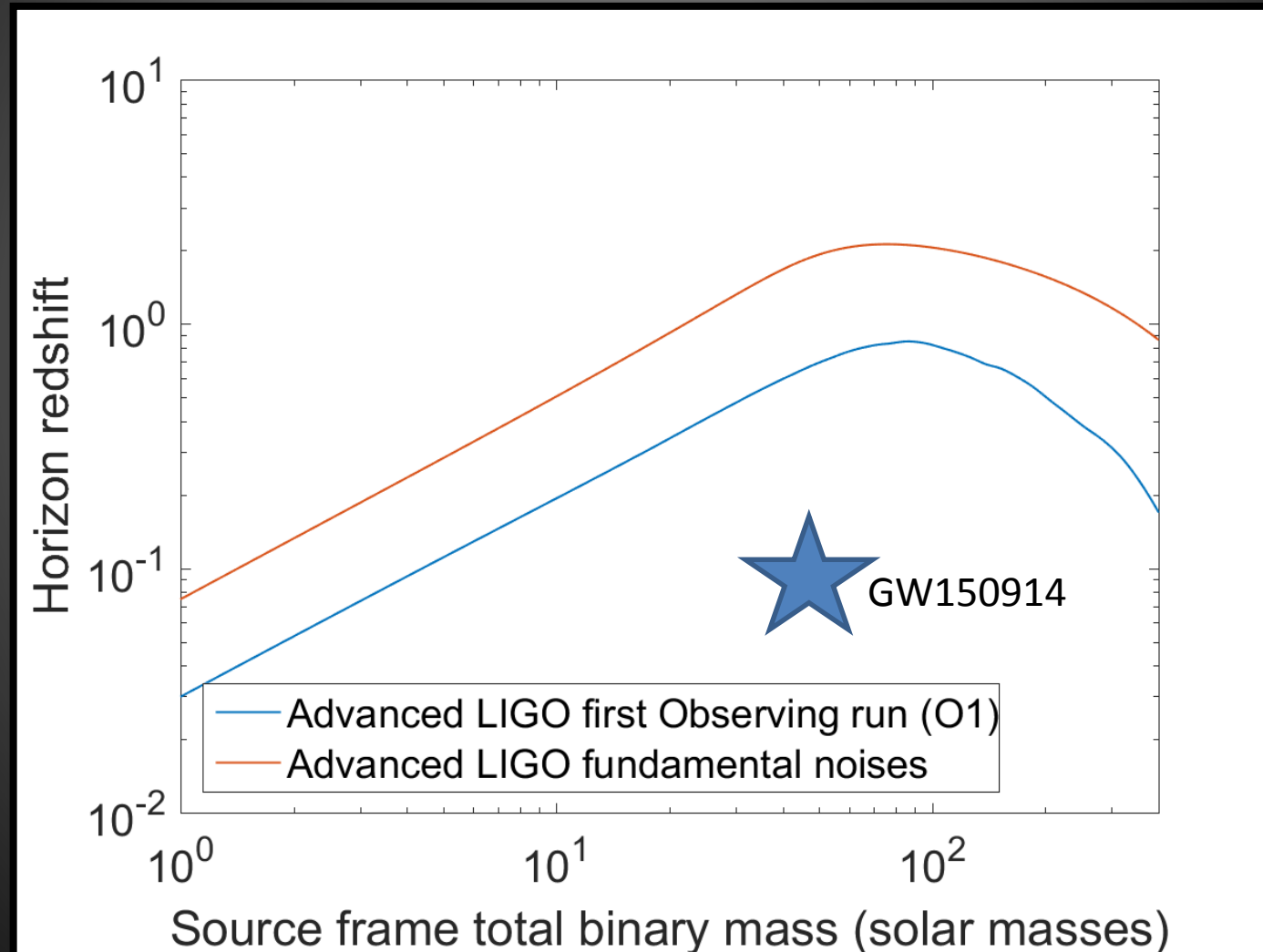


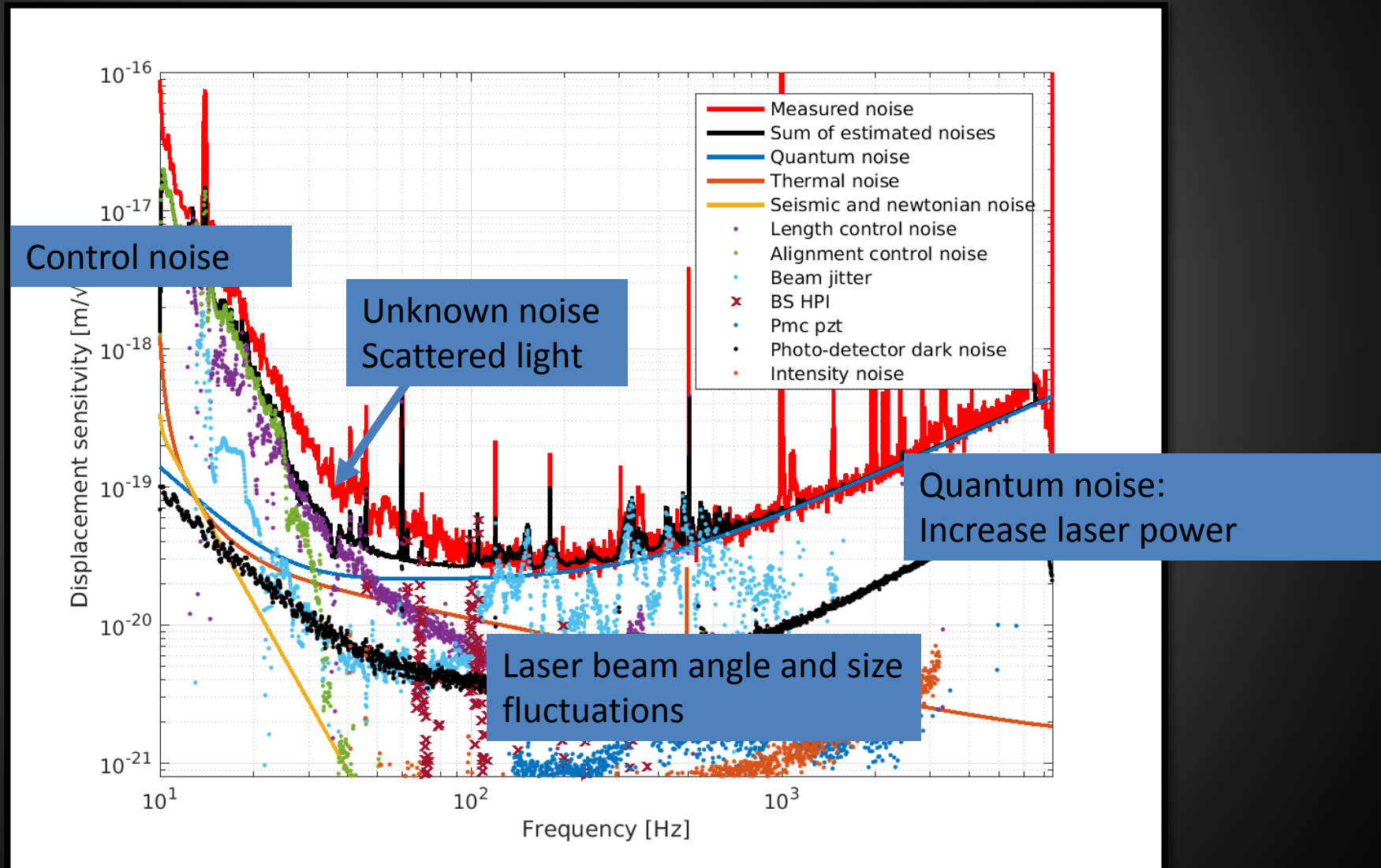
# Prospects for gravitational wave detections: update from LIGO

Sheila Dwyer for the LIGO Scientific collaboration  
LIGO DCC G1602072

# The reach of ground based gravitational wave detectors



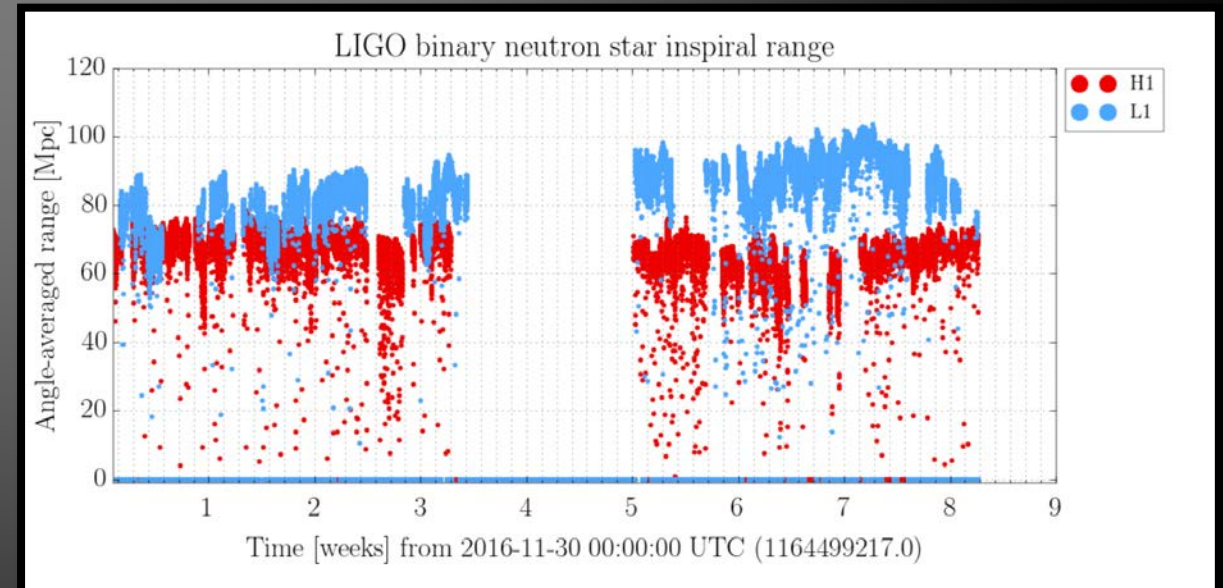
# Displacement sensitivity



# Observing break for instrument work

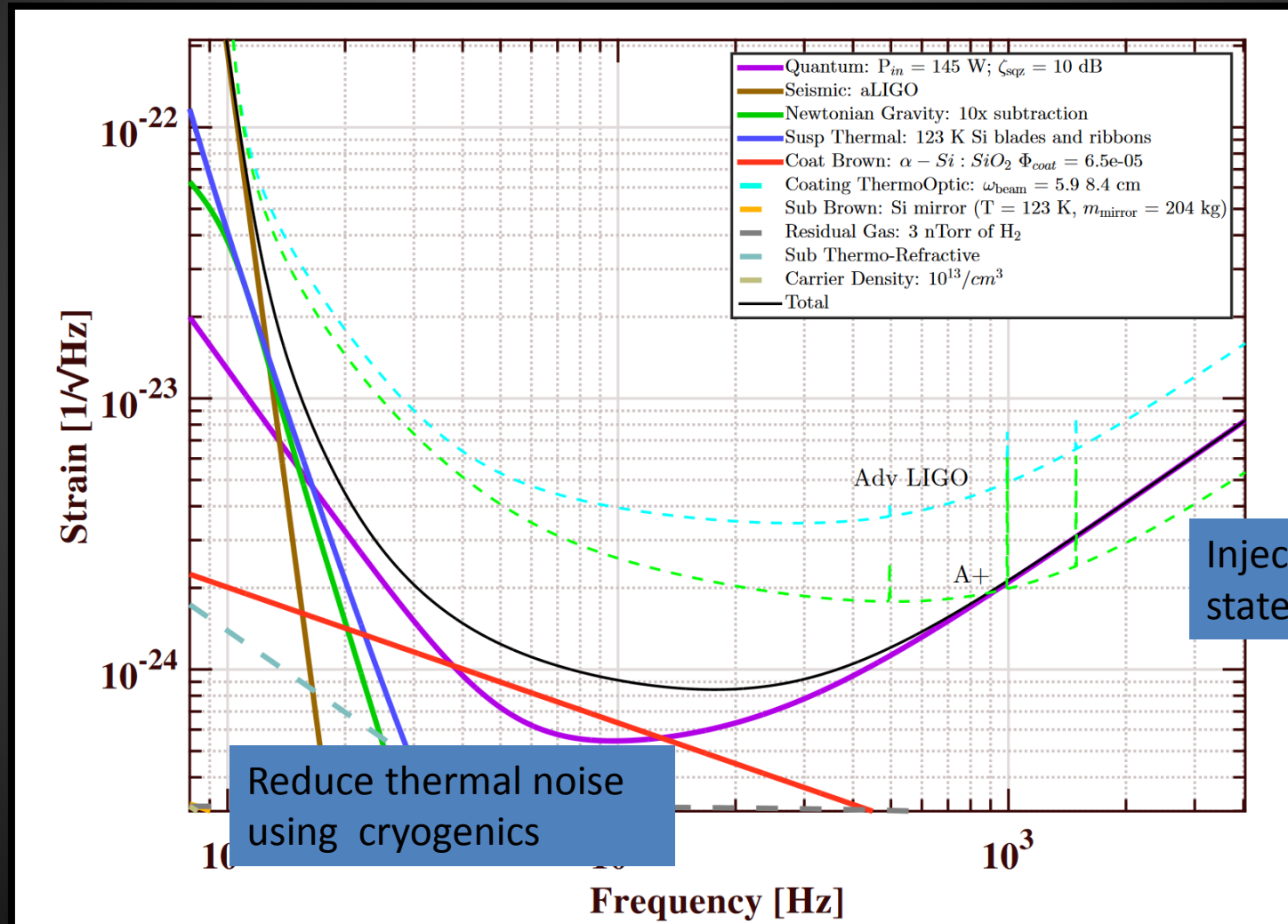
## February-November 2106

- High power operation –one of Advanced LIGO’s largest technical risks
  - Hanford turned on the high power laser, increased laser intensity, jitter, and beam size noise
  - Stable operation at 50 Watts achieved, not an improved noise performance
- Livingston-
  - In vacuum hardware upgraded
  - Reduced scattered light noise
- Started our second observing run in November





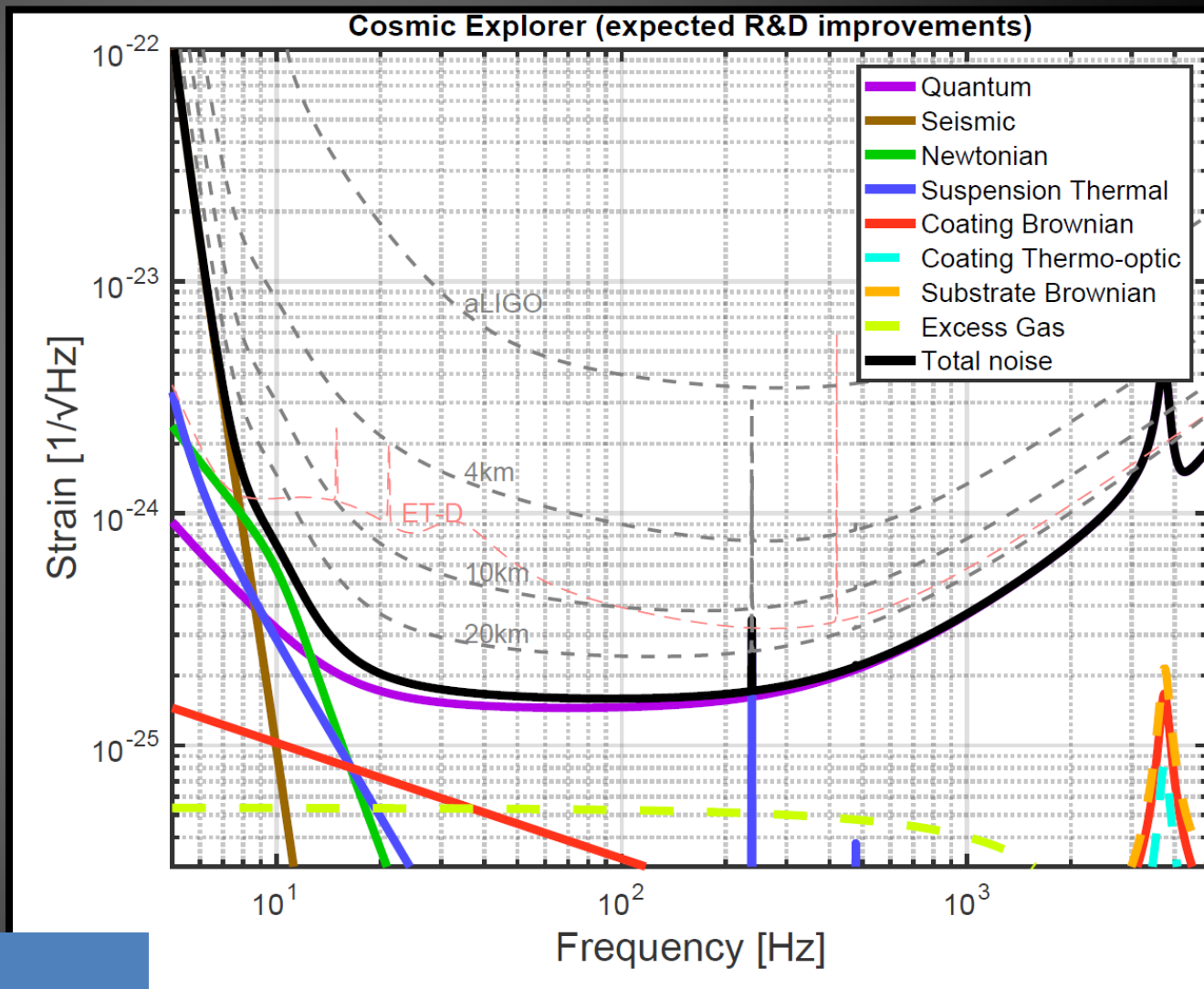
# Ultimate performance in current facilities: LIGO Voyager



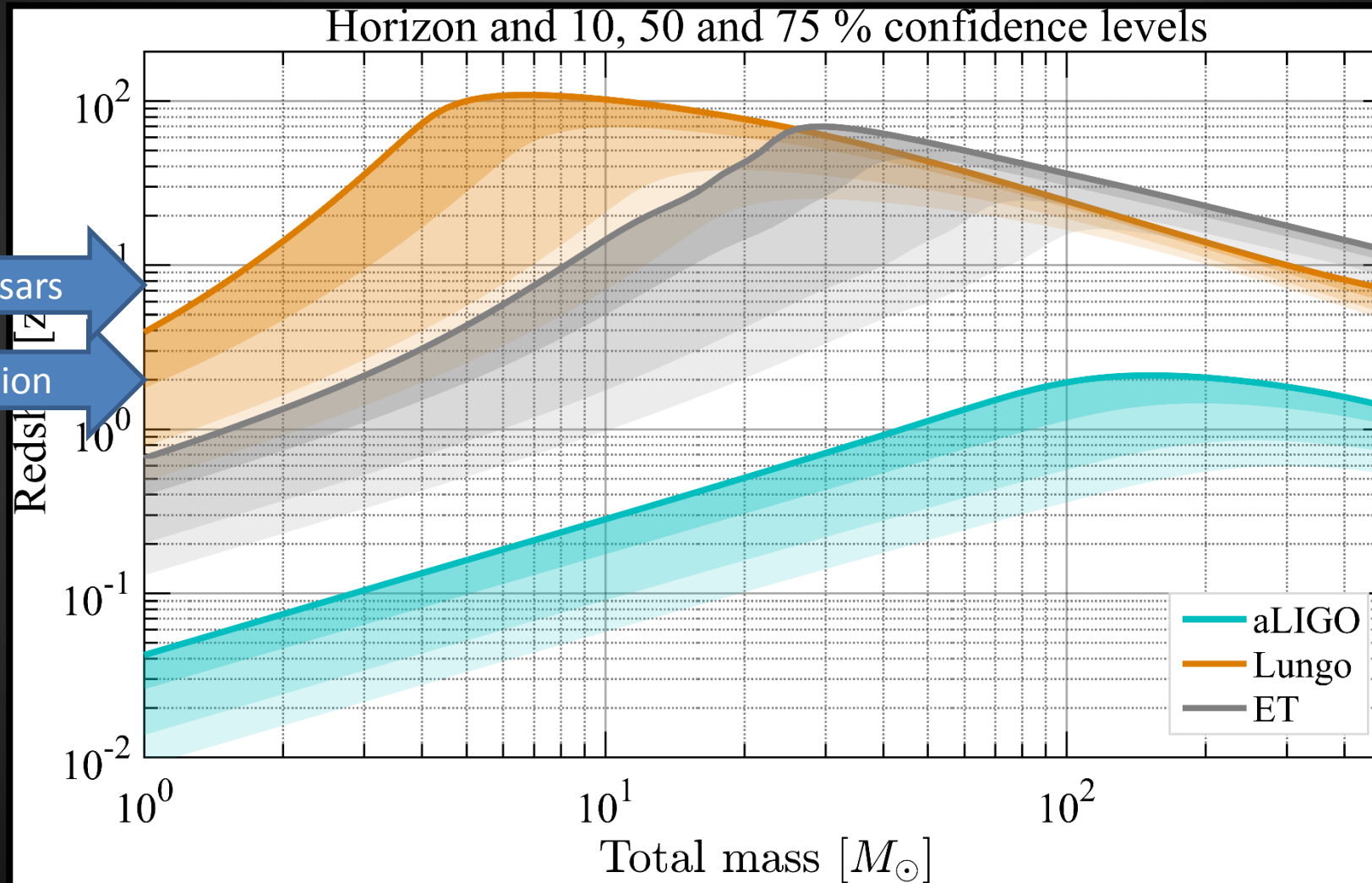
# Combining improved technology and new facility

$$\delta h = \delta L / 2L$$

Sensitivity to displacement noises  
are reduced as arm length increases



# Extending the reach of GW observatories



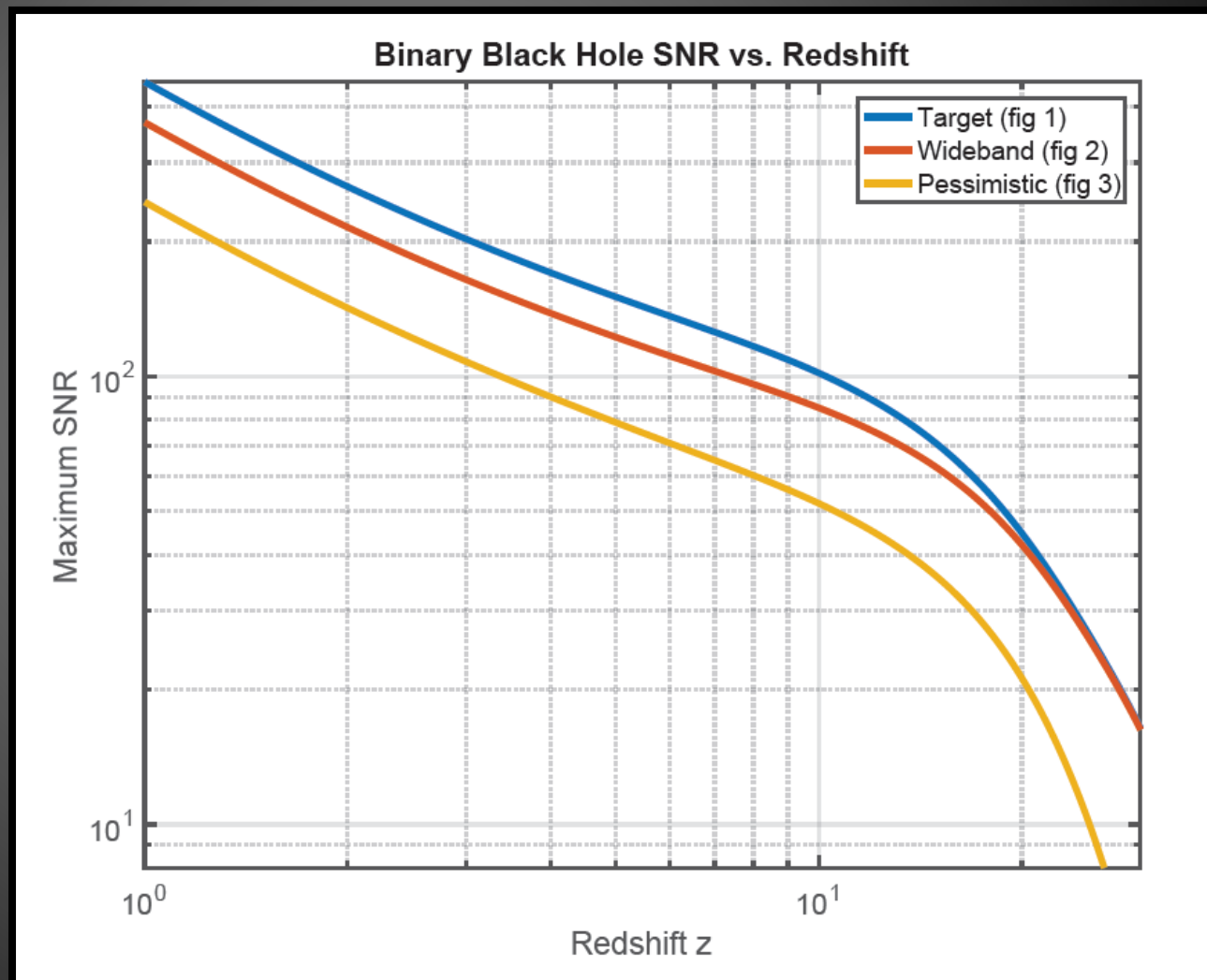
# Summary

- You can expect a trickle of detections as we work towards aLIGO design sensitivity over next few years
- It is possible to observe black hole binaries from the entire history of star formation, with high SNR

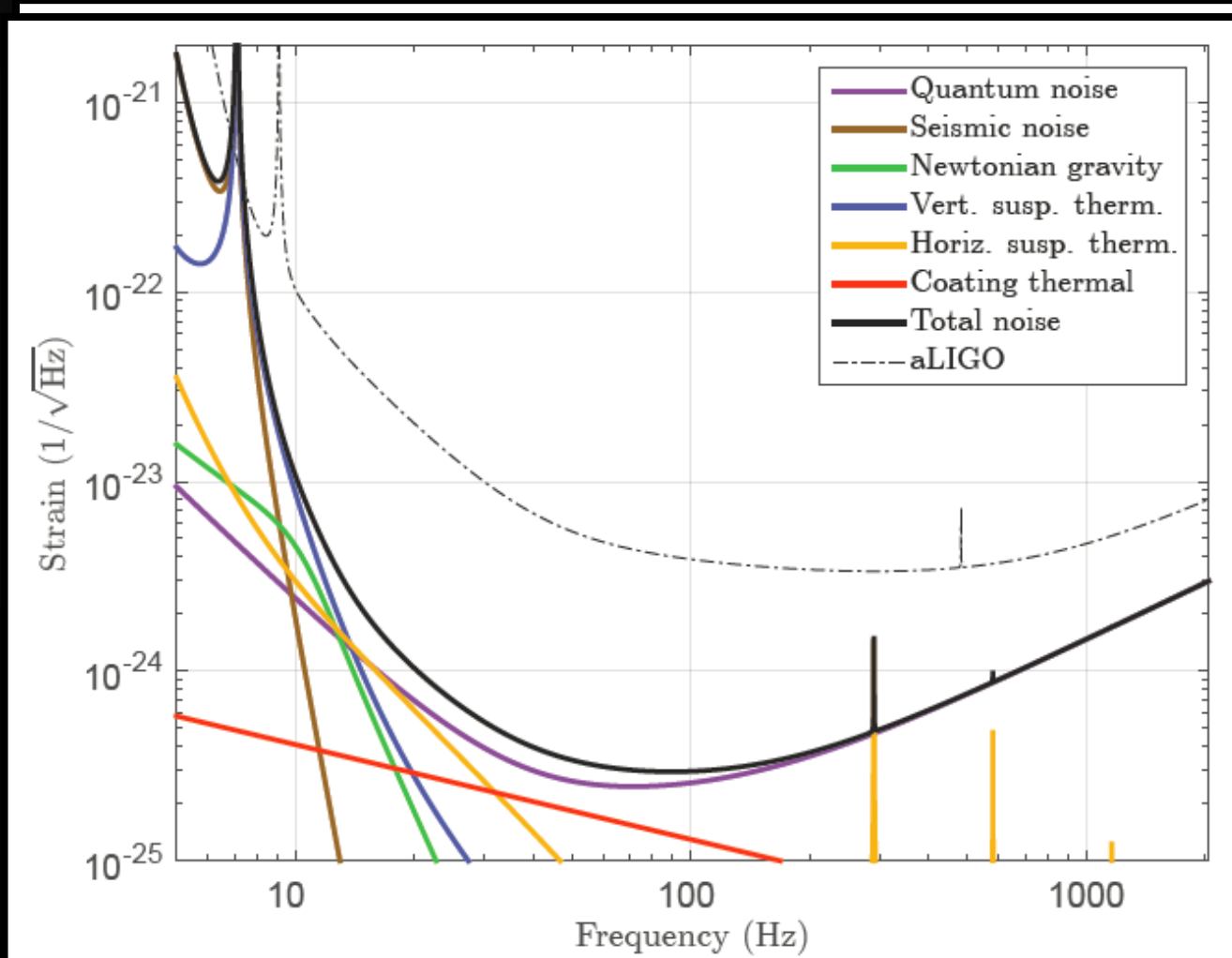
Sensitivity curves publicly available:

<https://dcc.ligo.org/LIGO-P1600143>





# Ultimate displacement noise reduction technique: Long arms



$$\delta h = \delta L / 2L$$

Sensitivity to displacement noises are reduced as arm length increases  
Even if **most** noise reduction techniques don't work out, 40km arms can improve over advanced LIGO sensitivity by a factor of  $\sim 10$

Dwyer *et al* 2014 Phys. Rev. D **91**, 082001

# Towards design sensitivity

