Constants/Conversions:

$$c = 3 \times 10^{10}\, cm/s \qquad G = 6.67 \times 10^{-8}\, cm^3/g/s^2$$
$$\sigma = 5.67 \times 10^{-5}\, erg/cm^2/K^4/s \qquad 1\, yr \approx \pi \times 10^7\, sec$$
$$M_\odot = 2 \times 10^{33}\, g \qquad R_\odot = 6.96 \times 10^{10}\, cm$$
$$M_\oplus = 6 \times 10^{27}\, g \qquad R_\oplus = 6.37 \times 10^8\, cm$$

Formulae:

$$T_{disk} = \left(\frac{3GM\dot{M}}{8\pi\sigma R^3}\right)^{1/4} \qquad T_{max} = 0.488\, T_{disk}$$
$$R_s = \frac{2GM}{c^2} \qquad \lambda_{peak} T = 0.2 \text{ cm K}$$
$$F_{self\, grav} \approx \frac{GM^2}{x^2} \qquad F_{tidal} = \frac{2GMmx}{r^3}$$
$$\Omega_{kep} = \sqrt{\frac{GM}{r^3}} \qquad v_{kep} = \Omega_{kep} r$$

Accretion disks around different objects: In general, the amount of matter falling onto an accretion disk and into the central object varies by several orders of magnitude, but a good approximate value for objects accreting in a binary system is $\dot{M} = 10^{-8}\, M_\odot/yr$, and for super-massive black holes is $\dot{M} = M_\odot/yr$.

Results:

Obj.	M ($M_\odot$)	R	$\dot{M}$ ($\frac{M_\odot}{yr}$)	$T_{max}\,(K)$	λ_{peak}	λ range	$F_{\oplus,tidal}$ (dyn)	Ω_{kep} ($\frac{rad}{s}$)	v_{kep} ($\frac{cm}{s}$)
WD	.85	$.0095 R_\odot$	10^{-8}	7.4×10^4	$27\, nm$	UV	6×10^{36}	0.6	4×10^8
NS	1.4	10 km	10^{-8}	10^7	$0.2\, nm$	X-ray	3×10^{45}	1.4×10^4	1.4×10^{10}
BH	3	$3R_s$	10^{-8}	6×10^6	$0.3\, nm$	X-ray	3×10^{44}	4.6×10^3	1.2×10^{10}
SMBH	10^8	$3R_s$	1	10^5	$18\, nm$	UV	2.8×10^{29}	10^{-4}	1.2×10^{10}
MS*	1	1	10^{-8}	2300	$.85\, \mu m$	IR	6×10^{30}	6×10^{-4}	4.4×10^7

$F_{\oplus, self\, gravity} \approx 6 \times 10^{30}$ dynes