

# 2-m Himalayan Chandra Telescope (HCT)

## Proposal for Observations

DEADLINES: 1 March; 1 July; 1 November

MAIL TO:

HCT Time Allocation Committee,  
Indian Institute of Astrophysics  
Bangalore 560 034, INDIA  
e-mail: htac@iiap.res.in

Proposal Code:

Received:

Cycle applying for: 2015-Cycle2

Date:

**1. Title of the proposal : Late time observations of the nearby SN2014J in optical and near infrared**

☐ Short term

☐ Long term

Number of cycles/nights:

☐ Ongoing proposal

Previous proposal code(s):

☒ Thesis topic

Expected year of thesis submission: 2016

**If proposal is intended to support a Ph. D. project, please include, in addition to the Scientific Justification, a brief outline of the Ph. D. project and the relevance of the proposal to the Ph. D. project**

**2. List of Proposers:** *indicate PI(s)*

Proposer	Affiliation	e-mail	Will be present for observations?
S. Dhawan	ESO	sdhawan@eso.org	no
B. Leibundgut	ESO	bleibund@eso.org	no
S. Taubenberger	ESO	tauben@mpa-garching.mpg.de	no
W. Kerzendorf	ESO	wkerzend@eso.org	no
K. Maguire	ESO	maguirek@eso.org	no
J. Spyromilio	ESO	jspyromi@eso.org	no

**3. Contact Name & Address:**

European Southern Observatory, Karl Schwarzschild Strasse, 2 Garching bei München, 85748, Germany

Telephone:

Fax:

**4. Abstract:** Type Ia supernovae (SN Ia) have been used as exceptional cosmological distance indicators. They are thermonuclear explosions of white dwarfs (WDs) in a binary system. However, there are several unanswered questions regarding the physics of SN Ia, for e.g. the progenitor mass, explosion mechanism and the nature of the companion. The late time photometry, in the optical and near infrared (NIR), is a crucial tool in answering some of these questions. At late times, the ejecta are transparent to  $\gamma$  rays and the light curve is powered by positron kinetic energy. The late time (pseudo-) bolometric light curve allows us to determine the positron escape fraction. The extent of trapping depends on the strength and configuration of the magnetic field. A strong, tangled magnetic field would arise from a Chandrasekhar mass WD whereas a weak magnetic field would indicate an edge-lit detonation of a sub-Chandrasekhar WD. Thus, the late time bolometric light curve is an efficient tool in constraining the explosion scenario. The proposed observations are part of an observing campaign to obtain the first detailed nebular phase dataset for an SN Ia. We have been awarded Gemini-North time (on GNIRS) to obtain NIR spectra for SN2014J, the closest SNIa in over 4 decades.

**5. Status of ongoing / previous proposals:**

1. Please give a brief status report of any previous HCT proposals, and attach any preprint/reprint based on these HCT observations
2. If your proposal is long-term / on-going, briefly state the status of the proposal, mentioning the progress with respect to the science goals.

**NOTE: Incomplete proposals are likely to be given low priority or rejected**

N/A

---

**For official purpose only**

*Referee's comments:*

*Science feasibility:*

*Technical feasibility:*

*Grade of the proposal:*

*Dates allotted:*

**6. Scheduling request:**

- ☐ Dark night is essential                      ☐ Grey night is all right  
☐ Bright night is all right                      ☒ Time-critical observations  
☐ Target of Opportunity                      ☒ Other (specify) We request dark night for optical observations only, NIR can be conducted in grey or bright conditions

No. of nights requested: 6 hrs (total exposure time and calculations in section 10)

Preferred dates: May 1st to June 30th

Impossible dates: June 30 to August 31

**7. Justification for scheduling request:**

SN2014J has the lowest visibility between end of June and end of September, hence, we request for the 3 observations in May and June, with a cadence of  $\sim 25$ -30 days. Hence, we have requested for time-critical observations

We request dark nights for the optical observations since they significantly reduce exposure times in the optical. With 1 mag sky brightness above dark sky, the exposures increase by a factor of 2. We would like to observe the SN in 1" or better seeing.

**8. Instrument:** *check all that apply*

- ☒ HFOSC  
☐ Optical CCD Imager  
☒ TIFR Near-IR Spectrometer (TIRSPEC)

**9. Mode of Observation:** *check all that apply*

- ☒ Imaging   ☐ Spectroscopy

**10. Brief description of observations:**

We request observations of the target at intervals of 25-30 days, with the first epoch in May. For each observation date, we would like to observe the SN in the B to K filters with the HFOSC(BVRI) and TIRSPEC (JHK) instruments

The total number of observations requested is 3 epochs.

To calculate the exposures in the optical, we normalise the maximum light observations of SN2001el (a well-observed normal Ia) to the peak of SN2014J in the *BVRI* filters. As a result, we use the predicted fading of SN2001el to get the magnitudes at these late phases. We summarise the exposure times in table. Note: we use the Liverpool Telescope's (LT) exposure time calculator (ETC) to get these estimates since there isn't an ETC for HFOSC and the LT has a similar diameter of the primary mirror. We calculate the exposure times for a signal to noise that gives us an error of  $\sim 0.03$  mag for

	Filter	Magnitude	Exp. Time (s; dark, 1")
each observation.	B	21.3	300
	V	20.3	100
	R	20.6	150
	I	18.0	30

For the optical the total exposure time is close to 10 mins. We estimate overheads of 10 mins.

For calculations in the NIR we use the TIRSPEC exposure time calculator. For our desired signal to noise, we require 5 dithers of 6 frames with 15s exposures. Since, we require off source images for sky subtraction, we would like to split this observation into two sequences of 5 dithers with 3 frames. Hence, we would obtain a total exposure time of 450s We multiply this by a factor of 3 to get the overheads for the on-source and off-source sequences. Hence, the total time in each filter for the observation will be close to 25 minutes. Thus, we require 1.5 hours for each epoch.

The total time for each epoch, optical + IR is a little less than 2 hours.

Hence, the total time requested for the semester is  $3 \times 2 = 6$  hours.

We would like to mention that this proposal is part of an observing campaign to compile a detailed late time dataset of SN2014J at very late epochs, starting with observations at  $\sim +450$  days and extending out to epochs beyond +700 days, for which there are almost no observations of a normal SNIa.

For this proposal, the SN will in the phase range between +450-+580 days during the observing cycle. We request for 3 epochs of observation in order to strictly constrain the evolution in the optical and NIR, within this phase range.

**11. Plans for data reduction and analysis:** We plan to use the available reduction software for TIRSPEC and HFOSC to reduce the images . We have downloaded the photometry templates for M82 for accurate host galaxy subtraction. We currently have routines ready for the bolometric light curve calculation which have also been tested on data for other projects (Dhawan et al. in prep).

---

## 12. Instrument Resource Requirements:

### HFOSC

**Broad Band Filters:**  $\bigcirc$  U  $\otimes$  B  $\otimes$  V  $\otimes$  R  $\otimes$  I  $\bigcirc$  I<sub>c</sub>  $\bigcirc$  z

### TIRSPEC

**Broad Band Filters:**  $\otimes$  J  $\otimes$  H  $\otimes$  K<sub>s</sub>

## 13. List of objects: (essential)

Name	RA (hh mm ss)	Dec (dd mm ss)	Epoch	V mag	size*
SN2014J	09 55 42.12	+69 40 25.9	J2000	20.3	N/A

\*for extended objects

**14. Scientific Justification:** Type Ia supernovae (SN Ia) are thermonuclear explosions of white dwarfs in a binary system. It is known that their light curves are powered by the decay of  $^{56}\text{Ni}$  at early times and its daughter nuclide  $^{56}\text{Co}$  at late epochs (Colgate & McKee [1969]; Arnett [1982]). There are, however, several unanswered questions regarding the physics of SN Ia. For e.g., the nature of the binary companion is unknown, with two leading possibilities of a main sequence/red giant (single-degenerate) or another white dwarf (WD; double degenerate) companion. Recent studies have also shown that SN Ia do not all explode at the Chandrasekhar mass, as was originally thought, and a significant fraction of them are sub-Chandrasekhar explosions (Stritzinger et al. [2006], Scalzo et al. [2014]) which challenges the notion of all SN Ia being Chandrasekhar mass explosions and indicates a possible diversity in the progenitor channel.

#### **Progenitor Scenarios and Energy Deposition:**

At late times, the ejecta is transparent to the  $\gamma$  rays and most of the light curve is powered by positrons. The extent of positron trapping in the ejecta is determined by the strength and configuration of the magnetic field in the ejecta, with a stronger, tangled field yielding a higher extent of positron trapping. Detonation models of sub-Chandrasekhar WDs predict a weak, radially combed magnetic field, whereas Chandrasekhar mass explosions, due to high turbulence, predict a stronger, tangled magnetic field (Ruiz-Lapuente & Spruit 1998). Hence, we can use the positron escape fraction to constrain the explosion scenario for SN2014J. The late time (pseudo-)bolometric light curve is sensitive to the extent of positron trapping, with a flatter slope indicating more positrons are trapped. These observations, in principle, can also be used to constrain the contribution these positrons make to the galactic 511 keV line (Milne et al. [1999]).

The late light curve can also be used to estimate the total  $^{56}\text{Ni}$  mass by fitting an energy deposition function for the  $^{56}\text{Co}$  decay. This method for estimation has been used for very few SN Ia, for e.g. SN2003hv (Leloudas et al. [2009]). However, SN2003hv does not have NIR observations in the phase range of +60–+200 days, which are publicly available for SN2014J (e.g. Johansson et al. [2014]). Moreover, we propose to sample the late light curve for SN2014J at a higher cadence than for SN2003hv, in order to obtain a more accurate estimate of the mass of  $^{56}\text{Ni}$ . SN2014J also has independent estimates of  $^{56}\text{Ni}$  from near-maximum Ultraviolet (UV) to NIR photometry, the first ever  $\gamma$  ray detection of  $^{56}\text{Co}$  lines for an SN Ia (Churazov et al. [2014]), and an estimate using the timing of the NIR second maximum (Dhawan et al. in prep). The comparison of  $^{56}\text{Ni}$  measurements from early and late time data allow us to have a strong handle on the amount of flux emitted outside the UV to NIR region of the spectrum at late times.

#### **Infrared Catastrophe and Late light curve morphology:**

At late times the ejecta temperature is predicted to drop below a threshold value and hence, lead to a dramatic shift in the excitation state, such that only the ground state levels are populated. This can be observed as a significant decrease in optical and NIR flux, known as the Infrared Catastrophe (IRC; Axelrod 1980). Theoretical models predict an IRC to occur at  $\sim +500$  days. The proposed observations will allow us to decipher whether an IRC has occurred in the ejecta., which probes the ejecta distribution and properties like ionisation and excitation state. We have also been awarded time on GNIRS (Gemini-North) to obtain late time NIR spectra, contemporaneous with the proposed photometric observations. This will help us in understanding the ionisation evolution and the line morphologies will be sensitive to the asymmetries in the ejecta distribution. The late time NIR light curve flattens after  $\sim +200$  d out to  $\sim +700$  d. The brightness of the SN in the proposed observations will allow us to plan further optical and NIR photometry, in Cycle3 of 2015 to obtain data out to  $\sim +700$  days with a high cadence, an observation that has never previously been carried out for an SN Ia. This will allow us to understand the late NIR contribution to the bolometric light curve and the morphology at epochs  $\sim 2$  years post-maximum.

## References

- [1] Arnett W. D., 1982, *ApJ*, 253, 785
- [2] Axelrod T. S., 1980, *PhDT*
- [3] Colgate S. A., McKee C., 1969, *ApJ*, 157, 623
- [4] Johansson J., et al., 2014, *arXiv*, arXiv:1411.3332
- [5] Kerzendorf W. E., Taubenberger S., Seitenzahl I. R., Ruiter A. J., 2014, *ApJ*, 796, LL26
- [6] Leloudas G., et al., 2009, *A&A*, 505, 265
- [7] Milne P. A., The L.-S., Leising M. D., 1999, *ApJS*, 124, 503
- [8] Ruiz-Lapuente P., Spruit H. C., 1998, *ApJ*, 500, 360
- [9] Scalzo R., et al., 2014, *MNRAS*, 560
- [10] Stritzinger M., Leibundgut B., Walch S., Contardo G., 2006, *A&A*, 450, 241

**14a. PhD project Outline:** *This proposal is intended to support part of a PhD project. The focus of the project is to understand the Near Infrared behaviour of Type Ia supernovae, with an emphasis on late time behaviour. The near infrared offers complementary information to the optical for understanding the physical processes in Type Ia SNe. They appear to be more uniform at longer wavelength with a scatter  $< 0.2$  mag. We have simulated light curves at high redshift for future space-based missions to forecast the constraints on cosmological parameters.*

*Another, very interesting feature in the NIR light curves is the appearance of a second maximum, around 2 weeks after maximum light in  $B$ -band. We investigated the relation between the second maximum (specifically the timing, termed as  $t_2$ ). We found correlations between  $t_2$  and the optical decline rate ( $\Delta m_{15}$ ). We also found that the  $t_2$  is related very strongly to the post-maximum colour evolution in the optical. These correlations point to a common origin of this diversity from the large range of  $^{56}\text{Ni}$  produced in the SNe. Hence, the NIR light curve at  $\sim 2 - 3$  weeks post maximum is a probe of the total radioactive nickel synthesised in the ejecta.*

*In the transitional phase (a few months after maximum light), the SN is getting progressively more transparent to the  $\gamma$  rays and the light curve has an increasing contribution from positrons. Along with our conducted study of the second maximum, we would like to learn about the SN from observations in the nebular phase. At these phases, we can probe the inner regions of the ejecta. While line profile measurements from spectroscopy can shed light on the ejecta distribution, the photometric measurements at  $> +400$  days are key in uncovering global parameters of the ejecta.*

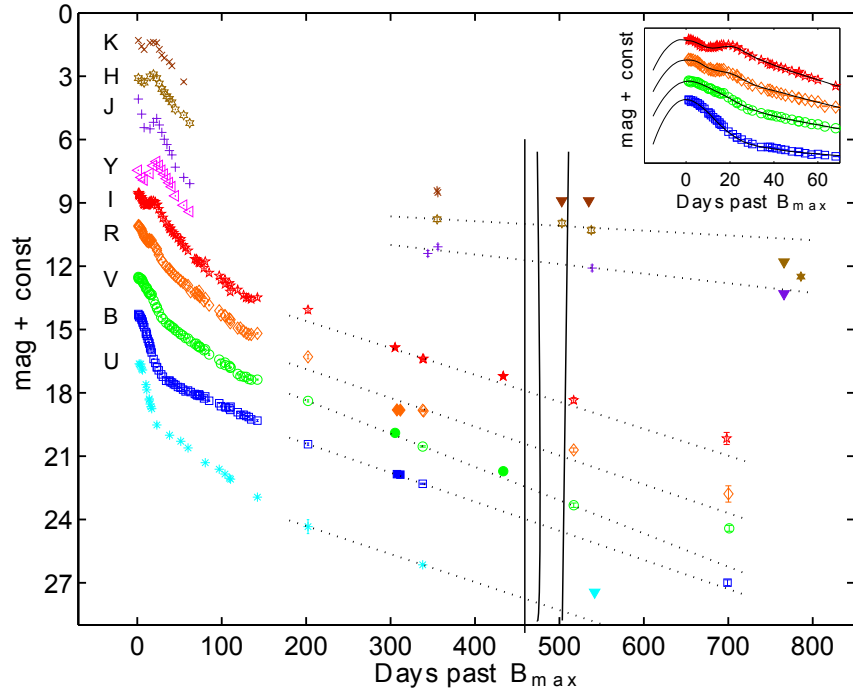


Figure 1: Multi-band observations of SN2003hv (Leloudas et al. [2009]). The NIR light curve starts to flatten at  $\sim +200$  days and there is an indication that the light curve declines again after  $\sim +700$  d. The black lines show the 3 epochs at which we propose to observe SN2014J



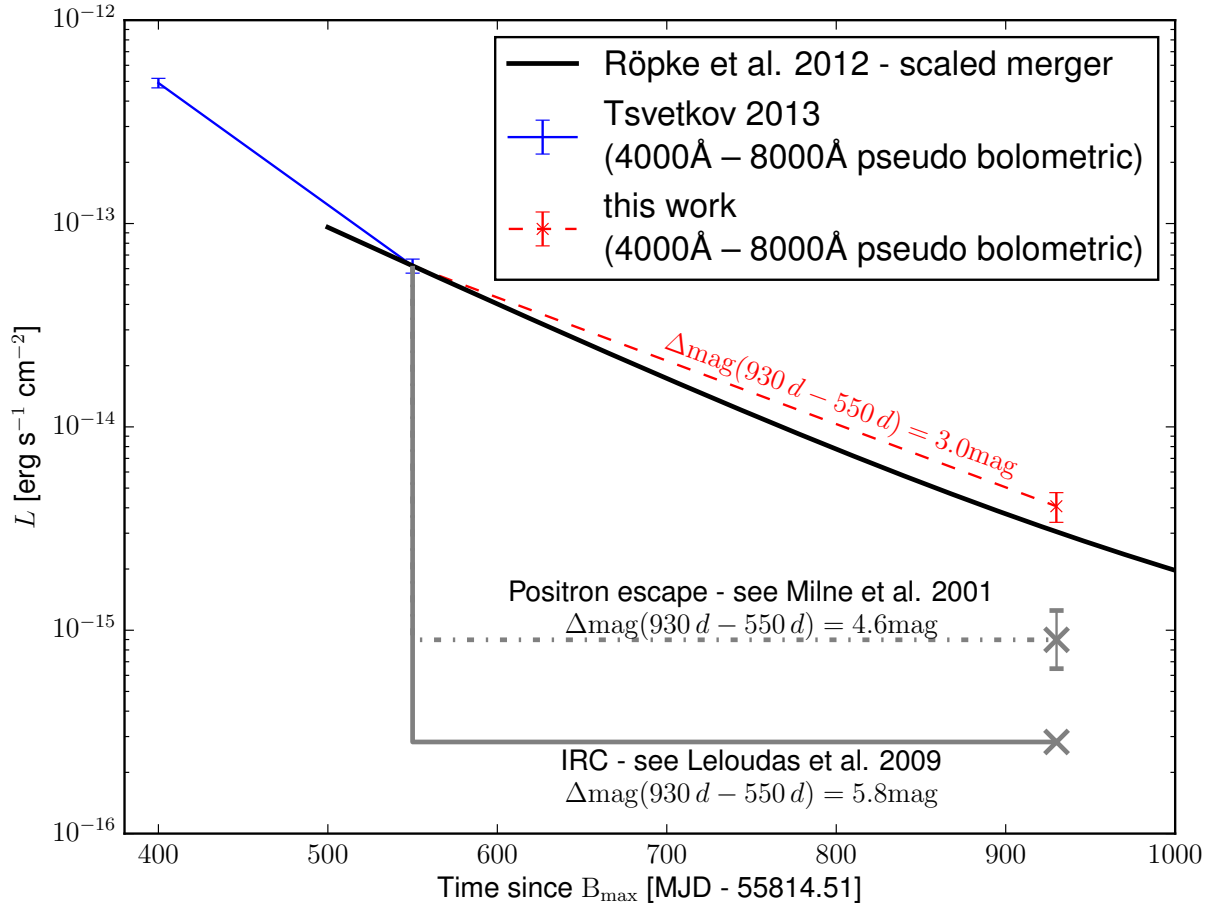


Figure 2: The (pseudo-) late time light curve of SN2011fe, the closest SN Ia before SN2014J with an excellent dataset (Kerzendorf et al. [2014]). Although models predict an Infrared Catastrophe (shift in flux from optical and NIR to mid-IR) after  $\sim +500$  d, there is no evidence in the observations for the presence of such a shift. We note that this is only an optical pseudo-bolometric light curve (*griz* filters). We propose to obtain optical and NIR observations for SN2014J, to better estimate the true bolometric flux.