

GRFolres: A code for modified gravity simulations in strong gravity

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The following brief overview has been prepared as part of the submission of the GRFolres code¹ to the Journal of Open Source Software.

Summary

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Gravitational waves (GWs) are generated by the mergers of dense, compact objects like black holes (BHs) and neutron stars (NSs). This provides an opportunity to study the strong field, highly dynamical regime of Einstein's theory of general relativity (GR) at higher curvature scales than previous observations ([Arun & others, 2022](#); [T. Baker et al., 2015](#); [Barack & others, 2019](#); [Barausse & others, 2020](#); [Gnocchi et al., 2019](#); [Perkins, Yunes, et al., 2021](#)). It is possible that at such scales modifications to GR may start manifest. However, in order to detect such modifications, we need to understand what deviations could look like in theories beyond GR, in particular in the merger section of the signal in near equal mass binaries, which are key targets of the LIGO-Virgo-KAGRA network of detectors (and their future 3G successors). Such predictions necessitate the use of numerical relativity (NR), in which the (modified) equations of GR are evolved from an initial configuration several orbits before merger, through the merger period and the subsequent “ringdown”, during which the gravitational wave signal can be extracted near the computational boundary.

Current waveforms are tested for consistency with GR by measuring parameterised deviations to the merger, inspiral and ringdown phases ([Abbott & others, 2021](#); [Carson & Yagi, 2020a](#); [Cornish et al., 2011](#); [Krishnendu & Ohme, 2021](#); [Maggio et al., 2022](#)), and not by comparison to any particular theories. If we obtain predictions for specific models, we can check whether such parameterised deviations are well-motivated and consistent in alternative theories of gravity ([Arun & others, 2022](#); [Carson & Yagi, 2020b, 2020c](#); [Johnson-McDaniel et al., 2022](#); [Okounkova et al., 2023](#); [Perkins, Nair, et al., 2021](#); [Shiralilou et al., 2022](#)), and the potential to extract model parameters from data.

There are many ways to modify GR, one of the simplest being to couple an additional scalar degree of freedom, which may (if certain conditions are satisfied) result in so-called “hairy” stationary black hole solutions; that is, black holes with a stable, non trivial configuration of the scalar field around them (see ([Doneva, Ramazanoğlu, et al., 2022](#)) for a review). An example of this is the class of Horndeski models ([Horndeski, 1974](#)). Cubic Horndeski theories

¹Folres (pronounced *fol-res*) is a word meaning covers or linings in the Catalan language. It has a specific application in the tradition of *Castells* (Human Towers), denoting the second layers of reinforcement above the base *pinya*. We use it here in analogy to our understanding of effective field theories (EFTs) of gravity as an infinite sum of terms organised as a derivative expansion, in which the first one corresponds to GR (with up to 2 derivatives), and the second one to modified theories up to 4 derivatives, which are those that we are able to simulate with GRFolres.

39 have been studied in Figueras & França (2022) and an implementation of this is included in
40 GRFolres. Another more general example within the Horndeski models is the four-derivative
41 scalar-tensor theory (4 ∂ ST), which is the most general theory with up to fourth powers of the
42 derivatives (but still second order equations of motion). Despite their relative simplicity, they
43 have lacked well-posed (and thus numerically stable) formulations until relatively recently.

44 An important breakthrough was made in 2020 by Kovács and Reall, who showed that Horndeski
45 theories are indeed well-posed in a modified version of the harmonic gauge (Kovács & Reall,
46 2020b, 2020a) – a particular coordinate system already used in NR. Subsequently, several specific
47 theories within these classes were probed in their highly dynamical and fully non-linear regimes
48 (Corman et al., 2023; East & Pretorius, 2022; East & Ripley, 2021b, 2021a). The extension
49 of the results of (Kovács & Reall, 2020b, 2020a) to the alternative “singularity avoiding”
50 coordinates in (Aresté Saló et al., 2022, 2023; Doneva et al., 2023) offers an alternative gauge
51 in which to probe questions of hyperbolicity, and may offer stability advantages for certain
52 cases such as unequal mass ratios, as studied in (Corman et al., 2023). Numerical work on
53 these theories is still in the early stages of development and many technical details on their
54 numerical implementation need to be further investigated. Equally, many scientific questions,
55 concerning our accurate understanding of binary black holes’ phenomenology in alternative
56 theories of gravity and their implications for tests of GR, also remain unanswered.

57 The goal of GRFolres is to meet this need for further research, and to provide a model code to
58 help others develop and test their own implementations. The code is based on the publicly
59 available NR code GRChombo (Andrade et al., 2022; Clough et al., 2015), which itself uses
60 the open source Chombo framework (Adams & others, 2015) for solving partial differential
61 equations (PDEs).

62 In the following sections we discuss the key features, motivations, and applications of the code.

63 Key features

64 GRFolres inherits many of the features of GRChombo and Chombo. Here we list the key
65 features.

- 66 ■ Stable gauge evolution - The code implements the modified moving puncture gauge that
67 ensures a well-posed evolution in the weak coupling regime, as proposed in (Aresté Saló
68 et al., 2022). The precise form of the gauge and its parameters can be changed and
69 the standard moving puncture gauge is safely recovered by setting certain parameters to
70 zero.
- 71 ■ Modified gravity theories - The currently available theories in the code are 4 ST and cubic
72 Horndeski. The code is templated over the theory (in the same way that GRChombo
73 is templated over a matter class) so that it can easily be changed without major code
74 modifications. The code also provides an implementation of 4 ST without backreaction
75 onto the metric (but including the possibility of using the new gauge), to enable
76 comparison with previous works using the decoupling limit approximation.
- 77 ■ Accuracy - The fields are evolved with a 4th order Runge-Kutta time integration and
78 their derivatives calculated with the same finite difference stencils used in GRChombo
79 (4th and 6th order are currently available).
- 80 ■ Boundary Conditions - GRFolres inherits all the available boundary conditions in GR-
81 Chombo, namely, extrapolating, Sommerfeld (radiative), reflective and periodic.
- 82 ■ Initial Conditions - The current examples use solutions that approximately or trivially
83 solve the modified energy and momentum constraints of the theory. An elliptic solver
84 for more general configurations is under development, using a modified CTTK formalism
85 (Aurrekoetxea et al., 2023; Brady et al., 2023).

- 86 ■ Diagnostics - GRFolres has routines for monitoring the constraint violation and calculating
87 the energy densities associated with the different scalar terms in the action, as discussed
88 in (Aresté Saló et al., 2022, 2023; Doneva et al., 2023). Other diagnostics can be added
89 as required. We also extract data for the tensor and scalar gravitational waveforms.
- 90 ■ C++ class structure - Following the structure of GRChombo, the GRFolres code is also
91 written in C++ and uses object oriented programming (OOP) and templating.
- 92 ■ Parallelism - GRChombo uses hybrid OpenMP/MPI parallelism with explicit vectorisation
93 of the evolution equations via intrinsics, and is AVX-512 compliant.
- 94 ■ Adaptive Mesh Refinement - The code inherits the flexible AMR grid structure of
95 Chombo, which provides Berger-Oliger style (M. J. Berger & Oliger, 1984) AMR with
96 block-structured Berger-Rigoutsos grid generation (M. Berger & Rigoutsos, 1991).
97 Depending on the problem, the user may specify the refinement to be triggered by the
98 additional degrees of freedom, i.e. the scalar field, or those of the metric tensor.

99 Statement of Need

100 As far as we are aware there is currently no other publicly available code that implements the
101 4∂ ST theory of modified gravity or the cubic Horndeski theory in (3+1)-dimensional numerical
102 relativity.

103 There is at least one private code, based on the PAMR/AMRD and HAD (East et al., 2012;
104 Neilsen et al., 2007) infrastructure, that was used in the first works to successfully implement
105 the modified general harmonic gauge for 4∂ ST (Corman et al., 2023; East & Pretorius, 2022;
106 East & Ripley, 2021b, 2021a). Since this code uses a Generalised Harmonic Coordinates (GHC)
107 formulation, it necessitates excision of the interior of black holes, which can be difficult to
108 implement in practice. As a consequence, many groups in the numerical relativity community
109 have opted to use singularity avoiding coordinates such as the BSSN (Baumgarte & Shapiro,
110 1998; Nakamura et al., 1987; Shibata & Nakamura, 1995), Z4C (Bernuzzi & Hilditch, 2010;
111 Bona et al., 2003) or CCZ4 (Alic et al., 2012, 2013) formulations in the puncture gauge (J. G.
112 Baker et al., 2006; Campanelli et al., 2006), which do not require the excision of the interior of
113 black holes from the computational domain. In GRFolres, we use the results of (Aresté Saló et
114 al., 2022, 2023; Doneva et al., 2023) to extend the well-posed formulations of modified gravity
115 to singularity avoiding coordinates. This provides an alternative gauge to the modified GHC
116 one used by other groups. Not only does this provide a valuable comparison to their work, but
117 also eliminates the need for excision.

118 There are also a number of (3+1)-dimensional codes that implement the equations for the
119 additional scalar degree of freedom in Einstein-scalar-Gauss-Bonnet without backreaction
120 onto the metric tensor, including one implementation using GRChombo (Evstafyeva et al.,
121 2023), which we have integrated into GRFolres to enable comparison between the methods. In
122 particular, Canuda (<https://bitbucket.org/canuda>) (Witek et al., 2019) which uses the Einstein
123 Toolkit (<http://einstein toolkit.org/>), with its related Cactus (<http://cactuscode.org>) (Loffler
124 & others, 2012; Schnetter et al., 2004) and Kranc (<http://kranccode.org>) (Husa et al., 2006)
125 infrastructure, was used in (Elley et al., 2022; R. et al., 2023; Richards et al., 2023; Silva
126 et al., 2021; Witek et al., 2019). Another implementation is based on the Spectral Einstein
127 Code or SpEC (<http://www.black-holes.org/SpEC.html>) (Pfeiffer et al., 2003), as used in
128 (Okounkova, 2020). A neutron star background was considered in (Kuan et al., 2023) with
129 a modification of SACRA-MPI code (Kiuchi et al., 2017; Yamamoto et al., 2008). Whilst
130 order-reduced methods like those in (Doneva, Vañó-Viñuales, et al., 2022; Elley et al., 2022;
131 Evstafyeva et al., 2023; Okounkova et al., 2019, 2020, 2023; Okounkova, 2020; R. et al., 2023;
132 Richards et al., 2023; Silva et al., 2021; Witek et al., 2019) provide an estimate of the scalar
133 dynamics and associated energy losses, they may miss information about the fully non-linear
134 impact on the metric and suffer from the accumulation of secular errors over long inspirals.

In spherical symmetry several codes have been developed that implement Einstein-scalar-Gauss-Bonnet (a subset of the 4∂ ST theory that we include as an example in GRFolres). In particular, using the NRPy framework (<http://astro.phys.wvu.edu/bhathome>) (Ruchlin et al., 2018) in (Doneva, Vañó-Viñuales, et al., 2022), and the private code of Ripley & Pretorius in (R et al., 2023; Ripley & Pretorius, 2019a, 2019b, 2020a, 2020b), and a modification of the GR1D code (Gerosa et al., 2016; O'Connor & Ott, 2010). There is also the fully nonlinear spherical code developed in (Corelli et al., 2023b, 2023a). Spherical codes provide a useful testing ground in which coordinate ambiguities can be avoided (R et al., 2023), but lack the generality required to study objects with angular momentum, or binary mergers.

Research projects to date using GRFolres

So far the code has been used to study a range of fundamental physics problems, as listed here.

- The test field case was used in (Evstafyeva et al., 2023) to model the scalar waves produced during the ringdown stage of binary black hole coalescence in Einstein-scalar-Gauss-Bonnet, and quantify the extent to which current and future gravitational wave detectors could observe the spectrum of scalar radiation emitted.

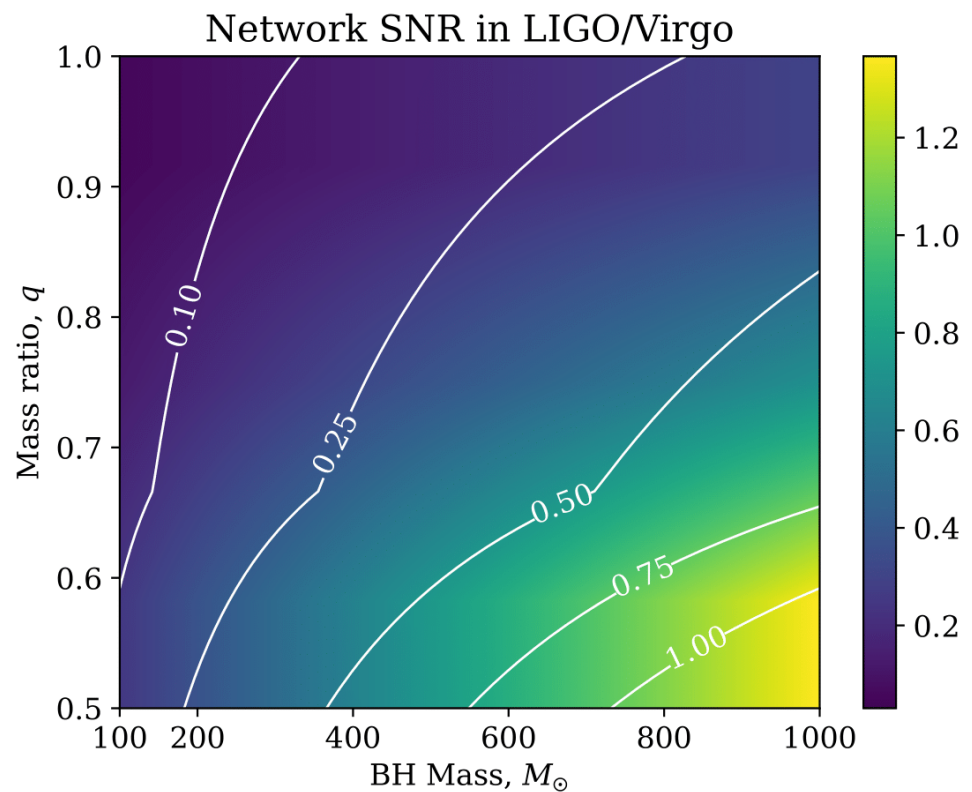


Figure 1: Contour plot of network signal-to-noise ratio (SNR) for the scalar ringdown of a binary black hole (BBH) in Einstein-scalar-Gauss-Bonnet gravity at 1 Gpc as observed by the Virgo, Livingston and Hanford network of detectors at design sensitivity. Taken from (Evstafyeva et al., 2023).

- The regime of validity of effective field theory in collapse and binary evolutions in cubic Horndeski theories were studied in (Figueras & França, 2020, 2022). It was found that the mismatch of the gravitational wave strain can be as large as 10%–13% in the Advanced LIGO mass range for such theories.

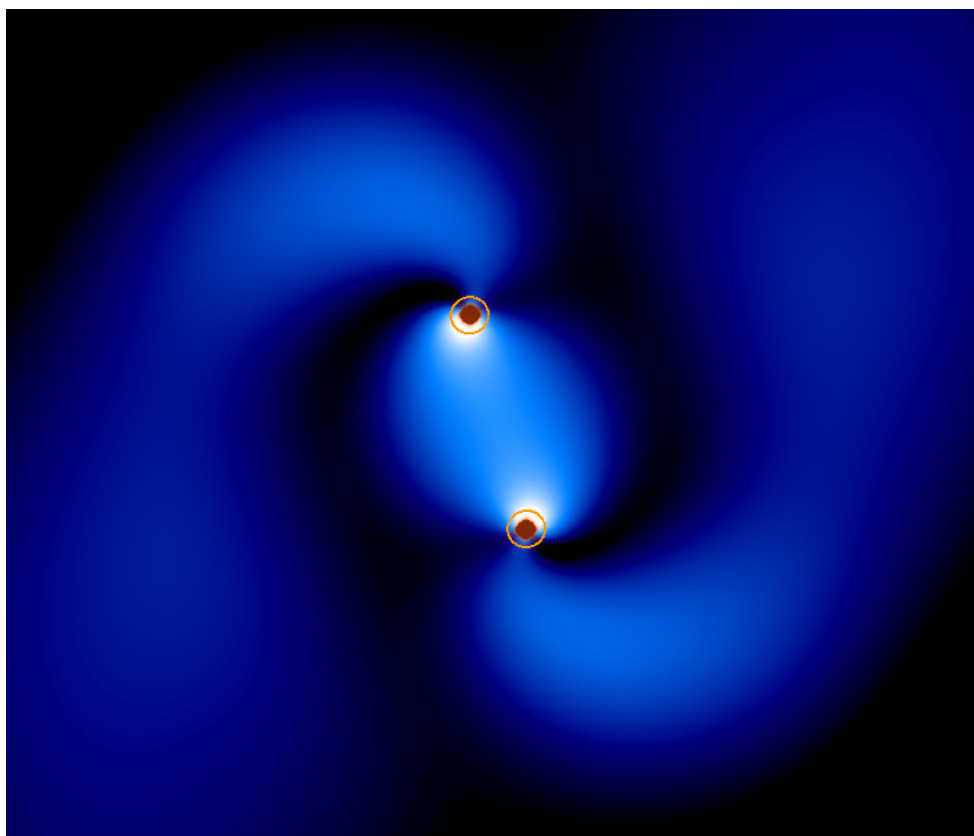


Figure 2: Energy density (in blue) of the scalar field surrounding the binary black holes for the Horndeski theory at a representative instant of time during the inspiral phase. The apparent horizon of the black holes is shown in orange. The region where the weak coupling conditions are larger than one is depicted in brown. Taken from (Figueras & França, 2022).

- 155 ■ In the work (Aresté Saló et al., 2022), the code was developed and tested, with waveforms
- 156 for shift-symmetric theories of Einstein-scalar-Gauss-Bonnet gravity produced for equal
- 157 mass binaries.

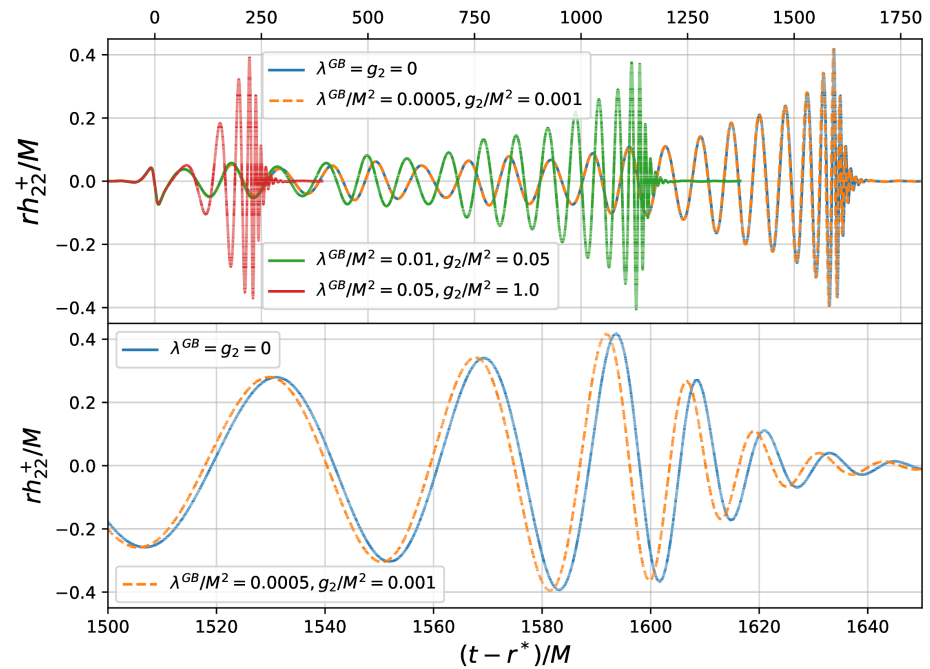


Figure 3: Modified gravity waveforms in 4dST with a shift-symmetric coupling. Taken from (Aresté Saló et al., 2022).

- In the work (Aresté Saló et al., 2023), the studies were extended to binary mergers in theories with spin-induced scalarisation. The clouds formed are dumbbell-like in shape.

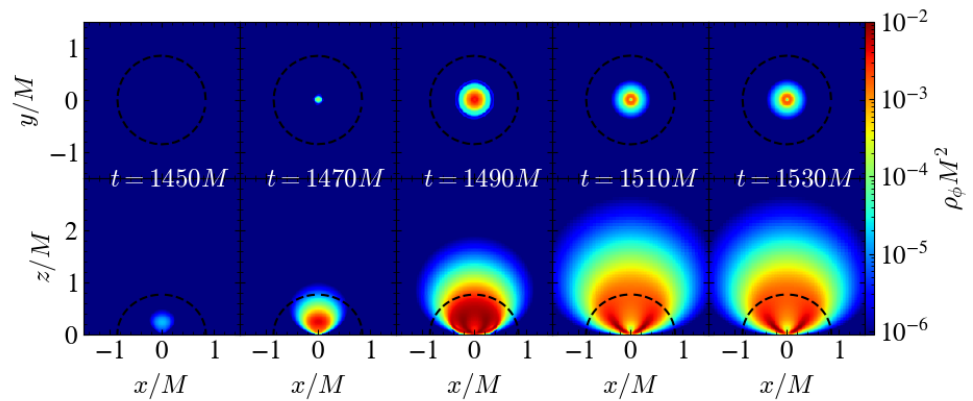


Figure 4: The time evolution of the density of the scalar cloud that develops in Einstein-scalar-Gauss-Bonnet gravity with an exponential coupling, resulting in spin-induced scalarisation. Taken from (Aresté Saló et al., 2023).

- In the work (Doneva et al., 2023), the dependence of the conditions for hyperbolicity and weak coupling were studied for spin-induced scalarisation, and the critical thresholds found for a number of cases.

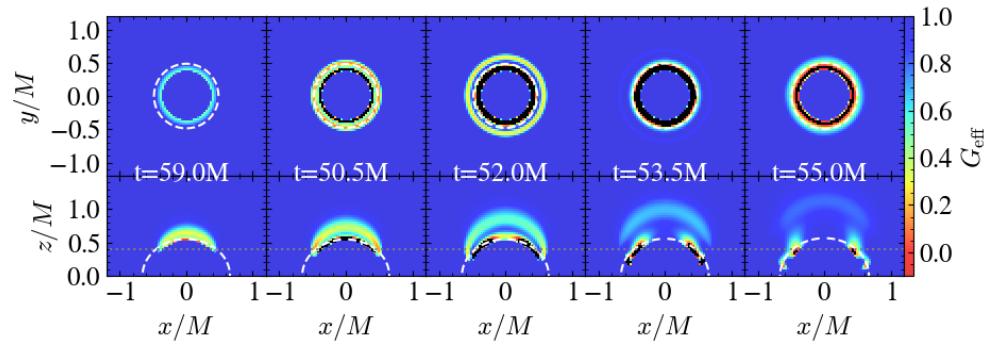


Figure 5: The time evolution of the determinant of the effective metric in a case of spin-induced scalarisation. When the determinant is negative (in black) outside the apparent horizon (depicted with a dashed white line), the theory has become ill-posed. Taken from (Doneva et al., 2023).

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