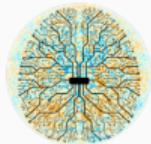


Machine Learning for Radiometer Calibration

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2nd Year PhD Candidate

With: Harry Bevins, Eloy de Lera Acedo, Will Handley, Rohan Patel, Kaan Artuc, Jiacong Zhu

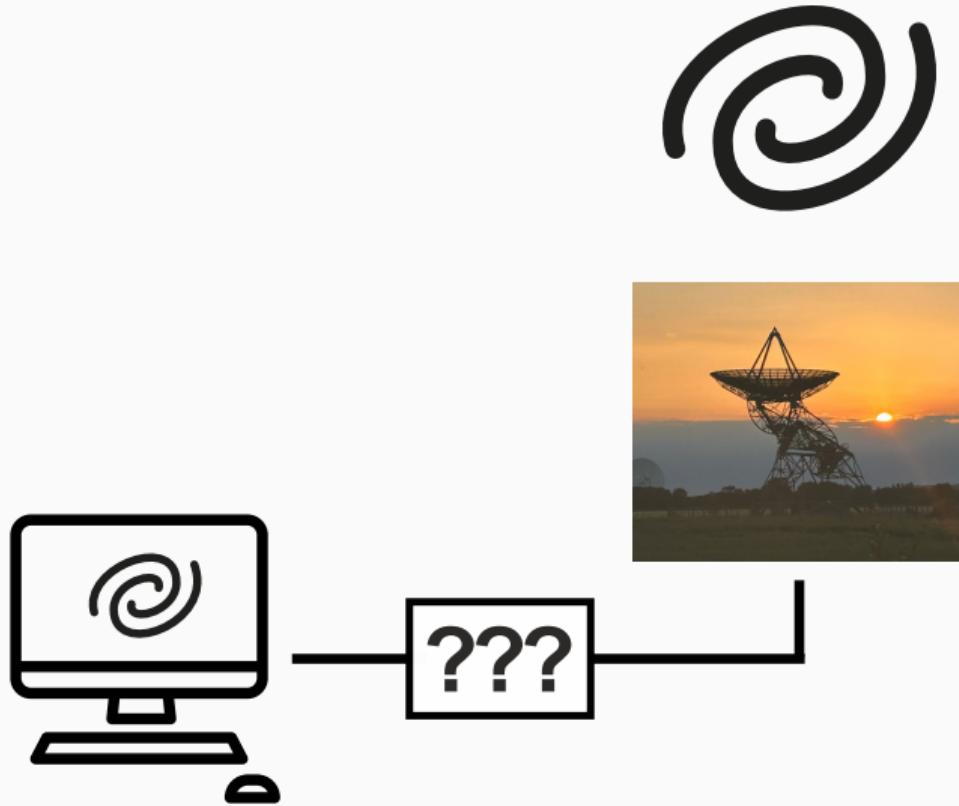


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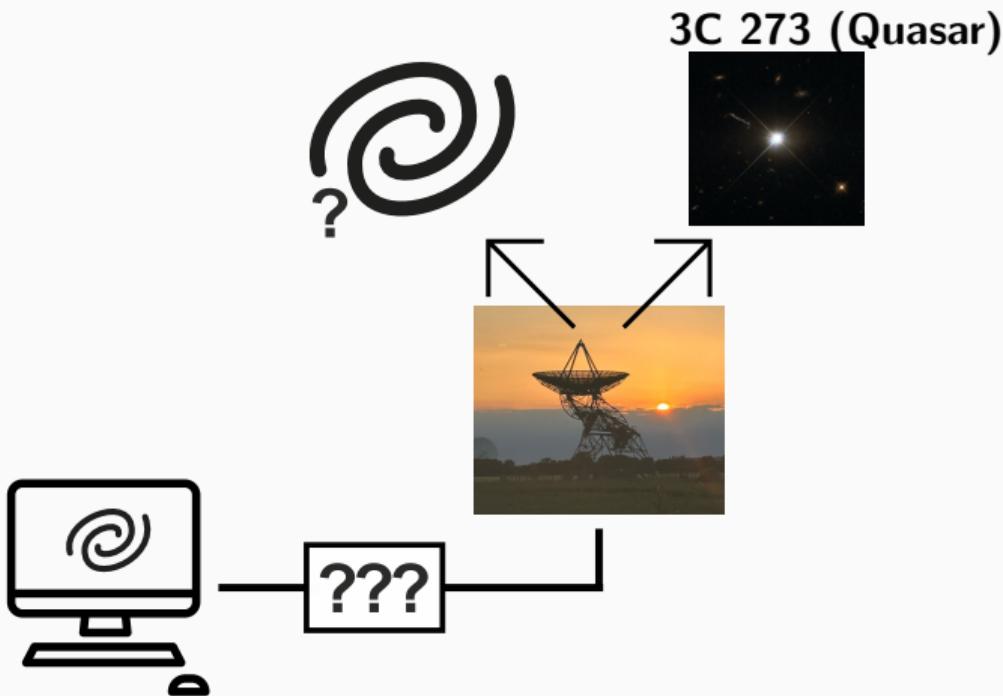


Radiometer calibration overview

What is calibration?



How to calibrate?



Why is calibration in Global 21cm Cosmology difficult?



We measure *sky averaged* signal.

Antenna LNA impedance mismatch

Very faint signal.

How to calibrate (in a bit more detail...)?

Objective: Map input temperature to output power.

Key Factors:

- LNA introduces time-dependent gain, $g(t)$.
- Impedance mismatch adds noise (T_{rec}) to the system.

Link Output Power to Input Temperature:

$$P_{\text{out}}^{\text{src}} = gM \times (T_{\text{in}}^{\text{src}} + T_{\text{rec}}) \quad (1)$$

$$M = \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{(1 - |\Gamma_{\text{cal}}|^2)} | \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_{\text{cal}} \Gamma_{\text{rec}}} |^2$$

Note: All parameters above are frequency-dependent, but the notation has been simplified here and thereafter for convenience.

Receiver/source mismatch

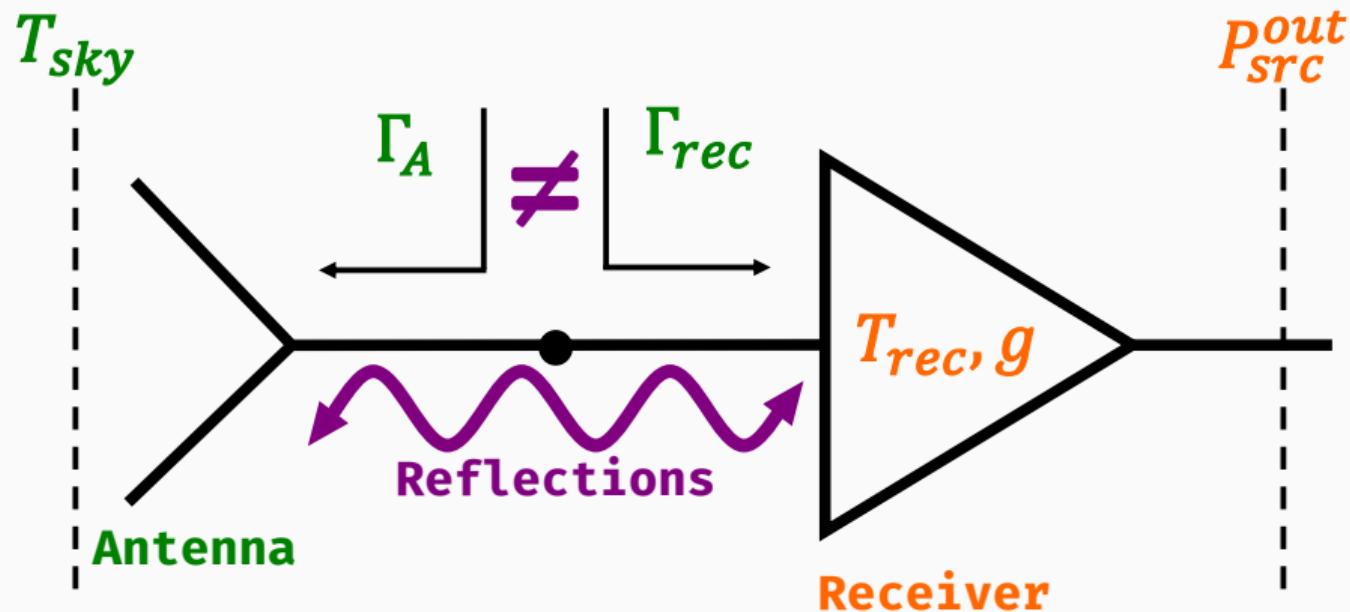


Figure 1: Schematic of the receiver system showing the signal path from antenna to digitizer

Dealing with reflections...

$$P_{\text{out}}^{\text{src}} = gM (T_{\text{in}}^{\text{src}} + T_{\text{rec}}) \quad (2)$$

Noise Parameter Equation:

$$P_{\text{out}}^{\text{src}} = gM \left(T_{\text{in}}^{\text{src}} + T_{\text{min}} + T_0 \frac{4R_N}{Z_0} \frac{|\Gamma_{\text{src}} - \Gamma_{\text{opt}}|^2}{(1 - |\Gamma_{\text{src}}|^2)(1 + |\Gamma_{\text{opt}}|^2)} \right) \quad (3)$$

Noise Wave Equation:

$$\begin{aligned} P_{\text{out}}^{\text{src}} = & g \left[T_0 + T_{\text{unc}} |\Gamma_s|^2 \left| \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_s \Gamma_{\text{rec}}} \right|^2 \right. \\ & + T_s (1 - |\Gamma_s|^2) \left| \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_s \Gamma_{\text{rec}}} \right|^2 + T_{\cos} \Re \left(\Gamma_s \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_s \Gamma_{\text{rec}}} \right) \\ & \left. + T_{\sin} \Im \left(\Gamma_s \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_s \Gamma_{\text{rec}}} \right) \right] \quad (4) \end{aligned}$$

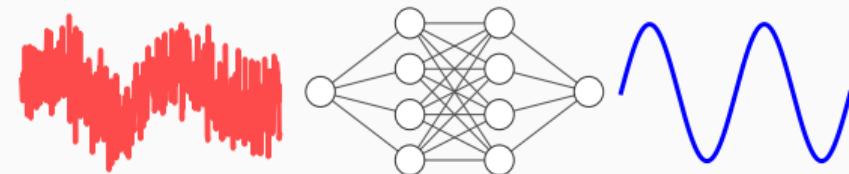
- g - Gain
- T_{min} - Minimum Noise Temperature
- R_N - Noise Resistance
- Γ_{opt} - Optimum Reflection Coefficient
- Γ_s - Source Reflection Coefficient
- Γ_{rec} - Receiver Reflection Coefficient
- T_s - Source Temperature
- $P_{\text{out}}^{\text{src}}$ - Power out
- $T_{\text{unc}}, \cos, \sin$ - Noise wave parameters
- T_0 reference temperature

Radiometer calibration with machine learning

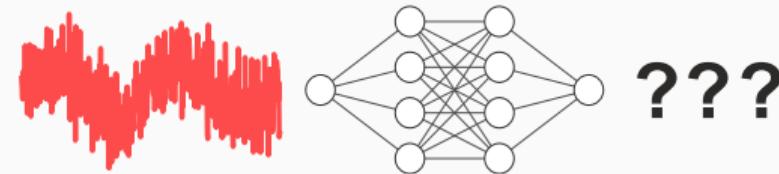
How/why?

- Can predict noise wave parameters using machine learning.
- Could apply method to noise parameters.
- Predict **directly** on noise parameters on frequency by frequency basis.
- Maleable to environmental changes.

Train



Predict



Machine learning calibration steps

1. Define the Loss Function

Regress over measured power and predicted power.

$$\mathcal{L} = \frac{1}{n} \sum_{i=1}^n (\mathcal{P}_{\text{measured},i} - \mathcal{P}_{\text{pred},i})^2 \quad (5)$$

2. Write Down the Equation for $\mathcal{P}_{\text{pred}}$

Using the noise wave formalism, relate $\mathcal{P}_{\text{pred}}$ to T_{src} .

$$\begin{aligned} \mathcal{P}_{\text{pred}} &= \mathbf{g} \cdot M(T_{\text{in}}^{\text{src}} \\ &+ T_{\text{min}} + T_0 \frac{4R_N}{Z_0} \frac{|\Gamma_{\text{src}} - \Gamma_{\text{opt}}|^2}{(1 - |\Gamma_{\text{src}}|^2)(1 + |\Gamma_{\text{opt}}|^2)} \Big) \end{aligned} \quad (6)$$

3. Minimise the Loss Function

$$\theta^* = \arg \min_{\theta} \mathcal{L}(\theta) \quad (7)$$

parameter vector θ includes all tunable parameters in the model:

$$\theta = \{\mathbf{g}, T_{\text{min}}, R_N, \Gamma_{\text{opt}}^\phi, |\Gamma_{\text{opt}}|\} \quad (8)$$

4. Rearrange and Predict (T_{src}) using θ^*

$$\begin{aligned} T_{\text{src}} &= \frac{\mathcal{P}_{\text{pred}}}{\mathbf{g}^* \cdot M} \\ &- \left(T_{\text{min}}^* + T_0 \frac{4R_N^*}{Z_0} \frac{|\Gamma_{\text{src}} - \Gamma_{\text{opt}}^*|^2}{(1 - |\Gamma_{\text{src}}|^2)(1 + |\Gamma_{\text{opt}}^*|^2)} \right) \end{aligned} \quad (9)$$

ML calibration overview

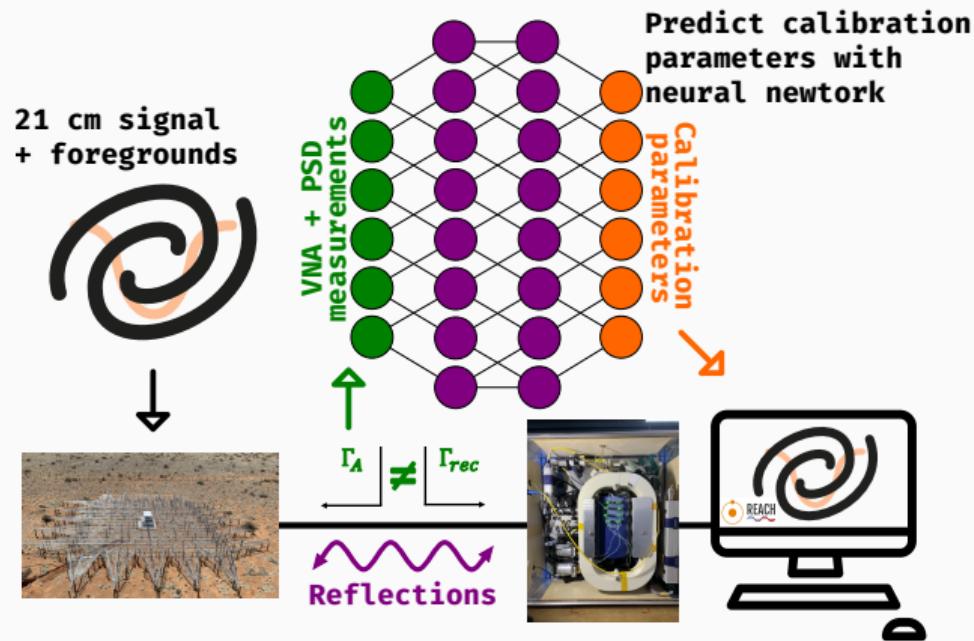
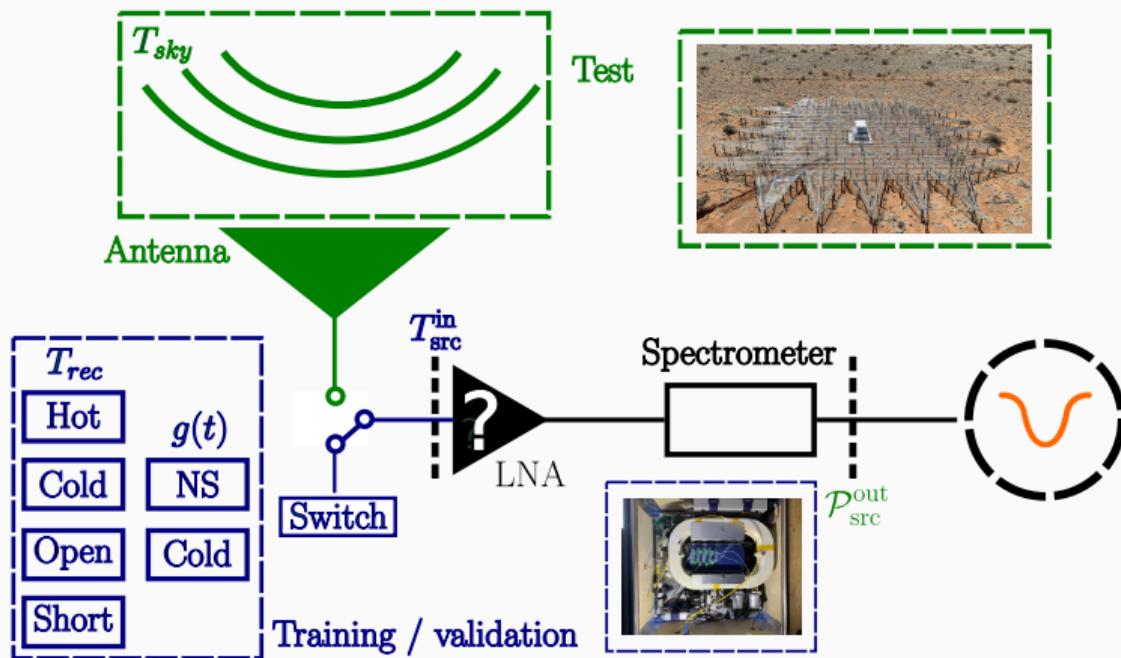


Figure 2: High-level overview of the machine learning-based calibration framework

The physical system

Process:

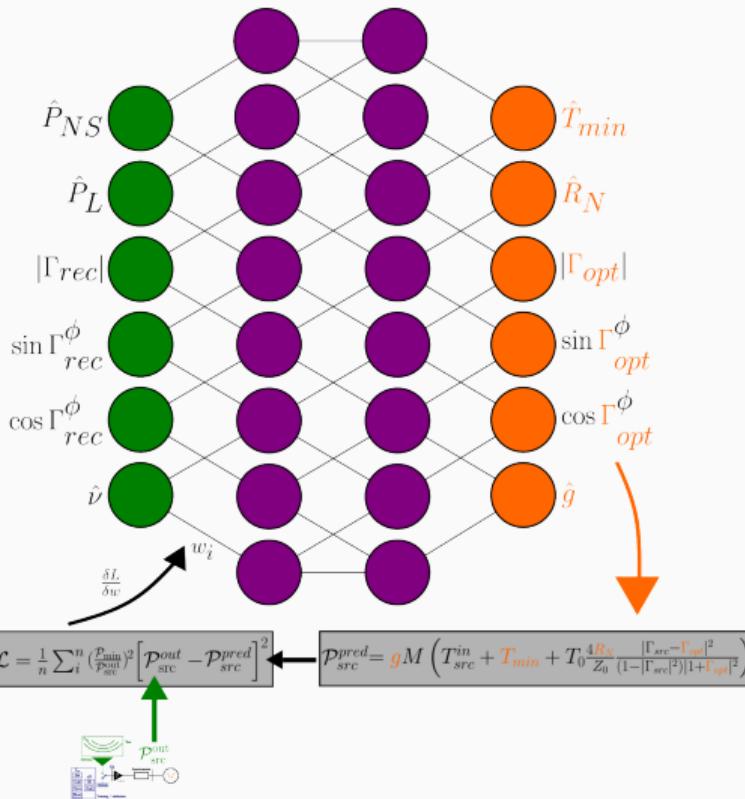
- Switch between sources to generate training data.
- Calibration sources with known temperature train neural net.
- Predict T_{src} of antenna.



Network Architecture

Structure:

- Input thermistor and VNA measurements.
- Also input frequency.
- Predict noise parameters.
- Regress over loss function.



Testing on internal validation source

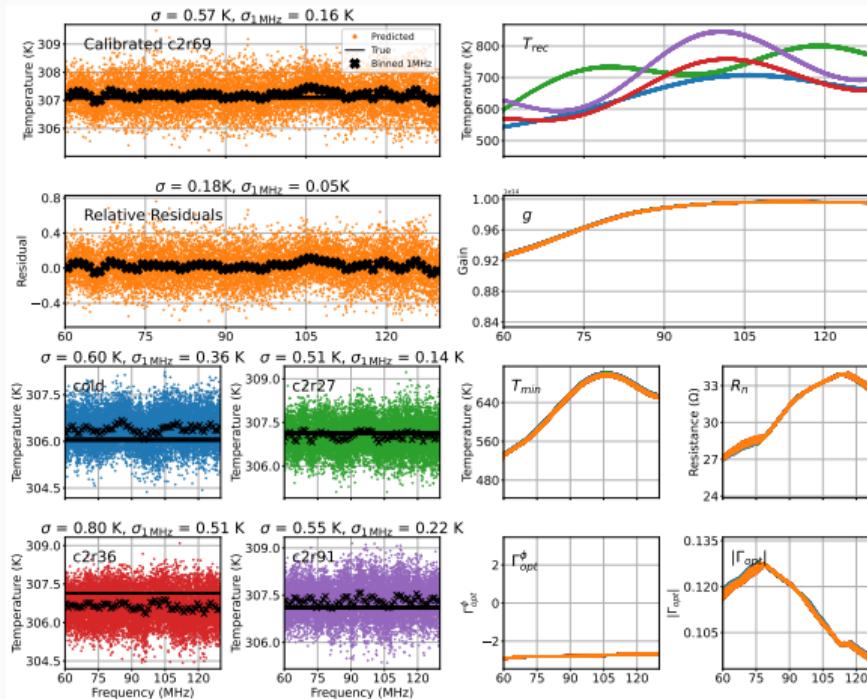
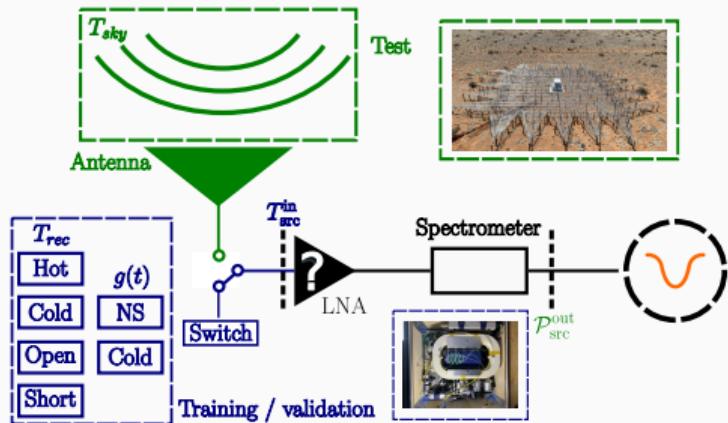
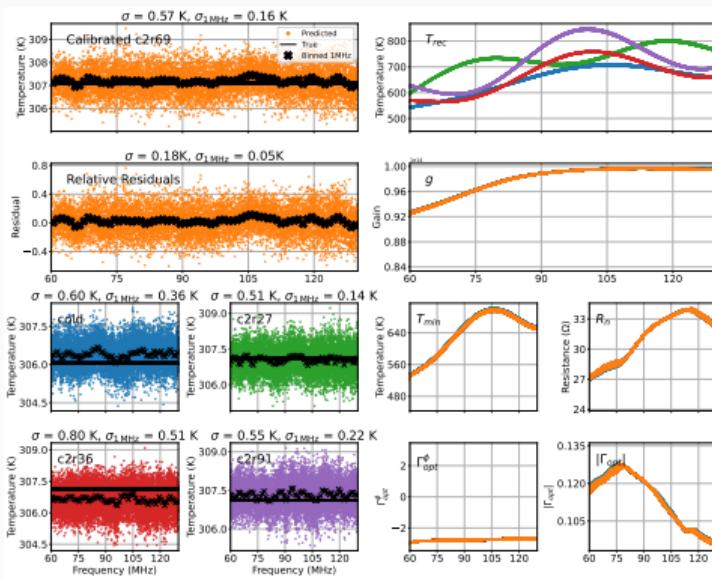


Figure 3: Temperature calibration using internal sources on the REACH receiver

Does not require cable/switch corrections?



End to end 21cm experiment simulation

Simulation pipeline for radiometer calibration

 **calibration-simulation** Private

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 .gitignore	cleanup	4 months ago
 CLAUDE.md	Add is_spectra feature for preserving frequency-dependent ...	2 hours ago
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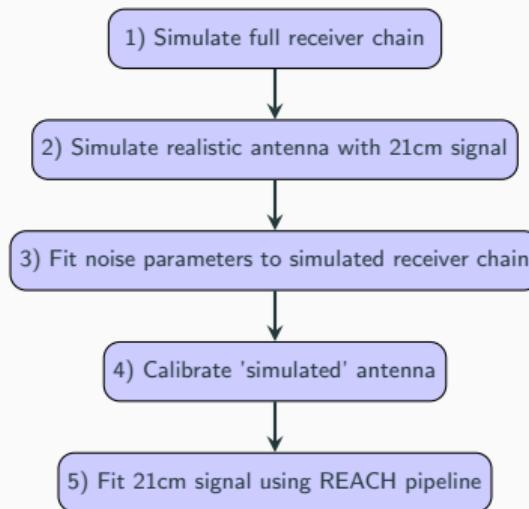
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End to end simulation

We have shown we can calibrate an internal source, we now test the method on as part of the broader system (simulated).



Predicted vs True Antenna Temperature

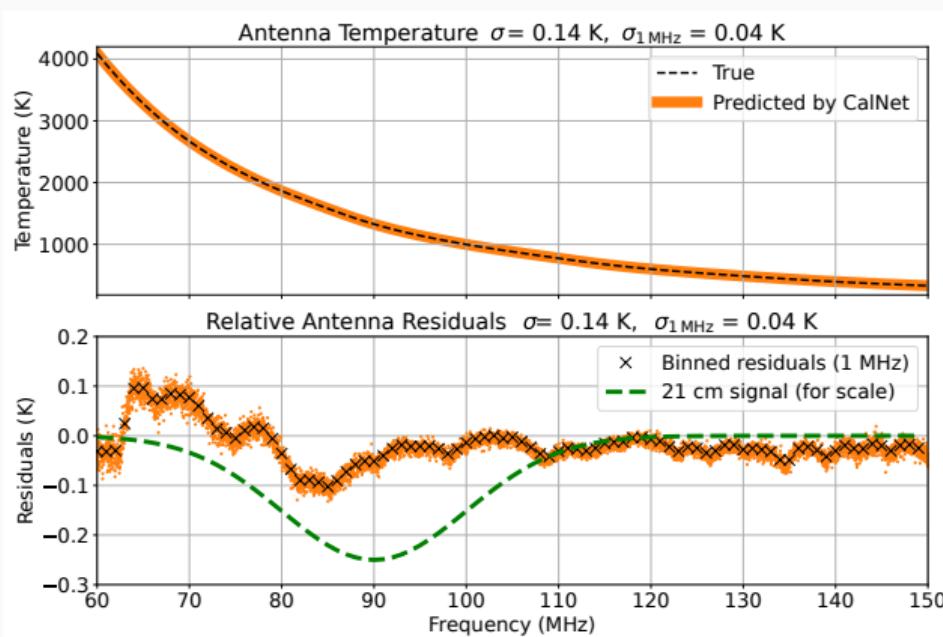


Figure 4: The top panel shows the predicted antenna temperature in orange, with the true temperature overlaid in black dashes

Inferred 21cm Signal

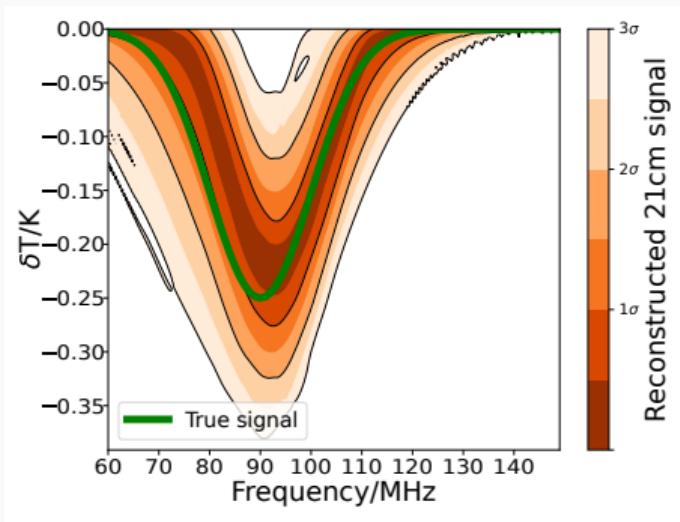


Figure 5: Left: Reconstructed 21cm signal with posterior (orange) and true signal (green)

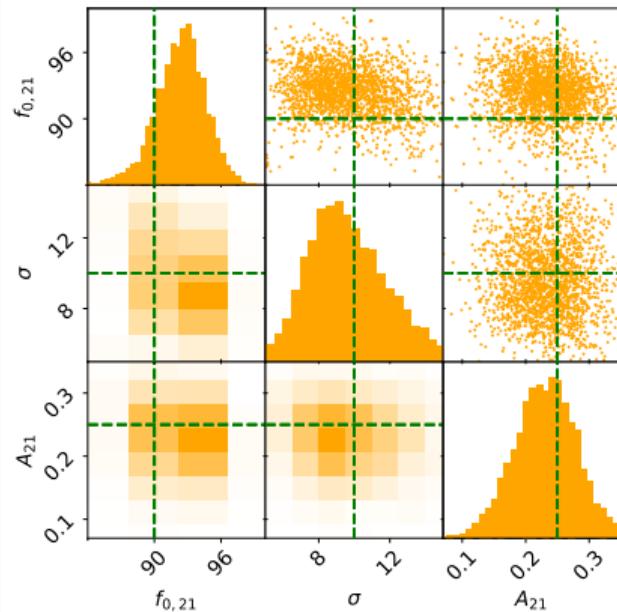
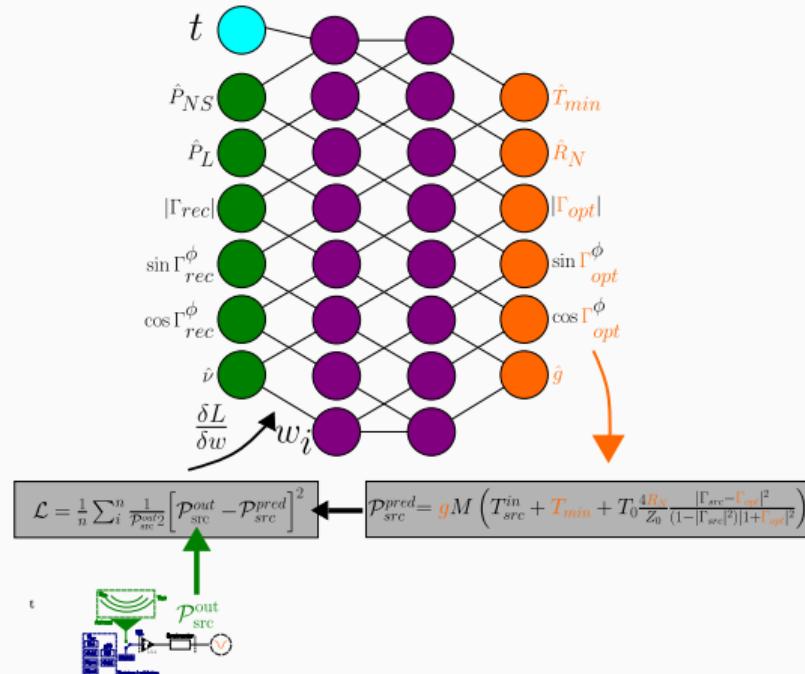


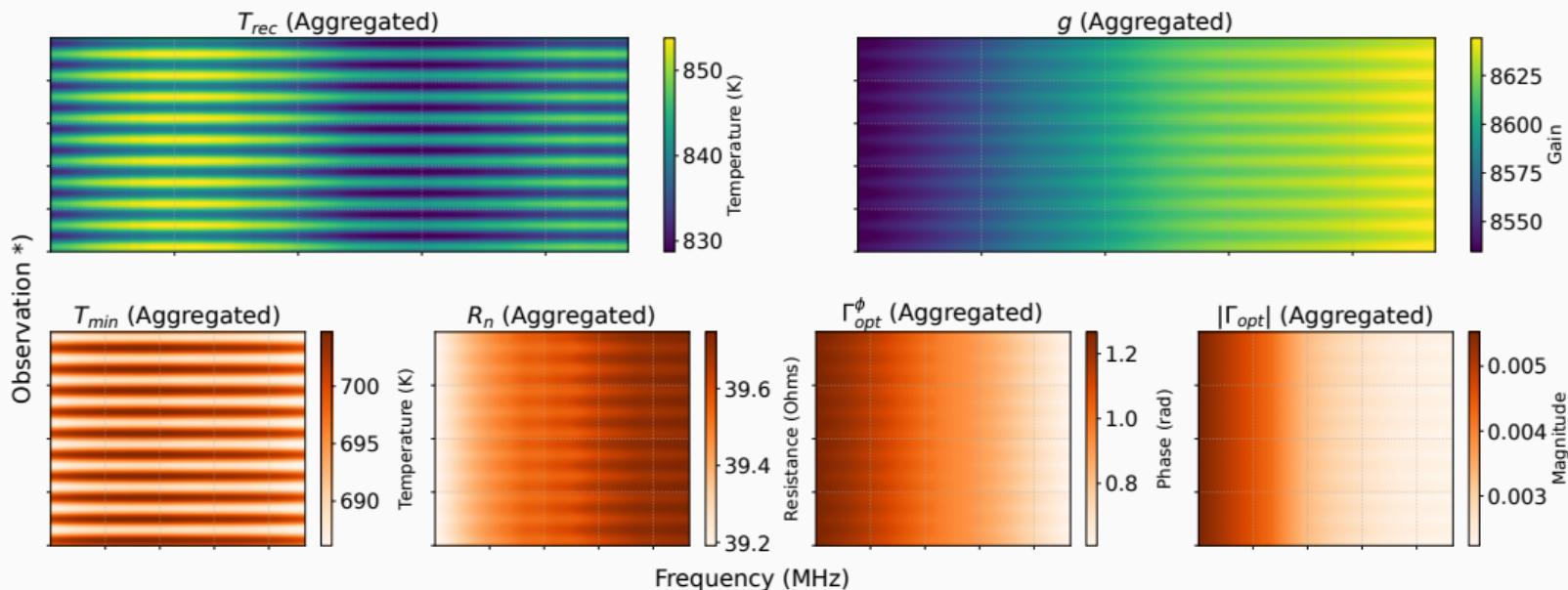
Figure 6: Right: Posterior plots of recovered 21cm signal parameters

Learning non-linear time-dependent system drift

Neural Network Time Evolution

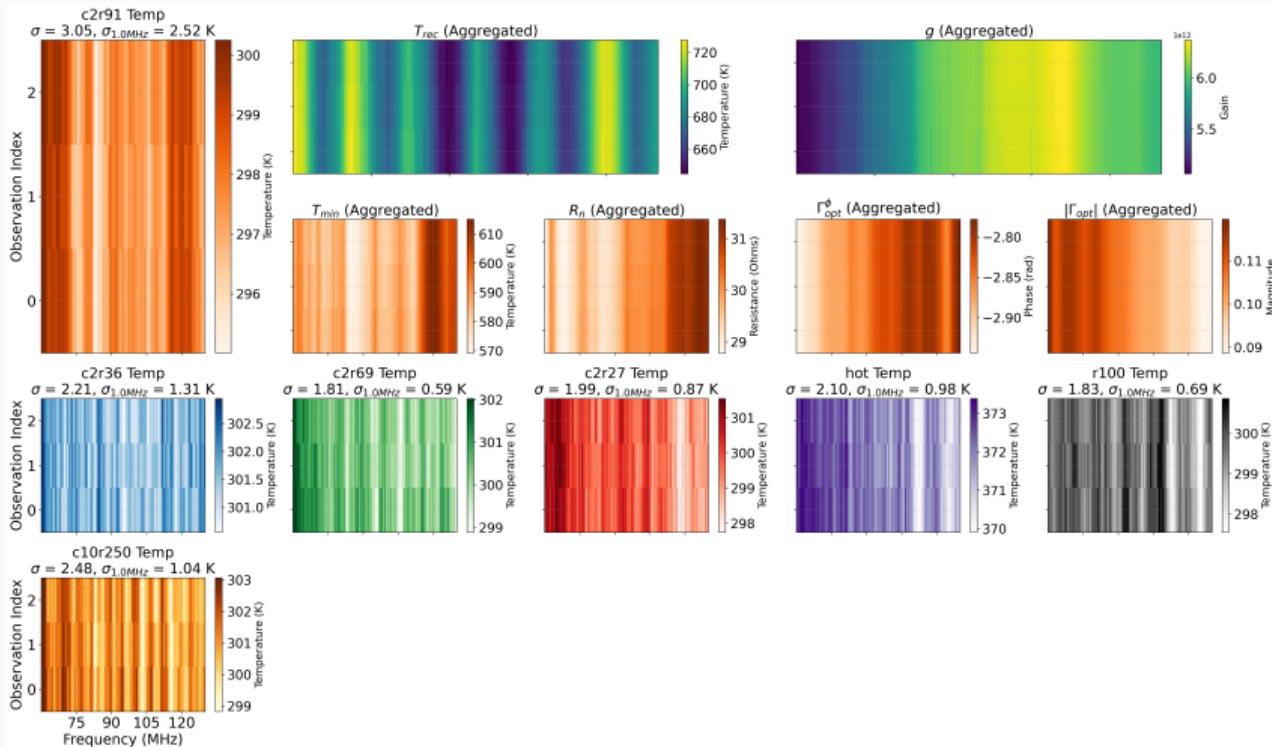


Inject system drift

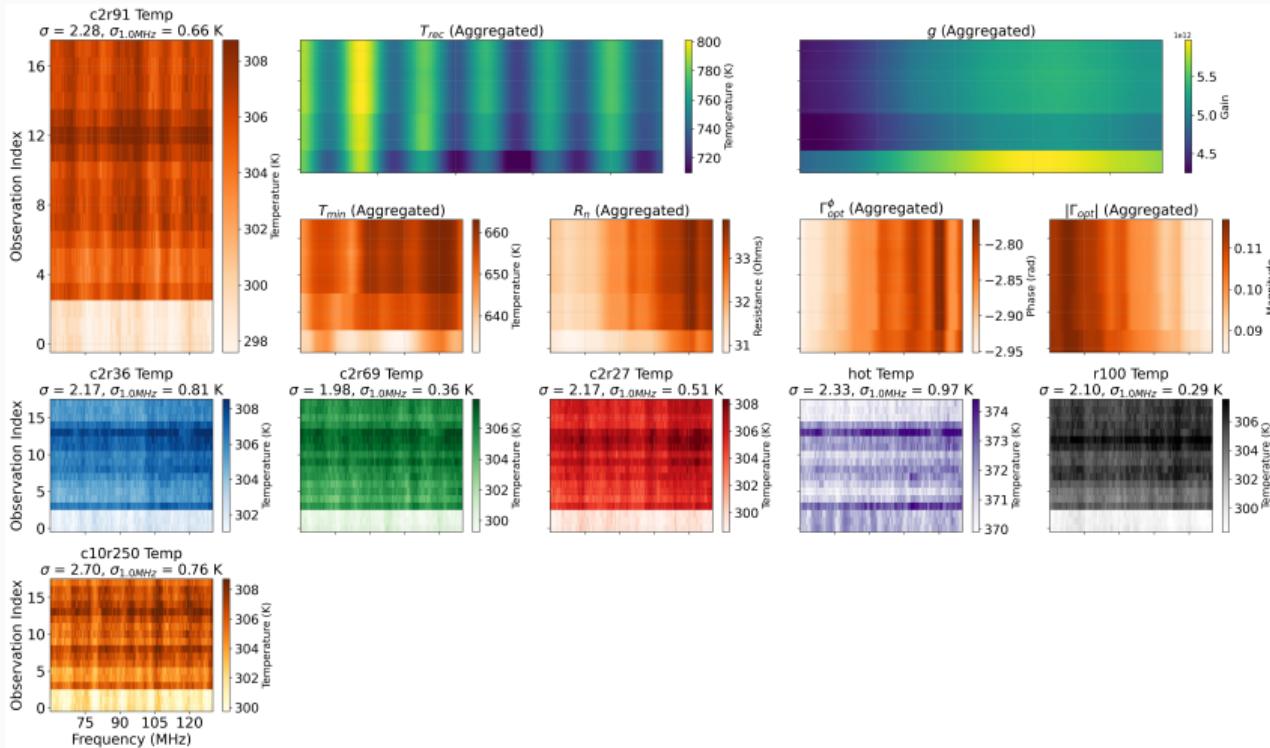


Inject a time varying sinusoid into T_{min} and predict the noise parameters → the network recovers this 'system drift'

Single night, many observations



Many nights, many observations



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[Submitted on 23 Apr 2025]

Radiometer Calibration using Machine Learning

S. A. K. Leeney, H. T. J. Bevins, E. de Lera Acedo, W. J. Handley, C. Kirkham, R. S. Patel, J. Zhu, D. Molnar, J. Cumner, D. Anstey, K. Artuc, G. Bernardi, M. Bucher, S. Carey, J. Cavillot, R. Chiello, W. Croukamp, D. I. L. de Villiers, J. A. Ely, A. Fialkov, T. Gessey-Jones, G. Kulkarni, A. Magro, P. D. Meerburg, S. Mittal, J. H. N. Pattison, S. Pegwal, C. M. Pieterse, J. R. Pritchard, E. Puchwein, N. Razavi-Ghods, I. L. V. Roque, A. Saxena, K. H. Scheutwinkel, P. Scott, E. Shen, P. H. Sims, M. Spinelli

Radiometers are crucial instruments in radio astronomy, forming the primary component of nearly all radio telescopes. They measure the intensity of electromagnetic radiation, converting this radiation into electrical signals. A radiometer's primary components are an antenna and a Low Noise Amplifier (LNA), which is the core of the "receiver" chain. Instrumental effects introduced by the receiver are typically corrected or removed during calibration. However, impedance mismatches between the antenna and receiver can introduce unwanted signal reflections and distortions. Traditional calibration methods, such as Dicke switching, alternate the receiver input between the antenna and a well-characterised reference source to mitigate errors by comparison. Recent advances in Machine Learning (ML) offer promising alternatives. Neural networks, which are trained using known signal sources, provide a powerful means to model and calibrate complex systems where traditional analytical approaches struggle. These methods are especially relevant for detecting the faint sky-averaged 21-cm signal from atomic hydrogen at high redshifts. This is one of the main challenges in observational Cosmology today. Here, for the first time, we introduce and test a machine learning-based calibration framework capable of achieving the precision required for radiometric experiments aiming to detect the 21-cm line.

Comments: Under peer review for publication in Nature Scientific Reports as part of the Radio Astronomy collection

Subjects: Instrumentation and Methods for Astrophysics (astro-ph.IM); Cosmology and Nongalactic Astrophysics (astro-ph.CO); Artificial Intelligence (cs.AI)

Cite as: arXiv:2504.16791 [astro-ph.IM]

(or arXiv:2504.16791v1 [astro-ph.IM] for this version)

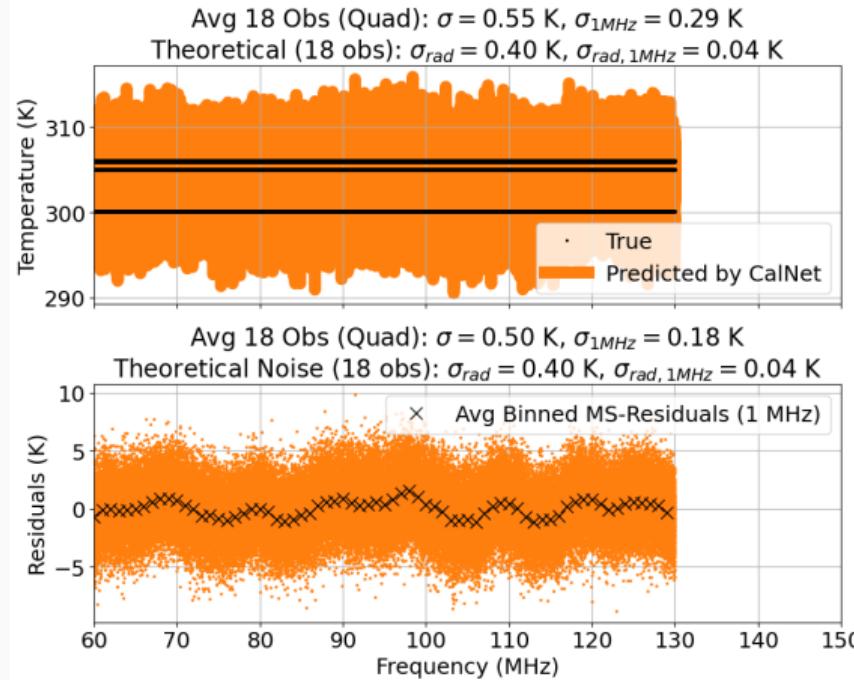
<https://doi.org/10.48550/arXiv.2504.16791> 

Submission history

From: Samuel Alan Kossoff Leeney [[view email](#)]

[v1] Wed, 23 Apr 2025 15:10:25 UTC (25,900 KB)

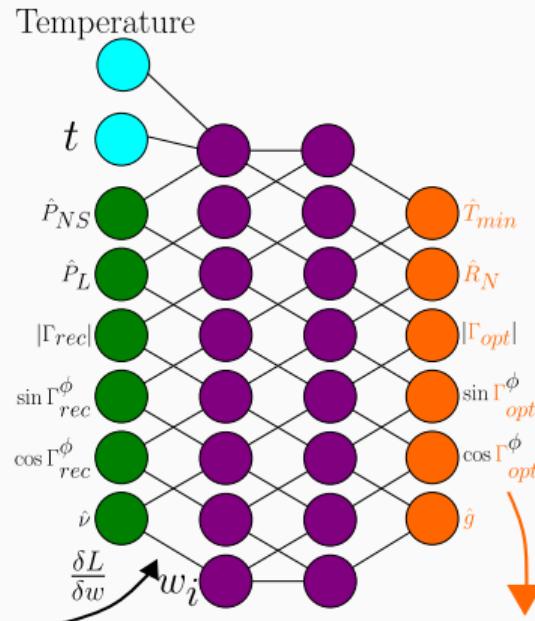
Calibration Performance



- Significant improvement when combining many nights
- Calibration down to 0.18K
- Getting close to theoretical noise

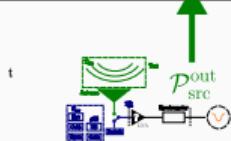
Modeling system drift from other complex system effects

Modeling environmental features



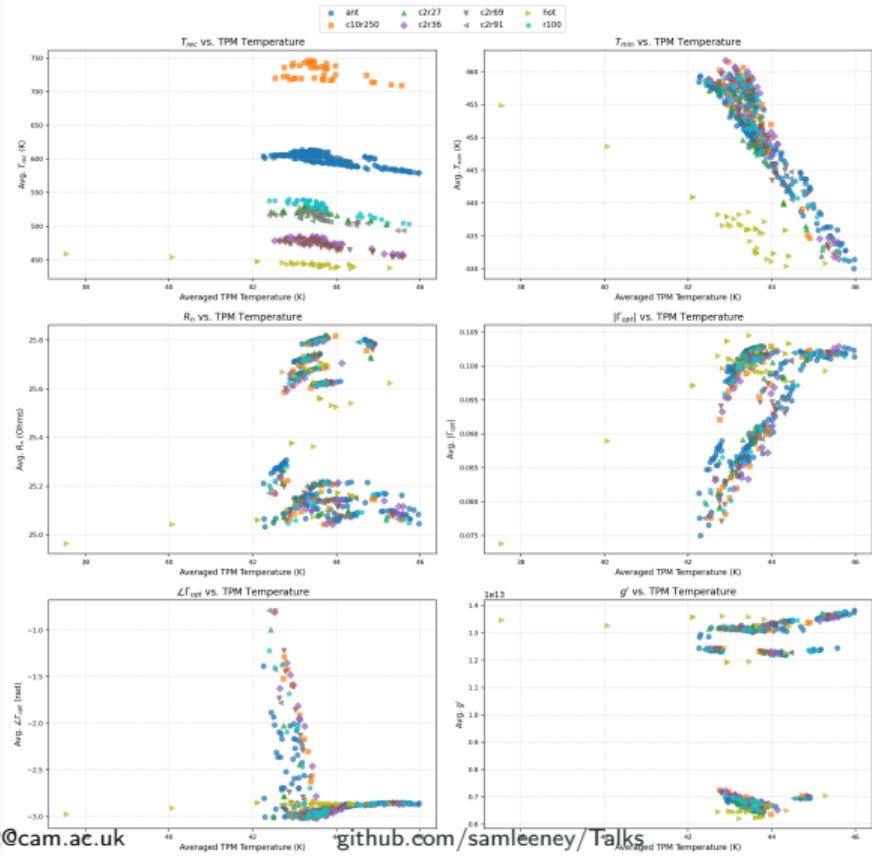
We can train on features that cannot be modeled analytically

$$\mathcal{L} = \frac{1}{n} \sum_i^n \frac{1}{\mathcal{P}_{src}^{out2}} \left[\mathcal{P}_{src}^{out} - \mathcal{P}_{src}^{pred} \right]^2 \quad \leftarrow \quad \mathcal{P}_{src}^{pred} = gM \left(T_{src}^{in} + T_{min} + T_0 \frac{4R_N}{Z_0} \frac{|\Gamma_{src} - \Gamma_{opt}|^2}{(1 - |\Gamma_{src}|^2)(1 + |\Gamma_{opt}|^2)} \right)$$



Fitting environmental features against fitted noise parameters

Correlate noise parameters with environmental data



Thank you!



SCAN ME

Appendix

Calibration Equation

Typically, substitute in the noise wave parameter equation here (gains cancel)

$$T_{\text{cal}}^* = T_{\text{NS}} \frac{P_{\text{cal}} - P_L}{P_{\text{NS}} - P_L} + T_L \quad (4)$$

Make some matching assumptions and re arrange:

$$\begin{aligned} T_s = & \color{red} T_{\text{NS}} \left(\frac{P_s - P_L}{P_{\text{NS}} - P_L} \right) \frac{|1 - \Gamma_s \Gamma_{\text{rec}}|^2}{1 - |\Gamma_s|^2} + \color{red} T_L \frac{|1 - \Gamma_s \Gamma_{\text{rec}}|^2}{1 - |\Gamma_s|^2} - \color{red} T_{\text{unc}} \frac{|\Gamma_s|^2}{1 - |\Gamma_s|^2} + \\ & - \color{red} T_{\text{cos}} \frac{\Re \left(\frac{\Gamma_s}{1 - \Gamma_s \Gamma_{\text{rec}}} \right) |1 - \Gamma_s \Gamma_{\text{rec}}|^2}{(1 - |\Gamma_s|^2) \sqrt{1 - |\Gamma_{\text{rec}}|^2}} - \color{red} T_{\text{sin}} \frac{\Im \left(\frac{\Gamma_s}{1 - \Gamma_s \Gamma_{\text{rec}}} \right) |1 - \Gamma_s \Gamma_{\text{rec}}|^2}{(1 - |\Gamma_s|^2) \sqrt{1 - |\Gamma_{\text{rec}}|^2}} \end{aligned} \quad (5)$$

Note: We end up with 5 parameters that need to be estimated to calibrate the system.

Calculating the error

By partial derivatives To find the error in T_s , we propagate the errors in Γ_s , Γ_{rec} , P_L , P_{NS} , and P_s :

$$(\Delta T_s)^2 = \left(\frac{\partial T_s}{\partial \Gamma_s} \Delta \Gamma_s \right)^2 + \left(\frac{\partial T_s}{\partial \Gamma_{rec}} \Delta \Gamma_{rec} \right)^2 + \left(\frac{\partial T_s}{\partial P_L} \Delta P_L \right)^2 + \left(\frac{\partial T_s}{\partial P_{NS}} \Delta P_{NS} \right)^2 + \left(\frac{\partial T_s}{\partial P_s} \Delta P_s \right)^2. \quad (10)$$

Calculating the error

$$(\Delta T_s)^2 = \left(\frac{\partial T_s}{\partial \Gamma_s} \Delta \Gamma_s \right)^2 + \left(\frac{\partial T_s}{\partial \Gamma_{rec}} \Delta \Gamma_{rec} \right)^2 + \left(\frac{\partial T_s}{\partial P_L} \Delta P_L \right)^2 + \left(\frac{\partial T_s}{\partial P_{NS}} \Delta P_{NS} \right)^2 + \left(\frac{\partial T_s}{\partial P_s} \Delta P_s \right)^2. \quad (11)$$

$$\frac{\partial T_s}{\partial P_L} = T_{NS} \frac{|1 - \Gamma_s \Gamma_{rec}|^2}{1 - |\Gamma_s|^2} \cdot \frac{P_s - P_{NS}}{(P_{NS} - P_L)^2}, \quad (12)$$

$$\frac{\partial T_s}{\partial P_{NS}} = -T_{NS} \frac{|1 - \Gamma_s \Gamma_{rec}|^2}{1 - |\Gamma_s|^2} \cdot \frac{P_s - P_L}{(P_{NS} - P_L)^2}, \quad (13)$$

$$\frac{\partial T_s}{\partial P_s} = T_{NS} \left(\frac{1}{P_{NS} - P_L} \right) \frac{|1 - \Gamma_s \Gamma_{rec}|^2}{1 - |\Gamma_s|^2}. \quad (14)$$

$$\frac{\partial T_s}{\partial \Gamma_{rec}} = \frac{\partial A}{\partial \Gamma_{rec}} + \frac{\partial B}{\partial \Gamma_{rec}} + \frac{\partial D}{\partial \Gamma_{rec}} + \frac{\partial E}{\partial \Gamma_{rec}}. \quad (15)$$

$$\frac{\partial T_s}{\partial \Gamma_s} = \frac{\partial A}{\partial \Gamma_s} + \frac{\partial B}{\partial \Gamma_s} + \frac{\partial C}{\partial \Gamma_s} + \frac{\partial D}{\partial \Gamma_s} + \frac{\partial E}{\partial \Gamma_s}. \quad (16)$$

$$A = T_{NS} \left(\frac{P_s - P_L}{P_{NS} - P_L} \right) \frac{|1 - \Gamma_s \Gamma_{rec}|^2}{1 - |\Gamma_s|^2}, \quad (17)$$

$$B = T_L \frac{|1 - \Gamma_s \Gamma_{rec}|^2}{1 - |\Gamma_s|^2}, \quad (18)$$

$$C = -T_{unc} \frac{|\Gamma_s|^2}{1 - |\Gamma_s|^2}, \quad (19)$$

$$D = -T_{cos} \frac{\Re \left(\frac{\Gamma_s}{1 - \Gamma_s \Gamma_{rec}} \right) |1 - \Gamma_s \Gamma_{rec}|^2}{(1 - |\Gamma_s|^2) \sqrt{1 - |\Gamma_{rec}|^2}}, \quad (20)$$

$$E = -T_{sin} \frac{\Im \left(\frac{\Gamma_s}{1 - \Gamma_s \Gamma_{rec}} \right) |1 - \Gamma_s \Gamma_{rec}|^2}{(1 - |\Gamma_s|^2) \sqrt{1 - |\Gamma_{rec}|^2}}. \quad (21)$$

Calculating the error

Using $T_{NS} \frac{P_{cal}-P_L}{P_{NS}-P_L} + T_L$

$$(\Delta T_s)^2 = \left(\frac{\partial T_s}{\partial P_s} \Delta P_s \right)^2 + \left(\frac{\partial T_s}{\partial P_L} \Delta P_L \right)^2 + \left(\frac{\partial T_s}{\partial P_{NS}} \Delta P_{NS} \right)^2 \quad (22)$$

$$+ \left(\frac{\partial T_s}{\partial \Gamma_s} \Delta \Gamma_s \right)^2 + \left(\frac{\partial T_s}{\partial \Gamma_{rec}} \Delta \Gamma_{rec} \right)^2. \quad (23)$$

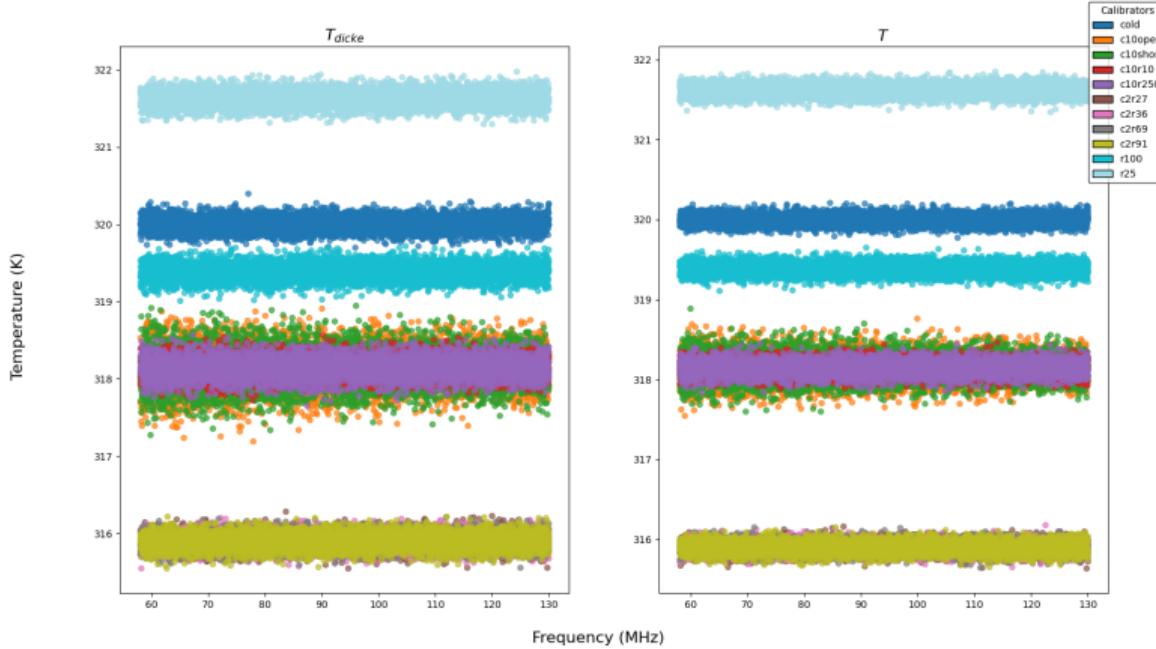
Using noise wave parameters only

$$(\Delta T_s)^2 = \left(\frac{\partial T_s}{\partial P_s} \Delta P_s \right)^2 \quad (24)$$

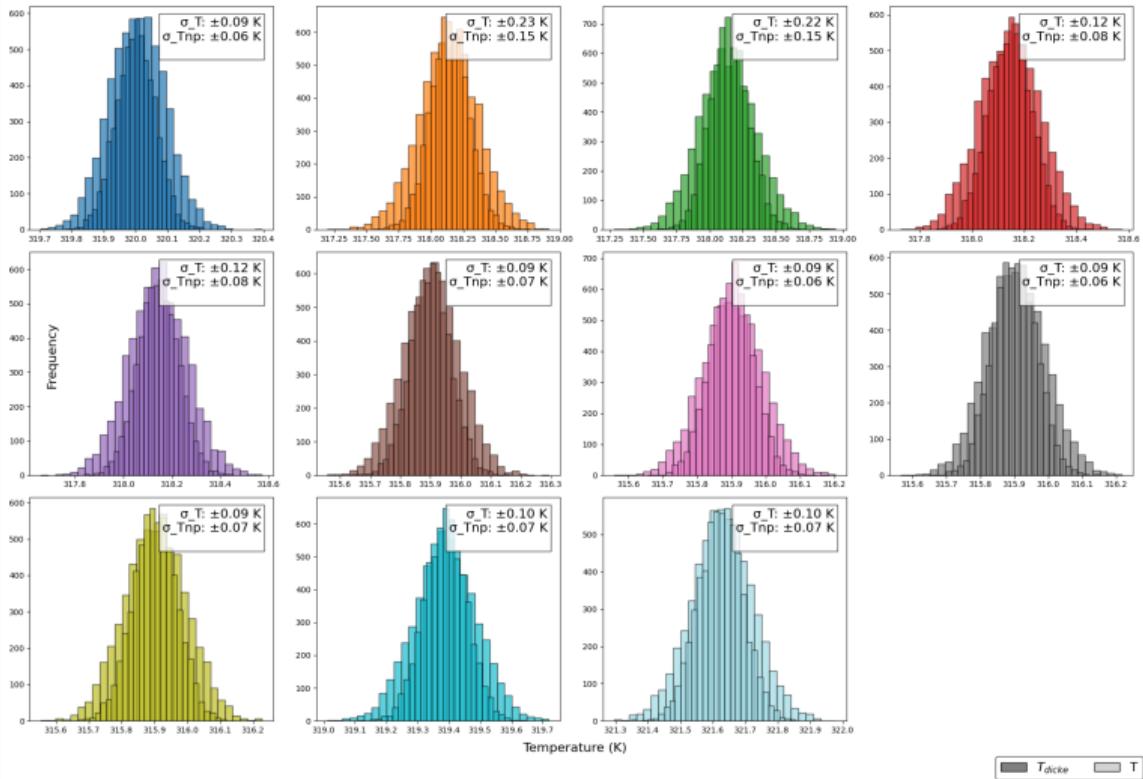
$$+ \left(\frac{\partial T_s}{\partial \Gamma_s} \Delta \Gamma_s \right)^2 + \left(\frac{\partial T_s}{\partial \Gamma_{rec}} \Delta \Gamma_{rec} \right)^2. \quad (25)$$

Note that this argument applies for both noise parameters and noise wave parameters.

There is more
noise when using
 $T_{\text{NS}} \frac{P_{\text{cal}} - P_L}{P_{\text{NS}} - P_L} + T_L$



Combined Histograms of T_{dicke} and T for Each Calibrator



Noise amplified
by **30%** when
using

$$T_{NS} \frac{P_{cal} - P_L}{P_{NS} - P_L} + T_L$$

Why not fit noise (wave) parameters directly?

Noise Parameter Equation:

$$P_{\text{out}}^{\text{src}} = \mathbf{g} M \left(T_{\text{in}}^{\text{src}} + \mathbf{T}_{\text{min}} + T_0 \frac{4R_N}{Z_0} \frac{|\Gamma_{\text{src}} - \Gamma_{\text{opt}}|^2}{(1 - |\Gamma_{\text{src}}|^2)(1 + |\Gamma_{\text{opt}}|^2)} \right) \quad (26)$$

Noise Wave Equation:

$$\begin{aligned} P_{\text{out}}^{\text{src}} = & \mathbf{g} \left[\mathbf{T}_0 + \mathbf{T}_{\text{unc}} |\Gamma_s|^2 \left| \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_s \Gamma_{\text{rec}}} \right|^2 \right. \\ & + T_s (1 - |\Gamma_s|^2) \left| \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_s \Gamma_{\text{rec}}} \right|^2 + \mathbf{T}_{\text{cos}} \Re \left(\Gamma_s \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_s \Gamma_{\text{rec}}} \right) \\ & \left. + \mathbf{T}_{\text{sin}} \Im \left(\Gamma_s \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_s \Gamma_{\text{rec}}} \right) \right] \end{aligned} \quad (27)$$

We still end up with 5 unknowns, as before.