

Interstellar Medium (ISM)

Week 9

2025 May 09 (Friday), 3PM

updated 05/08, 22:03

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Density and Metallicity Effects

- **Density Effect:** If the density is high, fewer of the possible cooling lines are above the critical density.

- ▶ Thus, cooling becomes less effective at higher densities, and **the equilibrium temperature of the nebula goes up.**
- ▶ For instance, the temperature of an Orion-like nebula increases from $T_{\text{gas}} = 6600 \text{ K}$ at $n_{\text{H}} = 100 \text{ cm}^{-3}$ to $T = 9050 \text{ K}$ at $n_{\text{H}} = 10^6 \text{ cm}^{-3}$.

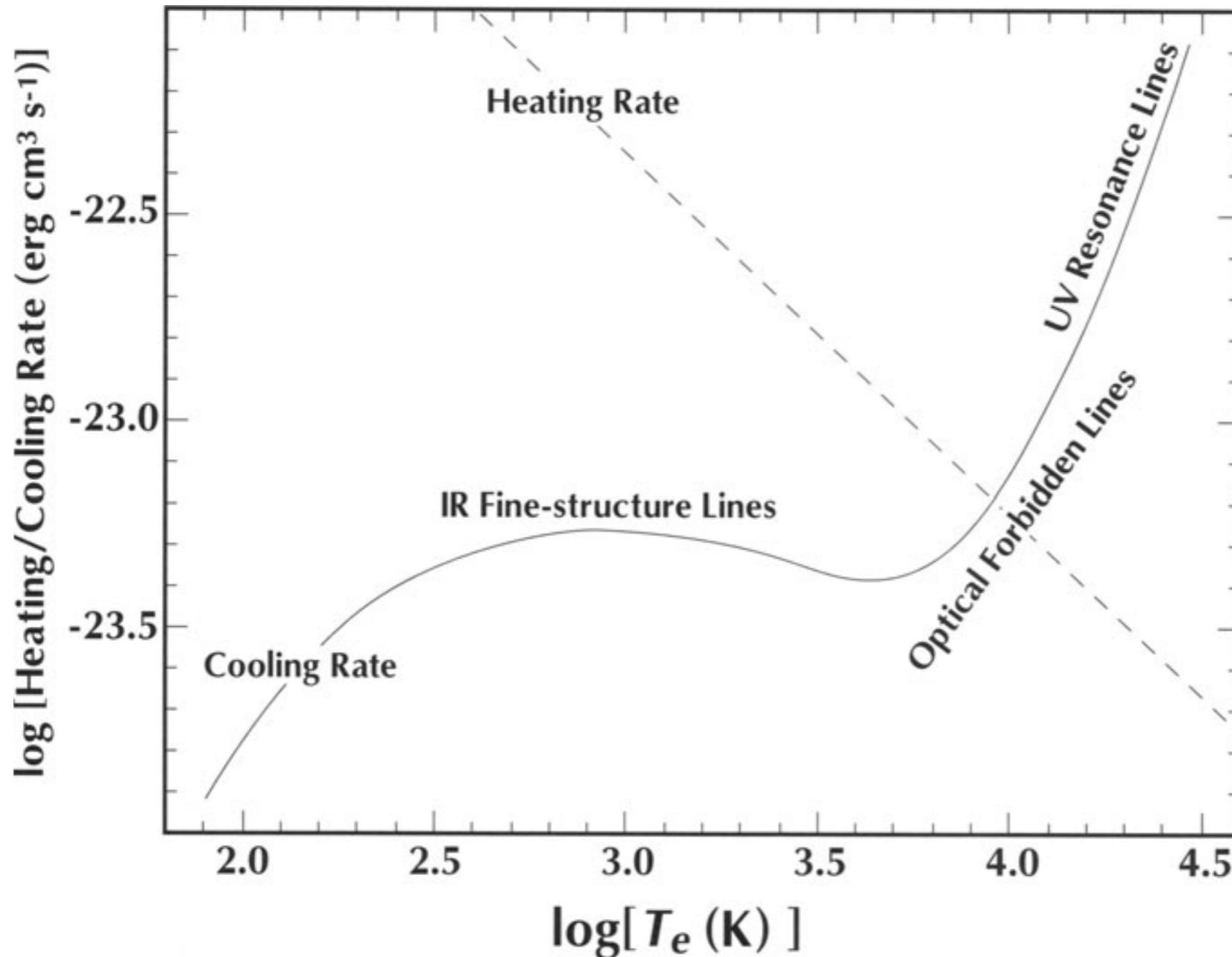
Main contributors to line cooling in H II regions
[Table 4.1 in Ryden]

Name	$\lambda [\text{\AA}]$	$A_{u\ell}$ [10^{-3} s^{-1}]	n_{crit} [10^4 cm^{-3}]
$[\text{O II}]^4\text{S} - {}^2\text{D}$	3726	0.16	1.5
	3729	0.036	0.34
$[\text{N II}]^3\text{P} - {}^1\text{D}$	6548	0.98	6.6
	6583	3.0	6.6
$[\text{O III}]^3\text{P} - {}^1\text{D}$	4959	6.8	68
	5007	20.	68

- **Metallicity Effect:**

- ▶ The equilibrium temperature of a nebula also depends on its metallicity. If the metallicity is lowered, its temperature raises.
- ▶ An Orion-like nebula (around a star with $T_{\text{eff}} = 35,000 \text{ K}$) has a gas temperature of $T_{\text{gas}} \sim 8050 \text{ K}$.
- ▶ If the metallicity is lowered to $Z = Z_{\odot}/10$, its temperature rises to 15,600 K.
- ▶ If the metallicity were zero, the gas temperature would be $T_{\text{gas}} \sim 25,000 \text{ K}$.
- ▶ If the metallicity were 3 times that of the Orion Nebula, its temperature would be $T_{\text{gas}} \sim 5400 \text{ K}$.

Heating and Cooling Function



[Figure 9.5, Dopita]

The cooling function for a fixed ionization state produced by an O star with $T_{\text{eff}} = 40,000$ K as a function of electron temperature.

The heating rate is related to the recombination rate. The equilibrium temperature is defined by the point at which these two curves cross.

At different electron temperatures, different collisionally excited emission lines dominate the cooling rate.

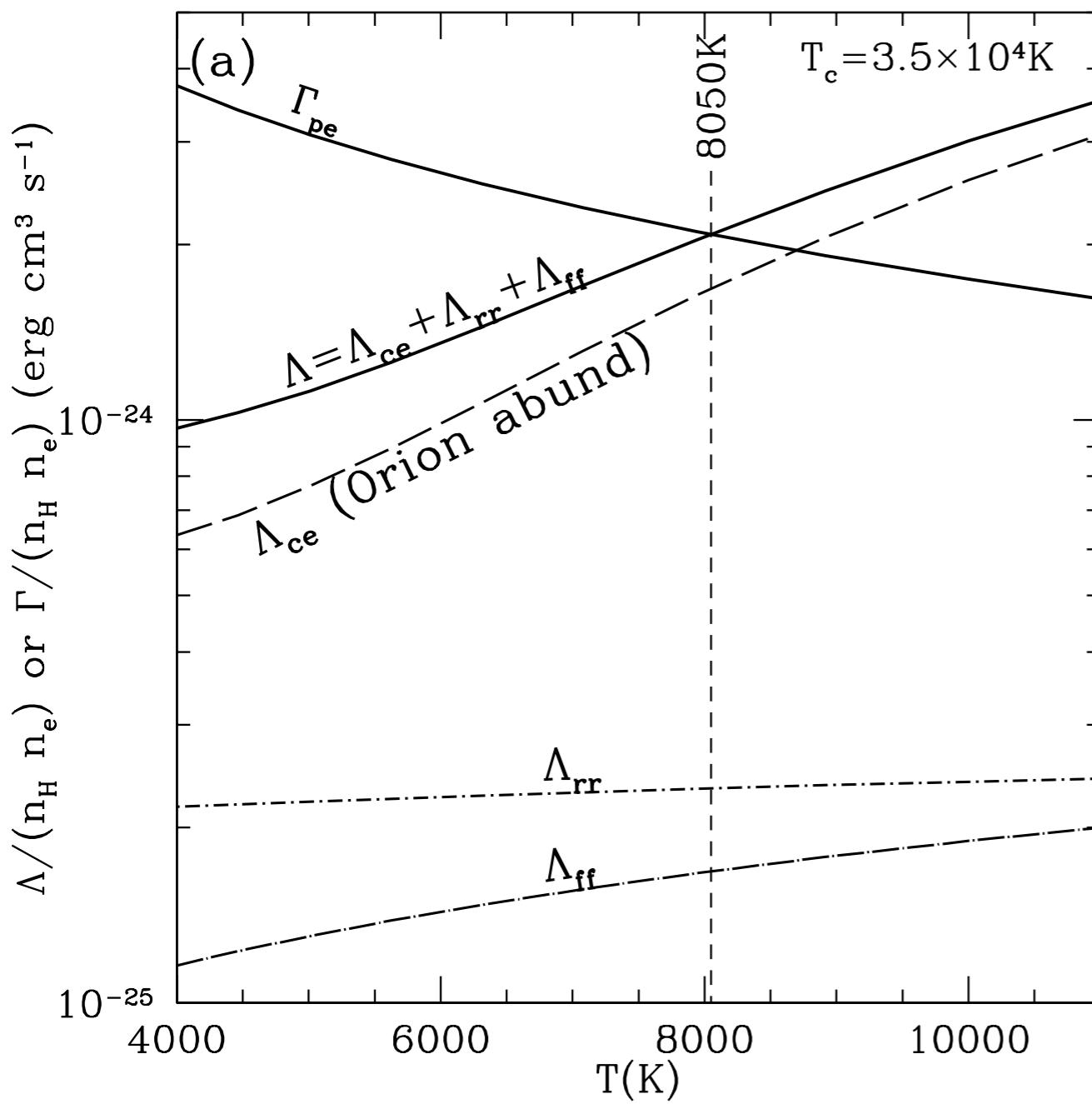
At $T \sim 1000$ K ($kT \sim 0.1$ eV), the cooling is dominated by infrared fine-structure lines, such as [S II] $18.7 \mu\text{m}$ line and, at lower temperatures, the [O I] $63 \mu\text{m}$ and [C II] $158 \mu\text{m}$ lines, which cool the CNM.

At $T \sim 25,000$ K, the cooling is dominated by ultraviolet lines such as Ly α .

At intermediate temperatures, $T \sim 8000$ K, optical forbidden lines from O II, N II, and O III dominate the cooling rate.

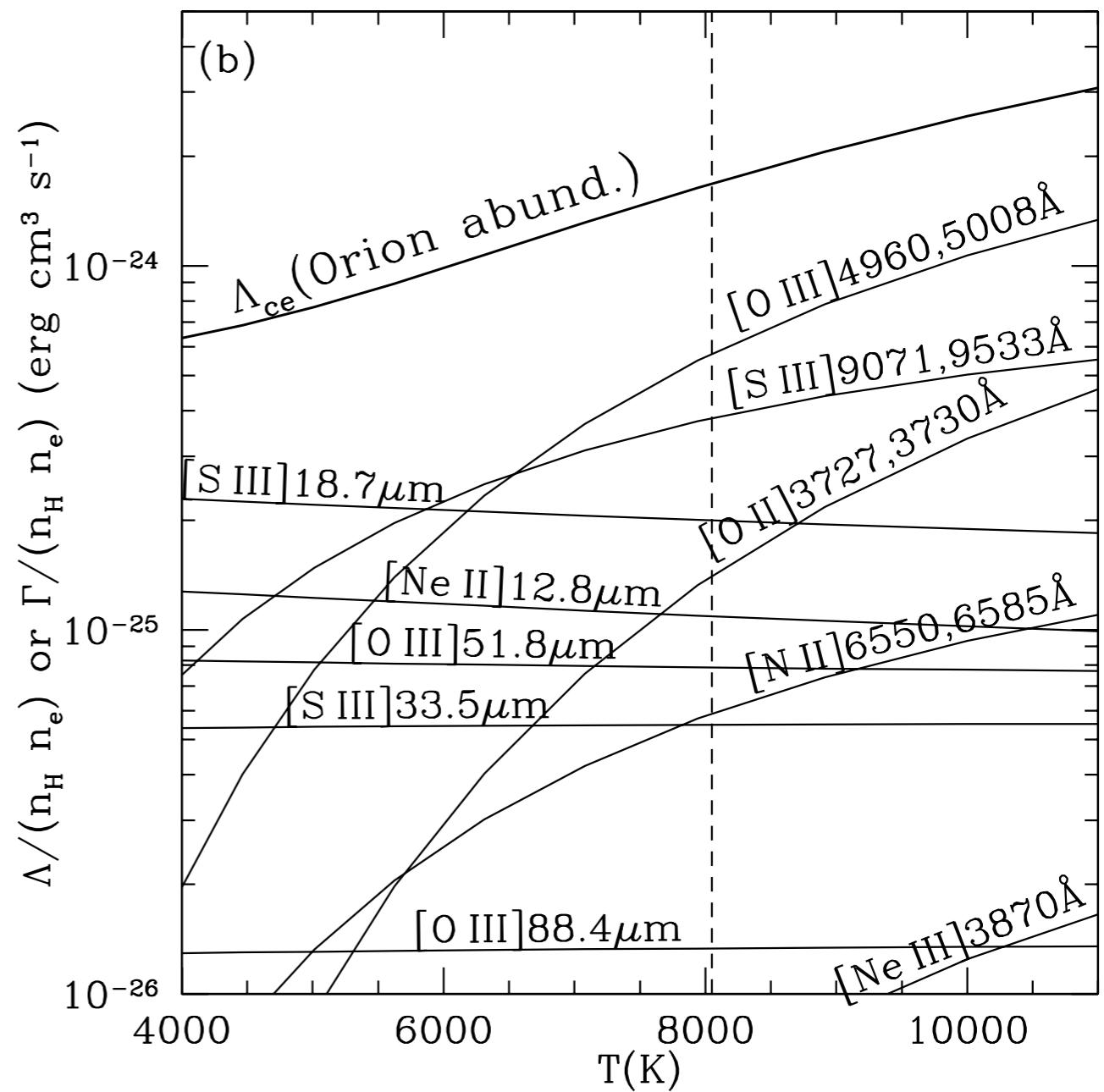
Heating and Cooling function as a function of gas temperature in an H II region with Orion-like abundances and density $n_{\text{H}} = 4000 \text{ cm}^{-3}$.

Heating and cooling balance at $T_{\text{gas}} \sim 8050 \text{ K}$.
pe = photoelectron



Contributions of individual collisionally-excited lines to the cooling function.

[Figure 27.1 in Draine]



Heating and Cooling - Dependence on Metallicity

Heating and Cooling function for different metal abundances

- (a) For an abundance of 10% of that of the Orion Nebula
- (b) For 3 times higher abundance

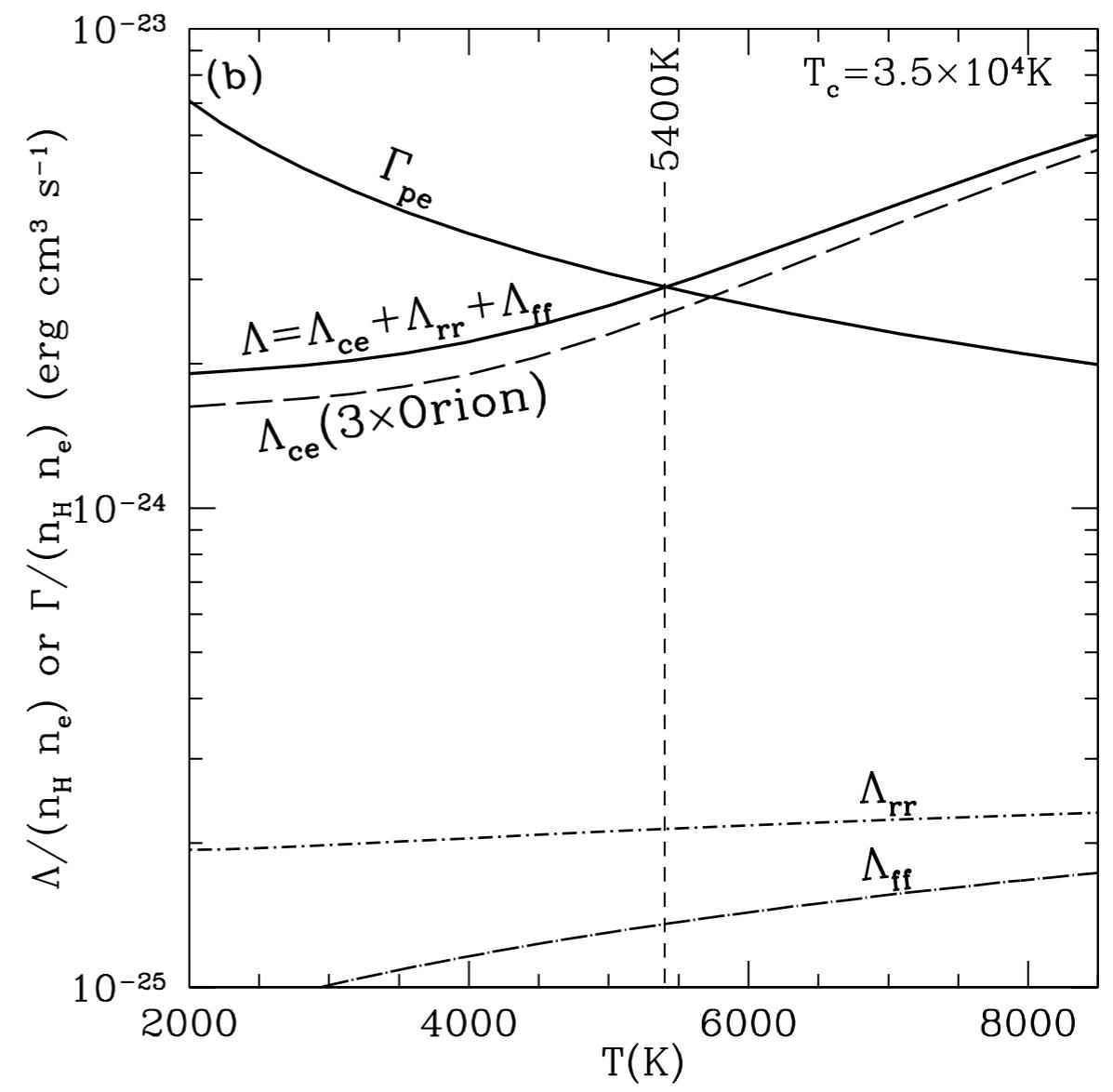
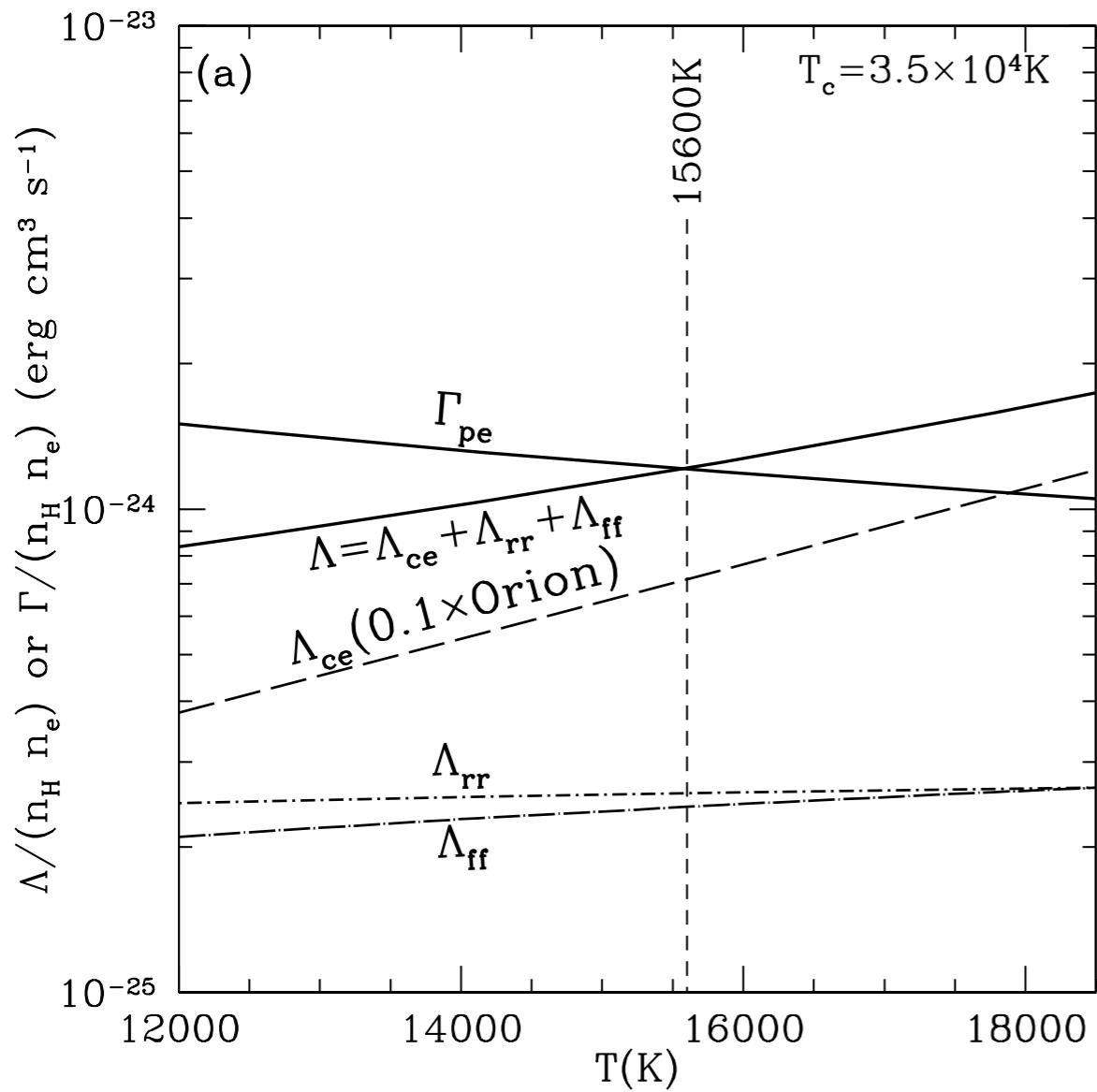


Figure 27.2 in Draine

Heating and Cooling - Dependence on Density

Cooling function for different densities.

The gas is assumed to have Orion-like abundances and ionization conditions.

As the gas density is varied from 10^2 to 10^5 cm^{-3} , the equilibrium temperature varies from 6600 K to 9050 K, *because of the contribution of collisional de-excitation*.

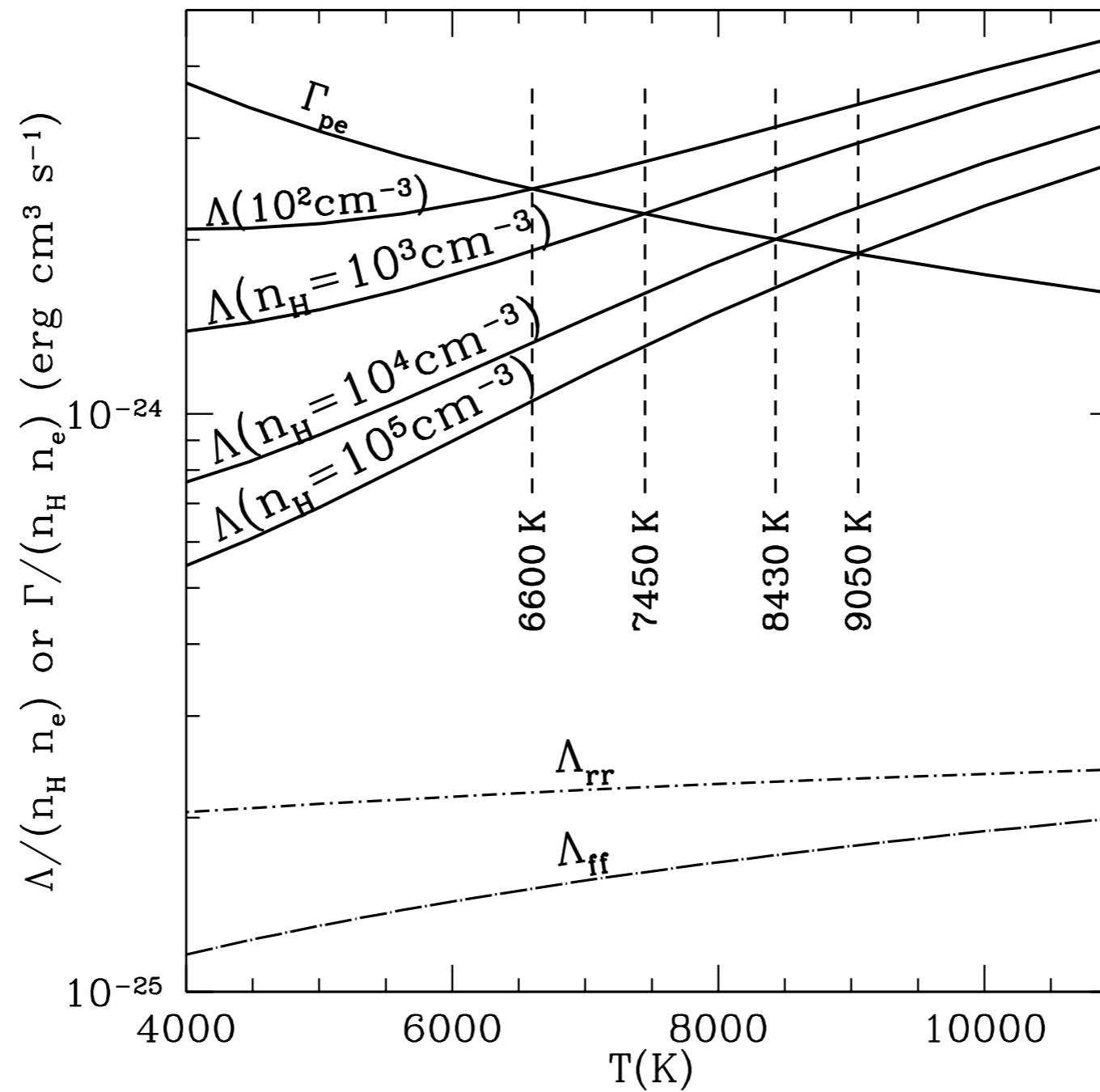
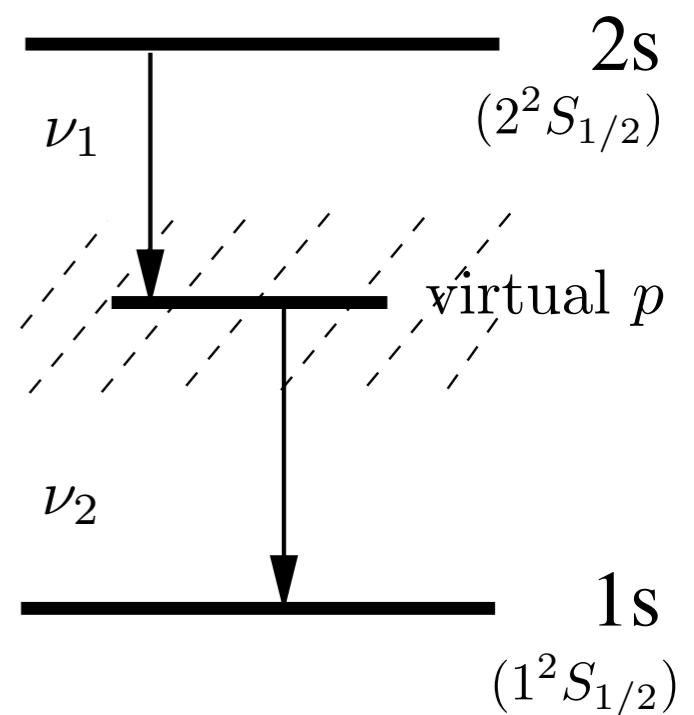


Figure 27.3 in Draine

Additional Cooling Mechanisms

- Thermal Emission of Dust
 - H II regions contain dust which scatters the light of the exciting stars. *Dust grains also absorbs some of the photons emitted by the stars and some of the Lyman α emission that fills the H II region. They re-emit the absorbed energy in the mid- and far-infrared, producing thermal continuum.*
- Two-Photon (Continuum) Emission
 - The emission of radiation from an atomic level can arise through the intermediate of a virtual state. In this case, two photons are emitted, the sum of their energies being equal to the energy of the transition.
 - The probability of this 2-photon emission is small, but it can become the main channel for the de-excitation of a metastable level if collisions are negligible.
 - *This is the case for neutral hydrogen and helium.*

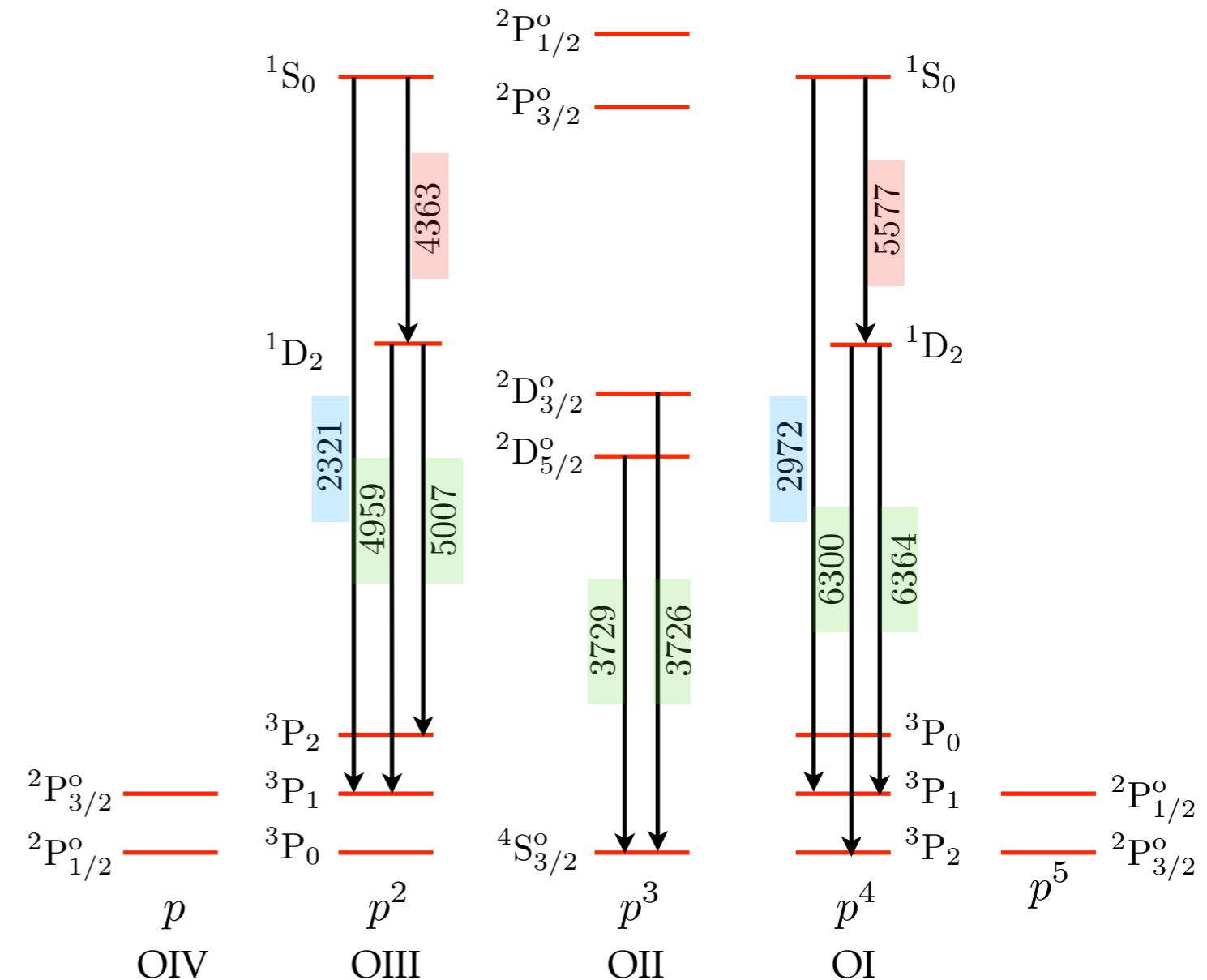


Nebular, Auroral Lines

- Definitions:

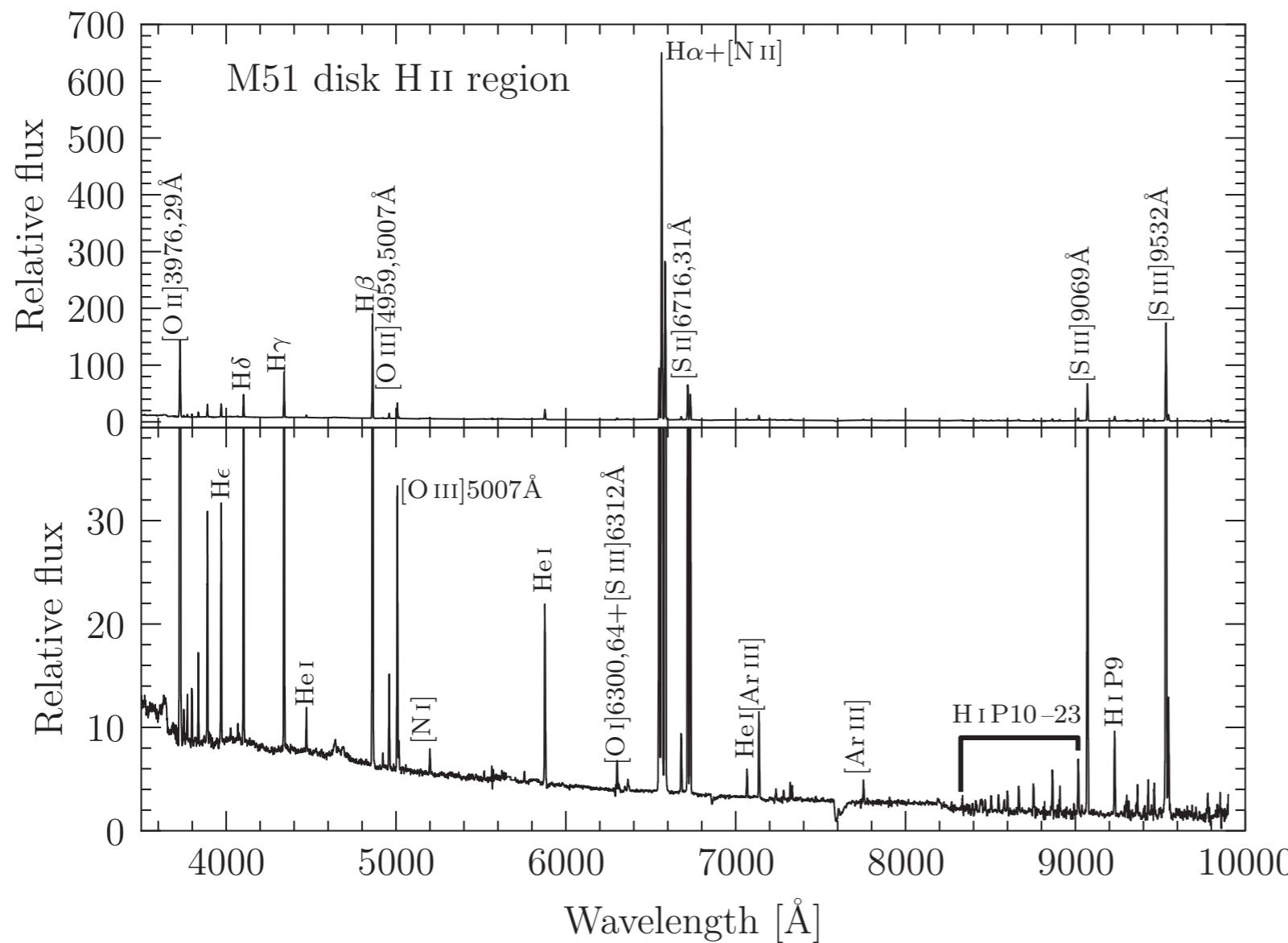
- **Auroral**: the transitions between ***two higher terms*** of configurations p^2 , p^3 , and p^4 are named auroral.
- **Nebular**: the transitions between ***the middle and the lowest terms*** give nebular lines.
- **Transauroral**: the transitions between ***the highest and the lowest terms*** give the transauroral lines.

The term structure for the ground configurations with p , p^2 , p^3 , p^4 , and p^5 outermost shells.
(not to scale)



Temperature, Density & Abundance Diagnostics

- In the figure, the continuum is a mixture of free-bound continuum (from radiative recombination), free-free emission (thermal bremsstrahlung), and two-photon emission.
- If we know enough about the temperature dependence of these continuum radiation, we can estimate the nebula temperature. However, the ***collisionally excited emission lines*** are much stronger than the continuum spectrum.



Spectrum of a disk HII region in the Whirlpool galaxy (M51).

(top) bright lines
(bottom) scaled to show faint lines.

Figure 4.5 [Ryden]
Data from Croxall et al. (2015)

Temperature Determination

- The key to using emission lines to estimate temperature is finding ***two excited states of the same ion whose energy differs by $\sim kT$.*** For nebulae with $T \sim 10^4$ K, this implies energy differences of order of 1 eV or so.
- Atoms and ions with six (or 14) electrons have $2p^2$ (or $3p^2$) as their lowest configuration, and have 3 terms, as shown in the figure. Because the energy difference between the terms are very different. ***The relative strengths of the emission lines will be very sensitive to temperature.***
- The measured intensity ratio can be used to determine the nebula temperature.

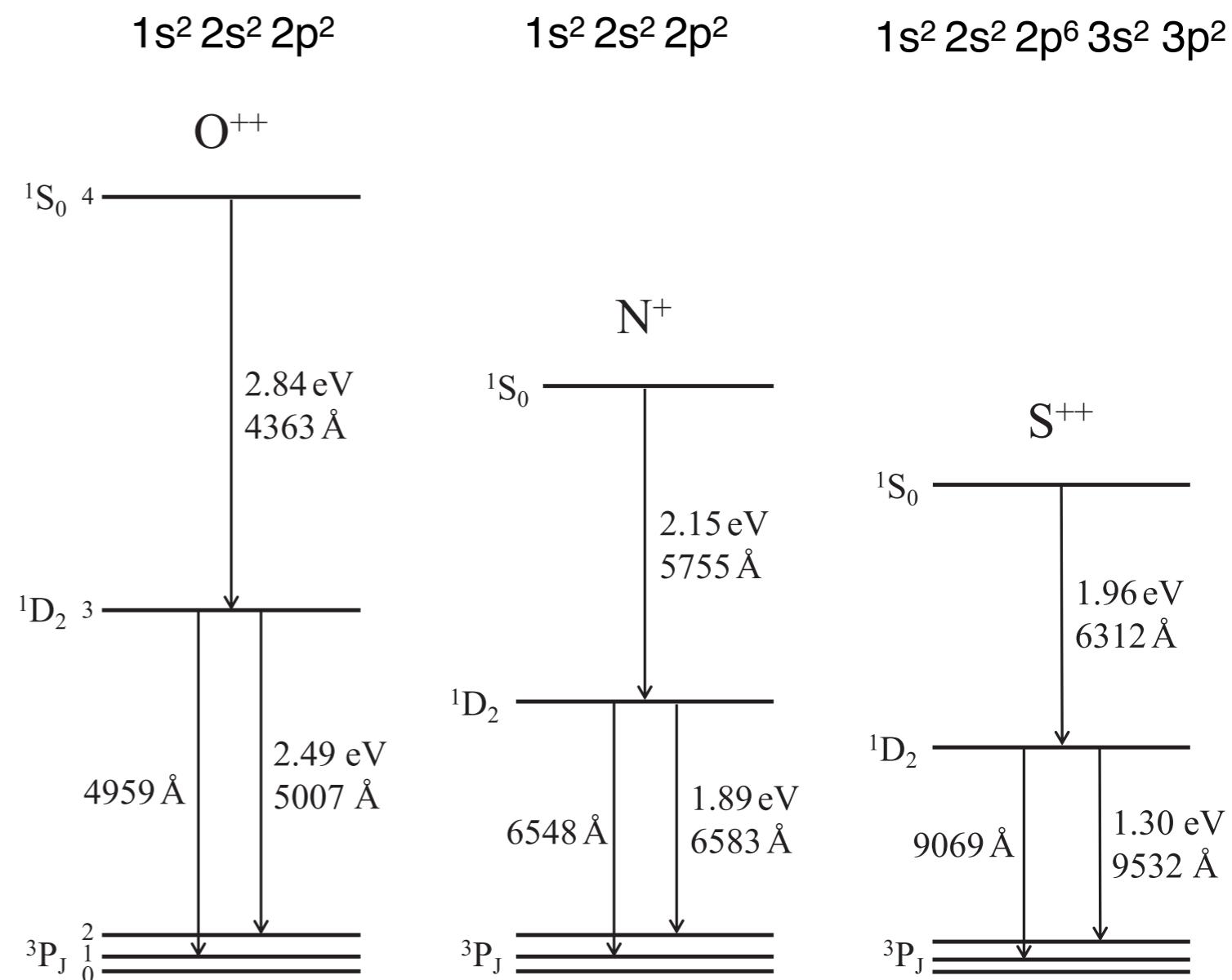


Figure 4.6 [Ryden]

- Candidate 2p² ions are C I, N II, O III, Fe IV, Ne V, and so on.
 - C I is easily photoionized, and will have very low abundance in an H II region.
 - Fe IV, Ne VI,... have an ionization potential exceeding 54.8 eV, and we do not expect such high ionization states to be abundant in H II regions.
 - This leaves N II and O III as the only 2p² ions that will be available in H II regions.
- The lowest excited states of singly ionized nitrogen (N II) and doubly ionized oxygen (O III) are useful tools for estimating the temperatures of H II regions and planetary nebulae.
- N II and O III have six bound electrons, and thus their fine structure energy levels in their lowest configuration are very similar in structure.

Temperature-sensitive nebular lines (Å).

p ² Ions	[N II]	[O III]	[Ne V]	[S III]
$^1S_0 \rightarrow ^1D_2$	5755	4363	2974	6312
$^1D_2 \rightarrow ^3P_2$	6583	5007	3426	9532
$^1D_2 \rightarrow ^3P_1$	6548	4959	3346	9069
p ⁴ Ions	[O I]	[Ne III]	[Ar III]	
$^1S_0 \rightarrow ^1D_2$	5577	3343	5192	
$^1D_2 \rightarrow ^3P_2$	6300	3869	7136	
$^1D_2 \rightarrow ^3P_1$	6363	3968	7751	

Element	I→II	II→III	III→IV	IV→V	V→VI	VI→VII	VII→VIII
1 H	13.5984						
2 He	24.5874	54.416					
3 Li	5.3917	75.640	122.454				
4 Be	9.3227	18.211	153.894	217.719			
5 B	8.2980	25.155	37.931	259.375	340.226		
6 C	11.2603	24.383	47.888	64.494	392.089	489.993	
7 N	14.5341	29.601	47.449	77.474	97.890	552.072	667.046
8 O	13.6181	35.121	54.936	77.414	113.899	138.120	739.293
9 F	17.4228	34.971	62.708	87.140	114.243	147.163	185.189
10 Ne	21.5645	40.963	63.423	97.117	126.247	154.214	207.271
11 Na	5.1391	47.286	71.620	98.91	138.40	172.183	208.50
12 Mg	7.6462	15.035	80.144	109.265	141.270	186.76	225.02
13 Al	5.9858	18.829	28.448	119.992	153.825	190.477	241.76
14 Si	8.1517	16.346	33.493	45.142	166.767	205.267	246.481
15 P	10.4867	19.769	30.203	51.444	65.025	220.422	263.57
16 S	10.3600	23.338	34.790	47.222	72.594	88.053	280.948
17 Cl	12.9676	23.814	39.911	53.465	67.819	97.030	114.201
18 Ar	15.7596	27.630	40.735	59.686	75.134	91.00	124.328
19 K	4.3407	31.628	45.806	60.913	82.66	99.4	117.6
20 Ca	6.1132	11.872	50.913	67.27	84.51	108.8	127.2
21 Sc	6.5615	12.800	24.757	73.489	91.69	110.7	138.0
22 Ti	6.8281	13.576	24.492	43.267	123.7	119.533	140.846
23 V	6.7462	14.655	29.311	46.709	65.282	128.125	150.641
24 Cr	6.7665	16.486	30.959	49.160	69.456	90.635	160.175
25 Mn	7.4340	15.640	33.668	51.2	72.4	95.60	119.203
26 Fe	7.9024	16.188	30.651	54.801	75.010	99.063	124.976
27 Co	7.8810	17.084	33.50	51.27	79.5	102.	129.
28 Ni	7.6398	18.169	35.187	54.925	76.06	107.87	133.
29 Cu	7.7264	20.292	36.841	57.380	79.846	103.031	138.862
30 Zn	9.3492	17.964	39.723	59.573	82.574	133.903	133.903

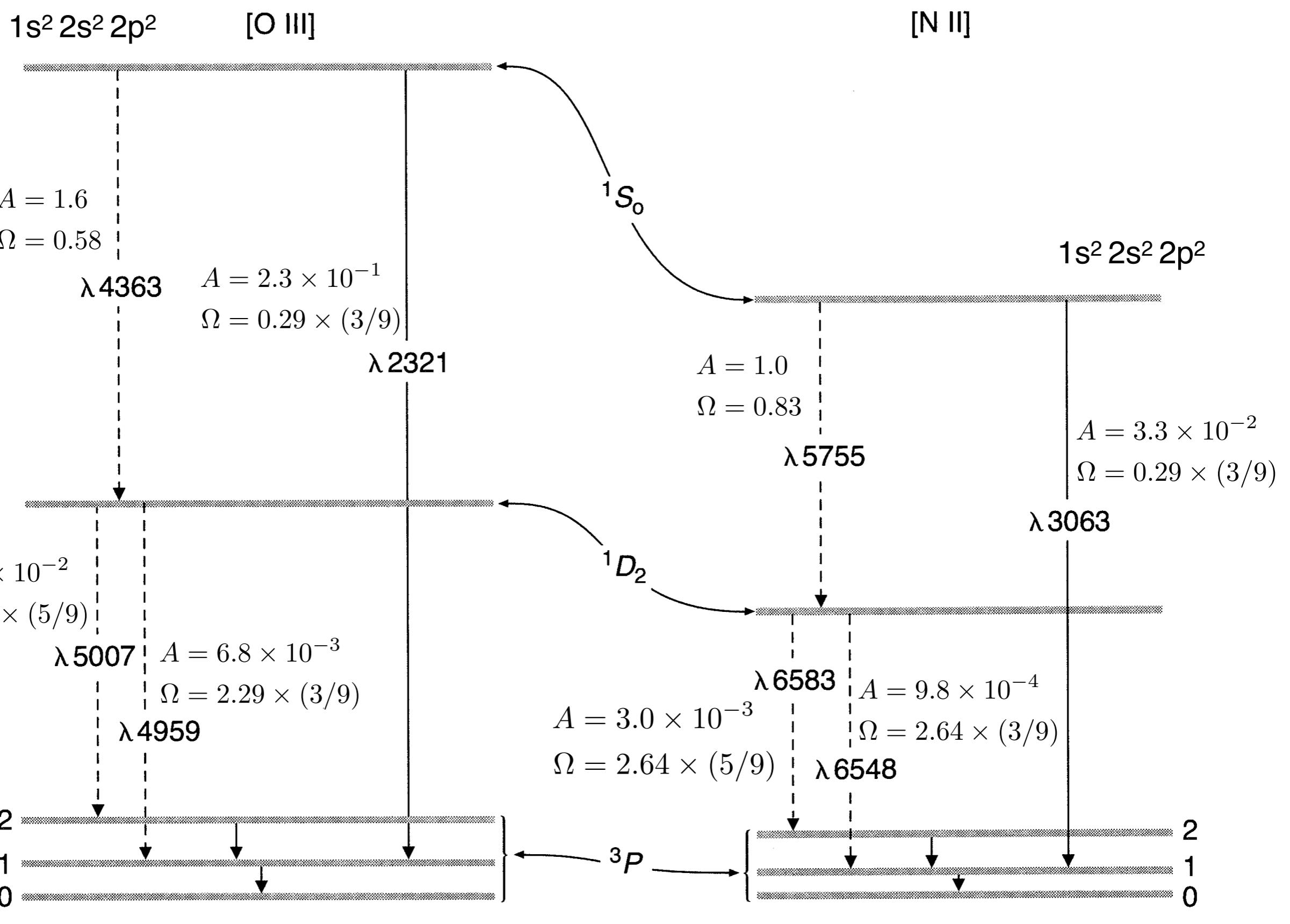


Figure 3.1 [Osterbrock]
Astrophysics of Gaseous Nebulae
and Active Galactic Nuclei

See Table 3.12 for A and
Table 3.6 for Collision Strength

Ion	$^3P, ^1D$	$^3P, ^1S$	$^1D, ^1S$	Table 3.6 Collision Strength
N^+	2.64	0.29	0.83	
O^{+2}	2.29	0.29	0.58	

- Let's suppose that we are in ***the low-density limit***, so that the free electron density is less than the critical density for collisional de-excitation of each line.

- In this case, every collisional excitation will be followed by radiative decays returning the ion to the ground state, with branching ratios that are determined by the Einstein coefficients.

- $4 \rightarrow 3$ transition:

- The emissivity of the $4 \rightarrow 3$ transition, integrated over the entire line width is:

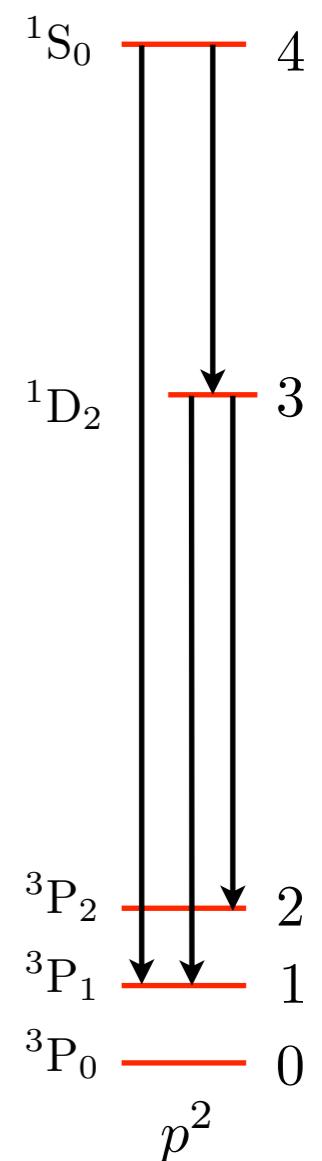
$$4\pi j(4 \rightarrow 3) = n_4 A_{43} h\nu_{43}$$

- The rate of collisional excitation from “0” to level “4” is balanced by radiative de-excitation from “4” to “3” and “1”:

$$n_e n_0 k_{04} = n_4 (A_{43} + A_{41})$$

- Therefore,

$$4\pi j(4 \rightarrow 3) = n_e n_0 k_{04} \frac{A_{43}}{A_{43} + A_{41}} h\nu_{43}$$



- $3 \rightarrow 2$ transition:

- The level “3” can be populated in two ways: (1) by collisional excitation directly from the ground state, and (2) by a collisional excitation from the ground to “4”, followed by radiative de-excitation to “3”.

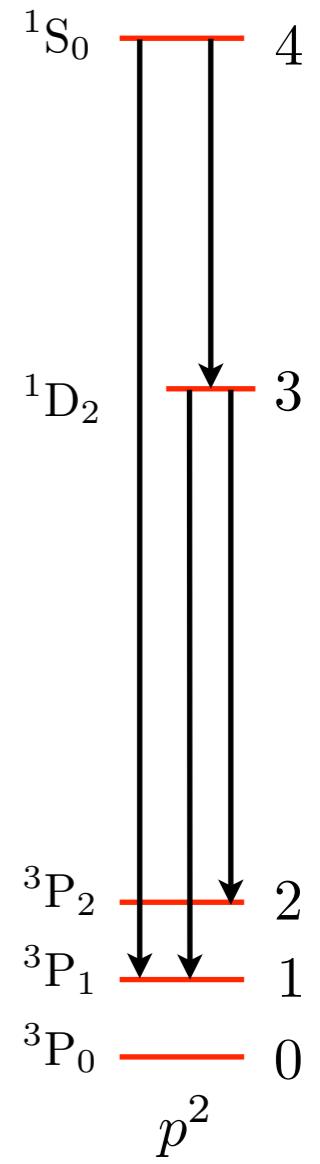
$$4\pi j(3 \rightarrow 2) = n_e n_0 \left(k_{03} + k_{04} \frac{A_{43}}{A_{43} + A_{41}} \right) \frac{A_{32}}{A_{32} + A_{31}} h\nu_{32}$$

- The relative strength between $4 \rightarrow 3$ and $3 \rightarrow 2$ emission line:

$$\frac{j(4 \rightarrow 3)}{j(3 \rightarrow 2)} = \frac{A_{43}}{A_{32}} \frac{\nu_{43}}{\nu_{32}} \frac{(A_{32} + A_{31})k_{04}}{(A_{43} + A_{41})k_{03} + A_{43}k_{04}}$$

- We notice that the temperature dependence in the above equation enters solely through the collisional rate coefficients k_{04} and k_{03} . Using the relation between the collisional excitation and de-excitation rate coefficients,

$$\begin{aligned} k_{0u} &= k_{u0} \frac{g_u}{g_0} e^{-h\nu_{u0}/kT} \\ &= \frac{\beta}{g_0} \frac{\langle \Omega_{u0} \rangle}{T^{1/2}} e^{-h\nu_{u0}/kT} \end{aligned}$$



- We obtain

$$\frac{j(4 \rightarrow 3)}{j(3 \rightarrow 2)} = \frac{A_{43}}{A_{32}} \frac{\nu_{43}}{\nu_{32}} \frac{(A_{32} + A_{31}) \langle \Omega_{40} \rangle e^{-h\nu_{43}/kT}}{(A_{43} + A_{41}) \langle \Omega_{30} \rangle + A_{43} \langle \Omega_{40} \rangle e^{-h\nu_{43}/kT}}$$

where $h\nu_{43} = h\nu_{40} - h\nu_{30}$

- Notice that all the temperature dependence, aside from the weak dependence of collision strengths on temperature, is contained in the exponential factor. Thus, the line ratio is sensitive to the temperature $kT \sim h\nu_{43}$ (2.15 eV for N II, 2.84 eV for O III).
- At the high and low temperatures, the ratio can be expressed as

$$\begin{aligned} \frac{j(4 \rightarrow 3)}{j(3 \rightarrow 2)} &\approx \frac{A_{43}}{A_{32}} \frac{\nu_{43}}{\nu_{32}} \frac{(A_{32} + A_{31}) \langle \Omega_{40} \rangle}{(A_{43} + A_{41}) \langle \Omega_{30} \rangle + A_{43} \langle \Omega_{40} \rangle} \quad \text{for } kT \gg h\nu_{43} \\ &\approx \frac{A_{43}}{A_{32}} \frac{\nu_{43}}{\nu_{32}} \frac{(A_{32} + A_{31}) \langle \Omega_{40} \rangle}{(A_{43} + A_{41}) \langle \Omega_{30} \rangle} e^{-h\nu_{43}/kT} \quad \text{for } kT \ll h\nu_{43} \end{aligned}$$

At high temperatures, the line ratio becomes more or less independent of temperature.

$$\frac{j(4 \rightarrow 3)}{j(3 \rightarrow 2)} = \frac{A_{43}}{A_{32}} \frac{\nu_{43}}{\nu_{32}} \frac{(A_{32} + A_{31}) \langle \Omega_{40} \rangle e^{-h\nu_{43}/kT}}{(A_{43} + A_{41}) \langle \Omega_{30} \rangle + A_{43} \langle \Omega_{40} \rangle e^{-h\nu_{43}/kT}}$$

Dependence of collision strength on temperature is very weak.
So, we will adopt a typical value.

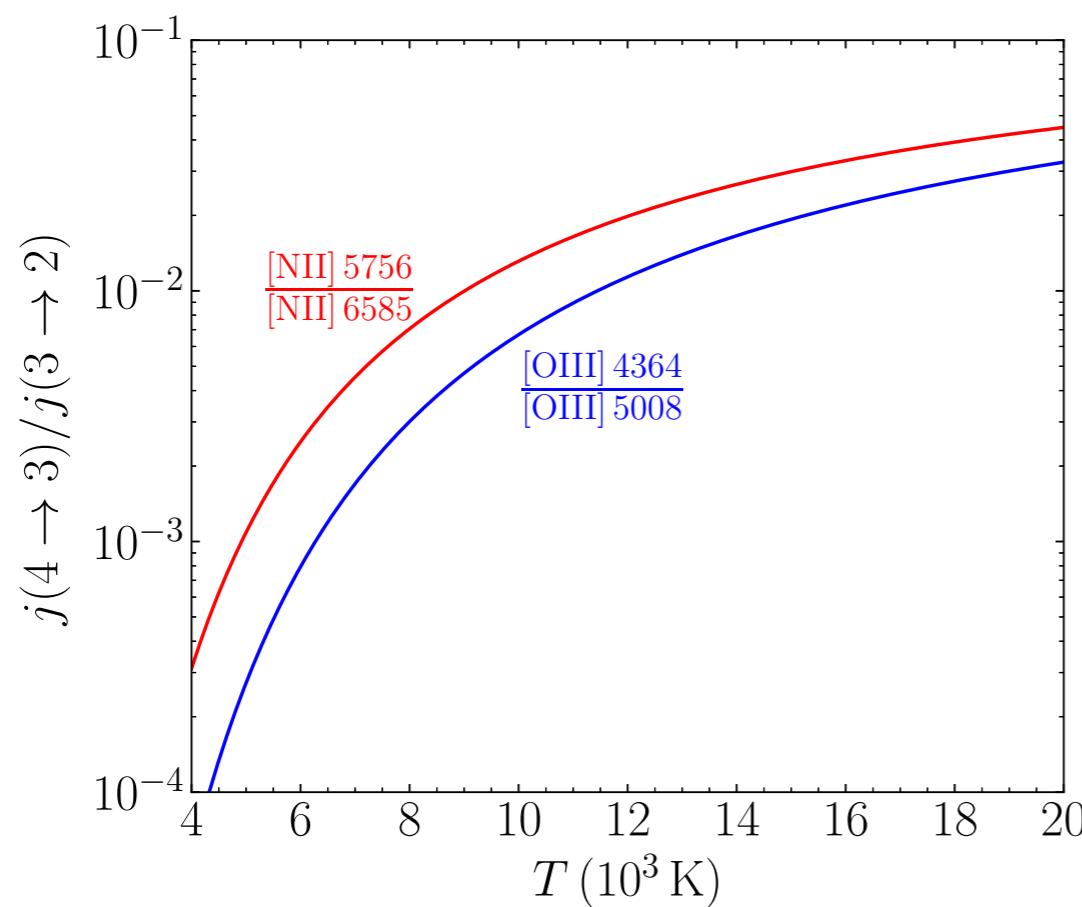
$$T_4 \equiv T/10^4 \text{ K}$$

[O III]	$\langle \Omega_{30} \rangle = 2.29 \times (1/9)$	$A_{32} = 2.0 \times 10^{-2} \text{ [s}^{-1}\text{]}$
	$\langle \Omega_{40} \rangle = 0.29 \times (1/9)$	$A_{31} = 6.8 \times 10^{-3} \text{ [s}^{-1}\text{]}$
	$E_{40}/k = 61207 \text{ [K]}$	$A_{43} = 1.6 \text{ [s}^{-1}\text{]}$
	$E_{30}/k = 29169 \text{ [K]}$	$A_{41} = 2.3 \times 10^{-1} \text{ [s}^{-1}\text{]}$
	$E_{20}/k = 441 \text{ [K]}$	
	$E_{10}/k = 163 \text{ [K]}$	

$$\frac{j(4364)}{j(5008)} = 0.1655 \frac{e^{-3.2038/T_4}}{1 + 0.1107e^{-3.2038/T_4}}$$

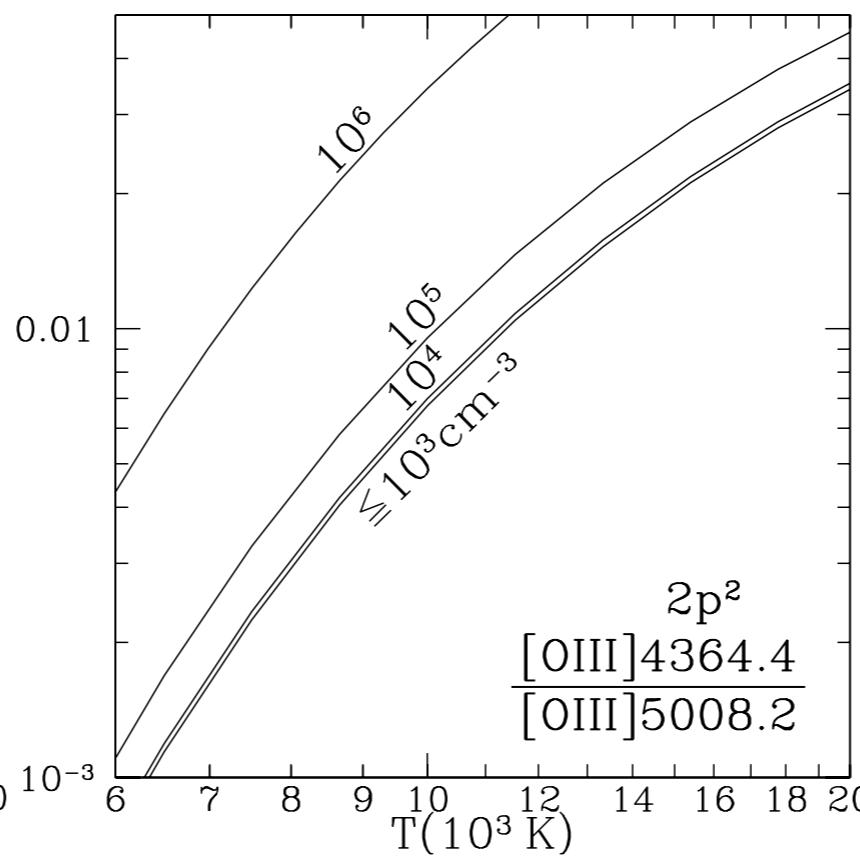
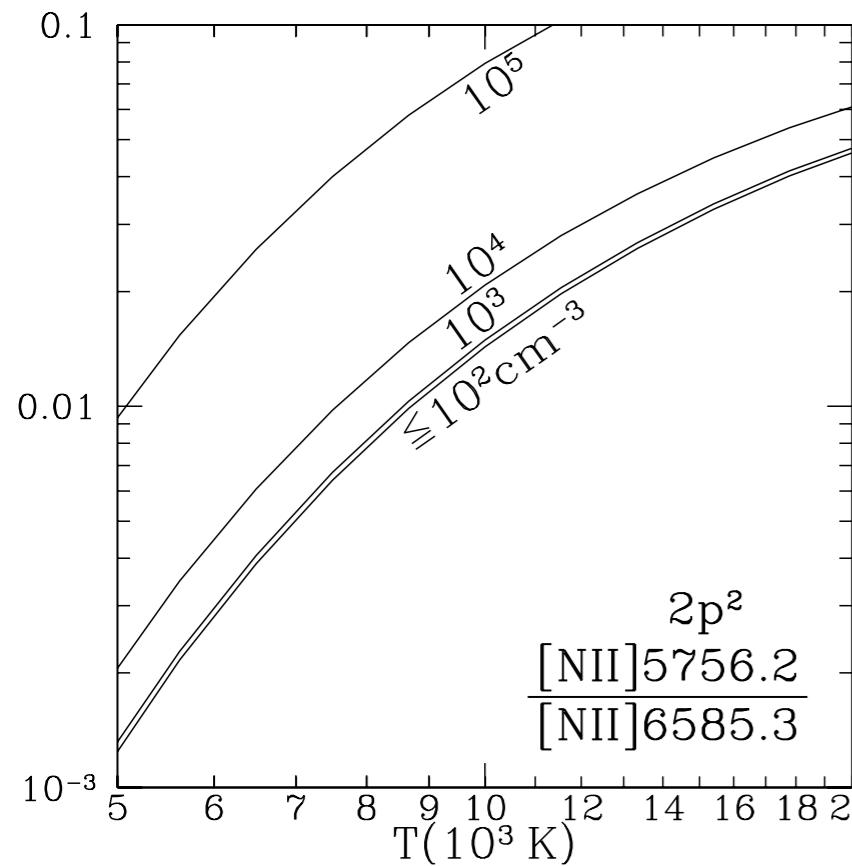
[N II]	$\langle \Omega_{30} \rangle = 2.64 \times (1/9)$	$A_{32} = 3.0 \times 10^{-3} \text{ [s}^{-1}\text{]}$
	$\langle \Omega_{40} \rangle = 0.29 \times (1/9)$	$A_{31} = 9.8 \times 10^{-4} \text{ [s}^{-1}\text{]}$
	$E_{40}/k = 47033 \text{ [K]}$	$A_{43} = 1.0 \text{ [s}^{-1}\text{]}$
	$E_{30}/k = 22037 \text{ [K]}$	$A_{41} = 3.3 \times 10^{-2} \text{ [s}^{-1}\text{]}$
	$E_{20}/k = 188 \text{ [K]}$	
	$E_{10}/k = 70 \text{ [K]}$	

$$\frac{j(5756)}{j(6585)} = 0.1614 \frac{e^{-2.4996/T_4}}{1 + 0.1063e^{-2.4996/T_4}}$$



Line ratios as a function of temperature, which is obtained using the equations in the previous slide.

Unfortunately, the auroral line at $[\text{O III}] 4364\text{\AA}$ can only be observed when the temperature is sufficiently high.



Line ratios for temperature diagnostics.

Curves indicate electron density. For each ion, the low density limit is shown, as well as results for higher densities.

Figure 18.2 [Draine]

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- Sometimes, it would be better to combine $3 \rightarrow 2$ and $3 \rightarrow 1$ transitions:

- $3 \rightarrow 1$ transition:

$$4\pi j(3 \rightarrow 2) = n_e n_0 \left(k_{03} + k_{04} \frac{A_{43}}{A_{43} + A_{41}} \right) \frac{A_{32}}{A_{32} + A_{31}} h\nu_{32}$$

$$4\pi j(3 \rightarrow 1) = n_e n_0 \left(k_{03} + k_{04} \frac{A_{43}}{A_{43} + A_{41}} \right) \frac{A_{31}}{A_{32} + A_{31}} h\nu_{31}$$

$$4\pi [j(3 \rightarrow 1) + j(3 \rightarrow 2)] = n_e n_0 \left(k_{03} + k_{04} \frac{A_{43}}{A_{43} + A_{41}} \right) h\bar{\nu} \quad \text{where } \bar{\nu} \equiv \frac{A_{32}\nu_{32} + A_{31}\nu_{31}}{A_{32} + A_{31}}$$

- Recall that

$$4\pi j(4 \rightarrow 3) = n_e n_0 k_{04} \frac{A_{43}}{A_{43} + A_{41}} h\nu_{43}$$

Combining these equations, we obtain

$$\frac{j(3 \rightarrow 1) + j(3 \rightarrow 2)}{j(4 \rightarrow 3)} = \frac{\bar{\nu}}{\nu_{43}} \frac{A_{43} + A_{41}}{A_{43}} \frac{k_{03}}{k_{04}} \left(1 + \frac{k_{04}}{k_{03}} \frac{A_{43}}{A_{43} + A_{41}} \right)$$

$$k_{0u} = \frac{\beta}{T^{1/2}} \frac{\langle \Omega_{u0} \rangle}{g_0} e^{-E_{u0}/kT_{\text{gas}}}$$

↓

$$\frac{k_{03}}{k_{04}} = \frac{\langle \Omega_{30} \rangle}{\langle \Omega_{40} \rangle} \frac{e^{-h\nu_{30}/kT}}{e^{-h\nu_{40}/kT}} = \frac{\langle \Omega_{30} \rangle}{\langle \Omega_{40} \rangle} e^{h\nu_{43}/kT} \quad (\text{where } \nu_{43} = \nu_{40} - \nu_{30})$$

$$\begin{aligned} \frac{j(3 \rightarrow 1) + j(3 \rightarrow 2)}{j(4 \rightarrow 3)} &= \frac{\bar{\nu}}{\nu_{43}} \frac{A_{43} + A_{41}}{A_{43}} \frac{\langle \Omega_{30} \rangle}{\langle \Omega_{40} \rangle} e^{h\nu_{43}/kT} \left(1 + \frac{\langle \Omega_{40} \rangle}{\langle \Omega_{30} \rangle} \frac{A_{43}}{A_{43} + A_{41}} e^{-h\nu_{43}/kT} \right) \\ &\simeq \frac{\bar{\nu}}{\nu_{43}} \frac{A_{43} + A_{41}}{A_{43}} \frac{\langle \Omega_{30} \rangle}{\langle \Omega_{40} \rangle} e^{h\nu_{43}/kT} \end{aligned}$$

Note $\langle \Omega_{40} \rangle < \langle \Omega_{30} \rangle$ and $e^{-h\nu_{43}/kT} \ll 1$.

Thus, the second term inside the parenthesis is negligible.

$$\frac{j(3 \rightarrow 1) + j(3 \rightarrow 2)}{j(4 \rightarrow 3)} \simeq \frac{\bar{\nu}}{\nu_{43}} \frac{A_{43} + A_{41}}{A_{43}} \frac{\langle \Omega_{30} \rangle}{\langle \Omega_{40} \rangle} e^{h\nu_{43}/kT}$$

Use the following data:

[O III]	$\langle \Omega_{30} \rangle = 2.29 \times (1/9)$	$A_{32} = 2.0 \times 10^{-2} \text{ [s}^{-1}\text{]}$
	$\langle \Omega_{40} \rangle = 0.29 \times (1/9)$	$A_{31} = 6.8 \times 10^{-3} \text{ [s}^{-1}\text{]}$
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[N II]	$\langle \Omega_{30} \rangle = 2.64 \times (1/9)$	$A_{32} = 3.0 \times 10^{-3} \text{ [s}^{-1}\text{]}$
	$\langle \Omega_{40} \rangle = 0.29 \times (1/9)$	$A_{31} = 9.8 \times 10^{-4} \text{ [s}^{-1}\text{]}$
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	$E_{30}/k = 22037 \text{ [K]}$	$A_{41} = 3.3 \times 10^{-2} \text{ [s}^{-1}\text{]}$
	$E_{20}/k = 188 \text{ [K]}$	
	$E_{10}/k = 70 \text{ [K]}$	

We obtain the line ratio as a function of temperature.

$$\frac{[\text{O III}] 4960 + 5008}{[\text{O III}] 4364} = 8.12 e^{3.20/T_4}$$

$$\frac{[\text{N II}] 6549 + 6585}{[\text{N II}] 5756} = 8.23 e^{2.50/T_4}$$

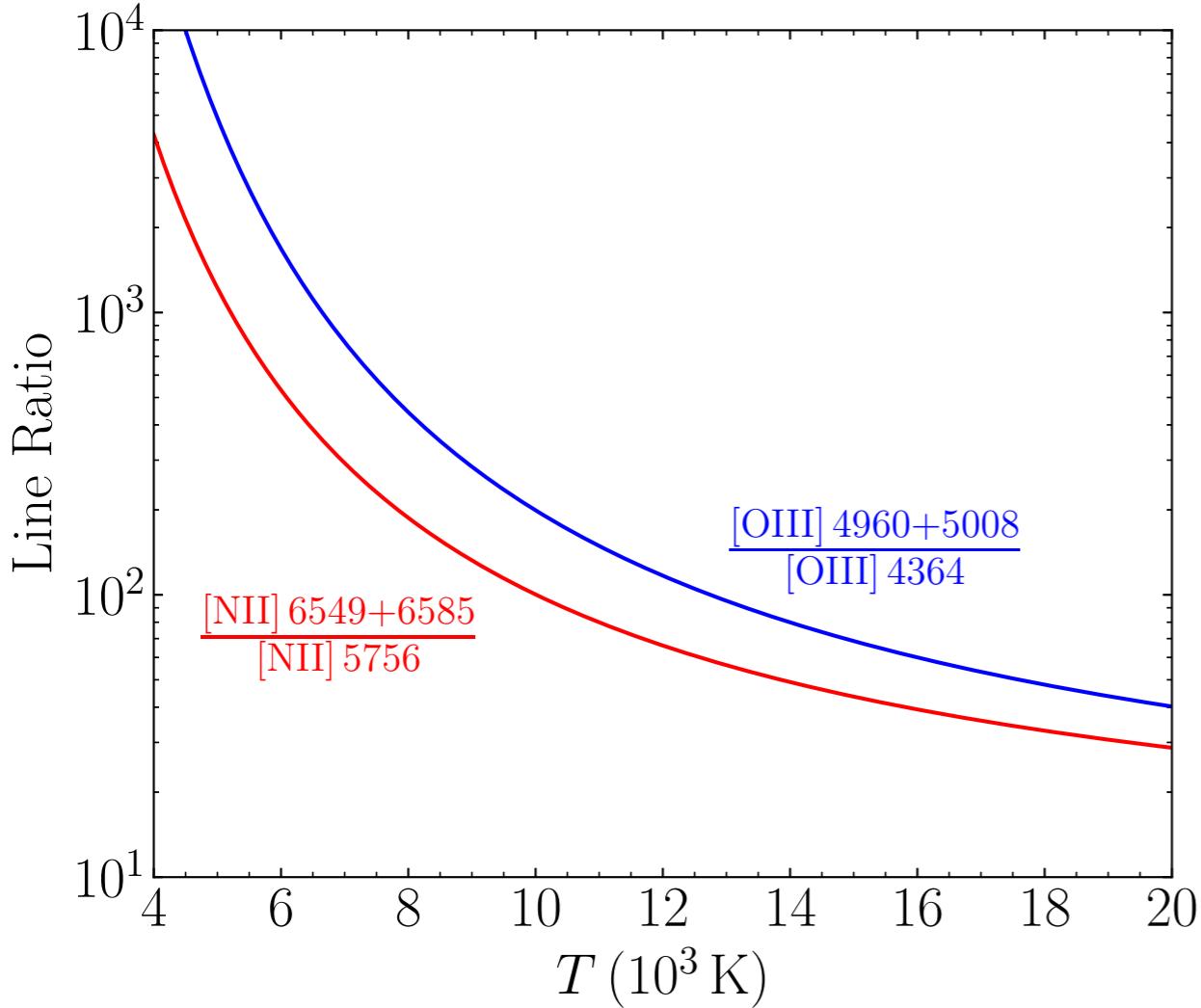
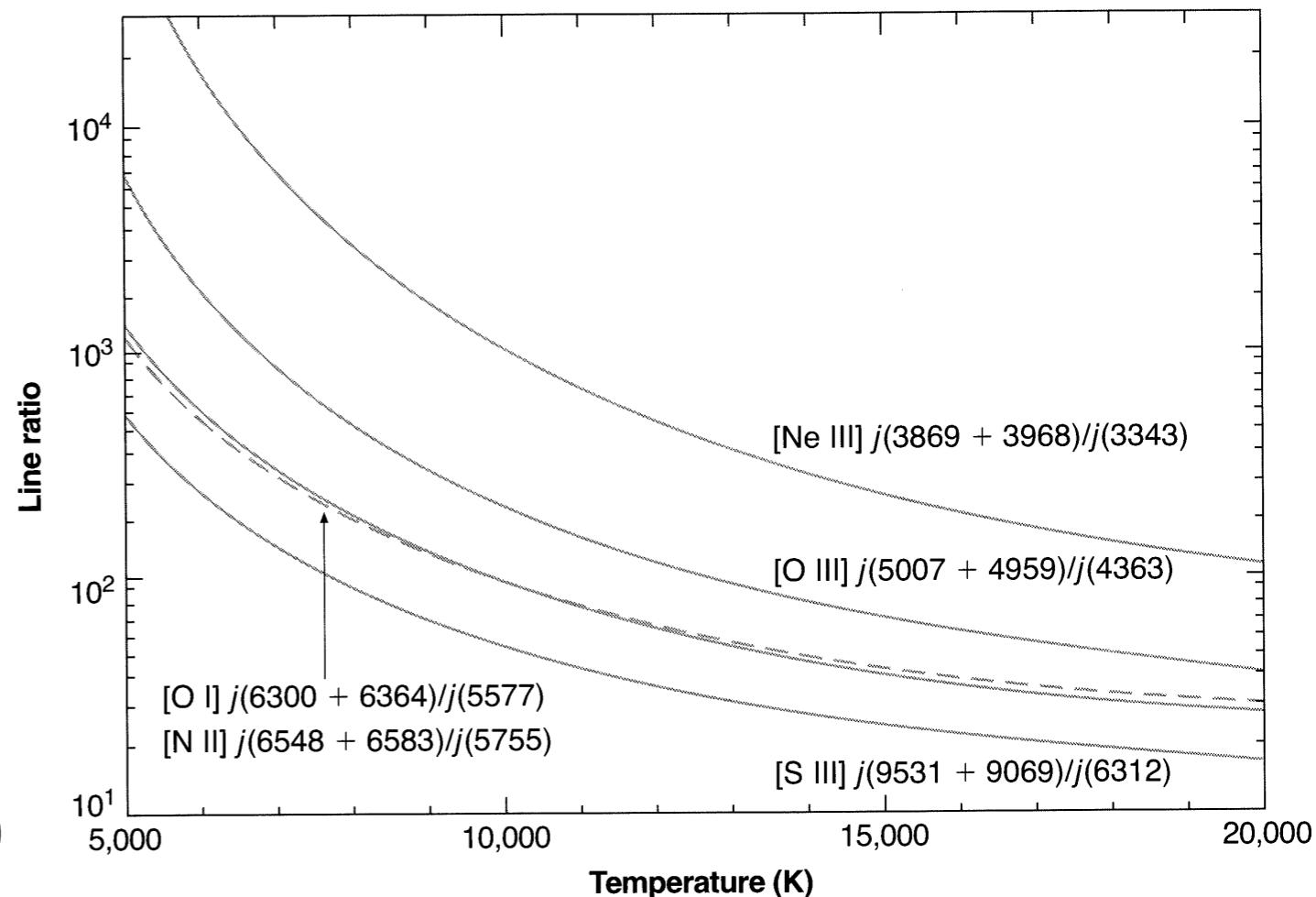


Figure 5.1 [Osterbrock]



See Equations (5.4)-(5.7) for a correction factor for the density effect.

$$[O III] \frac{j_{\lambda 4959} + j_{\lambda 5007}}{j_{\lambda 4363}} = \frac{7.90 \exp(3.29 \times 10^4/T)}{1 + 4.5 \times 10^{-4} n_e / T^{1/2}}$$

$$[N II] \frac{j_{\lambda 6548} + j_{\lambda 6583}}{j_{\lambda 5755}} = \frac{8.23 \exp(2.50 \times 10^4/T)}{1 + 4.4 \times 10^{-3} n_e / T^{1/2}}$$

$$[Ne III] \frac{j_{\lambda 3869} + j_{\lambda 3968}}{j_{\lambda 3343}} = \frac{13.7 \exp(4.30 \times 10^4/T)}{1 + 3.8 \times 10^{-5} n_e / T^{1/2}}$$

$$[S III] \frac{j_{\lambda 9531} + j_{\lambda 9069}}{j_{\lambda 6312}} = \frac{5.44 \exp(2.28 \times 10^4/T)}{1 + 3.5 \times 10^{-4} n_e / T^{1/2}}$$

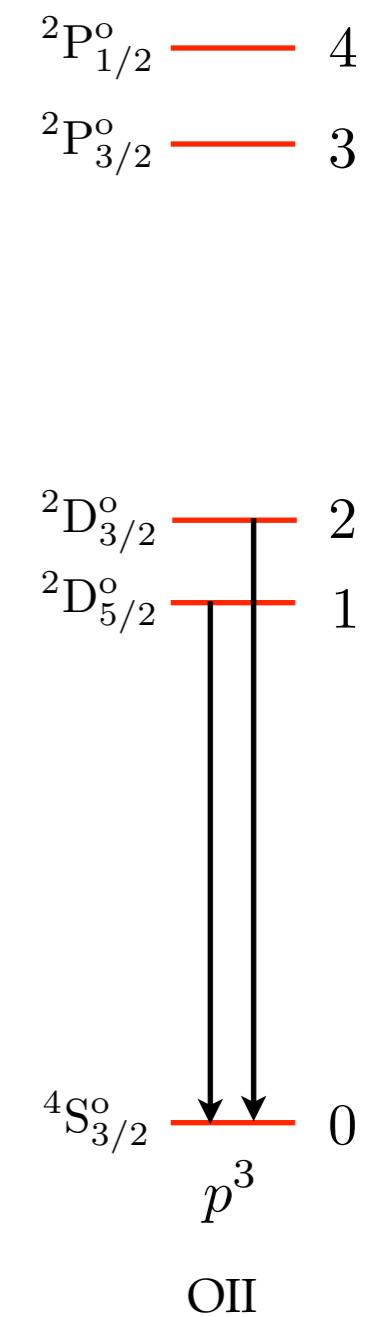
Density Determination

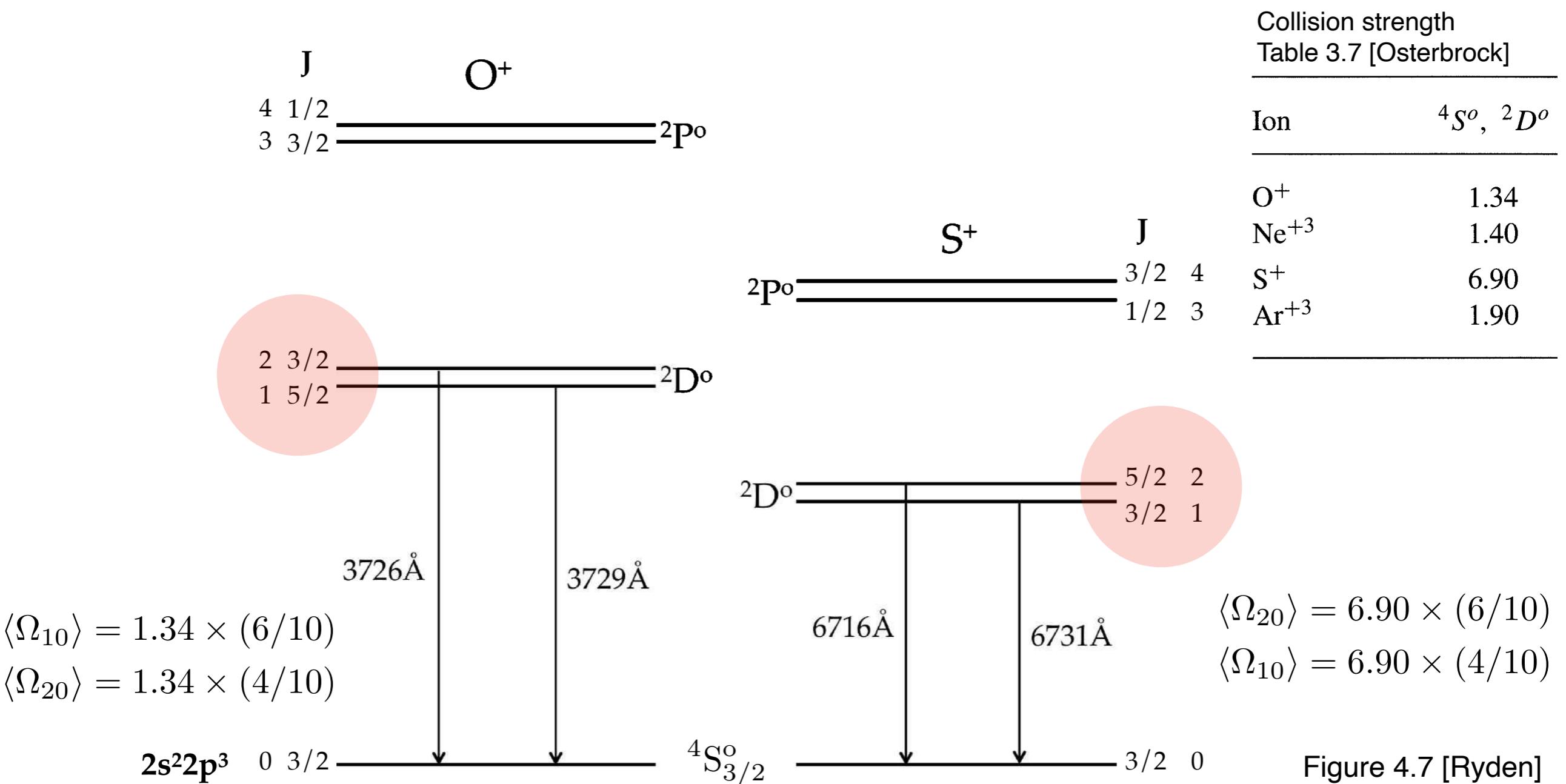
- Emission lines can also be used to estimate the free electron density of an ionized nebula. For this purpose, we need an ion which has ***two excited levels that are similar in energy, but which have different critical densities***. One of such systems is singly-ionized oxygen.

- Ions with 7 or 15 electrons have $2s^22p^3$ and $3s^23p^3$ configurations, with energy level structures that make them suitable for use as density diagnostics.

Density-sensitive nebular lines (Å).

p^3 Ions	[O II]	[S II]	[Ne IV]	[Ar IV]
$^2D_{3/2} \rightarrow ^4S_{3/2}$	3726	6731	2423	4740
$^2D_{5/2} \rightarrow ^4S_{3/2}$	3729	6716	2426	4711
	$2s^22p^2$	$3s^23p^2$	$2s^22p^2$	$3s^23p^2$
	Z = 8	Z = 16	Z = 10	Z = 18





Notice that energy ordering of the fine-structure levels are different between O^+ and S^+ . The p^3 configuration for the two ions are half-filled, and thus **Hund's rule for the energy ordering is not applicable**.

[There are three typos in J values for S^+ and in the notation for the lowest level in Figure 4.7 of Ryden.]

- Here, we will ignore the transition between 2 and 1 because the transition is very slow.
- $1 \rightarrow 0$ transition:
 - The emissivity of the $1 \rightarrow 0$ transition, integrated over the entire line width, is

$$4\pi j(1 \rightarrow 0) = n_1 A_{10} h\nu_{10}$$

- In statistical equilibrium, the rate of collisional excitation from the ground state will be balanced by radiative and collisional de-excitation:

$$n_e n_0 k_{01} = n_1 (A_{10} + n_e k_{10})$$

Then,

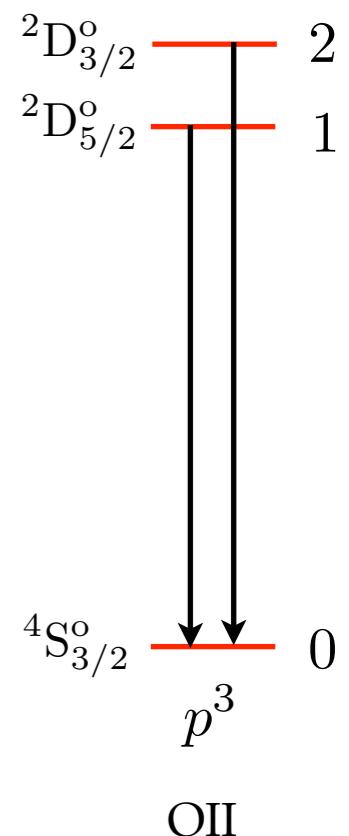
$$\begin{aligned} 4\pi j(1 \rightarrow 0) &= n_e n_0 \frac{k_{01}}{A_{10} + n_e k_{10}} A_{10} h\nu_{10} \\ &= n_e n_0 \frac{k_{01}}{1 + n_e / n_{\text{crit},1}} h\nu_{10} \end{aligned}$$

where $n_{\text{crit},1} \equiv A_{10}/k_{10}$

- $2 \rightarrow 0$ transition:

- Similarly, we obtain

$$4\pi j(2 \rightarrow 0) = n_e n_0 \frac{k_{02}}{1 + n_e / n_{\text{crit},2}} h\nu_{20} \quad \text{where } n_{\text{crit},2} \equiv A_{20}/k_{20}$$



- The ratio of the strength of the two lines in the doublet is

$$\begin{aligned}\frac{j(2 \rightarrow 0)}{j(1 \rightarrow 0)} &= \frac{\nu_{20}}{\nu_{10}} \frac{k_{02}}{k_{01}} \frac{1 + n_e/n_{\text{crit},1}}{1 + n_e/n_{\text{crit},2}} \\ &= \frac{\nu_{20}}{\nu_{10}} \frac{\langle \Omega_{20} \rangle}{\langle \Omega_{10} \rangle} e^{-h\nu_{21}/kT} \frac{1 + n_e/n_{\text{crit},1}}{1 + n_e/n_{\text{crit},2}}\end{aligned}$$

$$k_{0u} = \left(\frac{\beta}{T^{1/2}} \frac{1}{g_0} \right) \langle \Omega_{u0} \rangle e^{-h\nu_{u0}/kT}$$

$$\beta = 8.62942 \times 10^{-6}$$

- Thus, we can write the line ratio as

$$\frac{j(2 \rightarrow 0)}{j(1 \rightarrow 0)} \simeq \frac{\langle \Omega_{20} \rangle}{\langle \Omega_{10} \rangle} \frac{1 + n_e/n_{\text{crit},1}}{1 + n_e/n_{\text{crit},2}}$$

$h\nu_{20} \simeq h\nu_{10}$ Levels 1 and 2 are so close in energy.

$h\nu_{21} \equiv h\nu_{20} - h\nu_{10} \ll kT$

$h\nu_{21} \approx 2 \text{ meV}$ for O II ions

- In the low-density limit ($n_e \ll n_{\text{crit},1}, n_{\text{crit},2}$),

$$\frac{j(2 \rightarrow 0)}{j(1 \rightarrow 0)} \simeq \frac{\langle \Omega_{20} \rangle}{\langle \Omega_{10} \rangle} = \frac{g_2}{g_1}$$

$$\Omega_{(\text{SLJ}, \text{ S'L'J'})} = \frac{(2J' + 1)}{(2S' + 1)(2L' + 1)} \Omega_{(\text{SL}, \text{ S'L'})}$$

Recall the sum rule for the collision strength for the fine-structure transitions.

[However, it is not clear that the sum rule is valid even beyond the LS-coupling scheme. Recent QM calculations show that the proportionality relation is only an approximation.]

- In high-density limit ($n_e \gg n_{\text{crit},2}, n_{\text{crit},1}$),

$$\frac{j(2 \rightarrow 0)}{j(1 \rightarrow 0)} \simeq \frac{\langle \Omega_{20} \rangle}{\langle \Omega_{10} \rangle} \frac{n_{\text{crit},2}}{n_{\text{crit},1}} = \frac{\langle \Omega_{20} \rangle}{\langle \Omega_{10} \rangle} \frac{A_{20}/k_{20}}{A_{10}/k_{10}} = \frac{g_2}{g_1} \frac{A_{20}}{A_{10}}$$

$$k_{u0} = \frac{\beta}{T^{1/2}} \frac{\langle \Omega_{u0} \rangle}{g_u}$$

► For O II ion,

$$n_e \ll n_{\text{crit}} \rightarrow \frac{j(1 \rightarrow 0)}{j(2 \rightarrow 0)} \simeq \frac{g_1}{g_2}$$

$$\frac{j([\text{O II}] 3728.8)}{j([\text{O II}] 3726.1)} = 1.5$$

$$n_e \gg n_{\text{crit}} \rightarrow \frac{j(1 \rightarrow 0)}{j(2 \rightarrow 0)} \simeq \frac{g_1}{g_2} \frac{A_{10}}{A_{20}}$$

$$\frac{j([\text{O II}] 3728.8)}{j([\text{O II}] 3726.1)} = 0.3$$

$$g_1 = 6, \quad A_{10} = 3.59 \times 10^{-5} \text{ s}^{-1}$$

$$g_2 = 4, \quad A_{20} = 1.79 \times 10^{-4} \text{ s}^{-1}$$

$$\frac{j(3729)}{j(3726)} = 1.5 \frac{1 + (n_e / 1.55 \times 10^4 \text{ cm}^{-3}) T_4^{-1/2}}{1 + (n_e / 3.11 \times 10^3 \text{ cm}^{-3}) T_4^{-1/2}}$$

$$\langle \Omega_{10} \rangle = 1.34 \times (6/10)$$

$$\langle \Omega_{20} \rangle = 1.34 \times (4/10)$$

► For S II ion,

$$n_e \ll n_{\text{crit}} \rightarrow \frac{j(2 \rightarrow 0)}{j(1 \rightarrow 0)} \simeq \frac{g_2}{g_1}$$

$$\frac{j([\text{S II}] 6716)}{j([\text{S II}] 6731)} = 1.5$$

$$n_e \gg n_{\text{crit}} \rightarrow \frac{j(2 \rightarrow 0)}{j(1 \rightarrow 0)} \simeq \frac{g_2}{g_1} \frac{A_{20}}{A_{10}}$$

$$\frac{j([\text{S II}] 6716)}{j([\text{S II}] 6731)} = 0.44$$

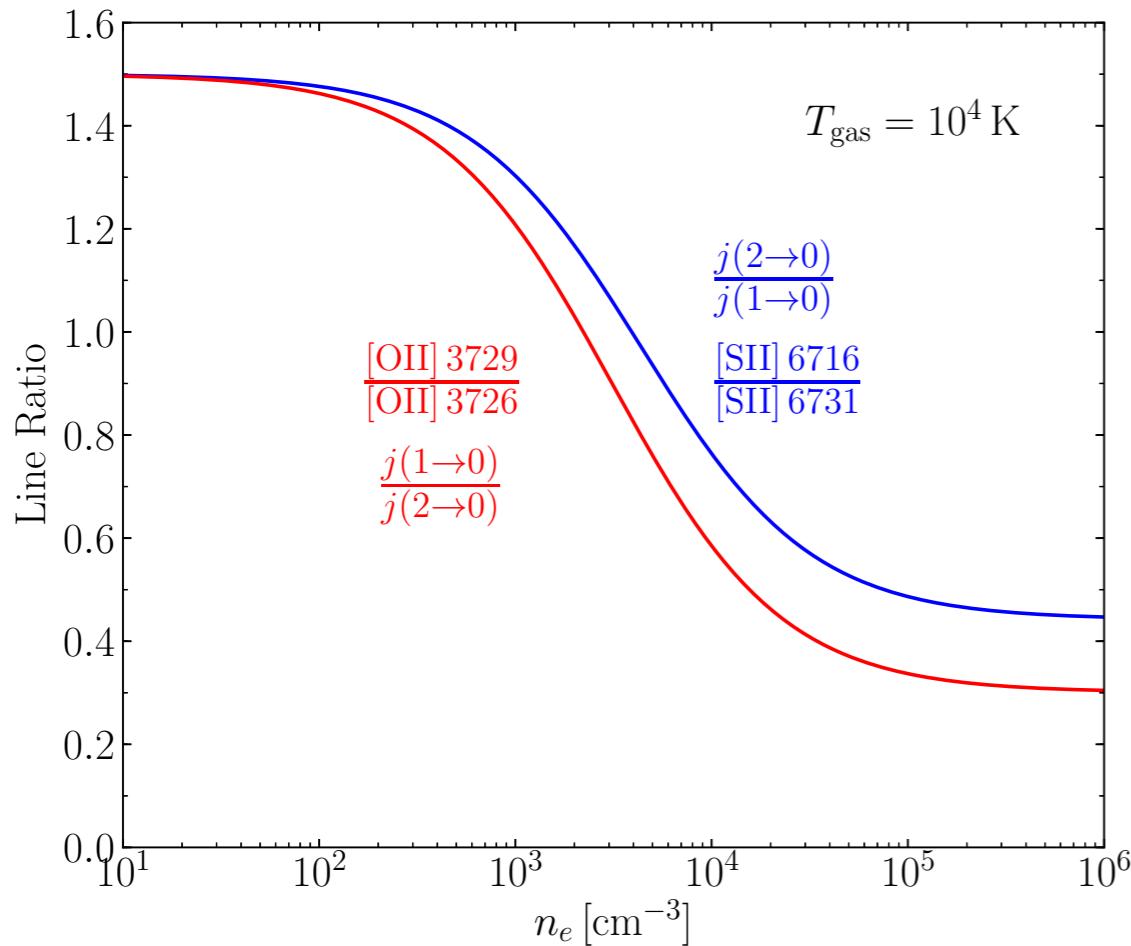
$$g_2 = 6, \quad A_{20} = 2.60 \times 10^{-4} \text{ s}^{-1}$$

$$g_1 = 4, \quad A_{10} = 8.82 \times 10^{-4} \text{ s}^{-1}$$

$$\frac{j(6716)}{j(6731)} = 1.5 \frac{1 + (n_e / 1.48 \times 10^4 \text{ cm}^{-3}) T_4^{-1/2}}{1 + (n_e / 4.37 \times 10^3 \text{ cm}^{-3}) T_4^{-1/2}}$$

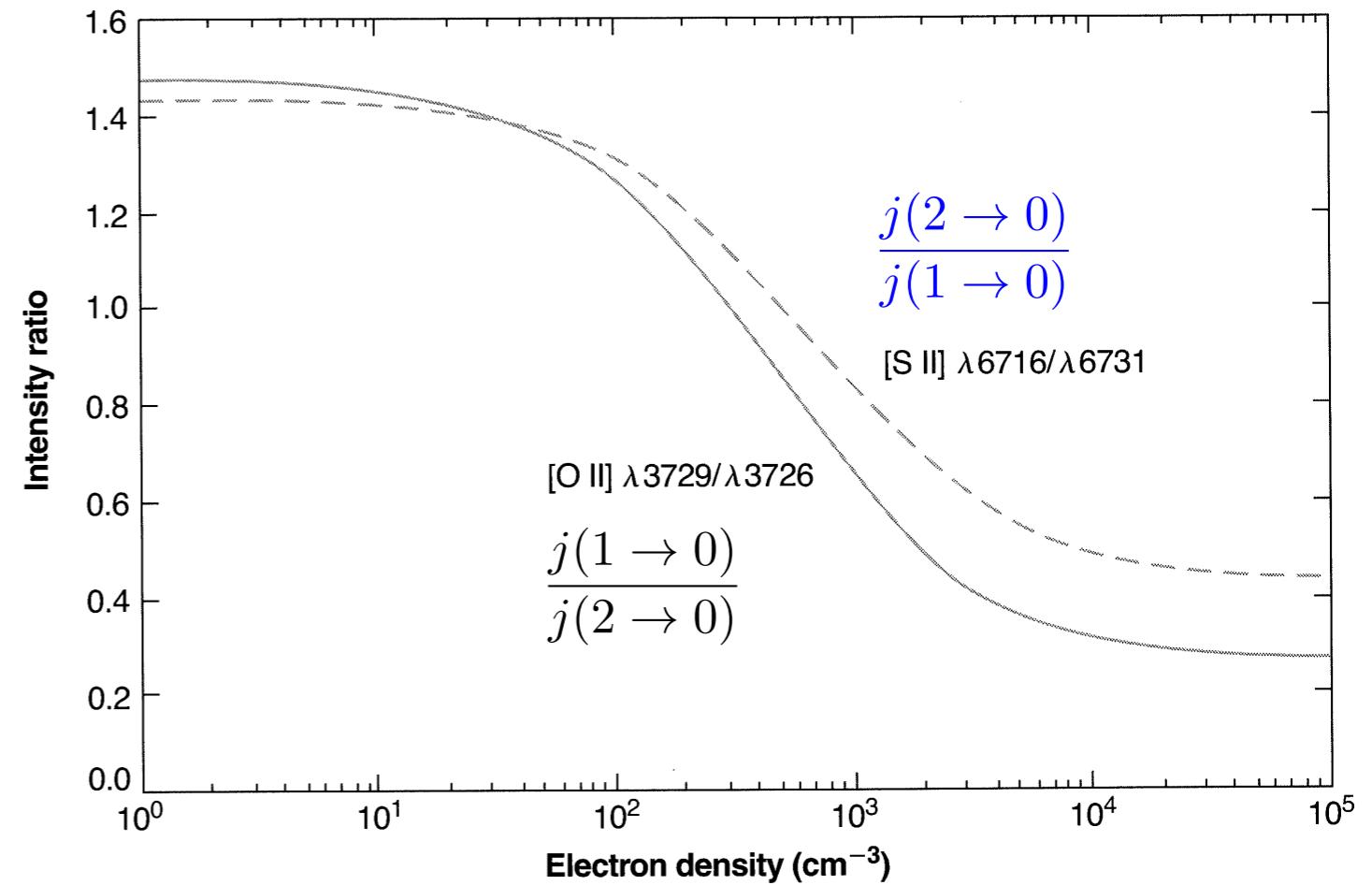
$$\langle \Omega_{20} \rangle = 6.90 \times (6/10)$$

$$\langle \Omega_{10} \rangle = 6.90 \times (4/10)$$



Obtained using the approximate equations in this lecture note.

Notice differences between two figures.



The full solution of the equilibrium equations, which also takes into account all transitions, including excitation to the ${}^2\text{P}^o$ levels with subsequent cascading downward.

Figure 5.8 [Osterbrock]

Abundance Determination

- Helium abundance
 - The abundance of He is determined from **comparison of the strengths of radiative recombination lines of H and He** in regions ionized by stars that are sufficiently hot ($T_{\text{eff}} \gtrsim 3.9 \times 10^4 \text{ K}$) so that He is ionized throughout the H II regions.
 - Heavy elements
 - The abundance of heavy elements can be inferred by **comparing the strengths of collisionally excited lines with recombination lines of H**.

- **Oxygen:** $(\lambda_{H\beta} = 4861.35\text{\AA})$

$$4\pi j(\text{[OIII]} 5008) = n_e n(\text{O}^{+2}) k_{03} \frac{A_{32}}{A_{31} + A_{32}} E_{32}$$

$$4\pi j(\mathrm{H}\beta) = n_e n(\mathrm{H}^+) \alpha_{\mathrm{eff},\mathrm{H}\beta} E_{\mathrm{H}\beta}$$

where

$$\alpha_{\text{eff}, \text{H}\beta} \approx 3.03 \times 10^{-14} T_4^{-0.874 - 0.058 \ln T_4} \text{ cm}^3 \text{ s}^{-1}$$

$$k_{03} = 8.62942 \times 10^{-8} T_4^{-1/2} \frac{\Omega_{30}}{q_0} e^{-E_{30}/kT} \text{ cm}^3 \text{ s}^{-1} \quad (g_0 = 1)$$

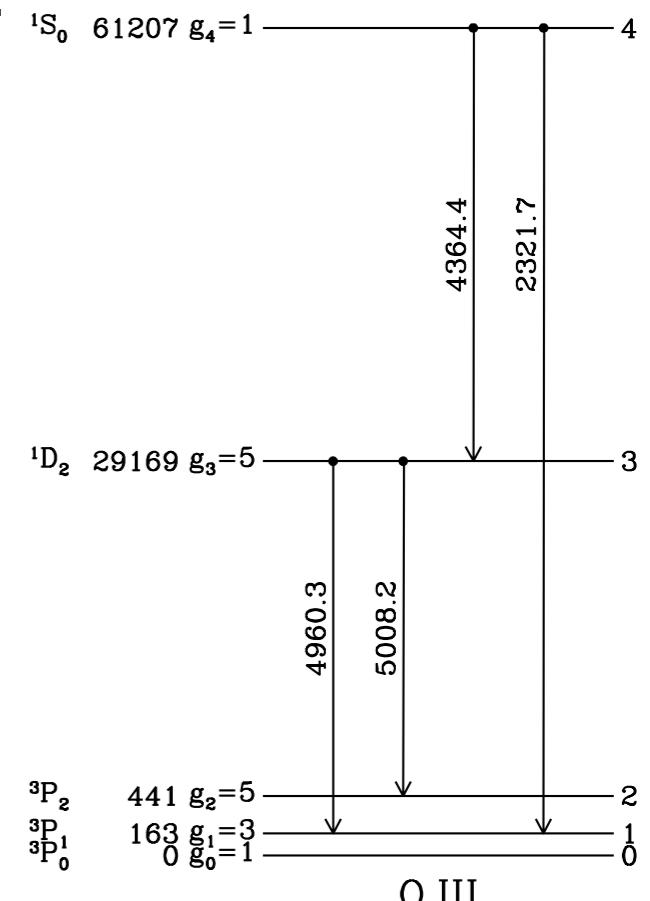
$$E_{32}/k = 29169 \text{ K}, \quad E_{\text{H}\beta}/k = 29588.5 \text{ K}$$

$$\Omega_{30} = 0.243 T_4^{0.120 + 0.031 \ln T_4}$$

$$A_{32} = 2.0 \times 10^{-2} \text{ [s}^{-1}\text{]}$$

$$A_{21} = 6.8 \times 10^{-3} \text{ [s}^{-1}\text{]}$$

$$\frac{[\text{O III}] \, 5008}{\text{H}\beta} = 5.091 \times 10^5 T_4^{0.494 + 0.089 \ln T_4} e^{-2.917/T_4} \frac{n(\text{O}^{+2})}{n(\text{H}^+)}$$



- **Nitrogen:** ($\lambda_{\text{H}\alpha} = 6562.79 \text{\AA}$)

$$4\pi j(\text{[NII]} 6585) = n_e n(\text{N}^+) k_{03} \frac{A_{32}}{A_{31} + A_{32}} E_{32}$$

$$4\pi j(\text{H}\alpha) = n_e n(\text{H}^+) \alpha_{\text{eff}, \text{H}\alpha} E_{\text{H}\alpha}$$

where

$$\alpha_{\text{eff}, \text{H}\alpha} \approx 1.17 \times 10^{-13} T_4^{-0.942 - 0.031 \ln T_4} \text{ cm}^3 \text{ s}^{-1}$$

$$k_{03} = 8.62942 \times 10^{-8} T_4^{-1/2} \frac{\Omega_{30}}{g_0} e^{-E_{30}/kT} \text{ cm}^3 \text{ s}^{-1} \quad (g_0 = 1)$$

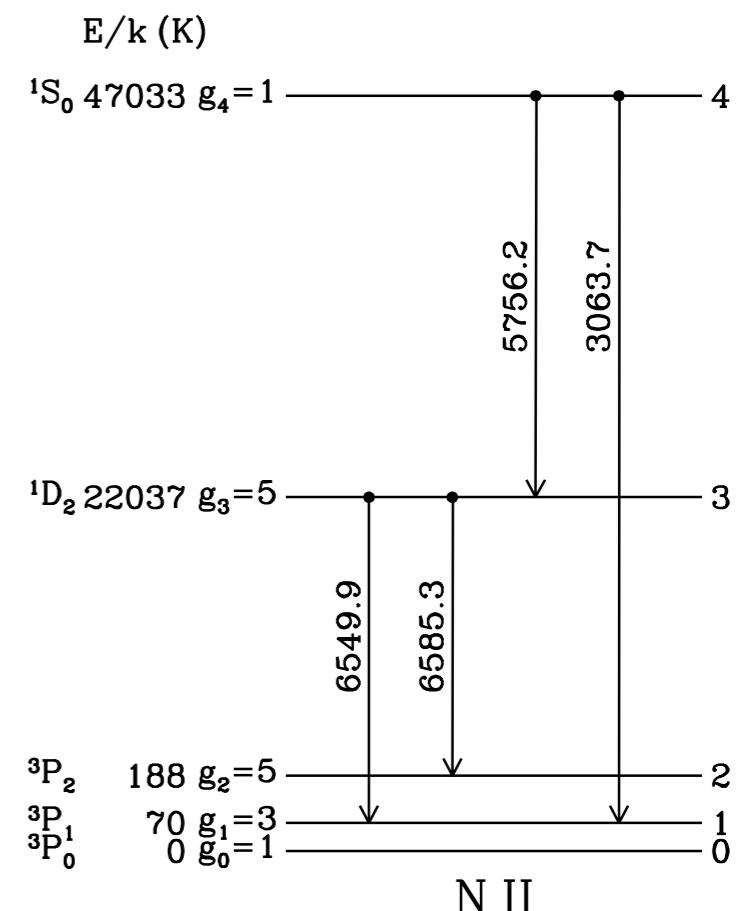
$$E_{32}/k = 21849 \text{ K}, \quad E_{\text{H}\alpha}/k = 21916.9 \text{ K}$$

$$\Omega_{30} = 0.303 T_4^{0.053 + 0.009 \ln T_4}$$

$$A_{32} = 3.0 \times 10^{-3} \text{ [s}^{-1}\text{]}$$

$$A_{31} = 9.8 \times 10^{-4} \text{ [s}^{-1}\text{]}$$

$$\frac{[\text{NII}]\ 6585}{\text{H}\alpha} = 1.679 \times 10^5 T_4^{0.495 + 0.040 \ln T_4} e^{-2.185/T_4} \frac{n(\text{N}^+)}{n(\text{H}^+)}$$



- Therefore, if the temperature T is known, the relative abundance of the ion can be obtained from the measured line ratio.
- The total elemental abundances are then obtained by applying **ionization correction factors (ICFs)**, which correct for abundances of unobserved ions.

General Multilevel Atom

- It is easy to generalize the equations of statistical equilibrium up to an arbitrary number of levels.
 - In statistical equilibrium, the rate of collisional and radiative population of any level j is matched by the collisional and radiative depopulation rates of that same level.
 - When combined with the population normalization equation (the sum of the populations of all levels must add up to the total number of ions), we have a linear set of simultaneous equations which may be solved in the standard way.

$$\sum_{i \neq j}^M n_i n_e k_{ij} + \sum_{i=j+1}^M n_i A_{ij} - n_j \left(\sum_{i \neq j}^M n_e k_{ji} + \sum_{i=1}^{j-1} A_{ji} \right) = 0 \quad j = 1, 2, 3, \dots, M$$

$$\sum_{j=0}^M n_j = 1 \quad (\text{normalization})$$

$M + 1 \equiv$ number of levels

Useful softwares to calculate line ratios.

(1) PopRatio: <http://www.ignacioalex.com/popratio/>
 Silva & Viegas (2001, Computer Physics Communications, 136, 319)

(2) PyNeb: <http://research.iac.es/proyecto/PyNeb/>; https://github.com/Morisset/PyNeb_devel
 Luridiana, Morisset, & Shaw (2015, A&A, 573, A42)

Ionization / Excitation Diagnostics: The BPT diagram

- Ionization / Excitation Mechanisms in galaxies
 - The optical line emission from star-forming galaxies is usually dominated by emission lines from H II regions.
 - Some galaxies have strong continuum and line emission from an active galactic nucleus (AGN). The line emission is thought to come from gas that is heated and ionized by X-rays from the AGN.
 - ▶ **Seyfert galaxies:** The AGN spectrum normally includes strong emission lines from high-ionization species like C IV and Ne V, which are presumed to be ionized by X-rays from the AGN. Seyfert (1943) discovered that some galaxies had extremely luminous, point-like nuclei, with emission line widths in some cases exceeding 4000 km/s.
 - ▶ **LINERs** (Low Ionization Nuclear Emission Region): In other cases, the nucleus has strong emission lines but primarily from low-ionization species. LINERs were first identified by Timothy Heckman (1980). [There are debates on the sources of ionization and line emission; AGN or star-forming regions, shock or photoionization]
- **BPT diagram:**
 - Baldwin, Phillips & Terlevich (1981) found that one could distinguish star-forming galaxies from galaxies with spectra dominated by AGNs by plotting the ratio of **[OIII]5008/H β vs. [NII]6585/H α .**
 - These lines are among the strongest optical emission lines from H II regions.
 - The line ratios employ pairs of lines with similar wavelengths so that the line ratios are nearly unaffected by dust extinction.

-
- Recall the structure of H and He ionization zone:
 - In H II regions where He is neutral (no photons with $E > 24.6 \text{ eV}$), N and O will be essentially 100% singly ionized throughout the H ionization zone.
 - On the other hand, for O stars that are hot enough ($24.6 < E < 54.6 \text{ eV}$) to have an appreciable zone of He ionization, the N and O in this zone can be doubly ionized.
 - Because **N and O have similar second ionization potentials** (29.6 and 35.1 eV for N and O, respectively), to a good approximation, HII regions will have:

$$\text{N}^+/\text{N} \approx \text{O}^+/\text{O} \quad \text{and} \quad \text{N}^{+2}/\text{N} \approx \text{O}^{2+}/\text{O}$$

- Essentially all of the gas-phase O and N in the H II region will be either singly or doubly ionized. Hydrogen will fully ionized in the H II region:

$$n(\text{N}) = n(\text{N}^+) + n(\text{N}^{+2}), \quad n(\text{O}) = n(\text{O}^+) + n(\text{O}^{+2}), \quad \text{and} \quad n(\text{H}) = n(\text{H}^+)$$

- Let's define the fraction of doubly-ionized ion.

$$\xi_{\text{N}} \equiv \frac{n(\text{N}^{+2})}{n(\text{N}^+) + n(\text{N}^{+2})} = \frac{n(\text{N}^{+2})}{n(\text{N})}$$

$$\xi_{\text{O}} \equiv \frac{n(\text{O}^{+2})}{n(\text{O}^+) + n(\text{O}^{+2})} = \frac{n(\text{O}^{+2})}{n(\text{O})}$$

Then, the fractions are approximately equal:

$$\xi_{\text{N}} \approx \xi_{\text{O}} \implies \xi$$

- We assume the N abundance to be solar, and the O abundance to be 80% solar (20% is presumed to be in silicate grains).
- In terms of the fraction,

$$\frac{n(\text{N}^+)}{n(\text{H}^+)} = (1 - \xi) \frac{n(\text{N})}{n(\text{H})}$$

$$\frac{n(\text{O}^{2+})}{n(\text{H}^+)} = \xi \frac{n(\text{O})}{n(\text{H})}$$

- Then, the line ratios can be written:

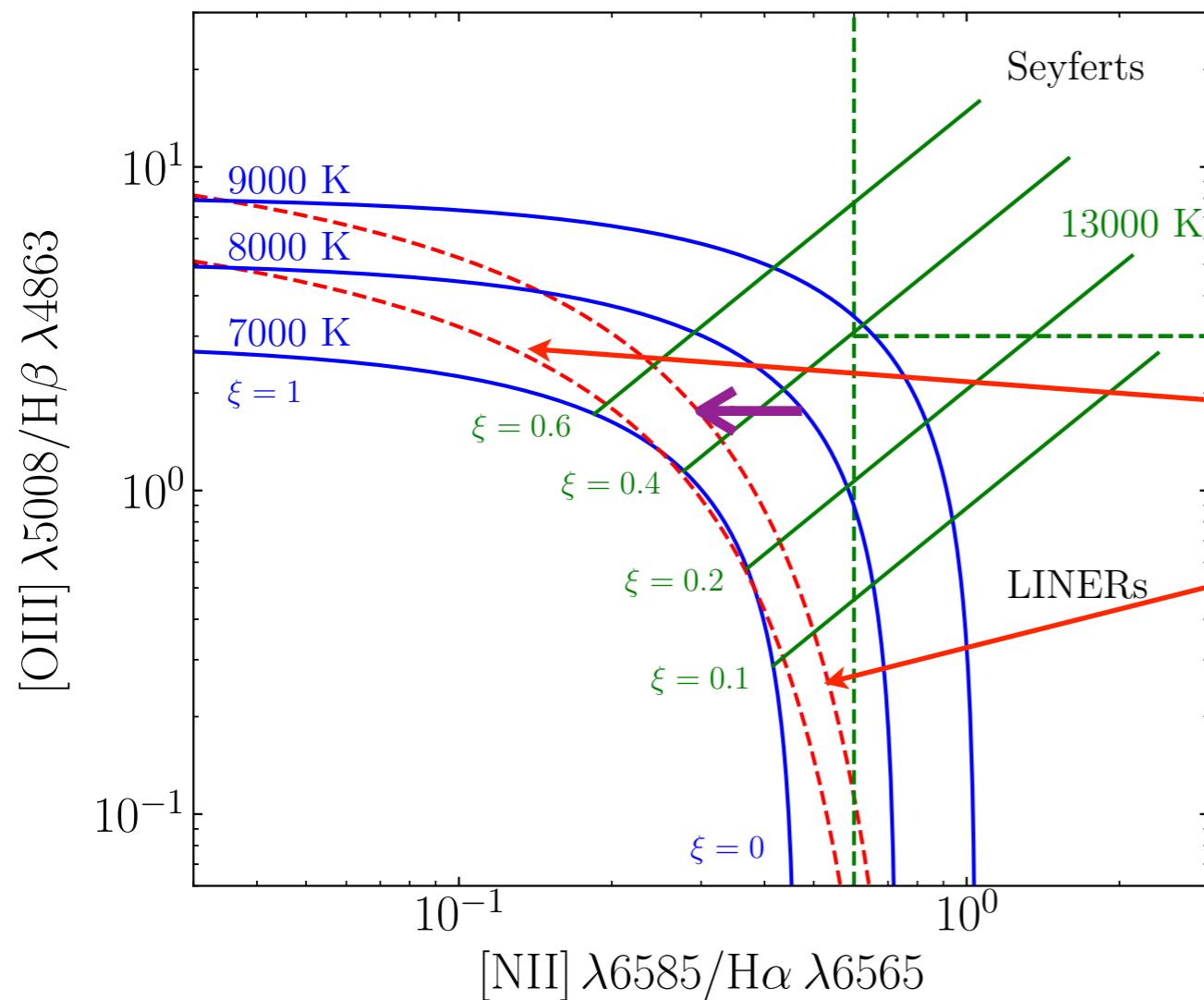
$$\frac{[\text{O III}] 5008}{\text{H}\beta} = 218.7 \xi T_4^{0.494 + 0.089 \ln T_4} e^{-2.917/T_4} \left(\frac{n_{\text{O}}/n_{\text{H}}}{0.8 \times 5.37 \times 10^{-4}} \right)$$

$$\frac{[\text{N II}] 6585}{\text{H}\alpha} = 12.44 (1 - \xi) T_4^{0.495 + 0.040 \ln T_4} e^{-2.185/T_4} \left(\frac{n_{\text{N}}/n_{\text{H}}}{7.41 \times 10^{-5}} \right)$$

- For an assumed temperature T, we can produce a theoretical curve of [OIII]5008/H β versus [NII]6585/H α by varying the fraction ξ of the N and O that is doubly ionized.

The theoretical curve of [OIII]5008/H β versus [NII]6585/H α that is obtained by varying the fraction is shown below:

Blue lines show the tracks of the equations calculated, by varying ξ from 0 to 1, for $T = 7000, 8000$, and 9000 K.



Empirical curves that discriminate the star-forming galaxies from AGNs.

$$\log_{10} ([\text{OIII}]/\text{H}\beta) = 1.10 - \frac{0.60}{0.01 - \log_{10} ([\text{NII}]/\text{H}\beta)}$$

$$\log_{10} ([\text{OIII}]/\text{H}\beta) = 1.3 - \frac{0.61}{0.05 - \log_{10} ([\text{NII}]/\text{H}\beta)}$$

Kauffmann et al. (2003, MNRAS, 346, 1055)

To derive the equation, we assumed that $\xi_N = \xi_O = \xi$. The discrepancy between the simple model with the observations would be due to the assumption. We need to note that $E(\text{N}^+ \rightarrow \text{N}^{2+}) < E(\text{O}^+ \rightarrow \text{O}^{2+})$, which implies that $\xi_N \gtrsim \xi_O$.

The arrow ← indicates the direction that is expected from $\xi_N > \xi_O$.

- Line ratios for 122,514 galaxies in SDSS DR7 with S/N > 5.

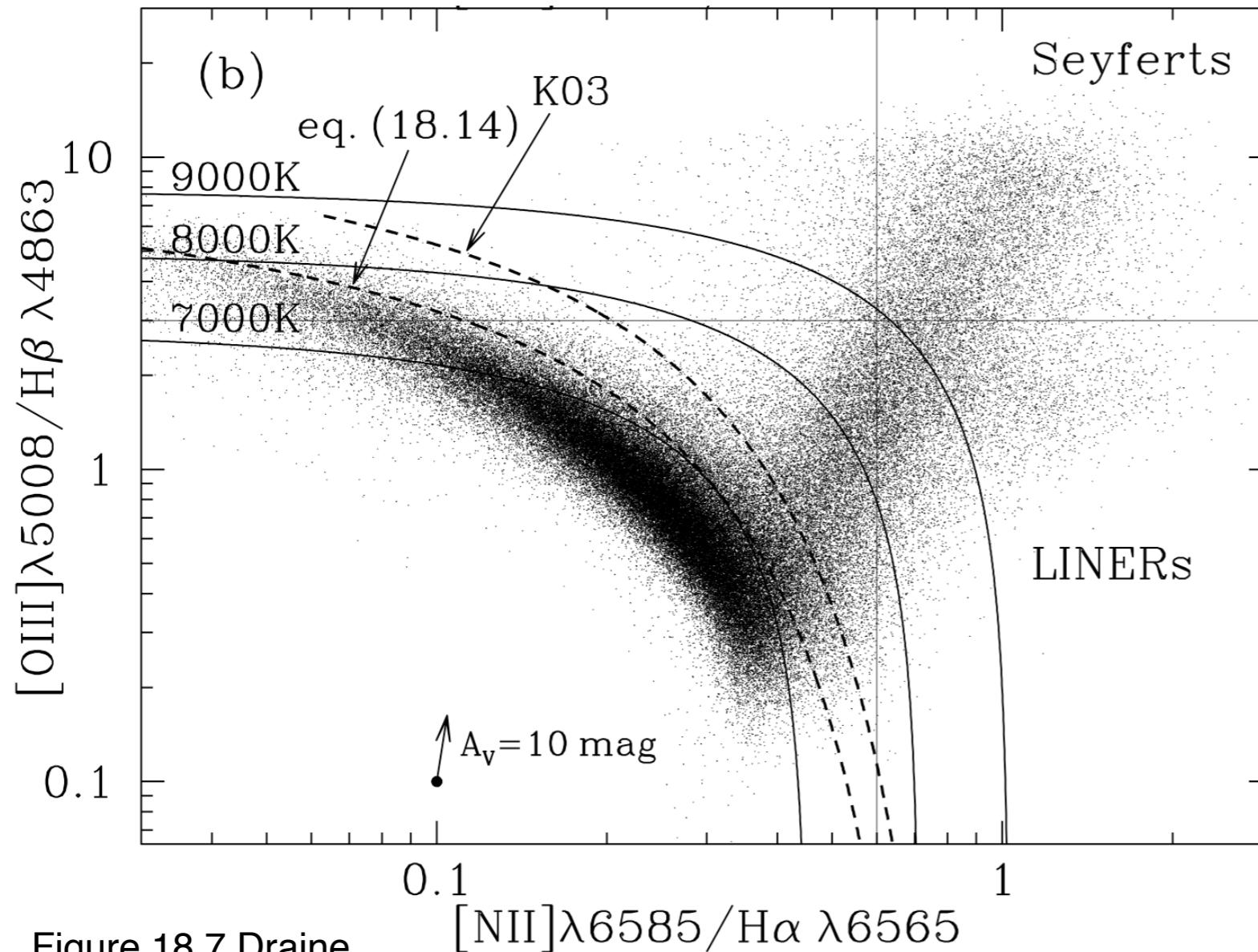


Figure 18.7 Draine



Flying seagull [Stasinska]

Photons from AGNs are harder than those from the massive stars that power H II regions.

They induce more heating, implying that optical collisionally excited lines will be brighter with respect to recombination lines than in the case of ionization by massive stars.

The heating by an AGN boosts the [N II] line and creates a clear separation of the two wings.

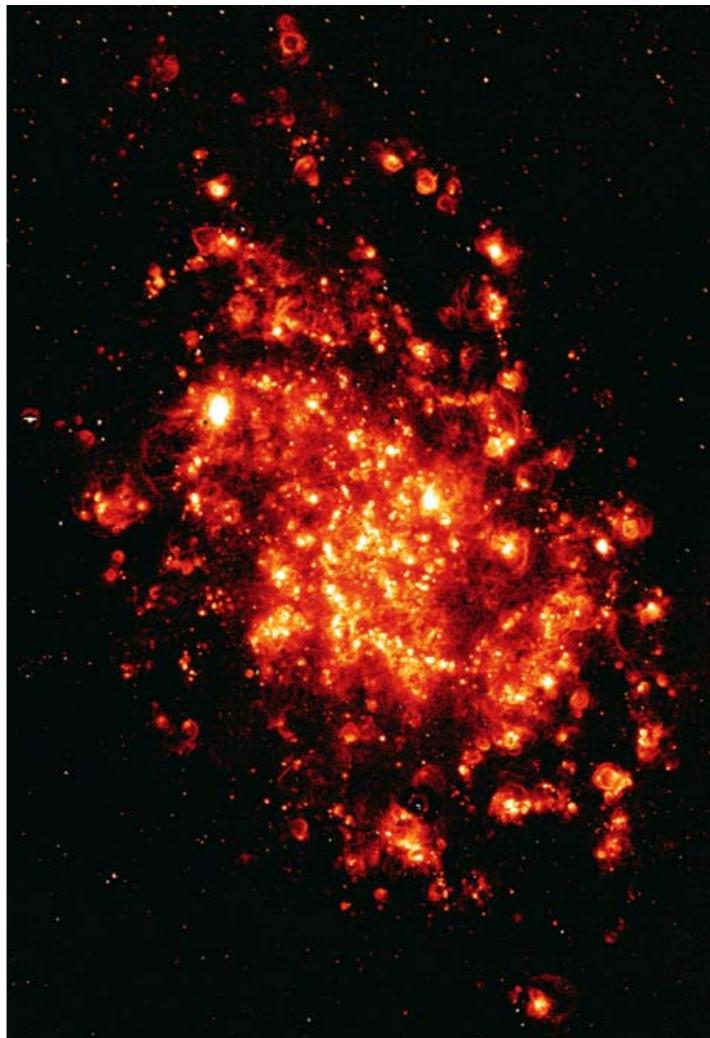
Warm Ionized Medium / Diffuse Ionized Gas

- ***Outside well-defined H II regions, the ISM is known to contain diffuse ionized gas***
 - which can originate either from leaks of ionized gas out of H II regions due to the champagne effect, or from ionization by the UV radiation of isolated hot stars, and perhaps from other mechanisms.
 - In our Galaxy, the DIG is known to contain much more mass than the H II regions. Its total mass is of the order of 1/3 of that of H I.



M51 (NGC5195)
Plate 1 [Lequeux]

B band - blue
V band - green
H α - red

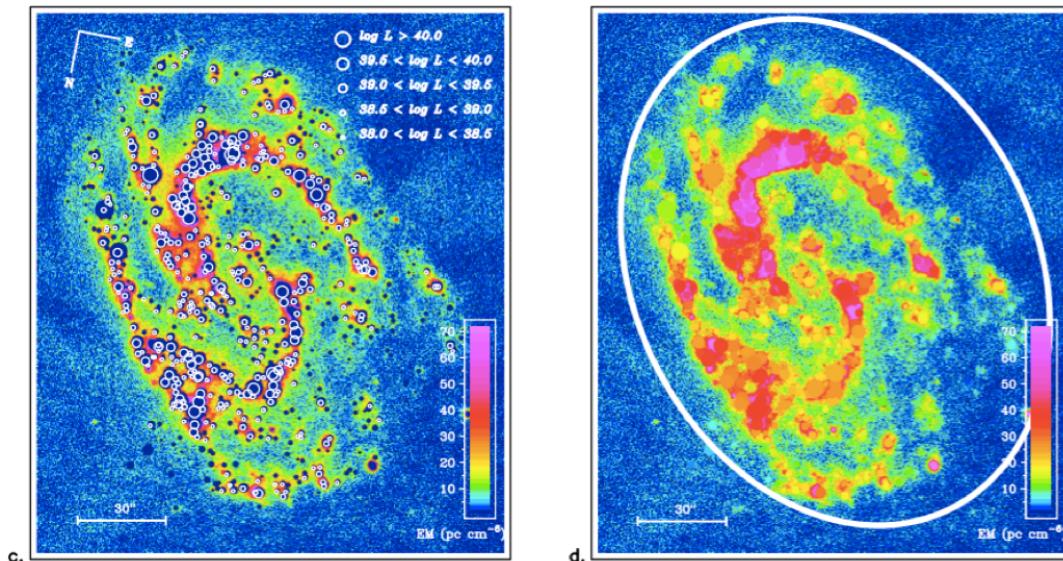


Ionized gas in M33
Plate 10 [Lequeux]

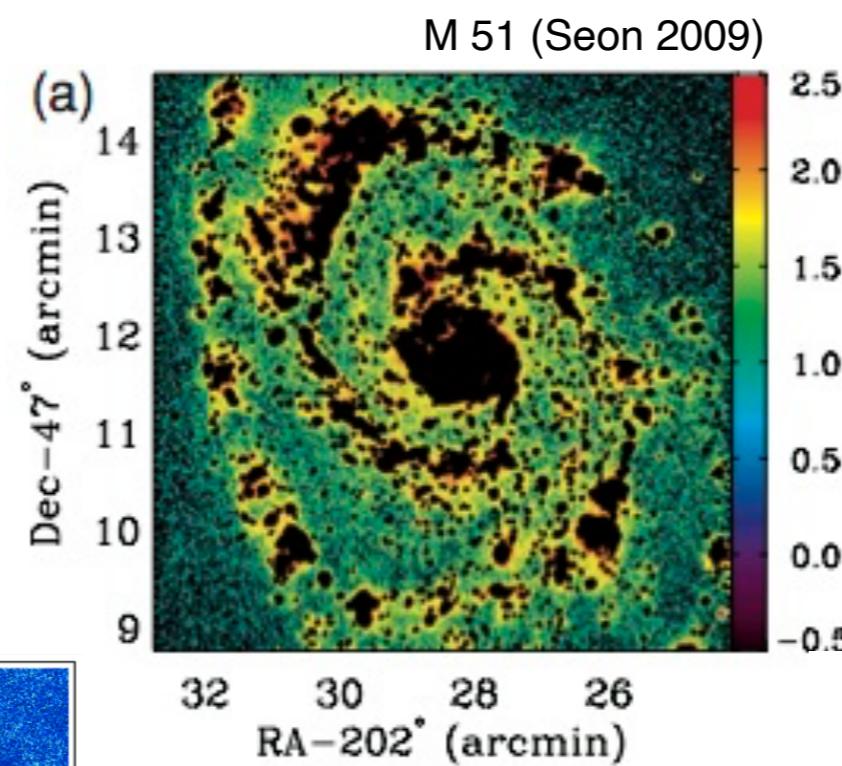
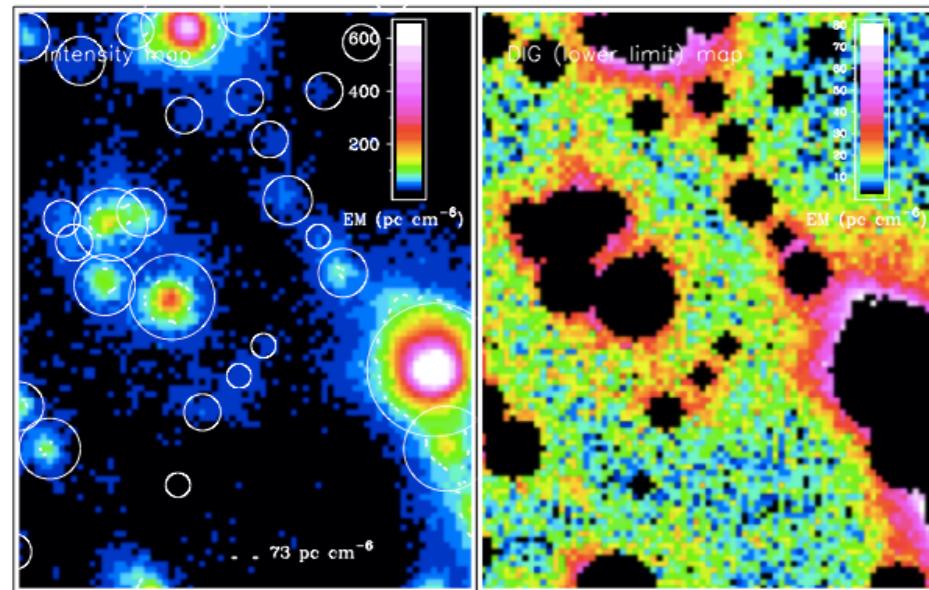
Many bubbles and
the DIG present
almost everywhere
in the central
region.

- ▶ Face-on & edge-on galaxies
- ▶ H α flux of WIM $\sim 20\%-60\%$ of the total H α flux.

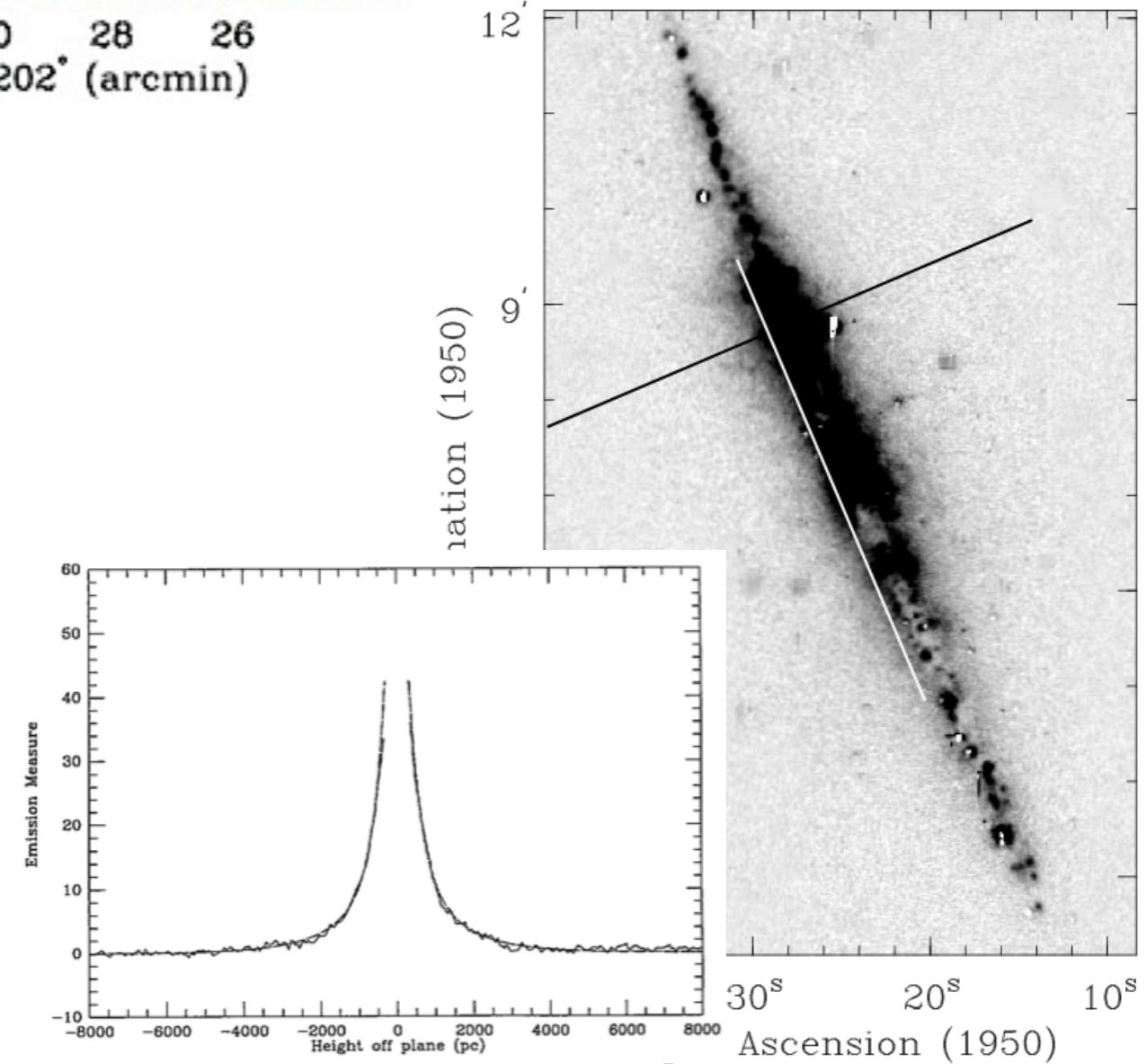
NGC 157 (Zurita e al. 2000)



NGC 157

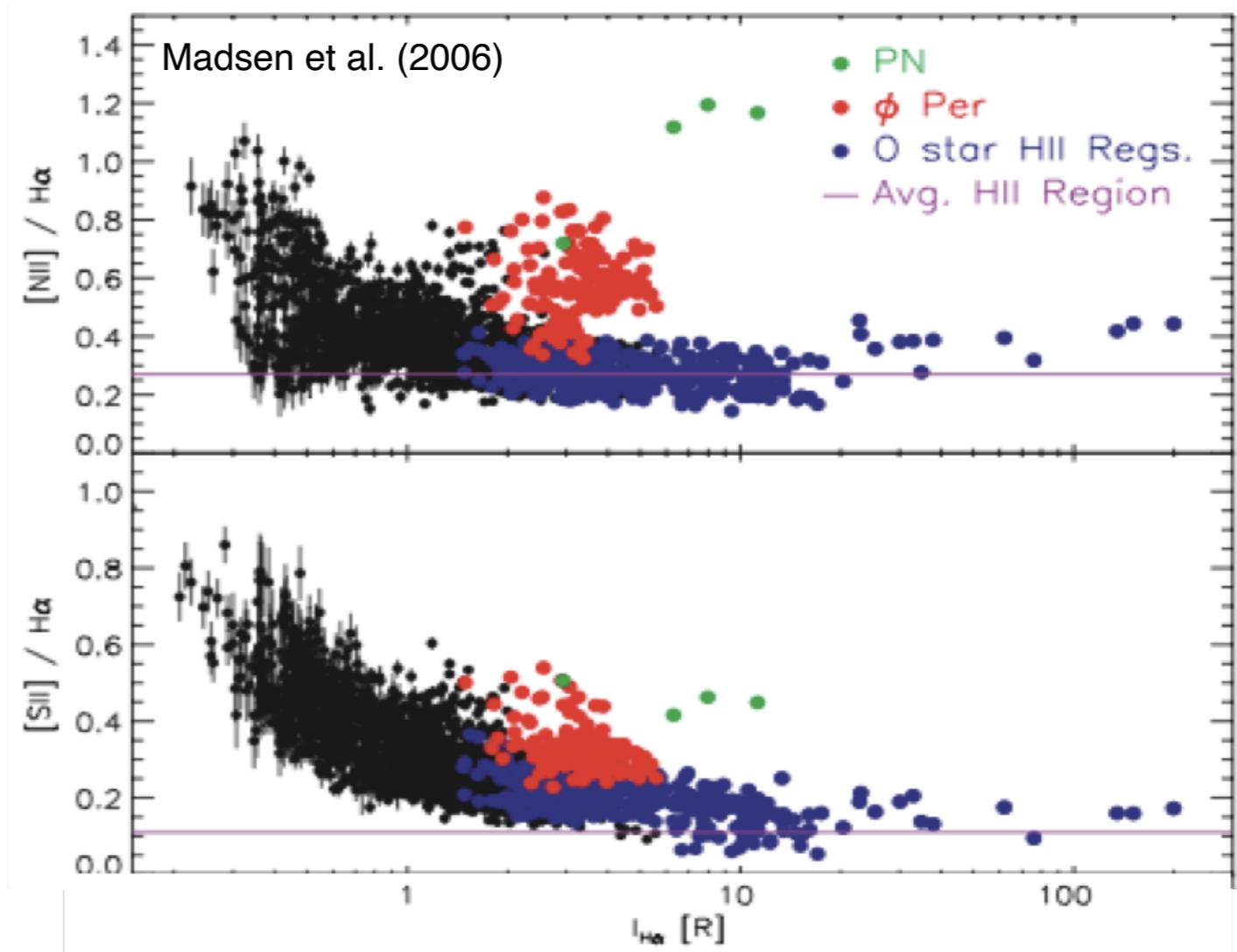
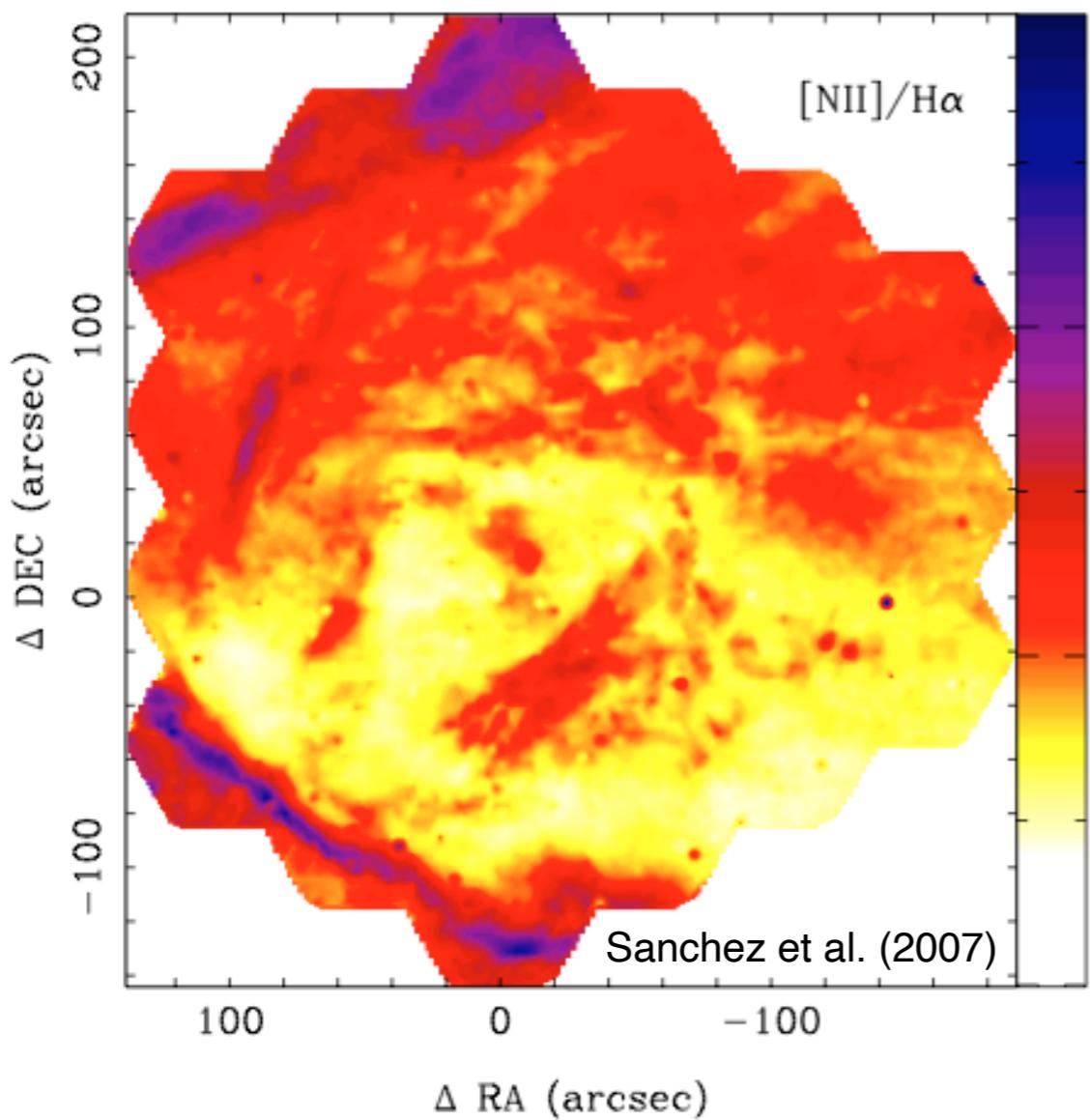


NGC 891 (Rand et al. 1998)

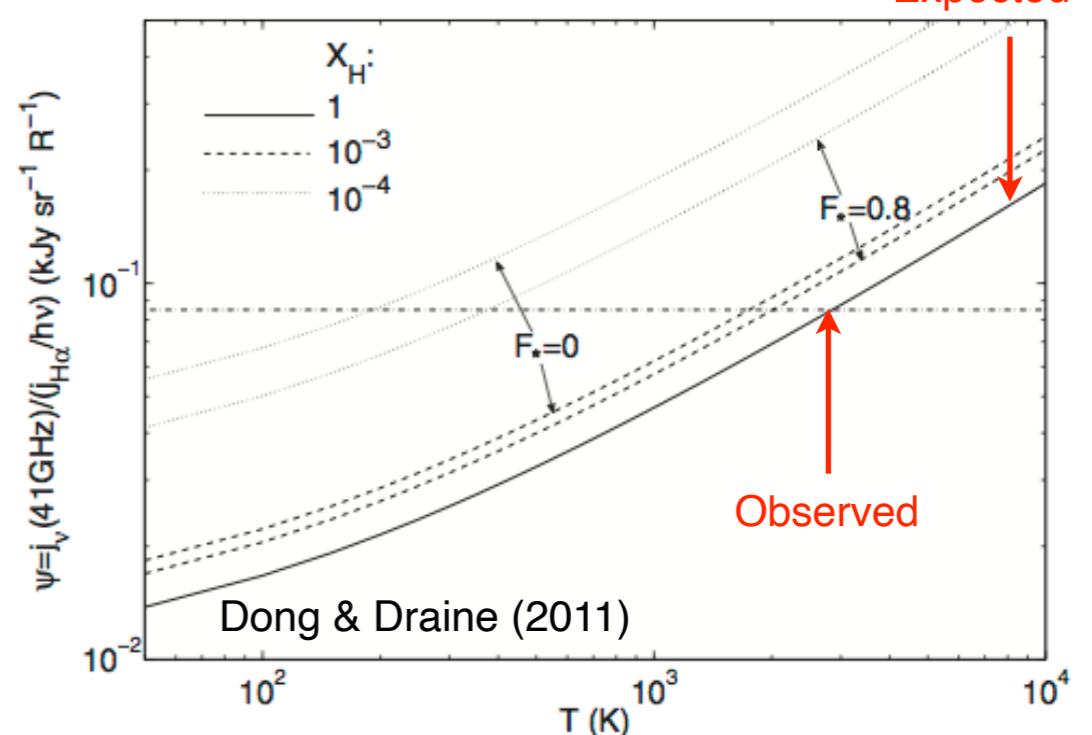
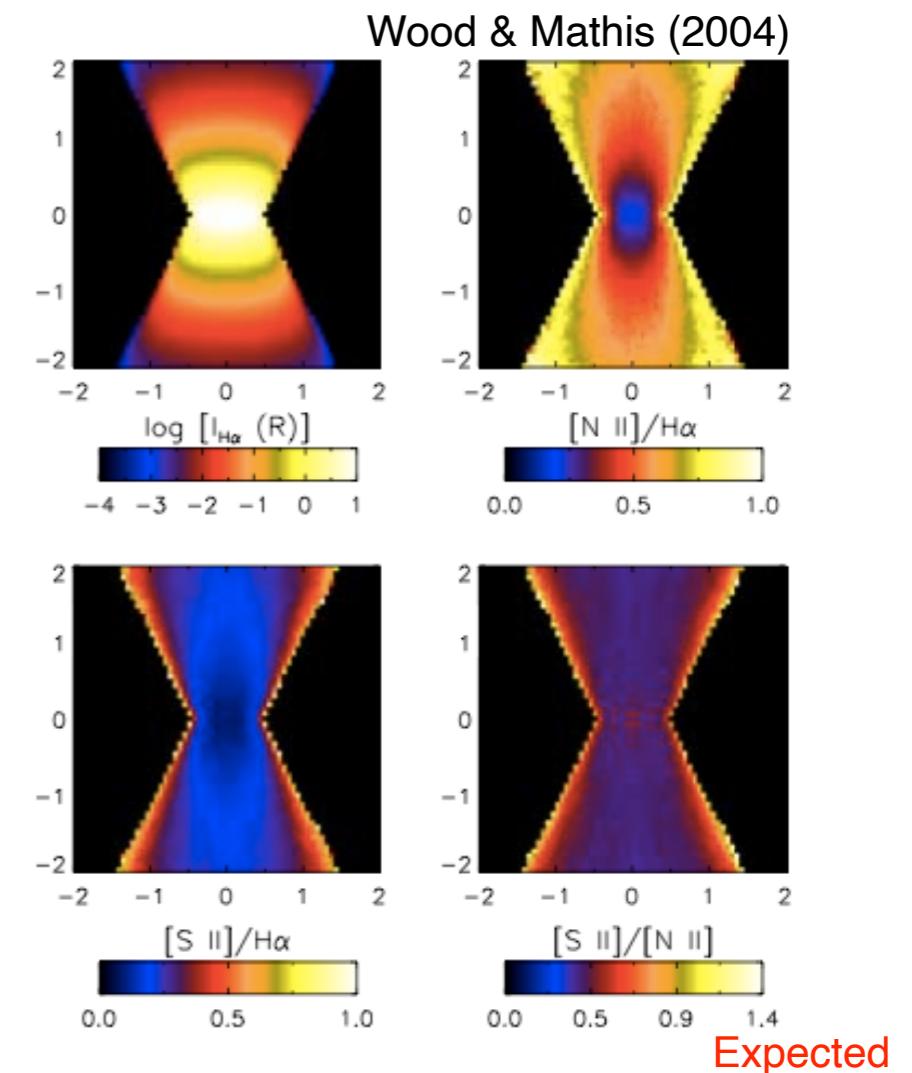


Optical Line Ratios

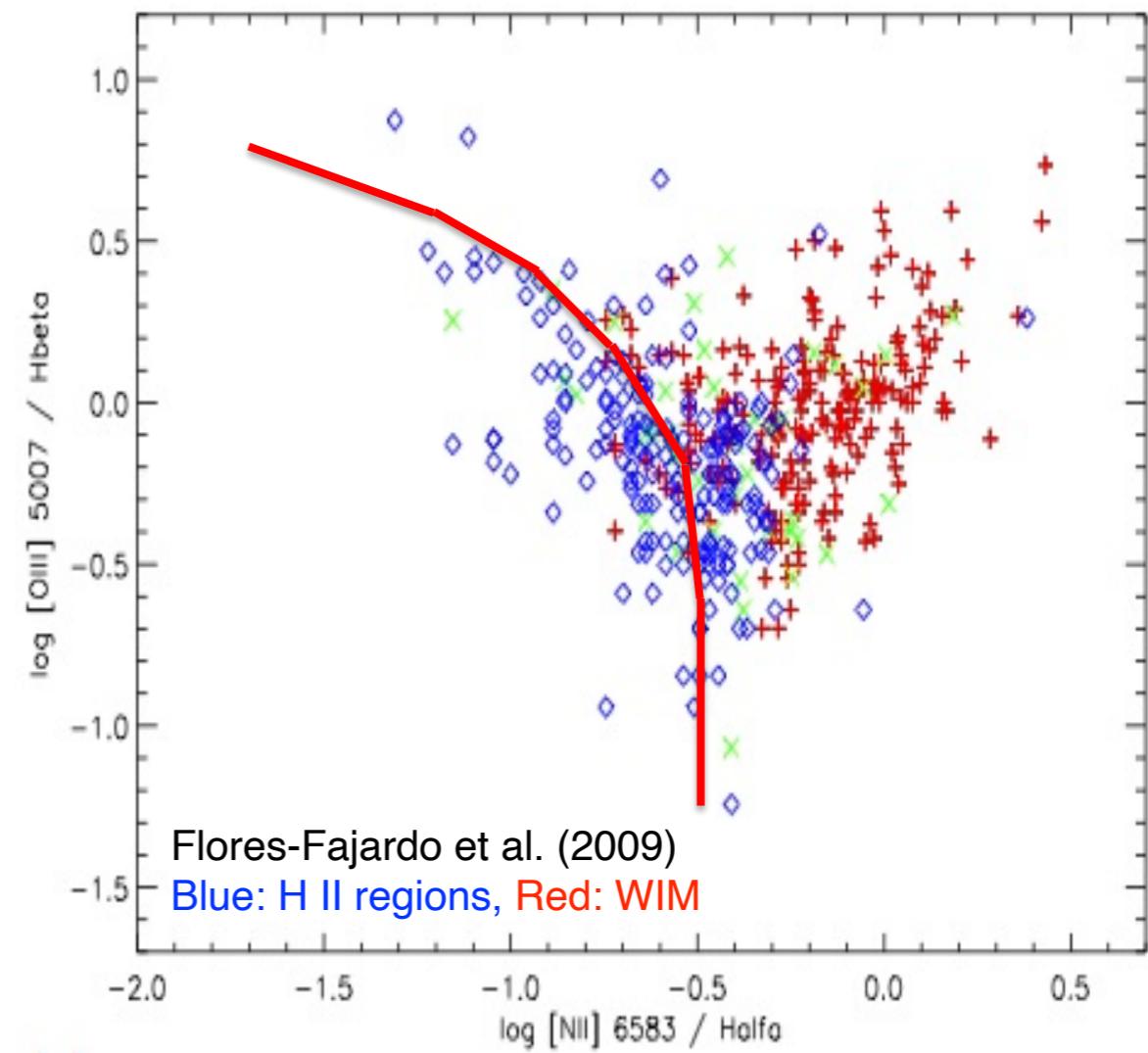
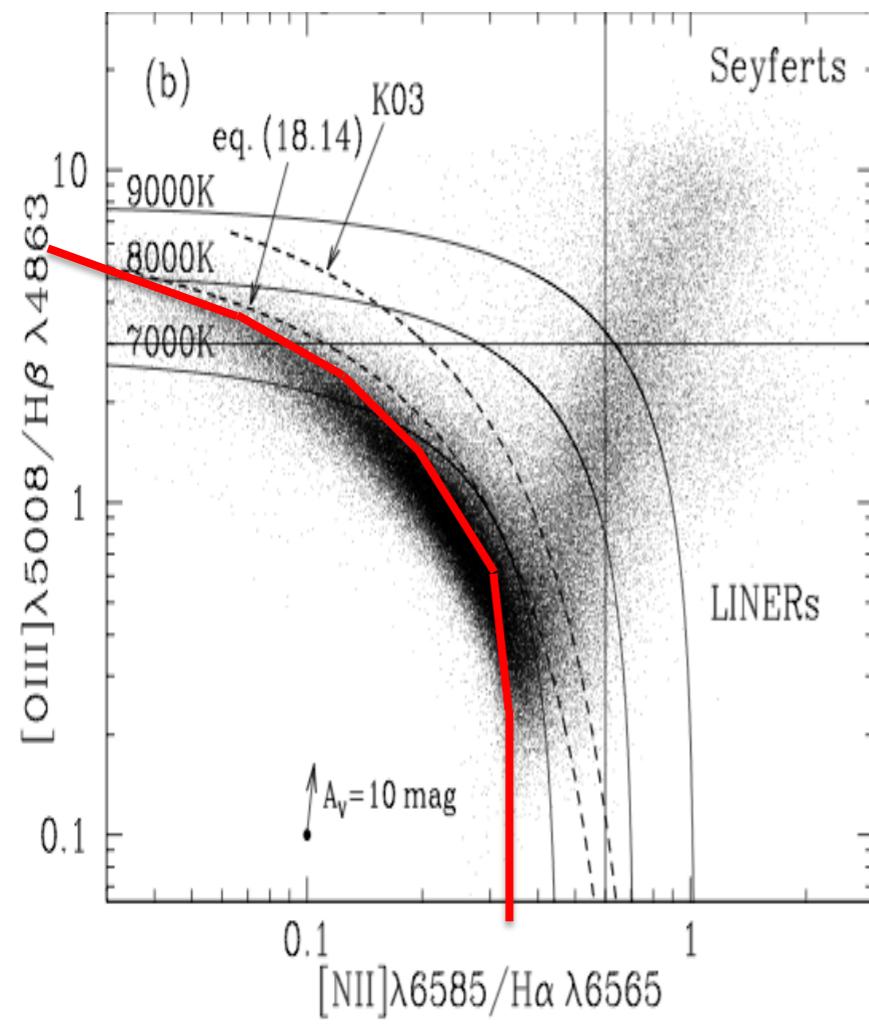
- $[\text{N II}] \lambda 6583/\text{H}\alpha$ and $[\text{S II}] \lambda 6716/\text{H}\alpha$ in the diffuse regions are generally higher than the ratios in bright H II regions.
 - $[\text{N II}]/\text{H}\alpha \approx 0.25$ and $[\text{S II}]/\text{H}\alpha \approx 0.1$ in bright H II regions
 - $[\text{N II}]/\text{H}\alpha \approx 0.3-0.6$ and $[\text{S II}]/\text{H}\alpha \approx 0.2-0.4$ in the diffuse ISM regions.



- Ionizing Mechanism
 - Only the O stars meet and surpass the power requirements to ionize the diffuse ISM.
 - Density-bounded (leaky) H II regions
 - ▶ Turbulent or clumpy morphology of the ISM
 - ▶ Existence of enormous, H I-free bubbles/holes surrounding the O stars
- However, this mechanism can not explain the free-free radio emission
 - The WMAP data shows that the observed ratio of free-free radio continuum to H α is at least twice smaller than the expected value.
 - Davies et al. (2006, 2009), Dobler & Finkbeiner (2008a,b), Gold et al. (2011)
- See Seon & Witt (2012), and Dong & Draine (2011) for the possibility of alternative explanations.



Heating source of the WIM



Detailed analyses of the line ratios $[N II]/H\alpha$ and $[S II]/H\alpha$ have been performed to obtain the temperature; it is found to be of the order of 8,000 K.

However, it is also known that WIM is being heated by additional sources. (see the above BPT diagram).

Dynamics of H II regions

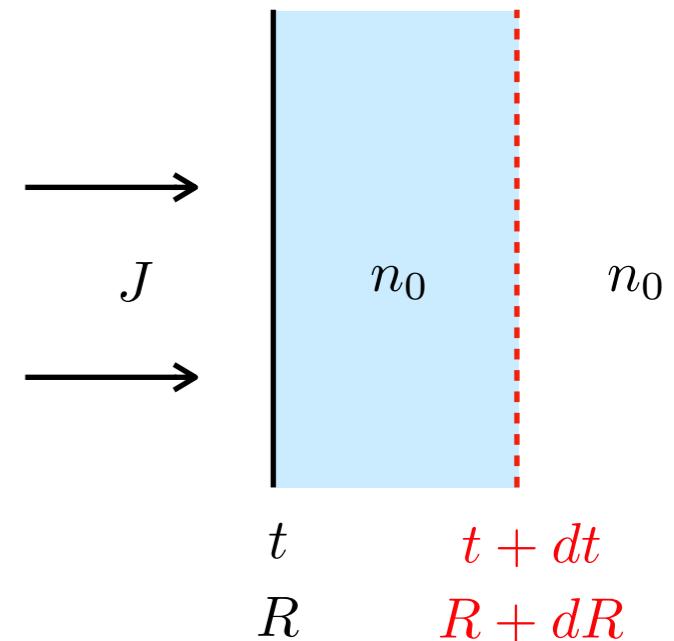
- A new O star is presumably born within clouds of relatively dense cold gas. The appearance of a source of UV photons will have two effects.
 - First, the gas surrounding the new star will become ionized. Since the mean free path of an UV photon is very short in neutral hydrogen, the photons will be absorbed in a relatively thin surrounding shell of neutral hydrogen, producing new ionization. Thus ***the ionized and neutral gases are separated by an ionization front, which moves rapidly outward*** as more and more atoms become ionized by the stream of photons.
 - Second, the process increases the gas temperature from $\sim 10^2$ K to $\sim 10^4$ K, by a factor of about a hundred. the ionization process itself increases the number of gas particles, by a factor two. As a result, ***the pressure in the ionized gas is ~200 times greater than that in surrounding neutral material.*** This ionized gas cannot be confined and will expand. The ionized and neutral gas are set in motion.
 - Since the expansion velocity is likely to exceed the sound velocity in the surrounding H I region, ***a shock front may be expected to form***, moving out through the neutral gas. The dynamical analysis of H II regions must consider the interactions between the ionization front and the shock front, together with the equations of motion of the gas behind the two fronts.
- This process is not the only way in which ISM is set in motion by means of interaction with stars.
 - There are effects produced by the very high speed continuous mass loss - a ***stellar wind***.
 - Many massive stars terminate their existence in a violent explosive event - a ***supernova***.

-
- Basic Assumptions:
 - Any disturbances to the cloud structure produced by the formation of a star are neglected. *After a relatively short time ($< 10^5$ yr), the star reaches a static configuration* in which it can remain for a much longer time ($> 3 \times 10^6$ yr). The stellar radiant energy output rate and the spectral distribution of the radiation are more or less constant during this phase. The star then produces Lyman continuum photons at a constant rate. Since *the star formation time scale is so short, we may take the star to be ‘switched on’ instantaneously.*
 - The gas around the star will be assumed to be at rest (in the frame of reference of the star).
 - The gas has initially assumed to be uniform in density and temperature.
 - Ionization front
 - The term “front” describes a more-or-less abrupt boundary between two regions of the ISM with very different properties.
 - An ionized nebula can be approximated as a region of highly ionized gas, separated from the surrounding neutral medium by a thin boundary region, of thickness $\lambda_{\text{mfp}} \approx 0.002$ pc. Thus, an H II region is surrounded by an ionization front.

The velocity of the ionization front

- Suppose that at time t the ionization front is located at a distance R from the star and at time $t + dt$ it is at a distance $R + dR$.

- Let n_0 = number density of the undisturbed neutral hydrogen
 - J = number of Lyman continuum photons incident normally on unit area of the ionization front per unit time.



- Ionization balance at the ionization front:** While the ionization front moves from R to $R + dR$, the photons will ionize all the neutral atoms lying between these two positions ($R, R + dR$).
 - We assume that only one photon is needed to ionize each atom as the front moves the distance dR . In other words, no recombination occurs within the distance interval dR . For unit area of the ionization front, the following relation must be satisfied:

$$J \Delta Adt = n_0 \Delta AdR$$

- Then, the velocity of the ionization front (in a fixed frame of reference) is:

$$\frac{dR}{dt} = \frac{J}{n_0}$$

The initial stage of evolution of an ionized region

- Suppose that the UV source has been suddenly turned on.
- ***Ionization balance for the ionized region:*** We consider two factors:
 - *The radiation field at the ionization front is diluted because of the spherical geometry.*
 - Recombination takes place continuously inside the ionized region, and *some of the UV photons produced by the central source must go to reionize the atoms that have recombined.*
- Inside the ionized sphere, the fractional ionization is near unity. Thus, $n_e = n_p = n_0$. Using this condition, we obtain an equation for the expansion velocity of the ionization front.

$$\frac{J}{n_0} = \frac{dR}{dt} = \frac{Q_0}{4\pi R^2 n_0} - \frac{1}{3} R n_0 \alpha_B$$

- Let's define the following dimensionless quantities:

$$\rho \equiv R/R_s \quad \text{where } R_s \equiv \left(\frac{3}{4\pi} \frac{Q_0}{\alpha_B n_0^2} \right)^{1/3}$$

$$\tau \equiv t/t_{\text{rec}} \quad \text{where } t_{\text{rec}} \equiv \frac{1}{\alpha_B n_0}$$

Then, the equation in dimensionless form is

$$\frac{d\rho}{d\tau} = \frac{1}{3} \left(\frac{1}{\rho^2} - \rho \right)$$

- ▶ The equation can be written:

$$\frac{d\rho}{d\tau} = \frac{1}{3} \left(\frac{1}{\rho^2} - \rho \right) \rightarrow \frac{d\rho^3}{d\tau} = 1 - \rho^3$$

- ▶ It's solution is

$$\rho^3 = 1 - e^{-\tau}$$

$$R(t) = R_s \left(1 - e^{-t/t_{\text{rec}}} \right)^{1/3}$$

initial condition: $R(t = 0) = 0$

$$\begin{aligned} \frac{dx}{d\tau} + x &= 1 \\ e^\tau \frac{dx}{d\tau} + e^\tau x &= e^\tau \\ \frac{d(e^\tau x)}{d\tau} &= e^\tau \end{aligned} \quad \rightarrow \quad \begin{aligned} e^\tau x &= \int_0^\tau e^{\tau'} d\tau' = e^\tau - 1 \\ x &= 1 - e^{-\tau} \end{aligned}$$

- Scale Parameters:

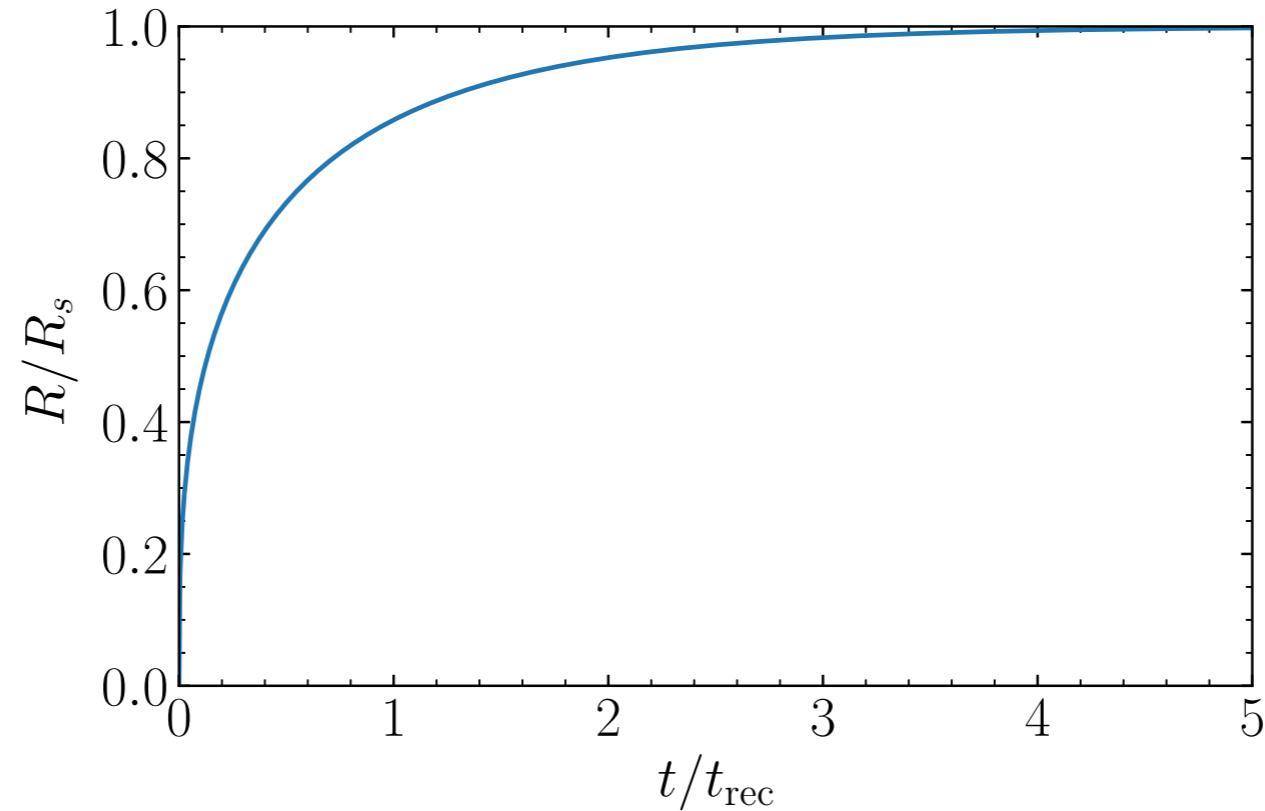
- The time scale introduced is the recombination time scale:

$$t_{\text{rec}} \equiv \frac{1}{\alpha_B n_0} \approx 4000 \text{ yr} \left(\frac{\alpha_B}{2.6 \times 10^{-3} \text{ cm}^3 \text{ s}^{-1}} \right)^{-1} \left(\frac{n_0}{30 \text{ cm}^{-3}} \right)^{-1}$$

the length scale introduced is the Strömgren radius:

$$R_s \equiv \left(\frac{3}{4\pi} \frac{Q_0}{\alpha_B n_0^2} \right)^{1/3} \approx 7 \text{ pc} \left(\frac{Q_0}{10^{49} \text{ s}^{-1}} \right)^{1/3} \left(\frac{\alpha_B}{2.6 \times 10^{-3} \text{ cm}^3 \text{ s}^{-1}} \right)^{-1/3} \left(\frac{n_0}{30 \text{ cm}^{-3}} \right)^{-2/3}$$

-
- Hence, the time required to create a Strömgren sphere after turning on a hot star is an order of ~ 4000 yr. This is also the time it takes the ionized Strömgren sphere to revert to neutral gas after the central UV source has been turned off.



- At times $t \gg t_{\text{rec}} \sim 4000$ yr , the gas medium will be fully ionized with radius $R \sim R_s \sim 7$ pc, surrounded by a partially ionized boundary of thickness $\sim \lambda_{\text{mfp}} = (n_{\text{H}}\sigma_{\text{pi}})^{-1} \sim 0.002$ pc $\ll R_s$.

- We can compute the ***rate of expansion of the ionization front***:

$$\frac{dR}{dt} = \frac{R_s}{3t_{\text{rec}}} \frac{e^{-t/t_{\text{rec}}}}{(1 - e^{-t/t_{\text{rec}}})^{2/3}}$$

where the characteristic expansion velocity is

$$v_* \equiv \frac{R_s}{3t_{\text{rec}}} \simeq 560 \text{ km s}^{-1} \left(\frac{Q_0}{10^{49} \text{ s}^{-1}} \right)^{1/3} \left(\frac{n_0}{30 \text{ cm}^{-3}} \right)^{1/3}$$

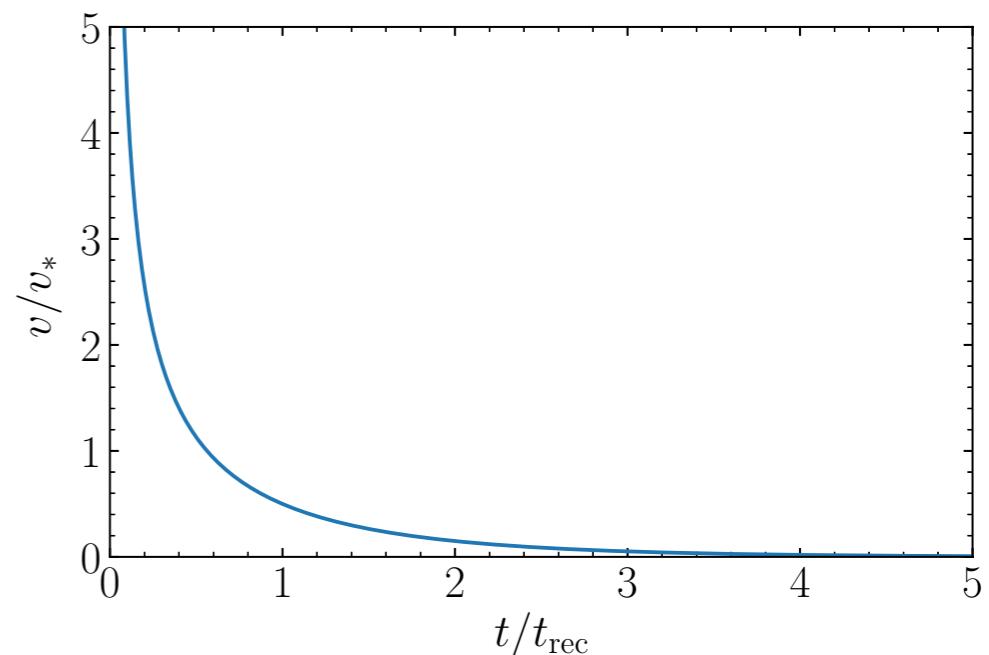
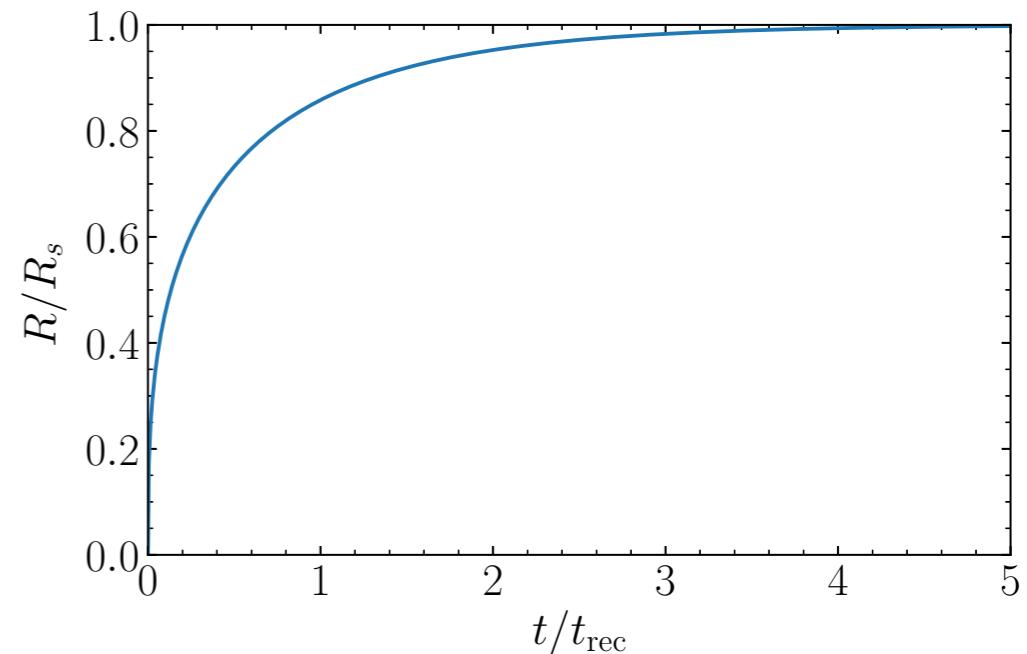
This is much larger than the sonic speed $c_s \approx 1 \text{ km s}^{-1}$ in the neutral medium as well as $c_s \approx 10 \text{ km s}^{-1}$ in the ionized medium.

- The expansion speed of the ionization front at two limits:

$$\frac{dR}{dt} \approx \frac{R_s}{3t_{\text{rec}}} \left(\frac{t}{t_{\text{rec}}} \right)^{-2/3} \quad \text{for } t \ll t_{\text{rec}}$$

$$\frac{dR}{dt} \approx \frac{R_s}{3t_{\text{rec}}} e^{-t/t_{\text{rec}}} \quad \text{for } t \gg t_{\text{rec}}$$

Note that the expansion speed diverges at $t = 0$.



-
- The ionization front will initially expand supersonically. When will the ionization front expand at subsonic speeds?

$$\frac{dR}{dt} = \frac{R_s}{3t_{\text{rec}}} e^{-t/t_{\text{rec}}} \lesssim c_i \quad c_i \approx 13 \text{ km s}^{-1} \quad \text{sound speed in the ionized medium}$$

$$t \lesssim t_{\text{sonic}} \equiv t_{\text{rec}} \ln \left(\frac{R_s}{3t_{\text{rec}}} \frac{1}{c_i} \right) \approx 3.8t_{\text{rec}} \simeq 15,000 \text{ yr}$$

- At this time, the ionization front will have a size of:

$$R(t = t_{\text{sonic}}) = R_s (1 - e^{-3.8})^{1/3} = 0.9925 R_s$$

- The ionization front will expand at a supersonic velocity until $t \approx t_{\text{sonic}}$ ($\sim 15,000$ yr). By that time, the ionized sphere has reached a radius $R \sim 0.99R_s$ and then it starts to expand at subsonic speed.
- ***At $t = R_s/c_s \sim 0.5$ Myr, the gas starts to flow outward as a result of the pressure gradient that has build up.***

The final stage of evolution of an ionized region

- Although the ionized sphere approaches ionization equilibrium at $t \gtrsim t_{\text{rec}}$, it would be still far from pressure equilibrium.
 - Outside the ionized zone, it will be embedded in the cold neutral medium with a temperature $T \sim 100$ K.
 - Inside the sphere, the heating and cooling processes yield a temperature of $T \sim 10,000$ K.
 - Also, the density of particles inside the ionized sphere will double when the hydrogen is ionized.
 - Thus, *the pressure inside the sphere will be ~ 200 times higher than the pressure outside, meaning that the ionized gas will begin to expand.*
 - The ionized gas expands as long as it has a higher pressure than its surroundings. This expansion produces a shock and will cease when the hot ionized gas reaches pressure equilibrium with the surrounding cold neutral gas.
- ***The condition of final pressure equilibrium*** can be written in the form:

$$2n_f k T_i = n_0 k T_n$$

n_f = number density of the ionized hydrogen.

T_i and T_n = temperatures of the ionized and neutral gas, typically $T_i = 10^4$ K, $T_n = 10^2$ K .

-
- The ionized gas sphere must still absorb all the stellar UV photons. Thus,

$$Q_0 = \frac{4}{3}\pi R_f^3 n_f^2 \alpha_B$$

Here, R_f is the final radius of the ionized gas sphere. From the pressure equilibrium condition, we obtain the final size:

$$n_f = (T_n/2T_i)n_0 \approx 0.005n_0 \quad \rightarrow \quad R_f = (2T_i/T_n)^{2/3} R_{s0} \approx 34R_{s0}$$

- The ratio of the mass of gas finally ionized to that contained within the initial Strömgren sphere is:

$$\frac{M_f}{M_s} = \frac{R_f^3 n_f}{R_{s0}^3 n_0} = \frac{2T_i}{T_n} \approx 200$$

- This indicates that *the initial Strömgren sphere contains only a very small fraction of the material which, in principle, a star could ultimately ionize.*

The intermediate stage of evolution of an ionized region

- Before the pressure equilibrium is established, the gas density and temperature will be

$$n_i \approx 2n_0 > n_f \quad \text{and} \quad T_i = 10^4 \text{ K}$$

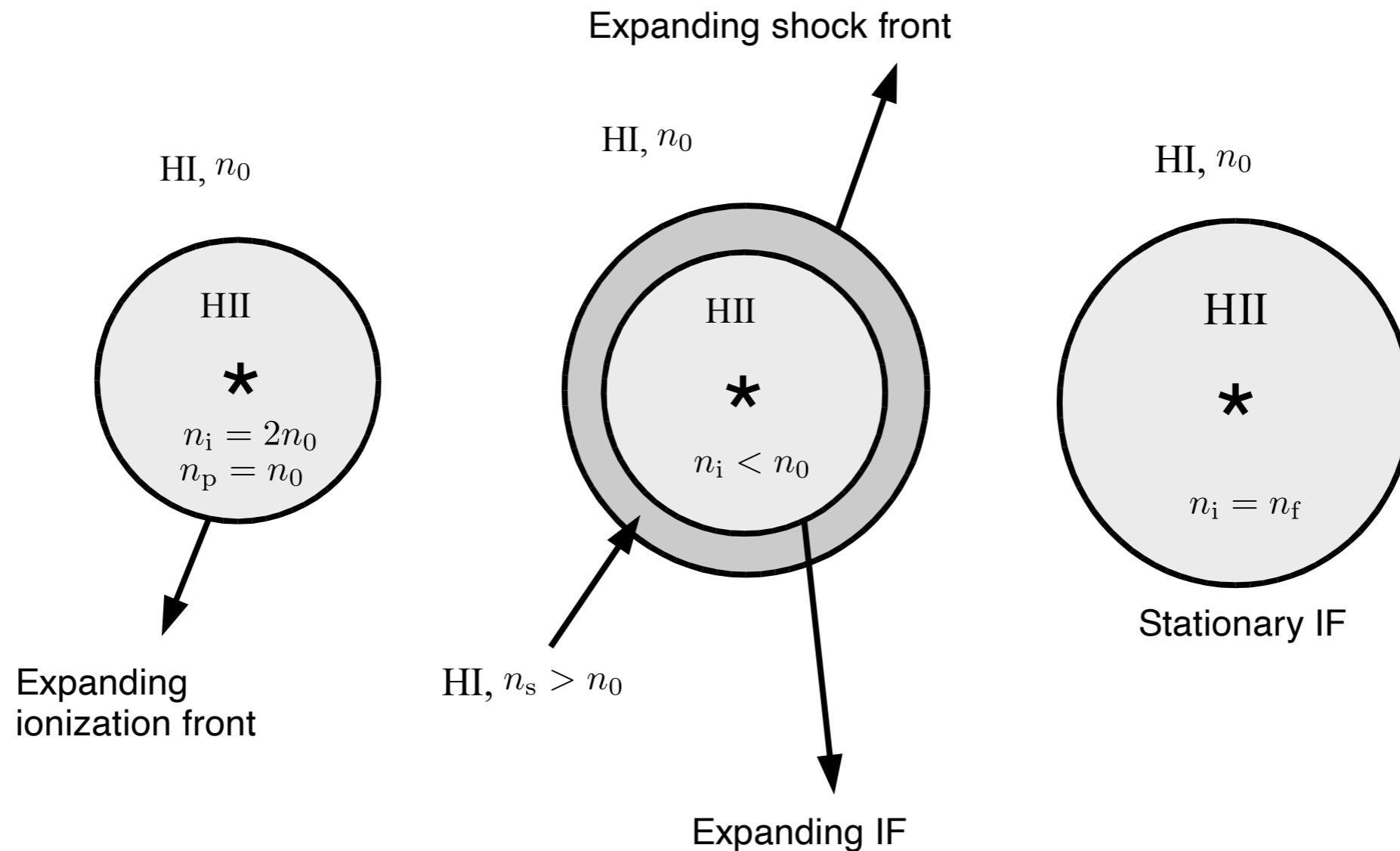
- Then the isothermal sound speeds of the ionized gas and neutral gas are, respectively:

$$c_i^2 = \frac{P_i}{\rho_i} \approx \frac{2n_0 k T_i}{n_0 m_H} \quad c_n^2 = \frac{P_n}{\rho_n} = \frac{n_0 k T_n}{n_0 m_H}$$

$$\frac{c_i}{c_n} = \left(\frac{n_i T_i}{n_0 T_n} \right)^{1/2} \approx \sqrt{200} = 14.14$$

- The sound speed of the ionized gas is much larger than that of the neutral gas.
- The ionized gas has a higher pressure and thus plays the role of a piston and pushes a shock wave into the neutral gas. *The expansion speed of the ionized gas is originally equal to about c_i , which is highly supersonic with respect to the sound speed in the neutral gas.*
- Note also that, at $t \gtrsim t_{\text{sonic}} \approx 3.8t_{\text{rec}}$, the expansion speed (c_i) of ionized gas is larger than that of the ionization front.

$\frac{dR}{dt} > c_i$ at $t \lesssim t_{\text{sonic}}$	\longrightarrow	$\frac{dR}{dt} \approx c_i$ at $t \approx t_{\text{sonic}}$
initial stage		intermediate stage



Evolutionary scheme of an expanding H II region. (a) The initial stage, (b) expansion with a shock in the neutral gas, (c) the final equilibrium state.

[Figure 7.2 Dyson]

-
- Sound crossing time
 - The ionized region will likely be overpressured relative to its surroundings, in which case it will expand on the sound crossing time.
 - The isothermal sound speed in fully ionized hydrogen is

$$c_s = (2kT/m_{\text{H}})^{1/2} = 13(T/10^4 \text{ K})^{1/2} \text{ km s}^{-1} \quad p = (n_{\text{HI}} + n_e)kT = 2n_{\text{H}}kT$$

- The time for a pressure wave to propagate a distance equal to Strömgren radius is

$$t_{\text{sound}} = \frac{R_s}{c_s} \approx 2.39 \times 10^5 \frac{Q_0/10^{49} \text{ s}^{-1}}{(n/10^2 \text{ cm}^{-3})^{2/3}} \text{ [yr]}$$

- This is about a hundred times longer than the recombination time (timescale of the expanding ionization front).

Introduction to Gas Dynamics

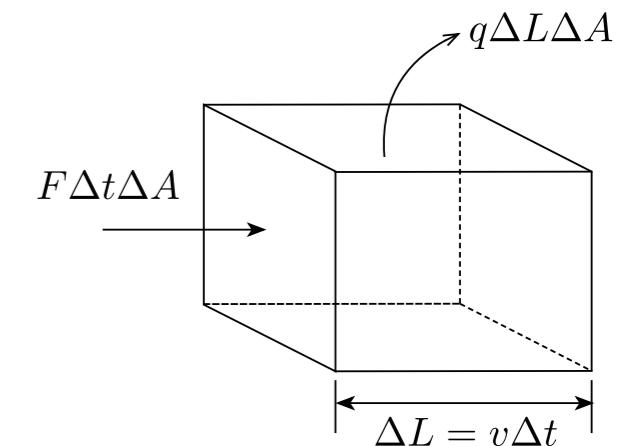
- Assumption for hydrodynamics:
 - particle mean free path << size of the region
 - We will describe the equations for conservation of mass, momentum and energy, in 1D space.

- ***Definition***

- Flux of a hydrodynamic quantity q (for instance, density):

Fluid moves a distance ΔL during a time interval Δt with a velocity v .

$$F\Delta t\Delta A = q\Delta L\Delta A \rightarrow F = qv$$



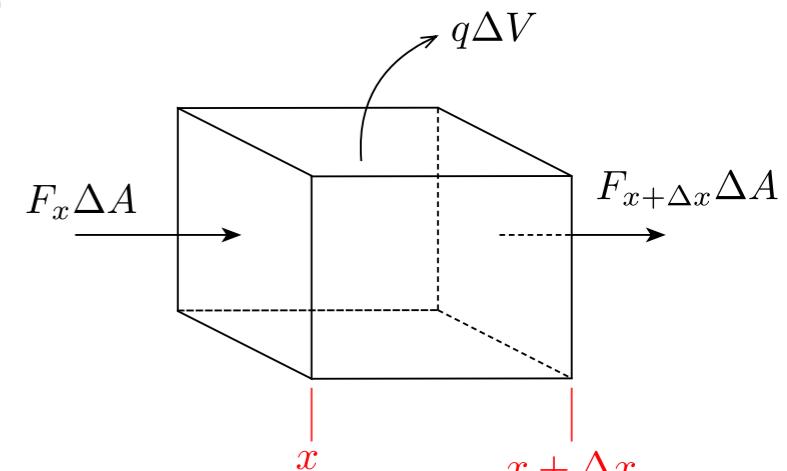
- ***Conservation equation for a quantity*** q

- change of the quantity within a volume ΔV for a time interval Δt :

Here, Δt and Δx are independent.

$$\frac{q\Delta V|_{t+\Delta t} - q\Delta V|_t}{\Delta t} = F\Delta A|_x - F\Delta A|_{x+\Delta x}$$

$$\frac{\partial q}{\partial t} = -\frac{\partial F}{\partial x} \rightarrow \frac{\partial q}{\partial t} = -\frac{\partial(qv)}{\partial x}$$



- Here, no sources or sinks of the quantity within ΔV were assumed. If any, the loss and gain terms should be added in the right-hand side.

Mass Conservation

- Conservation equations
 - ***Mass conservation (continuity equation)***
 - ▶ mass within a volume $dV = \rho dV$
 - ▶ no sources or sinks of material within dV
 - ▶ Consider the mass per unit area (dA), contained in the volume

$$\rho dV/dA = \rho dx \quad \longrightarrow \quad \frac{\partial}{\partial t}(\rho dx) = \overbrace{\rho u}^{\text{incoming}} - \overbrace{(\rho + d\rho)(u + du)}^{\text{outgoing}}$$

$$= -(\rho du + ud\rho + \cancel{d\rho du})$$

$$\frac{\partial \rho}{\partial t} = -\frac{\partial(\rho u)}{\partial x}$$

- ▶ Mass loss and gain terms should be added in the right-hand side, if necessary.

Momentum Conservation

- **Momentum conservation (Euler's equation)**

- ▶ momentum within dV (per unit area) = $(\rho dV)u/dA = \rho dx u$
- ▶ = change of momentum due to fluid flow and gas pressure acting on the surface of dV

$$\begin{aligned}\frac{\partial}{\partial t}(\rho u dx) &= \overbrace{\rho u^2}^{\text{incoming}} - \overbrace{(\rho + d\rho)(u + du)^2}^{\text{outgoing}} + \overbrace{P}^{\text{incoming}} - \overbrace{P + dP}^{\text{outgoing}} \\ &= \rho u^2 - \left(\rho u^2 + 2\rho u du + \cancel{\rho du^2} + u^2 d\rho + \cancel{2ud\rho du} + \cancel{d\rho du^2} \right) - dP\end{aligned}$$

$$\begin{aligned}\frac{\partial}{\partial t}(\rho u) &= -2\rho u \frac{\partial u}{\partial x} - u^2 \frac{\partial \rho}{\partial x} - \frac{\partial P}{\partial x} \\ \rho \frac{\partial u}{\partial t} + u \frac{\partial \rho}{\partial t} &= -\rho u \frac{\partial u}{\partial x} - u \left(\rho \frac{\partial u}{\partial x} + u \frac{\partial \rho}{\partial x} \right) - \frac{\partial P}{\partial x}\end{aligned}$$

Using mass conservation, $\frac{\partial u}{\partial t} = -\frac{\partial(\rho u)}{\partial x}$

$$\rho \frac{\partial u}{\partial t} = -\rho u \frac{\partial u}{\partial x} - \frac{\partial P}{\partial x}$$

or

$$\begin{aligned}\frac{\partial}{\partial t}(\rho u) &= -2\rho u \frac{\partial u}{\partial x} - u^2 \frac{\partial \rho}{\partial x} - \frac{\partial P}{\partial x} \\ &= -\frac{\partial}{\partial x}(\rho u^2) - \frac{\partial P}{\partial x}\end{aligned}$$

$$\frac{\partial}{\partial t}(\rho u) = -\frac{\partial}{\partial x} (\rho u^2 + P)$$

- ▶ Further terms could be added in the right-hand side, accounting for forces due to gravity, magnetic fields, radiation field, and viscosity.

- ▶ The following quantity is sometimes known as **Bernoulli's constant**.

$$\rho u^2 + P$$

One may use it to understand why, for example, fast winds engulfing a house causes it to **explode**, rather than **implode**, because the pressure external to the house becomes lower than its value inside it.

- ▶ Viscous force is due to “internal friction” in the fluid (resistivity of the fluid to the flow), as two adjacent fluid parcels move relative to each other.)

$$\text{viscous force} \propto \frac{\partial^2 u}{\partial x^2}$$

The viscous force is usually much smaller than force due to gas pressure, but important in high-speed flows with large velocity gradients, as in accretion disks.

Homework (due date: 05/15)

[Q12]

- The observed spectrum of an HII region has

$$\frac{I([\text{O III}]4364.4 \text{\AA})}{I([\text{O III}]5008.2 \text{\AA})} = 0.003 ,$$

$$\frac{I([\text{O II}]3729.8 \text{\AA})}{I([\text{O II}]3727.1 \text{\AA})} = 1.2 .$$

- If interstellar reddening is assumed to be negligible, estimate the electron temperature T and the electron density n_e .
- Now suppose that it is learned that there is reddening due to intervening dust with

$$A(4364.4\text{\AA}) - A(5008.2\text{\AA}) = 0.31 \text{ mag}$$

Re-estimate T and n_e . You may find it convenient to use the following equations.

$$\left. \frac{F_{\lambda_2}}{F_{\lambda_1}} \right|_{\text{observed}} = \left. \frac{F_{\lambda_2}}{F_{\lambda_1}} \right|_{\text{intrinsic}} \exp [-(\tau_{\lambda_2} - \tau_{\lambda_1})] \quad \frac{A_{\lambda}}{\text{mag}} = 1.086 \tau_{\lambda}$$