

# **Research Proposal for DAAD PhD scholarship**

## **The Gaseous Cosmic Web**

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# 1 General Information

## Information of the Applicant

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More detailed information of the applicant can be found here <http://goo.gl/BPZGzK>

## Information of the Project

<b>Title</b>	<b>The Gaseous Cosmic Web</b>
<b>Field</b>	Cosmology, Astrophysics, Physical Sciences
<b>Advisor 1</b>	Volker Springel, PhD. Heidelberg Institute for Theoretical Studies (HITS) & University of Heidelberg, Germany
<b>Advisor 2</b>	Jaime Forero-Romero, PhD. Universidad de los Andes, Colombia
<b>University</b>	University of Heidelberg, IMPRS PhD program
<b>Time Frame</b>	3 years

## 2 Abstract

## 3 Introduction

Since the filamentary nature of the large-scale matter distribution of the observable cosmos (the so-called cosmic web) was evidenced by the first compiled galaxy surveys (Chincarini & Rood, 1975; Gregory & Thompson, 1978; Einasto et al., 1980a,b; Kirshner et al., 1981, 1987), it has been identified as one of the most striking features of the Megaparsec Universe and an increasing interest in studying its dynamical properties and environmental influences on a plethora of different astrophysical phenomena has become evident. At present, a tremendous amount of observational data supports the cosmic web scenario at the point that it has become an essential part of the current standard paradigm in cosmology. Last generation galaxy redshift surveys, such as the *two-degree-Field Galaxy Redshift Survey* (2dFGRS) and the *Sloan Digital Sky Survey* (SDSS), do evince the intricate and complex structure of the cosmic

web at a level of detail never seen before. In addition, other valuable observational resources like X-ray emissions of hot intracluster gas embedded into large clusters of galaxies, Ly- $\alpha$  forest absorption lines in the spectra of shock-heated neutral hydrogen gas residing in filaments and clusters, and weak gravitational lensing and imprints in the CMB field produced by foreground structures have also validated this picture undoubtedly.

On the theoretical side, early descriptions of the evolution of the large-scale Universe, based on gravitational instabilities in primordial stages and leaded by the seminal work of Zel'dovich (1970), are highly consistent with the cosmic web picture, where planar pancake-like regions of matter enclose enormous sub-dense voids and are bordered, in turn, by thin filaments and high-density clumpy knots (Bond et al., 1996). Since then, our understanding of the structure and dynamics of the cosmic web has been dramatically improved as new and more powerful theoretical and computational tools and more refined observational data become available. In particular, N-body simulations, fuelled by last generation computing systems and ever more efficient numerical algorithms, are acquiring an increasingly important role in fathoming the complexity of the large-scale Universe.

Due to the poorly interacting nature of the dark matter component of the cosmic inventory, observations have been devoted to establish the underlying structure of the cosmic web entirely based on detecting baryonic matter (with the exception of non-direct inferences based on weak gravitational lensing). On the other hand, the highly complex *gastrophysical* processes involved in baryonic dynamics, i.e. shock heating, photoionization, supernova feedback, stellar wind, radiative cooling, star formation and others, make extremely difficult to obtain a completely consistent and reliable scenario from numerical simulations as many of these processes are not fully understood yet. Accordingly, most of the related numerical research has been made based on dark matter-only N-body simulations, where the gas dynamics has been neglected. Although this duality between observations and simulations can be thought as a complementary situation, actually it also makes quite hard to splice both, observational data and numerical predictions.

Although a visual naked-eye identification of single structures (i.e. voids, walls, filaments or knots) is always possible in galaxy surveys and simulations, translating this visual impression into a formal mathematical formulation is not such easy and much effort has been made in this direction.

## **4 Objectives**

## **5 Methods**

## **6 Expected Results**

## **7 Bibliography**

### **References**

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## **8 Timetable**