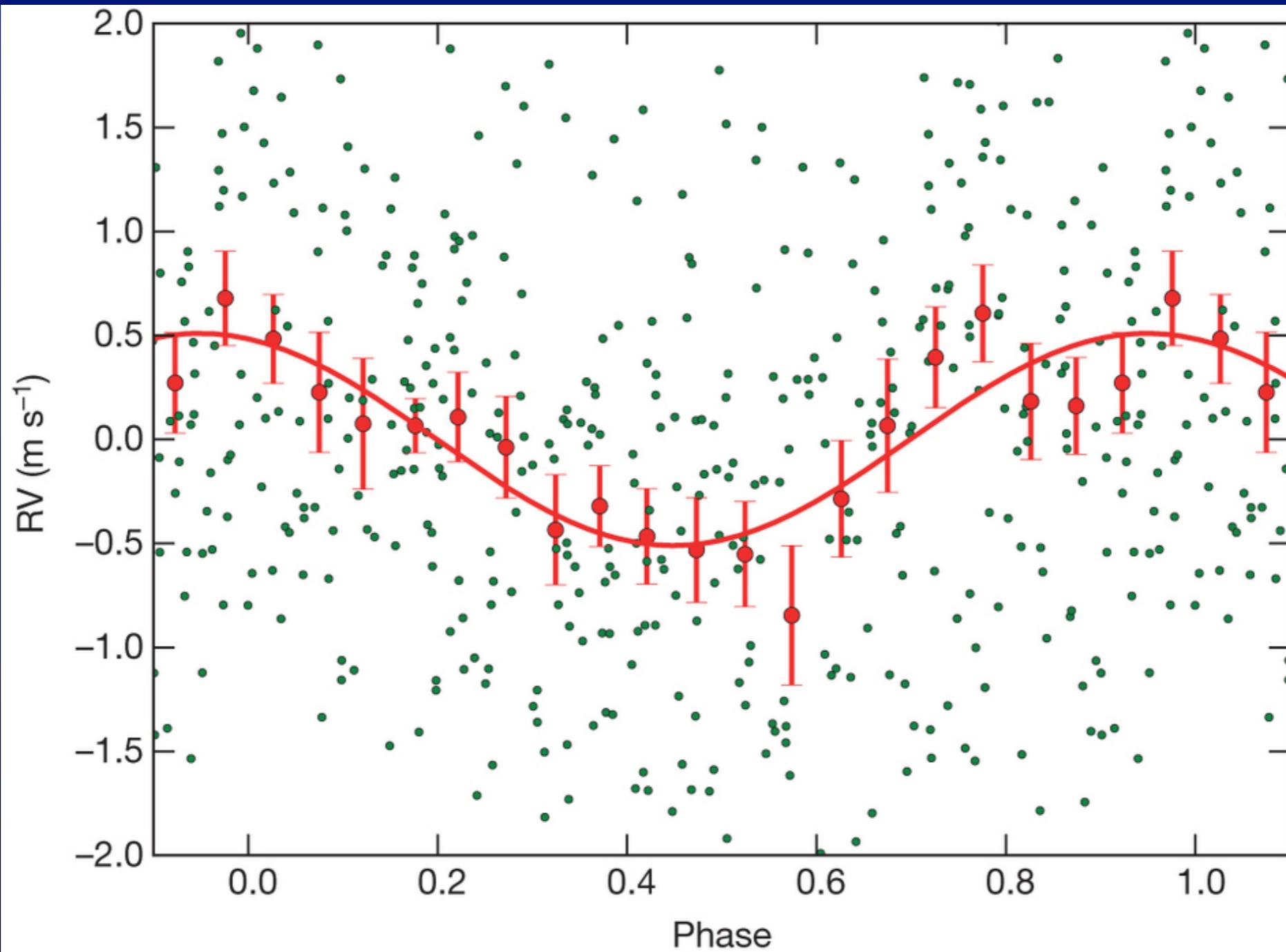


Typical RV data



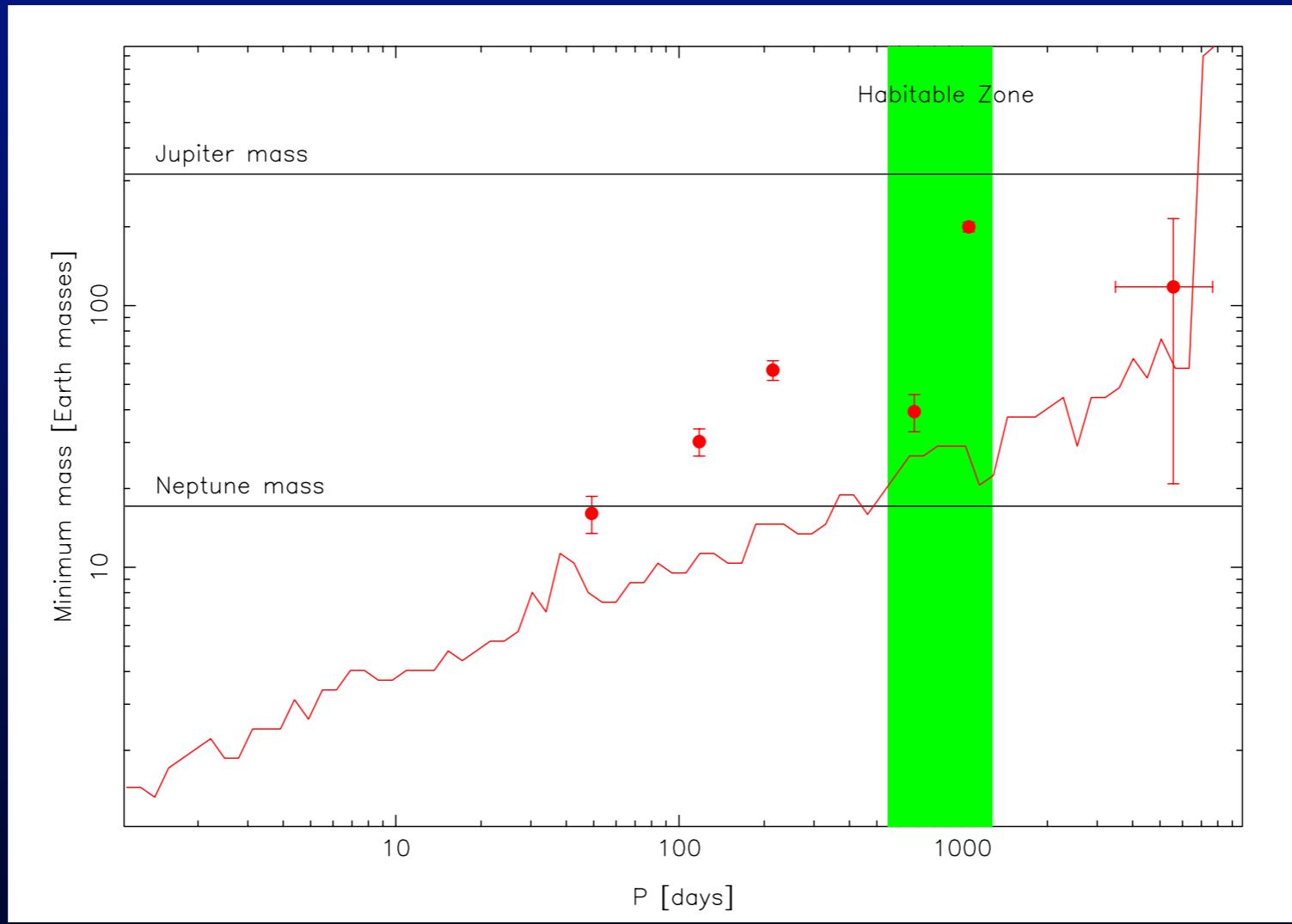
Data from Lovis et al. (2006); figure from Udry & Santos (2007)

The cutting edge??



a Cen Bb: Dumusque et al. (2012)
Claimed planet mass $1.1M_{\text{Earth}}$, $P=3.24\text{d}$, $K=51\text{cm/s}$
But actually an artefact! (see Rajpaul et al. 2016)

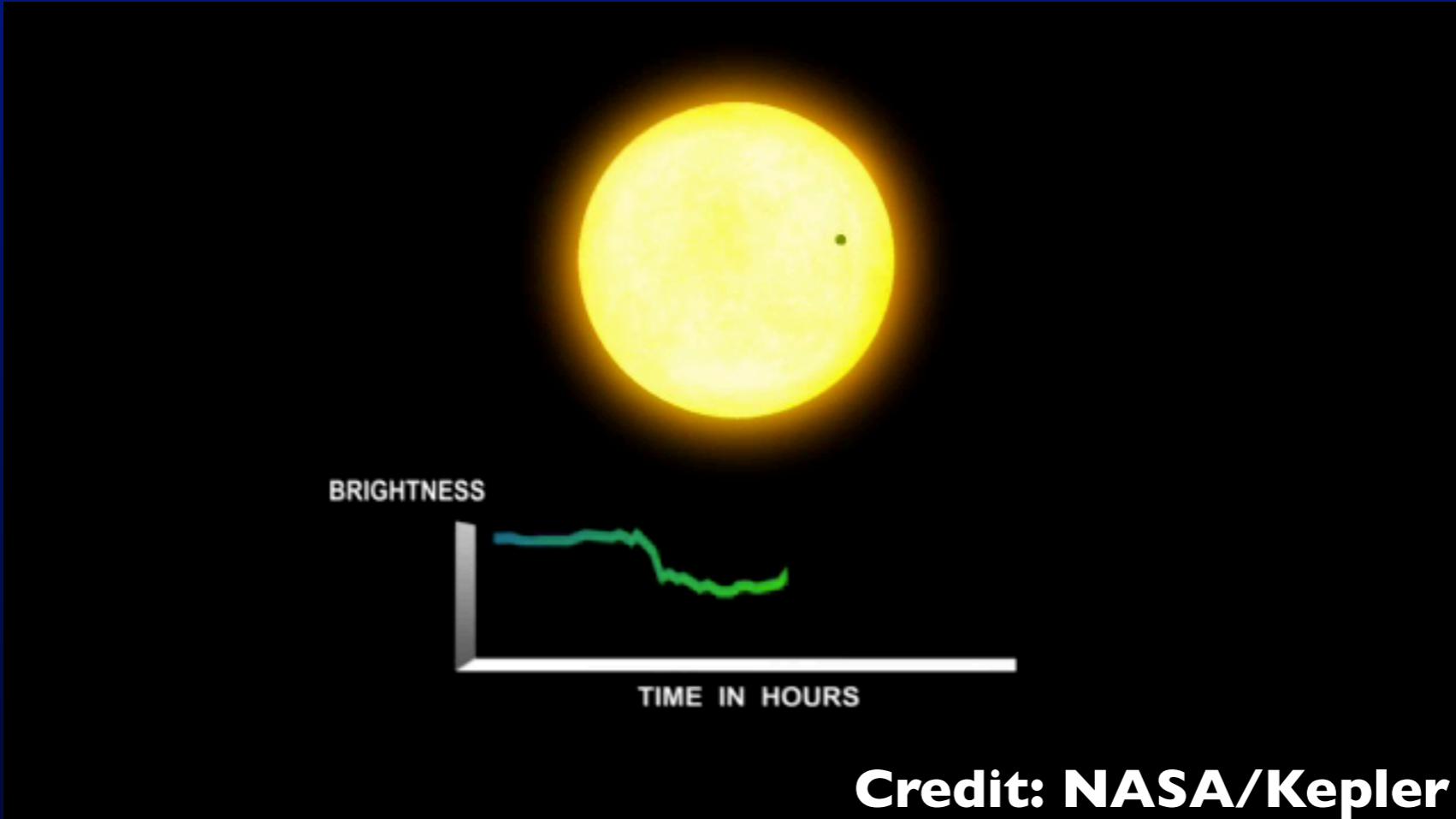
Long surveys, long periods...



Vogt+ (2017)

- 6-planet RV system around HD34445.
- 18 years of RV data; 333 Keck/HIRES spectra; \sim 1–2m/s precision.
- Periods range from 50–5700d; masses from 0.1–0.1 M_{Jup} ; semi-major axes from 0.26–6.4AU.

Transit method

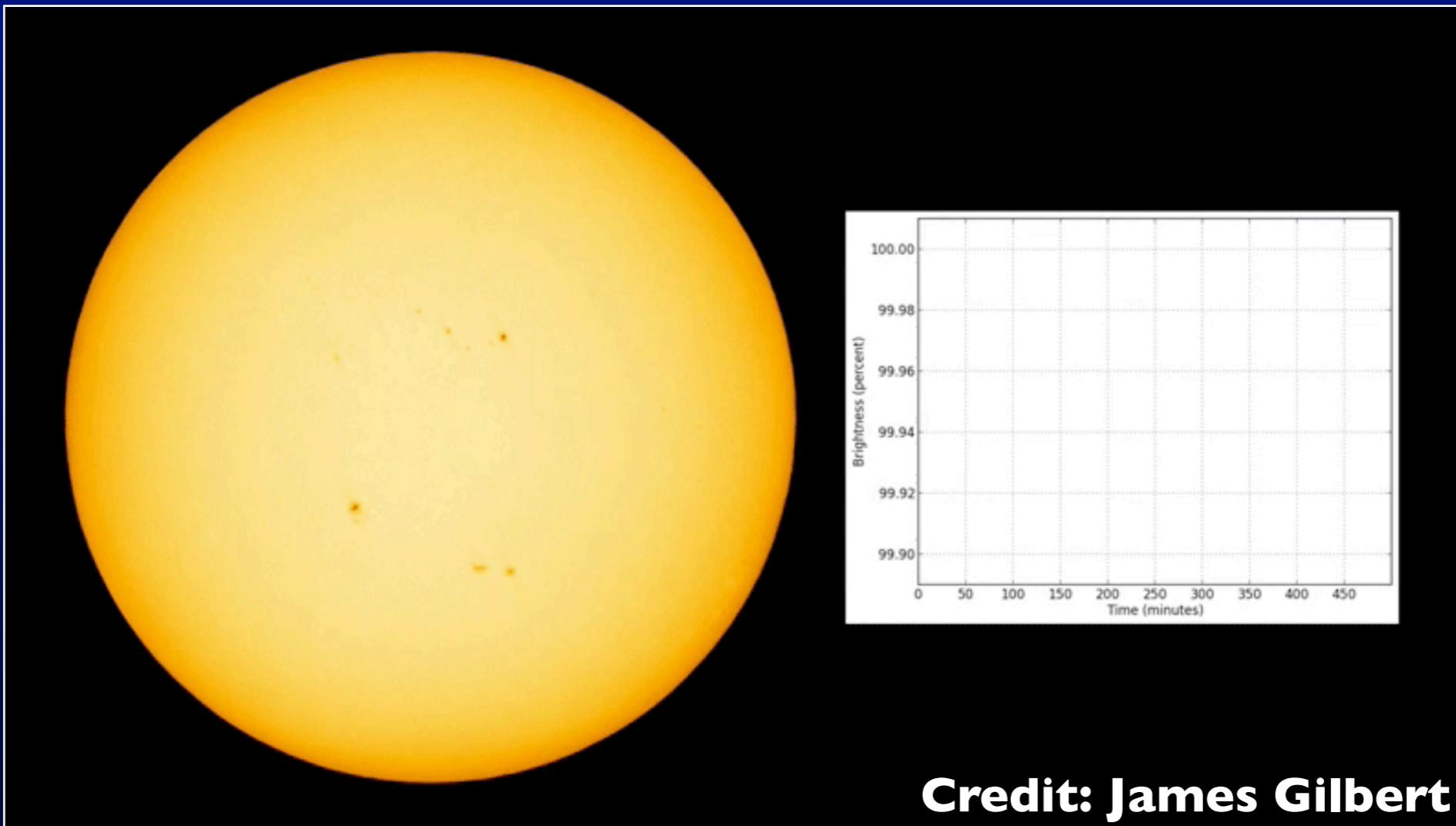


- Detect dimming of light as planet passes in front of star.
- Dimming fraction f depends on planet size:

$$f = \frac{\pi R_p^2}{\pi R_*^2}$$

$$f_{Jup} \simeq 0.01$$
$$f_{\oplus} \simeq 1 \times 10^{-4}$$

Transit method



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Transit method

- Detecting transits requires high precision:
 - <1% precision (~Jupiters) attainable from the ground.
 - 0.01% precision (~Earths) requires us to go to space.
- Detecting transits is very unlikely: requires edge-on orbits:
 - If every star had an Earth-like planet, we would observe transits in approximately 1 in 2000 stars.
- Searching for planets using transits requires us to observe *lots* of stars simultaneously.
- Transit depth tells us the planet's radius. Require follow-up RV measurements to determine mass and eccentricity.

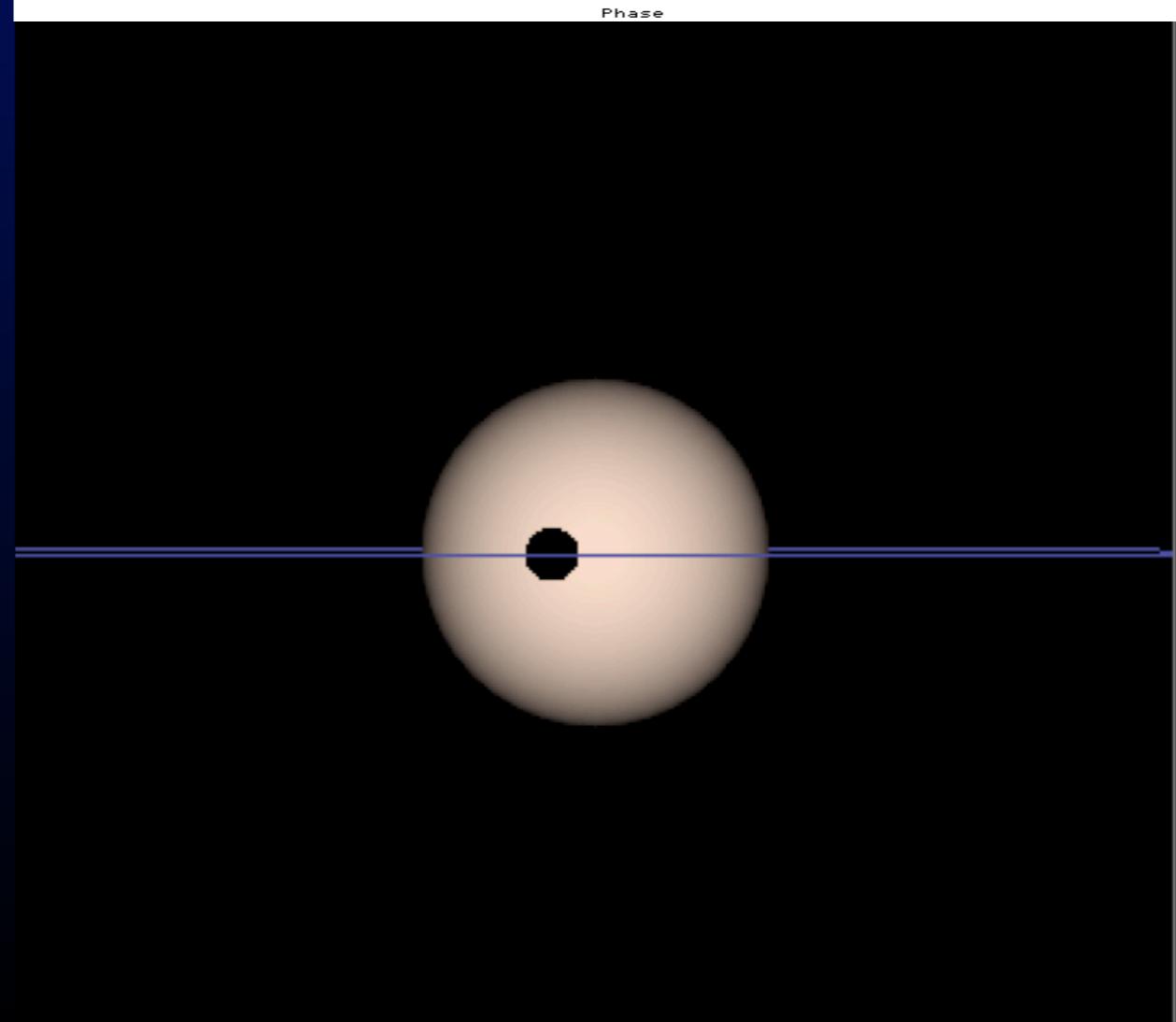
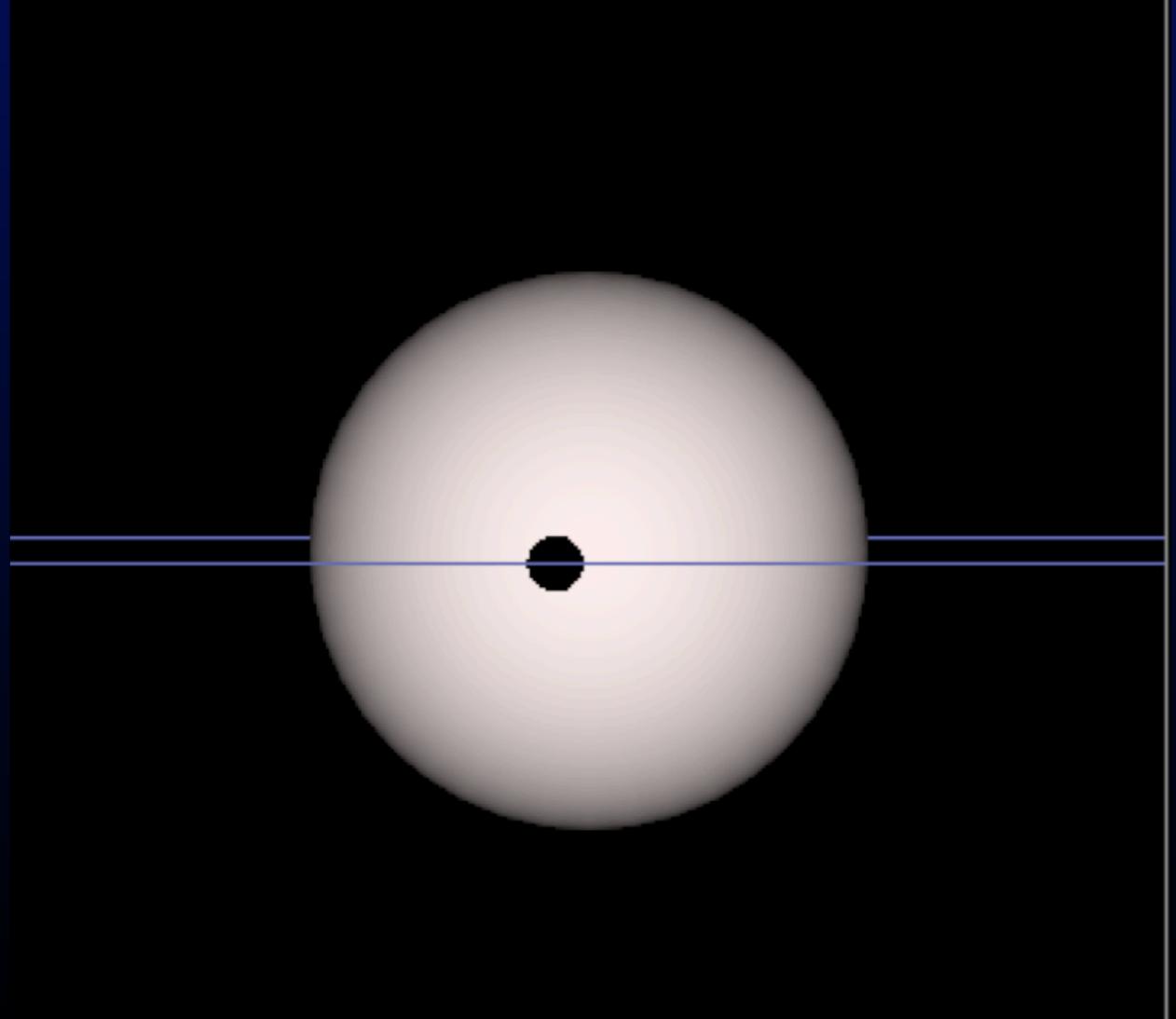
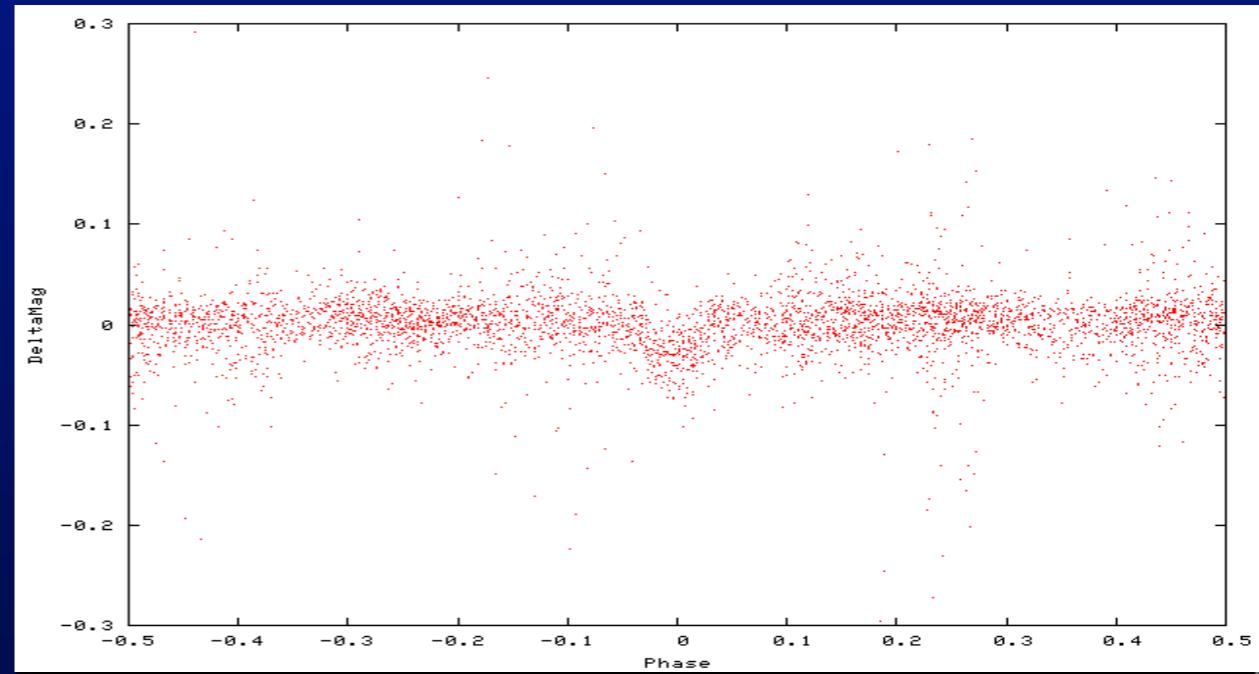
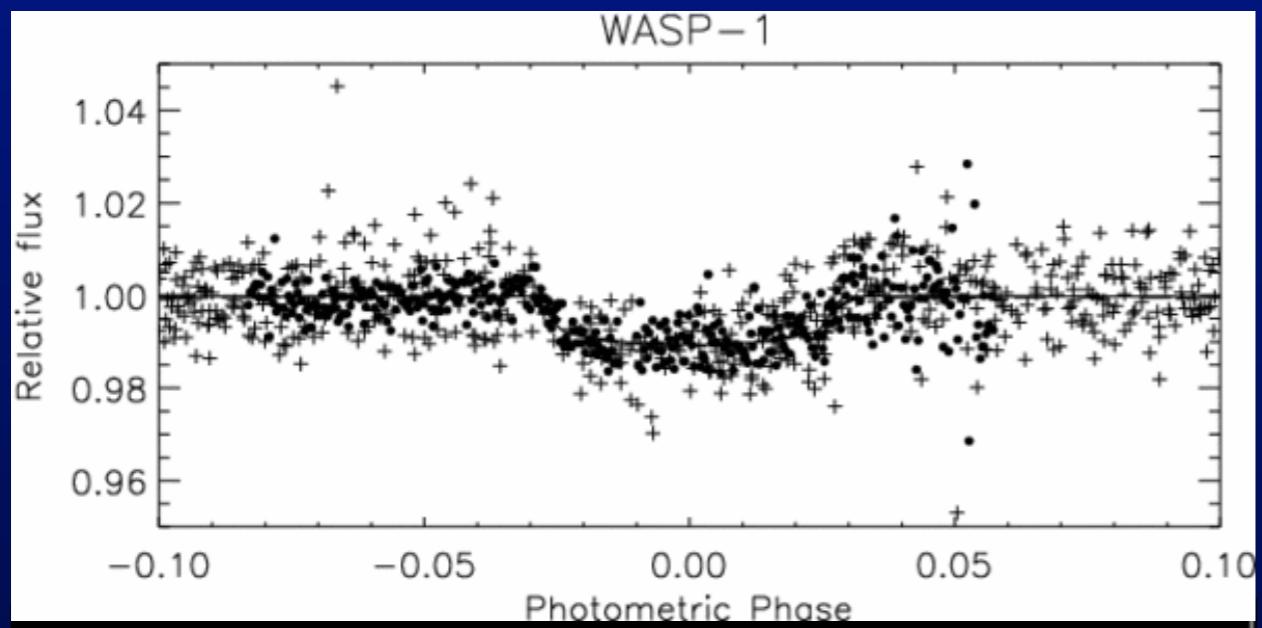
Transit method

- Many current searches using transit methods.
- Most successful ground-based programme is SuperWASP (**Wide-Angle Search for Planets**).
- SuperWASP surveyed 1/4 of the sky every night. Monitored several million stars every few minutes.
- Generates 50-100Gb of data per night.



Credit: Richard West

Ground-based transit lightcurves



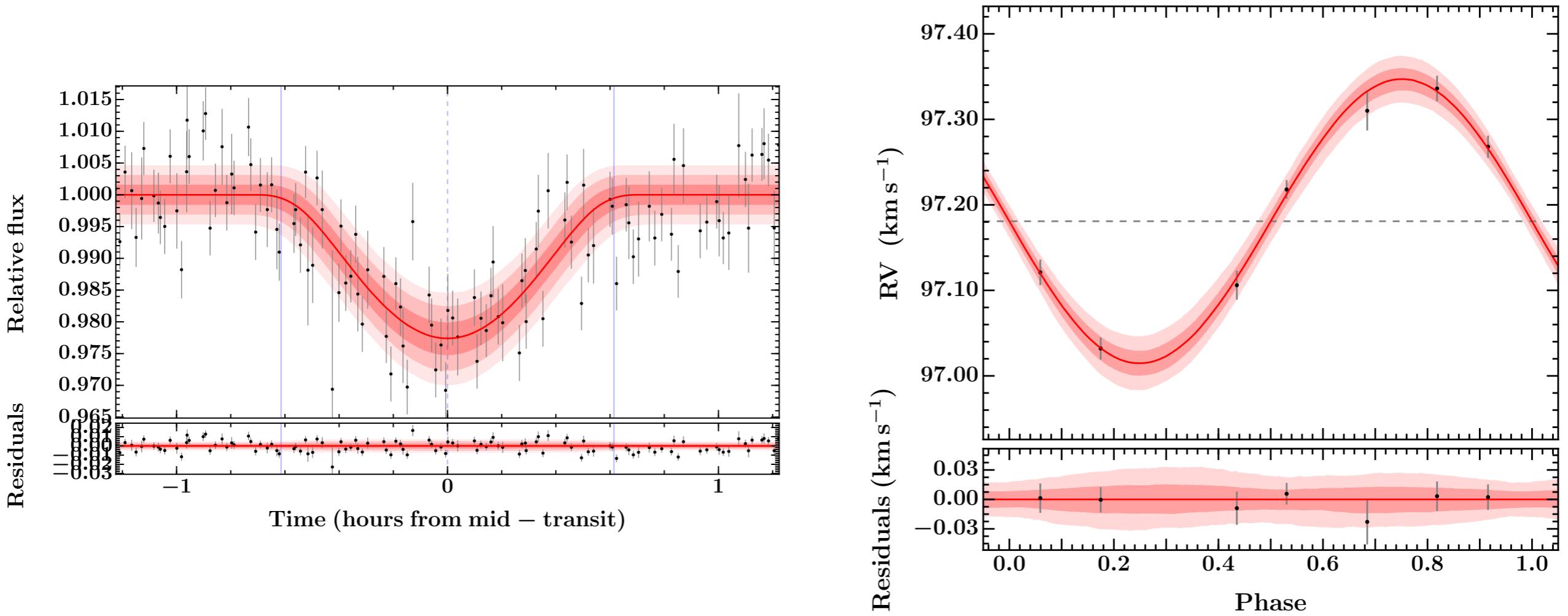
Next Generation...



Credit: ESO/Richard West

- Next Generation Transit Survey (NGTS) now operating at Paranal (first light Jan 2015).
- mmag precision; is yielding a large sample of super-Earths suitable for follow-up from the ground.

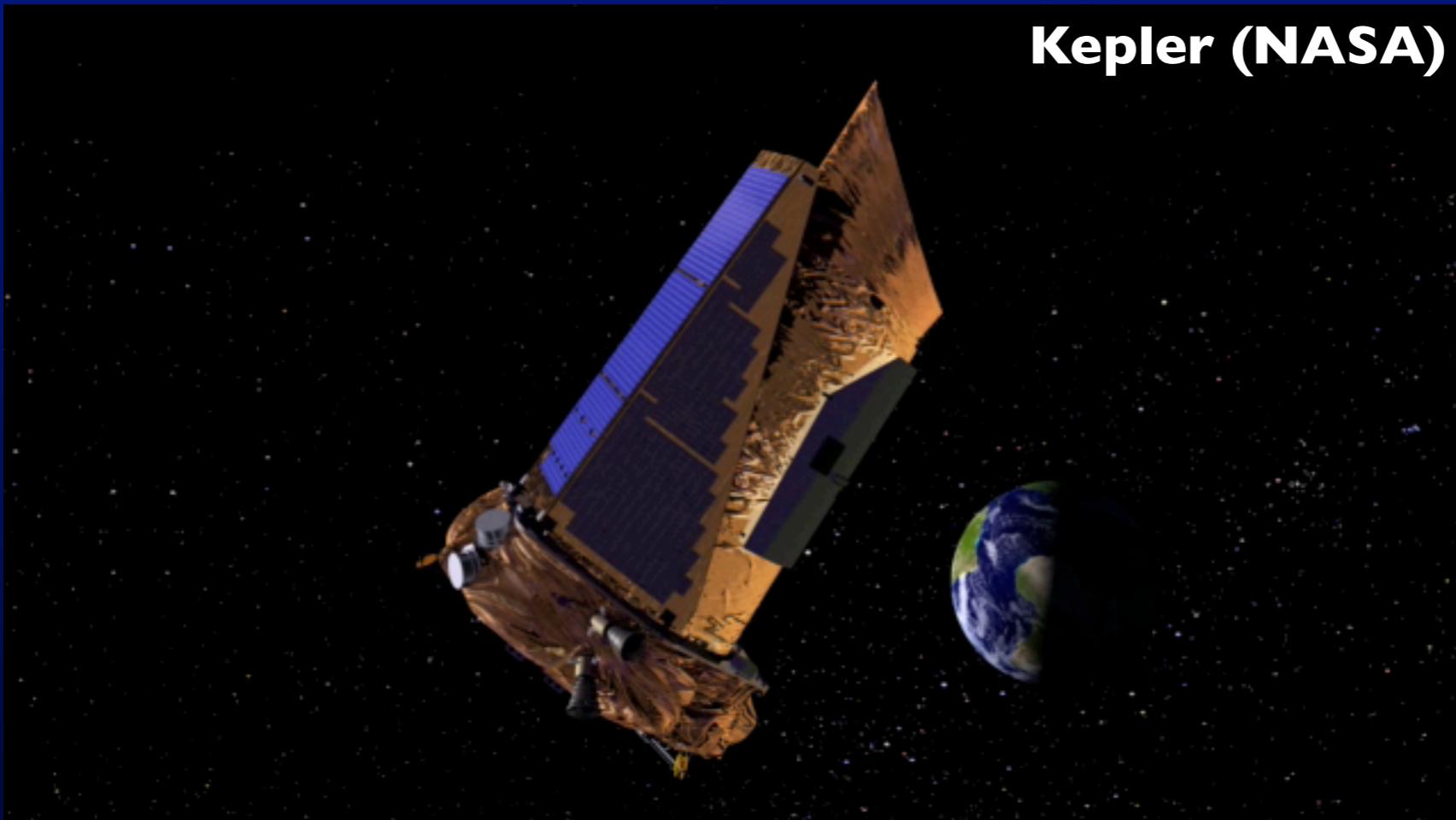
Ground-based cutting edge



Bayliss+ (2017)

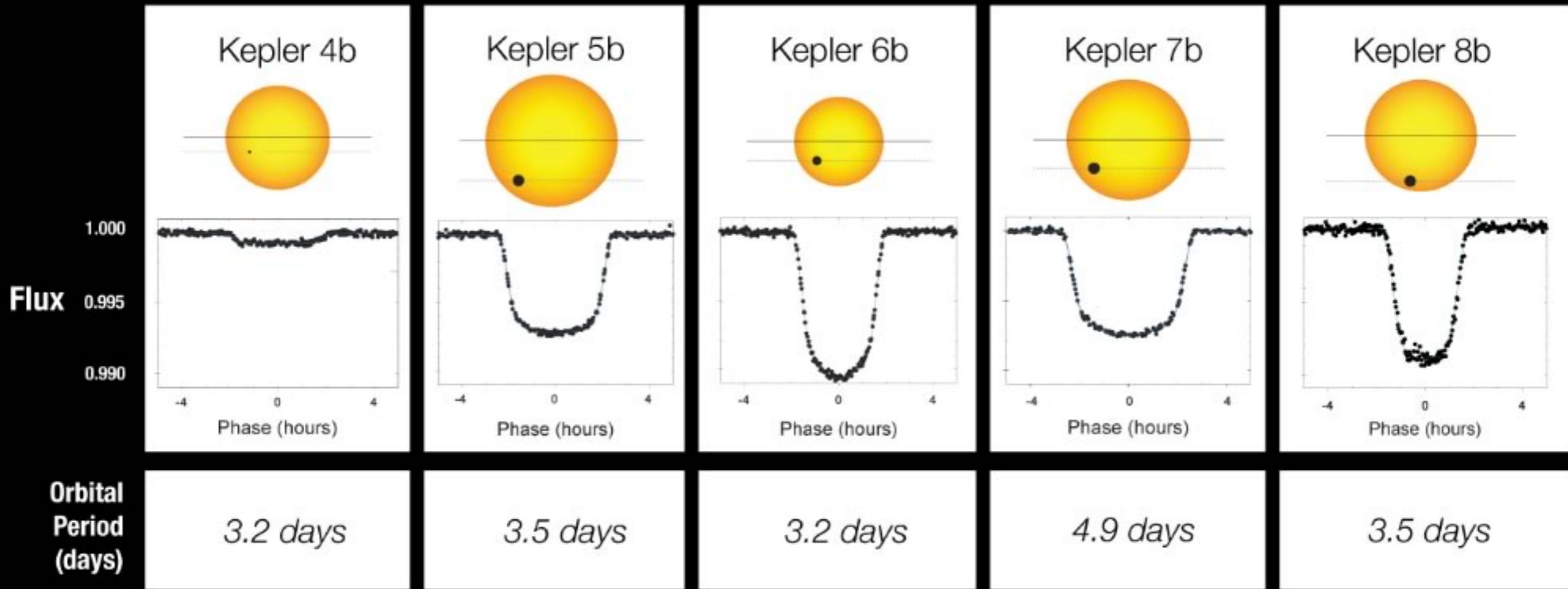
- First exoplanet discovery from NGTS.
- $0.8M_{Jup}$ planet in 2.65d orbit around a M0/M1-type host star.
- Most massive planet known around an M-dwarf. NGTS will give first large census of planets around low-mass stars.

Kepler



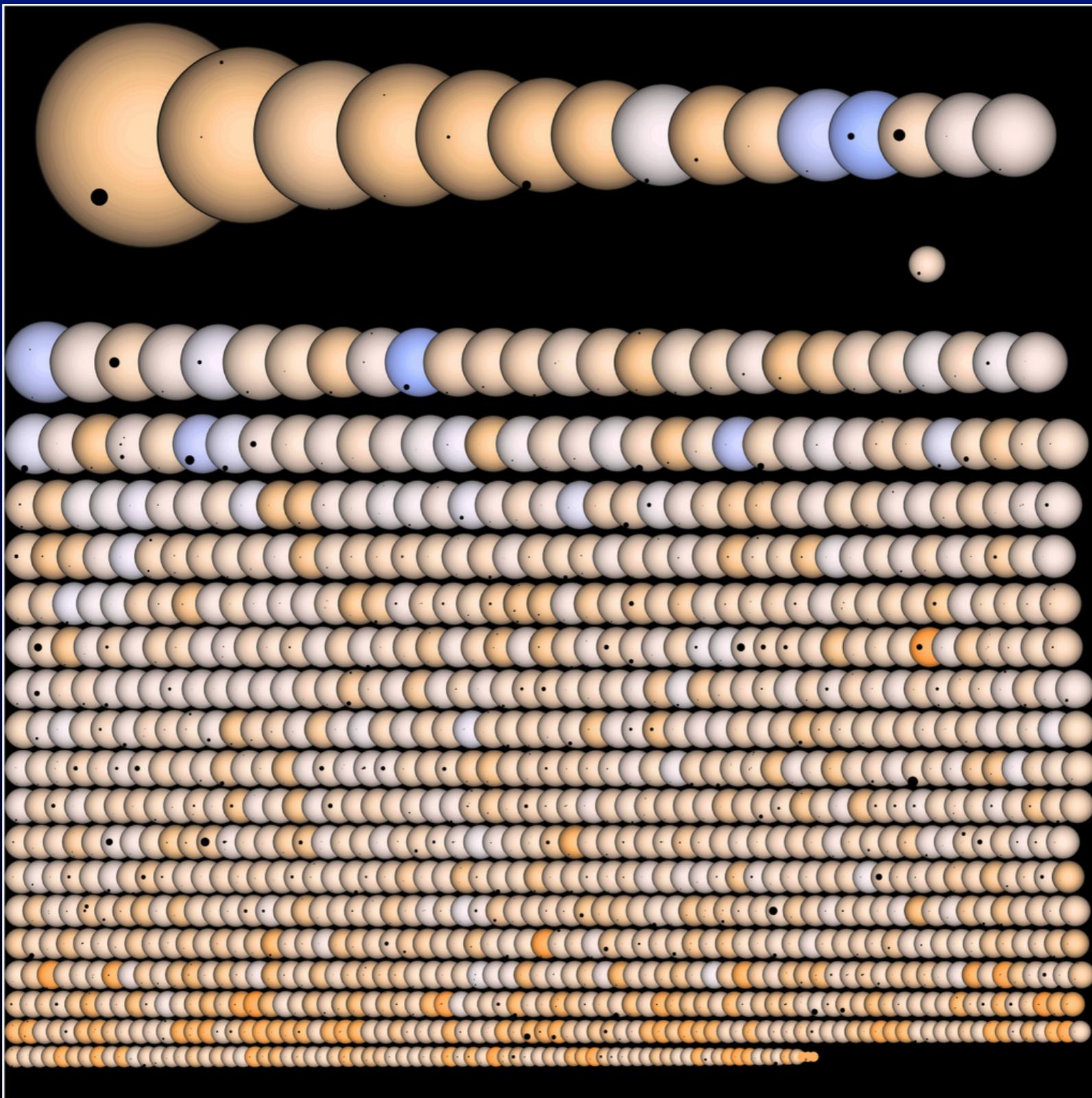
- Launched March 2009, 0.95m primary; “died” May 2013. Lived on as K2 until late 2018.
- 12° FOV, 42 CCD camera. “Stared” at fixed patch of (blank) sky to obtain light-curves for >150,000 stars.
- Photometric precision as good as ~10ppm (in some cases). Sensitive to sub-Earth-size planets.

Kepler light-curves

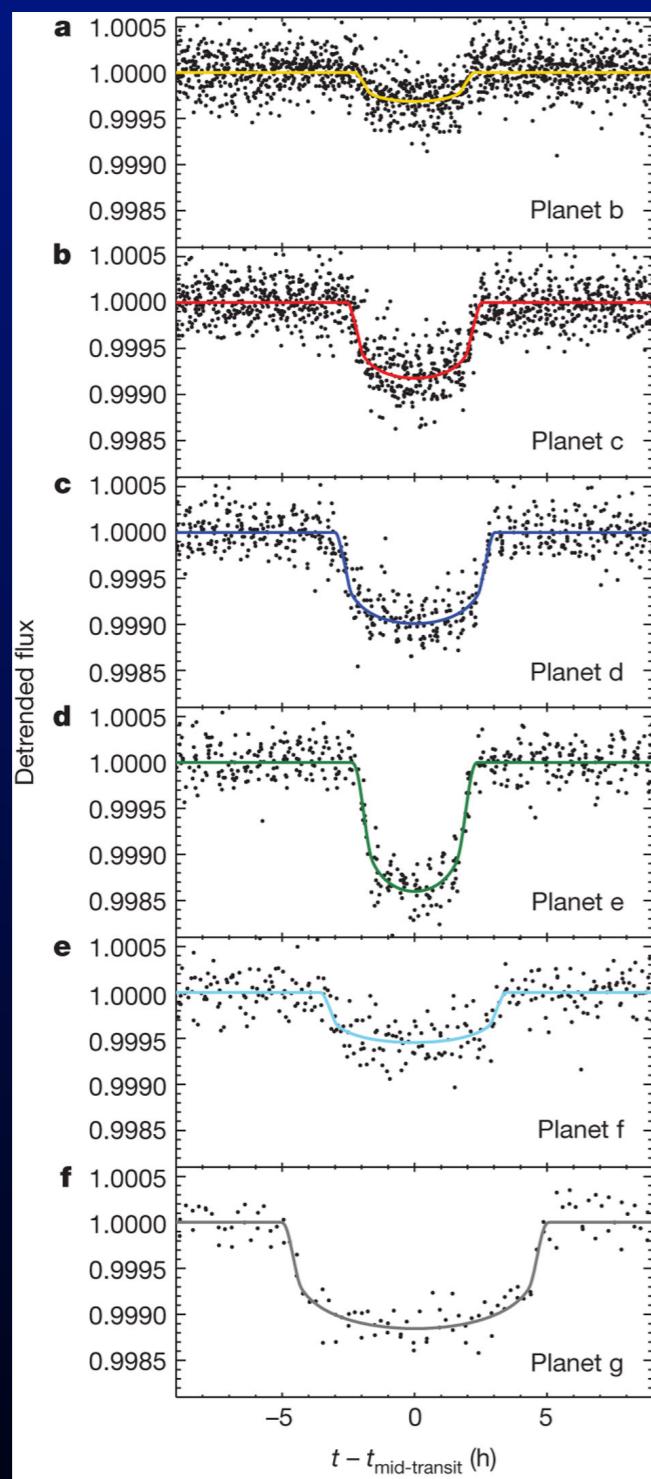


- Early data release (June 2010) focused on a few “hot Jupiters”, to demonstrate precision.
- Fourth (& final) major data release in January 2014. Total of ~4500 planet candidates, with >2000 now confirmed.

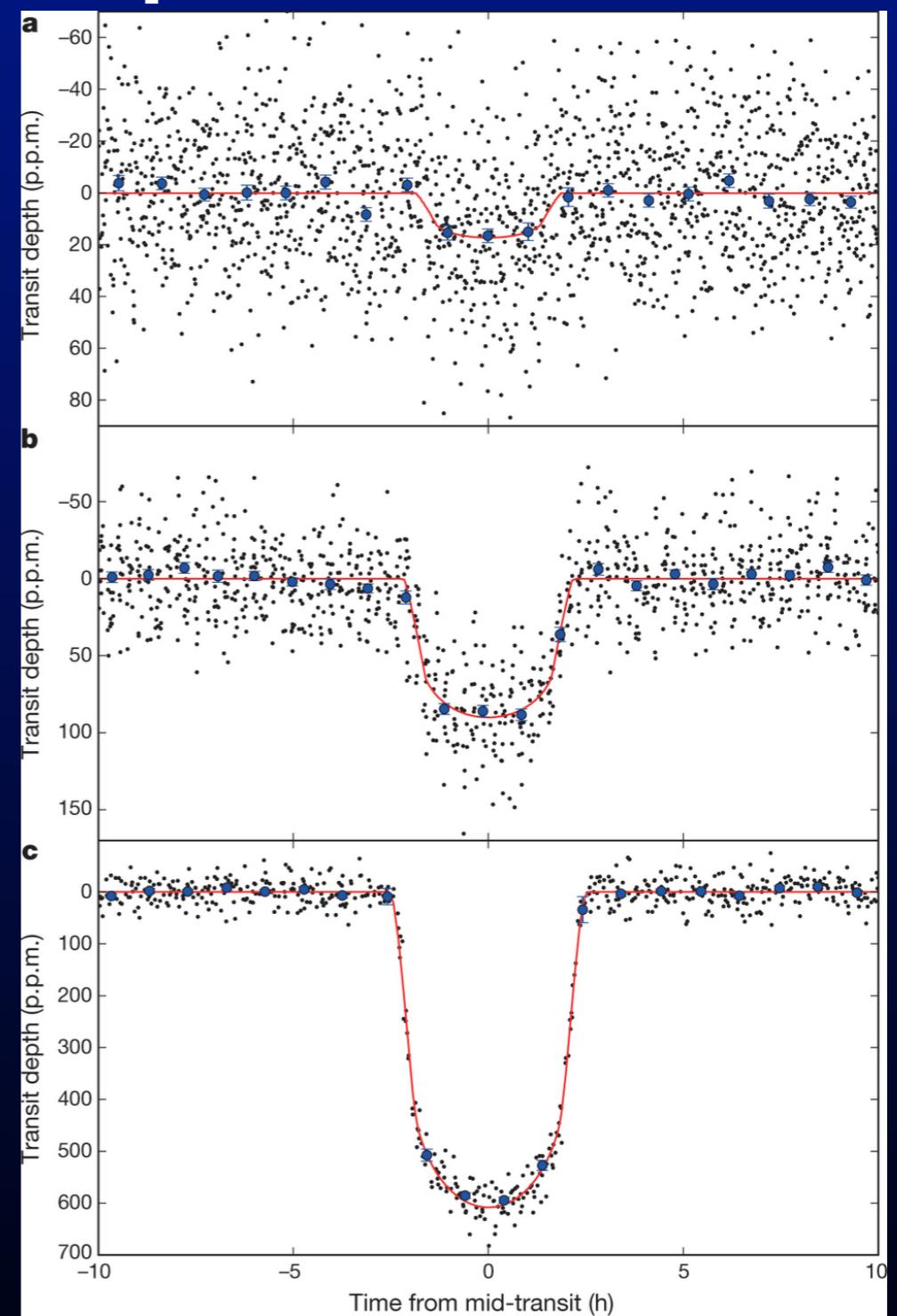
Kepler: first results



Kepler examples



Kepler-11: Lissauer et al. (2011)
6-planet system, periods 10–120d.
Masses range from $2–20M_{\text{Earth}}$.



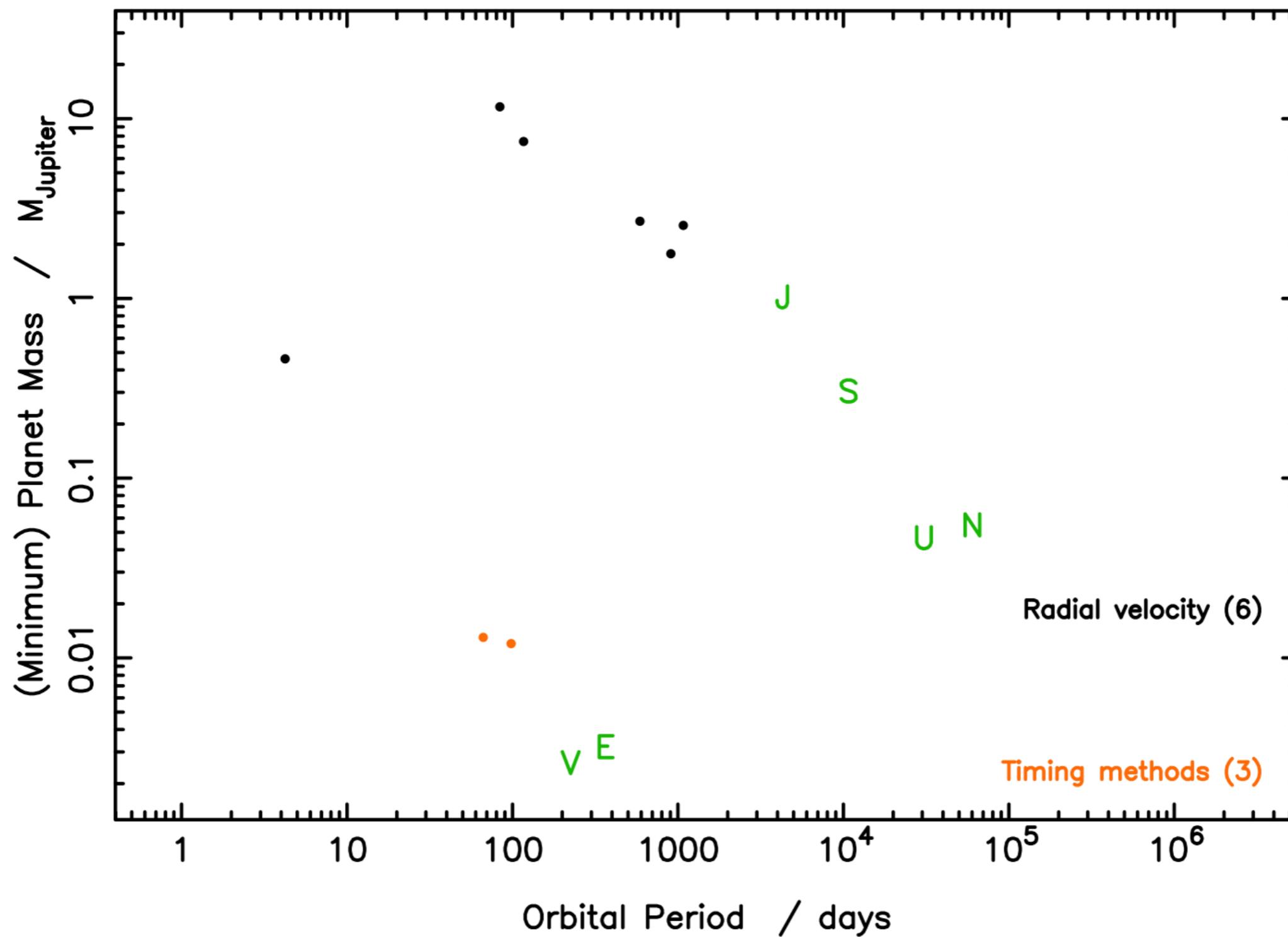
Kepler-37: Barclay et al. (2013)
3-planet system, periods 13.4, 21.3, 39.8d.
“b” is roughly the size of the Moon.

Summary of methods and biases

- First discoveries: 1995 (RV), 2005* (transit), 2008 (imaging).
- Now >3000 known exoplanets (+ ~2500 *Kepler* candidates):
- Direct Imaging
 - Easiest to detect bright (large R_p and/or massive) planets far from star (large a).
- Radial velocity
 - Easiest to detect massive planets close to star (short periods, small a).
- Transits
 - Easiest to detect large (large R_p) planets close to star (short periods, small a).

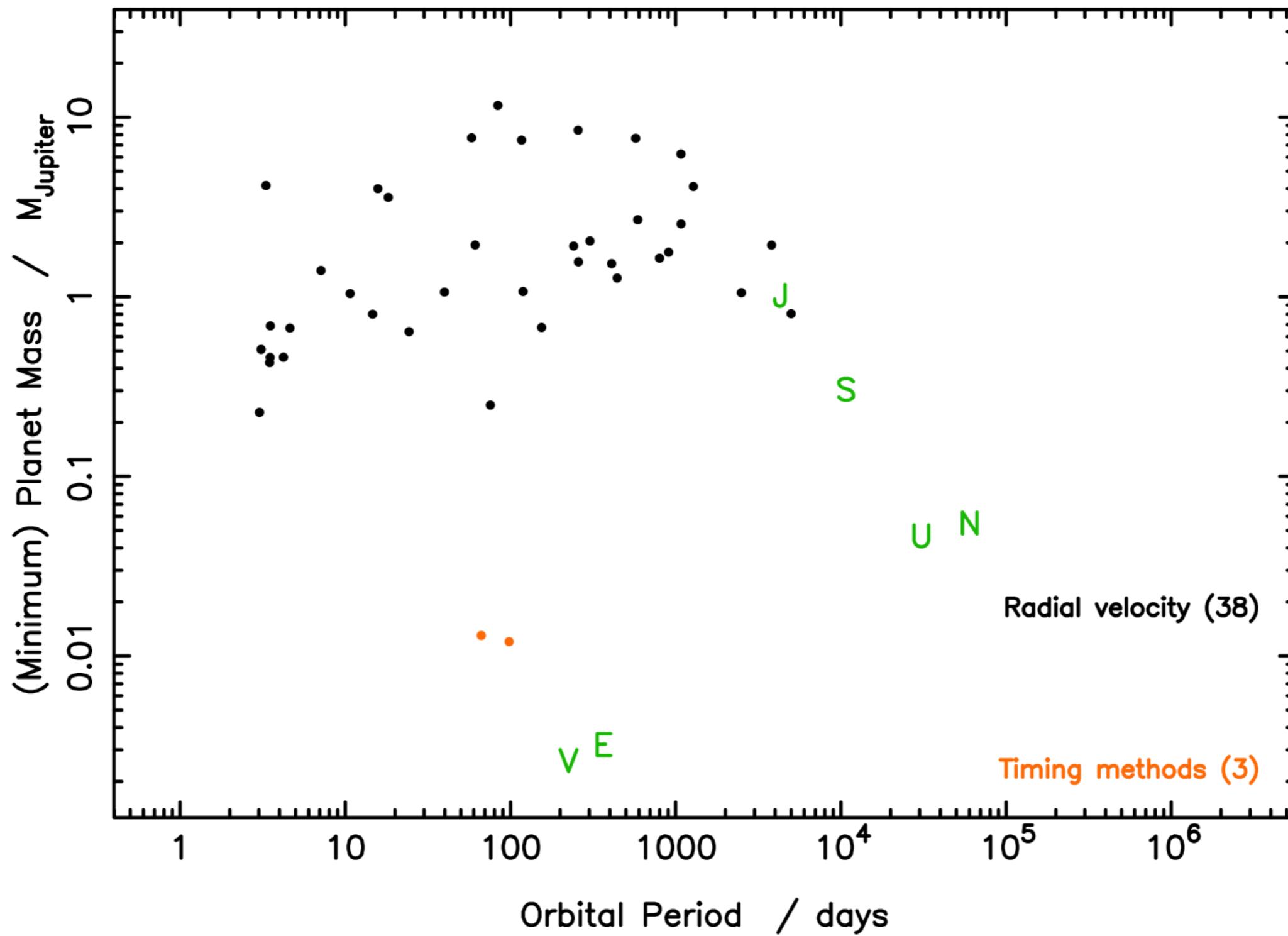
*The first transiting planet was found in 1999, but it was a known RV planet.

Known planets as of 1996



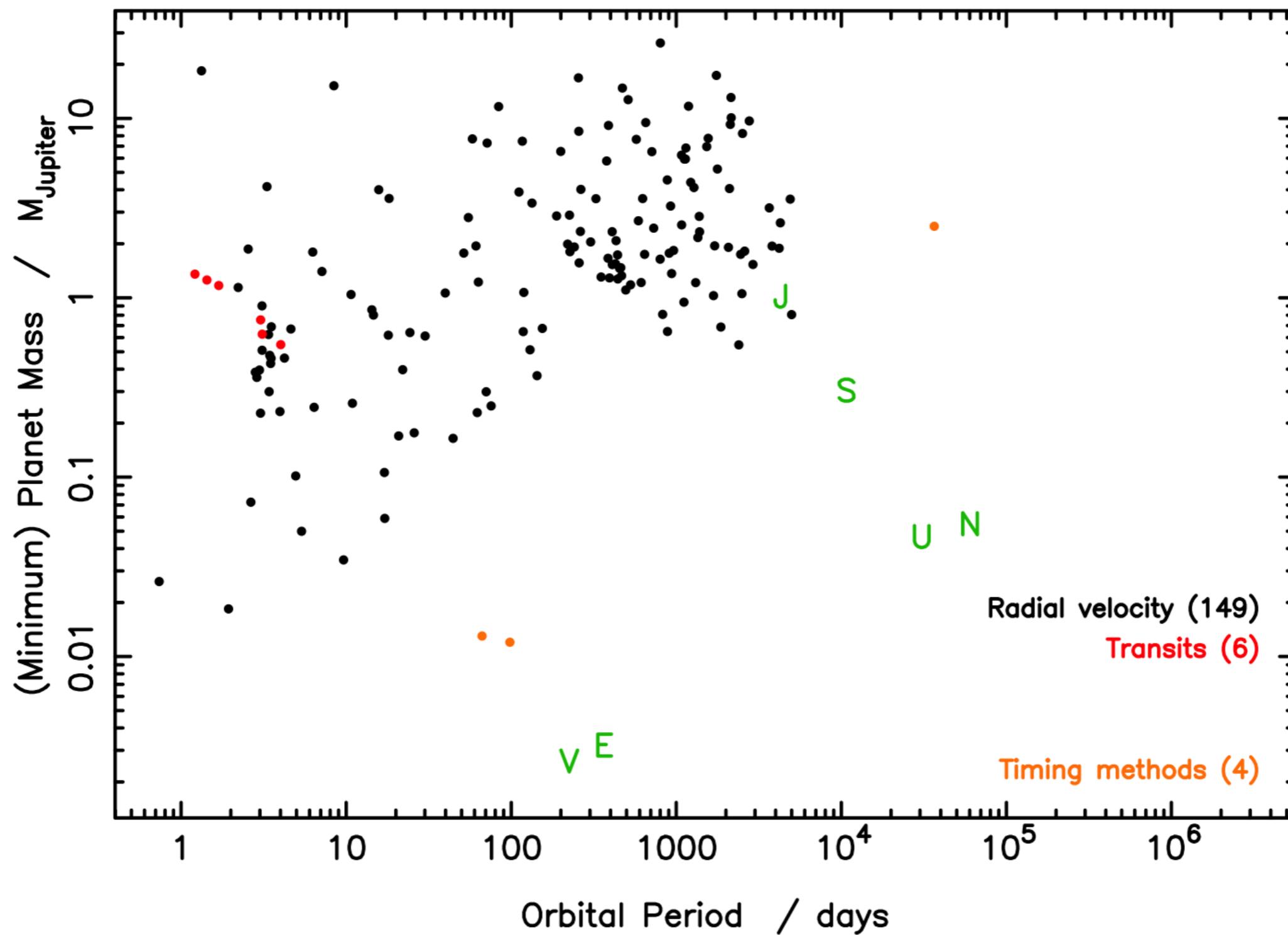
Data from exoplanet.eu

Known planets as of 2000



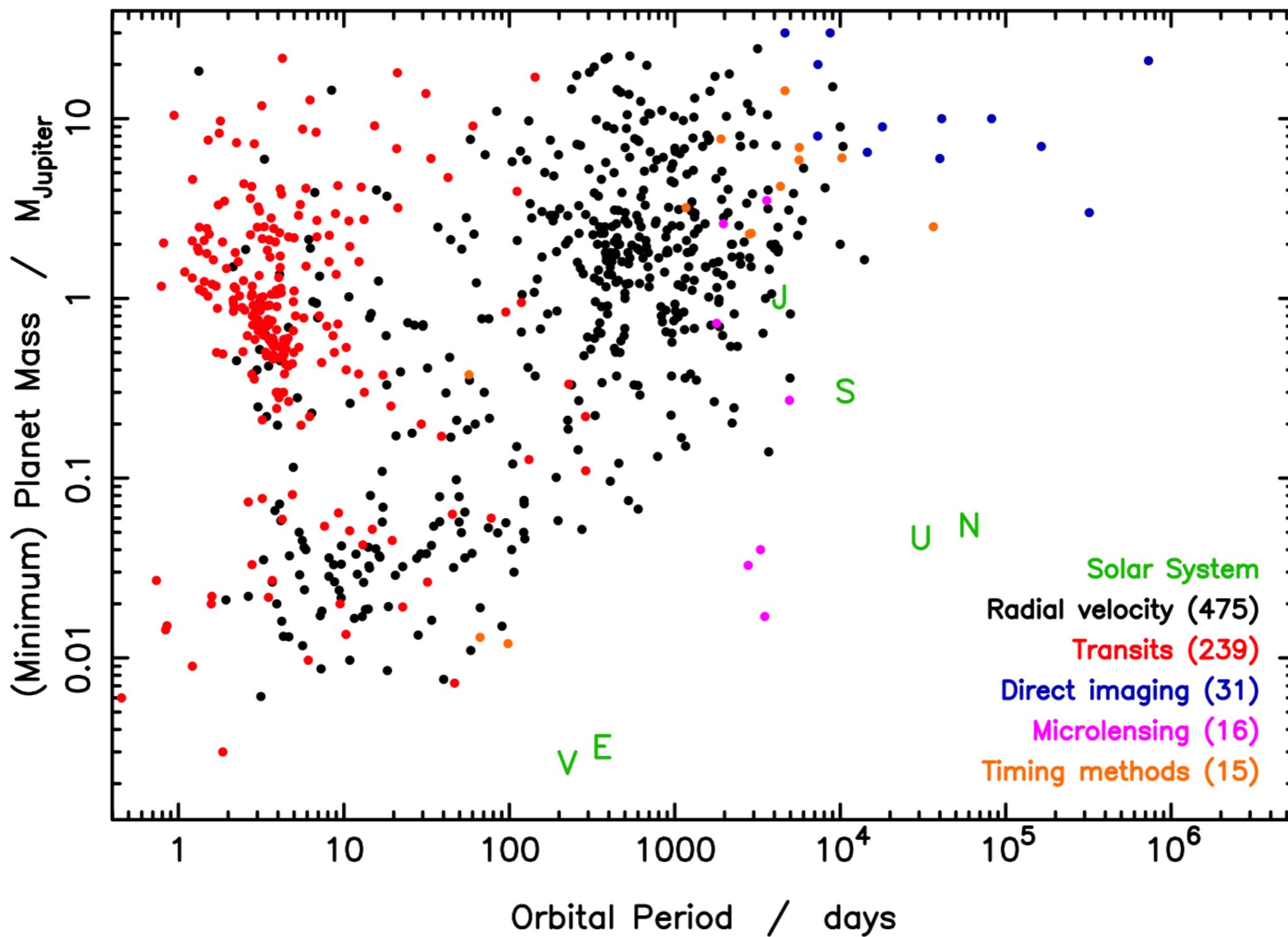
Data from exoplanet.eu

Known planets as of 2005



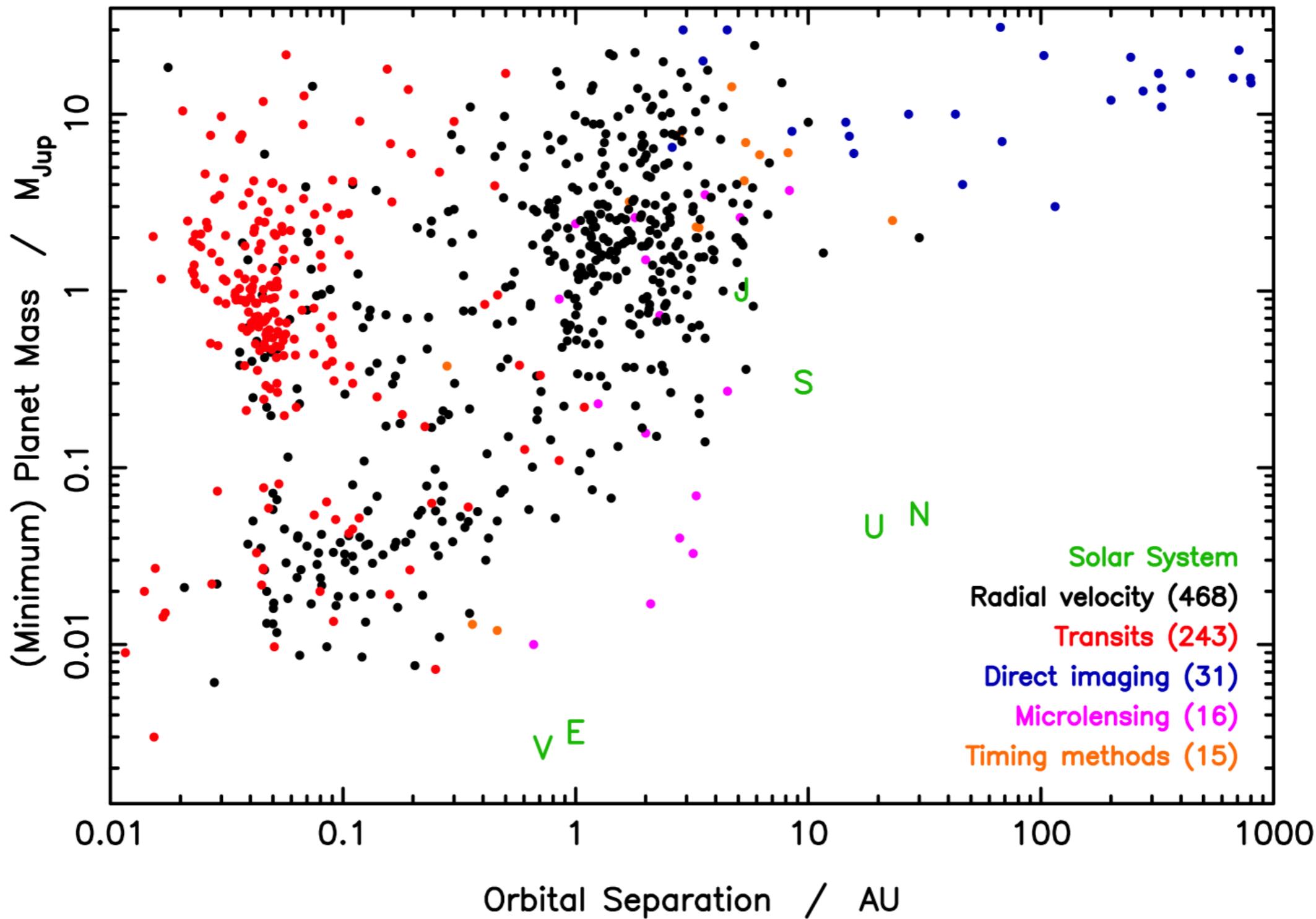
Data from exoplanet.eu

Known planets (as of 1 Oct 2012)



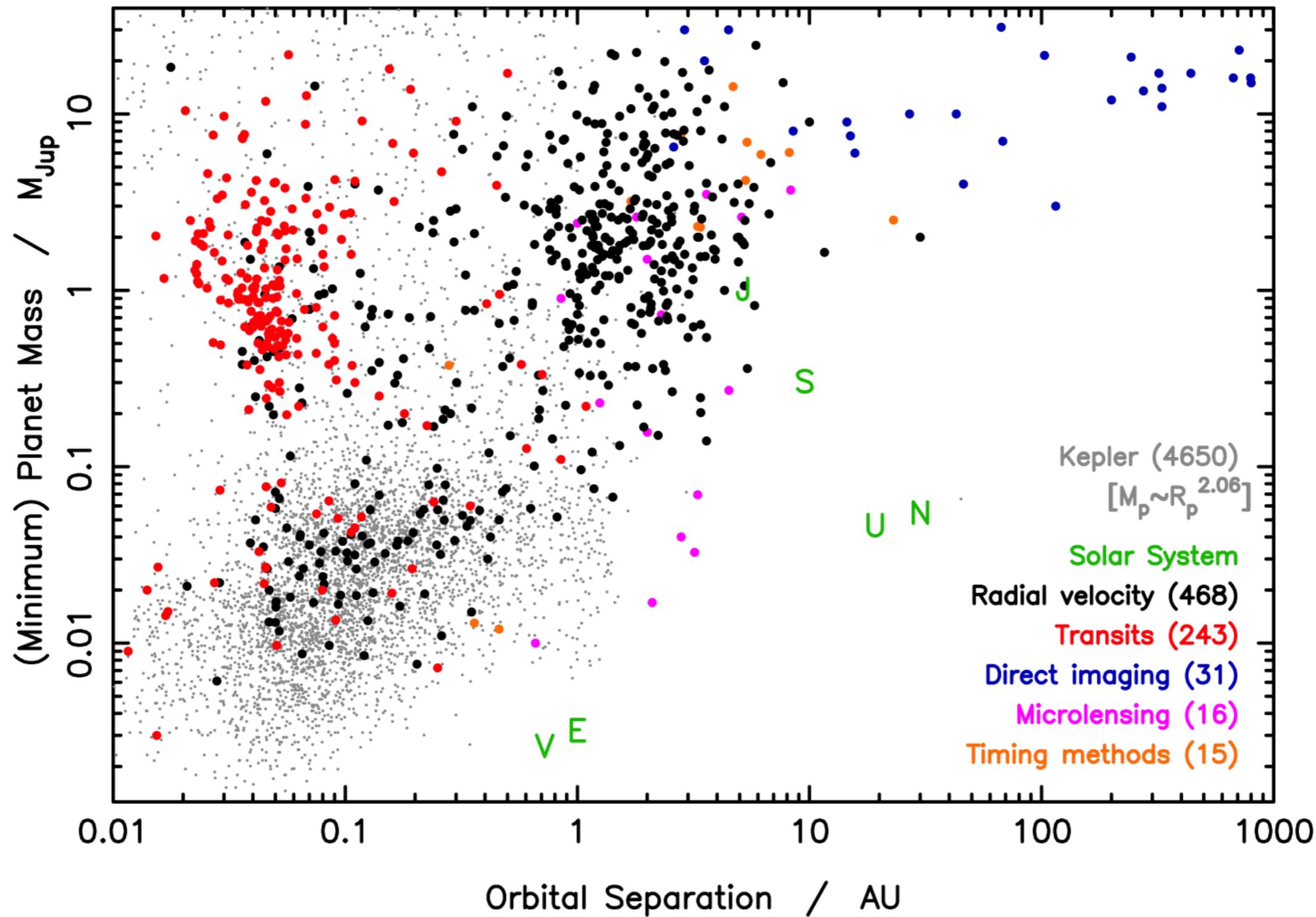
Data from exoplanet.eu

Known planets (as of 7 Oct 2015)



Data from exoplanets.org

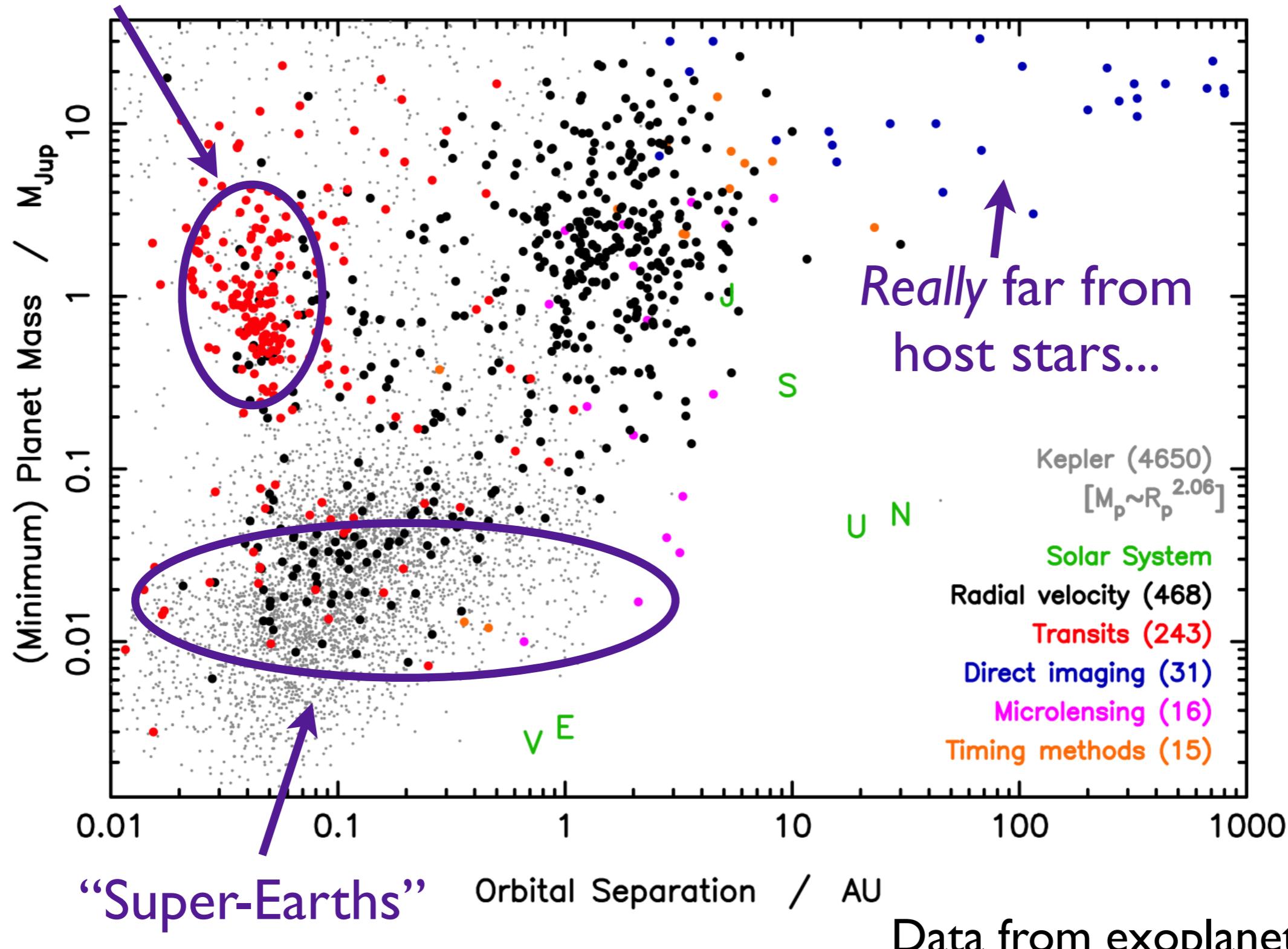
Known planets (as of 7 Oct 2015)



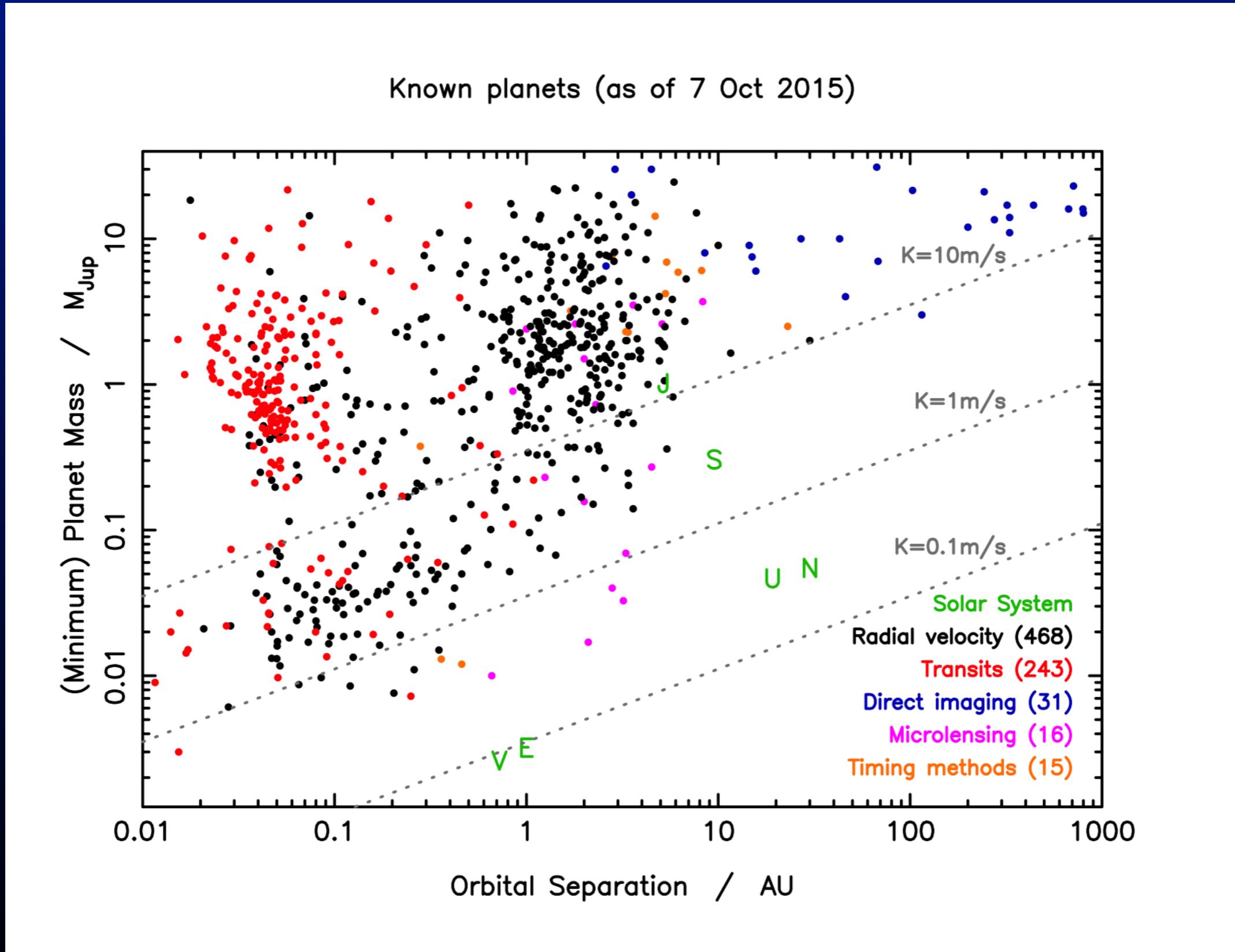
Data from exoplanets.org

“Hot Jupiters”

Known planets (as of 7 Oct 2015)

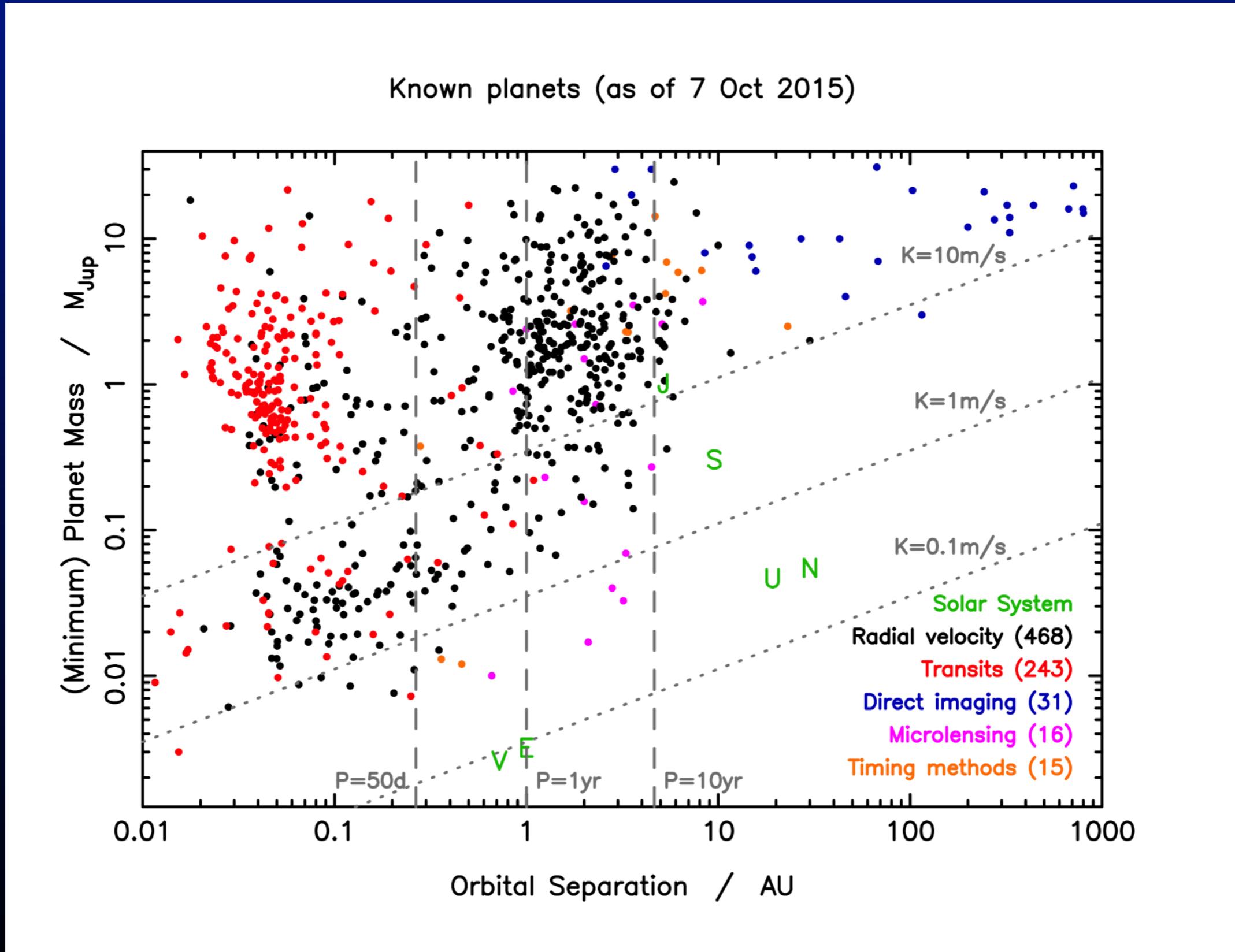


Selection biases

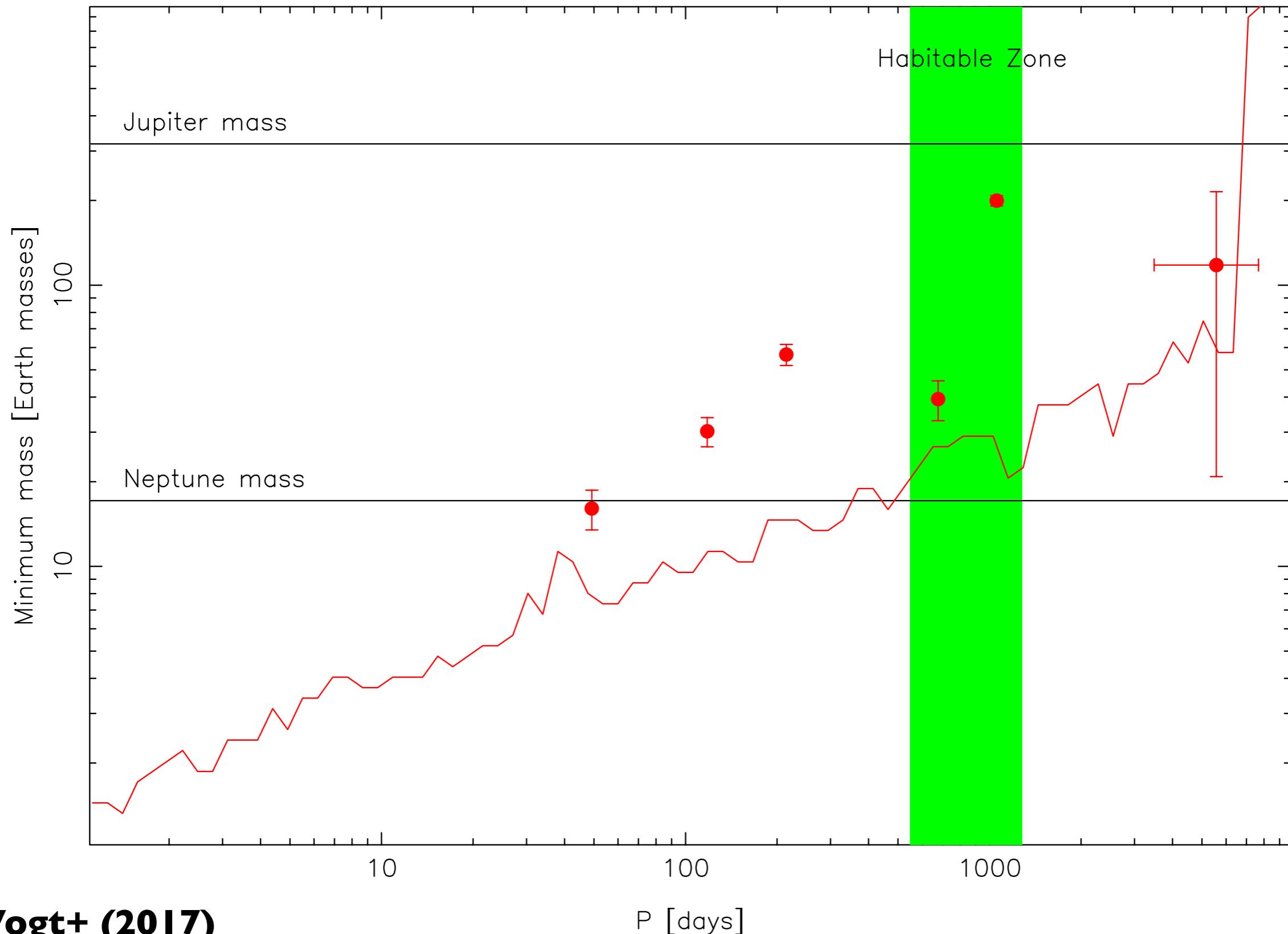


$$K \propto M_p \sin i a^{-1/2} \Rightarrow M_p \sin i \propto K a^{1/2}$$

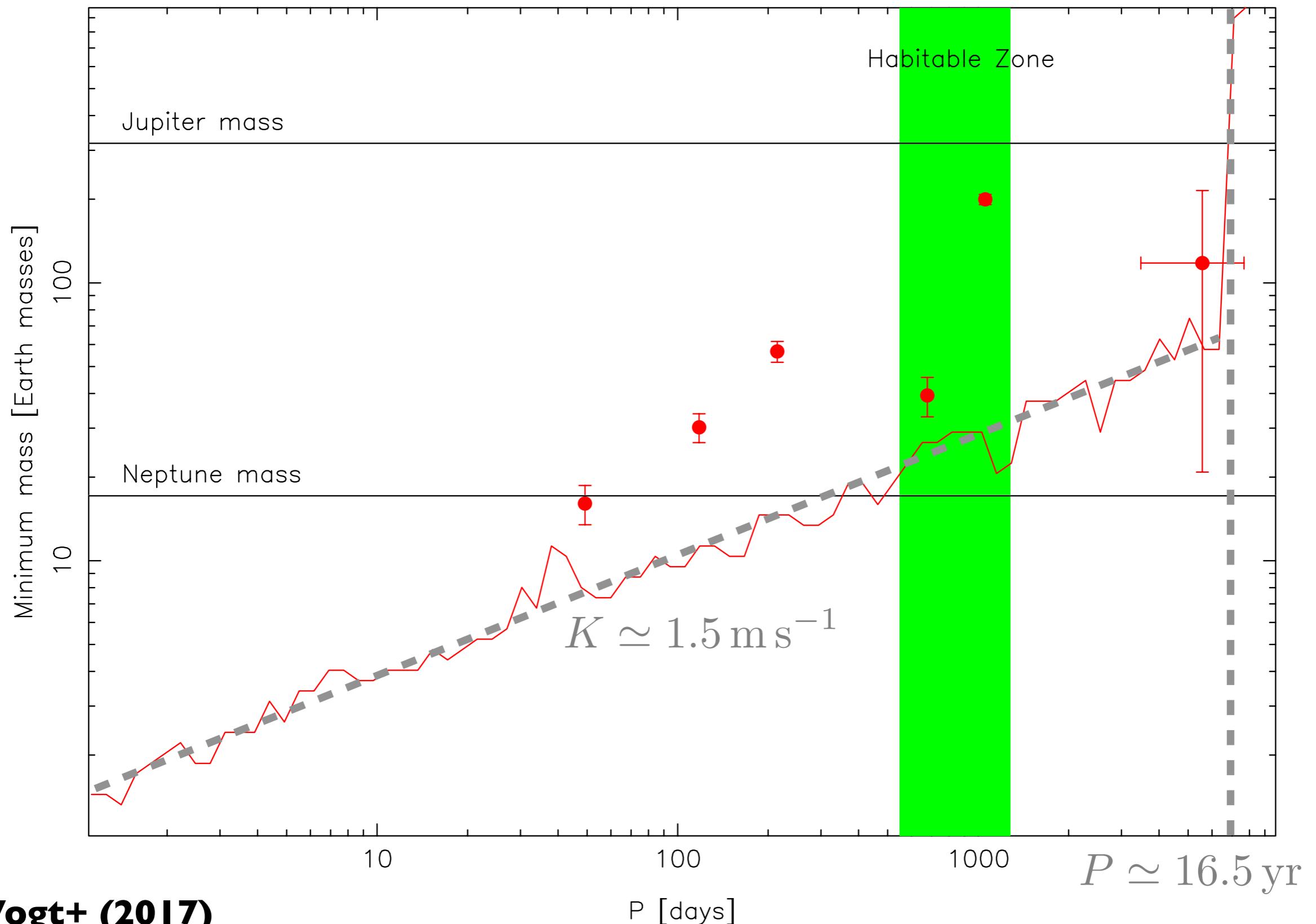
Selection biases



Selection biases



Selection biases



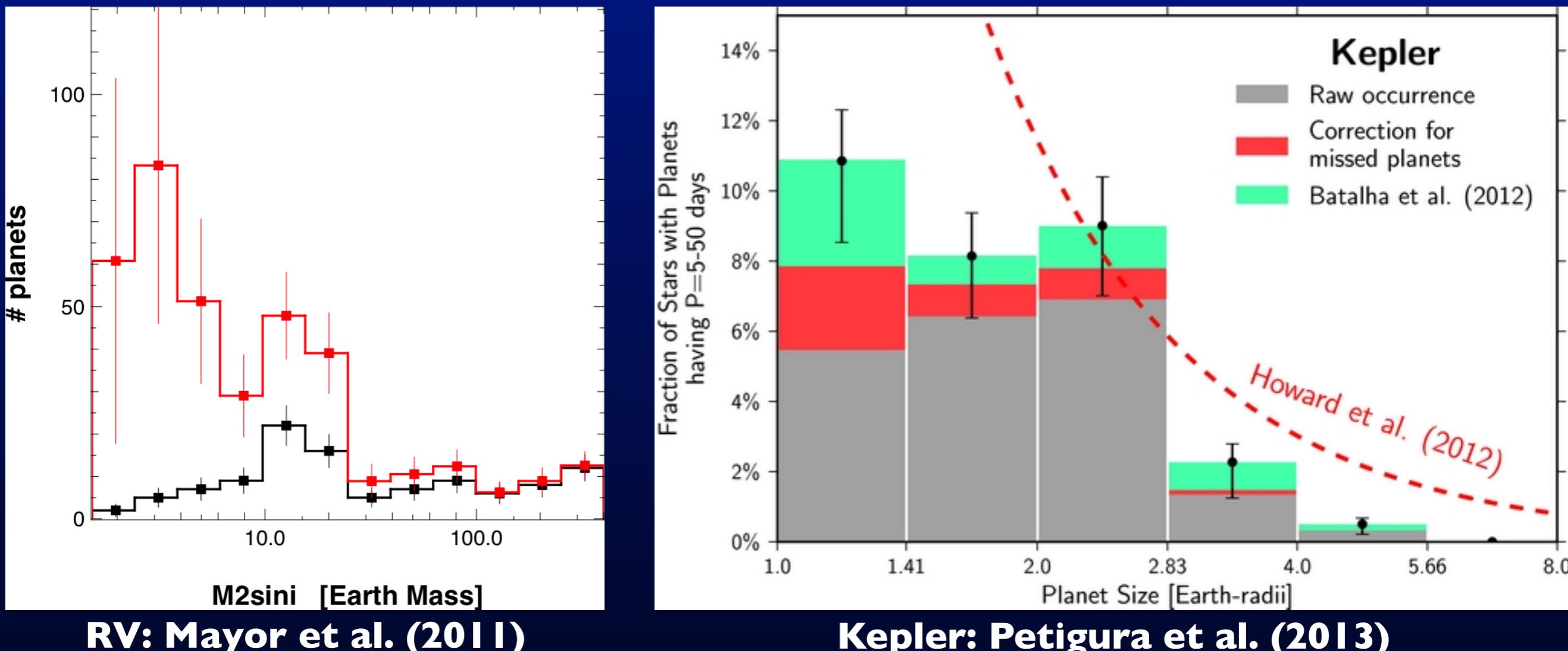
What fraction of stars host planets?

What fraction of stars host planets?

- Selection biases mean that measurements of f_p must be qualified (but detection methods are complementary).
- Current results:
 - 5-10% of FGK stars host a planet with $M_p \geq M_{Jup}$ at $a \leq 3\text{AU}$.
 - $>50\%$ of FGK stars host a planet with $M_p \geq 1M_\oplus$ and $P \leq 100\text{d}$.
 - $\sim 90\%$ of M stars host a planet with $R_p \geq 0.5R_\oplus$ and $P \leq 50\text{d}$.
- Extension of these results to larger radii will take time. Future missions will probe lower masses, but orbital periods at large ($>\text{AU}$) radii are long.
- Can currently say that $f_p \geq 0.5$ for sun-like stars. Seems likely that the true value is very close to 1.

Statistical properties of exoplanets

Planet mass function



- Distribution of planet masses increases to low M_p .
- Apparent “plateau” in mass (size) function below a few times the size of Earth.

Mass-radius relation

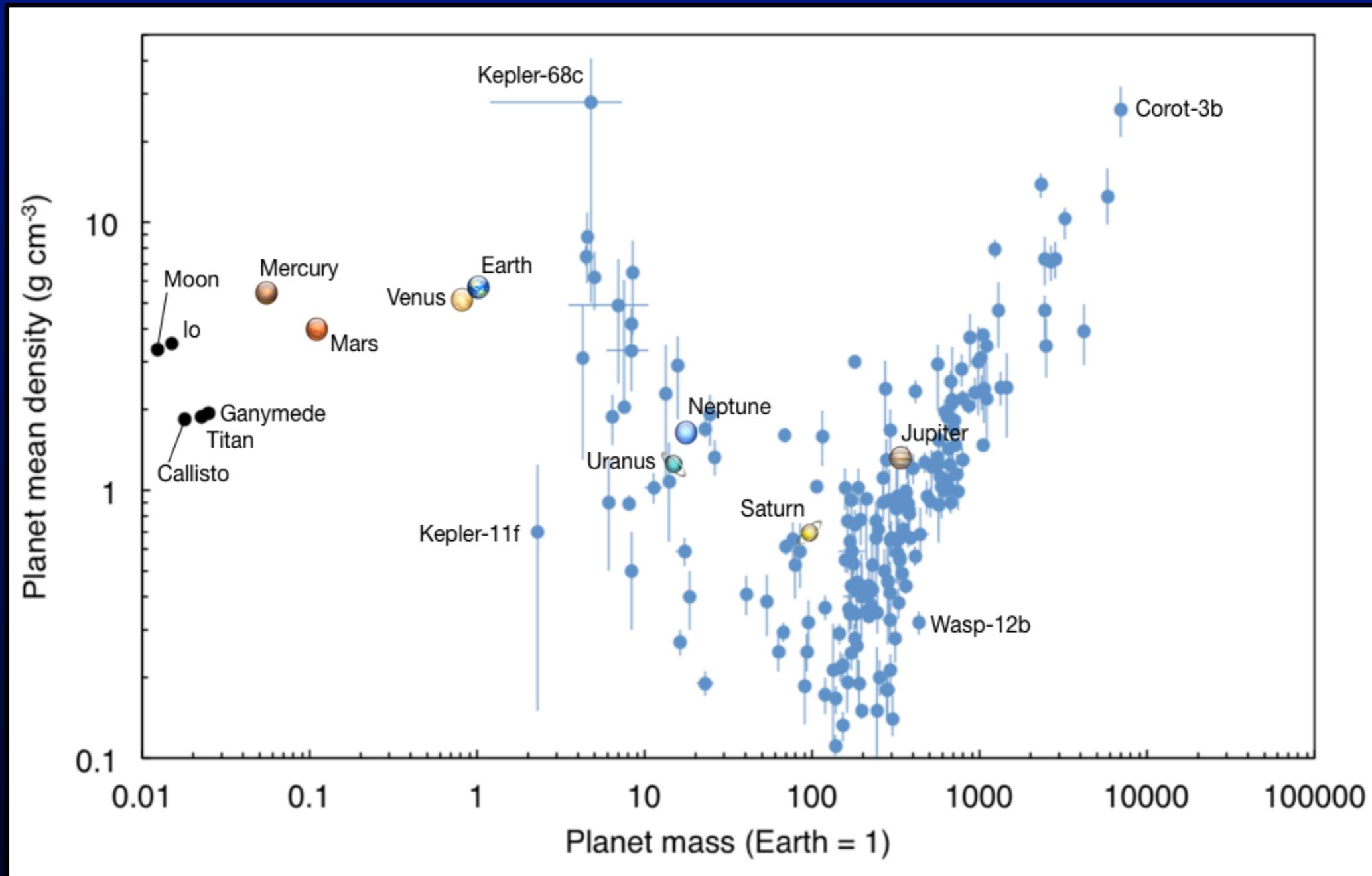
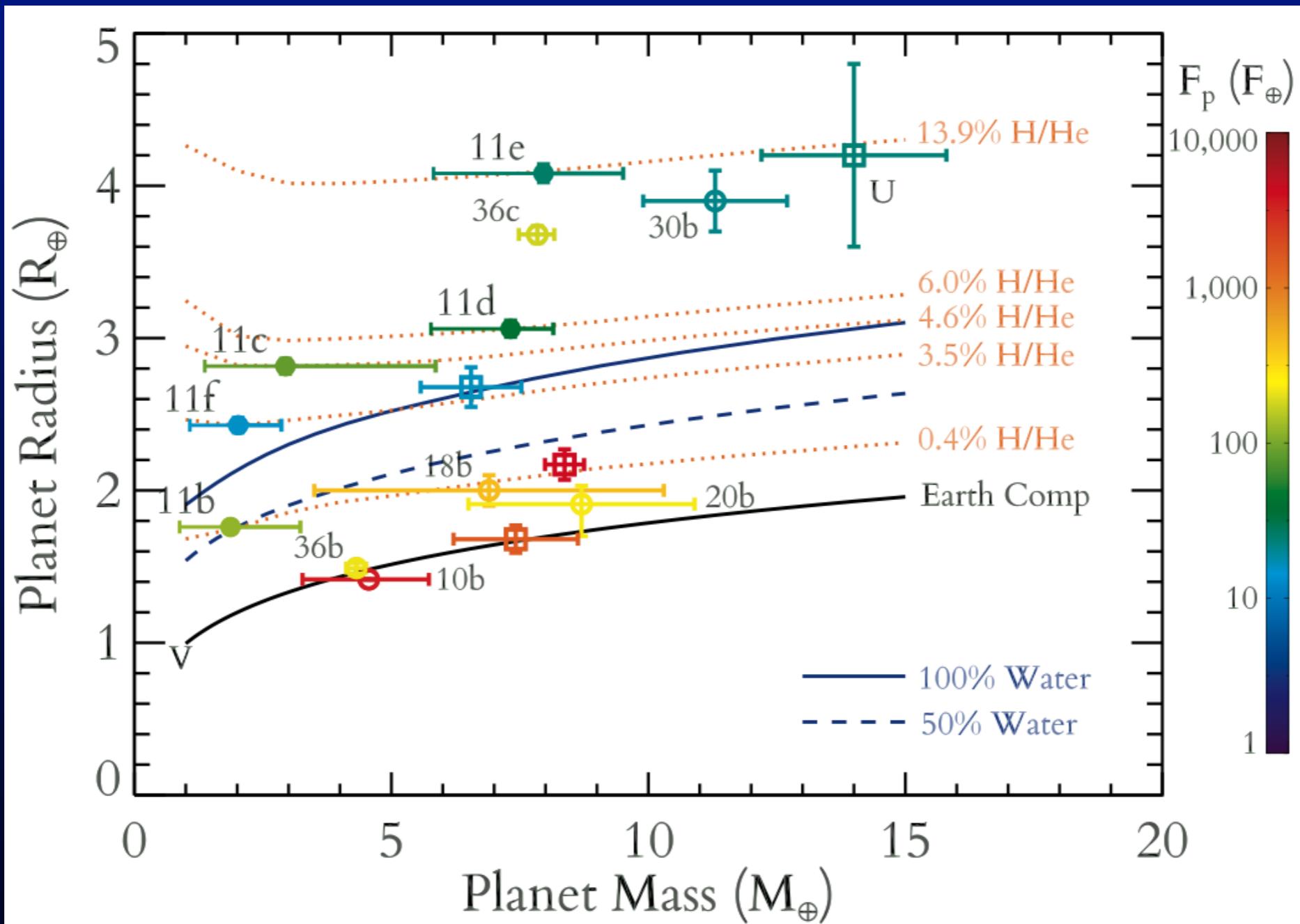


Figure courtesy of Didier Queloz

- Tight correlation for rocky & giant planets; large scatter in intermediate region.
- Dominant source of error is often stellar properties.

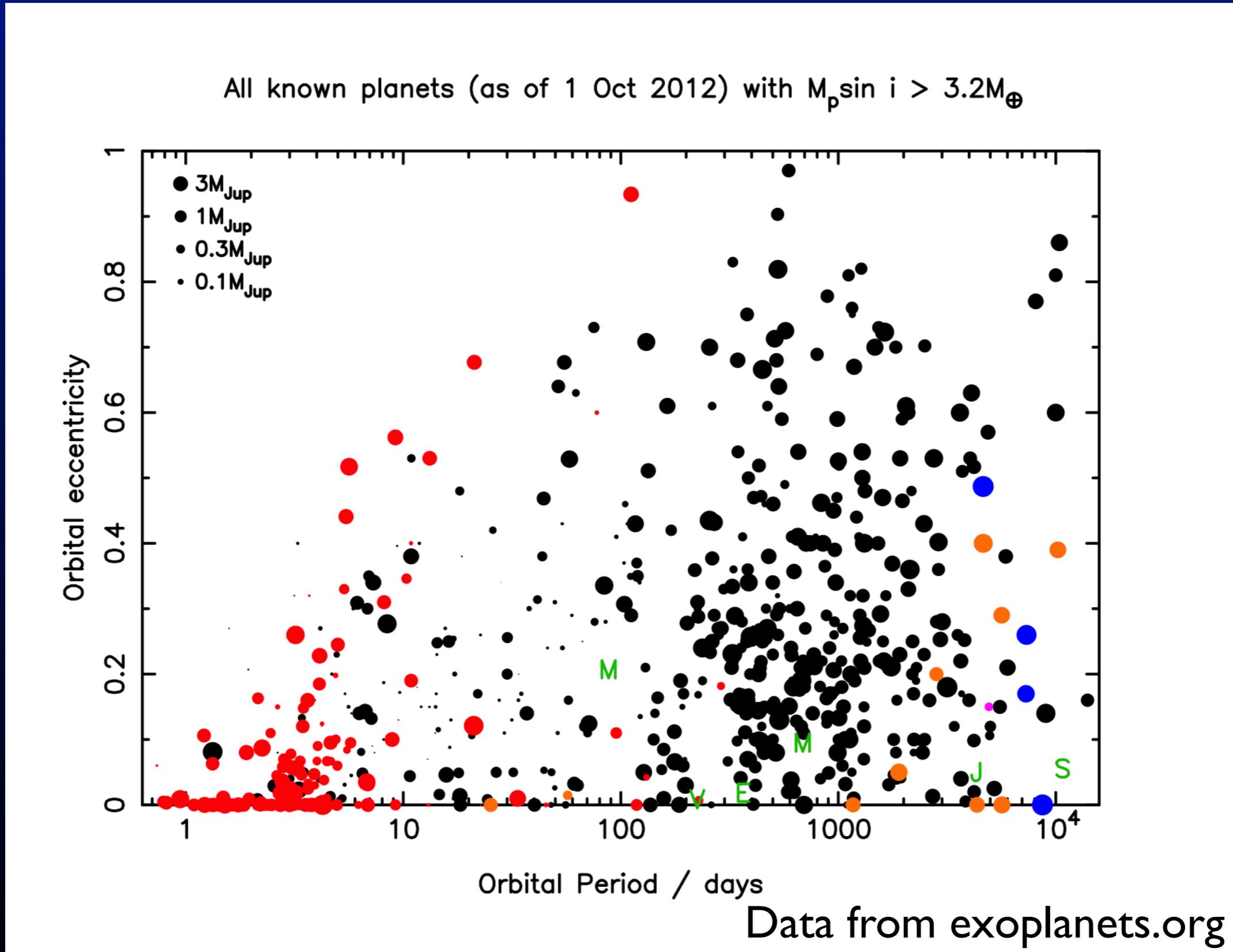
Mass-radius relation



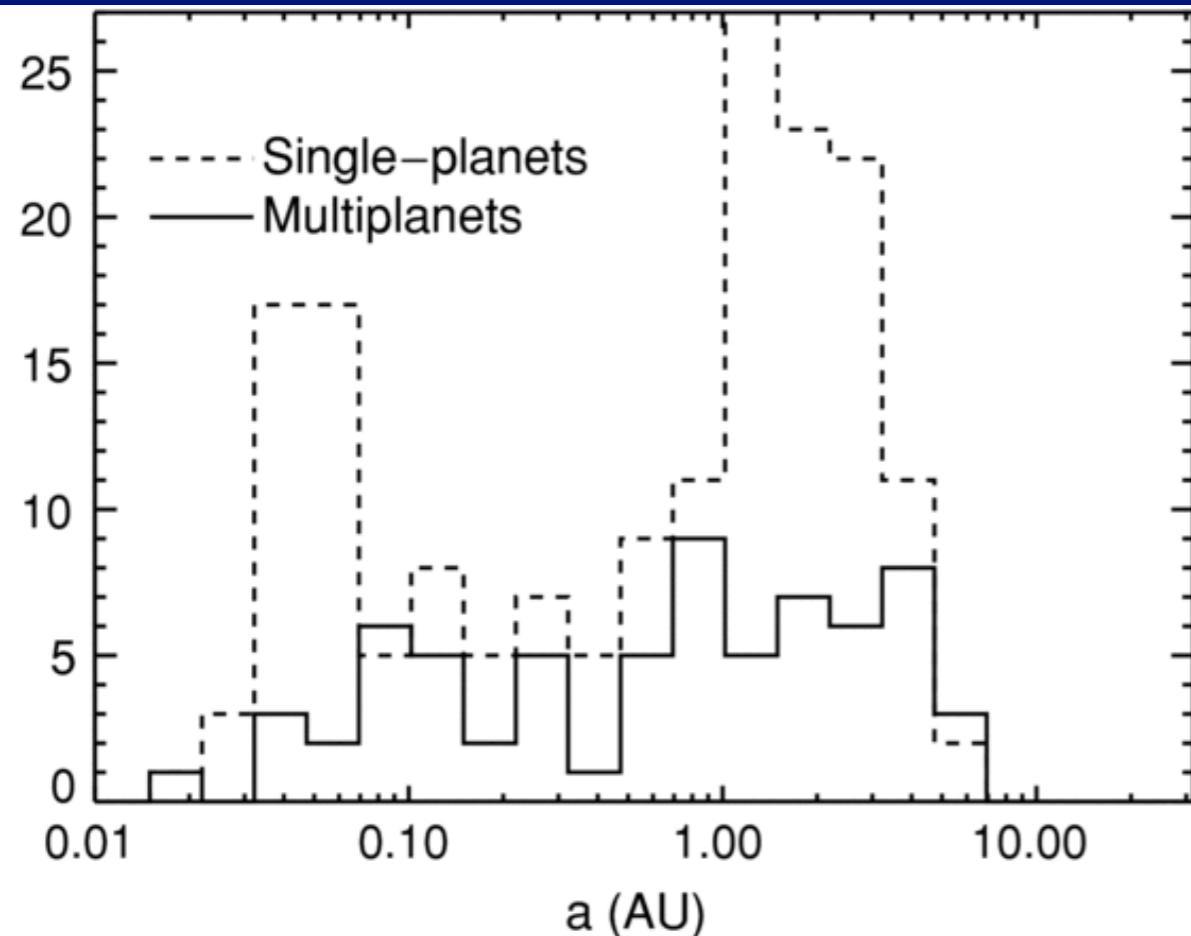
Lissauer et al. (2013)

- Comparison to models possible, but in many cases mean density not strongly constraining.
- However, some exoplanets are unambiguously rocky!

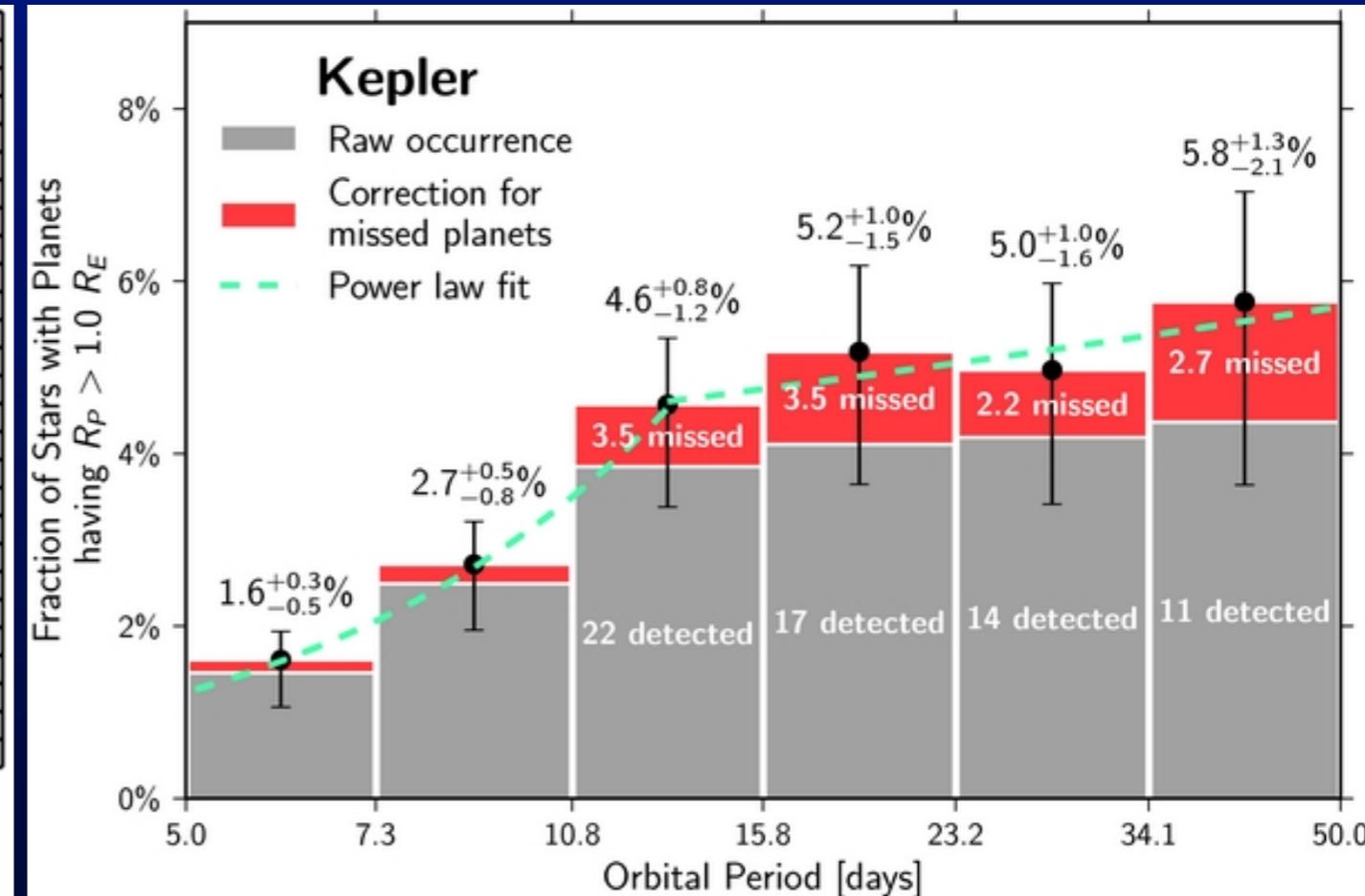
Eccentricities



Radial distribution



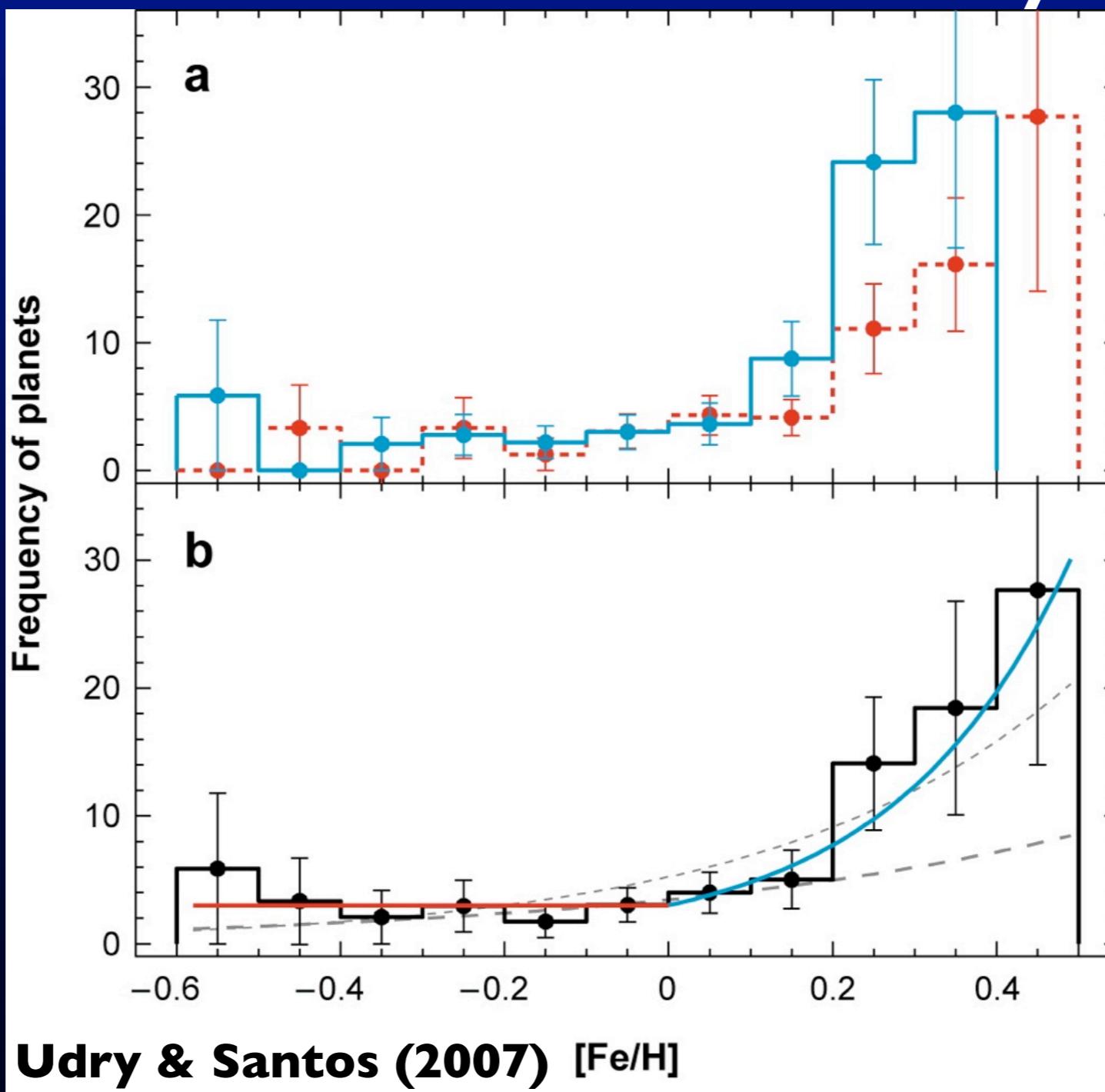
RV: Wright et al. (2009)



Kepler: Petigura et al. (2013)

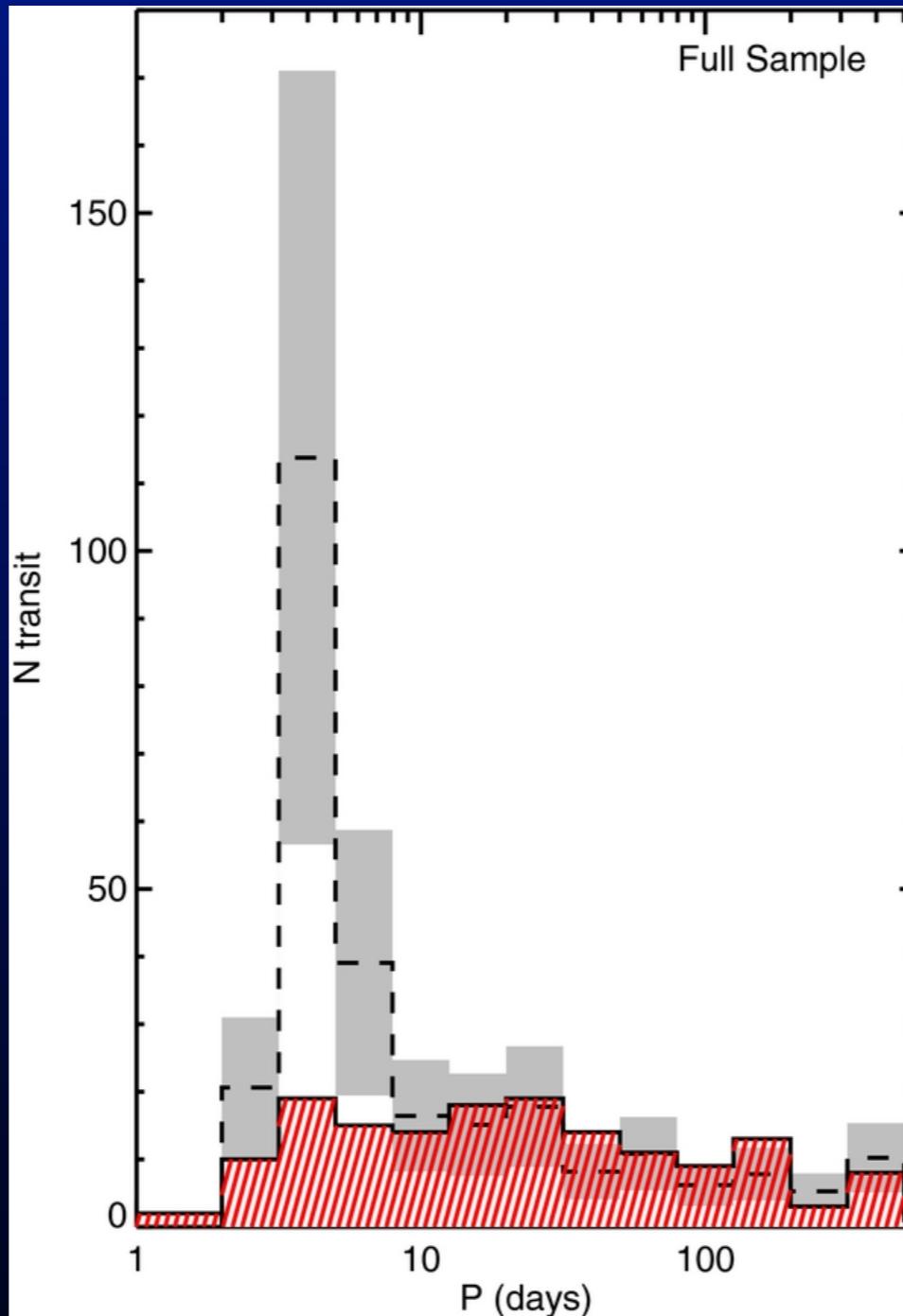
- Radial distribution is “smooth”, though data are limited.
- Evidence of excesses of \sim Jupiters at \sim 0.05AU and \sim 1-2AU in RV data.
- “Pile-up” of hot Jupiters only seen in metal-rich stars.

Host star metallicity



- Probability of hosting giant planets increases very sharply with host star metallicity.
- Appears not to hold for Neptune-mass planets.

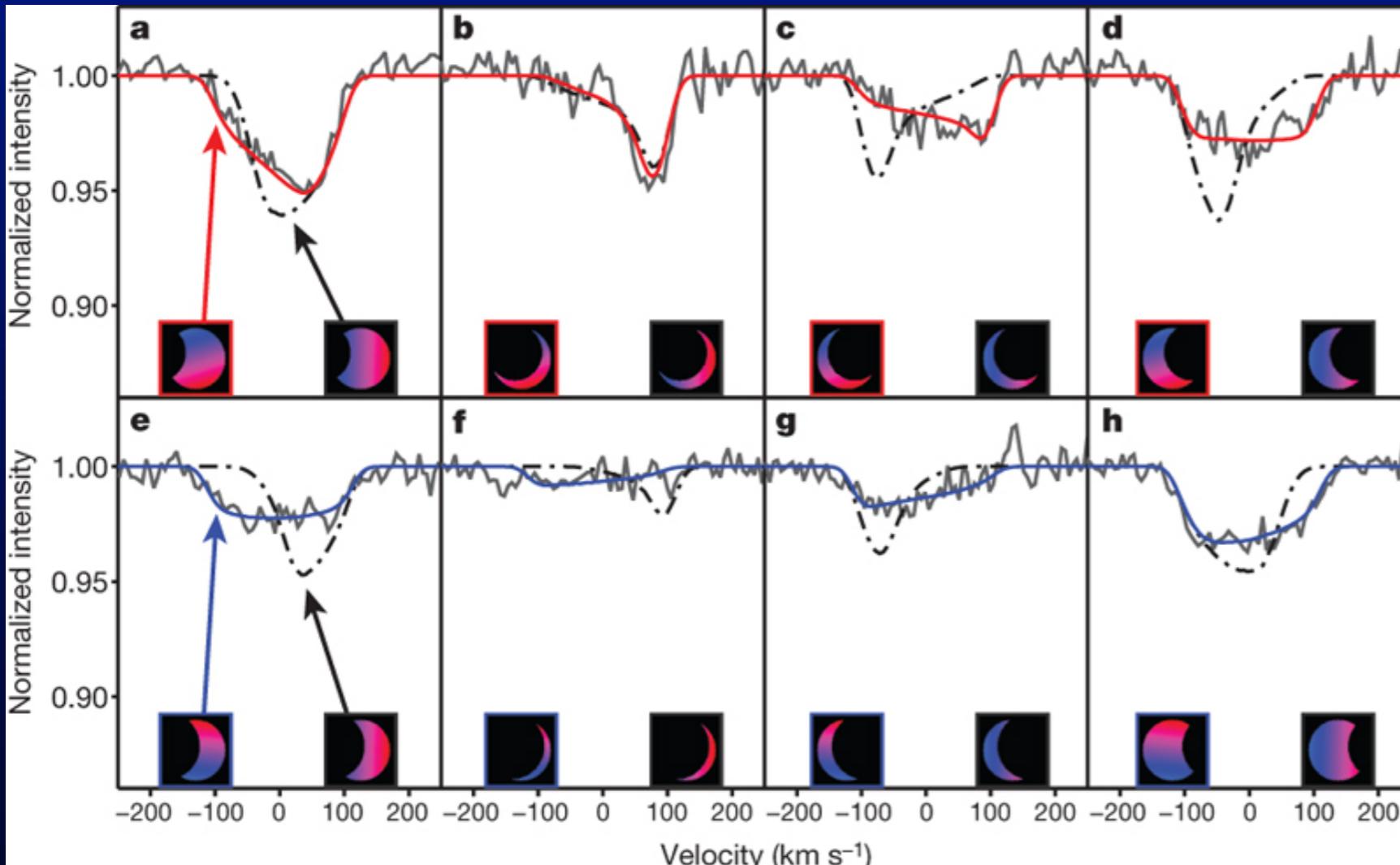
Host star metallicity



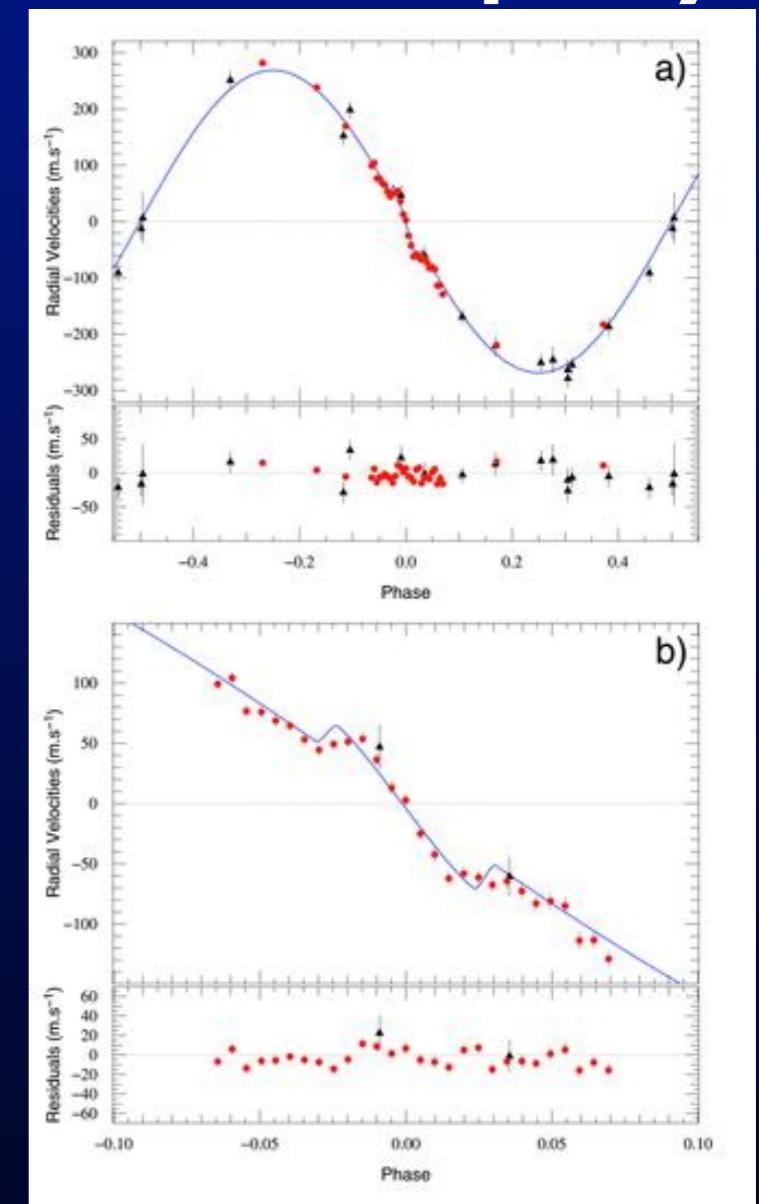
**Dawson & Murray-
Clay (2013)**

- Systematic differences between RV & Kepler samples.
- Most likely explanation is metallicity: Kepler stars are more distant than RV sample, with lower $\langle Z \rangle$.

Rossiter-McLaughlin effect & obliquity



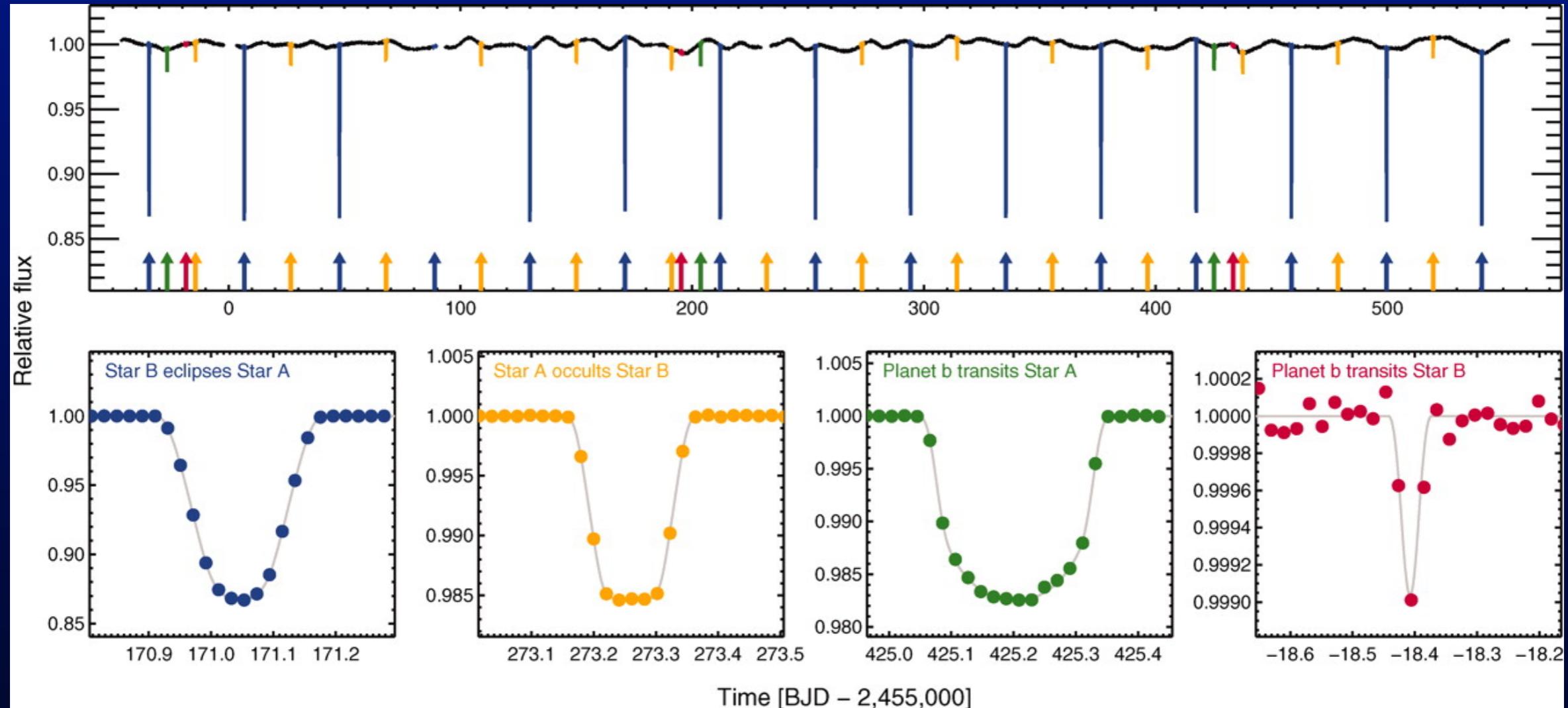
Albrecht et al. (2009)



Triaud et al. (2010)

- Line shifts during transit (R-M) allow us to measure relative inclination of orbit and stellar rotation axis.
- Significant fraction ($\sim 10\text{-}50\%$) of short-period gas giants show high (projected) obliquities.

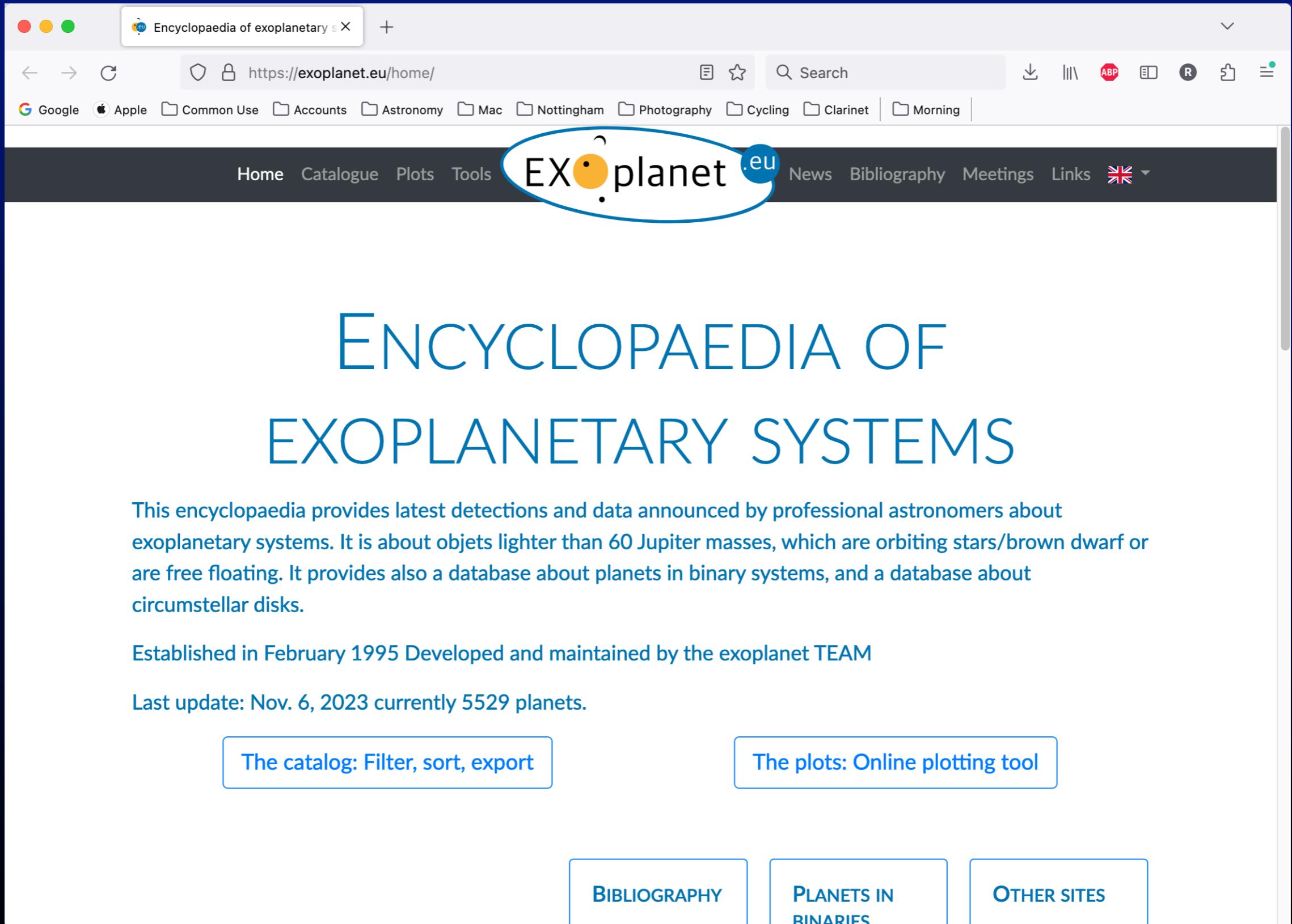
Kepler-16b: the first “Tatooine”



Kepler-16b: Doyle et al. (2011)



Exoplanet Resources



A screenshot of a web browser displaying the homepage of the Encyclopaedia of exoplanetary systems. The page features a large title 'ENCYCLOPAEDIA OF EXOPLANETARY SYSTEMS' in blue text. Below the title is a paragraph describing the purpose of the site. Further down, there are links for catalog filtering, plotting tools, and other resources.

Encyclopaedia of exoplanetary systems

https://exoplanet.eu/home/

Home Catalogue Plots Tools EXOplanet.eu News Bibliography Meetings Links 

ENCYCLOPAEDIA OF EXOPLANETARY SYSTEMS

This encyclopaedia provides latest detections and data announced by professional astronomers about exoplanetary systems. It is about objects lighter than 60 Jupiter masses, which are orbiting stars/brown dwarf or are free floating. It provides also a database about planets in binary systems, and a database about circumstellar disks.

Established in February 1995 Developed and maintained by the exoplanet TEAM

Last update: Nov. 6, 2023 currently 5529 planets.

The catalog: Filter, sort, export

The plots: Online plotting tool

BIBLIOGRAPHY

PLANETS IN BINARIES

OTHER SITES

Exoplanet Resources

NASA Exoplanet Archive

https://exoplanetarchive.ipac.caltech.edu

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EXOPLANET EXPLORATION Planets Beyond Our Solar System

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5,535 Confirmed Planets 10/24/2023 →

398 TESS Confirmed Planets 10/21/2023 →

6,875 TESS Project Candidates 10/20/2023 →

View more Planet and Candidate statistics →

Explore the Archive

Name or Coordinates Search

Optional Radius (arcsec) Advanced Search →

Transit Surveys 130,041,578 Light Curves

Tess

Confirmed Planets →

ExoFOP-TESS →

Project Candidates → | Community Candidates →

TESS Kepler K2 KELT UKIRT

Exoplanet Mass vs. Period

Mass – Period Distribution

02 Nov 2023 exoplanetarchive.ipac.caltech.edu

Mass [Jupiter Masses]

Period [days]

Legend:

- Radial Velocity
- Transits
- Microlensing
- Imaging
- Timing Variations
- Orbital Brightness Modulation
- Astrometry

News → 1 2 3 4 | Plots → 1 2 3 4

The screenshot shows the homepage of the NASA Exoplanet Archive. At the top, there's a navigation bar with links for Home, About Us, Data, Tools, Support, and Login. Below this, four boxes display key statistics: 5,535 confirmed planets (last updated 10/24/2023), 398 TESS confirmed planets (last updated 10/21/2023), 6,875 TESS project candidates (last updated 10/20/2023), and a link to view more planet and candidate statistics. To the right, there's a section titled 'EXOPLANET EXPLORATION' with a subtitle 'Planets Beyond Our Solar System' and social media links for Wikipedia, Facebook, Twitter, and YouTube. On the left, there's a 'Transit Surveys' section featuring the Tess logo and information about the 130,041,578 light curves collected. This section also includes links for Confirmed Planets, ExoFOP-TESS, Project Candidates, and Community Candidates, along with buttons for TESS, Kepler, K2, KELT, and UKIRT. On the right, there's a large, detailed scatter plot titled 'Exoplanet Mass vs. Period' showing the relationship between planetary mass (log scale from 10^-3 to 100 Jupiter masses) and orbital period (log scale from 0.1 to 10^10 days). The plot includes a legend for various detection methods: Radial Velocity (red circles), Transits (green squares), Microlensing (blue triangles), Imaging (purple stars), Timing Variations (yellow diamonds), Orbital Brightness Modulation (orange squares), and Astrometry (yellow stars). Below the plot, there are news and plot navigation links.