

Interstellar Medium (ISM)

Lecture 13

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Waves in a medium

In order to calculate scattering and absorption of electromagnetic waves by dust grains, we need to characterize the response of the target material to the local oscillating electric fields.

D electric displacement

Gauss

Macroscopic
Maxwell's equations

$$\nabla \cdot \mathbf{D} = 4\pi\rho$$

$$\nabla \cdot \mathbf{B} = 0$$

E electric field

B magnetic flux density
(magnetic induction)

H magnetic field strength
(magnetic field)

Faraday

$$\nabla \times \mathbf{E} = -\frac{1}{c} \frac{\partial \mathbf{B}}{\partial t}$$

Ampere

$$\nabla \times \mathbf{H} = \frac{4\pi}{c} \mathbf{J} + \frac{1}{c} \frac{\partial \mathbf{D}}{\partial t}$$

Constitutive
Relations

$$\mathbf{D} = \epsilon \mathbf{E}$$

$$\mathbf{B} = \mu \mathbf{H}$$

$$\mathbf{J} = \sigma \mathbf{E}$$

Continuity Equation

$$\frac{\partial \rho}{\partial t} + \nabla \cdot \mathbf{J} = 0$$

$$\nabla \cdot \left(\nabla \times \mathbf{H} = \frac{4\pi}{c} \mathbf{J} + \frac{1}{c} \frac{\partial \mathbf{D}}{\partial t} \right)$$

$$0 = 4\pi \nabla \cdot \mathbf{J} + \frac{\partial}{\partial t} \nabla \cdot \mathbf{D}$$

$$0 = \nabla \cdot \mathbf{J} + \frac{\partial \rho}{\partial t}$$

Assume a space and time variation of all quantities of the form $\exp [i(\mathbf{k} \cdot \mathbf{r} - \omega t)]$

\mathbf{k} = wave vector, ω = angular frequency

$$i\mathbf{k} \cdot \mathbf{D} = 4\pi\rho$$

$$i\mathbf{k} \cdot \mathbf{B} = 0$$

$$i\mathbf{k} \times \mathbf{E} = i\frac{\omega}{c} \mathbf{B}$$

$$i\mathbf{k} \times \mathbf{H} = \frac{4\pi}{c} \mathbf{J} - i\frac{\omega}{c} \mathbf{D}$$

$$-i\omega\rho + i\mathbf{k} \cdot \mathbf{J} = 0$$

$$\mathbf{k} \times (\mathbf{k} \times \mathbf{E}) = \frac{\omega}{c} \mathbf{k} \times \mathbf{B}$$

$$\mathbf{k}(\mathbf{k} \cdot \mathbf{E}) - \mathbf{E}(\mathbf{k} \cdot \mathbf{k}) = \frac{\omega\mu}{c} \mathbf{k} \times \mathbf{H}$$

$$k^2 \mathbf{E} = \frac{\omega^2 \epsilon \mu}{c^2} \left(1 + i \frac{4\pi\sigma}{\omega\epsilon} \right) \mathbf{E}$$

$$\mathbf{k} \times \mathbf{H} = -i\frac{4\pi}{c} \mathbf{J} - \frac{\omega}{c} \mathbf{D}$$

$$= -\frac{\omega\epsilon}{c} \left(1 + i \frac{4\pi\sigma}{\omega\epsilon} \right) \mathbf{E}$$

ϵ = permittivity

μ = permeability

σ = conductivity

(유전율)

(투자율)

(전도율)

Dispersion Relation

$$k^2 = \frac{\omega^2}{c^2} m^2$$

$$m^2 = \mu \left(\epsilon + i \frac{4\pi\sigma}{\omega} \right)$$

m = complex index of refraction

m is sometimes referred to as the optical constants, or simply the “n and k”.

- ▶ k is property of the wave, however, $\epsilon\mu$ is a property of the medium. Here, we will ignore the magnetic field, i.e., $\mu = 1$.
- ▶ These enter into the theory through the **complex index of refraction**, $m = n_r + in_i$, where the real and imaginary part are functions of the wavelength.
- ▶ Alternatively, the optical properties of a material can be expressed in terms of the **dielectric function** (or dielectric constant) $\epsilon = \epsilon_1 + i\epsilon_2$. The dielectric function and the complex index of refraction are related through.

$$m = n_r + in_i \quad (\text{or } m = n + ik)$$

$$\begin{aligned} \epsilon &= \epsilon_1 + i\epsilon_2 & \longrightarrow & \epsilon_1 = n_r^2 - n_i^2 \\ \epsilon &= m^2 & \epsilon_2 &= 2n_r n_i \end{aligned}$$

The electrical conductivity σ , if any, can be absorbed within the imaginary part of the dielectric function.

$$\mathbf{J} = \sigma \mathbf{E} \quad \epsilon \rightarrow \epsilon + \frac{4\pi i\sigma}{\omega}$$

- ▶ n (n_r refractive index) and k (n_i extinction coefficient) are often referred to as **optical constants**.
- Consider a plane wave traveling in the z direction represented by

$$E = E_0 \exp [i(kz - \omega t)]$$

- ▶ In free space, the wave vector is given by

$$k = \omega/c = 2\pi/\lambda \quad (\lambda = \text{wavelength in vacuum})$$

-
- ▶ In a material with the index of refraction m , the wave vector is:

$$k = m\omega/c$$

The electric field becomes:

$$E = E_0 \exp\left(-\frac{n_i\omega}{c}z\right) \exp\left[-i\omega\left(t - \frac{n_r z}{c}\right)\right]$$

Thus, ***the real part of the index of refraction introduces a phase shift*** while the ***imaginary part results in damping***. The power of electromagnetic wave will decrease as it propagates through the material, with

$$|E|^2 \propto e^{-2n_i\omega z/c}$$

The attenuation coefficient will be

$$\kappa = 2n_i \frac{\omega}{c} = \frac{4\pi n_i}{\lambda}$$

- Examples:

- ▶ For **transparent substances**, the imaginary part of the index of refraction is much smaller than one.

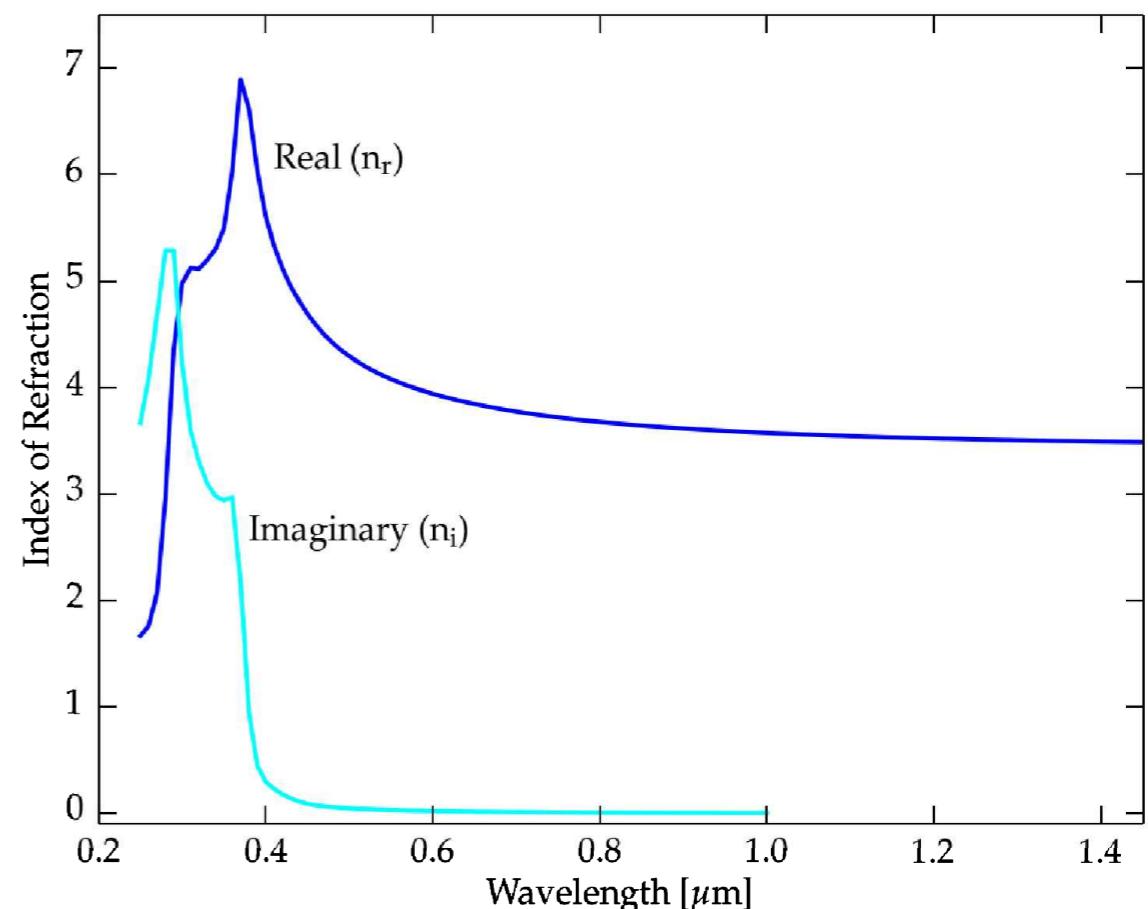
$$m = 1.31 + i(3.1 \times 10^{-9}) \quad \text{pure water ice, } \lambda = 5500\text{\AA}$$

- ▶ For highly **reflective substances**, the imaginary part of the index of refraction is comparable to or greater than one.

$$m = 0.36 + i2.69 \quad \text{gold, } \lambda = 5500\text{\AA}$$

- ▶ The index of refraction can be strongly dependent on wavelength.
- ▶ Silicon goes from being opaque in the UV to being transparent in the near IR.

The real (blue) and imaginary (cyan) components of the index of refraction for silicon at $T = 300$ K.



Mie Theory

- The derivation of the equations is somewhat elaborate.
 - See Chapter 4 of Bohren & Huffman [Absorption and Scattering of Light by Small Particles]
- Summary of the Results:
 - The interaction of an incident wave with a sphere of radius a causes the sphere to radiate electromagnetic waves. This outgoing wave can be written in terms of vector spherical harmonics. Like with spherical harmonics, this involves Legendre polynomials and Bessel functions.
 - The ***extinction and scattering cross sections*** can be written in terms of the scattering coefficients a_n and b_n :

$$Q_{\text{ext}} = \frac{2}{x^2} \sum_{n=1}^{\infty} (2n+1) \operatorname{Re}\{a_n + b_n\}$$

$$Q_{\text{sca}} = \frac{2}{x^2} \sum_{n=1}^{\infty} (2n+1) (|a_n|^2 + |b_n|^2)$$

Here, x is the **size parameter**:

$$x = \frac{2\pi a}{\lambda} = \text{ratio of the size of the particle over the wavelength}$$

- The ***asymmetry factor*** is given by:

$$g = \frac{4}{x^2 Q_{\text{sca}}} \sum_{n=1}^{\infty} \left[\frac{n(n+2)}{n+1} \operatorname{Re}\{a_n^* a_{n+1} + b_n^* b_{n+1}\} + \frac{2n+1}{n(n+1)} \operatorname{Re}\{a_n^* b_n\} \right]$$

The scattering coefficients are expressed in terms of Riccati-Bessel functions ψ and ξ .

$$a_n = \frac{m\psi_n(mx)\psi'_n(x) - \psi_n(x)\psi'_n(mx)}{m\psi_n(mx)\xi'_n(x) - \xi_n(x)\psi'_n(mx)}$$

Here, m = index of refraction

$$b_n = \frac{\psi_n(mx)\psi'_n(x) - m\psi_n(x)\psi'_n(mx)}{\psi_n(mx)\xi'_n(x) - m\xi_n(x)\psi'_n(mx)}$$

Recurrence relations:

$$\psi_n(x) = x j_n(x)$$

$$\psi'_n(x) = x j_{n-1}(x) - n j_n(x)$$

$$\xi_n(x) = x [j_n(x) + i y_n(x)]$$

$$\xi'_n(x) = x [j_{n-1}(x) + i y_{n-1}(x)] - n [j_n(x) + i y_n(x)]$$

The spherical Bessel functions satisfy the recurrence relation:

$$j_n(x) = -j_{n-2}(x) + \frac{2n-1}{x} j_{n-1}(x)$$

$$j_0(x) = \frac{\sin x}{x}$$

$$y_0(x) = -\frac{\cos x}{x}$$

$$y_n(x) = -y_{n-2}(x) + \frac{2n-1}{x} y_{n-1}(x)$$

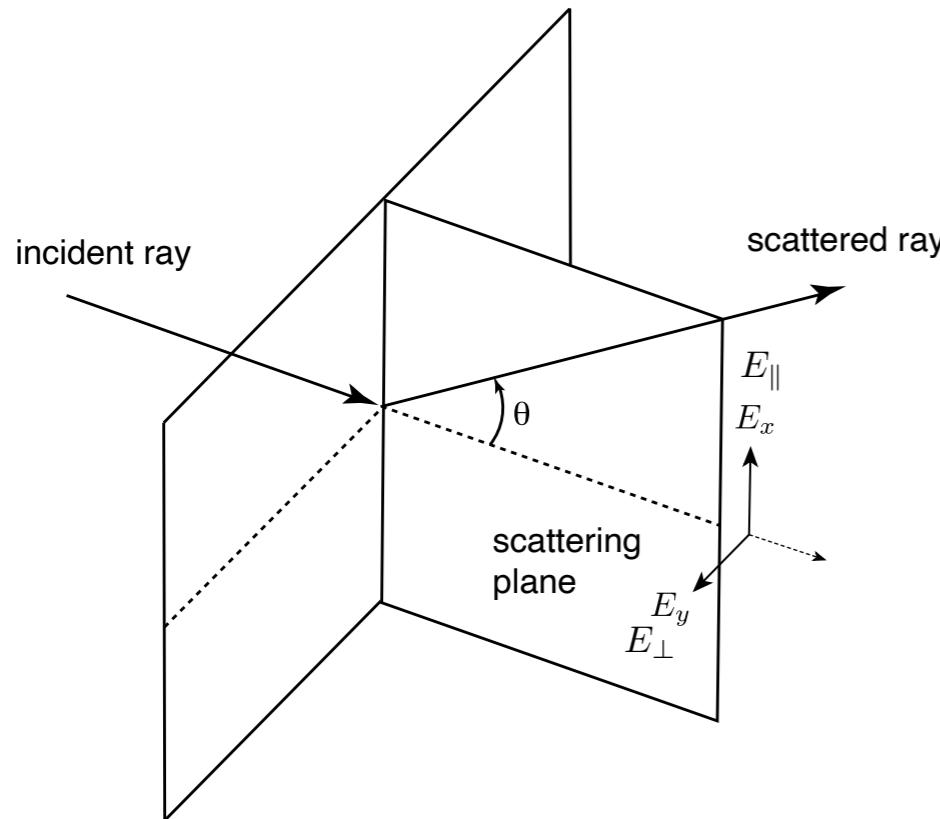
$$j_1(x) = \frac{\sin x}{x^2} - \frac{\cos x}{x}$$

$$y_1(x) = -\frac{\cos x}{x^2} - \frac{\sin x}{x}$$

The larger the particle is compared to the wavelength, the more terms have to be included in the sum. A good Mie code is BHMIE of Bohren & Huffman, a version of which can be downloaded from the website of Bruce Draine (<http://www.astro.princeton.edu/~draine/scattering.html>). A python code: <https://miepython.readthedocs.io/en/latest/>

Scattering Phase Function

Let's define the scattering geometry as follows:



Scattering plane = the plane containing the incident and scattered beams

The scattering angle is defined in the scattering plane.

The functions π_n and τ_n are defined by

$$\pi_n(\cos \theta) = \frac{P_n^1(\cos \theta)}{\sin \theta}$$

$$\tau_n(\cos \theta) = \frac{dP_n^1}{d\theta}$$

Then, the scattered electric fields are given by:

$$\begin{pmatrix} E'_{\parallel} \\ E'_{\perp} \end{pmatrix} = \frac{e^{ik(r-z)}}{-kr} \begin{pmatrix} S_2 & 0 \\ 0 & S_1 \end{pmatrix} \begin{pmatrix} E_{\parallel} \\ E_{\perp} \end{pmatrix}$$

Here, the elements of the amplitude scattering matrix are

$$S_1 = \sum_n \frac{2n-1}{n(n+1)} (a_n \pi_n + b_n \tau_n)$$

$$S_2 = \sum_n \frac{2n-1}{n(n+1)} (a_n \tau_n + b_n \pi_n)$$

Definition of Bohren & Huffman

$E_x = E_{\parallel}^{\text{BH}}$ parallel to the scattering plane

$E_y = -E_{\perp}^{\text{BH}}$ perpendicular to the scattering plane

Recurrence relations:

$$\pi_n(\mu) = \frac{2n-1}{n-1} \mu \pi_{n-1} - \frac{n}{n-1} \pi_{n-2}$$

$$\tau_n(\mu) = n \mu \pi_n - (n+1) \pi_{n-1}$$

$$\pi_0 = 0 \quad \text{and} \quad \pi_1 = 1$$

$$\mu = \cos \theta$$

For an unpolarized incident light ($|E_{\parallel}| = |E_{\perp}|$) , the intensities of the incident and scattered radiation into the direction θ are related by

$$\begin{aligned} I &\equiv |E_{\parallel}|^2 + |E_{\perp}|^2 \\ I' &\equiv |E'_{\parallel}|^2 + |E'_{\perp}|^2 \end{aligned} \quad \longrightarrow \quad I'(\theta) = S_{11}I \quad \text{where } S_{11} = \frac{1}{2} (|S_1|^2 + |S_2|^2)$$

$S_{11}(\cos \theta)$ is the scattering phase function, after a proper normalization.

When integrated over all directions, the S_{11} is related to the scattering efficiency:

$$\int_0^\pi \bar{S}_{11}(\cos \theta) \sin \theta d\theta = \frac{1}{2} x^2 Q_{\text{sca}}$$

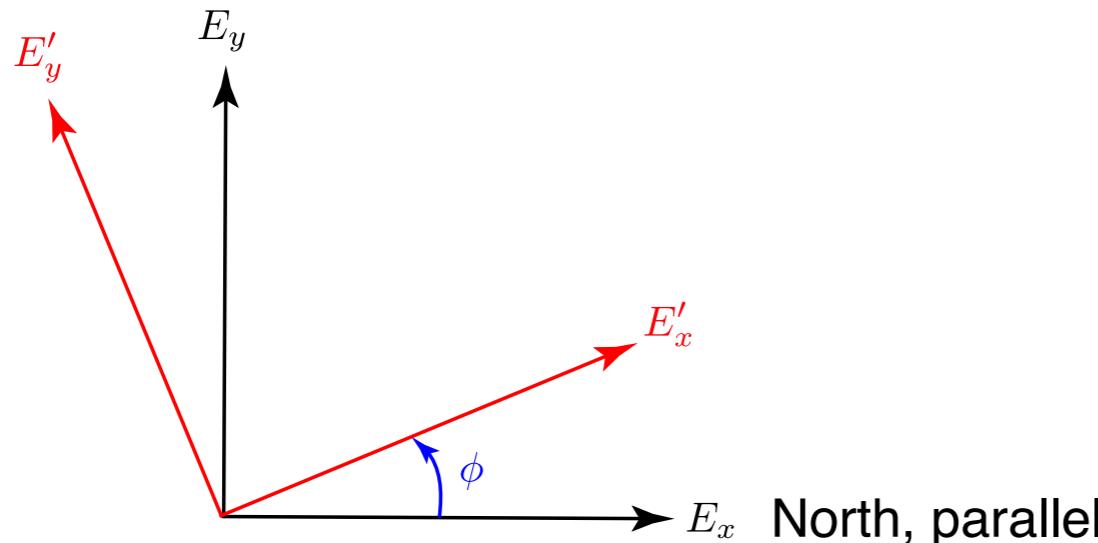
Then, the normalized phase function is given by

$$\mathcal{P}(\cos \theta) = \frac{2}{x^2 Q_{\text{sca}}} \bar{S}_{11}(\cos \theta) = S_{11}(\cos \theta) \quad \int_0^\pi \mathcal{P}(\cos \theta) \sin \theta d\theta = 1$$

Stokes Parameters

Definition of the Stokes parameters, according to the IAU recommendation:

East, perpendicular



$$\begin{aligned} I &= E_x E_x^* + E_y E_y^* \\ Q &= E_x E_x^* - E_y E_y^* \\ U &= E_x E_y^* + E_y E_x^* \\ V &= i(E_x E_y^* - E_y E_x^*) \end{aligned}$$

The Stokes parameters for scattered light for a spherical grain is then given by

$$\begin{pmatrix} I' \\ Q' \\ U' \\ V' \end{pmatrix} = \begin{pmatrix} S_{11} & S_{12} & 0 & 0 \\ S_{12} & S_{11} & 0 & 0 \\ 0 & 0 & S_{33} & S_{34} \\ 0 & 0 & -S_{34} & S_{33} \end{pmatrix} \begin{pmatrix} I \\ Q \\ U \\ V \end{pmatrix}$$

$$S_{11} = \frac{1}{2} (|S_1|^2 + |S_2|^2)$$

$$S_{12} = \frac{1}{2} (|S_2|^2 - |S_1|^2)$$

$$S_{33} = \frac{1}{2} (S_1 S_2^* + S_2 S_1^*)$$

$$S_{34} = \frac{i}{2} (S_1 S_2^* - S_2 S_1^*)$$

$$\begin{pmatrix} I'_{\text{BH}} \\ Q'_{\text{BH}} \\ U'_{\text{BH}} \\ V'_{\text{BH}} \end{pmatrix} = \begin{pmatrix} S_{11}^{\text{BH}} & S_{12}^{\text{BH}} & 0 & 0 \\ S_{12}^{\text{BH}} & S_{11}^{\text{BH}} & 0 & 0 \\ 0 & 0 & S_{33}^{\text{BH}} & S_{34}^{\text{BH}} \\ 0 & 0 & -S_{34}^{\text{BH}} & S_{33}^{\text{BH}} \end{pmatrix} \begin{pmatrix} I_{\text{BH}} \\ Q_{\text{BH}} \\ U_{\text{BH}} \\ V_{\text{BH}} \end{pmatrix}$$

Comparison with the definition of Bohren & Huffman

$$I = I_{\text{BH}} \qquad \qquad S_{11} = S_{11}^{\text{BH}}$$

$$Q = Q_{\text{BH}} \qquad \qquad S_{12} = S_{12}^{\text{BH}}$$

$$U = -U_{\text{BH}} \qquad \qquad S_{33} = S_{33}^{\text{BH}}$$

$$V = -V_{\text{BH}} \qquad \qquad S_{34} = S_{34}^{\text{BH}}$$

Extinction Efficiency

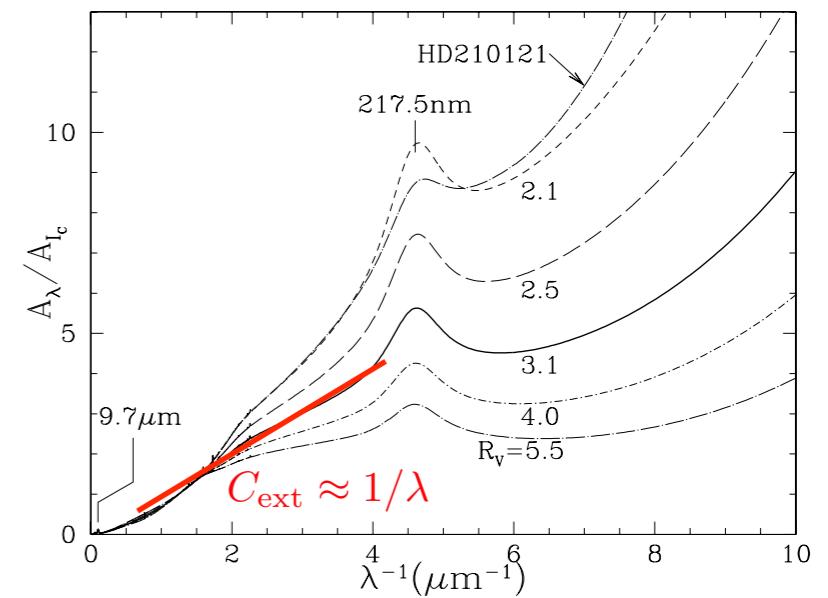
- **Rayleigh Limit:**

$$x = 2\pi a/\lambda \ll 1, \quad |m|x \ll 1$$

- Using only the lowest term of x , we obtain

$$\begin{aligned} a_1 &= -i \frac{2x^3}{3} \frac{m^2 - 1}{m^2 + 2} & Q_{\text{ext}} &\simeq \frac{6}{x^2} \operatorname{Re}(a_1) \\ b_1 &= -i \frac{x^5}{45} (m^2 - 1) & Q_{\text{sca}} &\simeq \frac{6}{x^2} |a_1|^2 \end{aligned}$$

$$\begin{aligned} Q_{\text{ext}} &\simeq 4 \left(\frac{2\pi a}{\lambda} \right) \operatorname{Im} \left(\frac{m^2 - 1}{m^2 + 2} \right) \simeq Q_{\text{abs}} \\ Q_{\text{sca}} &\simeq \frac{8}{3} \left(\frac{2\pi a}{\lambda} \right)^4 \left| \frac{m^2 - 1}{m^2 + 2} \right|^2 \end{aligned}$$



$$\begin{aligned} C_{\text{abs}} &= Q_{\text{abs}} \pi a^2 \propto \frac{a^3}{\lambda} \\ C_{\text{sca}} &\propto \frac{a^6}{\lambda^4} \quad C_{\text{sca}} \ll C_{\text{abs}} \end{aligned}$$

- ▶ The absorption coefficient is proportional to the fraction of the volume of space that is occupied by dust. ***Recall that $C_{\text{ext}} \approx 1/\lambda$ is approximately what is observed in the UV-to-NIR parts.***
- ▶ The scattering cross section is inversely proportional to $1/\lambda^4$, which is characteristic of Rayleigh scattering.
- ▶ ***As long as there is absorption ($n_i \neq 0$), absorption will dominate in the limit of small grain (in the infrared wavelengths).***

- Anomalous diffraction theory
 - When the refractive index of the particle is close to that of its surrounding medium, the particle is sometimes referred to as “soft” or “tenuous”.
 - For optically soft particles and large size parameter,

$$|m - 1| \ll 1, \quad x = 2\pi a/\lambda \gg 1$$

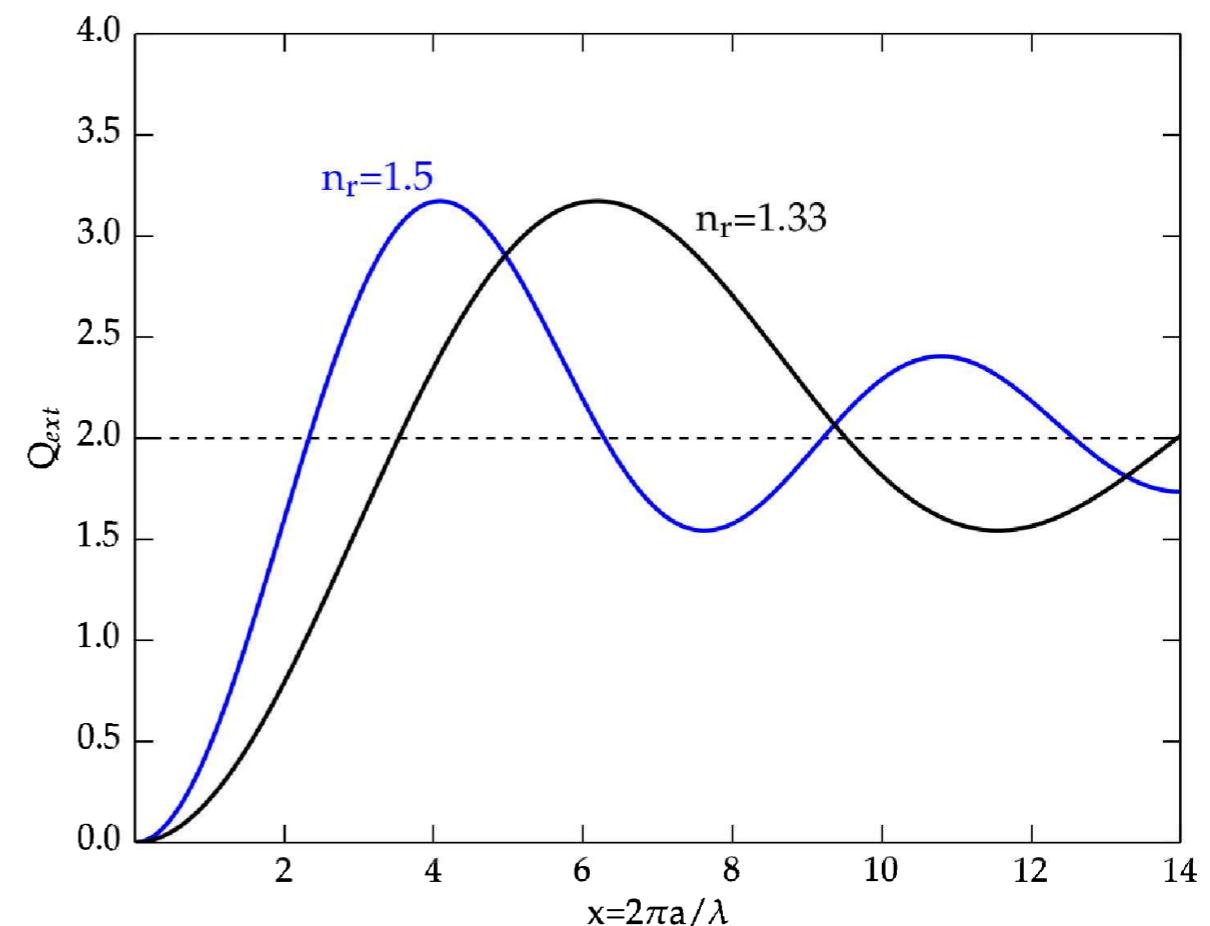
$$|m - 1| \ll 1 \rightarrow n_r \approx 1, \quad n_i \approx 0$$

This condition implies no significant attenuation.

- For the pure scattering case, the efficiency factor is

$$Q_{\text{ext}} = Q_{\text{sca}} = 2 - \frac{4}{\rho} \sin \rho + \frac{4}{\rho^4} (1 - \cos \rho)$$

$$\text{where } \rho \equiv 2 \left(\frac{2\pi a}{\lambda} \right) |m - 1|$$



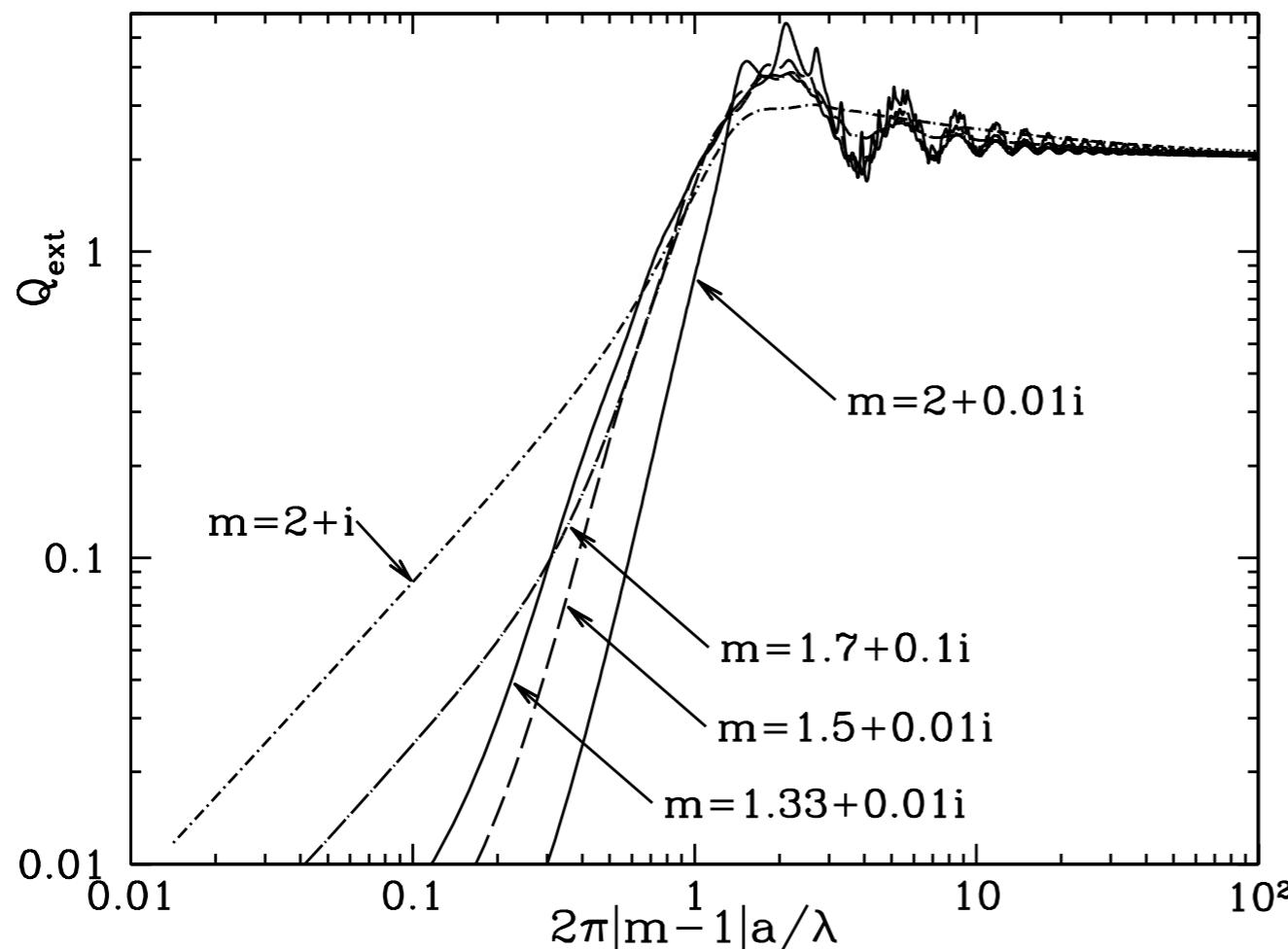
- ***In the short-wavelength limit,***

$$Q_{\text{ext}} \rightarrow 2 \quad (\text{as } |m - 1|x \rightarrow \infty)$$

$$C_{\text{ext}} = (\pi a^2) Q_{\text{ext}} \rightarrow 2(\pi a^2)$$

This is a general result.

- The “**extinction paradox**” (or **Babinet’s theorem**): The extinction cross section approaches twice the geometric cross section. For more realistic spheres, that absorb as well as scatter, it is also found that $Q_{\text{ext}} \rightarrow 2$.



Extinction efficiency factors for spheres with various refractive indices.

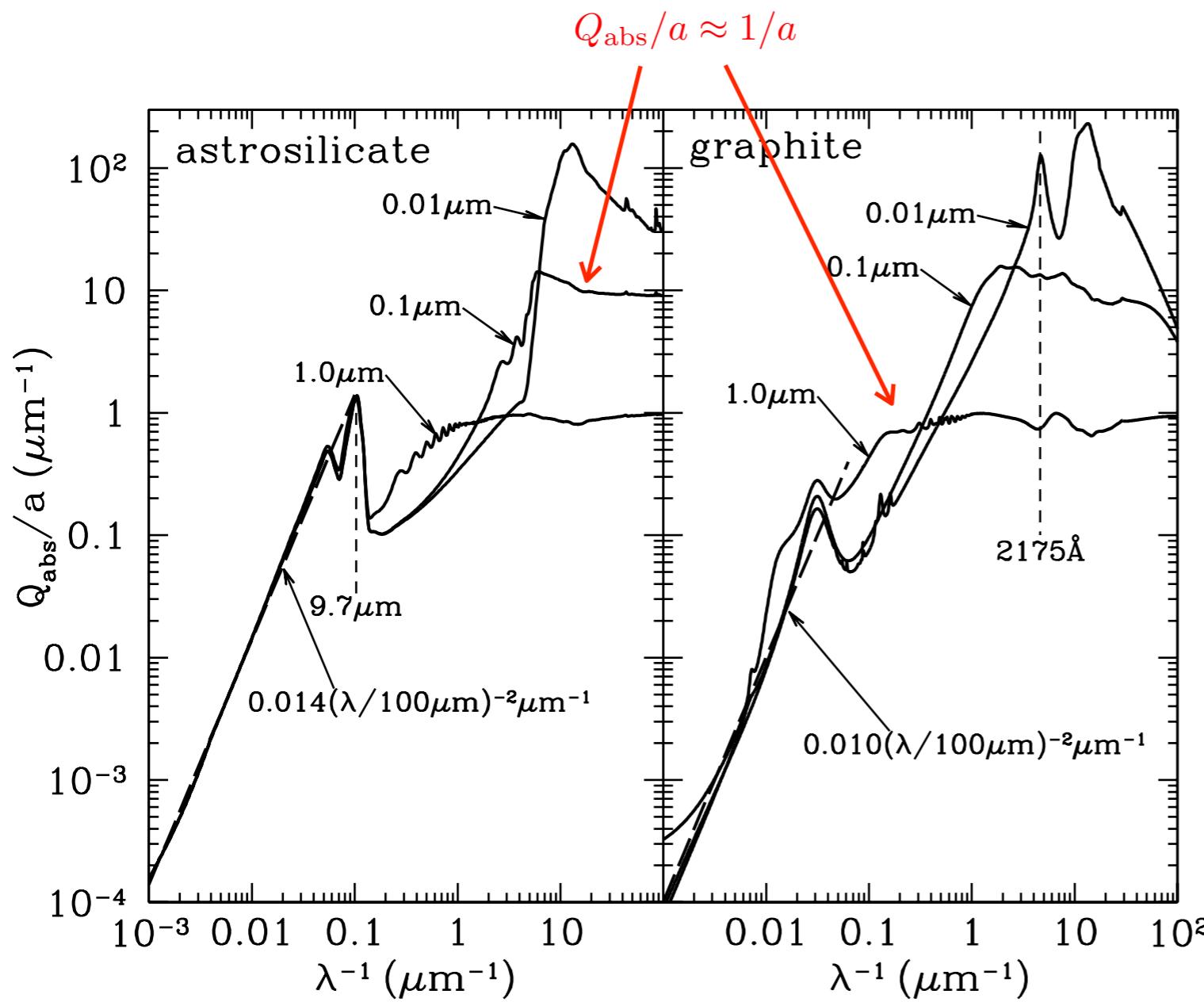
Fig 22.3 [Draine]

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- Ray-tracing arguments would lead us to expect the extinction cross section to be equal to the geometric cross section in the limit $\lambda \ll a$.
 - ▶ However, *diffraction around the target leads to additional small-angle scattering*, with the total extinction cross section equal to twice the geometric cross section.
 - ▶ **Babinet's principle:** The diffraction pattern produced by an opaque object is identical to that from a hole of the same size and shape.
 - ▶ Scattering at the edge of a large obstacle is predominantly forward. At short distances, it will not be possible to recognize the diffraction effect from everyday experience so that a brick appears to remove only as much sunlight as falls onto its projected surface, and not twice as much. Therefore, $Q_{\text{ext}} = 2$ can only be verified at far distances; it is always valid for interstellar grains.
 - ▶ In the limit of opaque bowling balls, light is dimmed with efficiency factors:

$$\begin{aligned} Q_{\text{abs}} &= 1 \\ Q_{\text{sca}} &= 1 \end{aligned}$$

Absorption Efficiencies for Silicate and Graphite

The figure shows the absorption efficiencies for the astrosilicate and graphite grains, with a size of $a \geq 0.01 \mu\text{m}$.



At short wavelengths ($\lambda \lesssim a$):

$$Q_{\text{abs}} \approx 1 \quad (x \gtrsim 2\pi)$$

At long wavelengths ($\lambda \gtrsim 20\mu\text{m}$):

$$Q_{\text{abs}} \approx 0.01(a/\mu\text{m})(\lambda/100\mu\text{m})^{-2}$$

[Fig 24.1, Draine]

Models for Interstellar Dust

- A model for interstellar dust must specify the **composition** of the dust as well as the **geometry (shape and size)** of the dust particles.
 - If the model is to reproduce the polarization of (extinguished) starlight, at least some of the grains should be nonspherical and aligned.
 - From the observational data available to us, **it is not yet possible to arrive at a unique grain model.**
- A class of models that has met with some success assumes the dust to consist of two materials: **(1) amorphous silicate, and (2) carbonaceous material.**
 - ***Mathis, Rumpl, and Nordsieck (1977; MRN)*** found that models using two components, silicate and graphite spheres with power-law size distributions, could reproduce the observed extinction from the near-IR to the UV ($\lambda = 0.11\mu\text{m} - 1\mu\text{m}$).

$$\frac{dn_{\text{gr}}}{da} da = A_i n_{\text{H}} a^{-3.5} da \quad \text{for } a_{\min} \leq a \leq a_{\max}$$

$a_{\min} \approx 0.025\mu\text{m}$

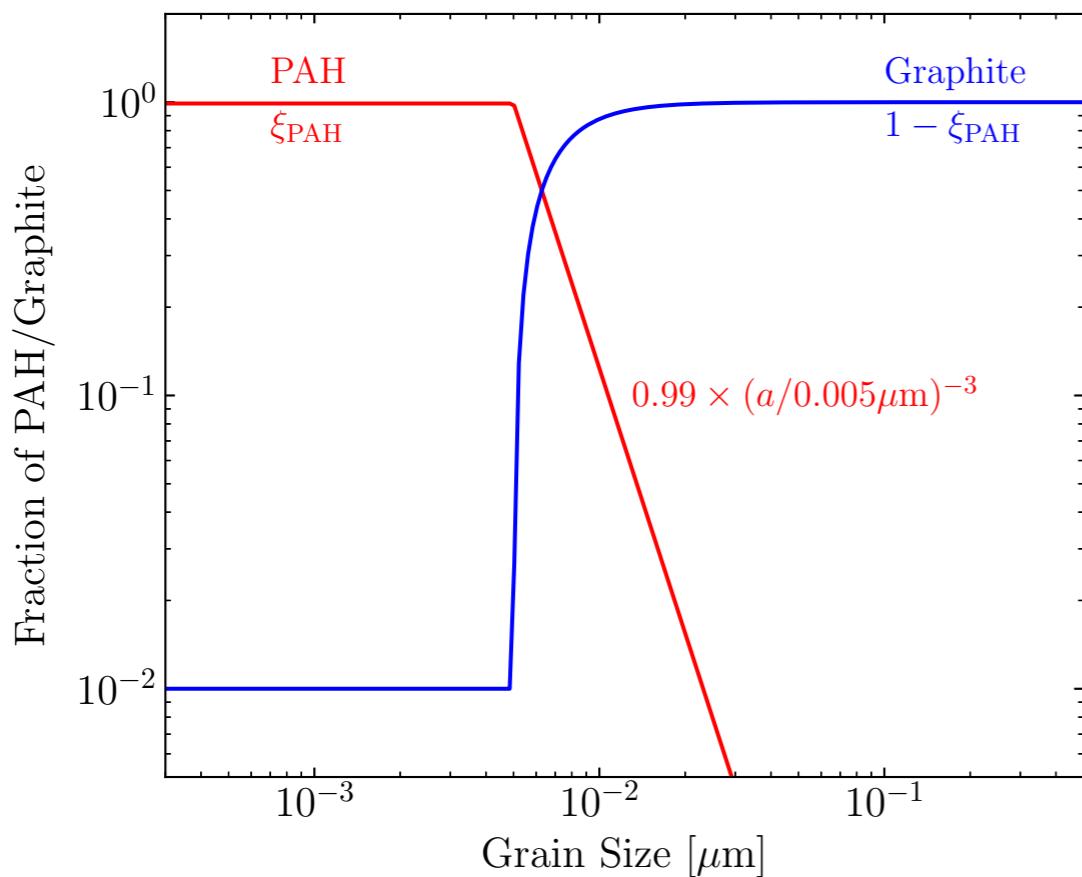
$$A_{\text{sil}} = 7.8 \times 10^{-26}, \quad A_{\text{gra}} = 6.9 \times 10^{-26} \text{ cm}^{2.5} (\text{H atom})^{-1}$$

$a_{\max} \approx 0.25\mu\text{m}$

- ▶ Graphite was a necessary component. The other could be silicon carbide (SiC), magnetite (Fe_3O_4), iron, olivine, or pyroxene.
- ***Draine and Collaborators***
 - ▶ Draine & Lee (1984) presented self-consistent dielectric functions for graphite and silicate, and showed that the graphite-silicate model appeared to be consistent with what was known about dust opacities in the Far-IR. (extended the MRN model to the Far-IR).

- They also included PAHs at small grain sized. PAHs should be added to the graphite-silicate model, either as a third component, or as the small-particle extension of the graphite model.
- The carbonaceous material is assumed to be PAH-like when the particles are small*, but when the particles are large, the carbonaceous material is approximated by graphite.

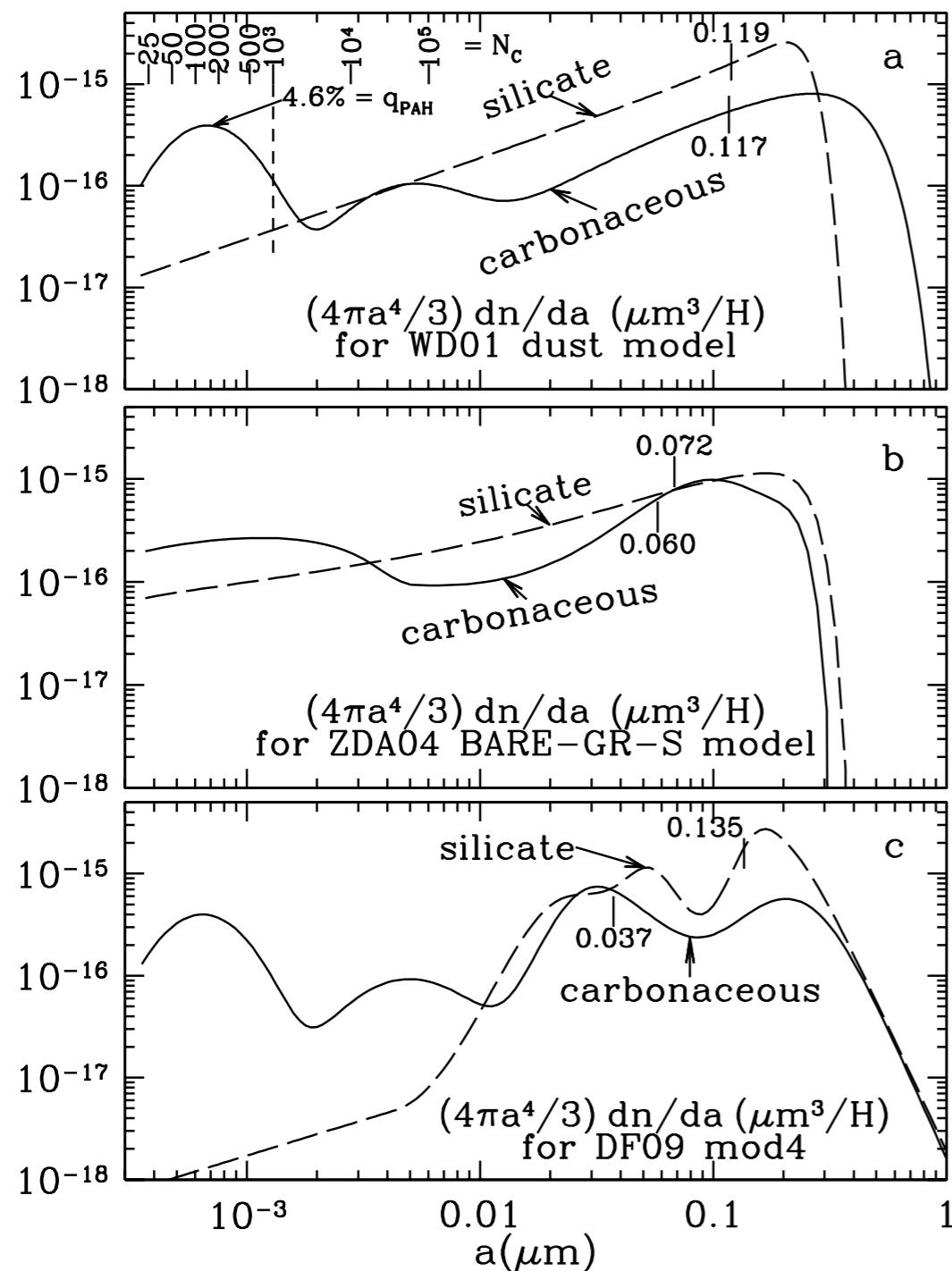
$$\sigma_{\text{abs}}^{\text{carb}}(a, \lambda) = \xi_{\text{PAH}} \sigma_{\text{abs}}^{\text{PAH}} + (1 - \xi_{\text{PAH}}) \sigma_{\text{abs}}^{\text{gra}}(a, \lambda)$$



- The size distributions should reproduce the observed extinction curve, using amounts of grain material that are consistent with the abundance limits.
- The model of **Weingartner & Draine (2001; WD01)** does a good job of reproducing the observed extinction, but the assumed mass in dust exceeds estimates based on elemental abundances and observed depletions.
- WD01 model contains 0.011g of dust per a gram of hydrogen. **In the WD01 model, 75% of the dust mass comes from silicate grains, while other 25% comes from carbon grains.**
- Draine & Fraisse (2009) used grain models with spheroidal graphite and silicate grains to reproduce both the observed extinction and polarization. The non spherical grains are sometimes able to account for the observed extinction using less mass.

- **Zubko et al. (2004)**

- The size distribution of the “BARE-GR-S” model of Zubko et al. (2004), composed of bare graphite grains, bare silicate grains, and PAHs, differs significantly from the WD01 size distribution.



A “typical” grain size may be taken as the half-mass grain size $a_{0.5}$, defined so that half the mass of dust is in grains of radius $a_{0.5}$ or greater.

In WD01 model, both carbonaceous and silicate grains have $a_{0.5} \approx 0.12 \mu\text{m}$.

In Zubko et al. model, $a_{0.5} \approx 0.60 - 0.72 \mu\text{m}$.

Size distributions for silicate and carbonaceous grains for dust models from (a) Weingartner & Draine (2001), (b) Zubko et al. (2004), and (c) Draine & Faisse (2009).

In each case, *tick-marks indicate the “half-mass” radii for the silicate grains and carbonaceous grains.*

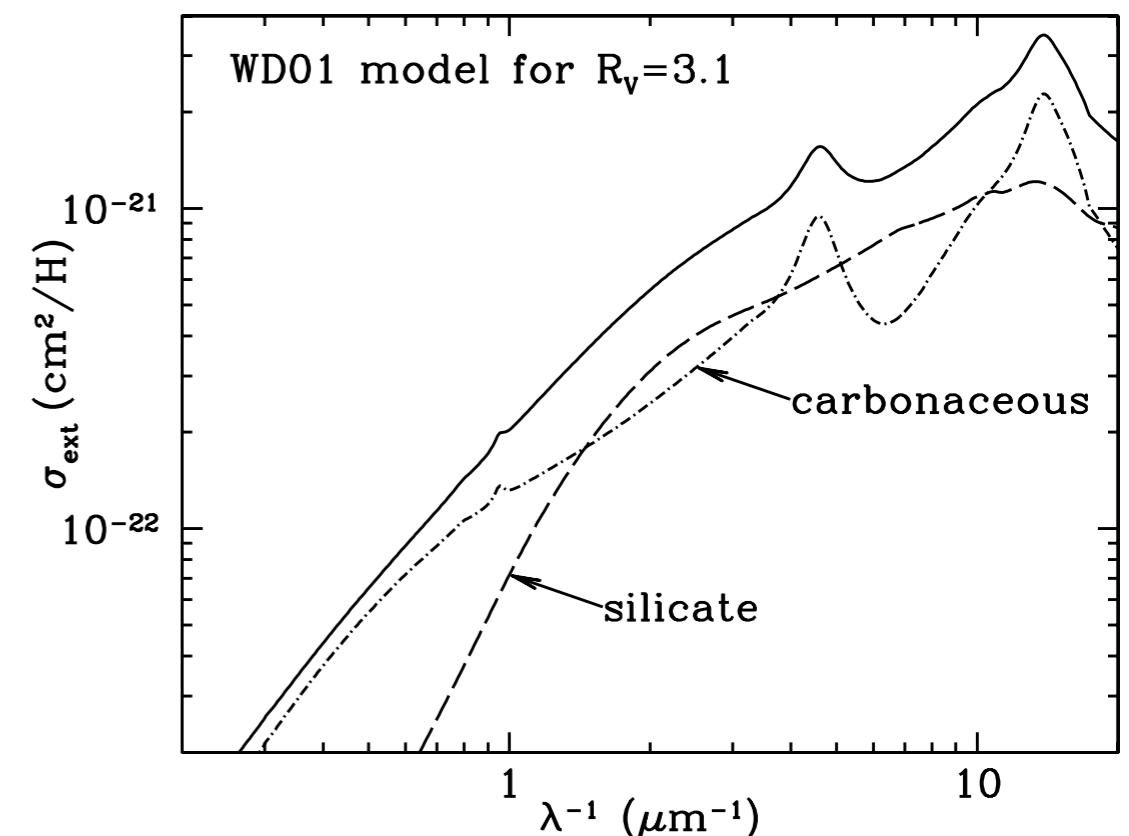
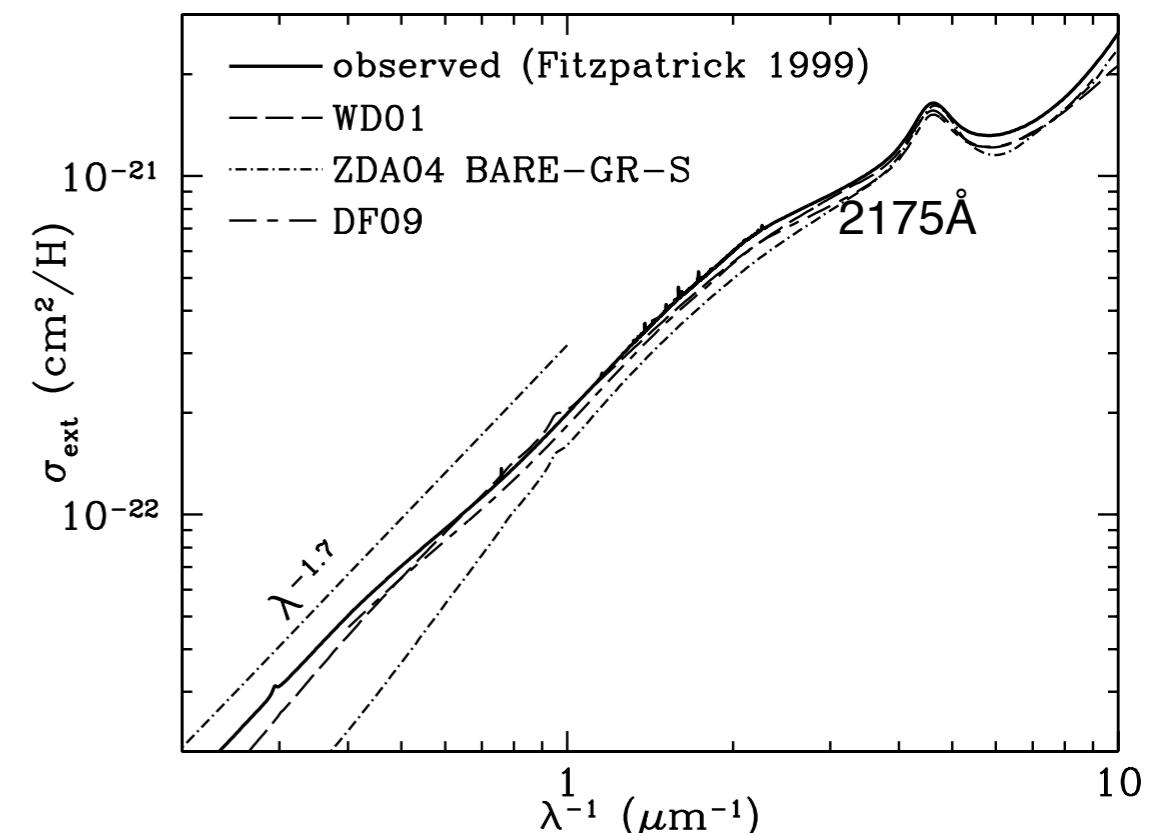
[Fig 23.10 Draine]

[Figure 23.11 Draine]

Averaged observed extinction for $R_V = 3.1$ (Fitzpatrick 1999) and extinction curves calculated for the Weingartner & Draine (2001) model and for the BARE-GR-S model of Zubko et al. (2004).

The Weingartner & Draine model provides considerably more extinction in the infrared ($1-4 \mu\text{m}$) than the Zubko et al. model.

Separate contributions of silicate and carbonaceous grains.



- Constraints on the small dust grains:
 - There is little constraint on the grain size distribution at the small end, because for small grains ($< 200 \text{ \AA}$) extinction is in the Rayleigh limit and independent of grain size. *The presence of very small grains and even large molecules is not derived from extinction measurements but rather inferred from studies of the mid-IR emission of the Galaxy.*
 - For very large grains ($\sim 1\mu\text{m}$), extinction in the visible is gray and dust abundances are mainly derived from abundance constraints on the elements making up these grains.
 - Despite these caveats, all dust models agree that ***the total dust volume is dominated by the large grains while the number density and surface area are dominated by the small grains.***

Dust models

- Li & Draine (2001a): pre-Spitzer model
amorphous silicate grains + graphite + PAHs
spherical grains: no polarization
- Draine & Li (2007) (DL07):
amorph. sil. and graphite from Li & Draine (2001a)
PAHs adjusted slightly to match early Spitzer results
spherical grains: no polarization
- Draine & Fraisse (2009):
DL07 materials
spheroidal grains grains with partial alignment
- Compiègne et al. (2011) (“DUSTEM” model)
amorph. silicate + amorph. C + PAHs
spherical grains: no polarization
- Jones et al. (2013):
amorph. silicate + Fe nanoparticles + amorph. C + PAHs
spherical grains: no polarization
- Hensley & Draine (2015):
amorph. silicate (new dielectric fn.) + Fe + graphite + PAHs
spheroidal grains with partial alignment

Temperatures of Interstellar Grains

- The “temperature” of a dust grain is a measure of ***the internal energy E_{int} present in vibrational modes and possibly also in low-lying electronic excitations.***
- Grain Heating
 - In diffuse regions, where ample starlight is present, grain heating is dominated by absorption of starlight photons.
 - In dense dark clouds, grain heating can be dominated by inelastic collisions with atoms or molecules from the gas (grain-grain collisions are too infrequent).
- When an optical or UV photon is absorbed by a grain, an electron is raised into an excited electronic state; three cases can occur.
 - If the electron is sufficiently energetic, it may be able to escape from the solid as a **“photoelectron.”**
 - In rare cases, the grain will “luminesce” (***luminescence = fluorescence*** when it occurs promptly + **phosphorescence** when it occurs slowly from a metastable level). The excited state will decay radiatively, emitting a photon of energy less than or equal to the energy of the absorbed photon.
 - ***In most solids or large molecules, however, the electronically excited state will deexcite nonradiatively, with the energy going into many vibrational modes - i.e., heat.***

- Large Grains
 - Grains with radii $a \gtrsim 0.03 \mu\text{m}$, can be considered “classical.” These grains are macroscopic - *absorption or emission of single quanta do not appreciably change the total energy in vibrational or electronic excitations.*
 - The temperature of a large dust grain can be obtained by equating the heating rate to the cooling rate.
- Very Small Grains
 - For ultra-small particles, ranging down to large molecules, **quantum effects** are important (this include the “spinning” dust grains responsible for microwave emission).
 - When a dust particle is very small, its temperature will fluctuate. This happens because whenever an energetic photon is absorbed, *the grain temperature jumps up by some not negligible amount* and subsequently declines as a result of cooling.
 - To compute their emission, ***we need their (1) optical and (2) thermal properties.***
 - ▶ The optical behavior depends in a sophisticated way on the the complex index of refraction and on the particle shape.
 - ▶ The thermal behavior is determined more simply from the specific heat.
 - We need to calculate the distribution function of temperature.

Temperature - (1) Heating

- Radiative Heating rate (for a single particle):
 - Ignoring the small fraction of energy appearing as luminescence or photoelectrons, the rate of heating of the grain by absorption of radiation can be written.

$$\begin{aligned} \left(\frac{dE}{dt} \right)_{\text{abs}} &= \int \frac{u_\nu d\nu}{h\nu} \times c \times h\nu \times Q_{\text{abs}}(\nu) \pi a^2 \\ &= \int d\nu 4\pi J_\nu Q_{\text{abs}}(\nu) \pi a^2 \end{aligned}$$

Here, $u_\nu d\nu/h\nu$ is the number density of photons; the photons move at the speed of light c and carry energy $h\nu$.

- It is convenient to define a spectrum-averaged absorption cross section:

$$\langle Q_{\text{abs}} \rangle_* \equiv \frac{\int d\nu u_\nu Q_{\text{abs}}(\nu)}{u_*} = \frac{\int d\nu J_\nu Q_{\text{abs}}(\nu)}{J_*}, \quad u_* \equiv \int d\nu u_\nu, \quad J_* \equiv \int d\nu J_\nu$$

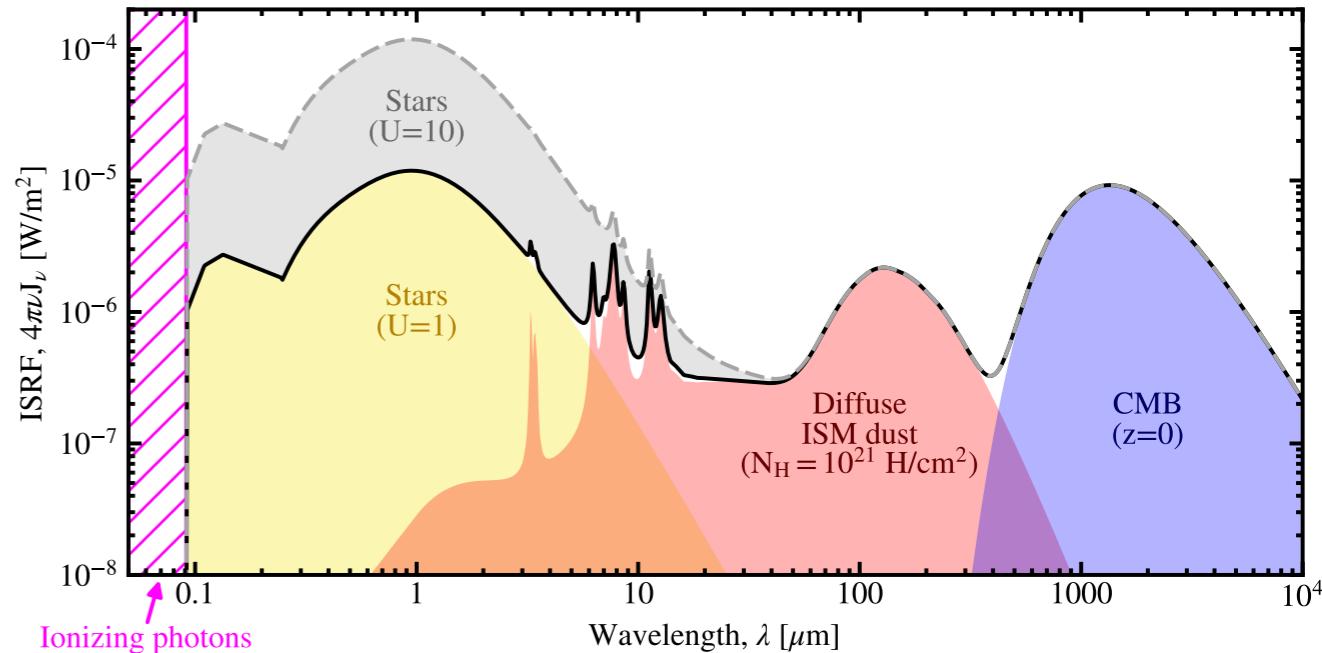
- The radiative heating rate is:

$$\left(\frac{dE}{dt} \right)_{\text{abs}} = \langle Q_{\text{abs}} \rangle_* \pi a^2 (u_* c) = \langle Q_{\text{abs}} \rangle_* \pi a^2 (4\pi J_*)$$

Interstellar Radiation Field (ISRF) of MMP83

- Mathis et al. (1983, MMP83)
 - In the energy range 1 to 13.6 eV, most of the photons are starlight. MMP83 have approximated the local starlight background as a sum of three dilute blackbodies and a piecewise power-law approximation for UV.
 - To improve agreement with the COBE-DIRBE observation, Draine has increased the dilution factor W_1 by 40%.

[Fig I.28, Galliano]



ISRF of the solar neighbor

$$\nu u_\nu = \sum_{j=1}^3 \frac{8\pi h\nu^4}{c^3} \frac{W_j}{e^{h\nu/kT_j} - 1} \quad \text{for } \lambda > 2450\text{\AA}$$

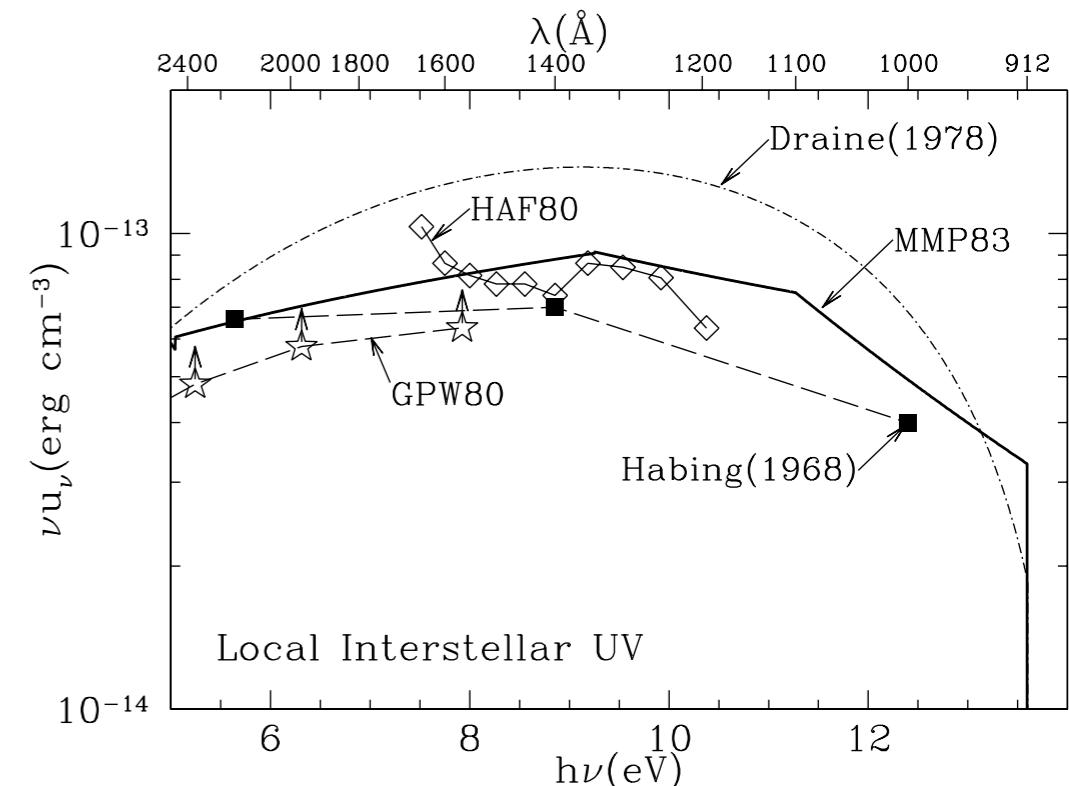
$$T_1 = 3000\text{ K}, \quad W_1 = 7 \times 10^{-13}$$

$$T_2 = 4000\text{ K}, \quad W_2 = 1.65 \times 10^{-13}$$

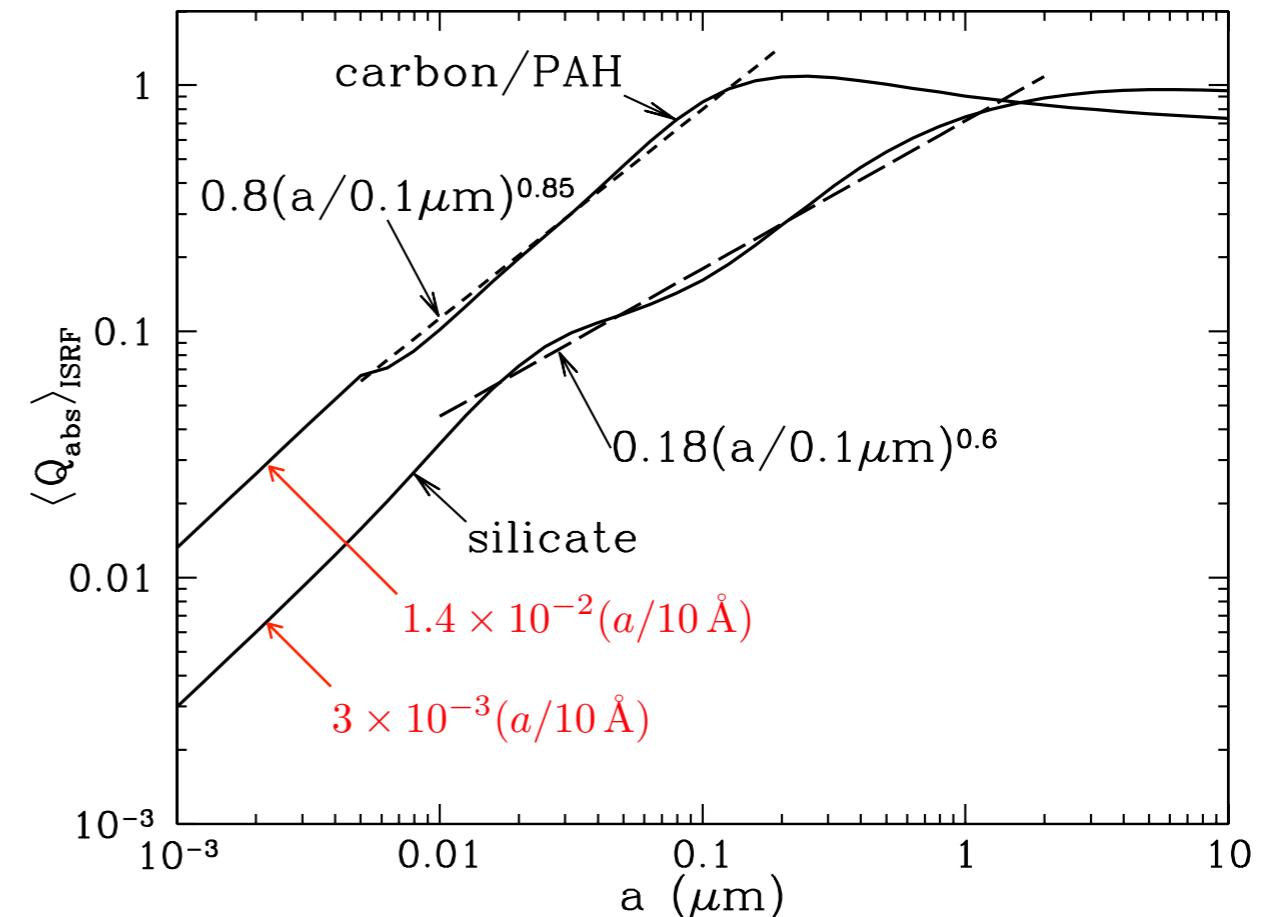
$$T_3 = 7500\text{ K}, \quad W_3 = 1 \times 10^{-14}$$

$$\nu u_\nu = \begin{cases} 2.373 \times 10^{-14} (\lambda/\mu\text{m})^{-0.6678} \text{ erg cm}^{-3} & 1340 - 2460 \text{\AA} \\ 6.825 \times 10^{-13} (\lambda/\mu\text{m}) \text{ erg cm}^{-3} & 1100 - 1340 \text{\AA} \\ 1.287 \times 10^{-9} (\lambda/\mu\text{m})^{4.4172} \text{ erg cm}^{-3} & 912 - 1100 \text{\AA} \end{cases}$$

[Fig 12.2, Draine]



- The spectrum-averaged absorption cross section as a function of radius for graphite and silicate grains, and the spectrum of the interstellar radiation field (ISRF) from Mathis et al. (1983).



[Fig 24.2, Draine]

The numerical results can be approximated by

$$\begin{aligned}\langle Q_{\text{abs}} \rangle_{\text{ISRF}} &\approx 0.18 (a/0.1\mu\text{m})^{0.6}, \quad \text{for silicate, } 0.01 \lesssim a \lesssim 1\mu\text{m} \\ &\approx 0.8 (a/0.1\mu\text{m})^{0.85}, \quad \text{for graphite, } 0.005 \lesssim a \lesssim 0.15\mu\text{m}\end{aligned}$$

$$u_* = 1.05 \times 10^{-12} U \text{ [erg cm}^{-3}\text{]} \quad U = \text{strength of the ISRF in units of that of MMP83.}$$

Temperature - (2) Cooling

- Radiative Cooling rate (for a single particle)

- Kirchhoff's Law in LTE

j_ν = emissivity per unit volume

κ_ν = absorption coefficient per unit length

j_ν/n_d = emissivity per particle

$\kappa_\nu/n_d = C_{\text{abs}}(\nu)$ = absorption cross section

n_d = number density of dust particles

$$\frac{j_\nu}{\kappa_\nu} = B_\nu(T) \Rightarrow \frac{j_\nu}{n_d} = C_{\text{abs}}(\nu)B_\nu(T)$$

[$B_\nu(T)$ = Planck function, $\kappa_\nu = n_d C_{\text{abs}}(\nu)$]

- Grains lose energy by infrared emission at a rate:

$$\begin{aligned} \left(\frac{dE}{dt} \right)_{\text{emiss}} &= \int d\nu 4\pi j_\nu / n \\ &= \int d\nu 4\pi B_\nu(T_d) C_{\text{abs}}(\nu) \\ &= 4\pi a^2 \langle Q_{\text{abs}} \rangle_{T_d} \sigma_{\text{SB}} T_d^4 \end{aligned}$$

The Stephan-Boltzmann constant is:

$$\sigma_{\text{SB}} = \frac{2\pi^5 k_B^4}{15c^2 h^3} = 5.67 \times 10^{-5} \text{ erg cm}^2 \text{ s}^{-1} \text{ K}^{-4} \text{ sr}^{-1}$$

Here, the Planck-averaged emission efficiency is defined by

$$\langle Q_{\text{abs}} \rangle_T \equiv \frac{\int d\nu B_\nu(T) Q_{\text{abs}}(\nu)}{\int d\nu B_\nu(T)}$$

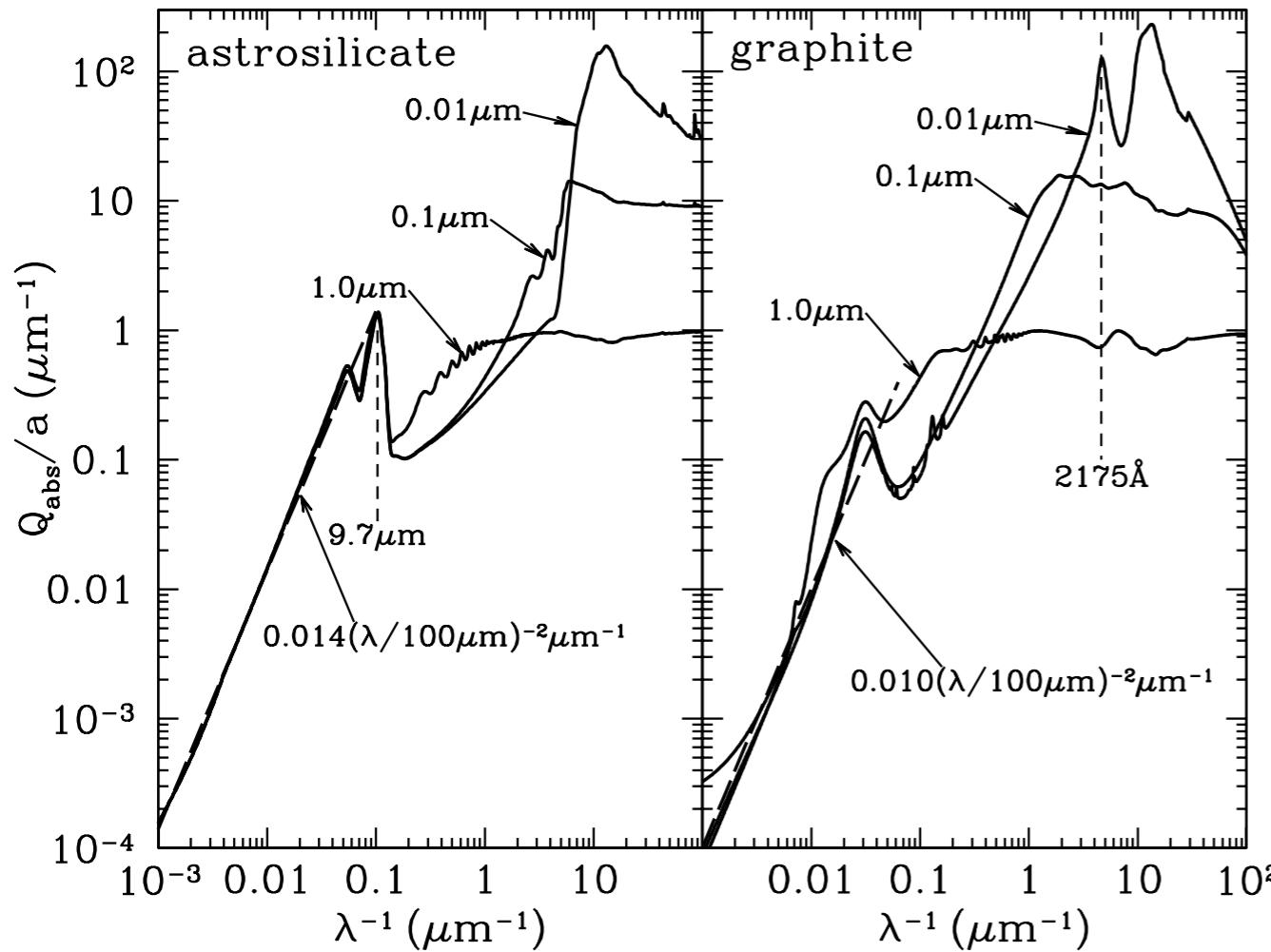
$$\int d\nu B_\nu(T) = \frac{\sigma_{\text{SB}}}{\pi} T_d^4$$

- In general, the absorption cross section in the far-IR can be approximated as a power-law in frequency,

$$Q_{\text{abs}}(\nu) = Q_0(\nu/\nu_0)^\beta = Q_0(\lambda/\lambda_0)^{-\beta} \quad (1 \lesssim \beta \lesssim 2)$$

in the far-IR

then the Planck average can be obtained analytically:



[Fig 24.1, Draine]

$$Q_{\text{abs}} \approx 1.4 \times 10^{-3} \left(\frac{a}{0.1\mu\text{m}} \right) \left(\frac{\lambda}{100\mu\text{m}} \right)^{-2}$$

silicate, $\lambda \gtrsim 20\mu\text{m}$

$$\approx 1.0 \times 10^{-3} \left(\frac{a}{0.1\mu\text{m}} \right) \left(\frac{\lambda}{100\mu\text{m}} \right)^{-2}$$

graphite, $\lambda \gtrsim 30\mu\text{m}$

Using the power-law approximation is valid because we are interested only in the Far-IR.

$$Q_{\text{abs}}(\nu) = Q_0 (\nu/\nu_0)^\beta$$

$$B_\nu(T) = \frac{2h\nu^3/c^2}{\exp(h\nu/k_B T) - 1}$$

$$\int_0^\infty d\nu B_\nu(T) Q_{\text{abs}}(\nu) = \left(\frac{2h}{c^2}\right) \left(\frac{kT}{h}\right)^4 \left(\frac{kT}{h\nu_0}\right)^\beta Q_0 \int_0^\infty dx \frac{x^{3+\beta}}{e^x - 1} \quad \leftarrow \quad x \equiv \frac{h\nu}{kT}$$

$$\begin{aligned} \int_0^\infty dx \frac{x^{3+\beta}}{e^x - 1} &= \int_0^\infty dx x^{3+\beta} e^{-x} (1 - e^{-x}) = \int_0^\infty dx x^{3+\beta} \sum_{n=1}^\infty e^{-nx} \\ &= \sum_{n=1}^\infty \frac{1}{n^{4+\beta}} \int_0^\infty dy y^{3+\beta} e^{-y} \quad \leftarrow \quad y \equiv nx \\ &= \zeta(4 + \beta) \Gamma(4 + \beta) \end{aligned}$$

$$\begin{aligned} \langle Q_{\text{abs}} \rangle_T &\equiv \frac{\int d\nu B_\nu(T) Q_{\text{abs}}(\nu)}{\int d\nu B_\nu(T)} \\ &= \frac{\left(\frac{kT}{h\nu_0}\right)^\beta Q_0 \zeta(4 + \beta) \Gamma(4 + \beta)}{\zeta(4) \Gamma(4)} \end{aligned}$$



$$\begin{aligned} \langle Q_{\text{abs}} \rangle_T &= \frac{15}{\pi^4} \Gamma(4 + \beta) \zeta(4 + \beta) Q_0 \left(\frac{kT}{h\nu_0}\right)^\beta \\ &= \left(\frac{40}{21}\pi^2\right) Q_0 \left(\frac{kT}{h\nu_0}\right)^2 \quad \text{for } \beta = 2 \end{aligned}$$

Riemann zeta-function	gamma function
$\zeta(4) = \sum_{n=1}^\infty \frac{1}{n^4} = \frac{\pi^4}{90}$	$\Gamma(4) = 3! = 6$
	$\Gamma(5) = 4! = 24$
$\zeta(5) = \sum_{n=1}^\infty \frac{1}{n^5} = 1.03692$	$\Gamma(6) = 5! = 120$
$\zeta(6) = \sum_{n=1}^\infty \frac{1}{n^6} = \frac{\pi^6}{945}$	

$$\begin{aligned} Q_0 &\approx 1.4 \times 10^{-3} (a/0.1 \mu\text{m}) \quad \text{silicate} \\ &\approx 1.0 \times 10^{-3} (a/0.1 \mu\text{m}) \quad \text{graphite} \\ \lambda_0 &= 100 \mu\text{m} \end{aligned}$$

$$\begin{aligned} \langle Q_{\text{abs}} \rangle_T &\approx 1.3 \times 10^{-6} (a/0.1 \mu\text{m}) (T/\text{K})^2 \quad (\text{silicate}) \\ &\approx 9.1 \times 10^{-7} (a/0.1 \mu\text{m}) (T/\text{K})^2 \quad (\text{graphite}) \end{aligned}$$

Equilibrium Temperature in the diffuse ISM

- Steady state temperature of large grains
 - The balance equation between the heating and cooling is:

$$\left(\frac{dE}{dt} \right)_{\text{abs}} = \left(\frac{dE}{dt} \right)_{\text{emiss}} \Rightarrow \pi a^2 \langle Q_{\text{abs}} \rangle_* (u_* c) = 4\pi a^2 \langle Q_{\text{abs}} \rangle_{T_d} \sigma_{\text{SB}} T_d^4$$

$$\Rightarrow \langle Q_{\text{abs}} \rangle_* (u_* c) = \frac{60}{\pi^4} \Gamma(4 + \beta) \zeta(4 + \beta) Q_0 \left(\frac{k}{h\nu_0} \right)^\beta \sigma_{\text{SB}} T_d^{4+\beta}$$

$$T_d = \left(\frac{h\nu_0}{k} \right)^{\beta/(4+\beta)} \left[\frac{\pi^4 \langle Q_{\text{abs}} \rangle_* c}{60\zeta(4+\beta)\Gamma(4+\beta)Q_0\sigma_{\text{SB}}} \right]^{1/(4+\beta)} u_*^{1/(4+\beta)}$$

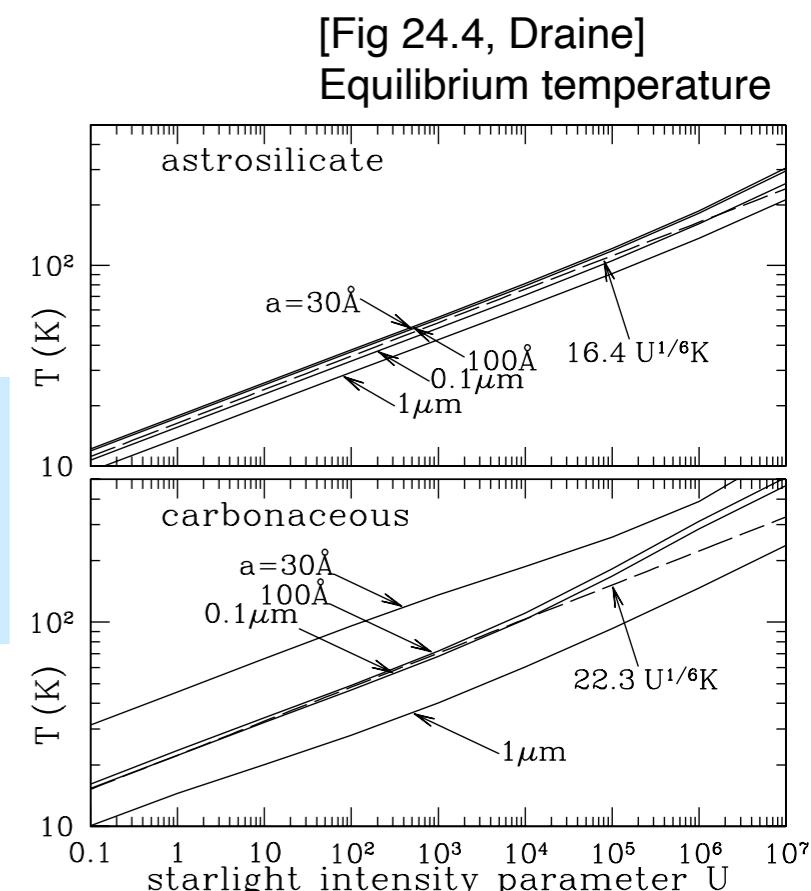
$$= \left(\frac{h\nu_0}{k} \right)^{1/3} \left[\frac{21 \langle Q_{\text{abs}} \rangle_* c}{160\pi^2 Q_0 \sigma_{\text{SB}}} \right]^{1/6} u_*^{1/6} \quad \text{for } \beta = 2$$

- Therefore, the temperature of a large grain is given by:

$$T_d \approx 16.4 (a/0.1 \mu\text{m})^{-1/15} U^{1/6} \text{ K, silicate} \quad (0.01 \lesssim a \lesssim 1 \mu\text{m})$$

$$\approx 22.3 (a/0.1 \mu\text{m})^{-1/40} U^{1/6} \text{ K, graphite} \quad (0.005 \lesssim a \lesssim 0.15 \mu\text{m})$$

for $Q_{\text{abs}} = Q_0(\lambda/\lambda_0)^{-2}$ ($\beta = 2$)



-
- Implications
 - ▶ If the ISRF is doubled, the grain temperature increases by $\sim 12\%$. In order to increase the temperature by a factor of 2, 64 times stronger radiation is required.
 - ▶ ***There is also little dependence of the grain temperature on grain radius. Therefore, large grains can be regarded to be grains with a single size.***
 - ▶ If we assume $\beta = 1$ for the absorption cross section in the FIR, we obtain

$$T_d = \left(\frac{h\nu_0}{k} \right)^{1/5} \left[\frac{\pi^4 \langle Q_{\text{abs}} \rangle_* c}{60\zeta(5)\Gamma(5)Q_0\sigma_{\text{SB}}} \right]^{1/5} u_*^{1/5} \quad \text{for } Q_{\text{abs}} = Q_0(\lambda/\lambda_0)^{-1}$$

$$\begin{aligned} T_d &\approx 14.6 (a/0.1 \mu\text{m})^{-2/25} U^{1/5} \text{ K, silicate} \quad (0.01 \lesssim a \lesssim 1 \mu\text{m}) \\ &\approx 21.1 (a/0.1 \mu\text{m})^{-3/100} U^{1/5} \text{ K, graphite} \quad (0.005 \lesssim a \lesssim 0.15 \mu\text{m}) \end{aligned} \quad \text{for } \beta = 1$$

This indicates that *the steeper absorption cross section ($\beta = 2$) give a higher dust temperature.*

Example: Temperature of Grains exposed to a Star

- **Blackbody Particle (perfect absorber, $Q_{\text{abs}} = 1$)**

- If a blackbody particle is heated by a star with a temperature T_* and bolometric L_* at distance d , the balance between heating and cooling gives the dust temperature:

$$\int \pi a^2 Q_{\text{abs}}(\nu) \frac{L_\nu}{4\pi d^2} d\nu = 4\pi \int \pi a^2 Q_{\text{abs}}(\nu) B_\nu(T_d) d\nu \quad \longrightarrow \quad \frac{L_*}{4\pi d^2} = 4\sigma_{\text{SB}} T_d^4$$

Since the stellar luminosity can be expressed in terms of the stellar effective temperature.

$$L_* = 4\pi R_*^2 \int \pi B_\nu(T_*) d\nu = 4\pi R_*^2 (\sigma_{\text{SB}} T_*^4)$$

Here, R_* is the radius of the star.

We can find the equilibrium temperature of dust grains.

$$T_d = T_* \left(\frac{R_*}{2d} \right)^{1/2}$$

The dust temperature therefore decreases as the inverse square root of distance.

- Calculating the temperature at the distance of 1 AU from the Sun,

$$T_* = 5778 \text{ K}$$

$$R_* = 6.985 \times 10^{10} \text{ cm} = 4.669 \times 10^{-3} \text{ AU}$$

$$T_d = 279 \text{ K} \left(\frac{d}{\text{AU}} \right)^{-1/2}$$

This value is close to the average temperature of the surface of the Earth, although our planet does not emit like a perfect blackbody.

- ***Interstellar dust grains directly exposed to stars***
 - We assume that the absorption efficiency can be approximated as follows:

$$Q_{\text{abs}} \approx 1 \quad \text{at visible/UV wavelengths}$$

$$Q_{\text{abs}} \approx 0.01(a/\mu\text{m})(\lambda/100\mu\text{m})^{-2} \quad \text{in the FIR}$$

$$Q_0 = 0.01(a/\mu\text{m}), \lambda_0 = 100\mu\text{m}$$
 - The energy balance is

$$\int Q_{\text{abs}}(\nu) \frac{L_\nu}{4\pi d^2} d\nu = 4\pi \int Q_{\text{abs}}(\nu) B_\nu(T_d) d\nu$$

$$\frac{L_*}{4\pi d^2} = 4\pi \int Q_{\text{abs}}(\nu) B_\nu(T_d) d\nu \quad \longrightarrow \quad \left(\frac{R_*}{d}\right)^2 \sigma_{\text{SB}} T_*^4 = 4 \langle Q_{\text{abs}} \rangle_{T_d} \sigma_{\text{SB}} T_d^4$$

$$\langle Q_{\text{abs}} \rangle_{T_d} = \left(\frac{40}{21}\pi^2\right) Q_0 \left(\frac{kT_d}{h\nu_0}\right)^2$$

 - Using the Planck-averaged absorption efficiency, we can obtain the dust temperature:

$$T_d = \left(\frac{21}{160\pi^2} \frac{hc}{k_B \lambda_0} \frac{1}{Q_0}\right)^{1/6} \left(\frac{R_*}{d}\right)^{1/3} T_*^{2/3}$$

$$= 2.4 \left(\frac{a}{\mu\text{m}}\right)^{-1/6} \left(\frac{d}{R_*}\right)^{-1/3} T_*^{2/3}$$

$$= 129 \text{ K} \left(\frac{a}{\mu\text{m}}\right)^{-1/6} \left(\frac{d}{\text{AU}}\right)^{-1/3}$$

If there is extinction along the line of sight, one must replace the luminosity L_ν with $L_\nu e^{-\tau_\nu}$.

Stochastic Heating of Very Small Grains

- Temperature History:
 - ▶ Two effects become increasingly important with diminishing grain size: (1) **the heat capacity of the dust becomes sufficiently small that single-particle hits can cause large spikes in the dust temperature** and (2) **the absorption rate with photons becomes sufficiently low that the cooling of the dust between successive collisions becomes important.**
 - ▶ ***Therefore, it is clear that one cannot speak of a representative grain temperature under these conditions - one must instead use a temperature distribution function.***
 - ▶ ***As the grain size is increased, however, photon absorption events occur more frequently, the temperature rise at each event is reduced by the increased heat capacity, and temperature varies over only a small range.***
- Energy content:
 - Heat capacity = heat energy required to raise the temperature of a material
 - When a grain is at a temperature T , the vibrational (internal) energy content of the grain (per unit volume) is

$$\frac{U(T)}{V} = \int_0^T C(T) dT$$

where $C(T)$ is the heat capacity of the grain (per unit volume) at temperature T .

- Small grains are subject to temporal fluctuations in temperature, because their steady state internal energies are small compared with the energies of absorbed UV photons.

$U_{\text{steady}}(T) \ll h\nu$	$VC\Delta T = h\nu : \Delta T \uparrow \text{ as } VC \downarrow \text{ for a given absorbed energy } h\nu$
---------------------------------	---

(1) Peak Temperature

- Peak Temperature

- **Heat Capacity:** We need to know the **heat capacity** for dust grains to estimate the peak temperature.

In the low temperature limit: For a plausible grain material, the expected grain temperature would be in general lower than the Debye temperature; we may then use the *Debye approximation* for the heat capacity:

$$C(T) \approx \frac{N_a}{V} k_B \frac{12\pi^4}{5} \left(\frac{T}{\theta_D} \right)^3$$

$V = (4\pi/3)a^3$ = volume of a grain particle

N_a = number of atoms forming the grain particle

$n_a = N_a/V$ = number density of atoms in the grain.

$$\left(n_a = \frac{\rho}{\mu m_H}, \quad N_a = \frac{\rho}{\mu m_H} \frac{4\pi}{3} a^3 \right)$$

Here, **Debye temperature θ_D** ~ temperature at which the highest-frequency mode is excited.

For graphite,

$$\mu = 12, \quad \rho = 2.24 \text{ g cm}^{-3}, \quad \theta_D = 420 \text{ K}$$

For silicate (olivine),

$$\mu = 170, \quad \rho = 3.8 \text{ g cm}^{-3}, \quad \theta_D = 665 \text{ K}$$

$$C \Rightarrow \mu = 6 \times 2$$

$$\text{Mg}_2\text{SiO}_4 \Rightarrow \mu = 12 \times 2 \times 2 + 14 \times 2 + 8 \times 2 \times 4$$

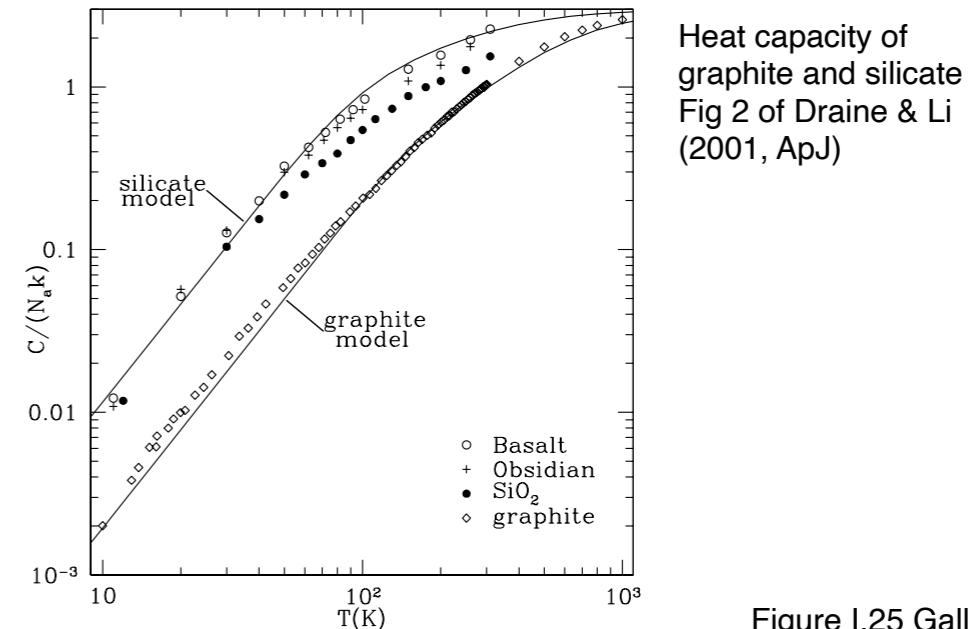
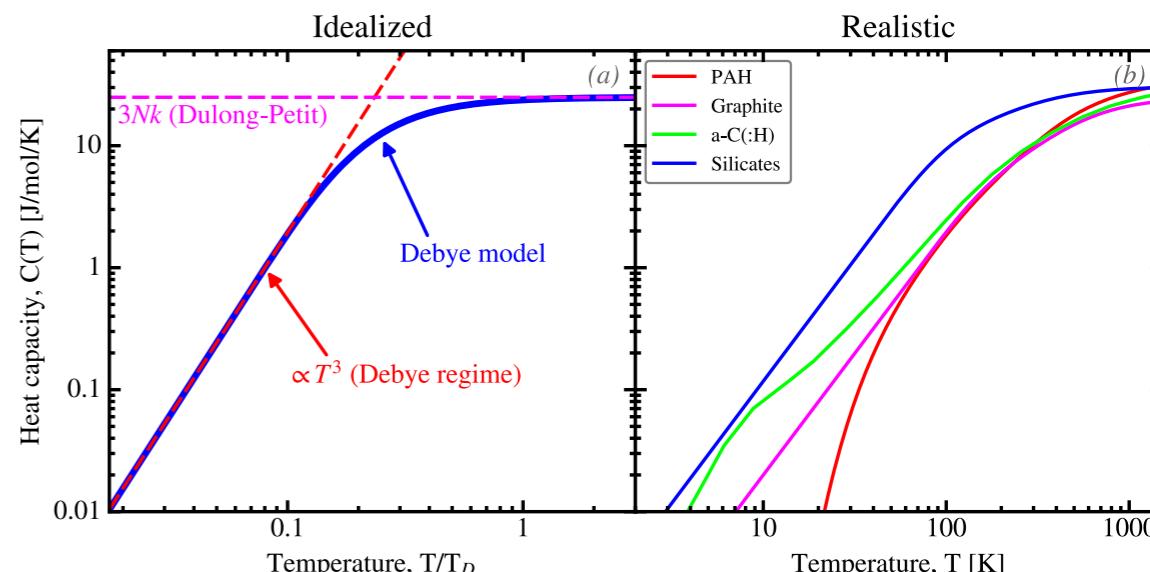


Figure I.25 Galliano (2022)



In the high temperature limit: the heat capacity is given by [the rule of Dulong-Petit](#):

$$C(T) \simeq 3 \frac{N_a}{V} k_B$$

- **Internal Energy:** The internal energy at T is then given by

$$U(T) \equiv V \int_0^T C(T') dT'$$

$$\begin{aligned} U(T) &\approx \frac{3\pi^4}{5} N_a k_B T \left(\frac{T}{\theta_D} \right)^3 \text{ at low temperature} \\ &\approx 3N_a k_B T \quad \text{at high temperature} \end{aligned}$$

The internal energy in the low temperature limit,

$$\begin{aligned} U(T) &\approx \frac{4\pi^5}{5} k_B \frac{a^3 \rho}{\mu m_H} \frac{T^4}{\theta_D^3} \\ &\approx 5.1 \times 10^{-16} [\text{erg}] \left(\frac{a}{10 \text{\AA}} \right)^3 \left(\frac{12}{\mu} \right) \left(\frac{\rho}{2.24 \text{ g cm}^{-3}} \right) \left(\frac{420 \text{ K}}{\theta_D} \right)^3 \left(\frac{T}{10 \text{ K}} \right)^4 \\ &\approx 1.5 \times 10^{-17} [\text{erg}] \left(\frac{a}{10 \text{\AA}} \right)^3 \left(\frac{170}{\mu} \right) \left(\frac{\rho}{3.8 \text{ g cm}^{-3}} \right) \left(\frac{665 \text{ K}}{\theta_D} \right)^3 \left(\frac{T}{10 \text{ K}} \right)^4 \end{aligned}$$

Therefore, ***the internal energy of a very small grain is much smaller than the absorbed photon energy.***

$$h\nu = 1.0 \times 10^{-11} \text{ [erg]} \left(\frac{2000 \text{ \AA}}{\lambda} \right) \longrightarrow \begin{cases} U(T_0) \ll h\nu & \text{for } a \ll 270 \text{ \AA} \\ U(T_0) \gg h\nu & \text{for } a \gg 270 \text{ \AA} \end{cases}$$

Now, the peak temperature can be estimated from:

$$h\nu = U(T_p) - U(T_0)$$

The peak temperature can be estimated as follows:

$$h\nu = U(T_p) - U(T_0) \longrightarrow U(T_p) \simeq h\nu$$

$$T_p \simeq 118 \text{ K} \left(\frac{2000 \text{ \AA}}{\lambda} \right)^{1/4} \left(\frac{a}{10 \text{ \AA}} \right)^{-3/4} \left(\frac{\mu}{12} \right)^{1/4} \left(\frac{\rho}{2.24 \text{ g cm}^{-3}} \right)^{-1/4} \left(\frac{\theta_D}{420 \text{ K}} \right)^{3/4}$$

The peak temperature for a small grain is in general much higher than the equilibrium temperature.

However, the internal energy of a large grain is much larger than the absorbed photon energy.
Therefore, **large grains will stay in a steady state with an equilibrium temperature.**

(2) Heating and Cooling Time Scales

- Time Scales
 - There are three time scales to understand the temperature fluctuation of very small grains.
 - ▶ **Absorption time scale** : How often a grain will absorb a photon?
 - ▶ **Diffuse time scale** : How fast the absorbed energy is distributed over the whole grain?
 - ▶ **Cooling time scale** : How fast the excited vibration modes radiates back the energy to the interstellar space?

- **Absorption Time Scale:**

- The absorption rate of starlight photons by very small grains is

$$\left(\frac{dN_{\text{ph}}}{dt} \right)_{\text{abs}} = \int \frac{u_\nu d\nu}{h\nu} c Q_{\text{abs}}(\nu) \pi a^2 = \int d\nu \frac{4\pi J_\nu}{h\nu} Q_{\text{abs}}(\nu) \pi a^2$$

B. T. Draine does not provide an approximate formula for the absorption rate, but we can roughly estimate it from the heating rate.

$$\left(\frac{dN_{\text{ph}}}{dt} \right)_{\text{abs}} \approx \frac{1}{\langle h\nu \rangle} \left(\frac{dE}{dt} \right)_{\text{abs}} = (\pi a^2) \langle Q_{\text{abs}} \rangle_* \frac{u_* c}{\langle h\nu \rangle}$$

Here, $\langle h\nu \rangle$ is a typical photon energy of the ISRF.

We will take $\langle h\nu \rangle \approx 1 \times 10^{-11} \text{ erg}$ for $\lambda \approx 2000 \text{ \AA}$.

$$u_*/\langle h\nu \rangle \approx 1 \times 10^{-1} U [\text{cm}^{-3}]$$

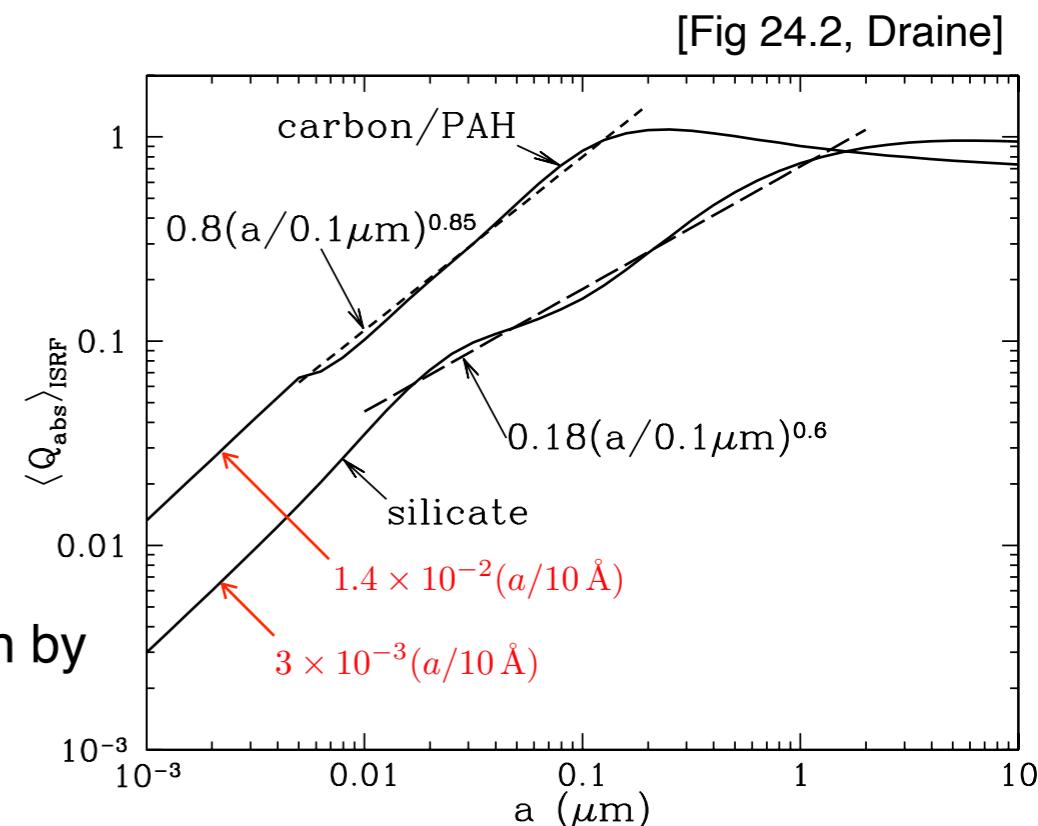
From the right figure, we take

$$\langle Q_{\text{abs}} \rangle_* \approx 1 \times 10^{-2} (a/10 \text{ \AA}) [\text{cm}^2]$$

Recall that the absorption efficiency is proportional to the grain size in the Rayleigh limit.

Hence, the photon absorption rate for small grains is given by

$$\left(\frac{dN_{\text{ph}}}{dt} \right)_{\text{abs}} \approx 1 \times 10^{-6} U (a/10 \text{ \AA})^3 [\text{s}^{-1}]$$



-
- Absorption Rate for Large grains (using the absorption cross-section in the figure)

$$\left(\frac{dN_{\text{ph}}}{dt} \right)_{\text{abs}} \approx (\pi a^2) \langle Q_{\text{abs}} \rangle_* \frac{u_* c}{\langle h\nu \rangle}$$

$$\approx 0.17 U(a/0.1\mu\text{m})^{2.6} [\text{s}^{-1}], \quad \text{for silicate}$$

$$\approx 0.75 U(a/0.1\mu\text{m})^{2.85} [\text{s}^{-1}], \quad \text{for graphite}$$

- ▶ Absorption time scale:

$t_{\text{abs}} \equiv (dN_{\text{ph}}/dt)_{\text{abs}}^{-1}$	
$\approx 1 \times 10^6 U^{-1} [\text{s}] \quad a \sim 10\text{\AA}$	small grain
$\approx 1.3 - 5.9 U^{-1} [\text{s}] \quad a \sim 0.1\mu\text{m}$	large grain

- ▶ ***In diffuse interstellar space, the smaller component of interstellar dust will absorb star light photons once in two hours (for $a \sim 50\text{\AA}$) - 11 days (for $a \sim 10\text{\AA}$). Large dust grains ($a \sim 0.1\mu\text{m}$) will absorb photons once in a couple of seconds.***

- **Diffuse Time scale:** How long does it take to heat up the whole grain?

- The macroscopic diffuse equation:

$$\frac{\partial T}{\partial t} = D \nabla^2 T \quad \text{where } D = \text{diffusion coefficient}$$

$$t_{\text{diff}} \approx a^2/D$$

$$\approx 10^{-12} [\text{s}] (a/10 \text{\AA})^2 (10^{-2} \text{ cm}^2 \text{ s}^{-1}/D) \quad \text{a typical } D \text{ value from Duley (1973, Ap\&SS, 23, 43)}$$

- Sound crossing time: Energy disturbance in a solid is transmitted to its lattice vibrational modes (photons) and phonons travel at the speed of sound, which is about a few 10^5 cm/s.

$$t_{\text{diff}} \approx 2a/c_s$$

$$\approx 2 \times 10^{-12} [\text{s}] (a/10 \text{\AA}) (10^5 \text{ cm s}^{-1}/c_s)$$

$c_s \sim 1 \text{ km s}^{-1}$ in a cold material.

- We may conclude that ***the grain is heated to a peak temperature instantaneously upon the absorption of a photon***, which is determined by:

$$h\nu = \frac{4\pi a^3}{3} \int_{T_0}^{T_p} C(T') dT' \quad C(T) = \text{heat capacity per unit volume}$$

where T_0 is the temperature just prior to the absorption event.

- ***The cooling time scale***

- ▶ The time evolution of grain temperature can be calculated by solving

$$\frac{dU}{dt} = -(4\pi)(\pi a^2) \int Q_{\text{abs}} B_\nu d\nu = -4\pi a^2 \langle Q_{\text{abs}} \rangle_T \sigma_{\text{SB}} T^4$$

$$\frac{dT}{dt} = -\frac{3}{aC(T)} \langle Q_{\text{abs}} \rangle_T \sigma_{\text{SB}} T^4$$

$$\begin{aligned} dU &= \frac{4\pi a^3}{3} C(T) dT \\ \leftarrow \quad \int Q_{\text{abs}} B_\nu d\nu &= \frac{\sigma_{\text{SB}}}{\pi} T^4 \end{aligned}$$

- ▶ The initial cooling time scale can be estimated by

$$\begin{aligned} t_{\text{cool}} &\approx \frac{U(T_p)}{|dU/dt|} \\ &\approx \frac{h\nu}{\Lambda} \end{aligned}$$

$$\begin{aligned} h\nu &= 1.0 \times 10^{-11} [\text{erg}] \left(\frac{2000 \text{ \AA}}{\lambda} \right) \\ \Lambda &= \left| \frac{dU}{dt} \right|_{T_p} = 4\pi a^2 \langle Q_{\text{abs}} \rangle_T \sigma_{\text{SB}} T_p^4 \end{aligned}$$

- ▶ We will assume an approximate, absorption efficiency, following S. S. Hong (홍승수, 1979, JKAS):

$$Q_{\text{abs}}(\nu) = \frac{4\pi a}{\lambda} = Q_0 (\lambda_0 / \lambda) \quad \leftarrow \quad Q_{\text{abs}}(\nu) = 4 \frac{2\pi a}{\lambda} \text{Im} \left(\frac{m^2 - 1}{m^2 + 2} \right) \quad \text{Rayleigh limit}$$

Here, $Q_0 \approx 6.283 \times 10^{-2} \chi(a/10 \text{ \AA})$

$\lambda_0 = 2000 \mu\text{m}$

$\beta = 1$

$\chi \simeq 1$

From the previous formula for $Q_{\text{abs}}(\nu) = Q_0(\lambda_0/\lambda)^{-\beta}$

the planck-averaged absorption efficiency is given by

$$\begin{aligned}\langle Q_{\text{abs}} \rangle_T &= \frac{15}{\pi^4} \Gamma(4 + \beta) \zeta(4 + \beta) Q_0 \left(\frac{kT}{h\nu_0} \right)^\beta \\ &= 3.83220 Q_0 \left(\frac{kT}{h\nu_0} \right) \quad \text{for } \beta = 1 \\ &\approx 3.950 \times 10^{-2} \chi(a/10 \text{\AA})(\lambda/2000 \text{\AA})(T/118 \text{K})\end{aligned}$$

The cooling rate is:

$$\begin{aligned}\Lambda &= 4\pi a^2 \langle Q_{\text{abs}} \rangle_T \sigma_{\text{SB}} T^4 \\ &= 5.46 \times 10^{-11} [\text{erg s}^{-1}] \chi(a/10 \text{\AA})^3 (\lambda/2000 \text{\AA}) (T/118 \text{K})^5\end{aligned}$$

The cooling time scale at the peak temperature is:

$$\begin{aligned}t_{\text{cool}} &\approx \frac{U(T_p)}{|dU/dt|} = \frac{h\nu}{\Lambda} \\ &\approx 0.18 [\text{s}] \chi^{-1} (10 \text{\AA}/a)^3 (118 \text{K}/T)^5\end{aligned}$$

Because of the strong temperature dependence of the cooling rate, **the grain actually needs much longer time (up to 10^4 sec) to radiate its excess (absorbed) energy back** to the interstellar space.

In summary,

$$t_{\text{abs}} \gg t_{\text{diff}}, \quad t_{\text{abs}} \gg t_{\text{cool}}, \quad t_{\text{cool}} \gg t_{\text{diff}}$$

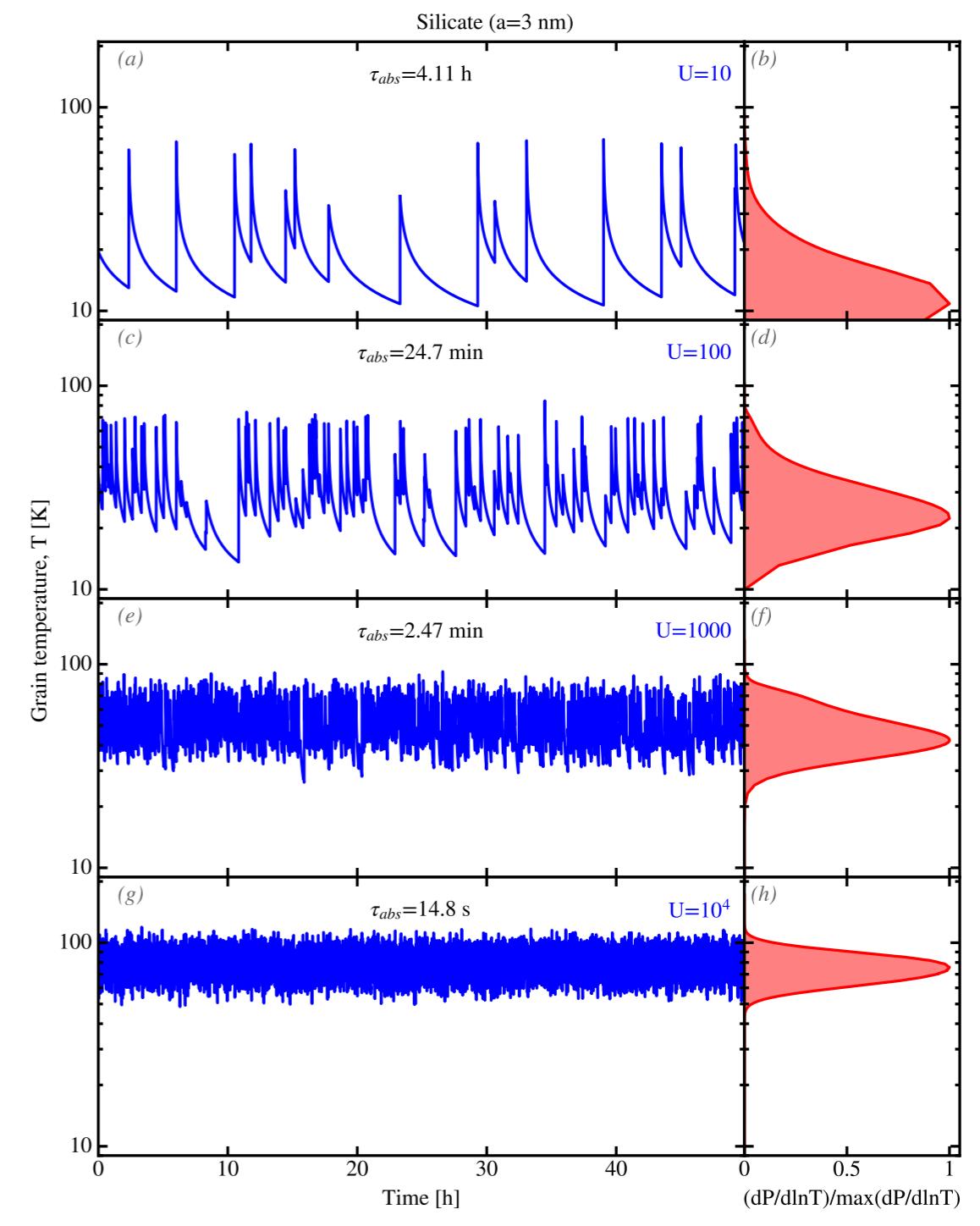
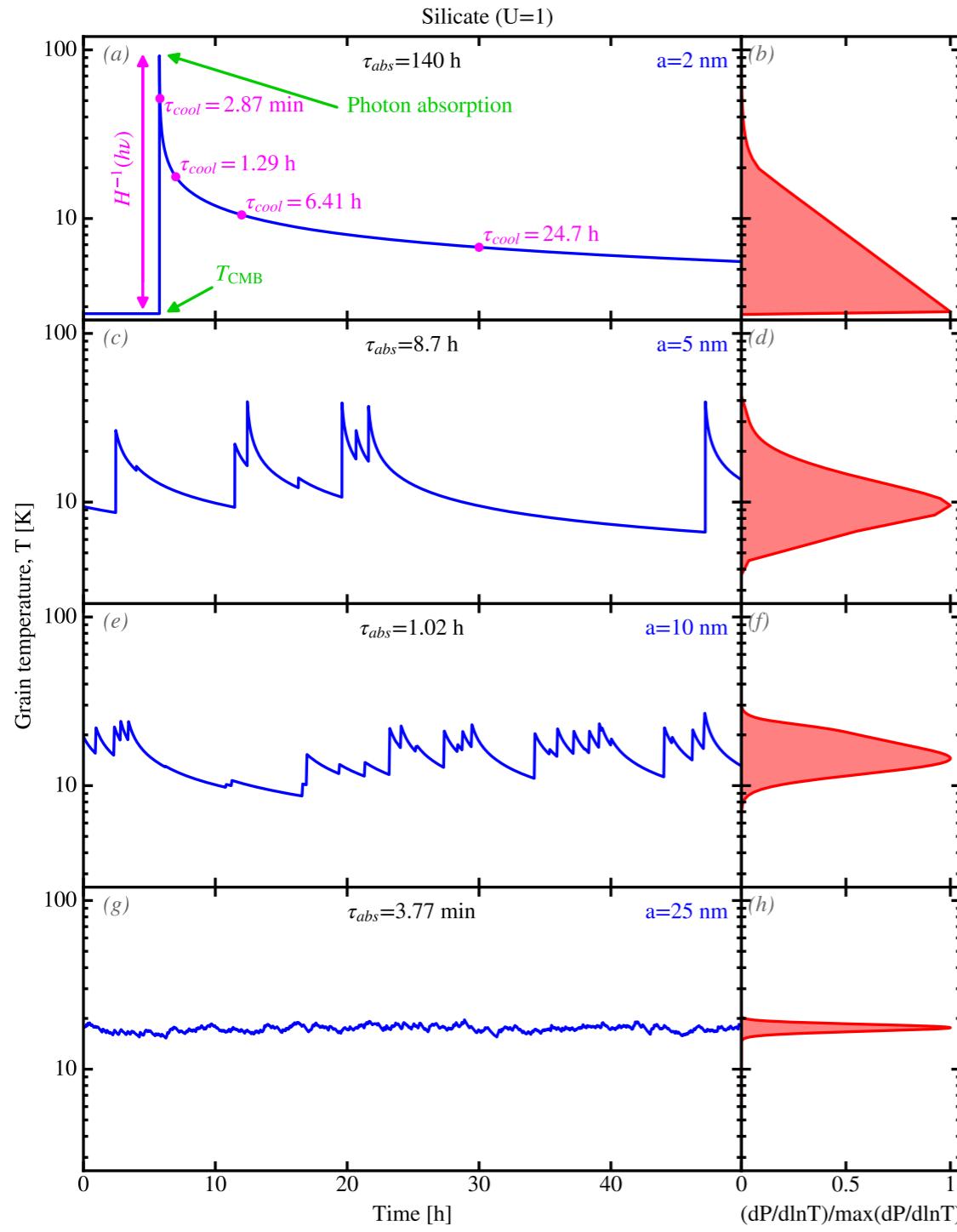
One may want to use the following formula for $\beta = 2$:

$$\langle Q_{\text{abs}} \rangle_T \approx 1 \times 10^{-6} (a/0.1 \mu\text{m}) (T/\text{K})^2$$

The Stochastic Time Evolution of Grain Temperature

Monte-Carlo simulations of the temperature fluctuation:

See Draine & Anderson (1985, ApJ, 292, 494) and Krugel (The Physics of Interstellar Dust, IoP).



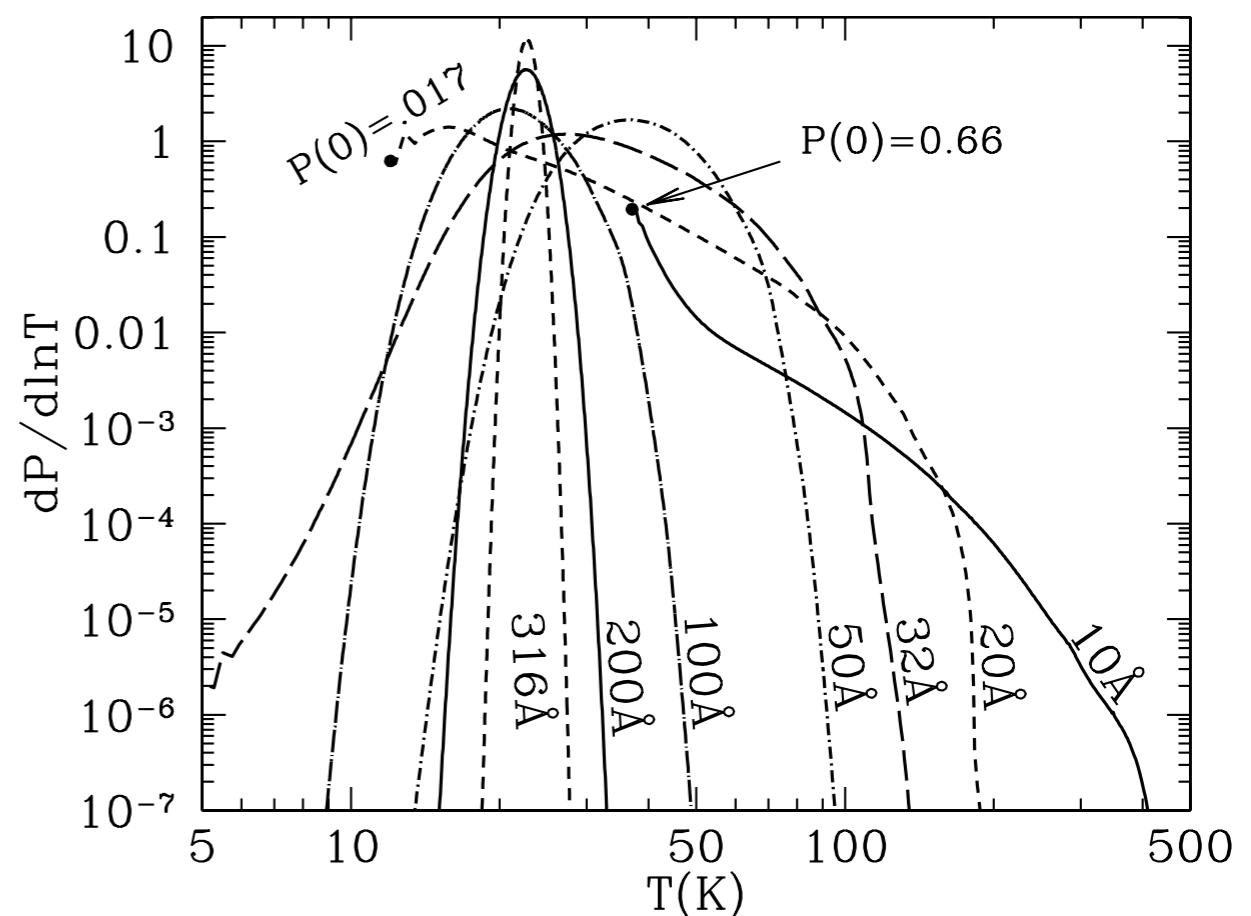
The importance of quantized stochastic heating is evident for the smallest sizes. [Figures I.30 and I.31, Galliano]

- Temperature History:
 - ▶ Two effects become increasingly important with diminishing grain size: (1) **the heat capacity of the dust becomes sufficiently small that single-particle hits can cause large spikes in the dust temperature** and (2) **the absorption rate with photons becomes sufficiently low that the cooling of the dust between successive collisions becomes important.**
 - ▶ ***For very small dust grains, one cannot speak of a representative grain temperature under these conditions - one must instead use a temperature distribution function.***
- Temperature Distribution Function:
 - ▶ Consider a large ensemble of identical grains in some interstellar environment. Let us define:
$$P(T)dT = \text{the probability that its temperature lies in the interval from } T \text{ and } T + dT.$$
The probability density is of course normalized:
$$\int_0^\infty P(T)dT = 1$$
 - ▶ The temperature distribution will depend on grain size, composition, and the intensity (and spectrum) of the radiation illuminating the grains.

- The following figure shows temperature distribution functions calculated for graphite/PAH dust grains of selected radii, exposed to the ISRF of Mathis et al. (1983).
 - ▶ We see that the distribution function for a grain with $a = 10\text{\AA}$ extends to $T = 400\text{ K}$ - this is the temperature that this grain will be heated to when it absorbs a single photon with energy just below the Lyman limit cutoff at 13.6 eV.
 - ▶ Most of the energy radiated by the grain is radiated while it is at temperature above 100 K - the typical absorbed photon raises the grain temperature to $T > 200\text{ K}$. Such grains radiate strongly in the PAH features at 7.7, 8.6, and 11.3 μm .

Temperature distribution function for seven carbonaceous grains in ISRF with $U = 1$. Curves are labeled by grain radius. For $a = 10$ and 20\AA curves, the dot indicates the first excited state and $P(0)$ is the fraction in the ground state.

[Fig 24.6, Draine]



- ▶ **For normal interstellar grains of average size**, the temperature oscillates only a little around an equilibrium value T_{eq} and in the limit of large grains, the distribution density function approaches the delta-function:

$$P(T) \rightarrow \delta(T - T_{\text{eq}})$$

where the equilibrium temperature follows from the steady-state balance between emission and absorption:

First, calculate the equilibrium temperature.

$$\int Q_{\nu}^{\text{abs}} J_{\nu} d\nu = \int Q_{\nu}^{\text{abs}} B_{\nu}(T_{\text{eq}}) d\nu$$

Second, the emissivity (per solid angle per area per time)

$$\rightarrow j_{\nu} = \pi a^2 Q_{\nu}^{\text{abs}} B_{\nu}(T_{\text{eq}})$$

- ▶ **Even for a very small particle** we will assume that its radiation obeys at any time Kirchhoff's law, so in the case of a sphere of radius a , we can express the average monochromatic emission per solid angle by

$$j_{\nu} = \pi a^2 Q_{\nu}^{\text{abs}} \int B_{\nu}(T) P(T) dT$$

Although the emission of a single such grain is not time-constant, the whole ensemble radiates at any frequency at a steady state.

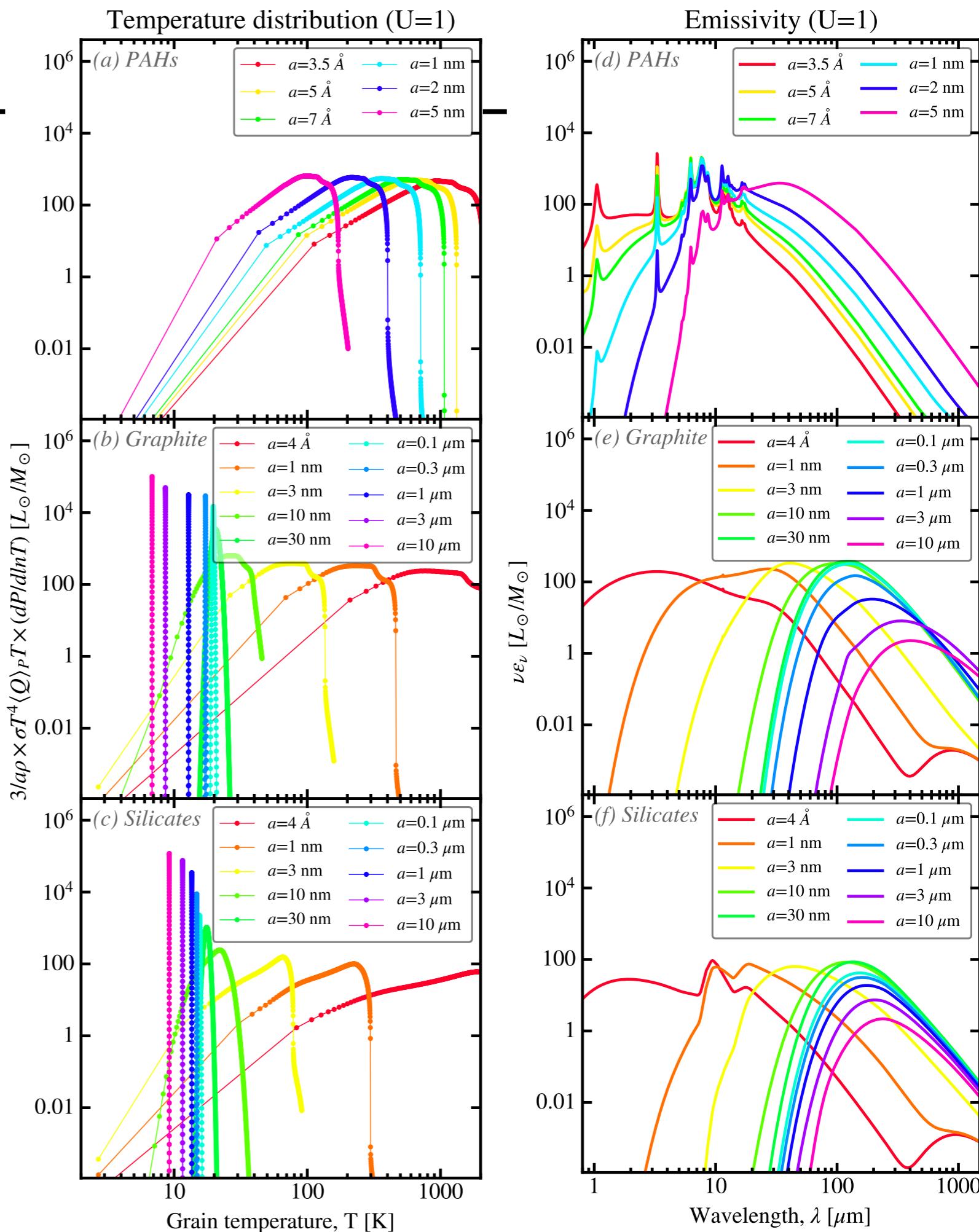
Stochastically heated grains

[Figure I.32, Galliano]

The left panels show the temperature distribution of grains with different radii.

The right panels show the corresponding emissivities.

The PAHs are a mixture of neutral (50%) and ionized (50%) molecules (Draine & Li 2007), the graphite and silicates are from Draine (2023a,b).



Method of Calculation

- **Monte-Carlo simulation** (Draine & Anderson 1985)

- The cooling flux of low energy infrared photons and heating by photons with $\lambda > \lambda_c$ ($= 1000\mu\text{m}$) are assumed to be continuous. However, ***heating by photons with $\lambda < \lambda_c$ must be treated as a sequence of stochastic absorptions of single energetic photons.***
- The probability a photon is absorbed by a grain in a time interval $(t, t + dt)$ is given by

$$P(t)dt = \dot{p}e^{-\dot{p}t}dt$$

$$\dot{p} = \int_0^{\lambda_c} \pi a^2 Q_{\lambda}^{\text{abs}} \frac{u_{\lambda} c}{h\nu} d\lambda \quad [\text{s}^{-1}] \quad \rightarrow$$

$$t = -\frac{1}{\dot{p}} \ln \xi \quad (0 \leq \xi < 1)$$

ξ = uniform random number

Heating of the grain by photons with $\lambda > \lambda_c = 1000\mu\text{m}$ is modeled as continuous.

- The energy (wavelength) of the absorbed photon is randomly determined according to

$$\frac{1}{\dot{p}} \frac{d\dot{p}}{d\lambda} d\lambda = \frac{Q_{\lambda}^{\text{abs}} (u_{\lambda} c / h\nu)}{\int Q_{\lambda}^{\text{abs}} (u_{\lambda} c / h\nu) d\lambda} d\lambda$$

- The temperature of the grain is then assumed to be immediately increased to that estimated from:

$$\frac{hc}{\lambda} = \frac{4\pi a^3}{3} \int_{T_0}^T C(T')dT'$$



$$dU = \frac{4\pi a^3}{3} C(T) dT$$

Here, T_0 is the temperature just prior to the absorption event.

- For each grain type and size, the temperature evolution can be followed by solving the differential equation:

$$\frac{dU}{dt} = 4\pi^2 a^2 \left[\int_{\lambda_c}^{\infty} Q_{\lambda}^{\text{abs}} J_{\nu}(t) d\lambda - \int_0^{\infty} Q_{\lambda}^{\text{abs}} B_{\lambda}(T(t)) d\lambda \right]$$

continuous heating continuous cooling

Heating of the grain by photons with $\lambda > \lambda_c = 1000 \mu\text{m}$ is modeled as continuous.

$$\frac{dU}{dt} = \frac{4\pi a^3}{3} C(T) \frac{dT}{dt}$$

$$\longrightarrow \frac{dT}{dt} = \frac{3\pi}{aC(T)} \left[\int_{\lambda_c}^{\infty} Q_{\nu}^{\text{abs}} J_{\nu}(t) d\nu - \int_0^{\infty} Q_{\nu}^{\text{abs}} B_{\nu}(T(t)) d\nu \right]$$

The right side describes the difference between the power absorbed from the radiation field J_{ν} and the cooling rate.

- Having followed $U(t)$ or $T(t)$ over a long period, one can obtain the probability density $P(T)$ from the fraction of time that the grain has spent in the temperature interval T and $T + dT$.

-
- Discretized steady state distribution function
 - It is far more efficient to solve directly for the discretized steady state distribution function.
 - ▶ **Guhathakurta & Draine (1989, ApJ, 345, 230)** developed a fast method to calculate the temperature distribution function.
 - ▶ **Draine & Li (2001, ApJ 551, 807)** presented a statistically exact, and quantum-mechanical treatment.
 - ▶ **Desert et al. (1986, A&A, 160, 295)** also developed a different method to compute the temperature distribution function, which is publicly available (**DustEM**; <https://www.ias.u-psud.fr/DUSTEM/>; Campiegne et al. 2010, A&A, 724, 44).
 - ▶ **Dwek (1986, ApJ, 302, 363)** presented a method to calculate the temperature fluctuations in dust particles when they collide with low-density, X-ray emitting electrons.
 - Guhathakurta & Draine (1989)
 - ▶ When a grain absorbs or emits a photon, its internal energy $U(T)$, which is a function of temperature, changes. We bin $U(T)$ into N states (bins) U_j of width ΔU_j .
Let P_j the probability for a grain to be within the j th state.
Let A_{fi} the transition probability that a single grain changes from state i to f .

-
- ▶ In statistical equilibrium, for each level j the number of populating and depopulating events must be equal.

$$\sum_{k \neq j} A_{jk} P_k = P_j \sum_{k \neq j} A_{kj}$$

population de-population

Note that P_j is proportional to the energy width ΔU_j

With the purely mathematical definition:

$$A_{jj} = - \sum_{k \neq j} A_{kj}$$

We can write the equilibrium equation as the following matrix equation:

$$\sum_{k=1}^N A_{jk} P_k = 0$$

- ▶ Only $N - 1$ of these N equations are linearly independent.

One may first put $P_1 = 1$ (or $P_N = 1$) and then solve for P_2, \dots, P_N (or P_1, \dots, P_{N-1}), and then rescale them:

$$\sum_{j=1}^N P_j = 1$$

- Matrix Elements

 - ▶ Heating:

$$A_{kj} = \frac{4\pi C_\nu^{\text{abs}} J_\nu}{h\nu} \Delta U_k \quad \text{where } h\nu = |U_k - U_j| \quad (j < k)$$

 - ▶ Cooling:

$$A_{kj} = \frac{4\pi C_\nu^{\text{abs}} B_\nu(T_j)}{h\nu} \Delta U_k \quad \text{where } h\nu = |U_j - U_k| \quad (j > k)$$

 - ▶ Above the main diagonal stand the cooling elements, below those for heating. The energy balance requires for the cooling and heating rate for each level j :

heating: $\sum_{k>j} A_{kj} |U_k - U_j| = \int_0^\infty 4\pi C_\nu^{\text{abs}} J_\nu d\nu$

cooling: $\sum_{k<j} A_{kj} |U_k - U_j| = \int_0^\infty 4\pi C_\nu^{\text{abs}} B_\nu(T_j) d\nu$

 - ▶ ***“Thermal Continuous” Cooling Approximation:*** As cooling proceeds via infrared photons which have low energy, their emission changes the grain temperature very little. ***This suggests that in cooling one needs to consider only the transitions to the levels immediately below*** ($k = j - 1$, $j \rightarrow j - 1$).

One can, therefore, put all matrix elements A_{fi} above the main diagonal to zero, except $A_{j-1,j}$. Then, in order to fulfill the energy equation, the ***cooling elements*** have to be written as:

$$\sum_{k>j} A_{kj} |U_k - U_j| = \int_0^\infty 4\pi C_\nu^{\text{abs}} J_\nu d\nu$$

$$A_{j-1,j} |U_{j-1} - U_j| = \int_0^\infty 4\pi C_\nu^{\text{abs}} J_\nu d\nu$$

$$\rightarrow A_{j-1,j} = \frac{1}{|U_j - U_{j-1}|} \int_0^\infty 4\pi C_\nu^{\text{abs}} B_\nu(T_j) d\nu$$

$$A_{kj} = 0 \text{ if } k < j - 1$$

- However, ***heating should not be reduced to single transitions*** $j \rightarrow j + 1$.
- Guhathakurta & Draine (1989) describe a simple trick to avoid numerical rounding errors in calculating the temperature distribution.

Setting $X_j = P_j/P_1$ ($X_1 = 1$), we can obtain the steady state solution for the distribution function:

$$\frac{d}{dt} \left(\sum_{j=1}^{f-1} P_j \right) = A_{f-1,f} P_f - \sum_{j=1}^{f-1} P_j \sum_{k=f}^N A_{k,j} \quad \downarrow$$

$$X_f = \frac{1}{A_{f-1,f}} \sum_{j=1}^{f-1} B_{f,j} X_j \quad \text{where} \quad B_{f,j} \equiv \sum_{k=f}^N A_{k,j} \quad (f > j)$$

However, in practice, even adopting the trick, calculating the temperature distribution function is a bit tricky; it is hard to define a proper interval for U or T .

- Infrared Emission from Grains
 - In a typical spiral galaxy, perhaps 30-50% of the energy radiated by stars is absorbed by dust grains, and reemitted in the IR. The spectrum of this emission is determined by the temperatures and composition of the dust grains.
 - IR emission is a quantum process - a radiative transition between an upper and lower vibrational level of the grain. However, it has been shown (Draine & Li 2001) that a “thermal” approach provides an excellent approximation.
 - In the approximation, the emissivity of a population of grains can be written:

$$4\pi j_\nu = 4\pi \sum_i \int da \frac{dn_i}{da} \int dT \left(\frac{dP}{dT} \right)_{i,a} C_\nu^{\text{abs}}(i, a) B_\nu(T)$$

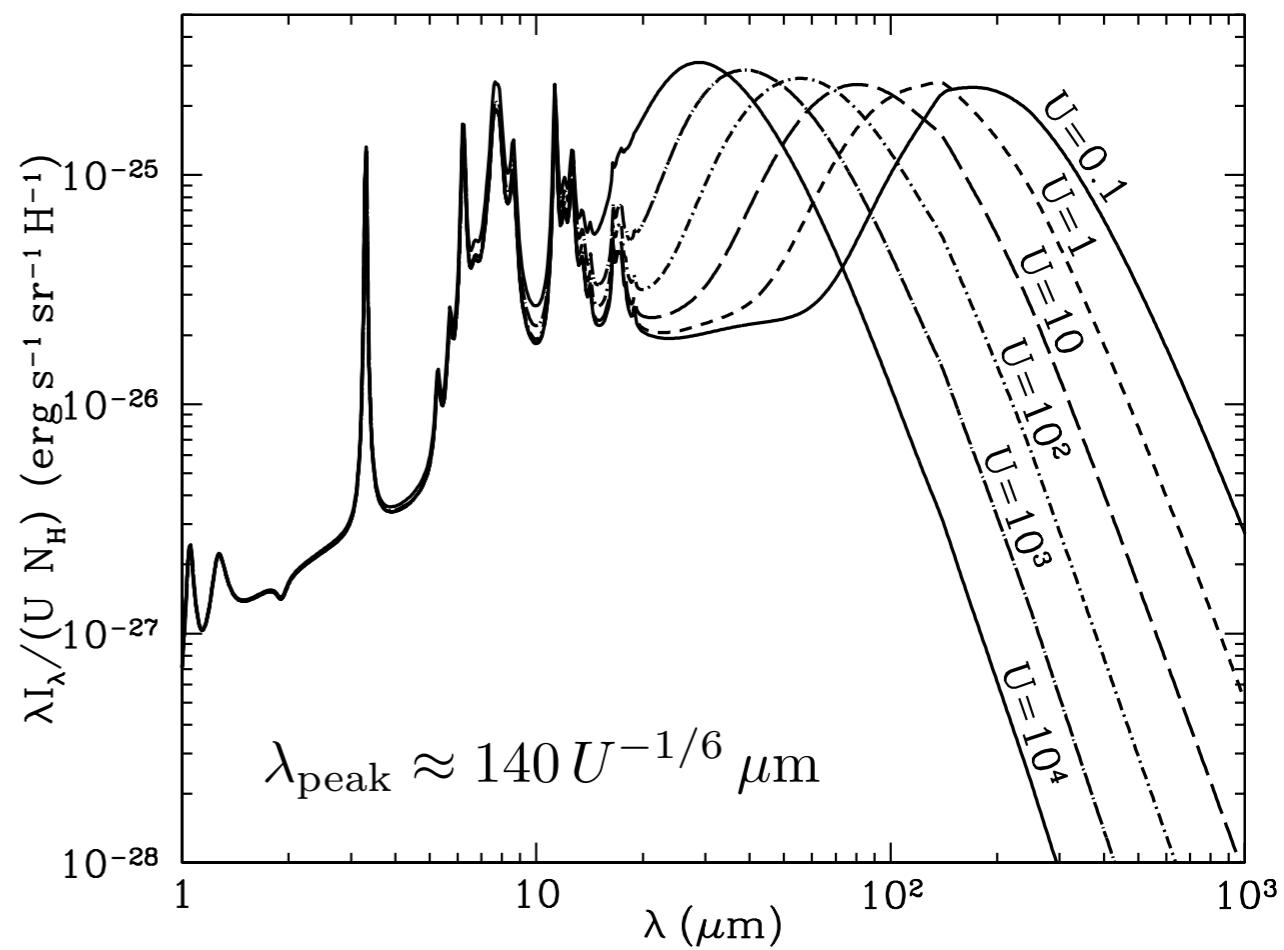
where $(dn_i/da)da$ is the number density of grains of type i with radii in $[a, a + da]$.

- ▶ We, therefore, need a grain model to provide the size distribution for each composition, the absorption cross sections and temperature distribution functions.
- ▶ For large grains, the temperature distribution is sufficiently narrow that it may be approximated by a delta function $dP/dT \rightarrow \delta(T - T_{\text{eq}})$, where T_{eq} is the steady-state temperature for which the time-averaged cooling equals the time-averaged heating.
- ▶ But for $a \lesssim 0.01 \mu\text{m} = 100 \text{ \AA}$, one should use a realistic temperature distribution dP/dT rather than a delta-function.

- Model IR emission spectra have been calculated by Draine & Li (2007).
 - ▶ The spectra shown in the following figure is normalized by the ISRF strength U .
 - ▶ The peak in the Far-IR corresponds to emission from “large” grains. The equilibrium temperature of the “large” grains are approximately:

$$T_{\text{eq}} \approx 20 U^{1/6} \text{ K}$$

- ▶ The peak shifts toward shorter wavelength as U increases and the grains become warmer. the peak $(\lambda I_\lambda)_{\text{peak}}$ occurs at $\lambda_{\text{peak}} \approx 140 U^{-1/6} \mu\text{m}$.
- ▶ The additional peaks at shorter wavelengths are due to vibrational modes of PAH grains; these features account for $\sim 25\%$ of the total power, but these features (normalized by the total power) hardly change as the radiation intensity is changed.
- ▶ The PAH emission occurs following single-photon heating of very small grains.



[Fig 24.7, Draine]

Evolution of Interstellar Dust in Galaxies

- **Sources**

- **Stellar sources**: AGB stars, Supernovae, Novae,...
- **Interstellar**: Growth of solid material in the ISM by accretion of gas onto grain surfaces
 - ▶ If stellar injection dominates: then interstellar dust is mainly “stardust”, with composition of stardust.
 - ▶ If interstellar growth dominates: then interstellar dust consists mainly of whatever materials can grow (and survive) in the ISM.
- Both processes are active. ***Which is dominant in the Galaxy?***

- **Sinks**

- Destruction of solid material in ISM
 - ▶ sputtering (in shock waves and shock-heated gas)
 - ▶ grain-grain collisions \Rightarrow vaporization
 - ▶ photodesorption
 - ▶ other (cosmic rays, Coulomb explosions, sublimation,...)
- Incorporation into protostars
- Removal in galactic outflow

From B. T. Draine's keynotes

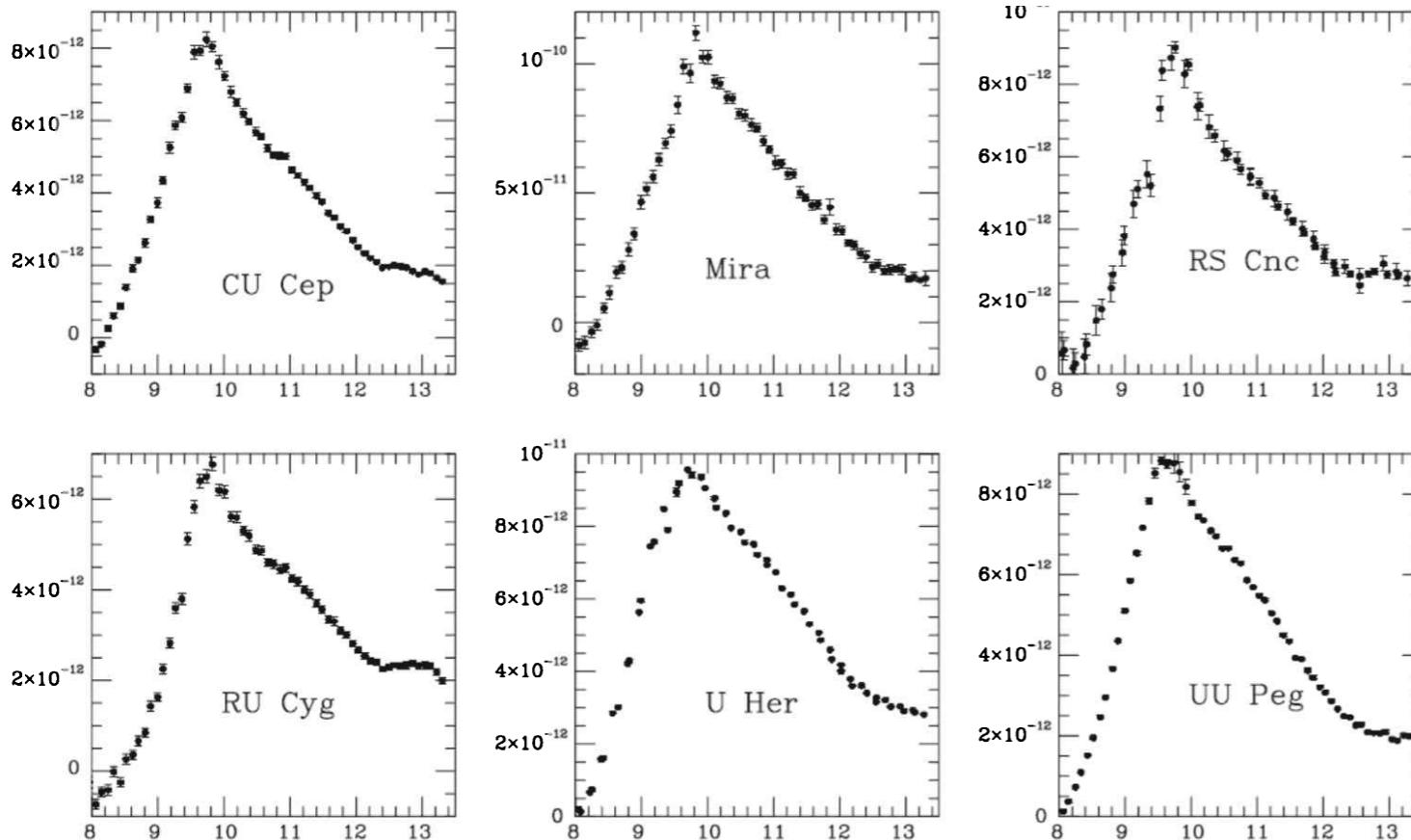
Sputtering: a phenomenon in which microscopic particles of a solid material are ejected from its surface, after the material is bombarded by energetic particles of a plasma or gas.

Coulomb explosion: when an intense electric field excites the valence electrons in a solid, they are ejected from the system and leaving behind positively charged ions. The chemical bonds holding the solid are weakened. The result is an explosion of ions and electrons by Coulomb repulsion.

Photodesorption: a quantum effect whereby a single photon causes the desorption of an atom or molecule from a solid surface.

Stellar Sources - Example

- Observationally, it is seen that these “dust nurseries” exist in the dense, cool stellar winds of AGB stars, old planetary nebulae, and the cooling envelopes of novae. In these winds, the densities are as much as $\sim 10^9 \text{ cm}^{-3}$, and the temperatures can drop below the condensation temperature of many heavy elements ($\sim 1000 \text{ K}$).
- For instance, Mira variable stars are AGB stars with substantial emission at mid-infrared wavelengths. In oxygen-rich Mira variables, most of carbon is locked up in gaseous CO, and the main condensates are silicates. Oxygen-rich Mira variables often have a strong silicate emission feature at $\lambda \sim 9.7\mu\text{m}$.



Continuum-subtracted spectra of oxygen-rich Mira variable stars.

The x-axis shows wavelength in μm and the y-axis flux in $\text{W m}^{-2} \mu\text{m}^{-1}$

[Speck et al. 2000]

-
- ▶ In carbon-rich Mira variables, the mid-IR thermal dust continuum is frequently accompanied by an emission feature at $\lambda \sim 11.3\mu\text{m}$, due to the presence of SiC. However, in the winds of carbon-rich Mira variables, much of the carbon condenses out in the form of PAHs rather than silicon carbide grains, because of the scarcity of silicon relative to carbon.
 - ▶ A typical mass loss rate for a Mira variable star is $\dot{M} \sim 10^{-6} M_{\odot} \text{ yr}^{-1}$. Given a dust-to-gas ratio of ~ 0.01 (by mass, considering only hydrogen), this implies that a Mira variable can produce as much as

$$\dot{M}_{\text{dust}} \sim 10^{-8} M_{\odot} \text{ yr}^{-1}$$

of dust during its $\sim \text{Myr}$ time scale for copious mass loss.

$$M_{\text{dust}} \approx \dot{M}_{\text{dust}} \times \Delta t_{\text{life}} \approx 0.01 M_{\odot}$$

The newly formed $\sim 0.01 M_{\odot}$ of stardust is spread into the ISM, where it undergoes competing processes of growth and destruction.

Growth in ISM

- How dust grains form?
 - Even in a fairly dense molecular cloud core, with 3×10^4 atoms per cm³, at cosmic abundances we expect only 1 Si atom and 8 C atoms per cm³. How can we make such widely separated atoms to form a dense solid grain?
 - **Gas accretion on grains** - How fast dust grains can grow?
 - ▶ Let's imagine a solid spherical grain made of element X. The number density of X atoms in the interstellar gas is n_X , and the rms thermal velocity of an X atom is

$$v_X = \left(\frac{3kT}{2m_X} \right)^{1/2}$$

T = the gas temperature
 m_X = the mass of an X atom

- ▶ If there exists a “seed” grain of radius a , it will accrete additional X atoms from the surrounding gas, and the grain’s mass will grow at the rate

$$\frac{dM_{\text{gr}}}{dt} = (n_X v_X)(\pi a^2) m_X \mathcal{P}_s$$

Here, \mathcal{P}_s is the “*sticking probability*”; that is, the probability that an X atom striking the grain will stick instead of bouncing off.

The mass of the spherical grain is

$$M_{\text{gr}} = \frac{4\pi}{3} a^3 \rho_{\text{gr}}$$

ρ_{gr} = the bulk density of element X.

- We can rewrite the mass growth rate as a radius growth rate:

$$\frac{da}{dt} = \frac{n_X v_X \pi a^2 m_X \mathcal{P}_s}{4\pi \rho_{\text{gr}} a^2} = \frac{n_X v_X m_X \mathcal{P}_s}{4\rho_{\text{gr}}}$$

- Thus, the growth rate da/dt is independent of the grain size a , unless the sticking probability \mathcal{P}_s is a function of grain size. For carbon atoms in the cold neutral medium ($T \sim 100$ K, $n_{\text{C}}/n_{\text{H}} \approx 3 \times 10^{-4}$.

$$\begin{aligned} \frac{da}{dt} &= \frac{n_{\text{C}}}{n_{\text{H}}} \frac{n_{\text{H}} m_{\text{C}} \mathcal{P}_s}{4\rho_{\text{gr}}} \left(\frac{3kT}{2m_{\text{C}}} \right)^{1/2} \\ &\approx 0.2 [\mu\text{m/Gyr}] \mathcal{P}_s \left(\frac{n_{\text{H}}}{30 \text{ cm}^{-3}} \right) \left(\frac{T}{100 \text{ K}} \right)^{1/2} \end{aligned}$$

$\rho_{\text{gr}} = 2.24 \text{ g cm}^{-3}$ for graphite
 See slide 62, the residence time scale of atoms (grains) in the ISM.

- This simple calculation suggests that ***grains can grow to the observed size during a time scale shorter than the age of our galaxy as long as atom-grain collisions are sticky.***
- It becomes much easier to grow grains, if there exist seed grains produced in regions that are denser and/or cooler than typical regions of the ISM.
- ***Grain-Grain collisions:*** Dust grains can grow as the result of grain-grain collisions. However, they can also be shattered into smaller fragments by high-speed collisions.
 [shattering: action of breaking something into small pieces]

Destruction of Grains: Sputtering in Hot Gas

- **Sputtering:** Grains can be vaporized by shock heating in supernova blast wave.
 - ▶ Gas atoms moving at slow speeds can be accreted onto a dust grain. However, *in a hotter gas, the collision of high-speed ions can erode the grain*, chipping away one or more atoms at a blow. This process of erosion by high-speed atoms or ions is called sputtering.
 - ▶ The molecular bonds holds together the atoms in a dust grain. They typically have a bond energy $E_{\text{bond}} \sim 5 \text{ eV}$. Thus, if a dust grain is surrounded by gas, the gas particles will cause sputtering if

$$T \gg \frac{2B_{\text{bond}}}{3k} \sim 40,000 \text{ K} \left(\frac{E_{\text{bond}}}{5 \text{ eV}} \right)$$

- ▶ Detailed studies: For temperatures in the range $10^6 \text{ K} < T < 10^9 \text{ K}$, the sputtering rate for a graphite or silicate grain can be approximated as

$$\frac{da}{dt} \simeq -1 \mu\text{m Myr}^{-1} \left(\frac{n_{\text{H}}}{1 \text{ cm}^{-3}} \right)$$

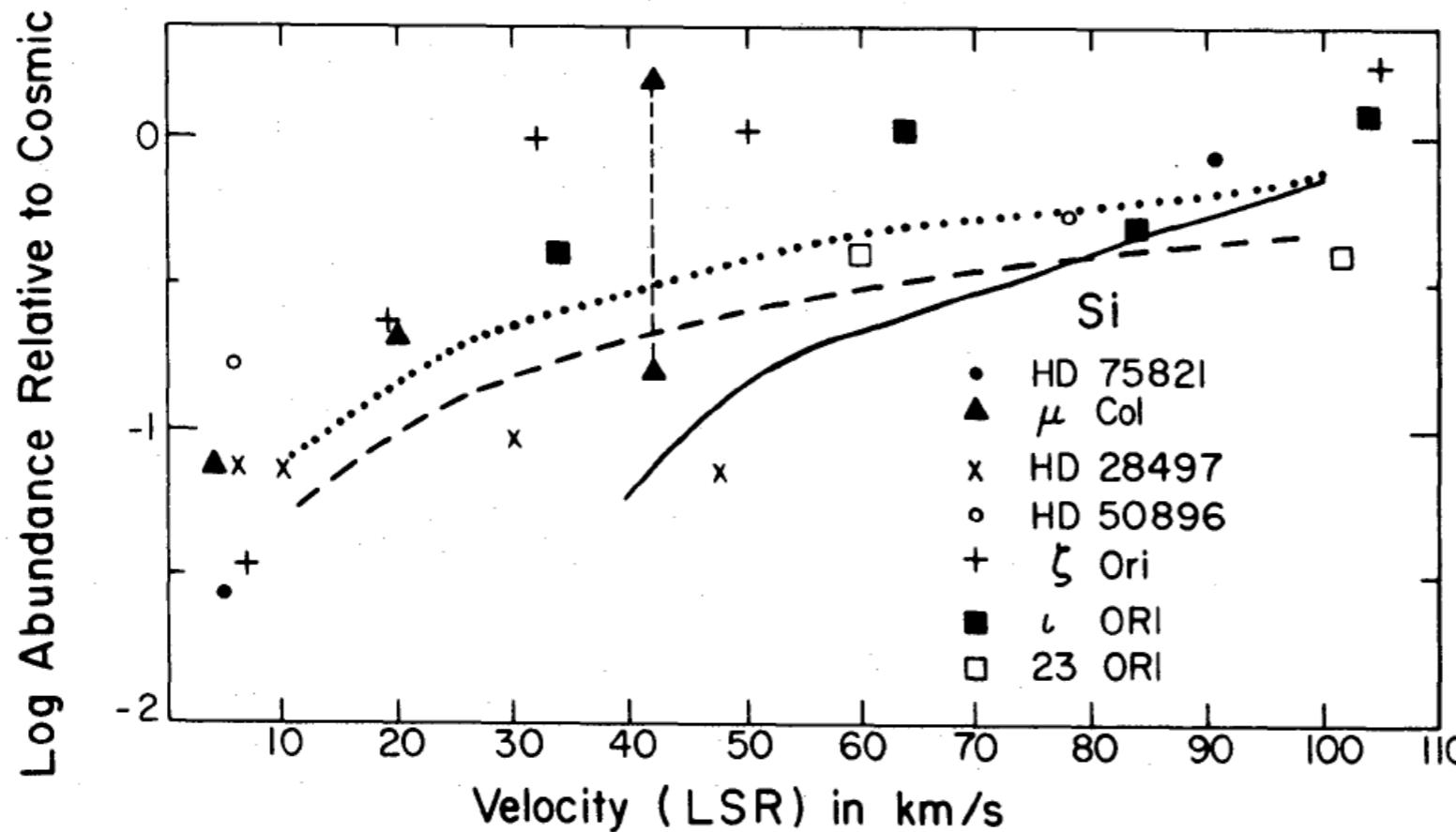
This sputtering rate corresponds to a lifetime:

$$t_{\text{sput}} = \frac{a}{|da/dt|} \approx 0.1 \text{ Myr} \left(\frac{a}{0.1 \mu\text{m}} \right) \left(\frac{n_{\text{H}}}{1 \text{ cm}^{-3}} \right)^{-1}$$

In a supernova remnant, $n_{\text{H}} \sim 1 \text{ cm}^{-3}$ and $T \gtrsim 10^6 \text{ K}$, giving a lifetime of 0.1 Myr for a typical dust grain in the ISM.

- ▶ In the intracluster medium of the Coma Cluster, $n_{\text{H}} \sim 3 \times 10^{-3} \text{ cm}^{-3}$ and $T \sim 10^8 \text{ K}$, the lifetime is ~ 30 Myr for a typical dust grain. Thus, we don't expect the intracluster medium to be dusty unless there are strong Galactic outflow sources. **[However, there is indirect evidence of the existence of dust in intracluster medium, from the study of statistical reddening of background galaxies/quasars.]**

- In low-velocity gas, Si is heavily depleted.
- Gas moving at high speeds in ISM
 - ▶ has been accelerated by a shock wave.
 - ▶ shows evidence of grain destruction: enhanced gas-phase abundances of Si and other species that are normally depleted.

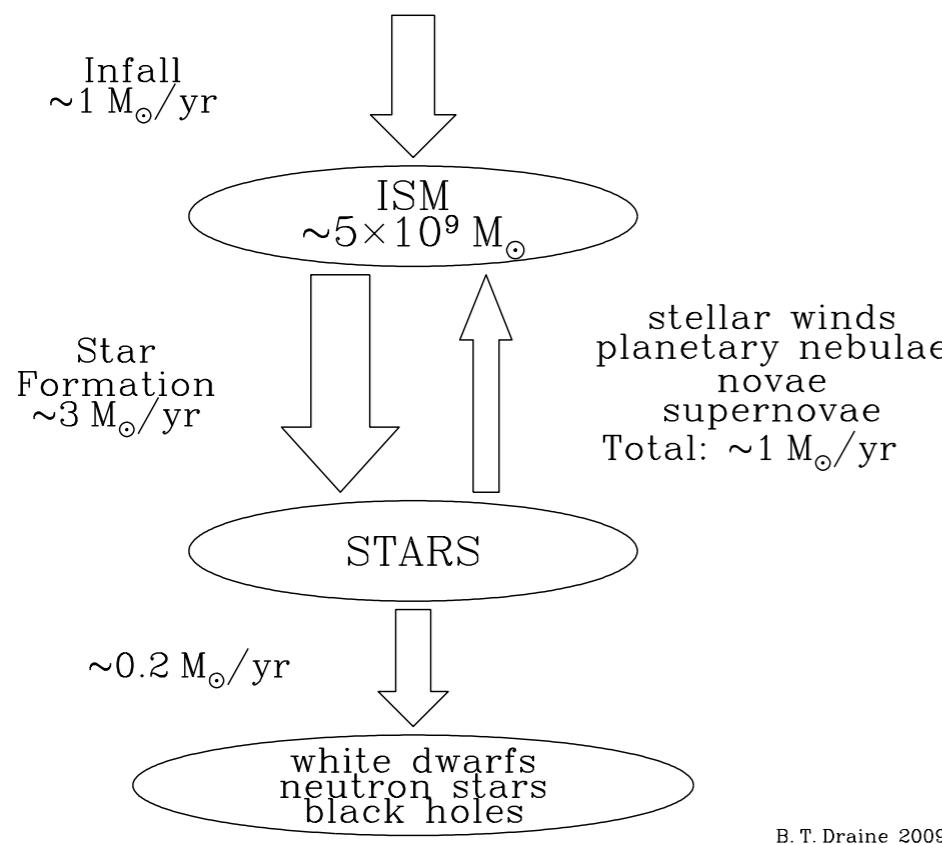


[Fig 1 in Cowie (1978, ApJ, 225, 887)]

- Survival time of random Si atoms in grain is $t_{\text{dest}} \approx (3 - 5) \times 10^8 \text{ yr}$.
(Barlow 1978; Draine & Salpeter 1979; Dwek & Scalo 1979, 1980; Jones et al. 1994, 1996).

Stardust or Growth in ISM: (1) Dust Mass Balance in the ISM

Flow of baryons in the Milky Way.



- MW: star formation is more-or-less steady for past ~ 8 Gyr (Rocha-Pinto et al. 2000):

$M > 1 M_{\odot}$ stars are dying at \sim same rates as being formed.

- $M_{\text{ISM}} \approx 5 \times 10^9 M_{\odot}$ in MW
- Sources and Sinks:

$$\begin{aligned}
 \dot{M}_{\text{ISM}} &\approx +1 M_{\odot}/\text{yr} : \text{Infall} \\
 &\approx -3 M_{\odot}/\text{yr} : \text{Star Formation} \\
 &\approx +1 M_{\odot}/\text{yr} : \text{Stellar Outflows} \\
 \hline
 \text{Net} &: -1 M_{\odot}/\text{yr}
 \end{aligned}$$

- ISM declining on timescale $M_{\text{ISM}} / |\dot{M}_{\text{ISM}}| \approx 5 \text{ Gyr}$
- Atom (or grain) in ISM incorporated in a star on timescale

$$\frac{M_{\text{ISM}}}{\text{SFR}} \approx \frac{5 \times 10^9 M_{\odot}}{3 M_{\odot}/\text{yr}} \approx 1.5 \text{ Gyr}$$

*This is the Residence Time scale of atoms (grains) in the ISM.
There is enough time that grain seeds can grow in the ISM.*

(2) Stellar Sources

Injection of Gas and Dust (“Stardust”) from Stellar Sources		
gas (M_{\odot} / yr)	dust (M_{\odot} / yr)	stellar source
0.4	0.002	Planetary Nebulae ($\sim 0.3/\text{yr}$)
0.5	0.0025	Red Giant, Red Supergiant, C star winds
0.06	< 0.0001?	OB, WR, other warm/hot star winds
0.25	0.0002?	Sne (1/100yr, $\sim 10^{-2} M_{\odot}$ dust/SN?)
0.01	0.00001	Nova (100/yr, $10^{-7} M_{\odot}$ dust/nova?)
~1.2	~0.005	All stellar sources

- O-rich stars ($\text{O/C} > 1$) produce (mainly) silicate dust (C is in CO)
- C-rich stars ($\text{C/O} > 1$) produce (mainly) carbonaceous dust and SiC (O is in CO)
C and O combine to form CO, and the remaining atoms are consumed for dust grains.

Core-collapse SNe (e.g., Cas A, SN87a)

- Cas A: silicates, etc.
 $M_{\text{dust}} \lesssim 0.14 M_{\odot}$ (Arendt et al. 2014)
 $M_{\text{dust}} \approx 0.6 M_{\odot}$ (Priestley et al. 2019)
- SN87a:
 $M_{\text{dust}} \approx 0.8 M_{\odot}$ (Matsuura et al. 2015)
- Cas A, SN87a: very young (pre-Sedov)
much ejecta still unshocked
reverse shock \rightarrow grain destruction
how much SN dust survives to mix into ISM?
some does: presolar grains in meteorites.
- SN presolar grains in meteorites rare
guess: < 10% of SN-formed dust enters ISM
- SN dust production: secondary importance today
possibly important in first $\sim 10^8$ yr of galaxy

Stardust Injection Rate

$$\dot{M}_{\text{dust}} \approx 0.005 M_{\odot}/\text{yr}$$

(3) Sputtering in Hot Gas

- Substantial grain destruction occurs in a shock of $V_s \gtrsim 220 \text{ km s}^{-1}$ and $T_s = 7 \times 10^5 \text{ K}$.
- A SN explosion ejects $\sim 10^{51} \text{ ergs}$ kinetic energy. The SNR blastwave shock-heats the ISM.
- In the energy-conserving Sedov-Taylor phase, the mass of the SNR would be

$$M_{\text{SNR}} V_s^2 \approx 10^{51} \text{ ergs}$$

$$V_s \gtrsim 220 \text{ km s}^{-1}$$

$$M_{\text{SNR}} \approx 10^3 M_\odot \left(\frac{220 \text{ km s}^{-1}}{V_s} \right)^2$$

- Detailed studies show that most dust grains materials are sputtered if $V_s \gtrsim 220 \text{ km s}^{-1}$.
- Every $\sim 10^2 \text{ yr}$, the SN blastwave destroys grains in $\sim 10^3 M_\odot$ of the ISM.
- Therefore, the grain lifetime against destruction is:

Lecture 11: The SN frequency in the Galaxy is estimated to be one event every 30-50 yr.

$$t_{\text{dest}} \approx \frac{5 \times 10^9 M_\odot}{10^3 M_\odot / 10^2 \text{ yr}} \approx 5 \times 10^8 \text{ yr}$$

- ***This lifetime is short compared to 1.5 Gyr residence time.***

MOST STARDUST WILL NOT SURVIVE

Balance of Formation and Destruction

Total dust mass in ISM: $M_{\text{dust}} \sim 0.007 \times M_{\text{ISM}} \approx 3.5 \times 10^7 M_{\odot}$

Balance:

Stardust Injection + Formation in ISM = Removal

$$\sim 0.005 M_{\odot} \text{ yr}^{-1} + \text{Formation in ISM} = \frac{\sim 3.5 \times 10^7 M_{\odot}}{\sim 5 \times 10^8 \text{ yr}}$$

$$\frac{\text{Formation in ISM}}{\text{Stardust injection}} = \frac{\sim 0.07 M_{\odot} \text{ yr}^{-1}}{\sim 0.005 M_{\odot} \text{ yr}^{-1}}$$

$$\frac{\text{Formation in ISM}}{\text{Stardust injection}} \approx 14$$

**Conclusion: Stardust accounts for only $\sim 5\text{-}10\%$ of interstellar dust.
 $\sim 90\text{-}95\%$ is grown from gas in the ISM.**