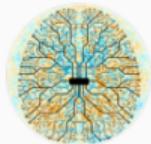


# Machine Learning for Radiometer Calibration

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2nd Year PhD Candidate

With: Harry Bevins, Eloy de Lera Acedo, Will Handley, Rohan Patel, Kaan Artuc, Jiacong Zhu



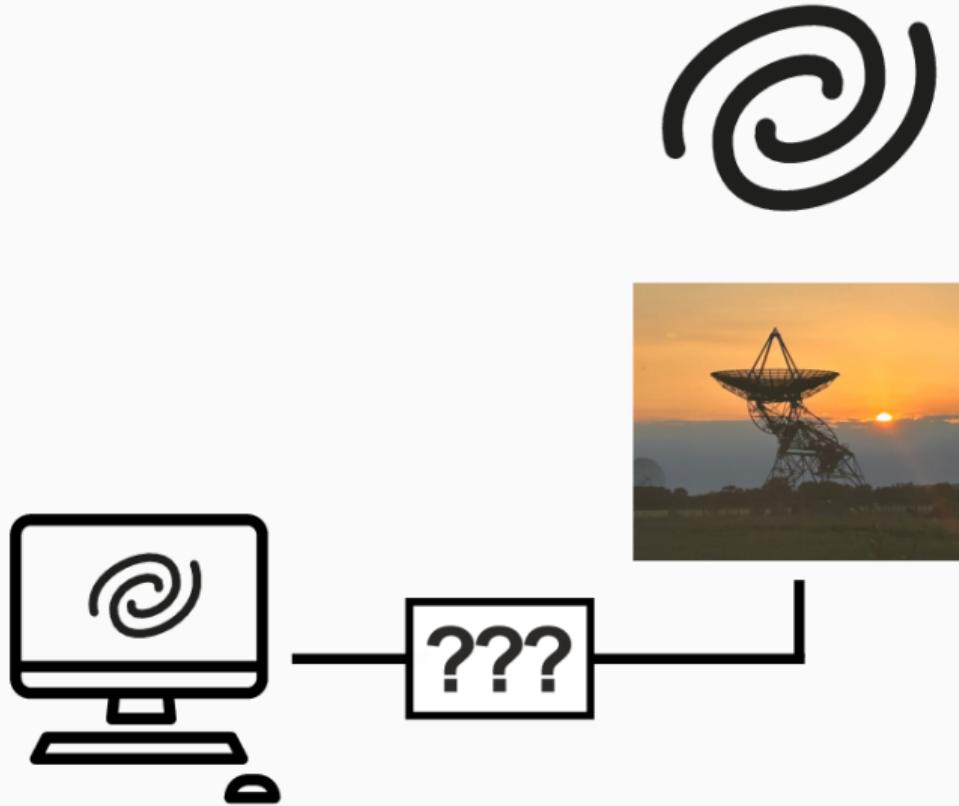
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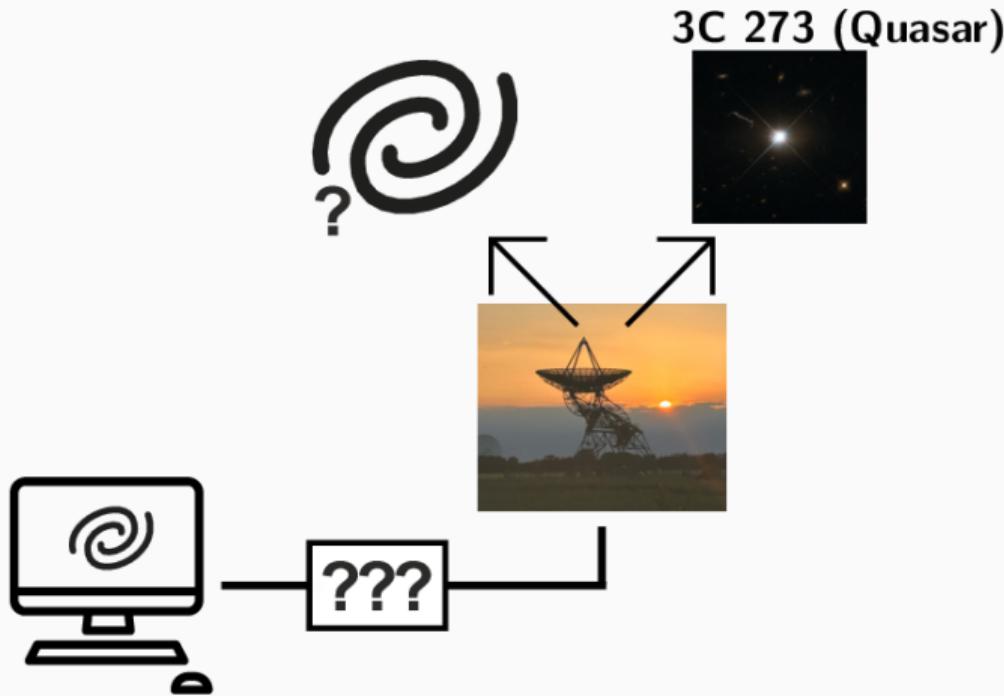
## **Radiometer calibration overview**

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# What is calibration?



## How to calibrate?



## Why is calibration in Global 21cm Cosmology difficult?



We measure *sky averaged* signal.

Antenna LNA impedance mismatch

Very faint signal.

## How to calibrate (in a bit more detail...)?

**Objective:** Map input temperature to output power.

### Key Factors:

- LNA introduces time-dependent gain,  $g(t)$ .
- Impedance mismatch adds noise ( $T_{\text{rec}}$ ) to the system.

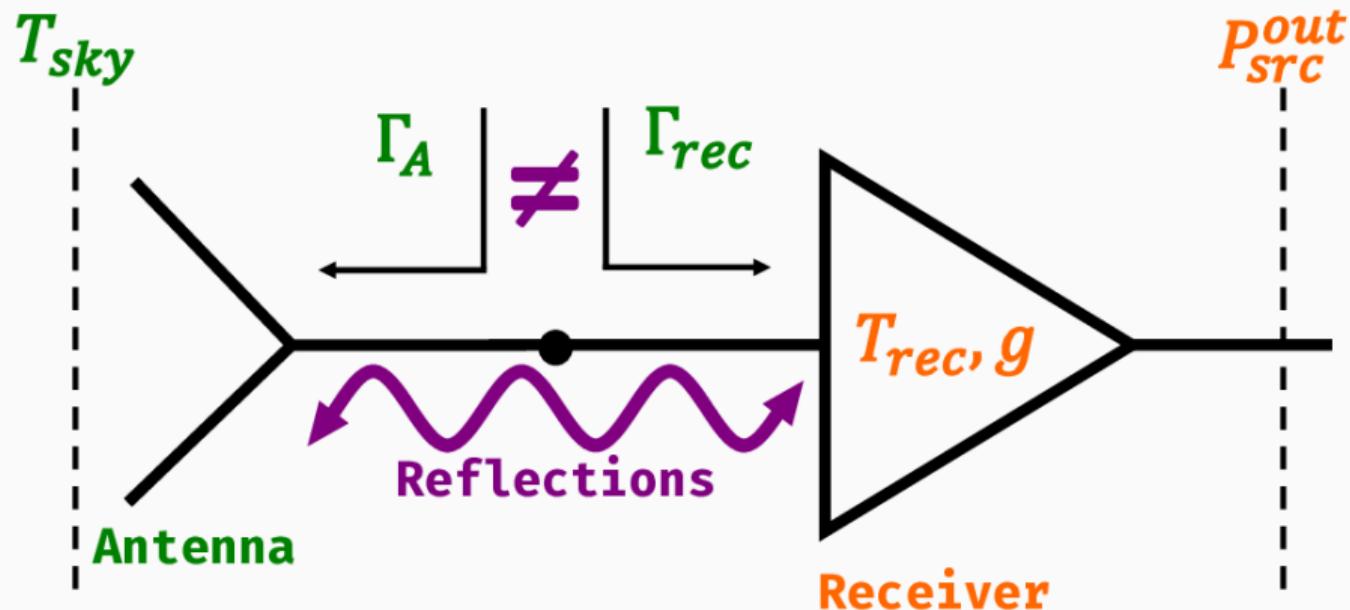
### Link Output Power to Input Temperature:

$$P_{\text{out}}^{\text{src}} = gM \times (T_{\text{in}}^{\text{src}} + T_{\text{rec}}) \quad (1)$$

$$M = \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{(1 - |\Gamma_{\text{cal}}|^2)} | \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_{\text{cal}} \Gamma_{\text{rec}}} |^2$$

*Note: All parameters above are frequency-dependent, but the notation has been simplified here and thereafter for convenience.*

## Receiver/source mismatch



**Figure 1:** Schematic of the receiver system showing the signal path from antenna to digitizer

## Dealing with reflections...

$$P_{\text{out}}^{\text{src}} = gM (T_{\text{in}}^{\text{src}} + T_{\text{rec}}) \quad (2)$$

### Noise Parameter Equation:

$$P_{\text{out}}^{\text{src}} = gM \left( T_{\text{in}}^{\text{src}} + T_{\text{min}} + T_0 \frac{4R_N}{Z_0} \frac{|\Gamma_{\text{src}} - \Gamma_{\text{opt}}|^2}{(1 - |\Gamma_{\text{src}}|^2)(1 + |\Gamma_{\text{opt}}|^2)} \right) \quad (3)$$

### Noise Wave Equation:

$$\begin{aligned} P_{\text{out}}^{\text{src}} = & g \left[ T_0 + T_{\text{unc}} |\Gamma_s|^2 \left| \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_s \Gamma_{\text{rec}}} \right|^2 \right. \\ & + T_s (1 - |\Gamma_s|^2) \left| \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_s \Gamma_{\text{rec}}} \right|^2 + T_{\cos} \Re \left( \Gamma_s \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_s \Gamma_{\text{rec}}} \right) \\ & \left. + T_{\sin} \Im \left( \Gamma_s \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_s \Gamma_{\text{rec}}} \right) \right] \quad (4) \end{aligned}$$

- $g$  - Gain
- $T_{\text{min}}$  - Minimum Noise Temperature
- $R_N$  - Noise Resistance
- $\Gamma_{\text{opt}}$  - Optimum Reflection Coefficient
- $\Gamma_s$  - Source Reflection Coefficient
- $\Gamma_{\text{rec}}$  - Receiver Reflection Coefficient
- $T_s$  - Source Temperature
- $P_{\text{out}}^{\text{src}}$  - Power out
- $T_{\text{unc}}, \cos, \sin$  - Noise wave parameters
- $T_0$  reference temperature

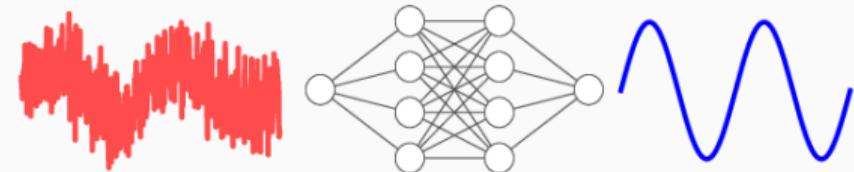
# **Radiometer calibration with machine learning**

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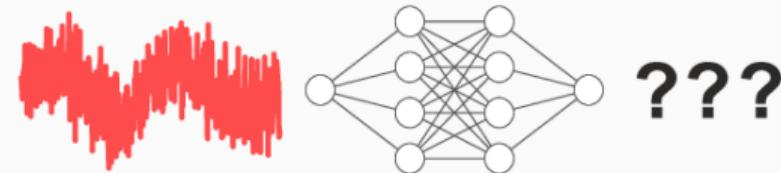
## How/why?

- Can predict noise wave parameters using machine learning.
- Could apply method to noise parameters.
- Predict **directly** on noise parameters on frequency by frequency basis.
- Maleable to environmental changes.

Train



Predict



# Machine learning calibration steps

## 1. Define the Loss Function

Regress over measured power and predicted power.

$$\mathcal{L} = \frac{1}{n} \sum_{i=1}^n (\mathcal{P}_{\text{measured},i} - \mathcal{P}_{\text{pred},i})^2 \quad (5)$$

## 2. Write Down the Equation for $\mathcal{P}_{\text{pred}}$

Using the noise wave formalism, relate  $\mathcal{P}_{\text{pred}}$  to  $T_{\text{src}}$ .

$$\begin{aligned} \mathcal{P}_{\text{pred}} &= \mathbf{g} \cdot M(T_{\text{in}}^{\text{src}} \\ &+ T_{\text{min}} + T_0 \frac{4R_N}{Z_0} \frac{|\Gamma_{\text{src}} - \Gamma_{\text{opt}}|^2}{(1 - |\Gamma_{\text{src}}|^2)(1 + |\Gamma_{\text{opt}}|^2)} \Big) \end{aligned} \quad (6)$$

## 3. Minimise the Loss Function

$$\theta^* = \arg \min_{\theta} \mathcal{L}(\theta) \quad (7)$$

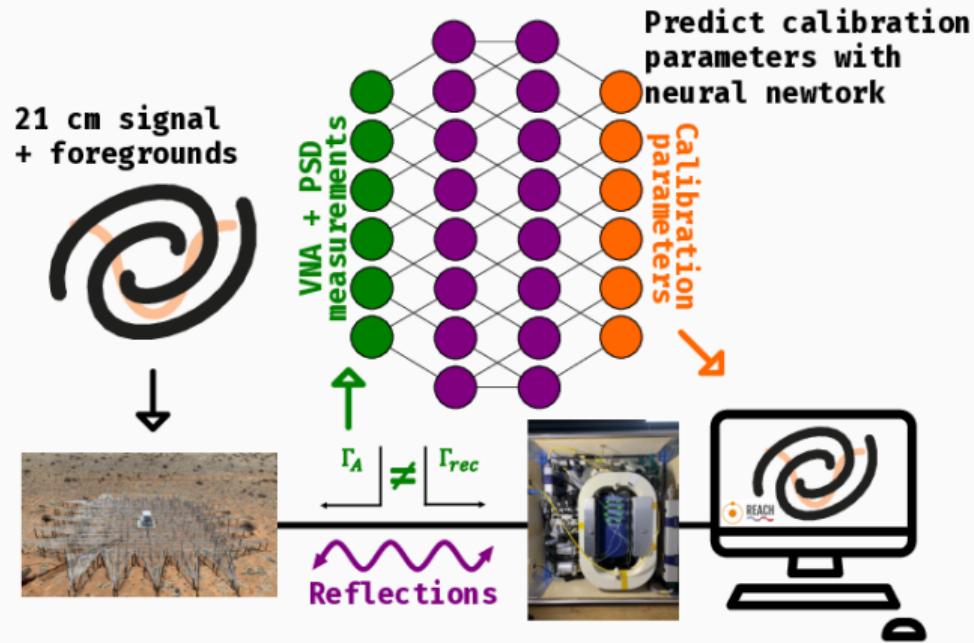
parameter vector  $\theta$  includes all tunable parameters in the model:

$$\theta = \{g, T_{\text{min}}, R_N, \Gamma_{\text{opt}}^\phi, |\Gamma_{\text{opt}}|\} \quad (8)$$

## 4. Rearrange and Predict ( $T_{\text{src}}$ ) using $\theta^*$

$$\begin{aligned} T_{\text{src}} &= \frac{\mathcal{P}_{\text{pred}}}{g^* \cdot M} \\ &- \left( T_{\text{min}}^* + T_0 \frac{4R_N^*}{Z_0} \frac{|\Gamma_{\text{src}} - \Gamma_{\text{opt}}^*|^2}{(1 - |\Gamma_{\text{src}}|^2)(1 + |\Gamma_{\text{opt}}^*|^2)} \right) \end{aligned} \quad (9)$$

## ML calibration overview

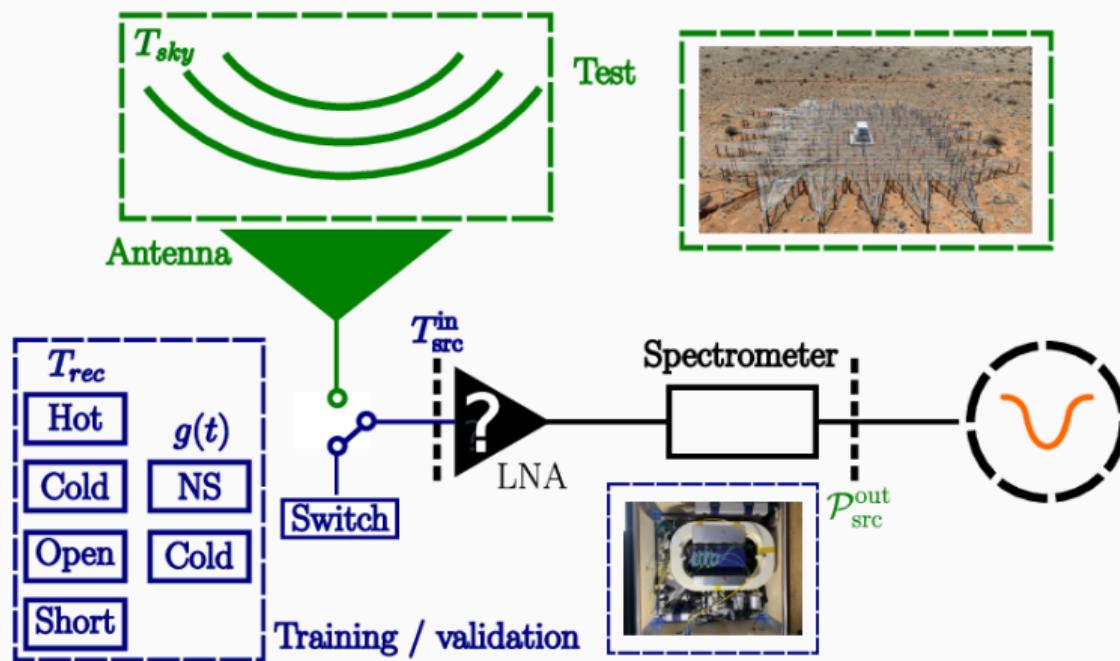


**Figure 2:** High-level overview of the machine learning-based calibration framework

# The physical system

## Process:

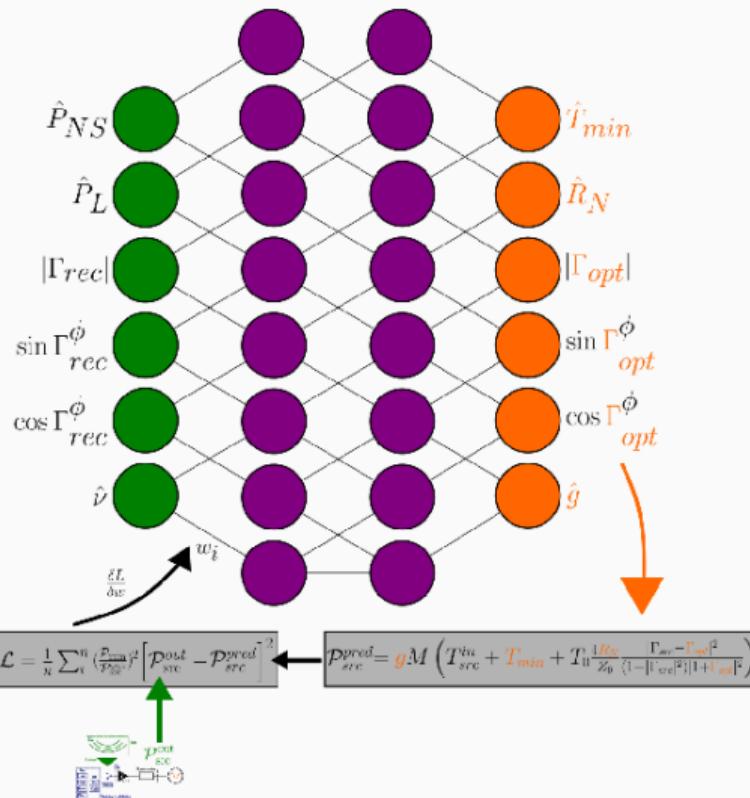
- Switch between sources to generate training data.
- Calibration sources with known temperature train neural net.
- Predict  $T_{src}$  of antenna.



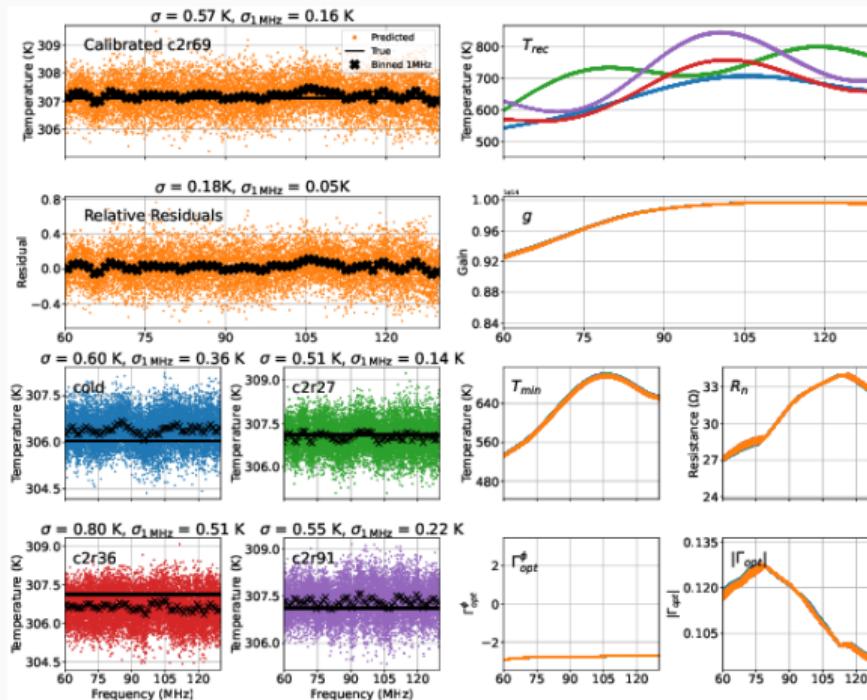
# Network Architecture

## Structure:

- Input thermistor and VNA measurements.
- Also input frequency.
- Predict noise parameters.
- Regress over loss function.

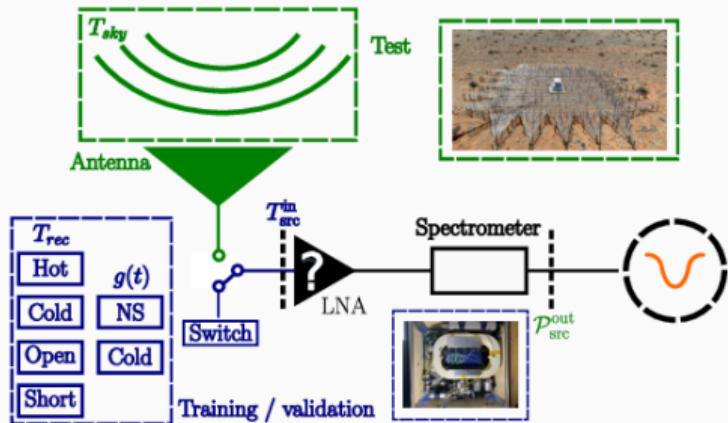
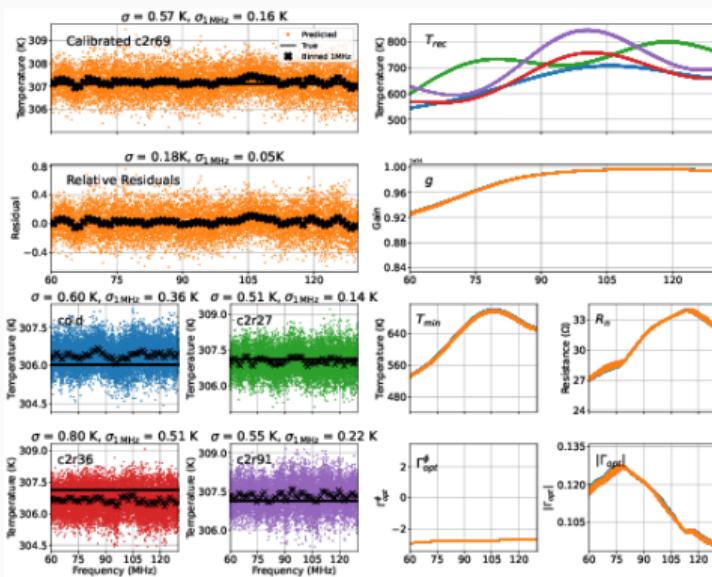


## Testing on internal validation source



**Figure 3:** Temperature calibration using internal sources on the REACH receiver

# Does not require cable/switch corrections?



## **End to end 21cm experiment simulation**

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# Simulation pipeline for radiometer calibration

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Calibration observation simulation codes for REACH

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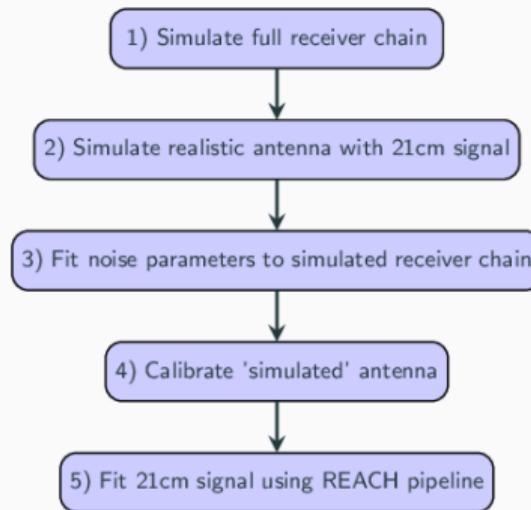
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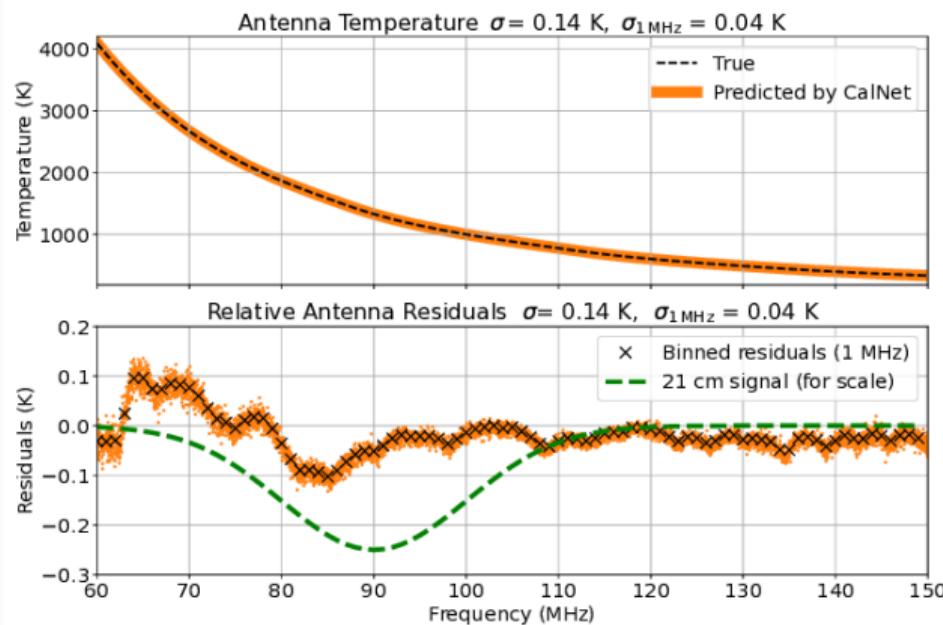
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## End to end simulation

We have shown we can calibrate an internal source, we now test the method on as part of the broader system (simulated).

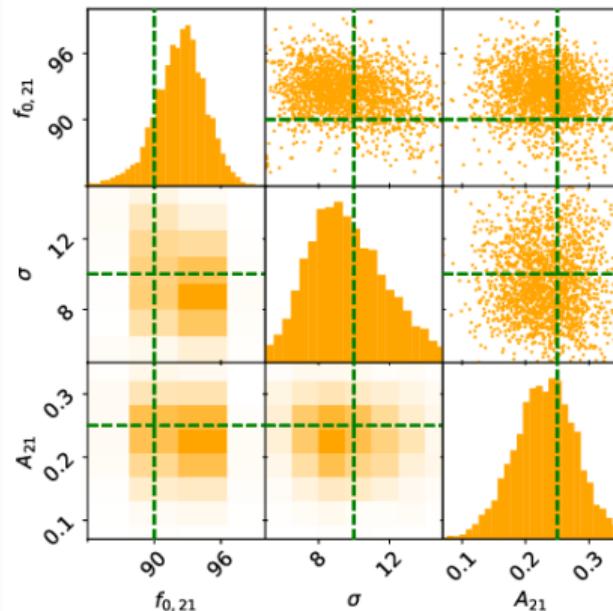
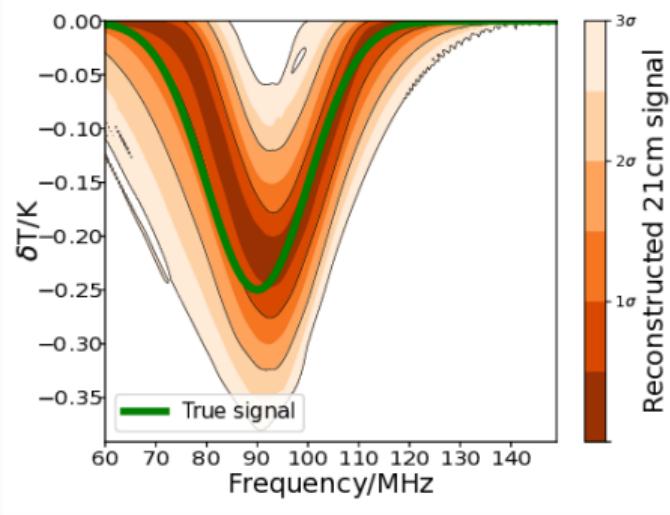


## Predicted vs True Antenna Temperature



**Figure 4:** The top panel shows the predicted antenna temperature in orange, with the true temperature overlaid in black dashes

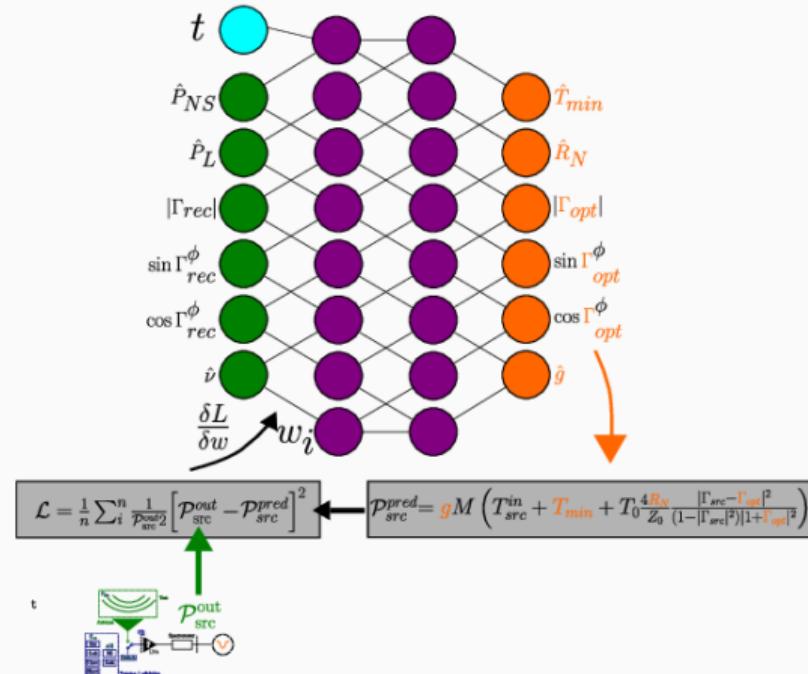
## Inferred 21cm Signal



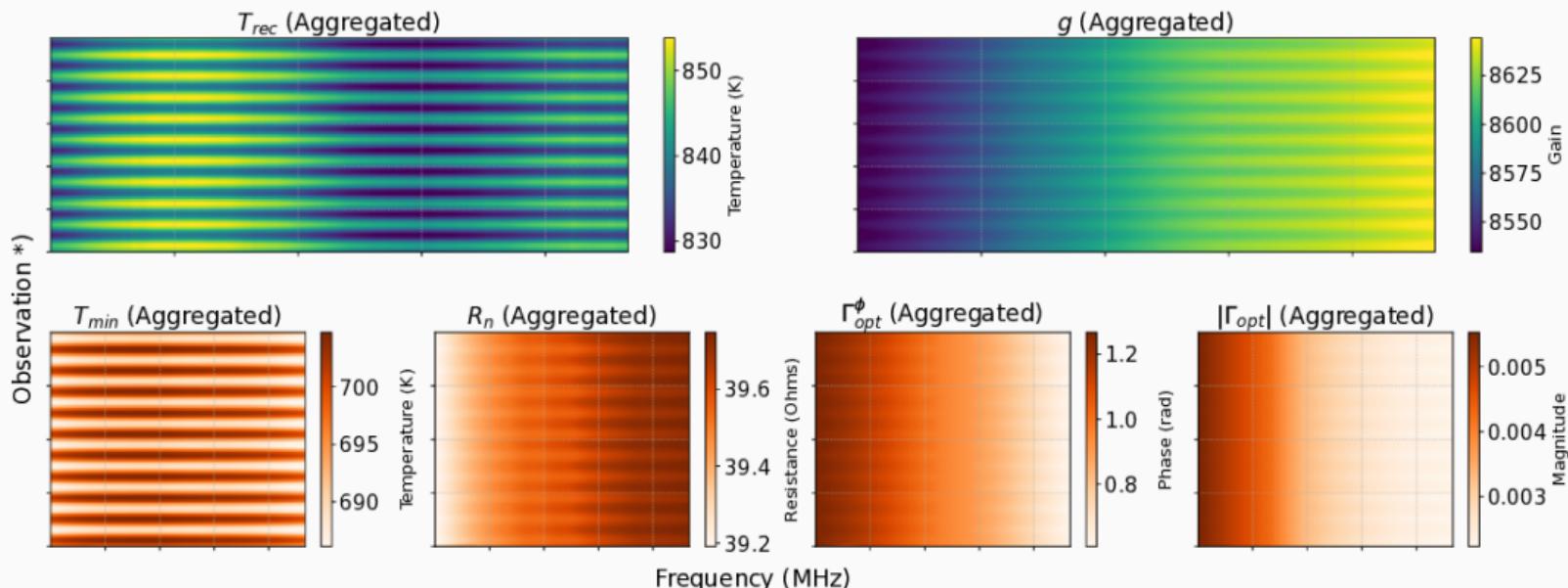
## **Learning non-linear time-dependent system drift**

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# Neural Network Time Evolution

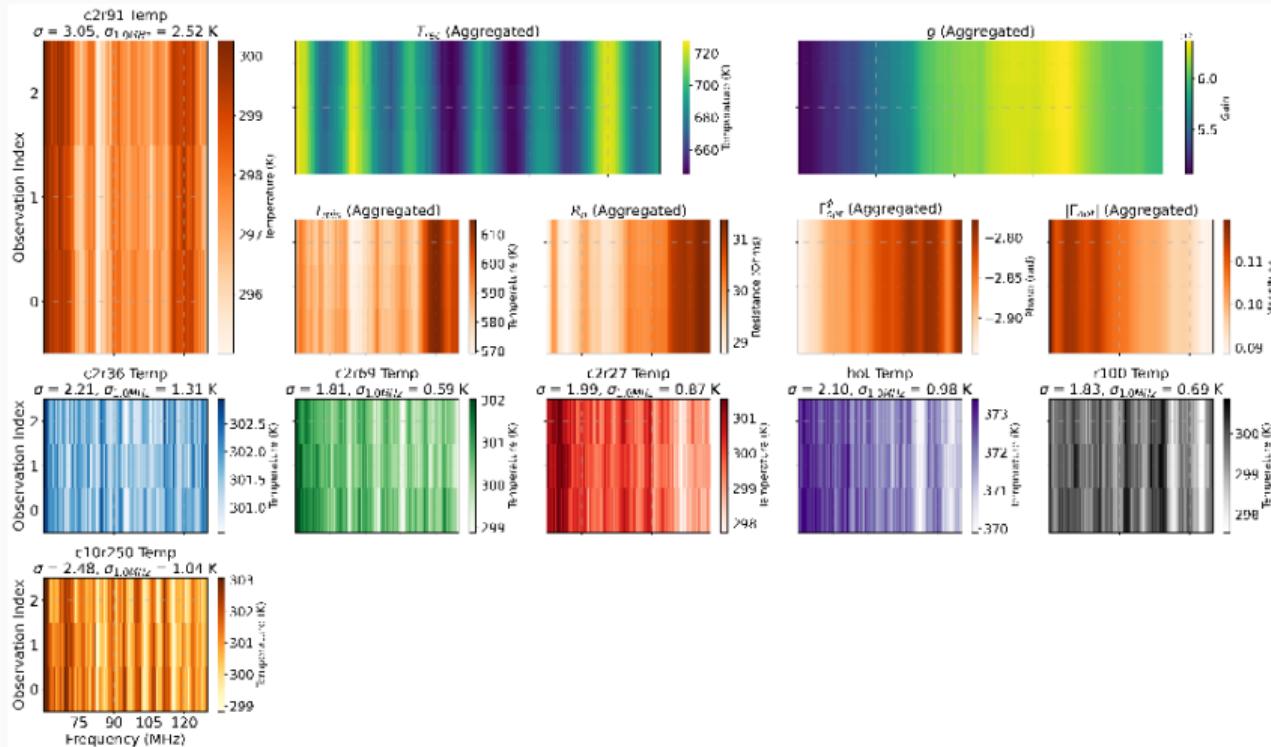


## Inject system drift

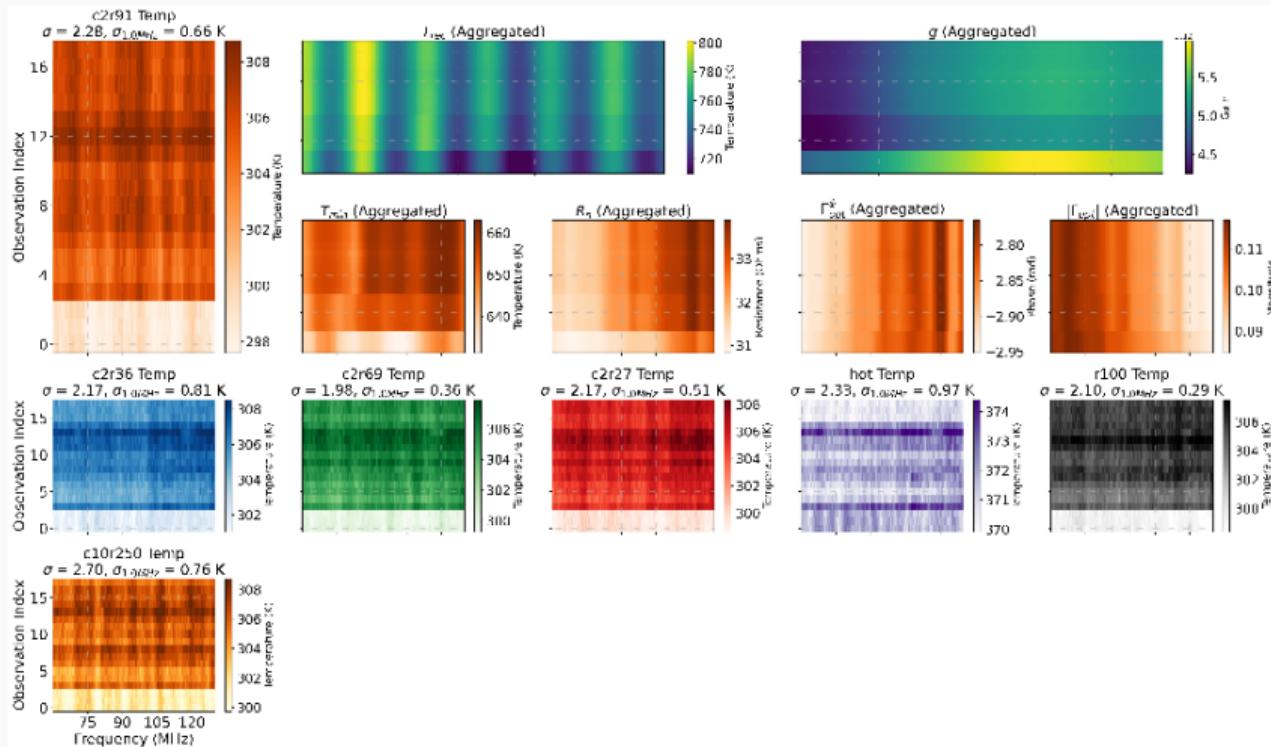


Inject a time varying sinusoid into  $T_{min}$  and predict the noise parameters → the network recovers this 'system drift'

# Single night, many observations



# Many nights, many observations



# Read the paper...

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(Submitted on 27 Apr 2025)

## Radiometer Calibration using Machine Learning

S. A. K. Leeney, H. T. J. Bovens, E. de Lera Acedo, W. J. Hardley, C. Kirkham, R. S. Patel, J. Zhu, D. Molnar, J. Currier, D. Ansley, K. Ariuc, G. Bernardi, M. Bucher, S. Carey, J. Cavillot, R. Chiole, W. Croukamp, D. E. L. de Villiers, J. A. Ely, A. Flalkov, T. Gesssey-Jones, G. Kulkarni, A. Magro, P. D. Meerburg, S. Millat, J. H. N. Pallison, S. Pegwal, C. M. Pieterse, J. R. Pritchard, E. Puchwein, N. Razavi Ghods, I. L. V. Roque, A. Saxena, K. H. Scheuwinkel, P. Scott, E. Shen, P. H. Sims, M. Spinelli

Radiometers are crucial instruments in radio astronomy, forming the primary component of nearly all radio telescopes. They measure the intensity of electromagnetic radiation, converting this radiation into electrical signals. A radiometer's primary components are an antenna and a Low Noise Amplifier (LNA), which is the core of the "receiver" chain. Instrumental effects introduced by the receiver are typically corrected or removed during calibration. However, impedance mismatches between the antenna and receiver can introduce unwanted signal reflections and distortions. Traditional calibration methods, such as Dicke switching, alternate the receiver input between the antenna and a well-characterised reference source to mitigate errors by comparison. Recent advances in Machine Learning (ML) offer promising alternatives. Neural networks, which are trained using known signal sources, provide a powerful means to model and calibrate complex systems where traditional analytical approaches struggle. These methods are especially relevant for detecting the faint sky-averaged 21-cm signal from atomic hydrogen at high redshifts. This is one of the main challenges in observational Cosmology today. Here, for the first time, we introduce and test a machine learning-based calibration framework capable of achieving the precision required for radiometric experiments along to detect the 21-cm line.

Comments: Under peer review for publication in Nature Scientific Reports as part of the Radio Astronomy collection

Subjects: Instrumentation and Methods for Astrophysics [astro-ph.IM]; Cosmology and Nongalactic Astrophysics [astro-ph.CO]; Artificial Intelligence [eess.AI]

Cite as: arXiv:2504.15791 [astro-ph.IM]

(or arXiv:2504.15791v1 [astro-ph.IM] for this version)

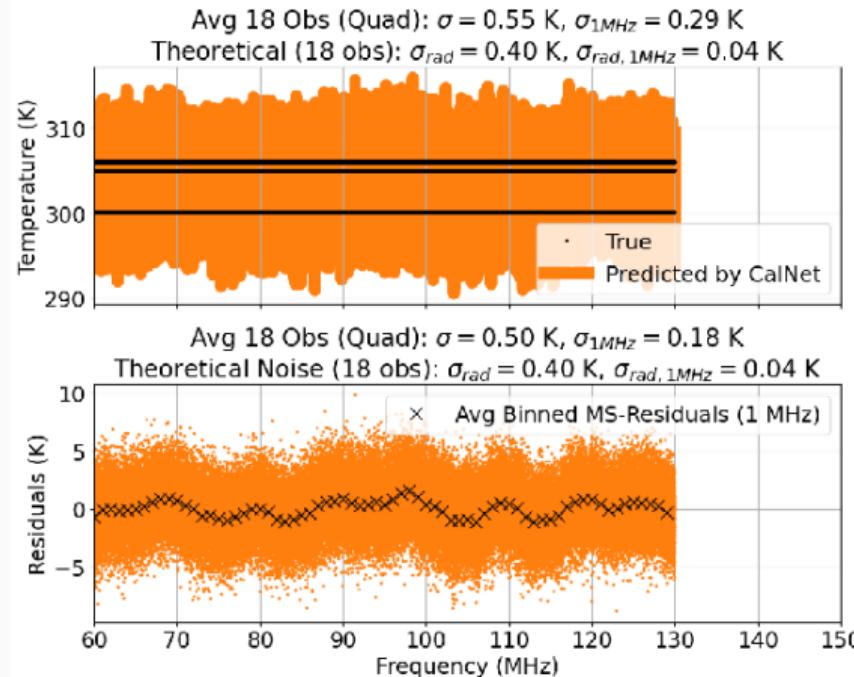
<https://doi.org/10.48550/arXiv.2504.15791> 

### Submission history

From: Samuel Alan Kassell Leeney [view email]

[v1] Wed, 27 Apr 2025 15:10:25 UTC (25,900 KB)

## Calibration Performance

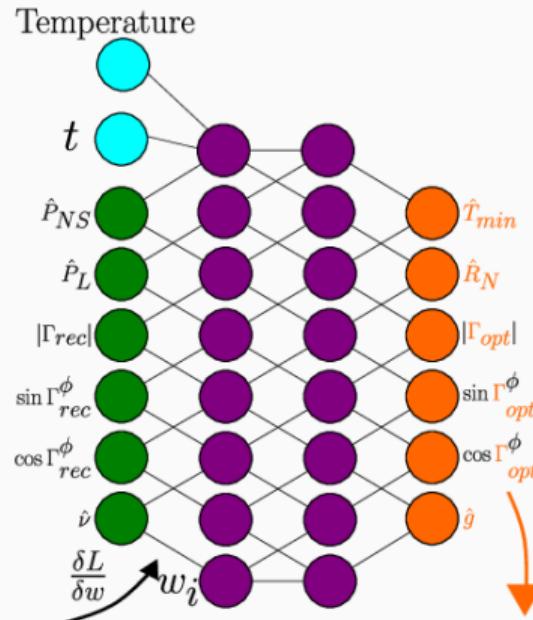


- Significant improvement when combining many nights
- Calibration down to 0.18K
- Getting close to theoretical noise

## **Modeling system drift from other complex system effects**

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# Modeling environmental features



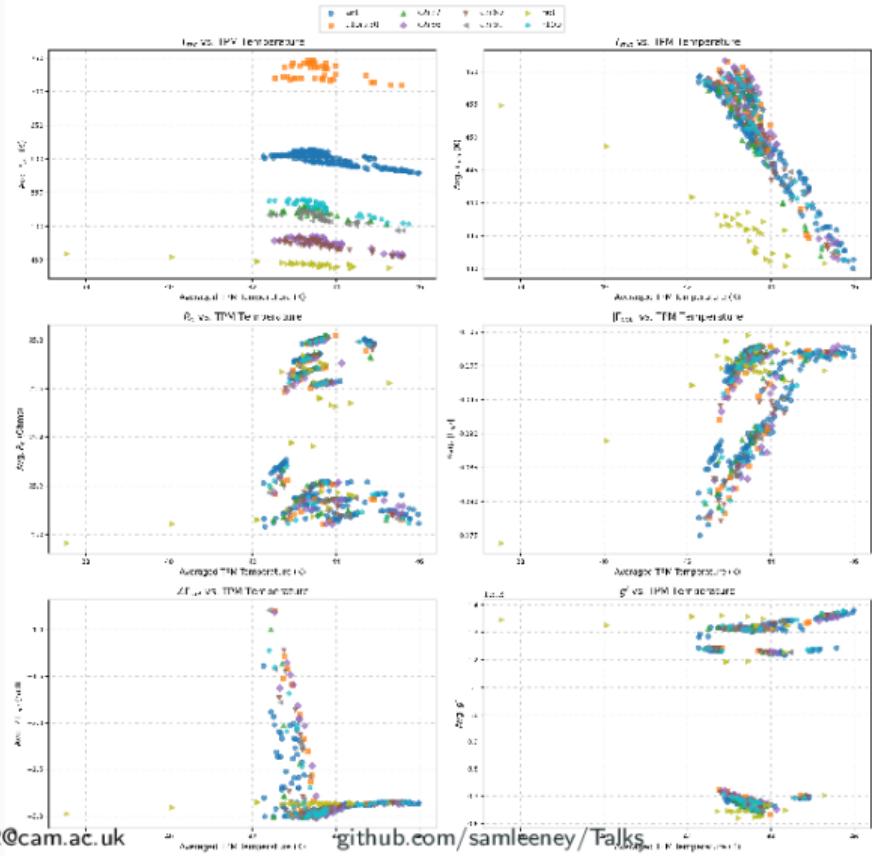
$$\mathcal{L} = \frac{1}{n} \sum_i^n \frac{1}{\mathcal{P}_{src}^{out2}} \left[ \mathcal{P}_{src}^{out} - \mathcal{P}_{src}^{pred} \right]^2 \quad \leftarrow \quad \mathcal{P}_{src}^{pred} = gM \left( T_{src}^{in} + T_{min} + T_0 \frac{4R_N}{Z_0} \frac{|\Gamma_{src} - \Gamma_{opt}|^2}{(1 - |\Gamma_{src}|^2)(1 + |\Gamma_{opt}|^2)} \right)$$



We can train on features that cannot be modeled analytically

# Fitting environmental features against fitted noise parameters

Correlate noise parameters with environmental data



Thank you!



SCAN ME

## Appendix

---

## Calibration Equation

Typically, substitute in the noise wave parameter equation here (gains cancel)

$$T_{\text{cal}}^* = T_{\text{NS}} \frac{P_{\text{cal}} - P_L}{P_{\text{NS}} - P_L} + T_L \quad (4)$$

Make some matching assumptions and re arrange:

$$\begin{aligned} T_s = & T_{\text{NS}} \left( \frac{P_s - P_L}{P_{\text{NS}} - P_L} \right) \frac{|1 - \Gamma_s \Gamma_{\text{rec}}|^2}{1 - |\Gamma_s|^2} + T_L \frac{|1 - \Gamma_s \Gamma_{\text{rec}}|^2}{1 - |\Gamma_s|^2} - T_{\text{unc}} \frac{|\Gamma_s|^2}{1 - |\Gamma_s|^2} + \\ & - T_{\text{cos}} \frac{\Re \left( \frac{\Gamma_s}{1 - \Gamma_s \Gamma_{\text{rec}}} \right) |1 - \Gamma_s \Gamma_{\text{rec}}|^2}{(1 - |\Gamma_s|^2) \sqrt{1 - |\Gamma_{\text{rec}}|^2}} - T_{\text{sin}} \frac{\Im \left( \frac{\Gamma_s}{1 - \Gamma_s \Gamma_{\text{rec}}} \right) |1 - \Gamma_s \Gamma_{\text{rec}}|^2}{(1 - |\Gamma_s|^2) \sqrt{1 - |\Gamma_{\text{rec}}|^2}} \end{aligned} \quad (5)$$

Note: We end up with 5 parameters that need to be estimated to calibrate the system.

## Calculating the error

**By partial derivatives** To find the error in  $T_s$ , we propagate the errors in  $\Gamma_s$ ,  $\Gamma_{rec}$ ,  $P_L$ ,  $P_{NS}$ , and  $P_s$ :

$$(\Delta T_s)^2 = \left( \frac{\partial T_s}{\partial \Gamma_s} \Delta \Gamma_s \right)^2 + \left( \frac{\partial T_s}{\partial \Gamma_{rec}} \Delta \Gamma_{rec} \right)^2 + \left( \frac{\partial T_s}{\partial P_L} \Delta P_L \right)^2 + \left( \frac{\partial T_s}{\partial P_{NS}} \Delta P_{NS} \right)^2 + \left( \frac{\partial T_s}{\partial P_s} \Delta P_s \right)^2. \quad (10)$$

## Calculating the error

$$(\Delta T_s)^2 = \left( \frac{\partial T_s}{\partial \Gamma_s} \Delta \Gamma_s \right)^2 + \left( \frac{\partial T_s}{\partial \Gamma_{rec}} \Delta \Gamma_{rec} \right)^2 + \left( \frac{\partial T_s}{\partial P_L} \Delta P_L \right)^2 + \left( \frac{\partial T_s}{\partial P_{NS}} \Delta P_{NS} \right)^2 + \left( \frac{\partial T_s}{\partial P_s} \Delta P_s \right)^2. \quad (11)$$


---

$$\frac{\partial T_s}{\partial P_L} = T_{NS} \frac{|1 - \Gamma_s \Gamma_{rec}|^2}{1 - |\Gamma_s|^2} \cdot \frac{P_s - P_{NS}}{(P_{NS} - P_L)^2}, \quad (12)$$

$$\frac{\partial T_s}{\partial P_{NS}} = -T_{NS} \frac{|1 - \Gamma_s \Gamma_{rec}|^2}{1 - |\Gamma_s|^2} \cdot \frac{P_s - P_L}{(P_{NS} - P_L)^2}, \quad (13)$$

$$\frac{\partial T_s}{\partial P_s} = T_{NS} \left( \frac{1}{P_{NS} - P_L} \right) \frac{|1 - \Gamma_s \Gamma_{rec}|^2}{1 - |\Gamma_s|^2}. \quad (14)$$

$$\frac{\partial T_s}{\partial \Gamma_{rec}} = \frac{\partial A}{\partial \Gamma_{rec}} + \frac{\partial B}{\partial \Gamma_{rec}} + \frac{\partial D}{\partial \Gamma_{rec}} + \frac{\partial E}{\partial \Gamma_{rec}}. \quad (15)$$

$$\frac{\partial T_s}{\partial \Gamma_s} = \frac{\partial A}{\partial \Gamma_s} + \frac{\partial B}{\partial \Gamma_s} + \frac{\partial C}{\partial \Gamma_s} + \frac{\partial D}{\partial \Gamma_s} + \frac{\partial E}{\partial \Gamma_s}. \quad (16)$$

$$A = T_{NS} \left( \frac{P_s - P_L}{P_{NS} - P_L} \right) \frac{|1 - \Gamma_s \Gamma_{rec}|^2}{1 - |\Gamma_s|^2}, \quad (17)$$

$$B = T_L \frac{|1 - \Gamma_s \Gamma_{rec}|^2}{1 - |\Gamma_s|^2}, \quad (18)$$

$$C = -T_{unc} \frac{|\Gamma_s|^2}{1 - |\Gamma_s|^2}, \quad (19)$$

$$D = -T_{cos} \frac{\Re \left( \frac{\Gamma_s}{1 - \Gamma_s \Gamma_{rec}} \right) |1 - \Gamma_s \Gamma_{rec}|^2}{(1 - |\Gamma_s|^2) \sqrt{1 - |\Gamma_{rec}|^2}}, \quad (20)$$

$$E = -T_{sin} \frac{\Im \left( \frac{\Gamma_s}{1 - \Gamma_s \Gamma_{rec}} \right) |1 - \Gamma_s \Gamma_{rec}|^2}{(1 - |\Gamma_s|^2) \sqrt{1 - |\Gamma_{rec}|^2}}. \quad (21)$$

## Calculating the error

Using  $T_{NS} \frac{P_{cal}-P_L}{P_{NS}-P_L} + T_L$

$$(\Delta T_s)^2 = \left( \frac{\partial T_s}{\partial P_s} \Delta P_s \right)^2 + \left( \frac{\partial T_s}{\partial P_L} \Delta P_L \right)^2 + \left( \frac{\partial T_s}{\partial P_{NS}} \Delta P_{NS} \right)^2 \quad (22)$$

$$+ \left( \frac{\partial T_s}{\partial \Gamma_s} \Delta \Gamma_s \right)^2 + \left( \frac{\partial T_s}{\partial \Gamma_{rec}} \Delta \Gamma_{rec} \right)^2. \quad (23)$$

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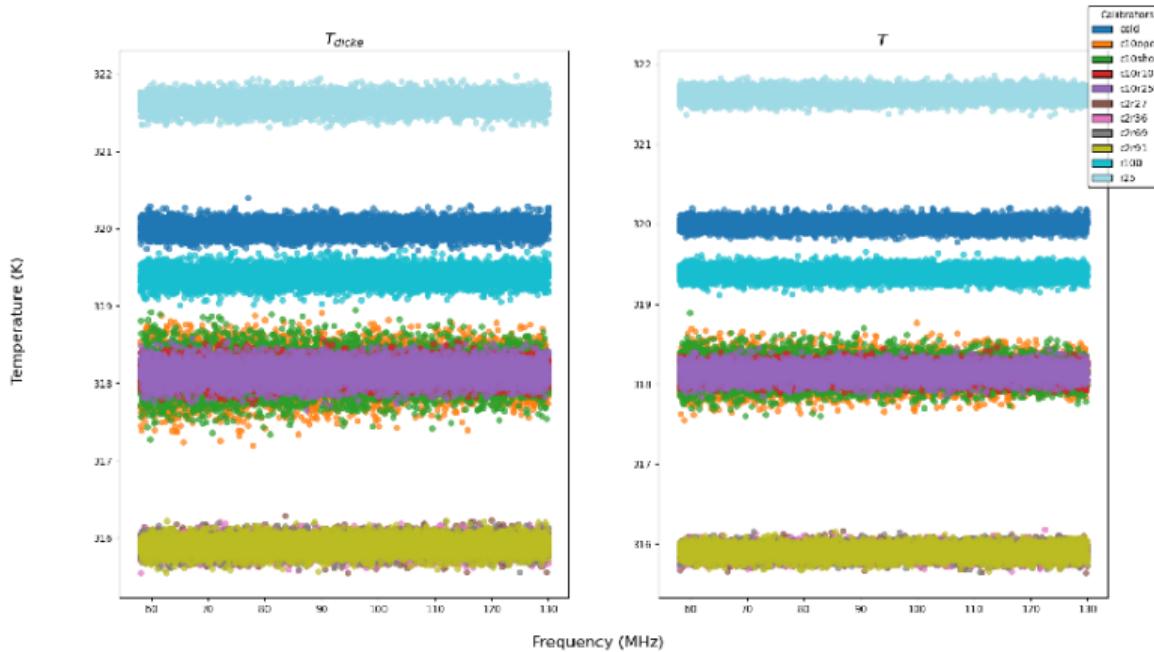
Using noise wave parameters only

$$(\Delta T_s)^2 = \left( \frac{\partial T_s}{\partial P_s} \Delta P_s \right)^2 \quad (24)$$

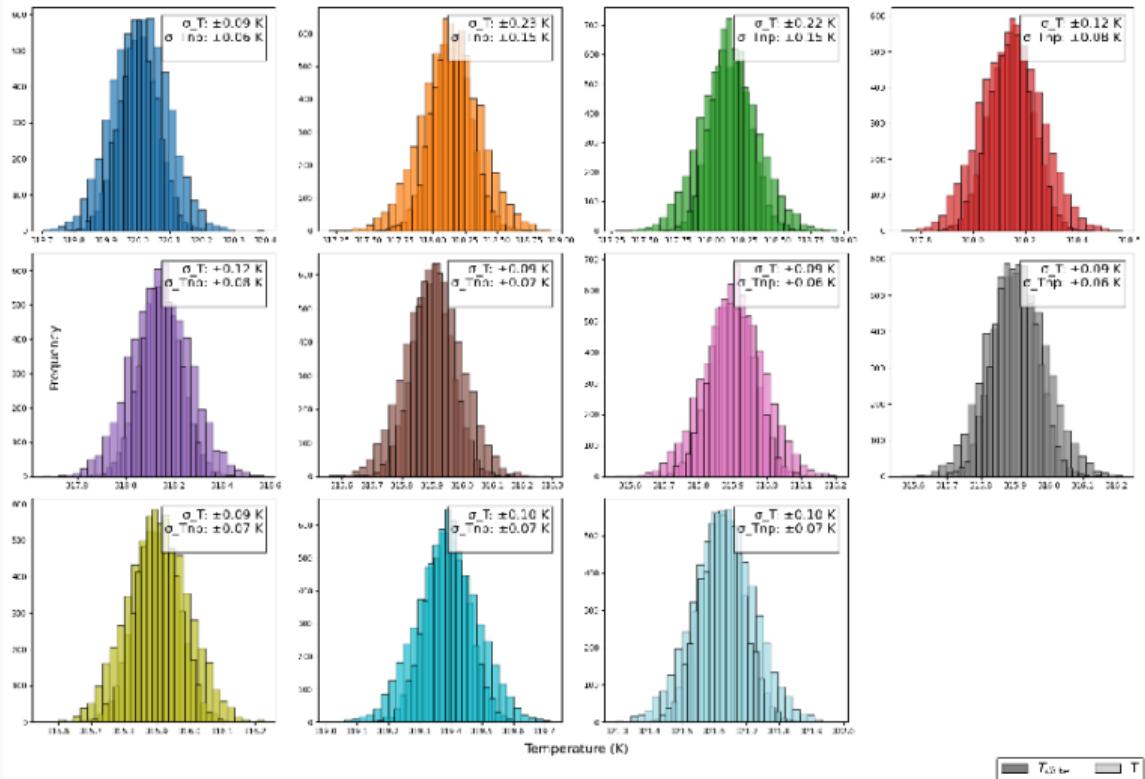
$$+ \left( \frac{\partial T_s}{\partial \Gamma_s} \Delta \Gamma_s \right)^2 + \left( \frac{\partial T_s}{\partial \Gamma_{rec}} \Delta \Gamma_{rec} \right)^2. \quad (25)$$

Note that this argument applies for both noise parameters and noise wave parameters.

There is more noise when using  
 $T_{NS} \frac{P_{cal} - P_L}{P_{NS} - P_L} + T_L$



Combined Histograms of  $T_{\text{dilute}}$  and  $T$  for Each Calibrator



Noise amplified  
by **30%** when  
using

$$T_{\text{NS}} \frac{P_{\text{cal}} - P_L}{P_{\text{NS}} - P_L} + T_L$$

Why not fit noise (wave) parameters directly?

Noise Parameter Equation:

$$P_{\text{out}}^{\text{src}} = \textcolor{red}{g} M \left( T_{\text{in}}^{\text{src}} + \textcolor{red}{T}_{\min} + T_0 \frac{4R_N}{Z_0} \frac{|\Gamma_{\text{src}} - \Gamma_{\text{opt}}|^2}{(1 - |\Gamma_{\text{src}}|^2)(1 + |\Gamma_{\text{opt}}|^2)} \right) \quad (26)$$

---

Noise Wave Equation:

$$\begin{aligned} P_{\text{out}}^{\text{src}} = & \textcolor{red}{g} \left[ T_0 + \textcolor{red}{T}_{\text{unc}} |\Gamma_s|^2 \left| \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_s \Gamma_{\text{rec}}} \right|^2 \right. \\ & + T_s (1 - |\Gamma_s|^2) \left| \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_s \Gamma_{\text{rec}}} \right|^2 + \textcolor{red}{T}_{\cos} \Re \left( \Gamma_s \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_s \Gamma_{\text{rec}}} \right) \\ & \left. + \textcolor{red}{T}_{\sin} \Im \left( \Gamma_s \frac{\sqrt{1 - |\Gamma_{\text{rec}}|^2}}{1 - \Gamma_s \Gamma_{\text{rec}}} \right) \right] \end{aligned} \quad (27)$$

We still end up with 5 unknowns, as before.