

## 1. Overview

Large scale structures in the universe form hierarchically. Massive galaxies are formed through mergers of smaller galaxies, groups grow through the accretion of galaxies, and clusters in turn accrete both groups and individual galaxies. Over the same epoch where these structures experience significant growth ( $z < 1$ ), the fraction of star forming galaxies within them decreases, and at a faster rate than for field galaxies (Saintonge et al. 2008; Finn et al. 2010). It is now widely accepted that there must be physical processes at work in dense environments to actively quench star formation (e.g. Lewis et al. 2002; Balogh et al. 2004), and many studies show that galaxy star-formation rates are suppressed at distances up to  $\sim 5$  virial radii from cluster centers (see Fig. 1 right panel; Lewis et al. 2002; Gómez et al. 2003; Bahé et al. 2013). This suggests that galaxies are affected by environmental processes before they fall into clusters (e.g. Poggianti et al. 1999; Cortese et al. 2006). However, despite significant effort to understand the quenching processes (e.g. Wetzel et al. 2013), state-of-the-art hydrodynamic (e.g. Davé et al. 2011) and semi-analytic (e.g. Guo et al. 2011; Hirschmann et al. 2014) models significantly overpredict the efficiency of quenching in dense environments. The result of this failure is that we still do not know which mechanisms dominate the environmental suppression of star formation and where these mechanisms are effective.

A major impediment to progress is the simplistic way in which environments have been defined. Galaxy environments are typically characterized by their local density or by their distance from the nearest group or cluster. However, this "spherical cow" approach to environments is clearly inadequate as the structure around clusters of galaxies is composed of a rich filamentary network through which galaxies move in their transition from regions of low to high global density (Fig. 1 left panel); a large range of local densities can exist at the same distance from a cluster. Indeed, this mixing of different densities may be a cause of the large scatter seen in star-formation properties of galaxies as a function of cluster-centric radius (Fig. 1 right panel).

To resolve this issue we have embarked on a multi-year program to examine how galaxies evolve in the filamentary structure around clusters and groups. These filaments are easily identified in the nearby universe using SDSS spectroscopy (Fig. 2). We propose a comprehensive survey of the warm ionized, neutral, and molecular gas and dust properties in galaxies that reside in 4 nearby filaments at approximately the distance of the Virgo cluster. We will create spatially-resolved maps of the gas and dust for a large sample of filament galaxies. We will combine our data with results from the literature and compare with in-house theoretical models. This will allow us to determine the role the filament environment plays an active role in quenching galaxies. We are well prepared to carry out this analysis as we: **1)** have extensive experience in the analysis of environmental effects on galaxies **@@references@@**, **2)** have already made significant progress on our observing campaign, **3)** can draw on extensive measurements from the literature (SDSS and HI) in the filaments and the full range of data on local clusters and groups, **4)** have institutional access to facilities necessary to complete our observing program, and **5)** have our own theoretical models with which we will be able to interpret our observational results. The area around the Virgo cluster is ideal for this experiment for two primary reasons. First, Virgo is one of the best studied regions of the sky, and our program directly benefits from the wealth of ancillary data. Second, the galaxies are close enough to allow CO and HI observations of a large sample of galaxies in a reasonable amount of observing time, yet the galaxies are far enough away to avoid significant aperture effect from

single-dish HI and CO measurements.

The combination of our multi-wavelength dataset, our well-defined parent sample of filament galaxies and comparison galaxies, and our theoretical models that we will use to compare with our results will allow us to accomplish the following *specific aims*:

- **Find the effect of the filament environment on the total contents of gas and dust:**

→ For 200 filament galaxies surrounding Virgo, we will characterize the content of the ISM, using H $\alpha$  and WISE for star formation rates, SDSS for stellar masses, HI for the neutral gas mass, and CO for the molecular gas mass. By comparing to field, group, and Virgo cluster core samples, we will determine the environmental density at which the ISM content of galaxies is first modified.

- **Find the effect of the filament environment on the spatial distribution of star formation and dust:**

→ In addition to measuring the ISM content of galaxies within Virgo filaments, we will take a step further and characterize its spatial distribution. We will use SDSS to measure the extent, morphology, and asymmetry of the stellar disk, H $\alpha$  for the star-forming disk, and WISE for the dust disk. This spatial analysis will give us a richer picture of how large scale structures affect the internal ISM of galaxies. We will compare these data to significant amounts of data taken by our team on the cores of local (but more distant than Virgo) clusters and groups.

- **Compare our data to theoretical models to constrain the physical mechanisms that alter gas supplies:**

→ We have in house semi-analytic models that make robust and accurate predictions for the spatial extent different gas phases for field galaxies. These models incorporate prescriptions for where and how starvation and ram pressure stripping affect the gas contents. By comparing our observations to these models we will be able to determine what mechanisms are dominant in what environmental regime.

- **Leverage our observations and literature data to calibrate WISE imaging as a probe of molecular gas contents:**

→ We will use the CO luminosities and spatially resolved WISE imaging of our target galaxies, as well as ample literature data on the dust-to-gas ratio in galaxies to calibrate the use of WISE 12 $\mu$ m spatial profiles as a tracer of the molecular gas extent, which would then allow us to extend our analysis to all nearby filaments found in SDSS.

Our project is novel and guaranteed to succeed for several reasons: **1)** we move away from the simple field/group/cluster trilogy and instead the complex network of galaxies around a well studied cluster; **2)** we will be able to characterize multiple phases of the ISM using extensive multi-wavelength data that either we have in hand or have a reliable path towards obtaining; **3)** we have in-house theoretical models that will help us interpret the observations in terms of physical processes; **4)** our team consists of experts in galaxy evolution and the effects of environment. We anticipate that this original approach will yield the following expected outcomes. We will learn whether filamentary structures host important environmental process that affect the gas in galaxies and how strongly these putative process act. We will determine which, if any, of the prescriptions currently employed in semi-analytic models correctly reproduces the observed characteristics of the galaxies. We will

calibrate techniques to extend our study to the large number of filaments covered by WISE and SDSS.

## 2. Scientific Motivation

**Studying the relative distribution and amount of gas and stars can help identify the dominant processes that deplete gas in galaxies.** Intrinsic processes may quench star formation through ejection or heating of the gas without impacting the distribution of existing stars (e.g. Springel et al. 2005; Croton et al. 2006; Dekel & Birnboim 2006). External or environmentally-driven processes such as tidal interactions and mergers can affect the distribution of both gas and stars (Springel et al. 2005; Croton et al. 2006; Dekel & Birnboim 2006), whereas pressure-driven interactions are expected to act primarily on the gas. For example, starvation, which results from a galaxy being cutoff from its supply of cold gas (Larson et al. 1980), is expected to result in truncated gas disks while the spatial distribution of the remaining disk gas is circularly symmetric and the stellar disk is unaffected (e.g. Kawata & Mulchaey 2008). The interaction of galactic disk gas with the intracluster medium via ram-pressure stripping can remove the gas and produce asymmetries in the remaining disk gas (e.g. Quilis et al. 2000; Cowl et al. 2005). The environmental effects thus have different signatures on the relative abundances and spatial distribution of the warm ionized gas, neutral, and molecular gas. By observing both phases, we can distinguish among these processes.

Large galaxy redshift surveys have revealed that galaxies are distributed in a complex network of matter with a large dynamic range of local density, called the cosmic web or filamentary structures (Kitaura et al. 2009; Darvish et al. 2014). These structures are seen in striking clarity in both simulations (left panel Fig 1) and around the Virgo cluster as shown in Figure 2.

Hydrodynamic simulations of cluster infall regions predict that the regions in filaments with the highest density gas can enhance the ram pressure in filaments by up to a factor of  $\sim 100$  (Bahé et al. 2013). This means that freshly infalling galaxies with  $\log(M_\star) < 9.5$  near a massive cluster can be stripped of their cold gas even well outside the virial radius. For more massive galaxies and at larger distances from the cluster, the ram pressure in filaments is still sufficient to strip off the hot gas that will replenish the dense star-forming gas, although it will likely not affect the densest cold gas. Filaments may therefore be the key sites where galaxies are affected by their environments before they fall into clusters.

## 3. Gas and Dust Content of Virgo Filament Galaxies

We propose to explore the physical properties of galaxies in the filaments around the Virgo cluster. Virgo is an ideal target as it is one of the best studied clusters. However, there is only sparse data on galaxies in the well-defined filaments leading into Virgo from large radii, making our project especially timely. We will complement our observations with the rich ensemble of data already obtained for the center of Virgo: the atomic gas with the VLA by Chung et al (2009), and also at large scale with ALFALFA (Giovanelli et al. 2005); the dust content with the Herschel Virgo Cluster Survey (Davies et al. 2010); the stellar mass with WISE (Ferrarese et al. 2012); the recent star formation from UV data with GALEX (Boselli et al. 2011); and CFHT  $H\alpha$  imaging.

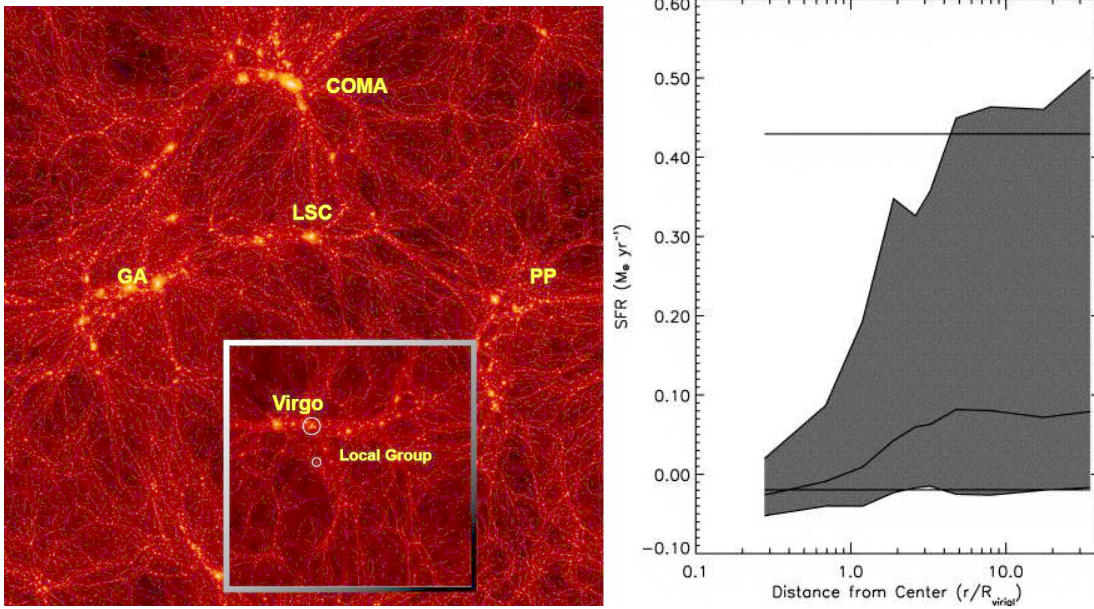


Fig. 1.— *(Left)* Simulation from CLUES showing filamentary structure surrounding clusters in the nearby Universe. Galaxies stream along filamentary structures and into clusters. *(Right)* Figure 6 from Gómez et al. (2003) showing that the upper envelope of star formation rates begins to drop at  $4R_{vir}$  from clusters; there is a very large scatter which complicates the interpretation of this trend. Much of the scatter may come from the large range in local densities at a given projected radius due to the filamentary structure surrounding clusters. Understanding the impact of environmental processes requires studying galaxies along filaments and not just as a function of projected radius - this is the goal of this proposal.

### 3.1. The Filament Galaxy Sample

Tully (1982) first studied the large-scale structure around the Virgo cluster and identified several concentrations of galaxies that he termed clouds. More recently, Kim et al. (2016) repeat this analysis with a much larger spectroscopic dataset that probes to lower luminosities, and they are able to identify multiple filaments around Virgo (see right panel of Figure 2), some that lead directly into the cluster and others that are nearby but not falling into Virgo. This proposal focuses on two of the filaments identified by Kim et al. (2016), and we highlight these in the left panel of Figure 2. The NGC5353 filament, the first in our study, feeds the NGC5353 group and extends over 20 Mpc, passing close to but not into the Virgo cluster (Kim et al. 2016). The second filament that we will study is the Leo filament. Kim et al. (2016) show that once you project galaxies into a Virgo-centric coordinate system (see right panel of Figure 2), galaxies in this region exist in three distinct filaments. We will target all of these. We focus on the NGC5353 and Leo filaments for two practical reasons: (1) these filaments extend the furthest north on the plane of the sky, and this allows for a longer observing window from northern-hemisphere telescopes, (2) these filaments are among the most distant of the Virgo filaments and thus the galaxies have smaller apparent size. This helps to minimize the aperture correction that we need to apply to CO observations when the galaxy size extends beyond the beam size.

The two filaments contain over 600 galaxies, and we show the  $NUV - r$  color versus stellar mass for these galaxies in Figure 3. The grey circles show all the filament galaxies that are in the NASA-

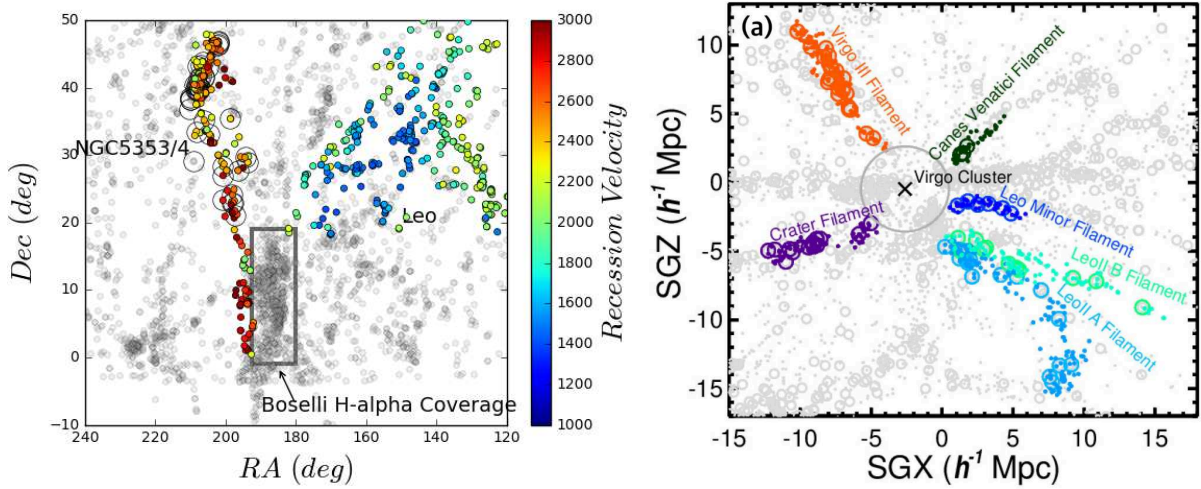


Fig. 2.— (Left) Filaments as identified in Figure 1(a) of Kim et al. (2016) showing filaments surrounding Virgo Cluster as observed on the plane of the sky in RA and Dec and coded by recession velocity. (Right) Figure 2(a) from Kim et al. (2016) showing the same galaxies but projected into Virgo-centric coordinates. The Leo filaments can now be seen as three distinct filaments. The NGC5353 filament, which does not intersect the Virgo Cluster, is not shown in the right panel.

Sloan Atlas (Blanton et al. 2005), which we use as our primary galaxy catalog. The green points show the 120 galaxies we will target with CO observations. We have restricted the mass range for the CO sample to  $9 < \log_{10}(M_{\star}/M_{\odot}) < 10$ . Galaxies below this mass limit are difficult to detect in CO due to lower metallicity and photodissociation of CO (e.g. Cormier et al. 2014). We set the upper limit to the mass range because we expect that galaxies with  $\log(M_{\star}/M_{\odot}) < 10$  will be most affected by the environment. We will push above this mass limit as telescope time permits. The blue circles in Figure 3 show galaxies that we will target with  $H\alpha$  imaging. The  $H\alpha$  subsample includes all galaxies with  $8.5 < \log_{10}(M_{\star}/M_{\odot}) < 10$ . A total of 95 of these are in the NGC5353 filament and the remaining 126 are associated with the Leo filaments. The open red circles in Figure 3 show 184 galaxies that are detected at  $12\mu\text{m}$  by WISE with a signal-to-noise ratio above 10. We set this as the lower SNR limit so that we have sufficient signal for fitting a two-dimensional *Sérsic* model, and we discuss the details of our image fitting in Section 4.3. HI observations are already available in the literature for 75% of these galaxies and we are proposing to get the remainder using the Nançay telescope, which has a low oversubscription. **The acquisition, data reduction, and analysis of the  $H\alpha$  imaging comprise one main thrust of this proposal. The second main goal of this project is to map the spatial distribution of dust within the filament galaxies using WISE  $12\mu\text{m}$  imaging.**

#### 4. Proposed Work

Several groups have shown the power of panchromatic galaxy surveys - SINGS, THINGS, LITTLE THINGS - for understanding the complicated interplay between warm and cold interstellar medium. We take a similar approach, providing a complete census of warm and cold gas and dust. While our data is not as high resolution as the nearby data, comparable data taken as part of the Herschel Reference Survey (Ciesla et al. 2012) and Herschel Virgo Cluster Survey (Davies et al. 2010) show

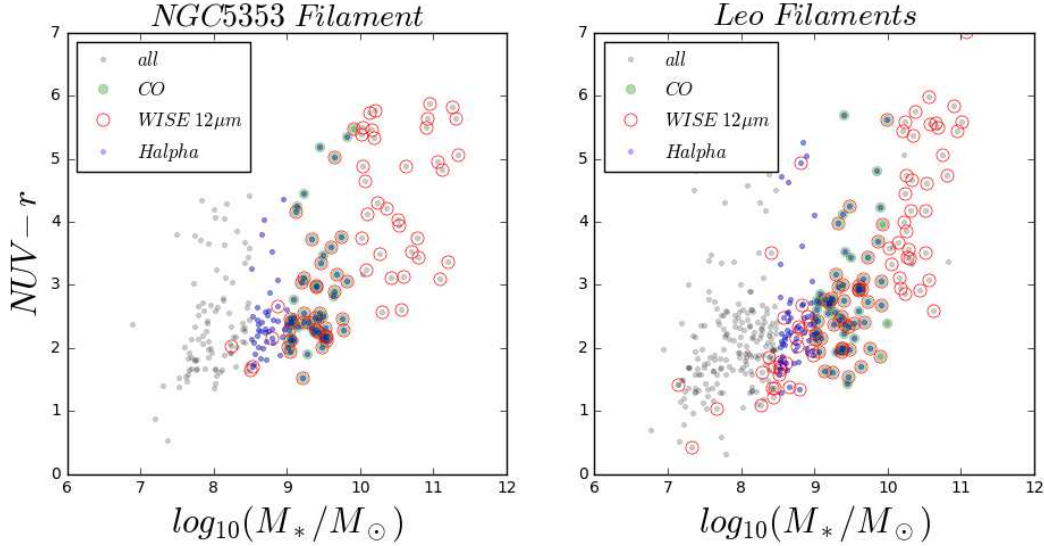


Fig. 3.—  $NUV - r$  color versus stellar mass for galaxies in the (left) NGC5353 filament and the (right) Leo filaments. The grey circles show all the filament galaxies that are in the NASA-Sloan Atlas, and the red circles show all galaxies that are detected at  $12\mu\text{m}$  by WISE with a signal-to-noise ratio above 10. The blue dots show galaxies that we will observe in  $H\alpha$ , and the green dots show the galaxies that we will observe in CO. We apply the upper-mass limit for both the CO and  $H\alpha$  samples because  $\log_{10}(M_*/M_\odot) < 10$  are expected to be most affected by environment. We apply the lower mass limit to the CO sample because CO becomes difficult to detect in galaxies with  $\log_{10}(M_*/M_\odot) < 9$  due to lower metallicity (e.g. Cormier et al. 2014).

the power of our measurements and will allow us to determine how the filament environment affects the many components a galaxy’s gas reservoir.

We will quantify the stellar and dust components using SDSS imaging, optical spectroscopy, and far infrared (WISE) fluxes. Cutting off the hot gas supply or stripping the diffuse molecular gas is a likely precursor to the suppression of star formation as the galaxy will use up its cold gas on a timescale of  $\sim 2.3$  Gyr (Bigiel et al. 2011). Likewise, direct stripping of the cold gas will result in a suppression of the CO and infrared luminosity and truncated  $H\alpha$  emission (e.g. Koopmann et al. 2004). Therefore, we expect a variety of gas effects to be observable by observing a sample of galaxies in  $H\alpha$  HI, the mid-infrared, and CO. Note that Virgo is the closest relatively massive cluster. There is no counterpart in which the major question of the origin of star formation quenching could be addressed in such exquisite details.

#### 4.1. Molecular and Atomic Gas

The CO luminosity can be used to estimate the molecular Hydrogen mass that is the direct fuel for star formation. The conversion from CO luminosity to  $\text{H}_2$  mass depends on metallicity, which in turn depends on stellar mass, as well as on parameters such as cloud size, which scale with the SFR density. This conversion has been calibrated both observationally and theoretically (see Bolatto et al. (2013) and references therein) and we will use it to estimate the molecular gas mass of our galaxies. This is crucial towards understanding how the moderately dense gas is affected by environmental gas processes.



Team members Jablonka and Combes are leading the observing campaign to measure CO(2-1) and (1-0) for all filament galaxies with  $9 < \log_{10}(M_*/M_\odot) < 10$  using the IRAM 30-m telescope. During October 2016, they successfully observed 40 galaxies in the NGC5353 filament, and 38 of these were detected. We show the CO spectra for three of these detections in the right panel of Figure 4. The team has an additional block of time in December 2016, and we expect to double the number of CO detections. Jablonka and Combes will continue to apply for IRAM time to complete the CO observations.

HI is the reservoir from which  $H_2$  is made and traces the more diffuse gas. As it is optically thin, it is also trivial to convert an HI luminosity to a neutral Hydrogen mass. Fairly, as it typically has much larger radial extents it is much more sensitive to environmental effects and may provide evidence of the external gas supply being cut off (starvation).

75% of our target galaxies already have HI observations and co-PI Combes is leading the effort to observe all of the filament galaxies in HI that don't already have existing observations. She has submitted a proposal to use the Nançay  $200 \times 35$  m<sup>2</sup> telescope, which has a low oversubscription rate and to which she has preferred access. She will continue to request time on this telescope until the HI observations are complete (see letter of collaboration).

With the HI and CO observations of our targets we will accomplish the following immediate goals: (1) We will measure their total molecular and neutral gas content; (2) we will compare the atomic and molecular content as different gas depletion mechanisms affect the diffuse and dense gas differently; (3) We will compare this to the galaxy sizes, SFRs, and stellar masses to determine the gas deficiency of filament galaxies (if any) compared to galaxies in the field, group, and clusters with similar data as obtained from the literature @@references@@.

#### 4.2. Ionized Gas and Current Star Formation Rates

$H\alpha$  is the standard for measuring star-formation in local galaxies (e.g. Kennicutt 1998), and the combination of  $H\alpha$ , UV, and far-infrared imaging provides a powerful measure of star-formation rates that is independent of extinction. The combination of precise star-formation rates and CO observations will allow us to calculate gas consumption timescales and characterize multiple phases of the galactic gas. Equally important, the spatial extent of  $H\alpha$ , when compared to the radial distribution of the underlying stellar population and the dust (§4.3), provides a powerful means to identify the physical processes that affect a galaxy's gas supply (e.g. Hodge & Kennicutt 1983; Dale et al. 2001; Gavazzi et al. 2012; Boselli et al. 2015). Studies of the Virgo cluster show evidence of cold gas stripping (e.g. Koopmann & Kenney 1998, 2004; Dale et al. 2001; Crawl et al. 2005; Chung et al. 2007; Corbelli et al. 2012; Gavazzi et al. 2012; Boselli et al. 2015), including truncated  $H\alpha$  emission of Virgo spirals compared with their field counterparts (Koopmann & Kenney 2004). We will be able to determine if environmental transformation starts in the filaments, before galaxies are accreted into the densest environments.

While extensive  $H\alpha$  imaging has been done in groups, clusters, and the field, little has been done to map the spatial extent of  $H\alpha$  in filament galaxies. The goal of our program is to obtain spatially resolved  $H\alpha$  maps for 222 star-forming galaxies in the NGC5353 and Leo filaments. Key to testing quenching mechanisms are that we are able to measure  $H\alpha$  profiles to low surface brightness and that we can probe galaxies at different positions along the filament out to large distances from

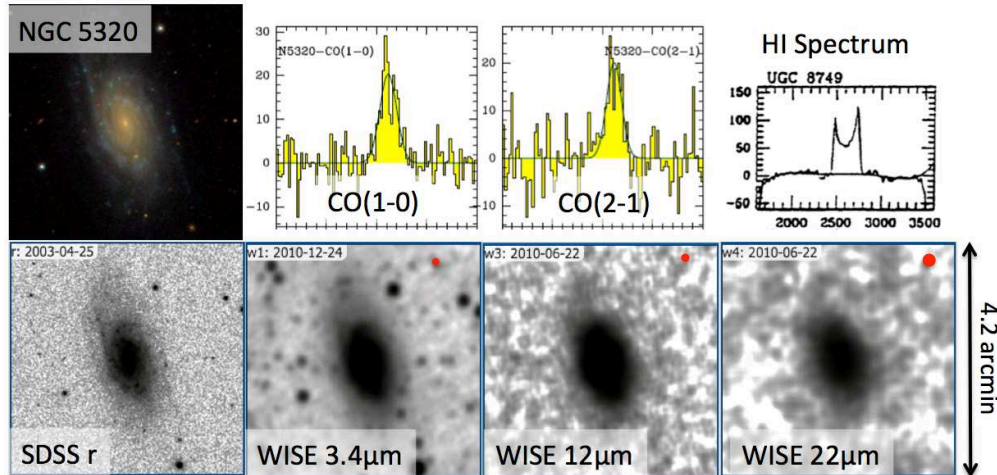


Fig. 4.— Multi-wavelength images of spiral galaxy NGC 5320. The 4 panels on the right are the WISE images. The resolution of WISE is sufficient to measure the spatial extent of dust emission using the  $12\mu\text{m}$  images. The  $22\mu\text{m}$  fluxes indicate the amount of dust-obscured star formation and thus provide an important complement to the  $H\alpha$ -derived star-formation rates (*Left*) Newly acquired CO (1-0) and (2-1) spectra from IRAM 30-m telescope for galaxies in the NGC5353 filament taken during October 2016. Of the 40 galaxies that were observed, 38 were detected in CO. Collaborator Jablonka will observe an additional 40 galaxies in December 2016, and she will continue observations until the CO sample is complete. (*Right*) SDSS color images showing 16 randomly selected galaxies in NGC5353 filament.

Virgo. To properly probe any environmentally-driven quenching, we will detect galaxies with star-formation rates 10 times lower than the star-forming main sequence.

We will complete the  $H\alpha$  observations at the WIYN 0.9 and the Phillips Claud Telescope (PCT) at the Mt. Laguna Observatory. For the WIYN 0.9 m, we will apply for time through NOAO (already submitted a proposal for Spring 2017 semester), requesting approximately 6 nights each spring for the three years covered by this proposal. Based on past  $H\alpha$  imaging experience with the WIYN 0.9 m, we need about 2 hours per target, and so we expect to be able to complete 4 objects per night and 24 pointings per run. Our yield will be slightly higher ( $\sim 30$ ) because we will be able to place multiple objects within the  $0.5^\circ \times 0.5^\circ$  field of view. Thus we expect to complete  $H\alpha$  imaging for  $\sim 100$  galaxies at the WIYN 0.9m.

We will observe the remaining 120 galaxies in the  $H\alpha$  sample using the PCT. This 1.25m telescope has a  $4k \times 4k$  detector with a  $22' \times 22'$  field of view. It is in the final phase of commissioning and should be ready for operation in Spring 2017. The University of Kansas is a partner on this telescope with a 15% share of the total time, of which co-PI Rudnick is entitled to a significant share. This telescope can be remotely controlled and so will allow for observing either from KU or from Siena College; remote observing will save substantially on travel costs for such a large observing program. We have included funds to purchase a narrow-band filter suitable for observing  $H\alpha$  at the redshift of our filaments.

PI Finn is involved in an  $H\alpha$  imaging survey of nearby galaxy groups as part of the Undergraduate ALFALFA Team. She and her undergraduate students have developed code to reduce  $H\alpha$  imaging



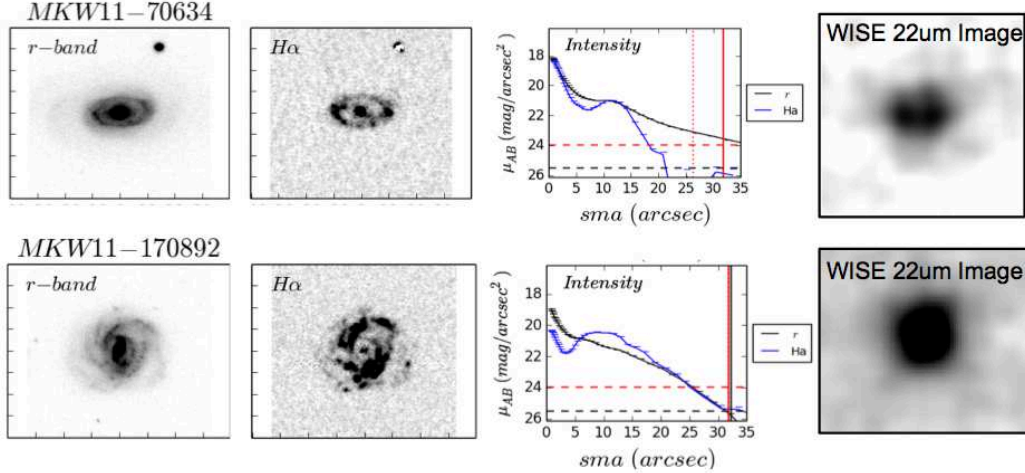


Fig. 5.— To illustrate our proposed  $H\alpha$  analysis, we show imaging taken with the KPNO 0.9-m+HDI of two galaxies within the nearby group MKW11 ( $v_r = 6900$  km/s). The left column shows a galaxy for which the star-forming disk probed by  $H\alpha$  is truncated along the semi-major axis (sma) relative to the stellar disk as probed in the  $r$ -band. The right column shows an example that is not truncated. The SFR is resolved at about  $1/10$  of a kpc in  $H\alpha$ . In contrast, the WISE PSF is 12 arcsec, corresponding to about 1 kpc. The image thumbnails in each row are the same size, demonstrating that the WISE data are not sufficient to measure the extent and morphology of the star formation. Indeed, the WISE photometry may underestimate the SFR for the lowest mass galaxies in our sample because of their low metallicity. As these galaxies are likely the most susceptible to stripping,  $H\alpha$  imaging is necessary to understand how environment affects their gas.

taken with the HDI camera on the WIYN 0.9 m telescope. These programs will be used to reduce the  $H\alpha$  imaging for this proposal, and we will adapt the code to accomodate the imaging data from Mount Laguna observatory.

#### 4.3. Dust Masses and Spatially-resolved Dust Maps

WISE 12-micron images are sensitive to emission from polycyclic aromatic hydrocarbons, and we will thus use these images to trace the spatial extent of dust in the filament galaxies. While the WISE PSF is rather large ( $6.5''$  at  $12\mu\text{m}$ ), it still probes down to kiloparsec scales for galaxies in our filaments, and is sufficient for this study (see Figure ??). Along with SDSS fiber spectroscopy, the WISE colors will be used to eliminate galaxies with infrared size measurements that are likely to be affected by an AGN.

To quantify the extent of the WISE  $12\mu\text{m}$  images, we will use *GALFIT* software (Peng et al. 2002) to fit two-dimensional *Sérsic* models to the galaxy images. We will use *unWISE* image products and PSFs (Lang 2014) because they are optimized for extended sources whereas the WISE image products have been optimized for point sources. Finn has completed a similar analysis using *GALFIT* to analyze MIPS  $24\mu\text{m}$  images for galaxies in 9 nearby galaxy clusters (Finn et al. in prep). The spatial resolution of the MIPS  $24\mu\text{m}$  images is comparable to the resolution of WISE at  $12\mu\text{m}$ , yet the filament galaxies are significantly closer than the galaxies that we analyzed with MIPS. We are thus confident in our ability to measure robust sizes at  $12\mu\text{m}$ .

The analysis of the WISE imaging will be the most computationally intensive part of the proposed

work because we plan to run multiple models for each galaxy to assess the impact of our initial conditions on the resulting size measurements. For each galaxy, we will stochastically sample across the parameter space of initial conditions, and each model has five free parameters. Nonetheless, using extensive testing, we estimate that it would take less than 2 days of computational time on the Siena College High Performance Computer Cluster (HPCC, see Facilities section of this proposal). We will compare the GALFIT results to those generated by other independent fitting programs such as GALPHAT (Yoon et al. 2011) and The Tractor (Lang et al. 2016).

Use our CO observations and those from the literature to calibrate the use of WISE profiles as the tracer of the dense gas.

#### 4.4. Comparison with Theory

An important aspect of our work is using theoretical models to understand what processes are affecting the gas distribution in our galaxies. Our collaborator De Lucia is an expert in semi-analytic models of galaxy formation and will produce models directly suited for comparison with our observations. The first scenario we must consider is whether our results are consistent with predictions of gas depletion through starvation and the consumption of the remaining gas rather than more extreme environmental processes such as ram-pressure stripping. In the left panel of Figure 6, we show the predicted size of the stellar and star-forming components of  $z = 0$  galaxies based on the semi-analytic models of Xie et al. (2016, in prep). These models include starvation but do not explicitly include environmental processes such as ram-pressure stripping. In the right panel of Figure 6, we show the *measured* size of the star-forming and stellar disks versus stellar mass for the *Local Cluster Survey* galaxies (Finn et al. 2016, in prep). While the size of the stellar disks are consistent with the model predictions, the size of the star-forming regions fall systematically below the model predictions. This indicates that additional environmental processes must be included to explain the small observed size of the star-forming region in these galaxies. However, the *Local Cluster Survey* sample is small and does not probe filamentary structures surrounding the clusters.

Theoretical models of ram-pressure stripping provide predictions that we can compare directly to our measurements. For example, numerical simulations of starvation and ram-pressure stripping of cold gas generically predict that star formation in the edges of galaxies will be affected more strongly than star formation near the centers of galaxies; the spatial extent of the star-forming disk should be smaller for galaxies that are undergoing stripping (e.g. Kawata & Mulchaey 2008; Bekki 2014). We find evidence of this in the *Local Cluster Survey*, but we also find a stronger correlation with bulge-to-total ratio. A large sample that extends beyond the virial radius of clusters is needed to disentangle these effects and this sample has to have low stellar mass, as those galaxies should be more vulnerable to having their gas removed because their gas is not as tightly bound (e.g. Kawata & Mulchaey 2008; McCarthy et al. 2007; Bekki 2014).

The SAMS implemented for many clusters in the Millennium simulation and in cluster zoom simulations will give us a statistical view of how galaxies are affected in filaments around clusters and groups of similar mass to Virgo and NGC 5353. However, none of these clusters actually looks like the surroundings of Virgo, which introduces a systematic error in the comparison. We will therefore implement our SAMS in the Constrained Local UniversE Simulations (CLUES) of the

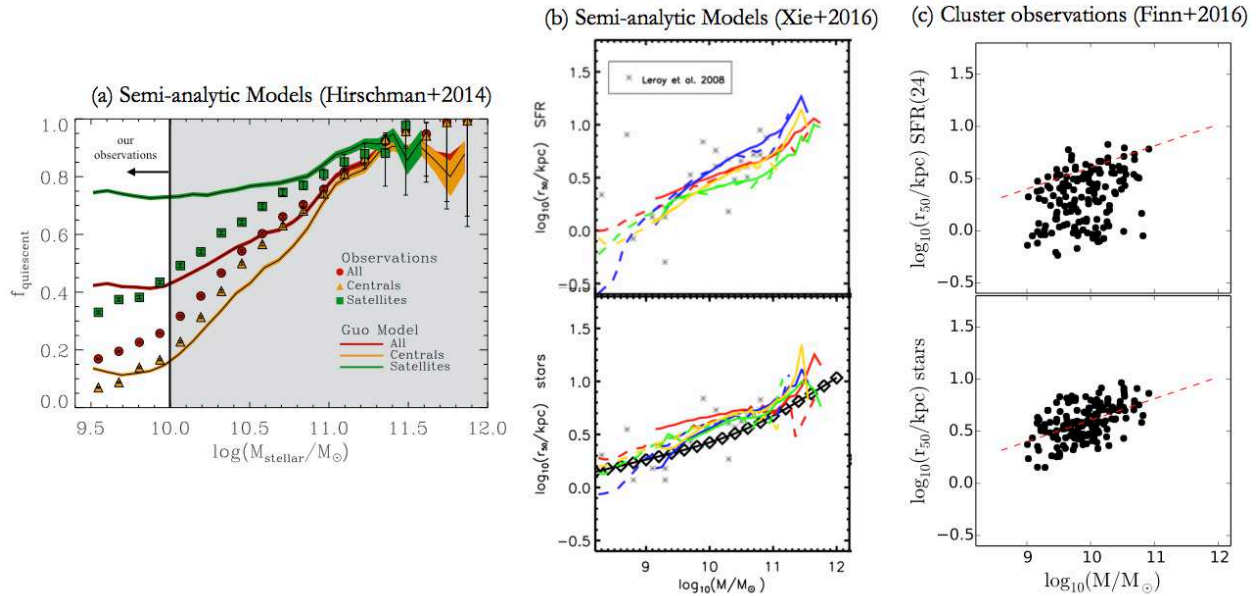


Fig. 6.— (a) (left) **Predicted** half-light radius of stars (bottom) and star formation (top) versus stellar mass for  $z = 0$  galaxies (Xie et al. 2016, in prep) that are undergoing starvation. The different color lines represent different models for partitioning atomic and molecular hydrogen, and the grey points show comparison with some existing observations. (right) **Measured** half-light radius of stars (bottom) and star formation (top) versus stellar mass for galaxies in the *Local Cluster Sample*. The red dashed line is a linear fit to the stellar half-light radius and is shown in the top panel for comparison. The size of the star forming region is smaller than predicted by semi-analytic models, suggesting that starvation is not sufficient to explain the observed size of the SF region in these galaxies. A larger sample of galaxies with more robust size measurements is needed to confirm this result.

local volume<sup>1</sup>. Specifically, Co-I De Lucia will work to integrate her SAMS with the CLUES dark matter merger trees so that she can predict the spatial distribution of star and star formation for galaxies in a simulation designed to look like the local volume.

#### 4.5. Workplan, Major Milestones, and Timeline for Completion

Justify collaboration - greg brings telescope access and resources, environment studies. Finn brings expertise of IR imaging fitting, spatially resolved SF and  $H\alpha$  imaging observations and reduction.

The division of observing and analysis responsibilities among the team members is outlined in Fig. 7. The expertise and roles of individual team members is described below.

Rose Finn, Ph.D., (PI) is a professor in the Department of Physics and Astronomy at Siena College (Loudonville, NY). She is the PI of the *Local Cluster Survey* and has extensive experience with running GALFIT on MIPS  $24\mu\text{m}$  imaging. She will lead the  $H\alpha$  imaging survey at KPNO and the analysis of the WISE  $12\mu\text{m}$  imaging, supervise the undergraduate students, and draft the paper.

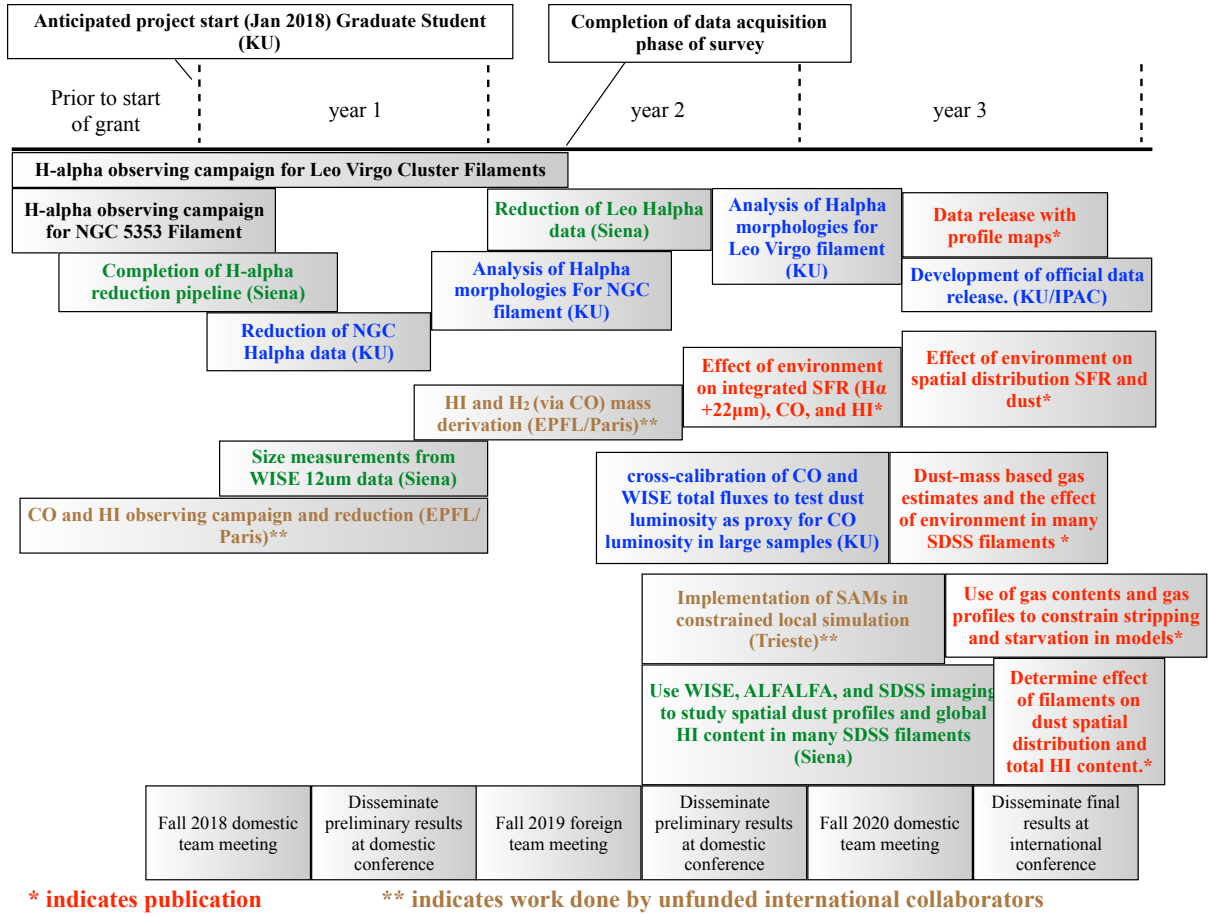
Gregory Rudnick, Ph.D., is an associate professor at the University of Kansas. He has expertise

<sup>1</sup><https://www.clues-project.org>

in galaxy evolution and the effects of environment on galaxy evolution, and he is leading an *HST* study to measure the spatial extent of star-formation in intermediate-redshift galaxies. He will lead the  $H\alpha$  imaging survey at the PCT and will supervise the graduate student whose Ph.D. dissertation will be on this project.

Graziano Vernizzi, Ph.D., is an associate professor at the Department of Physics and Astronomy of Siena College (Loudonville, NY). His research interests lie in computational and theoretical physics, biophysics, nanoscale science, soft condensed matter, and random matrix theory. Graziano will lead the computational aspects of this project.

Fig. 7.— Major Milestones and Timeline



#### 4.6. Data Sharing and Further Dissemination of Results

We will release all of our data with our final publication. This will include reduced, calibrated, and astrometrically aligned HI,  $H\alpha$ , and SDSS maps with integrated CO fluxes. We will produce a variety of measurements of the gas content of filament galaxies. We will produce a catalog of CO, HI, dust,  $H\alpha$  properties to be released with our final paper.

While this data set of derived information will not be exceptionally large in volume, it will be of value to other observational astronomers and theorists working on galaxy evolution. To facilitate the

widespread use of these data, we will publish the full catalog of measured and derived quantities on the web and in the on-line version of Astrophysical Journal Supplement. In addition, we will work with the NASA/IPAC Infrared Science Archive (IRSA) to ensure that our data are discoverable to a larger user base. Co-I Vandana Desai, resident at IPAC, will act as our NED liaison to ensure that this gets done in a timely manner.

## 5. Results from Previous NSF Projects

### 5.1. co-PI Rudnick

#### 5.1.1. Molecular Gas in Distant Galaxies

Co-PI Rudnick has recently completed his NSF project 1211358 “Characterizing the Molecular Gas Contents of High Redshift Galaxies” (\$306,754; 8/1/12-7/31/16).

**Intellectual Merit:** This study was based on a large body of JVL A data (200 hrs) on a  $z = 1.62$  galaxy cluster that was collected between 2012 and October of 2014. The goal of this study was to characterize the molecular gas content of high- $z$  galaxies by observing CO. As a result of the studies of this cluster and of the CO gas content of distant galaxies, Rudnick has authored or co-authored seven papers since 2012 with a total of 185 citations (Papovich et al. 2012; Rudnick et al. 2012; Lotz et al. 2013; Geach et al. 2013; Wong et al. 2014; Geach et al. 2014; Tran et al. 2015) as well as an ApJ paper that is in the resubmission process (Rudnick et al.). Since 2012, Rudnick has also given 28 oral presentations on this NSF project. Using the full JVL A data, Rudnick has securely detected CO(1–0) in two massive and gas-rich galaxies in the  $z = 1.62$  cluster. These galaxies have surprisingly low star formation efficiencies (SFE) for their high mass and gas fraction (e.g. Genzel et al. 2010). This may indicate the presence of environmental effects on the physical conditions of the molecular gas and on the accretion of gas from the cosmic web in a massive halo. These results appear in a paper that is being resubmitted to ApJ after a favorable referee report. The expected publication date is early 2017. As a direct result of this project Rudnick has also organized a large consortium of scientists who are seeking to use ALMA to make a census of the CO gas in distant cluster galaxies. They will resubmit a significantly sized proposal in April 2017.

**Broader Impact:** Rudnick has completed the third year (2013-2016) of an outreach program in close collaboration with Andrew Bricker, a Physics teacher at Lawrence High School (LHS). Rudnick developed and executed a year-long program in which the students receive an introductory calculus-based astronomy course and perform a bona fide research project. The goal of the class is to teach high school students research methods, computing skills, the electromagnetic spectrum, the nature of science, and science communication while also giving the teacher new tools to teach research-based activities in the classroom. The project involves using *Spitzer*/MIPS 24 $\mu$ m data to measure  $L_{\text{IR}}$  and SFRs for the galaxies in the infall regions of intermediate redshift clusters from the ESO Distant Cluster Survey (EDisCS). The students meet every day in a special class period. The teaching assistant funded by the grant performed most of the instructional duties and Rudnick attended class once a week. 30 students have gone through the program during these three years. This total was comprised of  $\sim 50\%$  underrepresented student groups: four African American, three Hispanic, one Native American, and 10 female students, two of which were also women of color. As described in §6.2 we employ extensive assessment to understand our success at meeting learning

goals. This program is continuing in 2016-2017 funded by another NSF project (see §5.1.2).

### 5.1.2. Galaxy Evolution in Distant Clusters

co-PI Rudnick is in the beginning of his second year for the NSF project 1517815, “Collaborative Research: The GOGREEN Survey - Caring About the Environment” (\$347,556; 8/1/15-7/31/18).

**Intellectual Merit:** This study funds the US analysis efforts for the international Gemini Observations of Galaxies in Rich Early Environments (GOGREEN) project. This project is based on the largest Gemini Long and Large Program (PI: Michael Balogh), which is comprised of 443 hours of Gemini imaging and spectroscopic observations conducted over a 4 year period starting in Fall 2014. The two main components of this project are very deep Gemini optical spectroscopy of a stellar mass ( $M > 10^{10} M_{\odot}$ ) limited sample of galaxies in 21 groups and clusters at  $1 < z < 1.5$  and a large multiwavelength imaging program.

The goals of the project are to: **1)** Find the dominant modes of satellite quenching at  $z < 1.5$ ; **2)** Determine how galaxies populate dark matter halos as a function of environment at  $z < 1.5$ ; **3)** Measure the relative timing of morphological transformation and star-formation quenching; **4)** Constrain the dominant driver of size growth in the early-type population at  $z < 1.5$ .

co-PI Rudnick is in charge of the imaging efforts for the whole collaboration, which involves gathering *UBVRIZYJK<sub>S</sub>*[3.6 $\mu$ m][4.5 $\mu$ m] data on a broad suite of telescopes including Subaru, CFHT, Magellan, VLT, and *Spitzer*. The imaging of the southern clusters is 95% complete and the northern clusters only lack their NIR data. We expect the imaging to be completed by the Fall of 2017.

The strategy of the project is to obtain deep spectra over many semesters on the faintest targets, and thus many of the science publications will appear at the end of the proposal period. However, an initial data paper based on the first 30% of the spectroscopic data is in preparation with a Dec. 2016 submission target.

**Broader Impact:** Rudnick has extended his LHS outreach program into the 2016-2017 academic year and will continue it through the 2017-2018 AY. Changes that we have made this year include a much more aggressive targeting of URM students, which we have accomplished by going to more junior students and not having as high of a math prerequisite for entry into the program. As a result we have our highest fraction of URM students yet, with one Native American, three African American, two Hispanic students and three women, one of which is a woman of color. We are currently attempting to expand the program by using seniors who have already completed the program as peer instructors. This will allow us to grow without additional personnel costs.

## 6. Broader Impact

### 6.1. Siena College

**Modeling Physics for High School Programs:** The modeling approach is an innovative and effective way to teach physics that is fundamentally different from traditional techniques. Students are led through carefully constructed experiments and exercises to clearly develop the conceptual, visual, and mathematical models of how physics works. Experienced physicists already have these mental models, but beginning physics students do not. These models are essential for understanding



physics. The modeling approach minimizes lectures, and instead students are actively engaged in collecting and analyzing real-time data that illustrate the fundamental concepts of physics. The students must then construct models to interpret these data.

As a former high school teacher, co-PI Finn knows first-hand the importance of bringing the more effective and engaging techniques to the front line. As part of a previous NSF grant (AST-08XX), we have offered modeling workshop for high school physics teachers for the past 9 summers. We are fortunate to have as an adjunct instructor an area high-school physics teacher who is an expert in using the modeling approach to teach physics, and he will continue to lead these workshops. We offer 6 hours of continuing education credit to participants that can be used toward teacher recertification. This grant will provide a stipend for Darren Broder to organize and lead the workshops. To assess the impact of the modeling curriculum, participating teachers will administer the force concept inventory and mechanics baseline test.

**Undergraduate Research :** The importance of undergraduate research is widely recognized in the science community. *Recent studies have shown that undergraduate research may be the pedagogy for the 21<sup>st</sup> century (e.g., Council on Undergraduate Research Statement and references therein).* Involvement in research projects fosters highly motivated, self-confident students with enhanced analytical and communication skills. co-PI Finn supervised 24 undergraduate students during the tenure of her previous NSF grant (AST-0847430).

Siena College undergraduate students will be involved in all aspects of this proposal. We have budgeted money to bring students on observing trips to Kitt Peak National Observatory, and they will help gather imaging data at PCT through remote observing once the telescope has been commissioned in the Spring of 2017. This grant provides funds for four undergraduates to complete 10 weeks of paid research for each year of the grant. Finn has experience supervising undergraduates on similar projects, and hiring a mix of freshmen through juniors has worked nicely as the older students can help train the freshmen. During the first year, two students will work with Finn to finalize the data reduction pipeline for the KPNO data and to adapt the pipeline for the TCP data. The students will measure radial profiles for the **H $\alpha$**  imaging and develop ways to quantify the **H $\alpha$**  morphologies. The students will work in collaboration with the graduate student at Kansas to compile the **H $\alpha$**  data products.

The other two summer students will work on the analysis of the WISE 12 $\mu$ m images. This part of the project will require more programming skills, so this will likely be assigned to the more experienced undergraduate students. The students will work with Finn and Vernizzi to analyze the images with GALFIT and visualize the results of the multiple models that we will fit to each galaxy.

An important part of the research experience is presenting results to the community. I will encourage all students to present their results at the fall meeting of the Astronomical Society of New York and Siena's Academic Celebration, which is held each spring. In addition, seniors will be encouraged to present a poster at the annual winter meeting of the American Astronomical Society.

To assess the impact of this project, I will track participation, papers, presentation, and post-graduate activity of all Siena students who are involved in this project. I will design and implement