

Scalar radiation from symmetron field oscillations

PIC1 Scientific project

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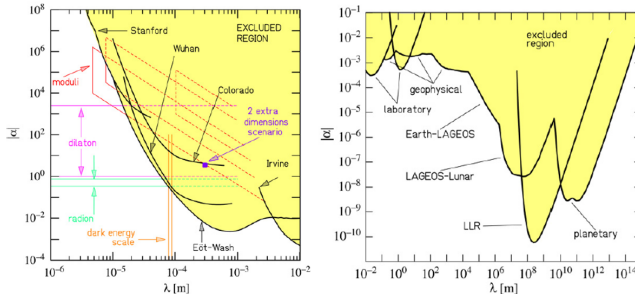
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Introduction

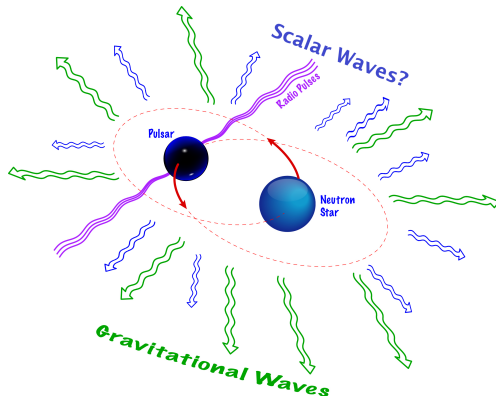
- In the standard model of cosmology, (Λ CDM), the acceleration of the universe is explained by a cosmological constant Λ
- Modified theories of gravity have emerged as an alternative by introducing a new scalar field mediating a fifth force, altering the behaviour of gravity
- General relativity has been tested with high precision, placing constraints on this new force



Adelberger et al. (2009)

Screening mechanisms and scalar waves

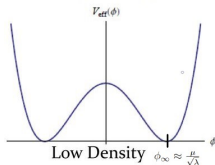
- Screening mechanisms suppress the field in regions of high matter density, hiding fifth forces from local gravity tests
- Due to the coupling of these fields to mass density, radially pulsating astrophysical objects can source scalar waves.
- These waves would cause deviations from General Relativity which could be tested by gravitational wave detectors



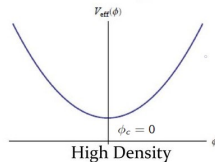
The symmetron

The symmetron screening mechanism works by combining a symmetry breaking potential with a coupling to the local matter density

$$V_{\text{eff}}(\phi) = \frac{1}{2} \left(\frac{\rho}{M_s^2} - \mu^2 \right) \phi^2 + \frac{1}{4} \lambda \phi^4$$



$$\rho \ll \mu^2 M_s^2$$



$$\rho \gg \mu^2 M_s^2$$

$$\phi_\infty \approx \mu / \sqrt{\lambda}$$

$$m_\infty \approx \sqrt{2} \mu$$

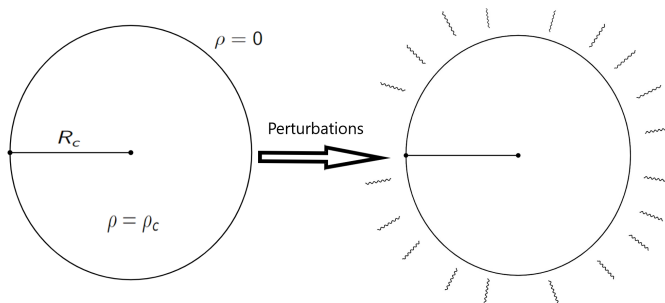
$$\phi_c = 0$$

$$m_c = \sqrt{\frac{\rho}{M_s^2} - \mu^2}$$

The model

Approximations:

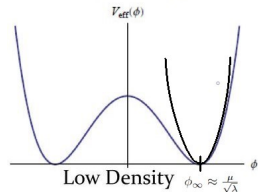
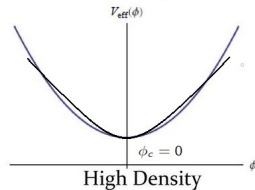
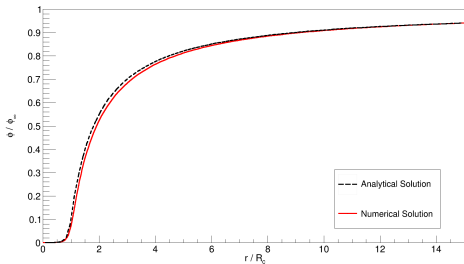
- 1 Newtonian limit
- 2 Constant density and radius
- 3 Radial pulsations with constant amplitude and frequency
- 4 The surrounding region is a perfect vacuum



Static profile

Static equation of motion in spherical coordinates

$$\frac{d^2\phi_0}{dr^2} + \frac{2}{r} \frac{d\phi_0}{dr} = \left(\frac{\rho_0}{M_s^2} - \mu^2 \right) \phi_0 + \lambda \phi_0^3$$



Perturbations profile

Introducing the perturbations into the system

$$\phi(r, t) = \phi_0(r) + \delta\phi(r, t)$$

$$R(t) = R_c + \Delta R \sin(\omega t)$$

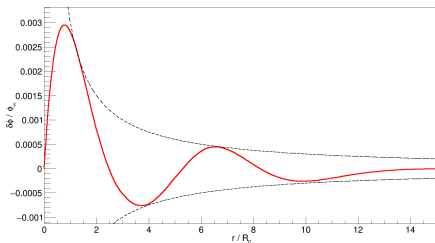
$$\Rightarrow \rho(r, t) = \rho_0(r) \left(1 + 3 \frac{\Delta R}{R_c} \sin(\omega t) \right)$$

$$x \equiv \frac{r}{R_c} \quad \tau \equiv \frac{t}{R_c} \quad \varphi \equiv \frac{\phi}{\phi_\infty}$$

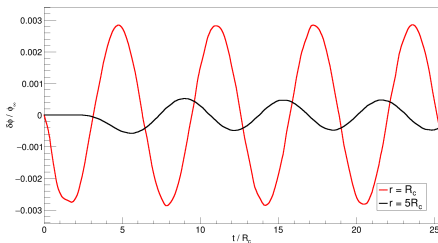
$$-\frac{\partial^2 \delta\varphi}{\partial \tau^2} + \frac{\partial^2 \delta\varphi}{\partial x^2} + \frac{2}{x} \frac{\partial \delta\varphi}{\partial x} = R_c^2 \left[\frac{\delta\rho}{M_s^2} \varphi_0 + \left(\frac{\rho_0}{M_s^2} - \mu^2 \right) \delta\varphi + 3\mu^2 \varphi_0^2 \delta\varphi \right]$$

Time dependent equation of motion

Perturbations profile (cont.)



Perturbations profile at $\tau = 11$ as a function of x

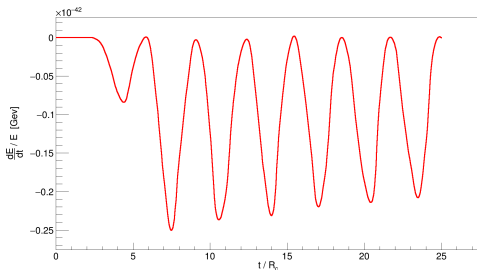


Perturbations profile at $x = 1$ and $x = 5$ as a function of τ

Scalar energy loss

$$\frac{dE}{dt} = \int_{\partial V} \dot{\phi} \partial_i \phi n_i dS = 4\pi R^2 \left. \frac{\partial \phi}{\partial t} \frac{\partial \phi}{\partial r} \right|_{r=R}$$

$$E = \frac{4}{3} \pi R_c^3 \rho_c$$

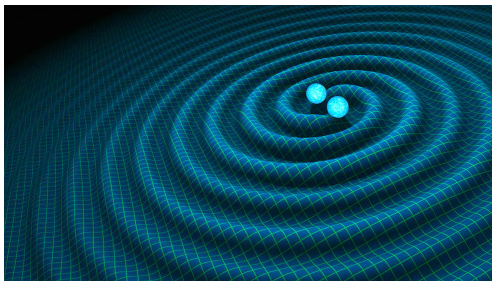


Relative power loss $\frac{1}{E} \frac{dE}{dt}$ over time

$$\left\langle \frac{1}{E} \frac{dE}{dt} \right\rangle \approx 10^{-19} s^{-1}$$

Conclusions

- Radial pulsations of astrophysical objects source scalar waves
- These waves carry energy with them in the form of scalar radiation
- Scalar radiation is subdominant when compared to gravitational waves, making them harder to detect
- A more realistic result can be obtained by considering GR and using a more accurate model of binary systems



Thank you!