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# HubbleCLIP: Associating Astronomical Observations and Natural Language with Multi-Modal Models

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## Abstract

We present a multi-modal model which associates astronomical observations imaged by the *Hubble* Space Telescope (HST) with natural language. The model is fine-tuned from a base CLIP model using successful proposal abstracts corresponding to HST observations, summarized via guided large language (LLM) model generation. We show that the model embodies a meaningful joint representation between observations and text through experiments targeting observation retrieval (i.e., retrieving the most relevant observations from a set using natural language queries) and description retrieval (i.e., querying the astrophysical object classes and science use cases most relevant to a given observation). The model demonstrates the potential for using generalist rather than task-specific models for astrophysics research, in particular by leveraging text as an interface.

## 1 Introduction

Machine learning (ML) is starting to have a significant impact on the sciences, with astrophysics being no exception. Machine learning methods have demonstrated promise when applied to every part of the astrophysics pipeline, from instrument design, to data acquisition, to its analysis. Until recently, most applications of ML within astrophysics have focused on replacing or augmenting traditional techniques with ML counterparts in order to improve performance on specific tasks.

The *Foundation Model* paradigm, in contrast, seeks to develop generalist models which can be deployed on a wide range of tasks. The paradigm has been highly successful in domains like computer vision and natural language processing, as demonstrated by the widespread adoption of tools like CLIP, ChatGPT, Dall-E, and Stable Diffusion. These models are typically pre-trained on massive amounts of unlabeled data using self-supervised learning techniques, enabling them to learn powerful representations which can be optionally fine-tuned to address domain-specific tasks. At the heart of the paradigm lies the triumph of scale – scaling up model size, dataset size, and compute. However, foundation models often benefit from fine tuning using a relatively small amounts of domain-specific data, increasing their usefulness when applied to those specific domains.

There is considerable interest in developing custom foundation models for the sciences, with astrophysics being ripe for such an effort due to several reasons. The first is the availability of large amounts of publicly-available data as a contingency of publicly-funded data-taking efforts. The second is the multi-modality

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inherent astrophysical observations, with different types of data (e.g., images, spectra, light curves, textual descriptions) often available for each observation. This multi-modality was recently exploited to train AstroCLIP (Lanusse et al., 2023) – a joint representation between multi-band images and optical spectra from the Dark Energy Spectroscopic Instrument (DESI). ASTROLLAMA (Nguyen et al., 2023) is another recent effort to fine-tune a publicly-available model (LLAMA-2) on astrophysics-specific textual data from the arXiv.

The CLIP (Radford et al., 2021) family of models has shown strong performance on a variety of downstream tasks including zero-shot classification and image retrieval.

GEOCLIP (Cepeda et al., 2023).

The success of the foundation model paradigm partly relies on the ability to flexibly leverage text as a *universal interface*. In this work, we take this outlook and train a joint representation between observations taken by the *Hubble* Space Telescope (HST) and natural language. We do so by using the associations between observation proposals and corresponding downstream observations. We show that fine-tuning a CLIP (Contrastive Language-Image Pre-training) model on this data enables learning meaningful joint representations.

The paper is organized as follows. In Sec. 2, we describe the dataset used in this work, including its curation and processing. In Sec. 3, we describe the methodology used to train and evaluate the model. In Sec. 4, we present the results of our experiments on image and text retrieval tasks. We discuss future prospects and conclude in Sec. 5.

## 2 Dataset Construction

We curate a dataset of images of *Hubble* observations and corresponding text descriptions. We rely on summarized versions of proposal abstracts from the Proposal Abstracts Catalog<sup>1</sup> – a catalog of all accepted *Hubble* proposals – as captions for the observations. The HST has been operational for 33 years, having been launched on April 24, 1990, and we use proposals and observations up to Cycle 31.

Examples of images and corresponding captions are shown in Tab. 1. It can be seen that these images have specific characteristics as well as artifacts particular to the nature of data-taking which distinguish them from the distribution of natural images typically used for large-scale pre-training of foundation models. This further motivates the need for fine-tuning on domain-specific data.

### 2.1 Data Selection and Pre-processing

Observations corresponding to individual proposal IDs are queried through the Mikulski Archive for Space Telescopes (MAST) via the `astroquery` interface. Products of type PREVIEW are filtered in, corresponding to preview postcard images. Note that these are not science-grade observations, but rather low-resolution images used for quick-look purposes; given the nature of associations we aim to learn, we deem this adequate for our current purposes. A maximum of 20 images are downloaded per proposal ID, selected at random, in order to avoid biasing the model towards proposals with a larger number of observations. Images are centered and resized to a resolution per side of 512 pixels before saving. Color previews (i.e., observations taken with multiple wavelength filters assigned to individual RGB channels) are manually excluded via a filename filter in order to maintain consistency across the dataset; models trained on datasets with color images included were observed to show worse performance on generalization metrics. If no appropriate images corresponding to an abstract are found, the abstract is excluded from the dataset.

In total 31,859 images corresponding to 4,438 abstracts are included in the fine-tuning dataset. 3,194 images are held out for validation, with no abstract being common between training and validation sets in order to ensure an independent set of text-image pairs for testing.

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<sup>1</sup>[https://archive.stsci.edu/hst/proposal\\_abstracts.html](https://archive.stsci.edu/hst/proposal_abstracts.html)

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## 2.2 Summarization via Guided Generation

Raw abstracts summarize the corresponding successful HST observing proposals, which intend to make the case for allocating *Hubble* telescope time towards a particular set of observations. These abstracts are written in a diversity of styles, formats, and lengths, being highly variable in the nature of content as well. Although the abstracts can be used as-is as image captions, we explore the use of summarization via guided LLM generation to standardize the captions used for fine-tuning the CLIP model. The goal is to summarize the objects and phenomena, as well as potential downstream science use cases corresponding to the HST observations in order to increase the signal between text and images.

The method from [Willard & Louf \(2023\)](#) is used to produce an LLM-generated summary of the abstract conforming to a particular schema, specified in JSON format. The schema is designed to represent a list of the objects (e.g., ‘Type Ia supernova’) and phenomena (e.g., ‘gravitational lensing’), as well as potential downstream science uses cases (e.g., ‘set constraints on supernova explosion models’) that could correspond to the eventual imaged observation given the abstract text.

The procedure guides the generation of LLM outputs while ensuring that the schema is respected at every point in the generation by masking out tokens that would violate the intended format. By framing the problem in terms of transitions between a set of finite-state machines, [Willard & Louf \(2023\)](#) showed that guided generation can be performed with negligible overhead compared to unconstrained generation. This ensures that the output of the LLM strictly conforms to the format of the following example:

```
1 {  
2     'objects_and_phenomena': ['star forming galaxy', 'lensed galaxy', ...],  
3     'science_use_cases': ['measure lensing magnification', 'probe spectral energy  
4         distributions', ...]  
}
```

which is then used to construct the summarized caption by combining the two key elements. Examples of raw abstracts and corresponding LLM-generated summaries are shown in Tab. 2. Further details on the summarization procedure, including the prompts used and a more detailed description of guided generation, are provided in App. A.

The open-weights, instruction-tuned model **MIXTRAL-8x7B**<sup>2</sup> is used to generate the summaries, with guided generation performed using the **outlines**<sup>3</sup> package.

The goal of summarization-via-guided-generation is to increase the signal between text and images by standardizing the captions used for fine-tuning the CLIP model.

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<sup>2</sup><https://huggingface.co/mistralai/Mixtral-8x7B-Instruct-v0.1>  
<sup>3</sup><https://github.com/outlines-dev/outlines>

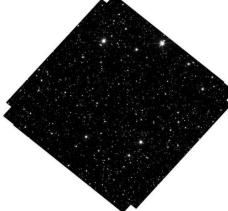
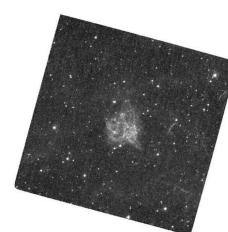
Image	Obs. cycle	Prop. ID	LLM-extracted summary
	26	15513	isolated black holes, background stars, Galactic bulge; constrain mass of isolated black holes, distinguish between black hole scenarios, analyze relative proper motions of stars
	19	12577	Cas A supernova, light echoes, interstellar dust, supernova outburst, shock breakout; Estimate radius of Cas A progenitor star, connect progenitor star to explosion to supernova to supernova remnant (SNR), analyze evolution of Cas A's spectrum over time, determine maximum-light characteristics of the supernova, probe properties of cooling envelope after shock breakout
	7	7340	young oxygen-rich supernova remnants SN0540–69.3, LMC, SMC, supernova debris, active pulsar, synchrotron nebula; characterize ionization structure and distribution of chemically peculiar debris in SN0540–69.3, determine ionization structure in the SN debris of E0102.2–7219, provide benchmarks for models of nucleosynthesis in massive stars, excitation mechanisms in extremely metal-rich plasmas, and supernova explosion dynamics, study the pulsar and synchrotron nebula in SN0540–69.3, investigate SN0540–69.3's proximity to SN 1987A in both space and time, and relation to the same extended complex of young stars
	22	13757	type Iax supernovae, white dwarfs, possible companion stars, accretion disks, luminous blue stars; constrain progenitor systems of type Iax supernovae, distinguish between explosion mechanisms, investigate mass transfer processes in accretion disks, determine if type Iax supernovae originate from massive stars

Table 1: Examples of images and corresponding captions, constructed using the LLM-extracted summaries. The CLIP model is fine-tuned on these text-image associations.

Prop. ID	Proposal abstract	LLM-extracted summary	
		Objects and phenomena	Science use cases
15513	<p>Category: Stellar Physics. A significant fraction of the mass of an old stellar population should be in the form of isolated black holes (BHs). Yet there has never been an unambiguous detection of a solitary BH. The only technique available to detect isolated BHs is astrometric microlensing-relativistic deflection of light from background stars. We have...</p>	isolated black holes, background stars, Galactic bulge	constrain mass of isolated black holes, distinguish between black hole scenarios, analyze relative proper motions of stars
12577	<p>Category: ISM AND CIRCUMSTELLAR MATTER. We propose to obtain time-resolved spectroscopy of the outburst of the enigmatic historical supernova Cas A using STIS spectroscopy of light scattered by a narrow filament of interstellar dust. Our group has identified recent, high-surface brightness filaments that are likely to provide high signal-to-noise reproduction of the evolving spectrum of...</p>	Cas A supernova, light echoes, interstellar dust, supernova outburst, shock breakout	Estimate radius of Cas A progenitor star, connect progenitor star to explosion to supernova to supernova remnant (SNR), analyze evolution of Cas A's spectrum over time, determine maximum-light characteristics of the supernova, probe properties of cooling envelope after shock breakout
7340	<p>Category: STELLAR EJECTA. We propose to use the WFPC2 and STIS CCD to obtain maximum spatial resolution emission-line images of the young, oxygen- rich supernova remnants SN0540–69.3 in the LMC and E0102.2– 7219 in the SMC. O IIILambda5007, S IIILambdaLambda6724 and O IIILambdaLambda3727 images of SN0540–69.3 will be used to characterize the ionization structure and...</p>	young oxygen-rich supernova remnants SN0540–69.3, LMC, SMC, supernova debris, active pulsar, synchrotron nebula	characterize ionization structure and distribution of chemically peculiar debris in SN0540–69.3, determine ionization structure in the SN debris of E0102.2–7219, provide benchmarks for models of nucleosynthesis in massive stars, excitation mechanisms in extremely metal-rich plasmas, and supernova explosion dynamics, study the pulsar and synchrotron nebula in SN0540–69.3, investigate SN0540–69.3's proximity to SN 1987A in both space and time, and relation to the same extended complex of young stars
13757	<p>Category: HOT STARS. Type Ia supernovae (SN Ia) have enormous importance to cosmology and astrophysics, but their progenitors and explosion mechanisms are not known in detail. Recently, observations and theoretical models have suggested that not all thermonuclear white-dwarf supernova explosions are normal SN Ia. In particular, type Iax supernovae (peculiar cousins to SN Ia), are...</p>	type Iax supernovae, white dwarfs, possible companion stars, accretion disks, luminous blue stars	constrain progenitor systems of type Iax supernovae, distinguish between explosion mechanisms, investigate mass transfer processes in accretion disks, determine if type Iax supernovae originate from massive stars

Table 2: Examples of the beginning of raw proposal abstracts (second column) and LLM (MIXTRAL-8x7B)-extracted summaries (rightmost two columns), separately extracting objects and phenomena as well as potential downstream science use cases. The LLM-extracted summaries are used for associating text with observations.

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### 3 Methodology

Our goal is to learn a semantically meaningful joint representation of HST image observations and natural language, with the intention of using it for a variety of downstream tasks. We leverage the strong generalization capabilities demonstrated by CLIP (Contrastive Language-Image Pretraining) and adapt these to work with domain-specific *Hubble* data via fine-tuning; we describe these below.

#### 3.1 Language-Image Pre-training

CLIP (Contrastive Language-Image Pretraining; Radford et al., 2021) is a multi-modal model pre-trained on a large corpus of image-text pairs via weak supervision using a contrastive loss. Given a minibatch  $\mathcal{B}$  of  $|\mathcal{B}|$  image-text pairs  $\{(I_i, T_i)\}$ , the goal is to align the learned representations of corresponding (positive) pairs  $(I_i, T_i)$  while repelling the representations of unaligned (negative) pairs  $(I_i, T_{j \neq i})$ . Image and text encoders  $f(\cdot)$  and  $g(\cdot)$  are used to map images and text to a common embedding space, and the standard softmax-based InfoNCE (Oord et al., 2018) contrastive loss function is given by

$$\mathcal{L}(\mathcal{B}) = -\frac{1}{2|\mathcal{B}|} \sum_{i=1}^{|\mathcal{B}|} \left( \overbrace{\log \frac{e^{tx_i \cdot y_i}}{\sum_{j=1}^{|\mathcal{B}|} e^{tx_i \cdot y_j}}}^{\text{image} \rightarrow \text{text softmax}} + \overbrace{\log \frac{e^{tx_i \cdot y_i}}{\sum_{j=1}^{|\mathcal{B}|} e^{tx_j \cdot y_i}}}^{\text{text} \rightarrow \text{image softmax}} \right) \quad (1)$$

where  $x_i = \frac{f(I_i)}{\|f(I_i)\|_2}$  and  $y_i = \frac{g(T_i)}{\|g(T_i)\|_2}$  are the normalized representations of the  $i$ -th image and text, respectively, and  $t$  is a (typically learnable) temperature hyperparameter.

As the base model, we use the CLIP-ViT-B/16 (Radford et al., 2021) variant trained by OpenAI. This model uses a (12-layer, 12-head, 768-embedding dimension) vision transformer as the image encoder and a (12-layer, 8-head, 512-embedding dimension) sequence transformer as the text backbone. The text encoder has a maximum length of 77-tokens and the image encoder a native resolution of  $224 \times 224$ . Linear projection layers map the outputs of the image and text encoders to a common embedding space of dimension 512. In total, the model has 149,620,737 trainable parameters. The model was trained on 400 million image-text pairs from the internet.

#### 3.2 Fine-tuning Procedure and Evaluation

The base CLIP model is fine-tuned using the dataset described in Sec. 2, using either the LLM-summarized data or the raw proposal abstracts. When using raw proposal abstracts, random chunks of the text delimited by periods are selected on the fly to fit within the maximum length of the text encoder. Images are downsampled to the native resolution of the image encoder and randomly rotated over training. Given the relatively modest size of the fine-tuning set, a batch size  $|\mathcal{B}| = 32$  is used throughout. We note that positive and negative image-text association is noisy and imperfect, since multiple images can be associated with the same abstract.

We explore three different methods of training the model on our domain dataset: (1) Fine-tuning the entire network, starting from the pre-trained base model; (2) Freezing the base image and text encoders, and training a small projection head on top of these; and (3) Training the entire model from scratch.

The model is trained over 50,000 steps with 5000 linear warmup steps and cosine decay using the AdamW optimizer (Loshchilov & Hutter, 2019; Kingma & Ba, 2015) with peak learning rate of either  $10^{-5}$  or  $10^{-6}$  and weight decay  $10^{-3}$ . Training takes approximately 6 hours on 4 Nvidia A100 GPUs.

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## 4 Results and Discussion

### 4.1 Fine-tuned Retrieval Accuracy

### 4.2 ‘Zero-shot’ Hypothesis and Object Retrieval

### 4.3 Text-to-Image Retrieval

## 5 Outlook and Conclusions

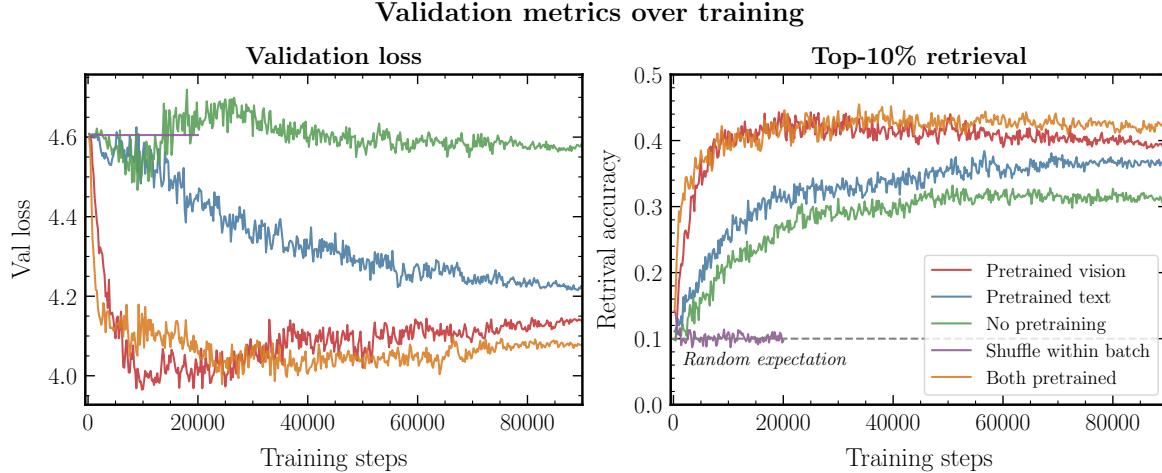


Figure 1: Retrieval accuracy

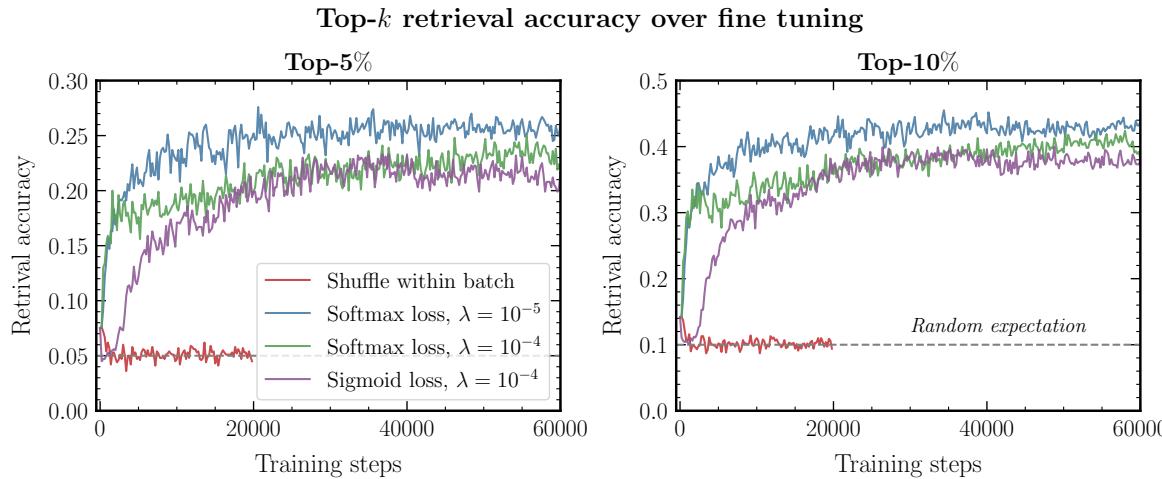


Figure 2: Retrieval accuracy

### Code

### Acknowledgments

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### Image-to-Text Retrieval

Image	Top classes (fine-tuned)	Top classes (base)	Abstract
	<ul style="list-style-type: none"> <li>1. dark matter</li> <li>2. Einstein rings</li> <li>3. galaxy mergers</li> <li>4. gravitational lensing</li> <li>5. galaxy clusters</li> </ul>	<ul style="list-style-type: none"> <li>1. galaxy clusters</li> <li>2. ultra diffuse galaxies</li> <li>3. dwarf galaxies</li> <li>4. gravitational lensing</li> <li>5. crowded stellar field</li> </ul>	<p><b>Category: COSMOLOGY.</b> We propose to study the physical nature of dark matter by using massive, merging clusters of galaxies. As shown with the Bullet Cluster (1E0657-56), such massive well-measured systems are critical for our understanding of dark matter. By more than doubling the number of clusters in the sample and obtaining systems at different observation angles, impact parameters, geometrical arrangements, and merger velocities, the systematic uncertainties in the dark matter cross section calculations can be improved substantially, allowing us to move from rough order of magnitude estimates to measurements with quantifiable uncertainties that can be compared usefully with the predictions from numerical simulations, and the constraints on alternate gravity models become unambiguous. Our proposed targets are three extraordinary, merging galaxy clusters with X-ray and optical offsets that are placed at ideal redshifts for such a study: A520, A1758N, and A2163. To pin down the position of the dark matter component we require high resolution, absolutely calibrated mass maps. High resolution gravitational lensing data is needed to attain this goal, which can only be achieved with the excellent resolving power of the HST.</p>
	<ul style="list-style-type: none"> <li>1. dark matter</li> <li>2. galaxy mergers</li> <li>3. Einstein rings</li> <li>4. gravitational lensing</li> <li>5. dark energy</li> </ul>	<ul style="list-style-type: none"> <li>1. ultra diffuse galaxies</li> <li>2. galaxy clusters</li> <li>3. gravitational lensing</li> <li>4. high-redshift quasars</li> <li>5. dwarf galaxies</li> </ul>	<p><b>Category: COSMOLOGY.</b> We propose to study the physical nature of dark matter by using massive, merging clusters of galaxies. As shown with the Bullet Cluster (1E0657-56), such massive well-measured systems are critical for our understanding of dark matter. By more than doubling the number of clusters in the sample and obtaining systems at different observation angles, impact parameters, geometrical arrangements, and merger velocities, the systematic uncertainties in the dark matter cross section calculations can be improved substantially, allowing us to move from rough order of magnitude estimates to measurements with quantifiable uncertainties that can be compared usefully with the predictions from numerical simulations, and the constraints on alternate gravity models become unambiguous. Our proposed targets are three extraordinary, merging galaxy clusters with X-ray and optical offsets that are placed at ideal redshifts for such a study: A520, A1758N, and A2163. To pin down the position of the dark matter component we require high resolution, absolutely calibrated mass maps. High resolution gravitational lensing data is needed to attain this goal, which can only be achieved with the excellent resolving power of the HST.</p>
	<ul style="list-style-type: none"> <li>1. crowded stellar field</li> <li>2. supernova remnants</li> <li>3. compact stellar remnants</li> <li>4. primordial black holes</li> <li>5. pre-main sequence stars</li> </ul>	<ul style="list-style-type: none"> <li>1. stellar abundances</li> <li>2. stellar populations</li> <li>3. interstellar medium</li> <li>4. pre-main sequence stars</li> <li>5. Cepheid variables</li> </ul>	<p><b>Category: RESOLVED STELLAR POPULATIONS.</b> Exploiting the full power of the Wide Field Camera 3 (WFC3), we propose deep panchromatic imaging of four fields in the Galactic bulge. These data will enable a sensitive dissection of its stellar populations, using a new set of reddening-free color indices we have constructed from multi-wavelength filters across UV, optical, and infrared wavelengths. This will provide the first multi-wavelength census for the stellar populations of thousands of individual bulge stars. Proper motions of these stars derived from multi-epoch observations will allow separation of the stellar populations into distinct kinematic components. Photometric and astrometric tools will support a wide range of bulge studies. Using these photometric and astrometric tools, we will reconstruct the detailed star formation history as a function of position within the bulge, and thus differentiate between rapid and gradual star formation. We will also measure the chemical evolution of the bulge, revealing how the characteristic mass of star formation varies with chemistry. Our sample of bulge stars with accurate metallicities will extend the knowledge of stellar populations in the bulge to higher metallicity, revealing the evolution of SF with metallicity. Our measurements will extend this knowledge to a remote environment with a very distinct chemistry. Our proposal also includes observations of six well-studied globular and open star clusters; these observations will serve to calibrate our photometric and astrometric products, and to validate the stellar populations in the bulge. Besides enabling our own program, these products will provide powerful new tools for a host of other stellar-population investigations with HST/WFC3. We will deliver all of the products from this Treasury Program to the community in a timely fashion.</p>
	<ul style="list-style-type: none"> <li>1. low surface brightness galaxies</li> <li>2. star formation histories</li> <li>3. galaxy formation</li> <li>4. ultra diffuse galaxies</li> <li>5. circumgalactic medium</li> </ul>	<ul style="list-style-type: none"> <li>1. gravitational lensing</li> <li>2. high-redshift quasars</li> <li>3. brown dwarfs</li> <li>4. trans-Neptunian objects</li> <li>5. Kuiper Belt objects</li> </ul>	<p><b>Category: Stellar Populations and the Interstellar Medium.</b> Observations of the ultra-faint dwarfs (UFDs), as relics of the epoch of reionization, allow us to probe the earliest epochs of star formation (SF). In particular, the UFDs in low density environments for most galaxies provide unique tools to probe the effects of early environmental conditions on the SF history of the UFDs and their stellar populations. We have been able to obtain a detailed record of the SF history of the UFDs in the local environment. We propose to obtain deep ACS and UVIS imaging in F606W and F814W for 2 LMC-like galaxies, and to obtain deep ACS imaging in F606W and F814W for 2 M31-like galaxies. This proposal is designed to identify systematic differences in the stellar populations of recently captured UFDs vs. long-term MW satellites (data available from previous programs) by using high-fidelity color-magnitude diagrams constrained by the latest stellar population models. The main goals of this program are: (1) to test whether the SF history of the UFDs is quenched at different times with different rate in UFDs in low density environment at early times, probing the patchiness of reionization by directly comparing with theoretical predictions; (2) identify variations in the sub-Solar IMF across UFDs born in different environments; (3) pave the way for a more accurate constraint on the MW halo mass.</p>
	<ul style="list-style-type: none"> <li>1. supernovae</li> <li>2. Cepheid variables</li> <li>3. star clusters</li> <li>4. star forming galaxies</li> <li>5. dust</li> </ul>	<ul style="list-style-type: none"> <li>1. high-redshift quasars</li> <li>2. gravitational lensing</li> <li>3. Kuiper Belt objects</li> <li>4. compact stellar remnants</li> <li>5. ultra diffuse galaxies</li> </ul>	<p><b>Category: RESOLVED STELLAR POPULATIONS.</b> We propose to test two of the clearest predictions of the theory of evolution of massive-star evolution: 1) the formation of Wolf-Rayet stars depends strongly on these stars' metallicity (Z), with relatively fewer WR stars forming at lower Z, and 2) Wolf-Rayet stars die as Type Ib or Ic supernovae. To carry out these tests we will use the Hubble Space Telescope (HST) to study the star-forming regions of the Large Magellanic Cloud (LMC). It is important, we will test the hypothesis that Superclusters like 3D Dorado are always richly populated with WR stars, and by implication that these WR stars are responsible for the special signatures of starburst galaxies. Our previous HST survey of the SF history of the SFH of the LMC and their stellar populations has shown that the SF history of the LMC is very similar to that of the recent star-forming history of these large complexes: young cores of O and WR stars are surrounded by older halos containing older stars. We propose to use the HST to study the SF history of the LMC and the WR stars in the supercluster. This proposal is to directly target this paradigm in a single WR in M101 being the ideal target. The abundance gradient across M101 (a factor of 20) suggests that relatively many WR will be found in the inner parts of this galaxy than in the outer parts. The WR population in M101 may be abundant enough for one to erupt as a Type Ib or Ic supernova within a generation. The clear WR stars in M101 are heavily obscured by dust, so we will use the HST to study the SF history of the LMC and the WR stars heavily populated by WR stars, are common in M101. It is widely claimed that such Superclusters produce the integrated luminosity of thousands of WR stars located in hundreds of M101 Superclusters, and correlate those numbers against the Supercluster sizes and luminosities. It is likely (but far from certain) that Supercluster sizes and emission-line luminosities are driven by their Wolf-Rayet star content. Our sample will be the largest and best-ever Supercluster/Wolf-Rayet sample, an excellent local proxy for characterizing starburst galaxies Superclusters.</p>

Figure 3: Retrieval accuracy

High Energy Physics of U.S. Department of Energy under grant Contract Number DE-SC0012567. The computations in this paper were run on the FASRC Cannon cluster supported by the FAS Division of Science Research Computing Group at Harvard University.

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Based on observations made with the NASA/ESA Hubble Space Telescope, and obtained from the Hubble Legacy Archive, which is a collaboration between the Space Telescope Science Institute (STScI/NASA), the Space Telescope European Coordinating Facility (ST-ECF/ESAC/ESA) and the Canadian Astronomy Data Centre (CADC/NRC/CSA).

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Text-to-Image Retrieval: Base Model

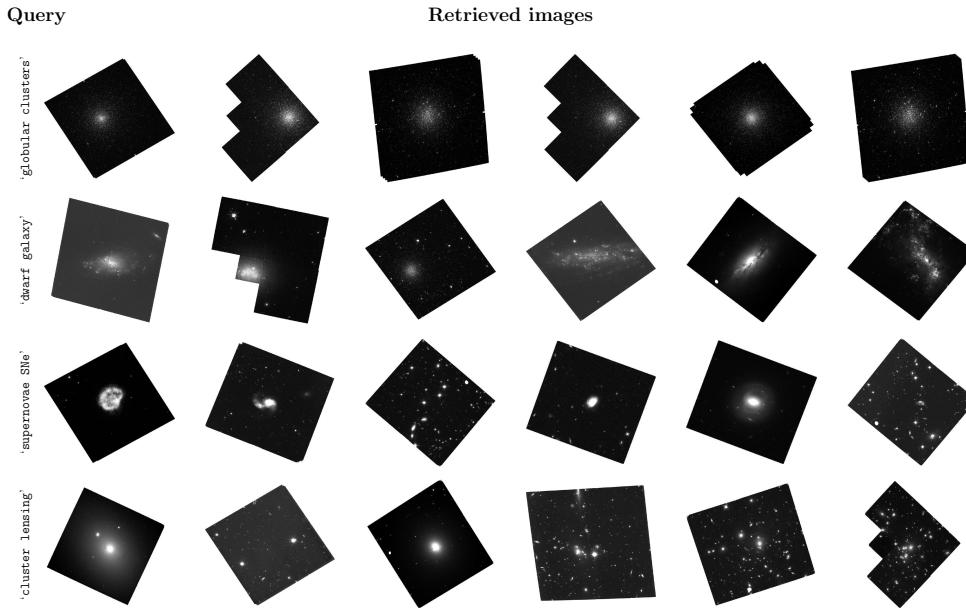


Figure 4: Retrieval accuracy

Text-to-Image Retrieval: Fine-Tuned Model

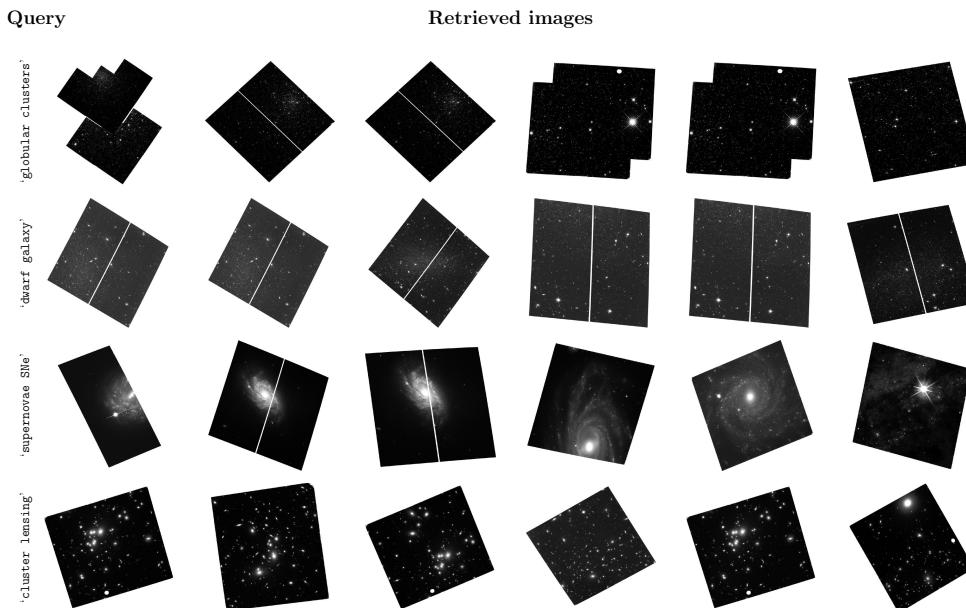


Figure 5: Retrieval accuracy

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Francois Lanusse, Liam Parker, Siavash Golkar, Miles Cranmer, Alberto Bietti, Michael Eickenberg, Geraud Krawezik, Michael McCabe, Ruben Ohana, Mariel Pettee, et al. Astroclip: Cross-modal pre-training for astronomical foundation models. *arXiv preprint arXiv:2310.03024*, 2023.

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## A Summarization via regex-guided generation

The following prompt is used to summarize the abstracts.

```
1 @outlines.prompt
2 def prompt_fn(abstract):
3
4     """<s>[INST] You are an expert astrophysicist, with broad expertise across
5        observational and theoretical astrophysics. You are able to extract core information
6        from astrophysical texts.
7
8 Abstract: "{{abstract}}"
9
10 Follow these instructions exactly:
11 - Mention up to 5 items for both categories; do not mention more than 5 items in either
12   category.
13 - Choose the most relevant ones if there are more than 5 items in a category.
14 - Never mention the Hubble Space Telescope, HST, or the HST archive.
15 - Mention the class (e.g., barred spiral galaxy) and not just the specific instance (e.g.,
16   Andromeda).
17 - Name the objects in the science use cases, if appropriate.
18 - Write out full names of objects in addition to acronyms.
19 - Do not list irrelevant objects which do not describe the eventual observation, such as
20   units or proposal Cycle numbers. List fewer but more relevant objects, if in doubt.
21 - Each science case listed must be self-contained but succinct.
22 - Only write in English.
23 - Do not list items that are too generic (e.g., galaxy, faint object, kinematics)
24 - The total length of text should not exceed 80 words.
25 - Present your lists in a comma-separated format; no dashed or numbered lists.
```

---

```
24 Example output: {'objects_and_phenomena':'spiral galaxies, galaxy clusters, supernova
25   remnants', 'science_use_cases':'model galactic structure and evolution, characterize
26   dark matter distribution in clusters, analyze expansion rates of supernova remnants'}
27
28 [/INST]
29 """
```