國　立　中　央　大　學

機械工程研究所

碩/博士論文

模版 ncuthesisCJK 使用說明

An example in LaTeX/XeLaTeX

研究生: 羅吉昌

指導教授: 羅吉昌

共同指導: 甲教授
乙教授

中 華 民 國 一百零二 年 六 月

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關鍵字： 碩博士論文，體裁檔，LaTeX，XeLaTeX

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謝誌

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- 感謝TeX/LaTeX網路社群內，龐大的TeX/LaTeX社群及其網頁提供無價資訊。
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- 2013/06/13碩士班 葉信麟同學發現目錄頁碼不正確。
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symbols	: 符號說明環境名，檔案則爲symbol.tex

1. Introduction

According to Hindu cosmological mythology, ancient people believe that a giant turtle bears the world on its back. Even after we stepped onto the moon at 1969, there are still plenty that we cannot explain. In the novel Lord of the Rings, the author named the path between hobbits as Mordor, which is also the name of the dark area on Pluto's moon, Charon. Recently, Mission New Horizons retrieved valuable data about Charon and Pluto. This thesis aims to explain the formation mechanisms of the red cap on the pole of Charon (fig. 1), especially during the long cold dark period, through observations in extreme ultra-violet (EUV) and vacuum ultra-violet (VUV) irradiation.

Composition of Charon

The main composition on the surface of Charon is H_2O. According to Infrared (IR) spectroscopy, it is a mixture of 90 % H_2O and 10 % tholin at millimetre depth. The second most dominant component is ammonia hydrate, which can be observed by earth-based telescopes (brown 2000, cook 2007). In far IR spectrum taken by LEISA camera on the New Horizons, concentrated ammonia is found on Organa crater (fig 2.) and throughout Charon (fig 3.) (Grundy 2016a). The third component which forms the dark red cap (tholin?) is cold-trapped methane from Pluto's atmosphere ejecta (Hoey 2017). The presence of nitrogen and other ejecta from Pluto are neglected in this thesis because according to the model of Hoey et al. (2017) (fig.4), during New horizons' approach, 98 % of the arrived ejecta is CH_4. Charon's atmospheric pressure is further constrained by New Horizons to be below 0.3 nano bars, which is 4×10^{-13} torr for all 14 atoms and molecules including CO, H_2, CH_4, Ne, Ar, etc. (fig. 5). CH_4 remains undetectable when we convert the momentum of CH_4 with 7 hops on the surface of Charon until deposited onto cold enough part is 1×10^{-11} Pa, which is 7.5×10^{-14} torr (Grundy

2016b).

VUV irradiation

Ly-α appears to be the largest source in the dark side of Charon, with attributions from both solar occultation (70 %) and resonance scattering by atomic hydrogen flow (30 %) in the solar system at flux 3.5×10^7 photons cm^{-2} s^{-1} onto the winter pole of Charon (Grundy 2016b). The flux is 50 % larger than expected before Mission New Horizons (Gladstone 2015). CH_4 deposits at temperature below 25 K at pressure 7.4×10^{-14} torr. The time for depositing CH_4 is 2 times longer at the pole (130 earth years) than at 45 lattitude according to the thermal model of Grundy et al. (2016b) (fig 6). In order to understand the formation of tholin at different latitudes of Charon, we performed VUV irradiation on CH_4+NH_3 and CH_4+NH_3+H_2O experiments with different ratios (including 3:2, 1:5, 1:10 and 1:20 for CH_4+NH_3 and 5:3:4, 1:5:5 and 1:10:10 for CH_4+NH_3+H_2O ice mixtures) to simulate the conditions at different latitudes on Charon with base pressure $3times10^{-10}$ torr, simulating atmosphere on Charon at 15 K, which corresponds to temperature on Charon at winter times (Grundy 2016b) in interstellar processing system (IPS) (Chen 2014).

EUV irradiation

Apart from VUV irradiation, EUV irradiation also took part. VUV irradiation is believed to be the main process to convert CH_4 into heavier molecules which remained on the surface of Charon until the temperature of Charon become 60 K, at which methane evaporates from the ice. The ice is then further processed by EUV, solar wind, coronal mass ejections and interstellar pickup ions, etc to produce the tholin on Charon (Grundy 2016b). The EUV irradiation (>12.4 eV) is $8.7 \times 10^7 eV cm^{-2} s^{-1}$ at mean heliocentric distance 39 A.U. whereas VUV irradiation (Ly-α) is $1.9 \times 10^9 eV cm^{-2} s^{-1}$. In order to investigate the effectiveness of EUV to VUV irradiation, we kept temperature of CH_4+NH_3 (3:2 & 1:5) and CH_4+NH_3+H_2O (5:3:4) ice mixtures at 15 K and use the monochromatic 30.4 nm (He II) light provided by High flux beamline at National Synchrotron Radiation Research Centre (NSRRC) in Taiwan to irradiate the ice mixtures.

H_2O involved?

We compared the conditions of CH_4+NH_3 and CH_4+NH_3+H_2O because tholin on Titan is believed to be formed by CH_4+N_2 and a similar colour was observed on Charon. Charon is different from Titan as H_2O dominates on Charon. What are the differences between tholin formed by CH_4+NH_3 and CH_4+N_2? What role does H_2O play on Charon in the formation of tholin? Is it just diluting the formation or new compounds are formed?

In this thesis, we will introduce the formation reaction mechanisms of CH_4+NH_3 ice mixtures in EUV and VUV irradiation (section 3), the formation reaction mechanisms of CH_4+NH_3+H_2O ice mixtures in EUV and VUV irradiation (section 4), and the residues of these mixtures and a brief comparison with tholin on Titan will be made (section 5). With these results, we will have a better understanding about Charon and some astrophysical implications will be presented (section 6).


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2. Methods

2.1 Laboratory Astrophysics

To study the chemical reactivity in astrophysical environment experimentally, we conducted our experiments in Interstellar photoprocessing system (IPS) (Chen et al. 2014), an ultrahigh vacuum chamber with base pressure 3×10^{-10} torr and 14 K, corresponds to a density of 10^6 cm^{-3}, similar to dense cloud interiors. The system will be introduced in detail in section 2.1.1. To simulate the irradiation in interstellar environments, we use a micro-wave discharge hydrogen lamp (MDHL) and monochromatic extreme-ultraviolet irradiation (EUV) 30.4 nm to irradiate our ice mixtures, and they will be introduced in section 2.1.2 and 2.1.3 respectively. The experimental protocols will be elaborated in section 2.2. In order to better understand the physics behind, some basic theories of Infrared spectroscopy and concepts of chemical kinetics used in data analysis are included in section 2.3 and 2.4 respectively. To demonstrate the ice mixtures in KBOs, we used different configurations of ice mixtures that refers to different sections in chapter 3 and chapter 4.

2.1.1 Experimental simulations by IPS system

We conducted our astrophysical simulations studied in chapter 3 to 4 in Interstellar Photo Processing System (IPS) (figure 2.1). IPS consists in three systems: the main chamber, where our experiments take places; the detection system, where we collect our data; and a gasline system, where we prepare our samples.

The main system consists of an ultrahigh vacuum chamber equipped with a closed-cycle helium cryostat (CTI-M350). It is pumped by a turbo molecular pump (KTKT FF – 160/620ZE, capacity 600 liters s^{-1}), which is backed up by a scroll pump, and a non –evaporation get-

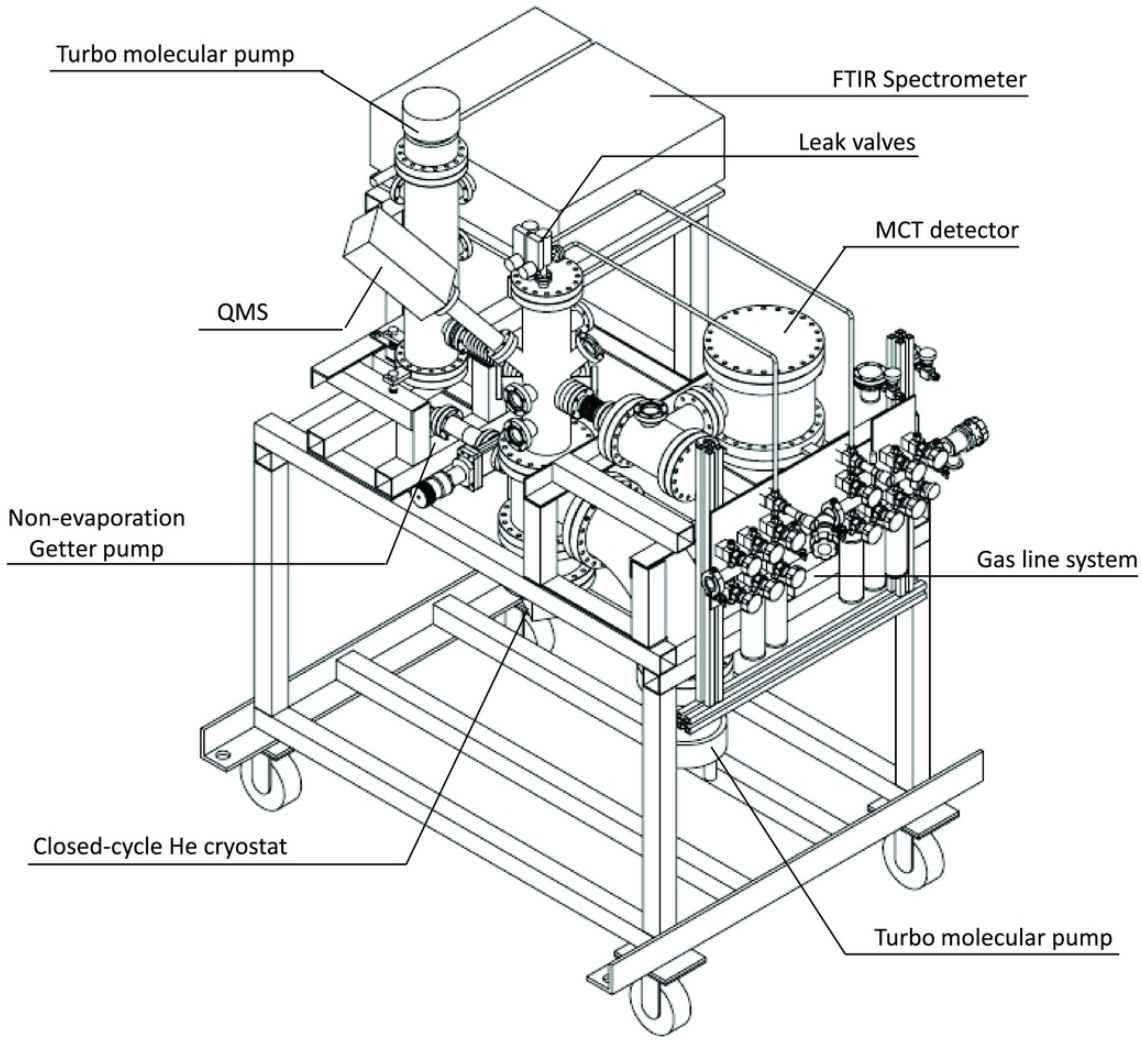


Figure 2.1: The schematic diagram of IPS system, mechanical pumps are not shown for clarity. (Quoted from Chen et al. 2014)

ter pump. The getter pump is a powerful tool to adsorb residue gases inside the main chamber, with a larger surface area, H_2, CO and N_2 are adsorbed to obtain a better base pressure. After baking, the base pressure of our main chamber can reach 1×10^{-10} torr at 14 K, monitored by a Granville-Phillips 370 Stabil-Ion gauge. This pressure can be used to demonstrate the dense cloud interior environments and star forming region. The substrate we have chosen is KBr, which can allow infra-red photons with 700 to 4000 cm^{-1} to penetrate. It is mounted by substrate holder made of oxygen-free copper, on the first stage of cold finger mounted on the tip of cryostat. Two silicon diodes and also a heater were placed onto the cold finger and one of the silicon diodes is near the substrate holder. They were connected to a temperature controller and PID system to achieve a warmup rate of 1K/min with an accuracy of 0.1 K.

The detection system consists in a mid-infrared Fourier transform spectrometer (mid-FTIR) (ABB FTLA2000-104) and a Quadrupole Mass Spectrometer (QMS). To prevent absorption bands of CO, CO_2 and H_2O gas in the atmosphere, the IR beam path was built inside vacuum, pumped by dry pump. The main chamber and the IR path are separated by ZnSe windows, which can allow infra-red penetration from 0.5 – 20 um with absorption less than 0.07 %. In this study, the infrared spectra are obtained with resolution of 4 cm^{-1} and averaged over 32 scans. The angle between the IR beam path and the substrate holder is 45 degrees. The QMS (MKS Microvision 2) consists of a controller and mechanical part sealed by a mounting flange in ultrahigh vacuum. It is mounted 10 cm from the substrate and run with a resolution 0.5 a.m.u. The Ionizer release 70 eV electron by filament and ionize incoming molecules to positive charged ions between anode grid and repeller. The ions were accelerated by focus plate and enters ion filter, which consists of four circular rods, with a combination of A.C and D.C. potential to sieve whole bandpass ions at millisecond timescale. The selected ions enter ion detector and are detected by either faraday cup and continuous dynode electron multiplier (CDEM) which can secondary multiply weak signals.

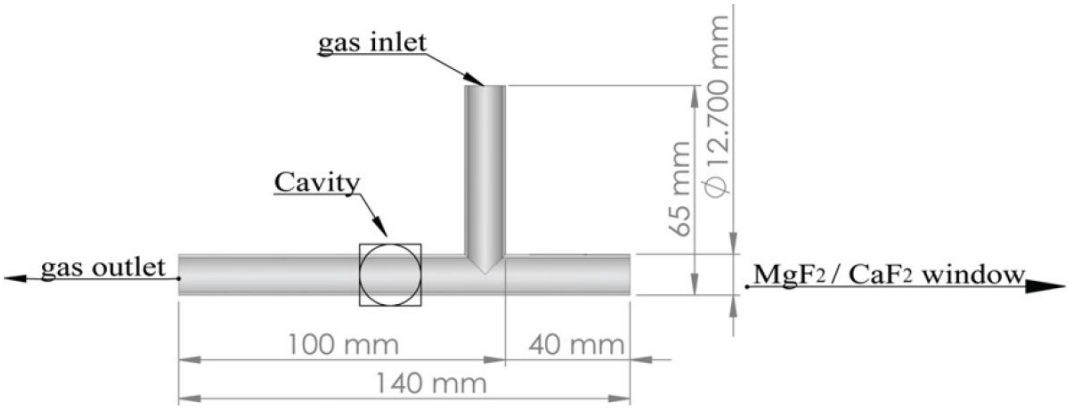


Figure 2.2: The cross-section of MDHL (T-type geometry) (Quoted from Chen et al. 2014).

The samples are prepared in situ in our gasline system. It contains four stainless steel bottles with the same volume, which is used to determine relative proportion of the gas mixtures by their partial pressures. The ammonia gas 99.99 % and methane 99.999 % are mixed with partial pressure measured by a Baratron with 0 - 100 torr range with a 0.25

2.1.2 Vacuum-UV source

In order to simulate the photoprocessing of vacuum ultraviolet (VUV) irradiation onto the interstellar ices and ices on planetary bodies, including KBOs, the ice mixtures are irradiated with a T-type Microwave-Discharged Hydrogen-flow lamp (MDHL). The molecular hydrogen with pressure 0.4 torr flows through the lamp with a support of a mechanical pump. Using a 2.4 GHz microwave generator and high voltage discharge, a low pressure plasma is produced in the Evenson cavity. Figure 2.2 shows a cross-section of T-type quartz tube; the middle part of the T-type quartz tube is being tunned by a ceramic rod that is called Evenson cavity. In order to measure the photon flux in situ, we use an 88 % transmittance nickel mesh with its photoelectric efficiency being obtained by high-flux beamline in National Synchrotron and a SXUV 100 photodiode calibrated by NIST. A MgF_2 window is placed between the lamp and the sample holder to prevent penetration of VUV photons with wavelength shorter than 114nm, leads to a cut off at 114nm. Figure 2.3 shows a VUV emission spectrum of a MDHL. It consists in Ly-α (121.6nm) and H_2 molecular emission in 110-180 nm

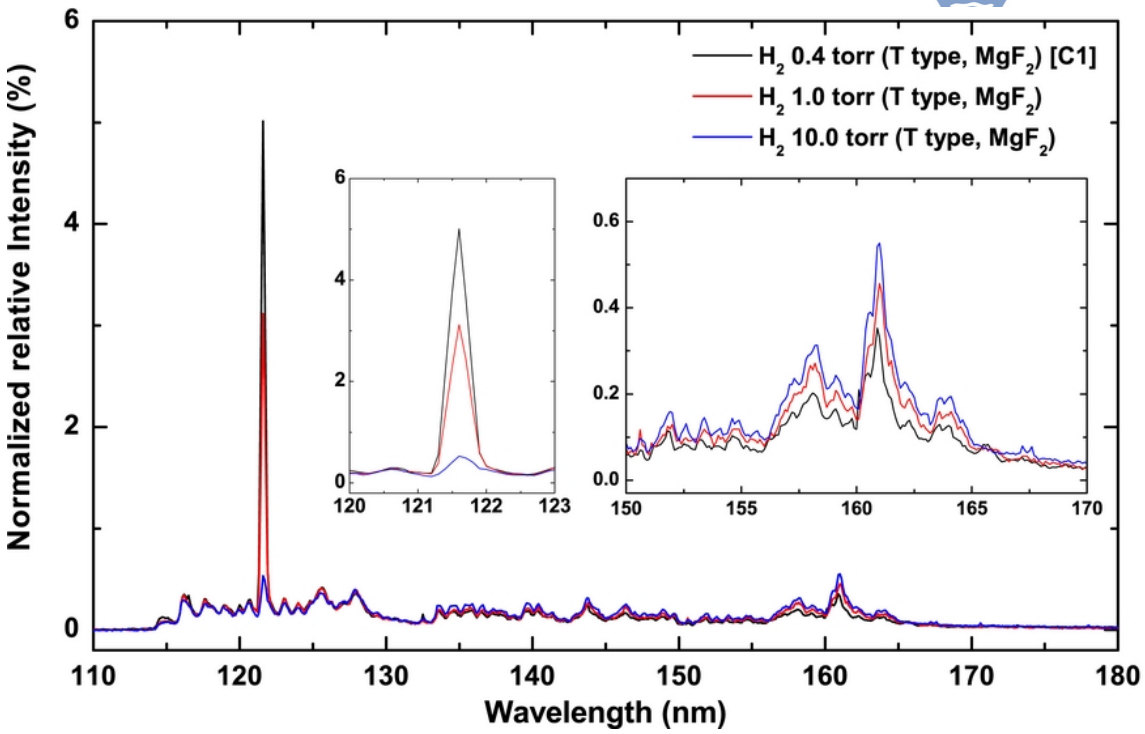


Figure 2.3: VUV spectra of MDHL (T-type geometry, 110-180 nm) with different H_2 pressure inside the lamp(Quoted from Chen et al. 2014).

range. Chen et al. (2014) showed that the spectral characteristics of the VUV light emitted in this range depends on the gas type (mixture of H_2 with He or Ar etc), pressure of H_2 and lamp geometry. Throughout those configurations stated there, we adopted 0.4 torr molecular hydrogen and T-type MDHL that produces VUV irradiation at 114-170 nm with 19.1 % of Ly-α and a mean photon energy of 9.27 eV. The photon flux is 6.4×10^{13} photons $cm^{-2}s^{-1}$ at sample position.

2.1.3 Extreme EUV source

To simulate the solar EUV irradiation reflected by IPM on both Charon and interstellar ices, we use the HF-CGM high – flux beam line of the National Synchrotron Radiation Research Center in Hsinchu, Taiwan. It provides a continuum EUV to VUV photons from 4 to 40 eV. The continuum is separated into monochromatic He II line (30.4nm) with a six-meter cylindrical grating monochrometer with an incident angle of 70 degrees. With the help of a movable entrance slit and movable curved exit slit, the energy resolving power can reach around 3×10^4 at 40 eV for grating 1600 l/mm with both slits movable and set opening to

10 μm (Hsieh 1998). Similar to VUV irradiation provided by MDHL, the light intensity was monitored by the same nickel mesh with photoelectric efficiency obtained by SXUV 100 photodiode calibrated by NIST. With the known photoelectric efficiency, the flux of monochromatic 30.4nm is measured to be 2.15×10^{14} photons $s^{-1}cm^{-2}$ with a spot size of 1 cm

2.2 Experimental Protocol

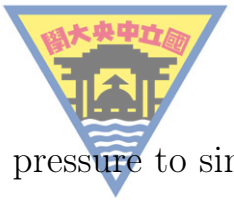
In this section, we will briefly introduce the procedures of how we performed our experiments. It is divided into four parts, preparation and cooling, deposition, irradiation and warmup.

Preparation of experiments and cooling

Before any of experiment is done, we bake our system at 100 oC for 48 hours to reduce the contamination of water and residue gases as much as possible. It was cooled to room temperature that the background pressure can reach routinely at 1×10^{-10} torr. The gasline were connected with the regulators of the gas tanks and bake to 100 ^oC and pumped by molecularturbo pump for two days before any experiment were done. Also, The water sample has been freeze thaw several times by liquid nitrogen until there is no pressure increase recorded by baratron when water is freezed. Before cooling the substrate to cryogenic temperature, we took an IR spectrum and started the monitoring of residue gases by QMS in order to compare the residue molecules and to verify any possible contaminations in the main chamber. We then start the cooling process thanks to the closed-cycle He cryostat.

Deposition

The gas mixtures are pre-mixed in our gasline system introduced in section 2.1.1. We used a leak valve to condense the gas from the stainless steel bottles onto pre-cooled KBr substrate at 14 K, which monitored by Fourier transformed Infra-red spectroscopy (FTIR) and Quadrupole mass spectrometer (QMS) during deposition. The pressure of deposition is fixed to 1×10^{-8} torr that the deposition rate is $4 \times 10^{16} molecules cm^{-2} min^{-1}$. After deposition, we placed the ice mixture at 14 K for 60 minutes and to allow pumping of residue gas, until



pressure of the main chamber reduce back to its base pressure to simulate the interstellar environment before irradiation.

Photon Irradiation

The total irradiation time is 270 to 450 minutes depend on experiment configurations; with time intervals varies from 2 to 30 minutes. After each irradiation, we waited for 10 minutes allowing pumping out of the photodesorpted gas molecules. During irradiation, the photon flux is monitored by a nickel mesh. After Irradiation, we place the sample for 30 minutes to observe if any thermal reaction was conducted.

Warmup

We use 1 K/min to warmup the substrate to 300 K to demonstrate effects of a new born star nearby an interstellar cloud. During warmup, we record the QMS from 1 to 100 a.m.u. to observe if there are low quantity of higher mass product formed during irradiation.

2.3 Infra-red spectroscopy and the Beer's Law

We used infra-red spectroscopy extensively in chapter 3 and 4, it is a powerful tool in studying molecular interactions during irradiation and warmup. We choose infra-red rather than Ramen spectroscopy because infra-red has lower energy that it would not change the structure of the ice mixture nor breaking any of the bonds. With different vibration modes, the energy absorbed by molecules are quantized. With the energy of absorption bands in infra-red spectrum, we may identify the functional group of the species. To simply classify, molecules can have, from less energetic, translational, rotational and vibrational motions. Generally, vibrational motions can be divided into stretching and bending. Stretching needs more energy than bending. For stretching, there exist Symmetric and Asymmetric stretching, while bending can be divided into In-plane Scissoring, rocking and out of plane Wagging and Twisting (Figure 2.4).

By Beer's Law, we may calculate the column density of the molecule

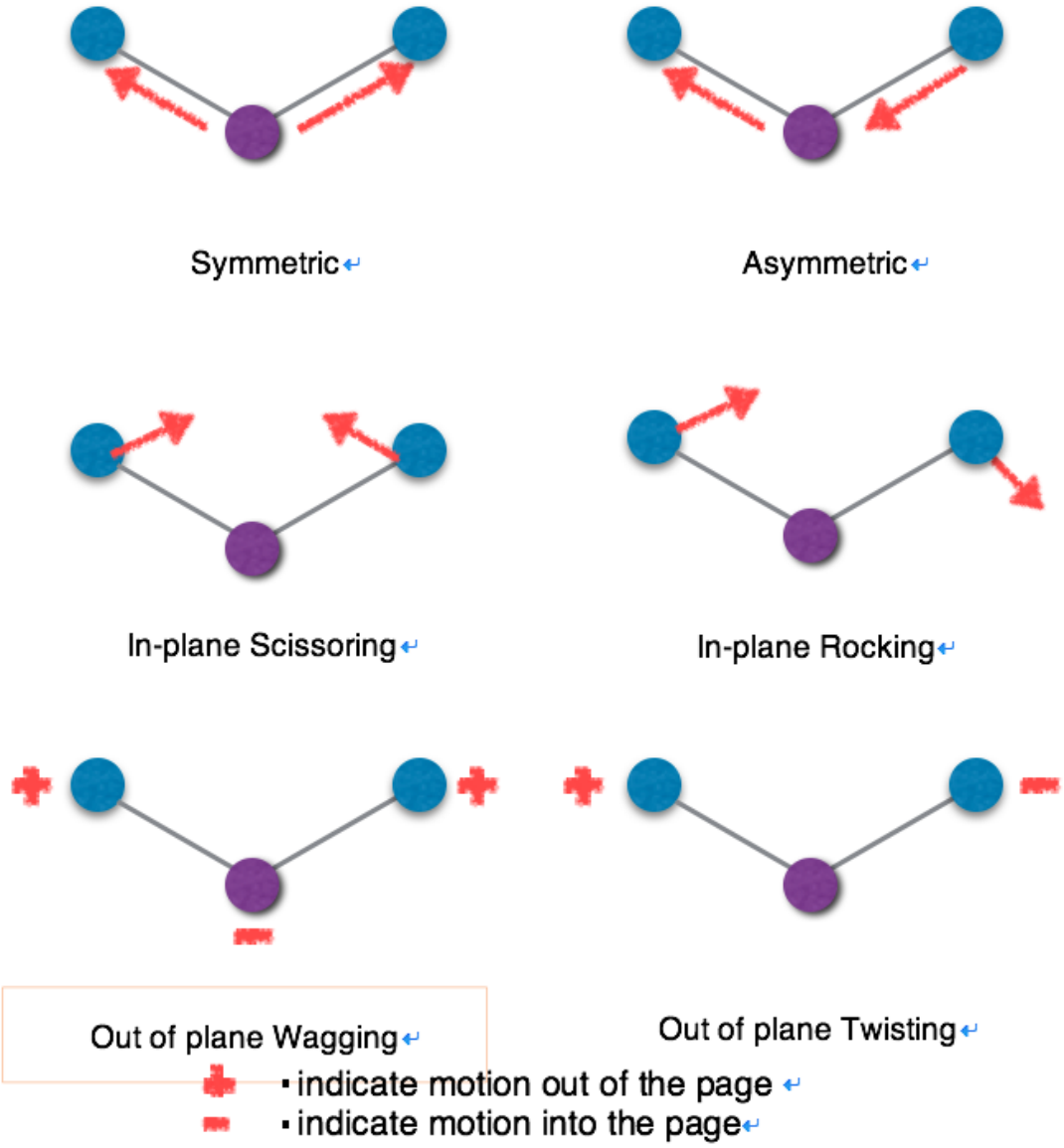


Figure 2.4: Different vibrational modes of a three atom molecule.

with its functional groups, which are used to plot figures in chapter 3 and 4. Beer Lambert's Law suggest that when light passes through a medium, amount of light absorbed is proportional to density and path length of the medium. Assume the known intensity beam $I_0(\nu)$ passes through the medium and beam intensity become $I(\nu)$. The transmittance $T(\nu)$ is defined by equation 2.1.

$$T(\nu) = \frac{I(\nu)}{I_0(\nu)} \tag{2.1}$$

Also, the absorbance $a(\nu)$ is defined by equation 2.2.

$$a(\nu) = -\ln T(\nu) = -\ln \frac{I(\nu)}{I_0(\nu)} = nl\sigma(\nu) \tag{2.2}$$

where n is number density (molecules/cm^3), l is the path length (cm), $\sigma(\nu)$ is the cross-section (cm^2/molecule) of corresponding frequency ν. This equation is known as Lambert Beer's Law.

As the ice mixture in our thesis are at 14K, the peaks of absorbance are often a broadband due to coupling between neighbor molecules. Therefore, we can integrate the whole band of the peak equation 2.2 with respect to frequency and use the absorbance strength (A value) in literatures to calculate the column densities N of the ices by equation 2.3.

$$N = \frac{\int a(\nu)\mathrm{d}\nu}{A(\nu)} \tag{2.3}$$

where N is the column density (molecule cm^{-2}), $A(\nu)$ is the absorbance strength (cm molecule^{-1}).

2.4 Reaction Rate Laws

In this section, we will introduce rate reaction of a consecutive reaction and the concept of pseudo first order which we used to fit our reaction product against irradiation time. The rate of a chemical reaction is the relation between change in concentration of a substance per unit of time. i.e. For a balanced chemical reaction, A $\rightarrow$ 2B, the rate of

reaction is $-\frac{\Delta[A]}{\Delta t}$. The formation rate of B is 2 times destruction rate of A.

When there are two reactants, with balanced equation 2A + B $\rightarrow$ 2C. The reaction is a third order overall, second order in A and first order in B. rate $= k[A]^2[B]$.

To determine the order of a reaction, we can only determine it experimentally. One way is method of initial rates. By changing concentration of initial reactants, and find out the initial reaction rate, we may find out the relation between two reactants and the rate. i.e. rate $= k[A]^x[B]^y$. For a reaction with only one reactant $[R]$, we may use the relation between time and reactant concentration to plot graphs to find out the order or reaction. For a zero order reaction, the rate is not depending on any reactant that it is a constant. The rate $= -\frac{\Delta[R]}{\Delta t} = k[R]^0$. By calculus, $[R]_0 - [R]_t = kt$.

For a first order reaction, rate $= -\frac{\Delta[R]}{\Delta t} = k[R]$. By calculus, $\ln[R]_t = -kt + \ln[R]_0$.

For a second order reaction, rate $= -\frac{\Delta[R]}{\Delta t} = k[R]^2$. By calculus, $\frac{1}{[R]_t} - \frac{1}{[R]_0} = kt$.

Hence, if we get a straight line in a plot between time as x-axis, and the concentration of reactant as y axis, it is a zeroth order reaction, similarly, in first order reactions, we get straight line in plots between $\ln[R]$ as y axis and t in x axis.

In a reaction with one reactant in excess, the rate of reaction is called pseudo first order reaction where pseudo means pretended. For A+B $\rightarrow$ C, rate $= k[A][B]$. As $[B]_0 \gg [A]_0$, change of $[B]$ is negligible that $[B] \sim [B]_0$. Therefore, $[B]$ is assumed to be a constant and included in the rate constant k.

For a consecutive reaction, where A $\rightarrow$ B $\rightarrow$ C that the produced product will not convert back as reactant. A simple example is radioactive decay. At $t = 0$, $[A] = [A]_0$, $[B] = 0$, $[C] = 0$ and at all times, $[A] + [B] + [C] = [A]_0$. The rate equations are as follows:

$$-\frac{\Delta[A]}{\Delta t} = k_1[A] \tag{2.4}$$

$$-\frac{\Delta[B]}{\Delta t} = k_1[A] - k_2[B] \tag{2.5}$$

$$-\frac{\Delta[C]}{\Delta t} = k_2[B] \tag{2.6}$$

By equation 2.4, we get

$$[A] = [A]_0 e^{-k_1 t} \tag{2.7}$$

By substituting equation 2.7 into equation 2.5, we get

$$-\frac{\Delta[B]}{\Delta t} + k_2[B] = k_1[A]_0 e^{-k_1 t} \tag{2.8}$$

After solving the differential equation 2.8 , we get

$$[B] = \frac{k_1}{k_2 - k_1}(e^{-k_1 t} - e^{-k_2 t})[A]_0 \tag{2.9}$$

Finally, since $[C] = [A]_0 - [B] - [A]$, by equation 2.7 and 2.9, we get

$$[C] = \left(1 + \frac{k_1 e^{-k_2 t} - k_2 e^{-k_1 t}}{k_2 - k_1}\right)[A]_0 \tag{2.10}$$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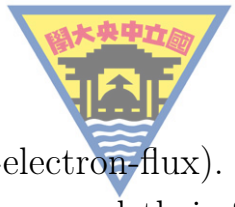
3. Experimental Results of CH_4 + NH_3 ice mixtures

According to Grundy et al. (2016), CH_4 from Pluto may accumulate by cold-trapping, onto surface of Charon. The amount of CH_4 varies throughout the surface of Charon because it depends on duration of temperature below 25 K. The duration depends on diurnal motion and thermal inertia of Charon. With a tilted axis of 112 degrees to the ecliptic, higher concentration of CH_4 will accumulate at the pole (see chapter 1 for details). Therefore, we investigate different concentrations of CH_4+NH_3 ice mixtures and answer several questions: Will different concentration of CH_4 mix with high concentration of ammonia observed on crater position and throughout the surface of Charon (Grundy et al. 2016) have structure difference in accumulation of tholin? Are there variations of photo-products when concentration of CH_4 differ during warm-up? Since both EUV and VUV irradiation irradiates onto Charon, are there any differences when we change the photon source from VUV to EUV to irradiate the ice mixtures?

The main source to irradiate the dark side of Charon is Lyα reflected by interplanetary medium (Grundy 2016). Other sources such as the energetic ions in solar wind, consists of mainly H^+, He^+, He^{++} and O^{2+} etc are originated from solar corona or IPM. These ions would also reflect solar irradiation to the dark side of Charon. Among these irradiations, we picked He II irradiation because He II is 3 – 20 times more intense then He I during a solar flare. As it varies, it is difficult to estimate the dose onto Charon. Besides, electronic flux is also present in solar wind but it is one order of magnitude lower than proton flux. The flux for energetic electrons observed at the 1 A. U. position is



available (http://www.swpc.noaa.gov/products/goes-electron-flux). Although electron flux is much less important than Lyα, and their flux varies, we also compare the electron irradiation experiment done by Kim and Kaiser (2011) on CH_4+NH_3 ice mixtures in this chapter.

When Charon is shine by direct sun light, the surface temperature increases and deliver the heat to the poles by conduction. From the model of Grundy et al. (2016), the surface temperature of the pole area would increase to 60 K that the heating rate depends on the thermal conductivity of Charon. To demonstrate the heating process, we warmup our ice mixture with a heating rate 1 K/min and monitor the ice by both QMS and scanning IR spectra with 5 K intervals. We will look into whether there are new species formed during warmup and monitor the gas phase desorption.

Finally, in this chapter, after we focus on the concentration effect of CH_4 on photo-products, photon energy effects, species detected during warmup phases, we present the residues accumulated by irradiating CH_4+NH_3 ice mixtures with different ratios. Since both tholin formed on Titan and Charon has similar colour, we also compare the IR spectra of MDHL, NSRRC with different configurations with the residues on Titan with experiments done by Imanaka et al.

3.1 3.1.1 The concentration effect of CH_4 on production of C_2H_6 and CN^-

We first look into the concentration effects of CH_4 by irradiation by VUV irradiation. Before and after deposition, we scanned an IR spectrum and plotted the absorbance of the ice mixtures. figure 3.1.1 plots the absorbance of the CH_4+NH_3 ice mixtures in different ratios. Due to the ice thickness, the infra-red spectrum of CH_4+NH_3 = 3:2 consist of 900 ML of CH_4 and 600 ML of NH_3 is tilted due to interference. Since the amount of ammonia is fixed (600 ML) in all the ratios, other ratios has less CH_4 and this problem is less serious. This is also not

observed in the ice mixtures after irradiation. Using the same reference spectrum, that is the infra-red spectrum recorded before deposition, we plotted the infra-red spectrum of irradiated ice mixture in figure 3.1.2. It shows the absorbance of each ratio after irradiation. We labelled the peaks which we used to calculate the column densities onto the graph. Main products we have detected are C_2H_6, CN^- and C_3H_8. The peak positions with the references are listed at 3.1.

Table 3.1: The peak positions of identified substances after irradiation in different configurations of ice mixtures.

Literture assignments		CH_4+NH_3 ratio (MDHL)				
Wavenumber (cm^{-1})	Carrier	1:5 (cm^{-1})	1:10 (cm^{-1})	1:20 (cm^{-1})	3:2 (cm^{-1})	Ref.
3375	ν_3 (NH_3)	3366	3366	3369	3367	1
3290	$2\nu_4$ (NH_3)	-	-	-	-	1
3210	ν_1 (NH_3)	3207	3208	3210	3205	1
3011	ν_3 (CH_4)	-	-	-	-	2
2972	ν_{10} (C_2H_6)	2975	-	-	2975	3
2960	C_3H_8	-	-	-	2960	7
2941	$\nu_8 + \nu_1 1$ (C_2H_6)	2940	-	-	2940	3
2904	ν_1 (CH_4)	2901	-	-	2901	5
2879	ν_5 (C_2H_6)	2882	2883	-	2882	3
2814	$\nu_2 + \nu_4$ (CH_4)	-	-	-	2815	5
2083	ν (CN^-)	2088	2087	2088	2088	2
1625	ν_4 (NH_3)	1625	1625	1626	1631	1
1514	δ (NH_2)	1509	1507	1505	1511	6
1465-1440	deform CH_2 scissor	1461	-	-	1463	3,4
1390-1370	CH_3 sym deform	1394	1394	1394	1372	4
1298	ν_4 (CH_4)	1301	1302	1305	1299	2
1075	ν_2 (NH_3)	1073	1072	1072	1072	1
820	$\nu_1 2$ (C_2H_6)	-	-	-	820	3

Reference: 1. Bossa et al 2008 2. Moore and Hudson 2003 3. Kim et al. 2010 4. Socrates 2001 5. Bennet and Kaiser 2007 6. Zheng et al. 2008 7. Hudson and Moore 2004

From infra-red absorption spectrum and their positions, we assigned the peak 2086 cm^{-1} to CN^- but not a combination of HCN and CN^- because of we cannot observe the HCN bending mode located at 848 cm^{-1}. Although in the case CH_4 + NH_3 = 3:2, we may observe a peak located at 820 cm^{-1}, the peak is with a FWHM half of HCN and it disappeared at 50 K. Since 50 K is the desorbing temperature of C_2H_6 and the position is v12 mode of C_2H_6, we assign the 820 cm^{-1} peak as C_2H_6. As the absence of HCN bending mode, we may assign our peak located at 2086 cm^{-1} as CN^-. Other assignments such as C_2H_6 and C_3H_8 are observed with multiple peaks.

We integrated the area and divided by the absorption strength stated in table 3.2. Although we understand that there is an average error in absorption strengths of no more than 10 % when the pure ice is diluted in N_2 and H_2O (Richey and Gerakines 2012). In our case, absorption strengths changes after CH_4 and NH_3 are mixed. For example, according to d' Hendecourt and Allamandola (1986), the band of NH_3 located at 1070 cm^{-1} would not change much (from 1.1×10^{-17} to 1.2×10^{-17}) when excess water is added to pure NH_3 and therefore, we may use the same absorption strength throughout our discussion to give a brief concept on what is the column density of the species and how is the absorption area changes when concentrations of ice mixtures and photon energy are changed. For the case of CN^-, we know that CN^- has a bond order =3 by its molecular orbitals which is different from CN stretching (bond order 2.5) which is very sensitive to the matrix environment. As an example, by Borget et al. (2012), the CN stretch in amino acetonitrile change by factor of 2 between the pure molecule itself and in a mixture of amino acetonitrile and H_2O (1:3). Here, we adopt the absorption strengths stated in Table 3.2 and neglect the error in absorption strengths.

Table 3.2: The strength of absorbance adopted in this thesis measured in literatures of pure ice samples

Wavenumber (cm^{-1})	Assignment	Vibration	FWHM	A value ($\times 10^{-17}$)	Reference
2976	C_2H_6	$-CH_3$	-	1.05	2
2960	C_3H_8	$-CH_2-$	-	2.58	2
2086	CN^-	CN	-	1.8	3
1297	CH_4	CH deformation	8	0.61	1
1070	NH_3	"umbrella mode"	68	1.7	1

Reference: 1. d'Hendecourt and Allamandola (1986) 2. Moore and Hudson (1998) 3. noble et al. (2013)

3.2 3.1.1 Reaction mechanisms

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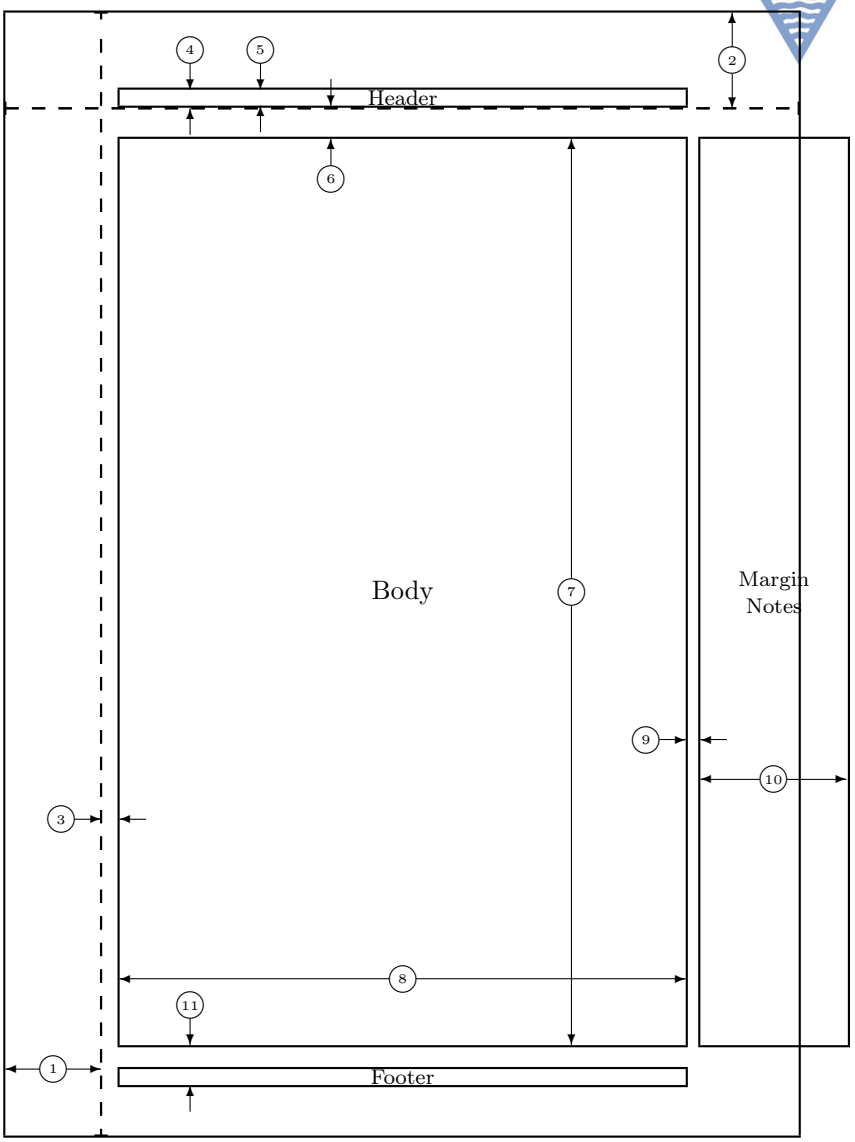

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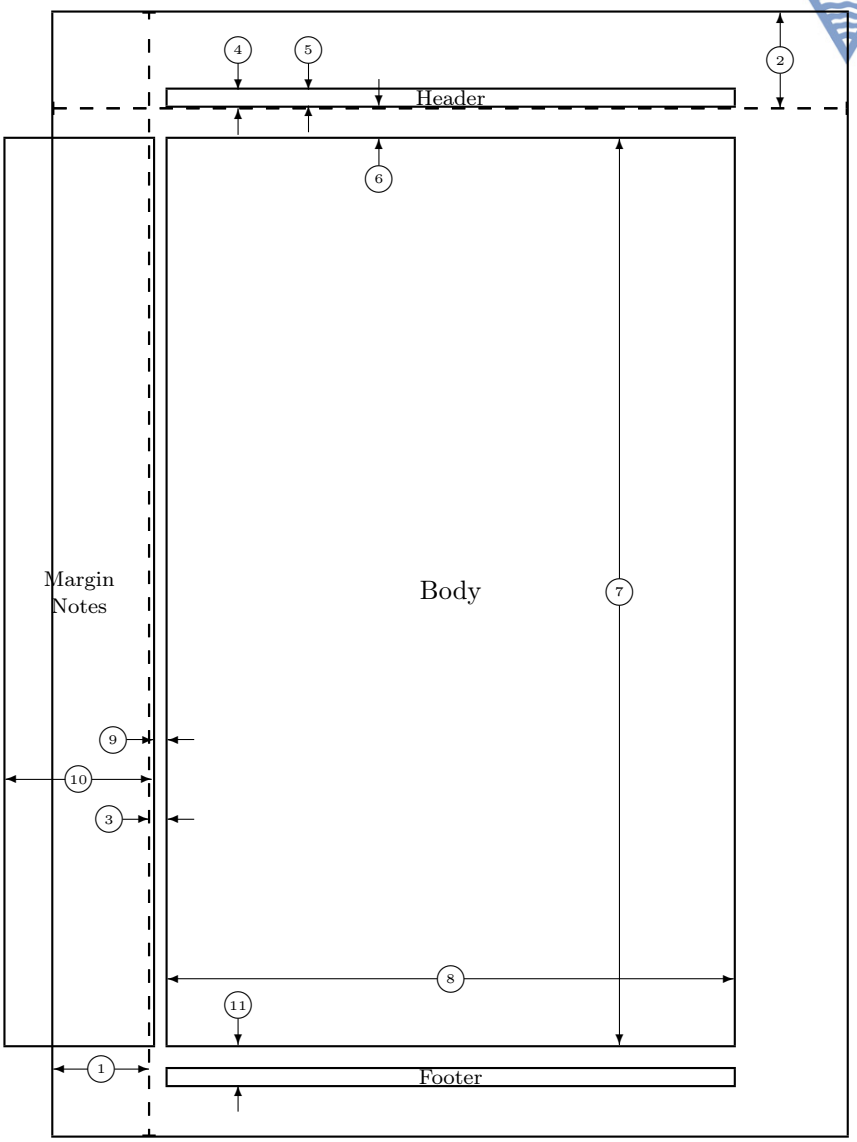

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