

Goodman HTS Pipeline User Manual

version 0.3

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Introduction

This is the User Manual for the *Goodman Spectroscopic Data Reduction Pipeline*. It provides an overview of the pipeline's main features, instructions on its use and how to run it on our dedicated *SOAR Data Reduction Server*, and installation instructions for those who wish to run it on their own computers.

Overview

The Goodman Spectroscopic Data Reduction Pipeline - GOODSPEC - is a Python-based package for producing science-ready, wavelength-calibrated, 1-D spectra. The goal of **goodspec** is to provide SOAR users with an easy to use, very well documented software for reducing spectra obtained with the Goodman spectrograph. Though the current implementation assumes offline data reduction, our aim is to provide the capability to run it in real time, so 1-D wavelength calibrated spectra can be produced shortly after the shutter closes.

The pipeline is primarily intended to be run on a data reduction dedicated computer. Instructions for running the software are provided in the [Running Pipeline](#) section of this guide. The Goodman Spectroscopic Data Reduction Pipeline project is hosted at GitHub at [it's GitHub Repository](#).

Currently the pipeline is separated into two main components. The initial processing is done by `redccd`, which trims the images, and carries out bias and flat corrections. The spectroscopic processing is done by `redspec` and carries out the following steps:

- Identifies multiple targets (spectra of more than one object in the slit)
- Trace the spectra
- Extract the spectra
- Estimate and subtract background
- Saves extracted (1D) spectrum, without wavelength calibration.
- Find the wavelength solution. Defaults to automatic wavelength solution, but can be done interactively
- Linearize data (resample)
- Write wavelength solution to FITS header
- Create a new file for the wavelength calibrated 1D spectrum

Features

- Self-contained, full data reduction package for the most commonly used spectroscopic setups with Goodman. Given the almost limitless number of possible configurations available with the Goodman instrument, only the most popular configurations will be supported, though we will try to add as many modes as possible.
- Python based, using existing Astropy libraries as much as feasible.
- Extensively documented, using general coding standards: PEP8 – Style Guide, PEP257 – Docstrings Convention (in-code documentation) – Google Style
- Multiplatform compatibility (tested on Linux Ubuntu, CentOS and MacOSX).
- Modular design. Could be used as a library within other Python applications.

Ways to run the pipeline

There are two ways to use the pipeline.

1. **Run it directly on a SOAR data reduction server** that you can access using VNC.
2. **Download and install the pipeline** (go to the [Installing Pipeline](#) section of this manual). Though we will try our best to provide answers to quick and simple installation issues, we cannot provide general installation support.

What the pipeline does not do

- In its current version the pipeline does not perform combination of individual spectra. If you obtained several individual exposures of the same object, they will be output as separate 1-D, wavelength-calibrated spectra.
- It does not combine lamps. If there are more than one comparison lamp associated with a spectrum they will be processed and saved separately.
- There is yet no flux calibration. We are working on a module that will do this. But this will be left for a later release.
- This pipeline does not evaluate nor select data by quality. It will simply try to run using all existing files. **Make sure you only have good data in the folder that will be reduced.**

General Considerations on using the pipeline

The Goodman Spectroscopic Pipeline is meant to work as a single package. However, the full process is split in two separate modules: `redccd` and `redspec`. The first does the basic 2D image reduction, applying bias, flat field corrections, and cosmic ray removal. The second module, `redspec`, takes the corrected 2D images output by `redccd` and produces wavelength-calibrated 1D spectra.

The pipeline is run from the command line in a terminal window. Each module is run separately, first `redccd` followed by `redspec`, however, you could run both sequentially from e.g. a shell script, just make sure you move to the the right directory.

In order to make things easier you should organize your data:

1. Make sure all the data in your folder corresponds to the same binning, readout mode, region of interest (ROI), and grating/wavelength mode combination.
2. You should have bias, flats (quartz or dome flats), and the appropriate comparison lamps. Other files like acquisition images, slit images and focus images should be deleted.
3. Do not mix dome flats with quartz lamp flats. As an example, suppose you took both quartz lamps and dome flats for your targets. You could create two folders, one with the science data and the dome flats, and another with the same science data and the quartz lamps. Then, if you run the pipeline in each folder you can compare the results and decide which type of flat works best for my particular case.

Command line arguments

For a list of the options and command line arguments type `--help` argument:

For `redccd`

```
usage: redccd [-h] [--auto-clean] [--cosmic <method>] [--combine]
              [--dcr-par-dir <dcr.par_directory>] [--debug]
              [--flat-normalize <normalization_method>]
              [--flat-norm-order <order>] [--ignore-bias] [--ignore-flats]
              [--keep-cosmic-files] [--raw-path <raw_path>]
              [--red-path <red_path>] [--saturation <value>]

Goodman CCD Reduction - CCD reductions for Goodman spectroscopic data.
Pipeline Version: 1.0b10

optional arguments:
  -h, --help            show this help message and exit
  --auto-clean          Automatically clean reduced data directory
  --cosmic <method>    Clean cosmic rays from all data. Options are: 'dcr',
                       'lacosmic' or 'none'. Default is 'dcr'. See manual for
                       full description of dcr.
  --combine             Combine compatible data (experimental)
  --dcr-par-dir <dcr.par_directory>
                       Directory of default dcr.par file
  --debug               Show detailed information of the process.
  --flat-normalize <normalization_method>
                       Choose a method to normalize the master flat
                       forspectroscopy. Choices are: mean, simple (model) and
                       full (fits model to each line).
  --flat-norm-order <order>
                       Defines the order of the model to be fitted. Default
                       to 15
  --ignore-bias         Ignore bias correction
  --ignore-flats        Ignore flat field correction
  --keep-cosmic-files   After cleaning cosmic rays with dcr, do not remove the
                       input file and the cosmic rays file.
  --raw-path <raw_path>
                       Path to raw data.
```

General Considerations on using the pipeline

```
--red-path <red_path>
                  Path to reduced data.
--saturation <value> Saturation limit. Default to 65.000 ADU (counts)
```

And for `redspec`

```
usage: redspec [-h] [--data-path <Source Path>]
                [--proc-path <Destination Path>]
                [--search-pattern <Search Pattern>]
                [--output-prefix <Out Prefix>] [--extraction <Extraction Type>]
                [--reference-files <Reference Dir>] [--interactive] [--debug]
                [--max-targets <max targets>] [--save-plots] [--plot-results]

Extracts goodman spectra and does automatic wavelength calibration. Pipeline
Version: 1.0b10

optional arguments:
-h, --help            show this help message and exit
--data-path <Source Path>
                  Path for location of raw data. Default <./>
--proc-path <Destination Path>
                  Path for destination of processed data. Default <./>
--search-pattern <Search Pattern>
                  Pattern for matching the goodman's reduced data.
--output-prefix <Out Prefix>
                  Prefix to add to calibrated spectrum.
--extraction <Extraction Type>
                  Only fractional pixel extraction is implemented.
--reference-files <Reference Dir>
                  Directory of Reference files location
--interactive         Interactive wavelength solution. Disabled by default.
--debug              Debugging Mode
--max-targets <max targets>
                  Maximum number of targets to be found in a single
                  image. Default 3
--save-plots          Save all plots in a directory
--plot-results        Show wavelength calibrated spectrum at the end.
```

Lists of Reference Lamps Available

The automatic wavelength calibration relies on having previously calibrated reference lamps obtained in the same configuration or mode. It is also important that the lamp names are correct, for instance HgAr is quite different than HgArNe. For interactive wavelength calibration, reference lamps are used as a visual aid only. It lets you find the matching laboratory lines values that will be used to fit a pixel to wavelength relation that we call *Wavelength Solution*. The list of lamps is the following.

| Grating | Mode | Filter | Lamp |
|---------|------|--------|--------|
| 400 | M1 | None | HgAr |
| 400 | M1 | None | HgArNe |
| 400 | M2 | GG455 | Ar |
| 400 | M2 | GG455 | Ne |
| 400 | M2 | GG455 | HgAr |
| 400 | M2 | GG455 | HgArNe |
| 400 | M2 | GG455 | CuHeAr |
| 400 | M2 | GG455 | FeHeAr |

Important

More lamps will be made public shortly.

Adding new reference lamps

It is possible to add new lamps, you just need a raw lamp that meets the following specifications with respect to your science project:

- Same instrument configuration or mode
- Same Grating
- Same order blocking filter if present
- Same binning
- Same lamp/combination that you use in your observations
- Smallest slit possible. Equal is OK too.

See [Creating Reference Lamps](#) for instructions on how to proceed. Now you have two options, identify the system folder where the lamps that come with the package are saved and simply put it there or put it in another directory and use the argument --reference-files

```
redspec --reference-files /path/to/ref-lamp-location
```

If you require that your lamps are made available as a part of the package you can contact storres [at] ctio [dot] noao [dot] edu and we will handle your request.

Headers Requirements

Goodman HTS spectra have small non-linearities on their wavelength solutions. They are small but big enough that they MUST be taken into account.

It was necessary to implement a custom way of storing non-linear wavelength solutions that at the same time allowed for keeping data as *untouched* as possible. The main reason is that linearizing the reference lamps made harder to track down those non-linearities on the new data being calibrated and also; The documentation on how to write non-linear solution to a FITS header is not good, besides it appears that nobody is trying to improve it neither trying to implement it. Below I compile a list of required keywords for comparison lamps if they want to be used as

Lists of Reference Lamps Available

reference lamps. The full list of keywords is listed under [Headers](#).

General Custom Keywords:

Every image processed with the *Goodman Spectroscopic Pipeline* will have the general custom keywords. The one required for a reference lamp is the following:

```
GSP_FNAM = file-name.fits // Current file name
```

Record of line centers in Pixel and Angstrom:

Every line detected in the reference lamp is recorded both in its pixel value and later (most likely entered by hand) in angstrom value. The root string is GSP_P followed by a zero-padded three digit sequential number (001, 002, etc). For instance.

```
GSP_P001= 499.5377036976768 / Line location in pixel value  
GSP_P002= 810.5548319623747 / Line location in pixel value  
GSP_P003= 831.6984711087946 / Line location in pixel value
```

The equivalent values in angstrom are then recorded with the root string GSP_A and the same numerical pattern as before.

```
GSP_A001= 5460.75 / Line location in angstrom value  
GSP_A002= 5769.61 / Line location in angstrom value  
GSP_A003= 5790.67 / Line location in angstrom value
```

GSP_P001 and GSP_A001 are a match. If any of the angstrom value entries have a value of 0 (default value) the equivalent pixel entry is ignored.

Important

Those keywords are used to calculate the mathematical fit of the wavelength solution and are not used on normal operation. Our philosophy here is that the line identification has to be done only once and then the model can be fitted several time, actually you can try several models if you want.

Non-linear wavelength solution:

The method for recording the non-linear wavelength solution is actually very simple. It requires: GSP_FUNC which stores a string with the name of the mathematical model from `astropy.modeling.models`. GSP_ORDR stores the order or degree of the model. GSP_NPIX stores the number of pixels in the spectral axis. Then there is N+1 parameter keywords where N is the order of the model defined by GSP_ORDR. The root string of the keyword is GSP_C and the rest is a zero-padded three digit number starting on zero to N. See the example below.

```
GSP_FUNC= Chebyshev1D / Mathematical model of non-linearized data  
GSP_ORDR= 3 / Mathematical model order  
GSP_NPIX= 4060 / Number of Pixels  
GSP_C000= 4963.910057577853 / Value of parameter c0  
GSP_C001= 0.9943952599223119 / Value of parameter c1  
GSP_C002= 5.59241584012648e-08 / Value of parameter c2  
GSP_C003= -1.2283411678846e-10 / Value of parameter c3
```

Warning

This method has been developed and tested to write correctly polynomial-like models. And ONLY reads Chebyshev1D models. Other models will just be ignored. More development will be done based on request, suggestions or needs.

Creating Your Own Reference Lamps

If you use a Custom mode you will need a Custom reference lamp too. You are welcome to try to do it, just make sure the [Header Requirements](#) are met. If not we might be able to help, depending on time availability, it will help a lot if you have the lines identified already.

Important

We are not ready to offer a simple tool to construct the reference lamps.

Cosmic Rays Removal

The default cosmic ray removal tool has not been written in Python (neither for Python). However it has been successfully integrated into the Pipeline's workflow using subprocess. Despite the integration being successful, the fact that is a c program, designed to be run *standalone* from a terminal, it comes with several challenges in terms of flexibility. The addition of keywords has been dealt with by creating a modified version of the software (see [Install DCR](#)). But one of the trickiest part, and probably even more important, is the *parameter configuration*. It uses an external ASCII file and that file HAS TO be named exactly `dcr.par`, also that file has to be located on the folder where the program is called from (usually the data location folder).

The default parameters work very good for the default configuration of the Goodman HT Spectrograph (Spectroscopic 1x1 for instance), but if you have binning 2x2 you will need to tune the parameters.

Since you can't parse the parameters as arguments the only option is to change the values on the `dcr.par` file.

If you use several configurations your best option is to store custom `dcr.par` files and use the argument `--dcr-par-dir <dcr-par-location>` to tell the pipeline where to look for a `dcr.par` file, with the limitation that is has to be called `dcr.par`.

The pipeline will copy the file to the folder where the data is stored.

Headers

The pipeline adds several keywords to keep track of the process and in general for keeping important information available. In the following table is a description of all Goodman Spectroscopic Pipeline keywords added, though not all of them are added to all the images.

General Purpose Keywords

These keywords are used for record purpose, except for `GSP_FNAM` which is used to keep track of the file name.

| Keyword | Purpose |
|-----------------------|--|
| <code>GSP_VERS</code> | Pipeline version. |
| <code>GSP_ONAM</code> | Original file name, first read. |
| <code>GSP_PNAM</code> | Parent file name. |
| <code>GSP_FNAM</code> | Current file name. |
| <code>GSP_PATH</code> | Path from where the file was read. |
| <code>GSP_TECH</code> | Observing technique. Imaging or Spectroscopy. |
| <code>GSP_DATE</code> | Date of processing. |
| <code>GSP_OVER</code> | Overscan region. |
| <code>GSP_TRIM</code> | Trim section. |
| <code>GSP_SLIT</code> | Slit trim section. From slit-illuminated area. |
| <code>GSP_BIAS</code> | Master bias file used. |

| | |
|----------|--|
| GSP_FLAT | Master flat file used. |
| GSP_NORM | Master flat normalization method. |
| GSP_COSM | Cosmic ray rejection method. |
| GSP_WRMS | Wavelength solution RMS Error. |
| GSP_WPOI | Number of points used to calculate RMS Error. |
| GSP_WREJ | Number of points rejected from RMS Error Calculation. |
| GSP_DCRR | Reference paper for DCR software (cosmic ray rejection). |

Non-linear wavelength solution

Since writing non-linear wavelength solutions to the headers using the FITS standard (reference) is extremely complex and not necessarily well documented. We came up with the solution of simply describing the mathematical model from `astropy.modeling.models`. This allows for maintaining the data *untouched* while keeping a reliable description of the wavelength solution.

The way it is currently implemented will work for writing for any polynomial kind of model. Reading is implemented only for `Chebyshev1D` which is the model by default.

| Keyword | Purpose |
|----------|--|
| GSP_FUNC | Name of mathematical model. <code>astropy.modeling.models</code> |
| GSP_ORDR | Order of the model used. |
| GSP_NPIX | Number of pixels. |
| GSP_C000 | Value of parameter c_0 . |
| GSP_C001 | Value of parameter c_1 . |
| GSP_C002 | Value of parameter c_2 . This goes on depending the order. |

Combined Images

Every image used in a combination of images is recorded in the header of the resulting one. The order does not have importance but most likely the header of the first one will be used

| Keyword | Purpose |
|----------|---------------------------------------|
| GSP_IC01 | First image used to create combined. |
| GSP_IC02 | Second image used to create combined. |

Detected lines

The *reference lamp library* maintains the lamps non-linearized and also they get a record of the pixel value and the equivalent in angstrom. In the following table a three-line lamp is shown.

| Keyword | Purpose |
|----------|--|
| GSP_P001 | Pixel value for the first line detected. |
| GSP_P002 | Pixel value for the second line detected. |
| GSP_P003 | Pixel value for the third line detected. |
| GSP_A001 | Angstrom value for the first line detected. |
| GSP_A002 | Angstrom value for the second line detected. |
| GSP_A003 | Angstrom value for the third line detected. |

Running the pipeline in the SOAR data reduction computer

The Goodman Spectroscopic Data Reduction Pipeline has been installed on a dedicated computer at SOAR. The procedure is to open a VNC session, for which you need to be connected to the SOAR VPN. The credentials for the VPN are the same you used for your observing run, provided by your *Support Scientist*, who will also give you the information for the data reduction computer VNC connection.

Note

IRAF is available in all three data servers. Running `iraf` will open an `xgterm` and `ds9` windows. `iraf-only` will not open `ds9`

Establish a VNC connection

Separately, you should receive a server hostname, IP, display number and VNC-password. If you don't you can ask for it. We have decided to use a similar organization of vnc displays as for `soaric7`:

VNC display number and working folder assigned to each partner.

| Display | Partner/Institution | Folder |
|---------|---------------------|---------------------------|
| :1 | NOAO | /home/goodman/data/NOAO |
| :2 | Brazil | /home/goodman/data/BRAZIL |
| :3 | UNC | /home/goodman/data/UNC |
| :4 | MSU | /home/goodman/data/MSU |
| :5 | Chile | /home/goodman/data/CHILE |

For the rest of this tutorial we will assume your host name is `vnc-server` the display number is 1 and your password is `password`. Though we recommend using RealVNC, most other VNC clients will work fine (e.g., Remmina in Linux). For GNU/Linux and Mac OSX machines we suggest the RealVNC Viewer client. For Windows machines, we suggest either the RealVNC Viewer client or the UltraVNC viewer client. We also know that Vinagre and vncviewer on GNU/Linux work fine.

VNC from the Terminal

Open a terminal, and assuming you have installed `vncviewer`.

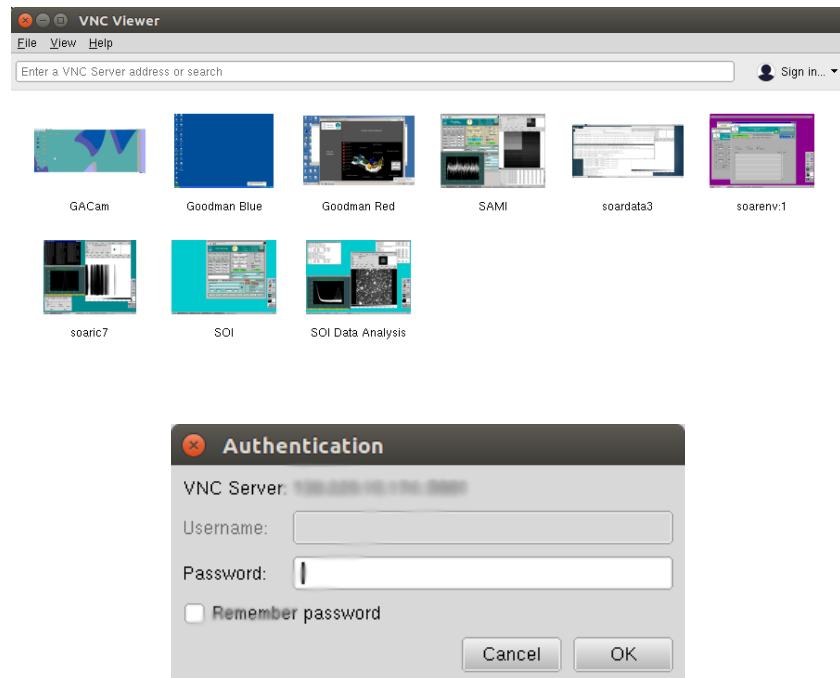
```
vncviewer vnc-server:1
```

You will be asked to type in the `password` provided.

VNC using a Graphical Client

Using a graphical VNC client is quite similar and intuitive

Dealing with Virtual Environments



In this case the *IP address* was used, which is equivalent and sometimes better.

Dealing with Virtual Environments

The Goodman Spectroscopic Pipeline uses virtual environments since they allow for easy portability and also give a higher level of safety for the host system. If you know nothing about them, we encourage you to do a quick search to get at least the basics.

All terminals using bash are configured to start with the appropriate virtual environment activated by default. The default virtual environment is called `astroconda`.

If for any reason it is not activated you will need to activate it in order to use the pipeline.

```
source activate astroconda
```

Or if for some reason you need to deactivate it, with the virtual environment activated you do:

```
source deactivate
```

Running the Pipeline

1. Open a Terminal

2. Go to /home/goodman/data

```
cd /home/goodman/data
```

3. Here you have a workspace to put your data according to your institution.

```
goodman@soardata3:~/data
File Edit View Search Terminal Help
[goodman@soardata3 ~]$ pwd
/home/goodman
[goodman@soardata3 ~]$ cd data
[goodman@soardata3 data]$ ls
BRAZIL CHILE MSU NOAO test UNC
[goodman@soardata3 data]$ df -h
Filesystem      Size  Used Avail Use% Mounted on
/dev/nvme0n1p1    24G   11G   12G  49% /
devtmpfs        16G     0   16G   0% /dev
tmpfs          16G  172K   16G   1% /dev/shm
tmpfs          16G   18M   16G   1% /run
tmpfs          16G     0   16G   0% /sys/fs/cgroup
/dev/nvme0n1p2   211G   61M  201G  1% /home_local
/dev/md127      5.5T   42G  5.4T  1% /home_
tmpfs          3.2G   20K   3.2G   1% /run/user/8143
tmpfs          3.2G     0   3.2G   0% /run/user/8142
[goodman@soardata3 data]$
```

4. Create a data folder inside your workspace.

```
cd NOAO
```

```
mkdir 2017-07-05
```

```
cd 2017-07-05
```

5. Copy your data from Goodman Computer

```
scp observer@soaric7:/home3/observer/GODMAN_DATA/NOAO/2017-07-05/ ./
```

6. Make sure you have a full data set. At this point your observing logs will become very useful, eliminate focus sequence, aquisition exposure and any other file present that will not be needed for the processing. The following list summarizes the kind of data that you need to fully process your data.

- BIAS: Bias
- FLAT: Flats
- COMP: Comparison Lamps
- OBJECT: Science Frames

Also make sure your data has the same *readout speed*, *binning*, and *ROI*. If you used different configurations during the same night, we recommend you to set up a separate folder for each.

7. Run redccd:

If you are running `redccd` for the first time you can use `redccd` alone but if it's a second or third time you will need to use `--auto-clean` which is a built-in protection for your data, in case you don't want to delete what has been done. Also you might want to consider `--saturation <new value>` to change the saturation level if you get all your flats rejected due to saturation. Sometimes there is a hot column at the end that produced very high values.

```
redccd --auto-clean
```

In case you want to use `--saturation` here is an example:

```
redccd --auto-clean --saturation 70000
```

This changes the saturation level to 70000 ADU in this context the saturation value works as a threshold for rejecting images and it varies from one instrument configuration to another.

Running the Pipeline

By default, `redccd` puts reduced data in a subdirectory `RED`, you can provide a different one by using `--red-path`.

An image `image_file.fits` that has been fully (and properly) processed should have the new name (including the reduced data folder `RED`):

```
RED/cfzsto_image_file.fits
```

8. Run `redspec`:

All data processed with the `redccd` (previous step) will be saved in a separate folder this was intentionally made this way in order to avoid overwriting your original data.

First, move to the folder were the data was stored.

```
cd RED
```

By default `redspec` will search for images with the prefix `cfzsto`, in case you have produced a different prefix you can change it by using `--search-pattern`

You can just run `redspec` in case everything is the default but if this is the first time you run the pipeline we suggest:

```
redspec --plot-results
```

In that way two important plots will be shown full screen, the comparison lamp fitted to a reference comparison lamp and some values for the wavelength solution fit and the extracted spectrum plotted with the wavelength solution.

Before the wavelength solution is calculated, the extracted spectrum (1D already) is saved with an `e` as prefix. The final image has a `w` added to the start of the name, following the above example your final 1D and wavelength calibrated image will be named:

```
wecfzsto_image_file.fits
```

9. Finally, review the results. Below is a table with the definition of all letters used in the construction of the prefix.

The meaning of every letter in `wecfzsto` is summarized in the following table:

| Letter | Meaning |
|--------|---|
| w | Wavelength calibrated. |
| e | Extracted spectrum, 1D, no wavelength solution. |
| c | Cosmic ray cleaned or mask created depending on the method. |
| f | Flat corrected. |
| z | Zero or Bias corrected. |
| s | Slit trimmed, trims off the non-illuminated sections of the detector. |
| t | Initial Image trimming. |
| o | Overscan corrected. |

Using the Interactive Mode

Warning

The automatic wavelength calibration method has become every time more reliable therefore is more likely that you will find bugs using the interactive mode.

If you are not convinced with the *automatic wavelength solution* you can try to improve it by using the *Interactive Mode*. Some of the key features are:

- Manually match spectroscopic lines.
- Zoom in to get a better reference.

Running the Pipeline

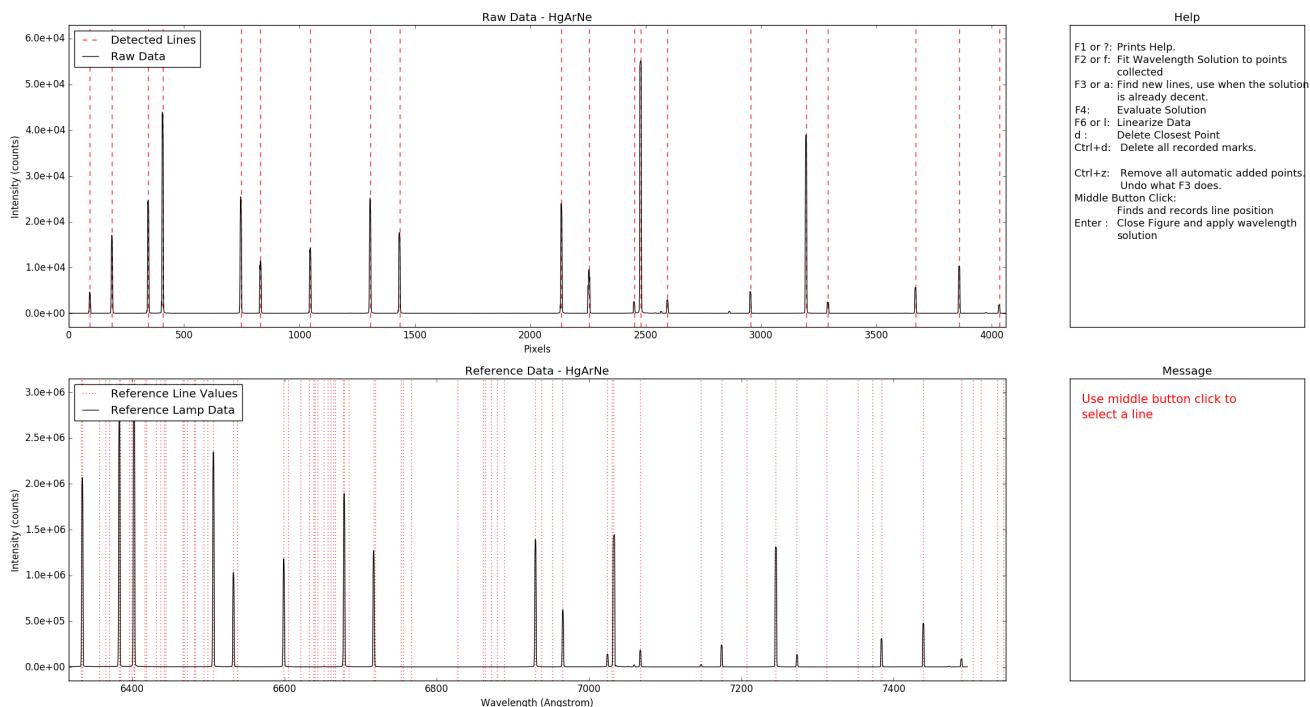
- Evaluate the quality of your solution as you progress.
- Uses a visual reference for matching the lines
- Uses laboratory values for reference lines

Right now it uses `matplotlib.pyplot` as the underlying tool to enable interaction.

Below you will see a sequence as well as a description of the procedure to use the interactive mode.

1. The interactive screen has four subplots. In the *Upper main panel* you have the raw comparison lamp, i.e. your comparison lamp 1-D-extracted but without wavelength solution, the dashed red lines represent the lines found by the pipeline. In the *Lower main panel* you have the reference lamp. This is a previously calibrated lamp that is distributed with the package (also you can use your own). In this case the red lines are the laboratory values obtained from the NIST Atomic Spectra Database site (https://physics.nist.gov/PhysRefData/ASD/lines_form.html). The *Upper right panel* is a static help, intended to give you a quick and easy-to-reach help. Finally, the *Lower right panel* is an information window, in fact it serves three main purposes:

- Display messages and warnings.
- Shows you a zoomed line.
- Displays the wavelength solution quality information.



2. The way you interact is by using your mouse and keyboard. The following table describes all the available keyboard commands with their respective keys and/or buttons, and a summary of their functionality.

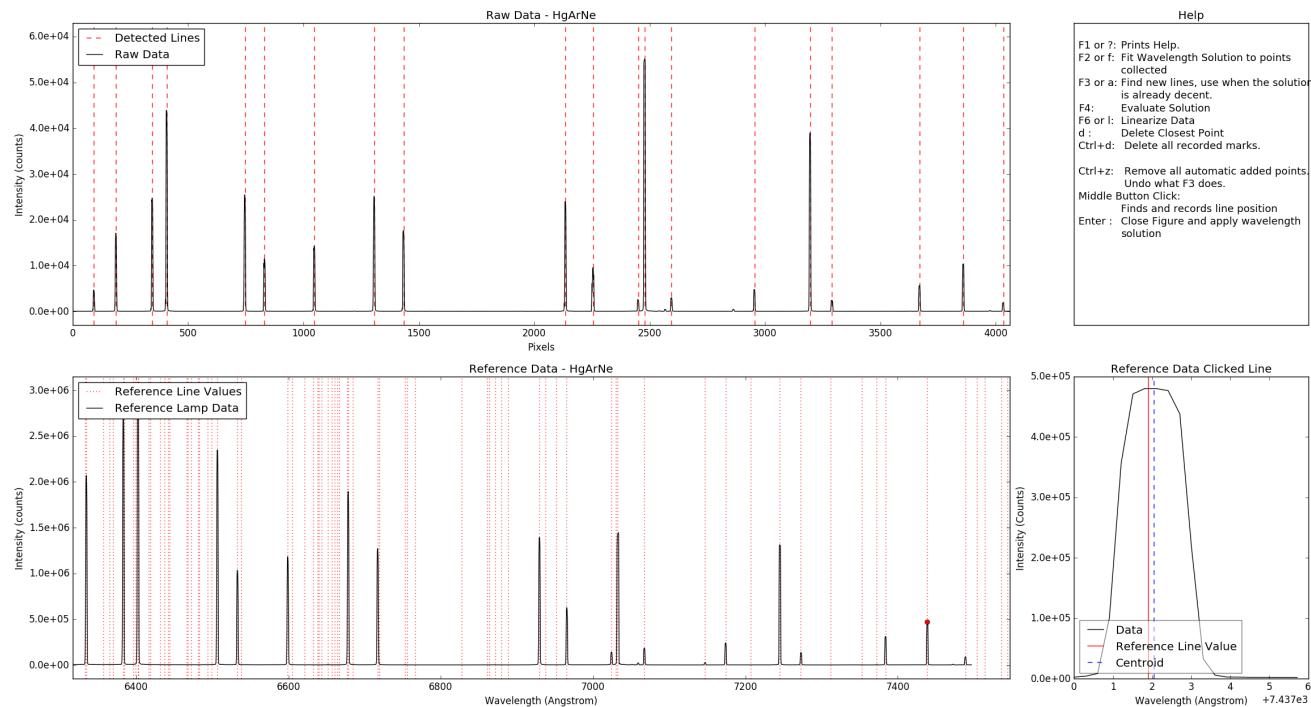
| Main Key | Alternative Key | Mouse Equivalent | Description |
|----------|-----------------|------------------|--|
| ? | F1 | None | Print Help message on the terminal |
| f | F2 | None | Fit a function to collected points |
| a | F3 | None | Find lines automatically |
| None | F4 | None | Evaluate wavelength solution |
| d | F5 | None | Remove point closest to the mouse pointer |
| l | F6 | None | Linearize and smooth spectrum |
| m | None | Middle Button | Register the line closest to the mouse pointer |
| ctrl+z | None | None | Deletes ALL automatically added points |

Running the Pipeline

| | | | |
|--------|------|------|---|
| ctrl+d | None | None | Deletes all recorded points |
| ctrl+q | None | None | Ends the program |
| Enter | None | None | Accepts wavelength solution and closes the figure |

It is advisable to practice a little bit to get familiar with this combination of keys and functions. Since we are still in the development stage, if you feel that the use of a key creates problems for you let us know.

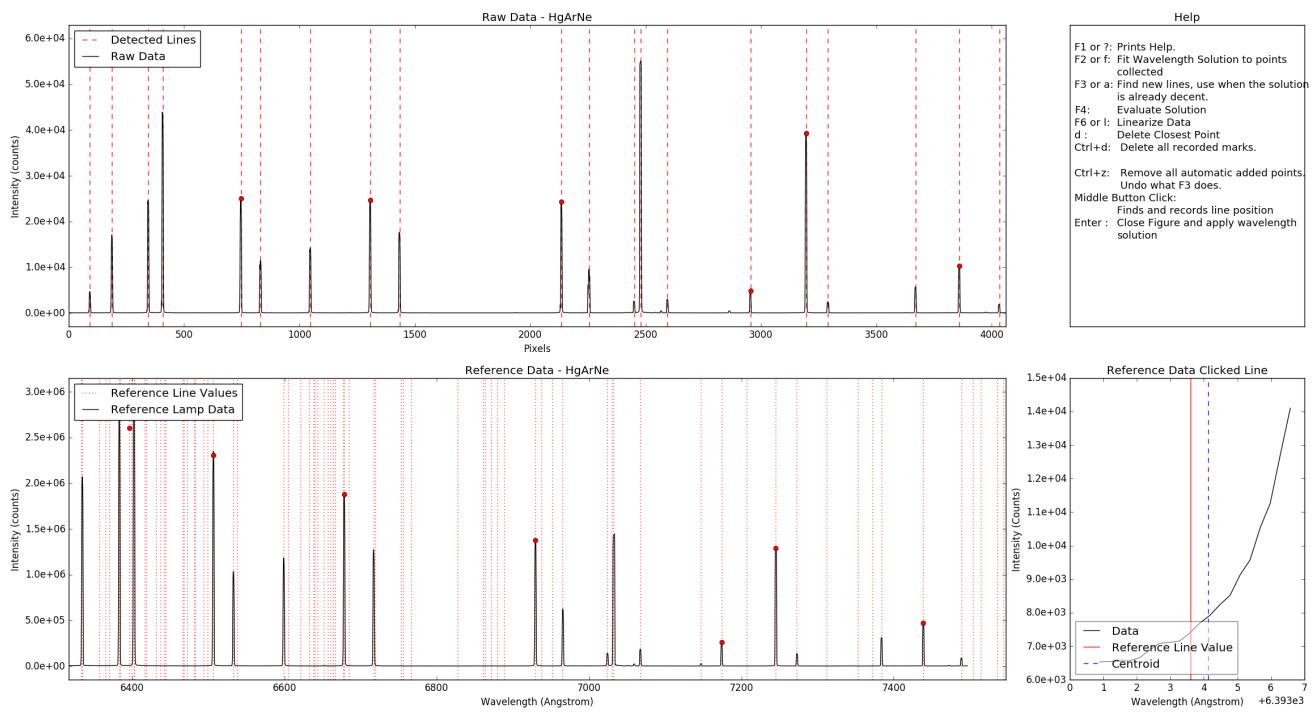
In the following image the mouse pointer was placed very close to the center of the line and then with a middle button click the line is marked with a red circle. The software will calculate a centroid (vertical dashed blue line) and then will try to find the closest reference value (vertical red line). Keep in mind that the reference data (red dotted lines) are not the *reference lamp* lines themselves but values obtained from the NIST site indicated above.



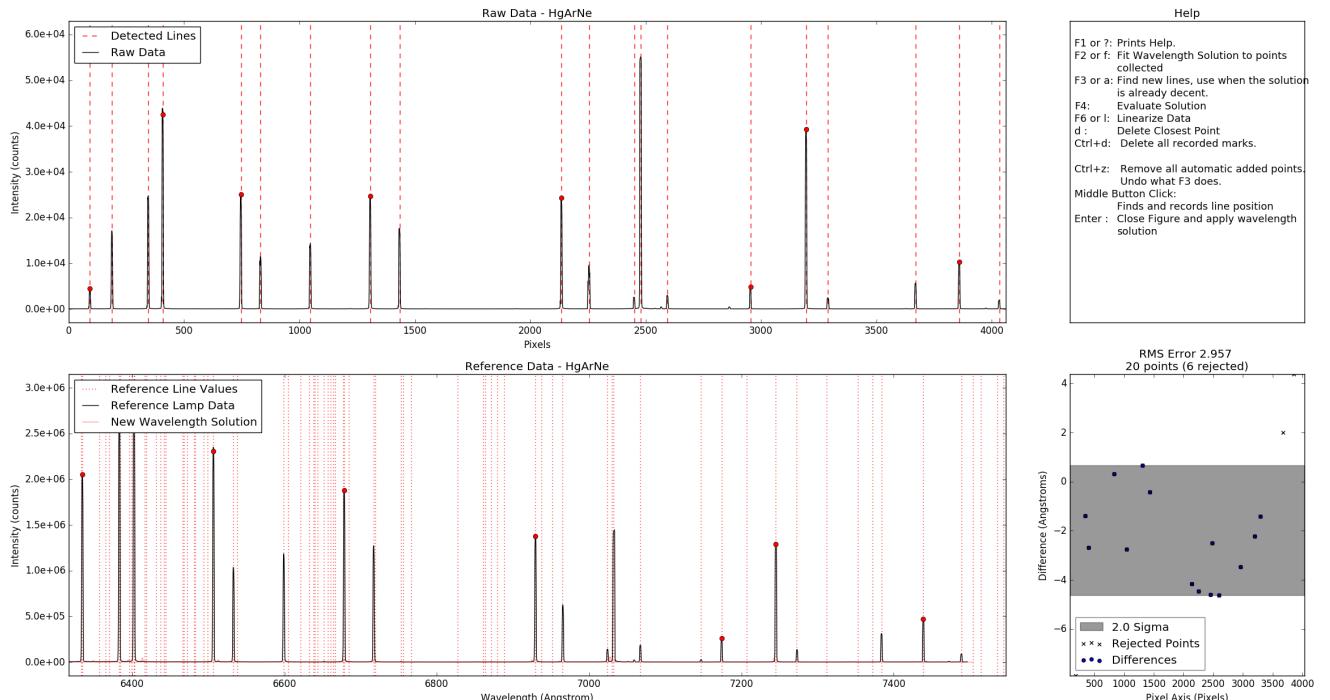
Then you have to find the corresponding line in the opposite plot. There is no preferred plot to start as long as you are consistent with the order in which you mark the lines.

3. In case you miss-identified a line, like in the image below, you can place the mouse pointer above the corresponding red circle and press **d** to delete it. It will search for the closest mark along the horizontal axis and remove it. If there is a counterpart in the opposite plot it will delete it too.

Running the Pipeline



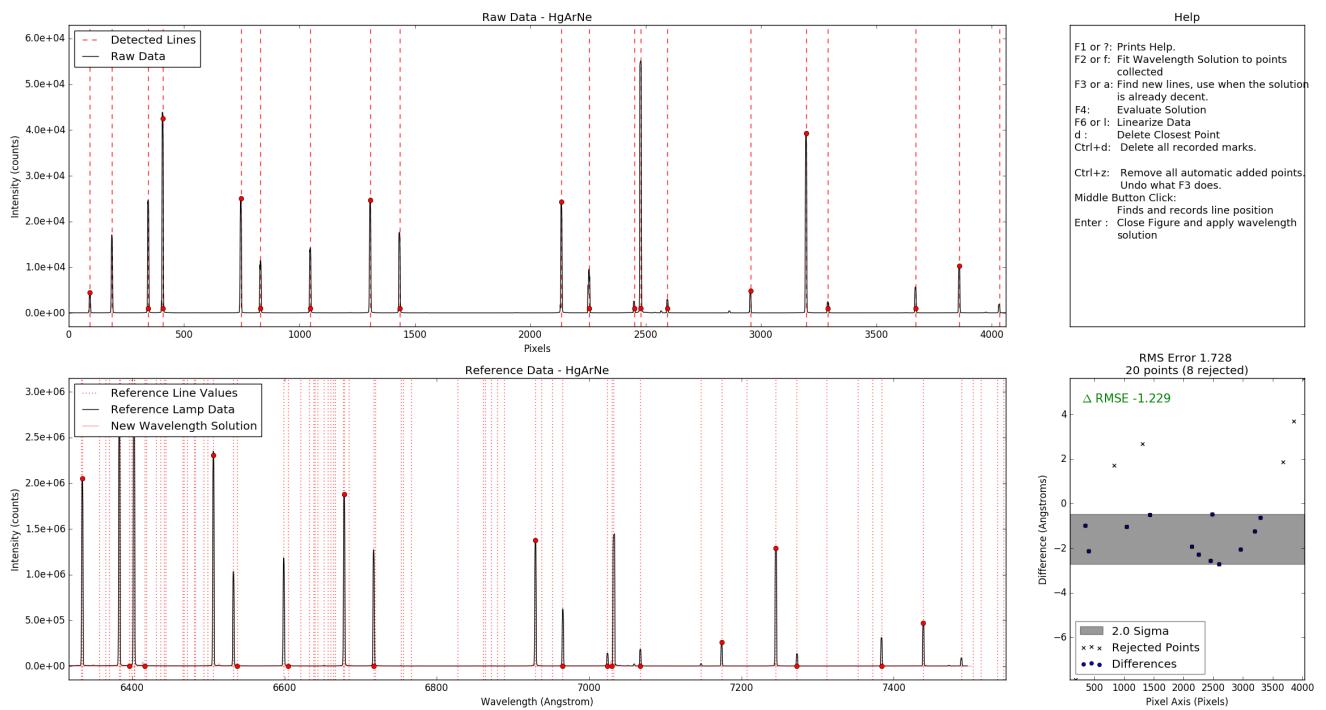
4. Once you matched a good number of lines, the minimum required by the fitting routines are four, you can either press F2 or f to make a fit of the pixels and angstrom values collected. Now the *Lower Right panel* will show the scatter plot of the fit. It is important to note here that the points in this plot do not represent the points you marked but rather each of the lines detected in your extracted 1-D comparison lamp spectrum (red dashed lines in the *Upper main plot*) It does one iteration of a 2-sigma clipping to reject outliers, and then it uses those values to calculate the Root Mean Square Error. In the example we show here the RMS is a bit high, but we will fix that below.



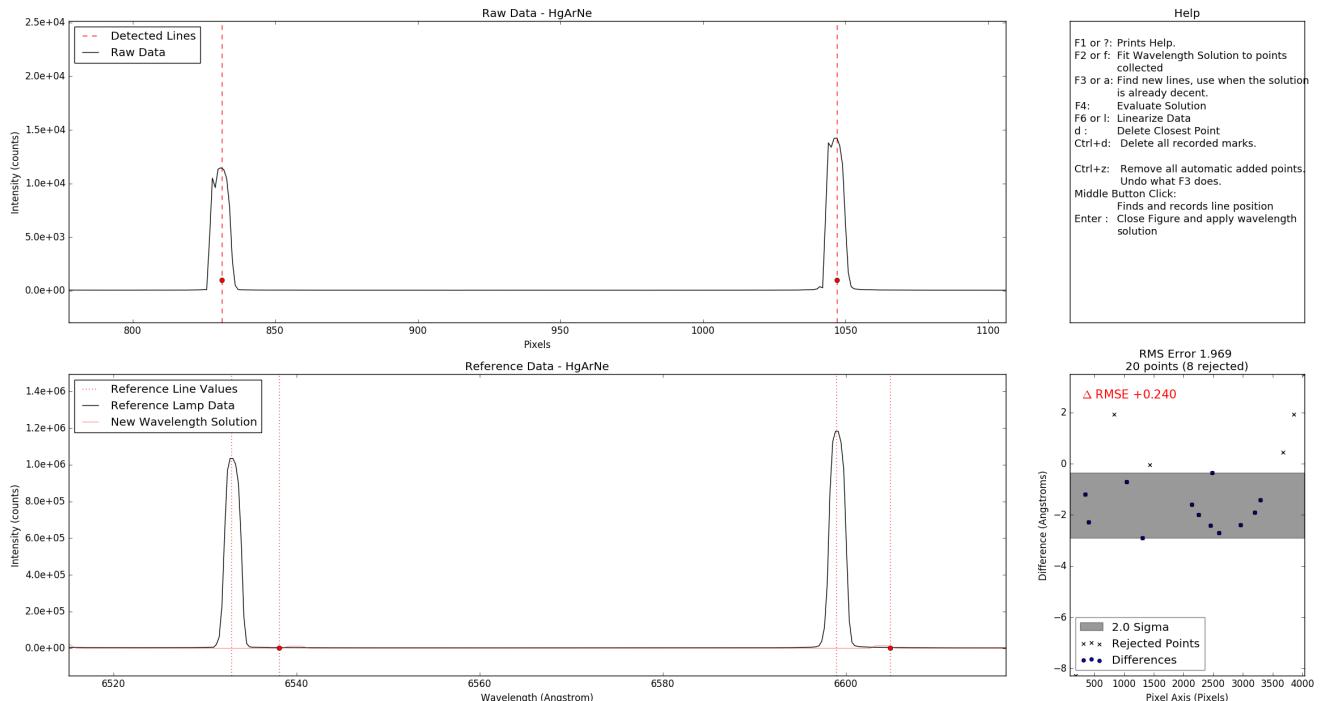
5. If you see that your current solution is decent you can press F3 or a and the pipeline will try to find more points automatically. The new matches found by the software are shown as red dots at the base of each line in the two main *Upper* and *Lower* panels.

The automatic finding routine is not perfect, and indeed it depends on the preliminary wavelength solution. It uses the detected lines in the uncalibrated 1-D lamp, applies the preliminary solution and tries to find a match in the reference line values. In most cases it improves the solution, but not always so keep that in mind. In this case the RMS error is reduced by almost half, which is good, but if you look closely you can see the mismatches; also the *Bottom right* panel will show you this.

Running the Pipeline

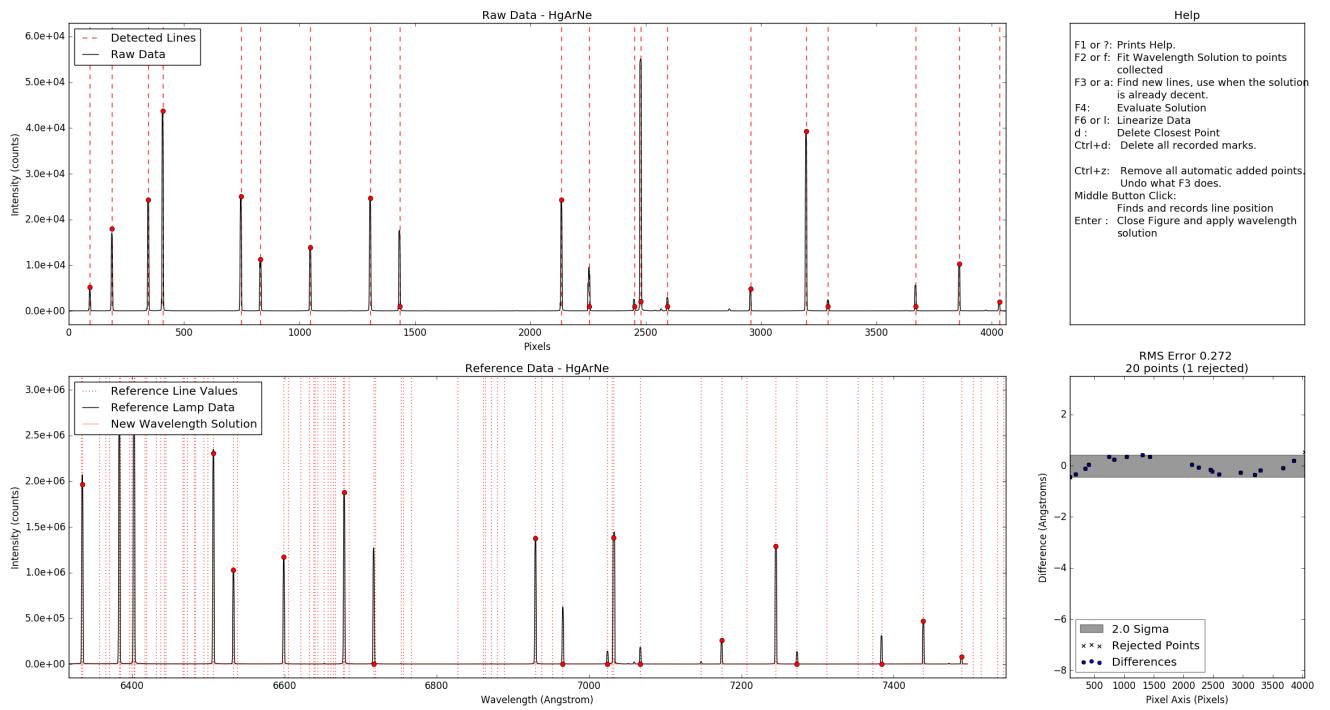


6. Now you we show an example of when the match is apparently good but in fact it's not. Here you need to zoom-in to see that there is an offset between the line centers and the matched laboratory line values, as shown in the figure below. You may have to apply different zoom values to your lamp and the reference lamp to get the plot to look like show it here. The solution in this case is locate the offending line(s), and delete it(them) pressing **d**. Then do a new match by clicking with the middle mouse button on the lines in the laboratory/reference plot, and the respective line in your 1-D uncalibrated lamp.



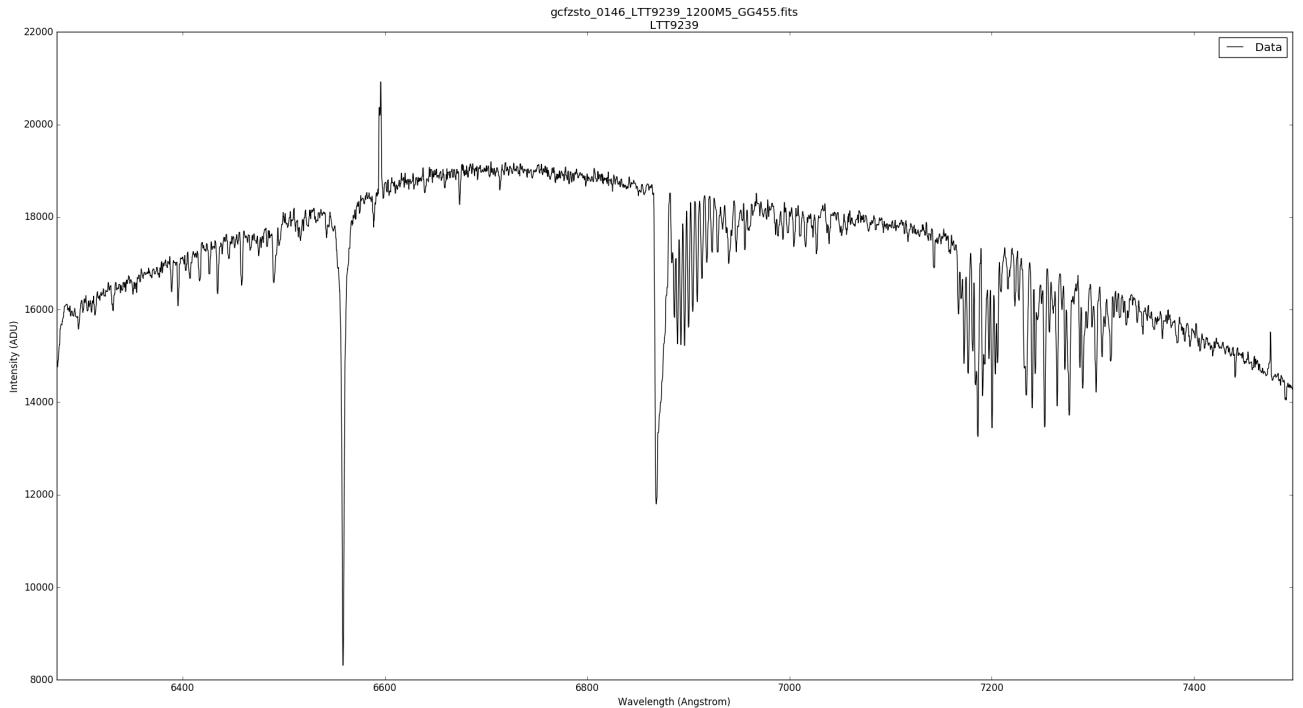
7. After you checked all the identifications and are happy with it, fit the solution again and you will obtain something like this:

Running the Pipeline

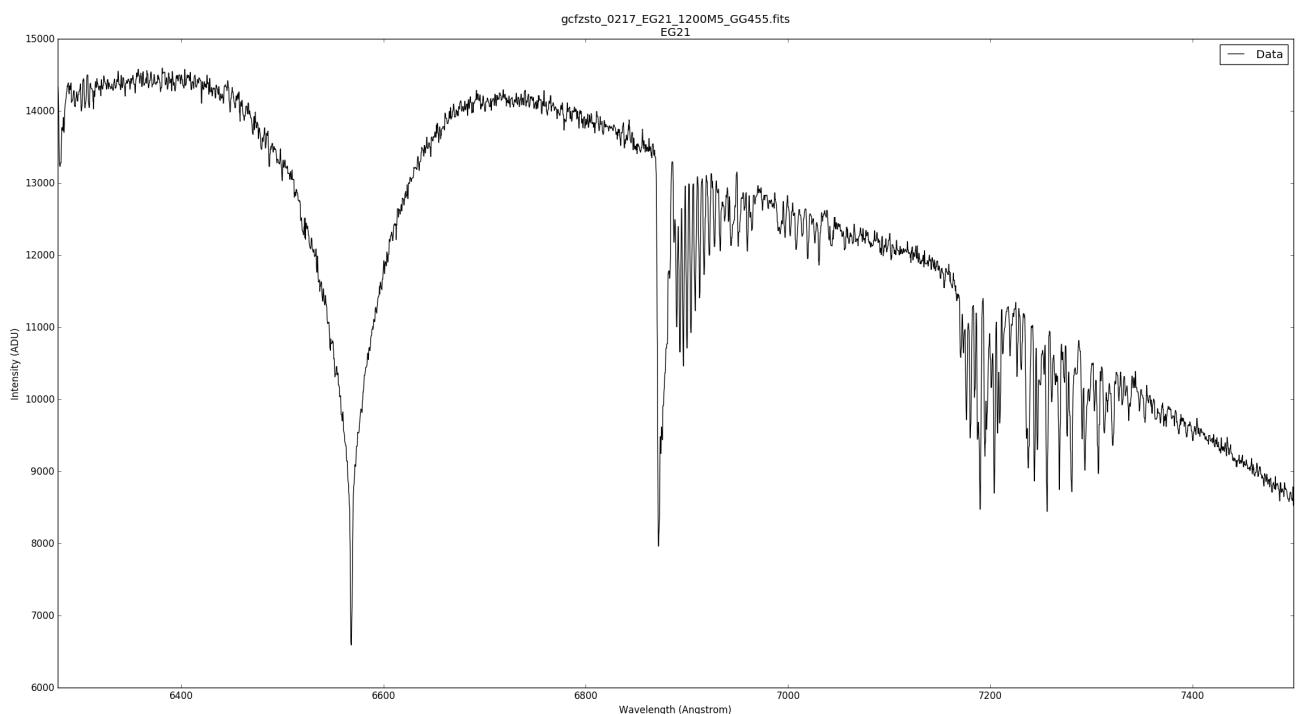
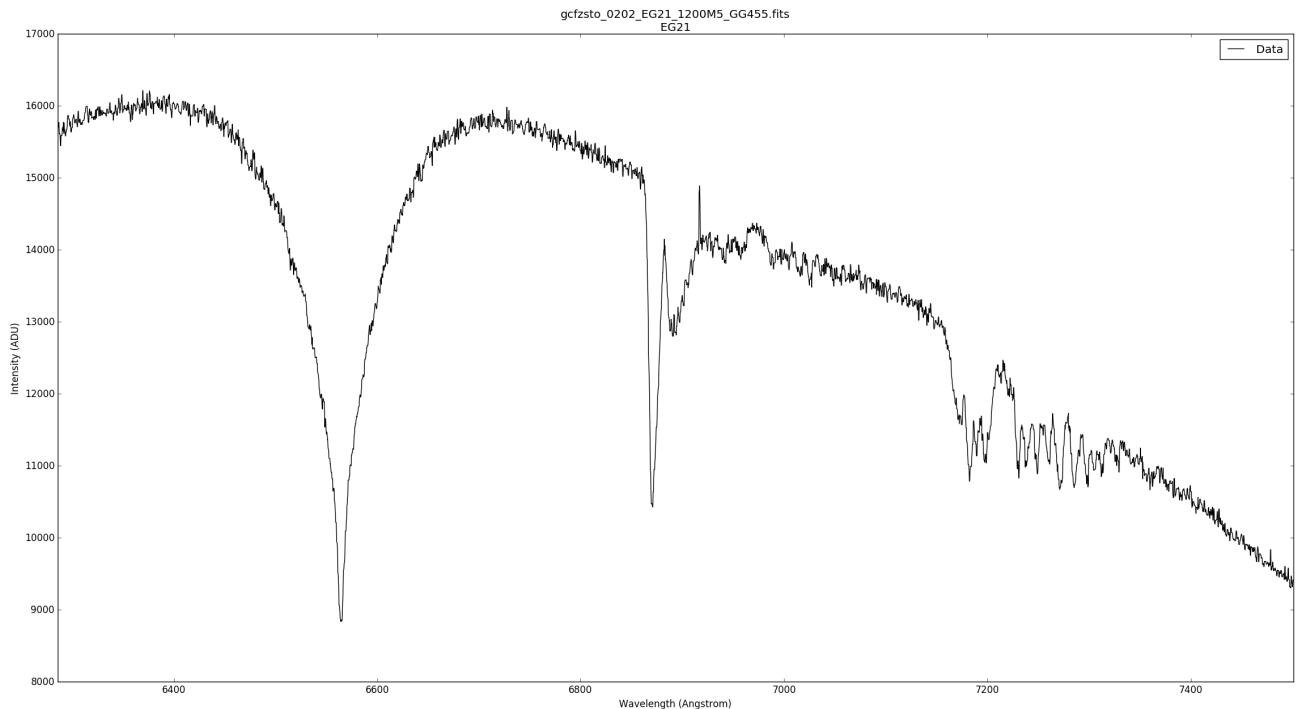


Here you can see that the *Bottom Right* panel shows the differences have a sinusoidal shape, which is also a sign that the solution can be improved. There are ways this can be implemented to refine the fit even further, but this is at present deferred to a later version.

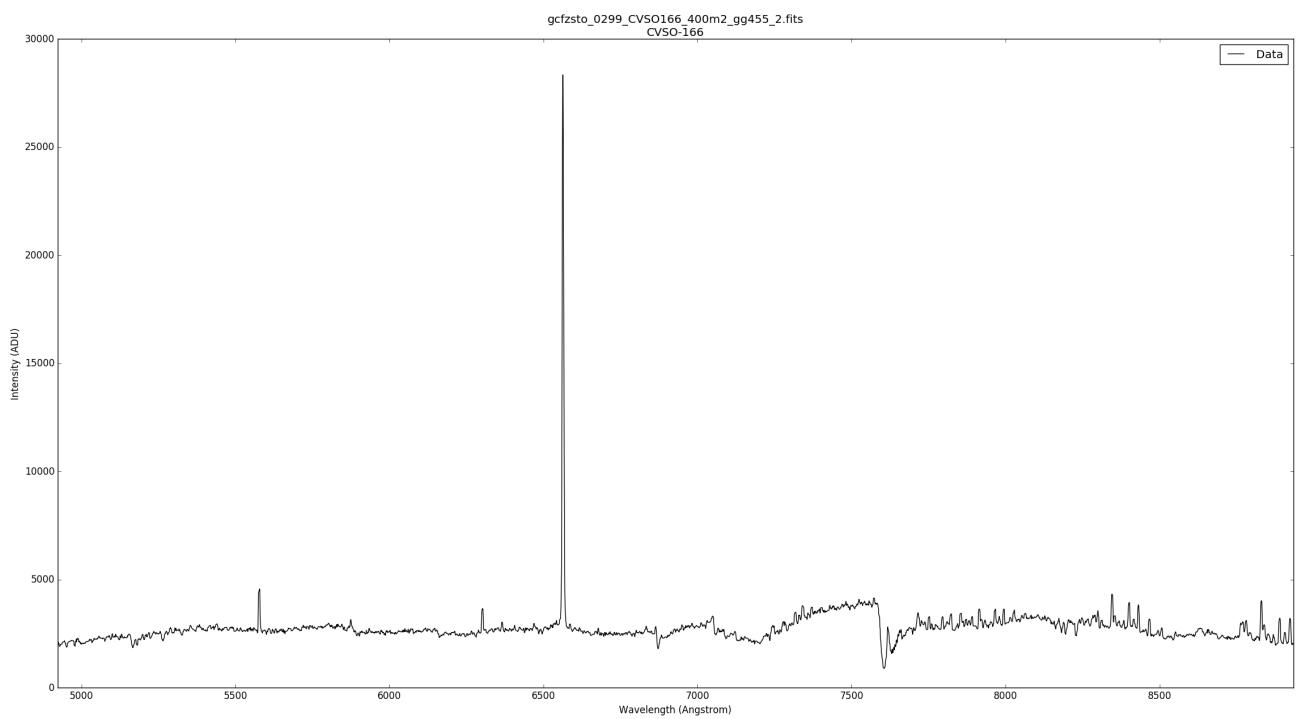
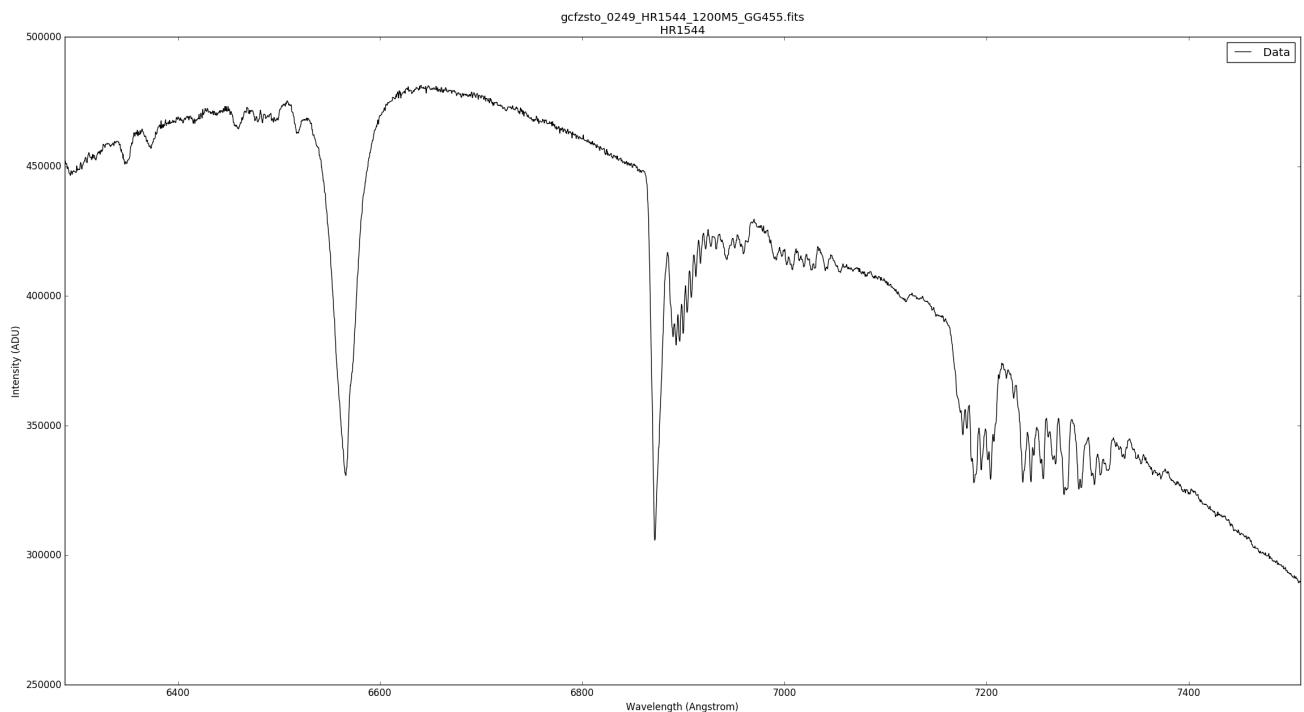
8. Finally, a few samples of the spectra extracted by the pipeline.



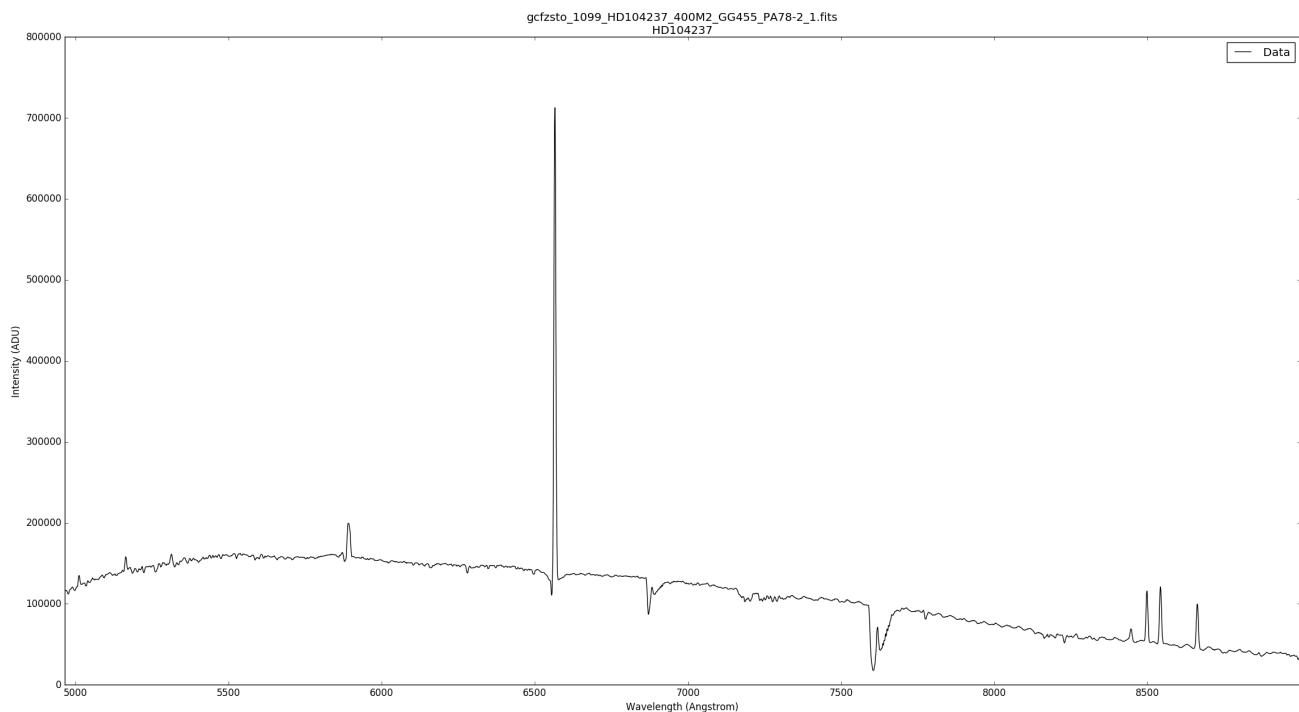
Running the Pipeline



Running the Pipeline



Troubleshooting



Troubleshooting

- The wavelength Solutions is way off: Check that the lamp was correctly registered in the header. Also check that the corresponding reference lamp exist. for instance is not the same to have HgArNe to HgAr
- Can't detect any objects: Check that the keyword OBSTYPE is correct.
- The reference data plot, in interactive mode, doesn't show anything or only vertical dotted lines: The reference lamp doesn't exist for that configuration, since this is used only for visual reference sometimes it will display the same lamp but in other instrument configuration, this will not affect the quality of the solution.

Appendix A: Installation Instructions

We strongly recommend installing the pipeline using *virtual environments*. Below you will find a summary of installation steps.

Warning

Remember that we are not providing any kind of support for installation. This documentation will be the only existing.

The following list provides a summary of all the steps.

- Install anaconda
- Add astroconda channel
- Create virtual environment
- Activate environment
- Install requirements
- Install pipeline

Anaconda and Virtual Environment

For anaconda installation we recommend you to check the [astroconda channel's documentation page](#). The instructions will be reproduced here but they might change for newer versions. Also they are limited to the *best case scenario*.

Warning

Anaconda installer requieres BASH. Don't try with other shell.

1. Installing anaconda - Go to <https://www.anaconda.com/downloads> and download the appropriate *anaconda installer* for your platform, most likely it has been automatically selected.

- Run the installer.

```
cd <download_directory>
bash <install_script>
```

- Once completed, check the bottom of `~/.bash_profile` or `~/.bashrc` there should be a new PATH definition with anaconda included.

2. Check anaconda installation

```
which conda
```

You should get a response similar to this:

```
~/bin/anaconda3/bin/conda
```

If you don't get this response check the detailed instructions on the astroconda site. Otherwise continue to the next step.

3. Configure Conda to use the *Astroconda Channel*

```
conda config --add channels http://ssb.stsci.edu/astroconda
```

4. Create a virtual environment. We have dropped support for python2.7 so you have to use >3.5.

```
conda create -n astroconda python=3 stsci
```

`astroconda` is the name of your environment, you can use any name you want.

5. Activate your environment.

```
source activate astroconda
```

Goodman Spectroscopic Pipeline

6. Get latest release of the *Goodman Spectroscopic Pipeline*

visit <https://github.com/soar-telescope/goodman/releases/latest> and download the *.zip or *.tar.gz

```
cd <download_location>
tar -xvf <pipeline_file>.tar.gz
or
unzip <pipeline_file>.zip
```

7. Install requirements from requirements.txt

```
cd <goodman_pipeline_unpacked_location>
pip install -r requirements.txt
```

8. Install the pipeline

```
pip install .
```

9. Upgrading the pipeline

```
pip install . --upgrade
```

Install DCR

Warning

Don't forget to cite: Pych, W., 2004, PASP, 116, 148

In terms of cosmic ray rejection we shifted to a non-python package because the results were much better compared to LACosmic's implementation in astropy. LACosmic was not designed to work with spectroscopy though.

The latest version of the Goodman Spectroscopic Pipeline uses a modified version of `dcr` to help with the pipeline's workflow. It is included under

```
<path_to_download_location>/goodman/pipeline/data/dcr-source/dcr/
```

`goodman` is the folder that will be created once you untar or unzip the latest release of the *Goodman Spectroscopic Pipeline*.

Important

The changes includes deletion of all `HISTORY` and `COMMENT` keywords, which we don't use in the pipeline. And addition of a couple of custom keywords, such as: `GSP_FNAM`, which stores the name of the file being created. `GSP_DCRR` which stores the reference to the paper to cite.

You are still encouraged to visit the official [Link](#) own by the author and let me remind you once more that you have to cite the paper mentioned several times in this manual.

Compiling DCR

Compiling `dcr` is actually very simple.

```
cd <path_to_download_location>/goodman/pipeline/data/dcr-source/dcr/
```

Then simply type:

```
make
```

This will compile *dcr* and also it will create other files. The executable binary here is *dcr*.

I have successfully compiled *dcr* in several platforms, such as:

1. Ubuntu 16.04
2. Centos 7.1, 7.4
3. MacOS Sierra
4. Solaris 11

Install binary DCR

This is a suggested method. If you are not so sure what you are doing, we recommend you following this suggestion. If you are more advanced user you just need the *dcr* executable binary in your \$PATH variable.

1. Open a terminal
2. In your home directory create a hidden directory *.bin* (Home directory should be the default when you open a new terminal window)

```
mkdir .bin
```

3. Move the binary of your choice and rename it *dcr*. If you compiled it, most likely it's already called *dcr* so you can ignore the renaming part of this step.

```
mv dcr.Ubuntu16.04 ~/.bin/dcr
```

Or

```
mv dcr ~/.bin/dcr
```

4. Add your \$HOME/.bin directory to your \$PATH variable. Open the file *.bashrc* and add the following line.

```
export PATH=$PATH:/home/myusername/.bin
```

Where /home/myusername is of course your home directory.

5. Close and reopen the terminal or load the *.bashrc* file.

```
source ~/.bashrc
```

Appendix B: On Goodman's Radial Velocity Precision

Important

The calculations presented here are for a single line only. Using the 0.45 arcsecond slit.

Here we present a summary of the best *radial velocity* precision that can be obtained with a given configuration. The equations used are listed below.

$$R = \frac{\lambda}{\Delta\lambda} = \frac{c}{v}$$

Then,

$$v = \frac{c}{R}$$

We can calculate the central wavelength for a given configuration which will correspond to λ DeltaLambda` is the dispersion in units of Angstrom/Pixel obtained from the Goodman Spectrograph Cheat Sheet.

The smallest grating available is the 0.45", then:

$$FWHM = \frac{slit-size}{pixel-scale} = \frac{0.45}{0.15} = 3.0$$

Then the limiting factor is not the spectrograph's dispersion but the *FWHM*,

$$\Delta\lambda = 3 * dispersion$$

| Grating | Mode | Central Wavelength | Dispersion | Resolving Power (R) | RV Limit |
|---------|------|--------------------|------------|---------------------|----------------|
| 400 | m1 | 505.281 | 1.00 | 1516 | 197.773 km / s |
| 400 | m2 | 700.154 | 1.00 | 2100 | 142.727 km / s |
| 600 | UV | 442.270 | 0.65 | 2041 | 146.867 km / s |
| 600 | Blue | 492.529 | 0.65 | 2273 | 131.881 km / s |
| 600 | Mid | 578.827 | 0.65 | 2672 | 112.218 km / s |
| 600 | Red | 777.885 | 0.65 | 3590 | 83.502 km / s |
| 930 | m1 | 384.521 | 0.42 | 2747 | 109.151 km / s |
| 930 | m2 | 469.125 | 0.42 | 3351 | 89.466 km / s |
| 930 | m3 | 554.787 | 0.42 | 3963 | 75.652 km / s |
| 930 | m4 | 639.418 | 0.42 | 4567 | 65.639 km / s |
| 930 | m5 | 724.936 | 0.42 | 5178 | 57.896 km / s |
| 930 | m6 | 809.084 | 0.42 | 5779 | 51.875 km / s |
| 1200 | m0 | 374.297 | 0.31 | 3622 | 82.765 km / s |
| 1200 | m1 | 424.181 | 0.31 | 4105 | 73.031 km / s |
| 1200 | m2 | 492.678 | 0.31 | 4768 | 62.878 km / s |
| 1200 | m3 | 561.804 | 0.31 | 5437 | 55.141 km / s |
| 1200 | m4 | 629.735 | 0.31 | 6094 | 49.193 km / s |
| 1200 | m5 | 699.087 | 0.31 | 6765 | 44.313 km / s |
| 1200 | m6 | 767.000 | 0.31 | 7423 | 40.389 km / s |
| 1200 | m7 | 835.851 | 0.31 | 8089 | 37.062 km / s |