

DUST AND GAS IN NGC3627

DUST AND GAS IN NGC3627 USING OBSERVATIONS FROM
SCUBA-2

By

JONATHAN H. NEWTON, B.A.

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AUTHOR: Jonathan Newton, B.A. (Western Kentucky University)

SUPERVISOR: Christine D. Wilson

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Abstract

Saw some dust and wanted to do something about it!

To my family and Poly.

Acknowledgements

When life looks like easy street, there is danger at your door... -Robert Hunter

Thank Chris and group members of course. Don't forget Christian!

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Chapter 1

Introduction

1.1 The Interstellar Medium

The interstellar medium (ISM) is the material between that fills the space between the stars in galaxies and is composed of cosmic rays, magnetic field lines, dust, and gas. Each one of these constituents plays a role in shaping overall physical and chemical behavior of a galaxy (Ferrière, 2001). This material is inhomogeneously spread about the galaxy's disk with dust and gas densities ranging from 1 particle per a thousand cubic centimeters to millions particles per one cubic centimeter, and the temperatures range from 10 K in the most dense regions to 10^6 K in the most diffuse (Ferrière, 2001). The constituents of the ISM will interact leading to the creating of molecular gas in giant molecular clouds (GMCs) and eventually lead to the formation of stars (Krumholz, 2014; Kennicutt & Evans, 2012). When the stars are formed, they will eject new material back into the ISM through stellar winds and supernovae repopulating the ISM creating a galactic ecosystem, where each successive iteration introduces heavier atoms creating a galactic ecosystem (Ferrière, 2001).

In this thesis we focus on the physical properties of the dust component of the ISM of the nearby barred spiral galaxy NGC3627 to learn more about its abundance of molecular gas.

1.1.1 GMC Formation

Intergalactic gas can be accreted into a host galaxy where it will condense and lead to the formation of molecular clouds (Kennicutt & Evans, 2012). The dominant process of GMC formation is divided between two camps, either a “bottom-up” or “top-down” formation scenario (McKee & Ostriker, 2007). The bottom-up scenario consists of small clouds of cold HI coagulating to eventually form a GMC (Field & Saslaw, 1965; Kwan, 1979). The major concern with the bottom-up scenario is if we include heating mechanisms in the cloud, coagulation will cease before the observed masses are reached (McKee & Ostriker, 2007). In fact, the time scales required to form a cloud typical of those we observe would take around 10^8 years to accumulate a minimum of $10^5 M_{\odot}$, which is much greater than the expected GMC lifetime (McKee & Ostriker, 2007).

The alternative formation scenario, top-down, postulates that GMC formation comes from instabilities in the diffuse ISM causing the clouds to collapse from their surrounding medium (McKee & Ostriker, 2007). Two main instabilities are thought to be responsible for the collapse. The first type of instability is a Parker instability, which involves distortion of magnetic field lines in the mid-plane of the galaxy, and at these distortions gas will begin to accumulate (Parker, 1966; Dobbs et al., 2013). The second instability responsible for col-

lapse can be determined by the amount of rotational shear present (McKee & Ostriker, 2007). If a strong rotational shear is present, such as in the inter-arm region of a spiral galaxy, then the accumulation of gas will be due to its shearing as it moves through the disk of a galaxy in a process known as swing amplification (McKee & Ostriker, 2007; Dobbs et al., 2013). If no rotational shear is present, such as in the inner regions of a galaxy or spiral arms, then the collapse can be attributed to a magneto-Jeans instability which will remove the angular momentum responsible for the rotational shear through magnetic fields in order for the normal Jeans instability to occur (Elmegreen, 1987; Kim & Ostriker, 2001).

1.1.2 Molecular Hydrogen Formation

Regardless of either collapse or coagulation, molecular hydrogen is being formed inside the cloud. The formation of molecular hydrogen can be summarized using either the two body reaction, three body reaction, formation using a free electron or proton, or surface formation (Krumholz, 2014). The two body formation scenario is the simplest reaction using two hydrogen atoms to produce molecular hydrogen,



However two body formation is not the major mechanism in the formation of H_2 due to the requirement of a forbidden photon that arises from the combination of hydrogen atoms in the ground state (Gould & Salpeter, 1963). If

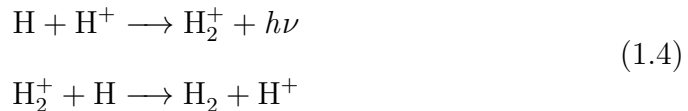
one of the hydrogen atoms is excited, the transition can occur and molecular hydrogen is formed, but the amount of excited hydrogen atoms in the temperature ranges typical of the cold and warm phases of the ISM are expected to be nearly nonexistent (Krumholz, 2014).

The second formation scenario listed, three body formation, involves three hydrogen atoms coming together to form molecular hydrogen with a spare hydrogen atom,



The required density for three body formation to occur is on the order of 10^8 cm^{-3} (Palla et al., 1983; Abel et al., 1997), while the typical GMC density is on the order of 300 cm^{-3} . The disparity between these densities eliminates any possibility of this being the primary mechanism to form molecular hydrogen in galaxies today.

An alternative to the two or three body reactions uses either an electron or proton to ionize the hydrogen forming either H^- or H_2^+ (Krumholz, 2014). The chemical reaction involving an electron is shown in equation 1.3, and the equation utilizing a proton is shown in reaction 1.4



The main limitation of the free electron/proton formation mechanism is an undersupply of free electrons and protons. Typical Milky Way conditions show that regions $>1 \text{ cm}^{-3}$ have free electron and proton densities $<10^{-4} \text{ cm}^{-3}$ (Wolfire et al., 2003). Secondly, the H^- ion in equation 1.3 is more likely to interact with a stray photon than an electron resulting in the ion returning to the atomic state (Glover, 2003). While we observe singly ionized hydrogen regions, HII regions, only a small fraction of the region will successfully produce molecular hydrogen.

While the free electron/proton formation method is suspected to be the primary H_2 catalyst in the early universe (Herbst et al., 2005), at low redshifts the formation of molecular hydrogen on the surface of dust grains is the dominant formation mechanism (Krumholz, 2014). Surface formation of molecular hydrogen will occur when a hydrogen atom strikes a dust grain and successfully sticks to the grain. The hydrogen will then interact with another hydrogen atom to form H_2 and be ejected from the dust particle after the reaction has occurred (Pirronello et al., 1997). The same process occurs on the dust grain as the two body formation in the gas phase, but the dust grain will act as a medium to absorb the energy that would create the forbidden photon (Krumholz, 2014).

With a dominant mechanism for molecular hydrogen formation, a reaction rate can be defined based on the cross section of the grain, Σ_{gr} , a striking probability dependent on the temperature of the observed dust, $S(T)$, the probability of molecular hydrogen forming on the grain, ϵ_{H_2} , the density of free hydrogen in the GMC, n_{H} , and the density of hydrogen attached to the grain

surface, n_{H_0} (Krumholz, 2014). Scaling this with the integrated collisional probability of a Maxwellian gas, the reaction rate is determined to be

$$\frac{dn_{H_2}}{dt} = \frac{1}{2} \left(\frac{8k_b T}{\pi m_H} \right)^{\frac{1}{2}} \Sigma_{gr} S(T) \epsilon_{H_2} n_H n_{H_0} \quad (1.5)$$

where k_b is the Boltzmann constant, T is the temperature, and m_H is the mass of the hydrogen atom. Equation 1.5 is often simplified to

$$\frac{dn_{H_2}}{dt} = \mathcal{R}_{gr} n_H n_{H_0} \quad (1.6)$$

by introducing a variable known as the formation rate, \mathcal{R}_{gr} , that is constrained using the column densities of CI, CII, HI, H₂. Using the column densities of CI and CII to determine a reaction rate is built around examining their ratios and the mechanisms that convert from ion to atom and atom to ion. The neutralization of CII involves its interaction with polycyclic aromatic hydrocarbons, and the ionization of C takes place by interacting with far-ultraviolet (FUV) photons (Wolfire et al., 2008). The FUV intensity will also dictate the amount of disassociation of H₂ into H, and by balancing the amount of CI and CII with the amount of H and H₂, the FUV intensity can be determined as well as the amount of H₂ being produced (Wolfire et al., 2008). Typical reaction rates for the Milky Way have been found to be $\mathcal{R}_{gr}=3 \times 10^{-17} \text{cm}^3 \text{s}^{-1}$ (Jura, 1975; Gry et al., 2002; Wolfire et al., 2008).

1.1.3 Dissociation of Molecular Hydrogen

When molecular hydrogen has formed, it is still susceptible to photodissociation from FUV photons. The energy required for a single photon to break the bonds of molecular hydrogen is 4.5 eV (Krumholz, 2014). Conveniently, this is also enough energy to excite atomic hydrogen, so photodissociation using a single photon is highly unlikely due to the abundance of HI in the ISM (Krumholz, 2014). However, a photon with an energy of 11-13.6 eV will not be able to ionize atomic hydrogen, but will be able to excite molecular hydrogen to its first and second excitation levels, the Lyman and Werner bands, respectively. The excited H_2 will eventually settle via photon emission to its ground state with a finite probability of returning to two hydrogen atoms rather than maintaining its molecular state (Krumholz, 2014).

A dissociation rate, ζ_{diss} , can be obtained by scaling the total excitation rate, ζ_{exc} , by the fraction of excited hydrogen molecules that will settle to an atomic ground state (Krumholz, 2014). The total excitation rate is found by summing each individual excitation from the ground state e.g. $\zeta_{exc,0-1}$, $\zeta_{exc,0-2}$. In the Milky Way's diffuse ISM, the interstellar radiation field is 6.9×10^{-14} erg cm $^{-3}$ over the range of 6-13.6 eV resulting in $\zeta_{exc} \approx 3 \times 10^{-10}$ s $^{-1}$ (Draine, 2011). The expected fraction of H_2 to dissociate is between 0.11 and 0.13 resulting in $\zeta_{diss} \approx 4 \times 10^{-11}$ s $^{-1}$ (Draine, 2011). Equating the formation and dissociation rates and then solving for a molecular hydrogen to atomic hydrogen ratio gives

$$\begin{aligned}
\zeta_{diss} n_{H_2} &= \mathcal{R} n_{H_0} n_H \\
\frac{n_{H_2}}{n_H} &= \frac{\mathcal{R} n_{H_0}}{\zeta_{diss}} \\
&= 8 \times 10^{-6} \left(\frac{4 \times 10^{-11} s^{-1}}{\zeta_{diss}} \right) \left(\frac{n_{H_0}}{10 cm^{-3}} \right)
\end{aligned} \tag{1.7}$$

(Krumholz, 2014). From equation 1.7 we can see, in the diffuse medium, atomic hydrogen is far more abundant than molecular hydrogen, as expected.

However, as the density of the gas increases via collapse or coagulation, the optical depth will also increase, limiting the number of FUV photons able to penetrate into the core of the cloud, a process known as shielding (Draine, 2011). This shielding will decrease the value of ζ_{diss} which will increase the n_{H_2}/n_H ratio in equation 1.7. The increase in the molecular to atomic hydrogen ratio signifies the accumulation of molecular gas reservoirs. The molecular gas build up will lead to fragmentation within the cloud, and will eventually lead to star formation.

1.2 Determining the H₂ Abundance

From the previous section, the formation of molecular hydrogen will result in the energy released during formation to be absorbed by a dust grain while the other methods resulted in the emission of a photon to account for the change in energy. The absorption of energy by the dust grain opposed to the energy being released through emission results in formation of H₂ being a dark process (§1.1.2). Furthermore, since the molecule is made of two hydrogen atoms, and its symmetry and low mass means it has no permanent dipole

moment making low energy rotational transitions nonexistent (Bolatto et al., 2013; Kennicutt & Evans, 2012). The high symmetry and low mass does not mean H_2 is unexcitable, but that the temperatures required to excite molecular hydrogen ($T \gtrsim 100K$) are above the temperature of a typical GMC (Bolatto et al., 2013). Also the returning of H_2 to the ground state will more often than not result in the molecule separating into two hydrogen atoms (§1.1.3). For our purposes, we can consider molecular hydrogen a dark molecule requiring special treatment to determine the amount present.

The amount of molecular hydrogen in a system can be calculated in several ways, and for the purpose of extragalactic sources they all involve using a molecular tracer to determine the amount of H_2 present. The molecule most commonly used in extragalactic studies is CO due to its abundance and ability to be easily observed (Bolatto et al., 2013). Other tracers that have been used are CO's photodissociated counterpart, CII, in order to trace regions that have little to no CO emission (Madden et al., 1997), and molecules such as OH that have been used in the past but have been limited to Milky Way targets (Barrett, 1964). The amount of CO can then be converted to H_2 using a conversion factor given as either

$$n_{H_2} = X_{CO} I_{CO} \quad (1.8)$$

where n_{H_2} is the column density, X_{CO} is the conversion factor in units of $cm^{-2} (K km s^{-1})^{-1}$, and I_{CO} is the CO intensity in units of $K km s^{-1}$. Typical values for X_{CO} in normal spiral galaxies tend to be around $1-4 \times 10^{20} cm^{-2} (K km s^{-1})^{-1}$, and a commonly used value for the Milky Way is $2 \times 10^{20} cm^{-2} (K$

$\text{km s}^{-1})^{-1}$ (Bolatto et al., 2013). Alternatively, the conversion factor can be used to determine the surface density of the molecular hydrogen, Σ_{H_2} , using

$$\Sigma_{H_2} = \alpha_{CO} I_{CO} \quad (1.9)$$

where α_{CO} is given in units of $M_\odot \text{ pc}^{-2} (\text{K km s}^{-1})^{-1}$. Values of α_{CO} for nearby extragalactic sources have shown a mean value of $3.1 M_\odot \text{ pc}^{-2} (\text{K km s}^{-1})^{-1}$ which is slightly lower than the assumed Milky Way value of $4.4 M_\odot \text{ pc}^{-2} (\text{K km s}^{-1})^{-1}$ (Sandstrom et al., 2013).

1.2.1 Methods for Determining CO-to-H₂ Conversion Factor

Several methods are available to determine the conversion factor for extragalactic sources, and each has its own caveats. One common method used to determine the conversion factor is to use the virial nature of GMCs (Bolatto et al., 2013). This method works best for well defined clouds; however in the case of more distant nearby galaxies, the issue of whether or not giant molecular associations (GMAs) display the same virialization as their constituent GMCs can limit concrete results (Bolatto et al., 2013). The second caveat to using the virial mass to determine the conversion factor is the method will only trace CO bright regions. Tracing only CO bright regions produced results that show no dependence between the conversion factor and metallicity, while other methods display a negative correlation between the conversion factor and metallicity (Bolatto et al., 2013). The lack of correlation with metallicity is believed to be due to the level of dust shielding present in the system. If

the amount of dust present in the region is large, then the region will display a large metallicity, and if little to no dust is present, the region will show a low metallicity. Regions with low metallicity will correspond to regions with little to no dust resulting in poor shielding. The lack of shielding will allow the CO to be dissociated resulting in large amounts of molecular hydrogen traced by ionized and atomic carbon gas instead of CO (Bolatto et al., 2013). While the virial method is suitable for determining conversion factors within CO bright regions, excluding any low metallicity regions will underestimate the total amount of H_2 present in the system as well as skew any possible relations of the conversion factor with metallicity (Bolatto et al., 2013).

A second method to determine the conversion factor is to incorporate observations of isotopologues of CO that are optically thin, commonly ^{13}CO (Bolatto et al., 2013). The temperature, density, and column or surface density of ^{13}CO can be used to restrict the possible outcomes of the physical conditions of the GMC or GMA being examined. The main problem with using this method is the degeneracy between the temperature and density of CO emission such that hot low density regions can resemble cool dense regions (Rosenberg et al., 2014). This method also shares the same problem as the virial technique in that it only probes CO bright regions of the target leaving any CO-faint H_2 gas untraced (Bolatto et al., 2013). The issue of CO-faint molecular hydrogen has been examined using the photodissociated tracer of CO, CII, instead of ^{13}CO . The CII ion was used as a tracer of molecular hydrogen in dwarf galaxies and was able to reveal large reservoirs of self-shielding H_2 (Madden et al., 1997).

The third way to determine the CO-to-H₂ conversion factor is by incorporating the emission from dust to determine the amount of molecular gas present. This method assumes that the gas is well mixed and has a constant ratio of dust mass to gas mass present in the galaxy (Leroy et al., 2011), which has been shown to be true for the Milky Way (Boulanger et al., 1996). A suitable conversion factor is found by solving

$$\begin{aligned}\delta_{GDR}\Sigma_{dust} &= \Sigma_{H_2} + \Sigma_{HI} \\ &= \alpha_{CO}I_{CO} + \Sigma_{HI}\end{aligned}\tag{1.10}$$

where Σ_{dust} , Σ_{H_2} , and Σ_{HI} are the respective surface densities in $M_\odot \text{ pc}^{-2}$, I_{CO} is the CO line intensity, α_{CO} is the conversion factor in $M_\odot \text{ pc}^{-2} \text{ K}^{-1} \text{ km}^{-1} \text{ s}$, and δ_{GDR} is the total mass of the gas divided by the total mass of the dust known as the gas-to-dust ratio (Leroy et al., 2011; Sandstrom et al., 2013). In equation 1.10, we can measure Σ_{dust} , Σ_{HI} , and I_{CO} leaving only the conversion factor and gas-to-dust ratio free to vary. An appropriate α_{CO} value will generate a molecular gas mass that produces a constant gas-to-dust ratio over the galaxy being studied. This method has been carried out extensively by Sandstrom et al. (2013) on kpc scales on nearby galaxies as well as on both of the Magellenic clouds by Leroy et al. (2011).

Determining a conversion factor in via the dust emission provides the capability to trace the CO faint regions of a target (Israel et al., 1996). Despite this advantage over the other two methods, using equation 1.10 leaves any gas not associated with atomic hydrogen to be treated as H₂. This is shown in equation 1.10 given that the total amount of gas, is $\alpha_{CO}I_{CO} + \Sigma_{HI}$, so any gas

not associated with atomic hydrogen is assumed to be molecular hydrogen. This means that the calculated amount of H_2 will include contributions from CO, OH, and CII as well as any other chemical species present in the region. This will effectively increase the conversion factor and the overall amount of molecular hydrogen reported (Bolatto et al., 2013), but given the dominance of H_2 in the dense regions of the ISM, any increase will be negligible. Another caveat of this method rests with the assumption of a constant dust-to-gas ratio (Bolatto et al., 2013). Issues such as the gas-to-dust ratio's dependence with metallicity (Draine et al., 2007) can render this assumption null if not treated properly. Nevertheless, the agreement of conversion factors between the dust based method and other methods used for more local targets suggests that any gas present that has been incorrectly assumed to be H_2 and local fluctuations in the metallicity of the target galaxy have very little effect on the final results (Bolatto et al., 2013).

1.3 Determining Dust Mass

Calculating a conversion factor using dust emission requires a knowledge of the amount of dust present in our system to determine a gas-to-dust ratio. Given that a significant portion of the dust mass lies within the cold phase we can calculate a dust mass using a modified blackbody fit over the cold portion of the dust's spectral energy distribution (SED) (Galametz et al., 2012). A slight modification needs to be introduced to our blackbody resulting in what is known as a greybody or modified blackbody (MBB). The modification is necessary because the dust does not absorb and re-emit all of the light incident

on its surface. If we assume an isolated optically thin medium ($\tau_\nu \ll 1$) with a blackbody source function, $B_\nu(T)$, the radiative transfer equation can be written as

$$I_\nu = (1 - e^{-\tau}) B_\nu(T), \quad (1.11)$$

and simplified using a first order Taylor expansion with $\tau_\nu \ll 1$ to get

$$I_\nu = \tau_\nu B_\nu(T) \quad (1.12)$$

The optical depth can be expanded to incorporate the surface density of the dust, Σ_{dust} , as

$$\tau_\nu = \kappa_\nu \Sigma_{dust} \quad (1.13)$$

such that the units on Σ_{dust} are in kg m^{-2} opposed to $\text{M}_\odot \text{pc}^{-2}$. The dust emissivity cross-section per unit mass is represented by κ_ν and is often referred to as the opacity. It is important to note that the opacity will reflect the chemical makeup and grain structure of the dust, but will not indicate the grain size (Abergel et al., 2013). The behavior of the opacity has been well fit with a power law such that opacity shows a dependence with the frequency of the observed emission (Hildebrand, 1983). The opacity is typically written as

$$\kappa_\nu = \kappa_{\nu,0} \left(\frac{\nu}{\nu_0} \right)^\beta \quad (1.14)$$

where ν_0 is a reference frequency, $\kappa_{\nu,0}$ is a reference opacity based on the reference frequency, and β is the dust emissivity index. Values for κ_ν have been calculated for several models of varying dust compositions and can range from 0.2 to 2 m² kg⁻¹ over the wavelengths used to model the cold component (100 μ m–850 μ m)(Li & Draine, 2001). The emissivity index has commonly been fit between ranges of 1.0 to 2.0 suggested from laboratory experiments (Walcher et al., 2011). However, with the influx of Herschel data the ability to fill in the sub-millimeter portion of the SED between 100 μ m and 500 μ m has allowed the emissivity index to be determined based on observations, which have shown a range from 1 to 2.5 (Galametz et al., 2012). It should be noted that the lower limit of $\beta=1$ is a hard cutoff due to limits determined by the Kramers-Kronig relationship in the optically thin case (Li, 2005), but in extreme instances when the dust becomes optically thick β will approach 0.

With the opacity well defined, substituting equation 1.14 into equation 1.13 and then into the original equation for the specific intensity of our MBB, equation 1.12, gives the formula for the specific intensity of a greybody as

$$I_\nu = \Sigma_{dust} \kappa_{\nu,0} \left(\frac{\nu}{\nu_0} \right)^\beta B_\nu(T) \quad (1.15)$$

Since we measure the flux of our target, it is useful to convert from the specific intensity to the flux using

$$F_\nu = \pi I_\nu \left(\frac{r}{D} \right)^2 \quad (1.16)$$

given the target is a uniform sphere with radius, r , and distance, D (Rybicki & Lightman, 1986). Substituting equation 1.16 into the greybody equation, 1.15, gives

$$S_\nu = \frac{\Sigma_{dust} \pi r^2}{D^2} \kappa_\nu \left(\frac{\nu}{\nu_0} \right)^\beta B_\nu(T). \quad (1.17)$$

Converting from surface density to overall mass is done using the relationship between mass and surface density of $M = \Sigma * \text{Area}$ where the target's projected area is πr^2 . This gives a final modified blackbody equation of

$$S_\nu = \frac{M_{dust}}{D^2} \kappa_\nu \left(\frac{\nu}{\nu_0} \right)^\beta B_\nu(T). \quad (1.18)$$

The dust mass can be solved from equation 1.18 by fitting the mass, temperature, and possibly the emissivity index. The fitted mass will act as a normalizing factor for the SED and correspond with the mass associated with the peak flux. The fitted dust mass should be used cautiously due to a sensitivity in mass with temperature fluctuations (Draine et al., 2007). An alternative to using the fitted mass is to isolate the mass from equation 1.18 to get

$$M_{dust} = \frac{S_\nu D^2}{\kappa_\nu B_\nu(T)} \left(\frac{\nu}{\nu_0} \right)^{-\beta} \quad (1.19)$$

and use the parameters returned from fitting. This equation is best used with a longer wavelength observation that will explore the Rayleigh-Jeans tail of the modified blackbody. This portion of the SED will see less of a dependence on the mass and temperature fluctuations (Draine et al., 2007). The main concern with both methods to determine the mass arises from fitting both

the temperature and dust emissivity index and their strong anti-correlation (Galametz et al., 2012; Tabatabaei et al., 2014). A common method to help break the degeneracy is to fix the emissivity index to a reasonable value based upon the opacity model (Tabatabaei et al., 2014).

1.4 NGC 3627

In this thesis, we are using the new SCUBA-2 instrumentation on the JCMT, which observes both $450\mu\text{m}$ and $850\mu\text{m}$ emission simultaneously. Utilizing multiple 60 minute scans in high quality weather we can better constrain the emissivity index in our fits, as well as increase the resolution compared to previous work. With these new observations from the NGLS, we can better fit an SED and determine a more reliable dust mass to explore the amount of molecular gas present. To take full advantage of the benefits of SCUBA-2, we have selected the target NGC3627 for its relatively small angular size, pronounced spiral features, and relatively large flux.

NGC 3627 is the most prominent member of the interacting trio of galaxies called the Leo triplet (NGC 3623, NGC 3627, and NGC3628). An optical image is shown in Figure 1.1 while the data we used can be seen in sections §2.2 and §2.3. The interaction between NGC 3627 and NGC 3628 has been well documented via HI emission in the form of a large HI debris tail extending from NGC 3628 in the direction of NGC 3627 (Rots, 1978; Haynes et al., 1979). The interaction between these galaxies has resulted in a high star formation rate of over $1.7 \text{ M}_{\odot} \text{ yr}^{-1}$ (Calzetti et al., 2010), making NGC 3627 a prime candidate to study the current star forming conditions. The morphology of NGC 3627 is

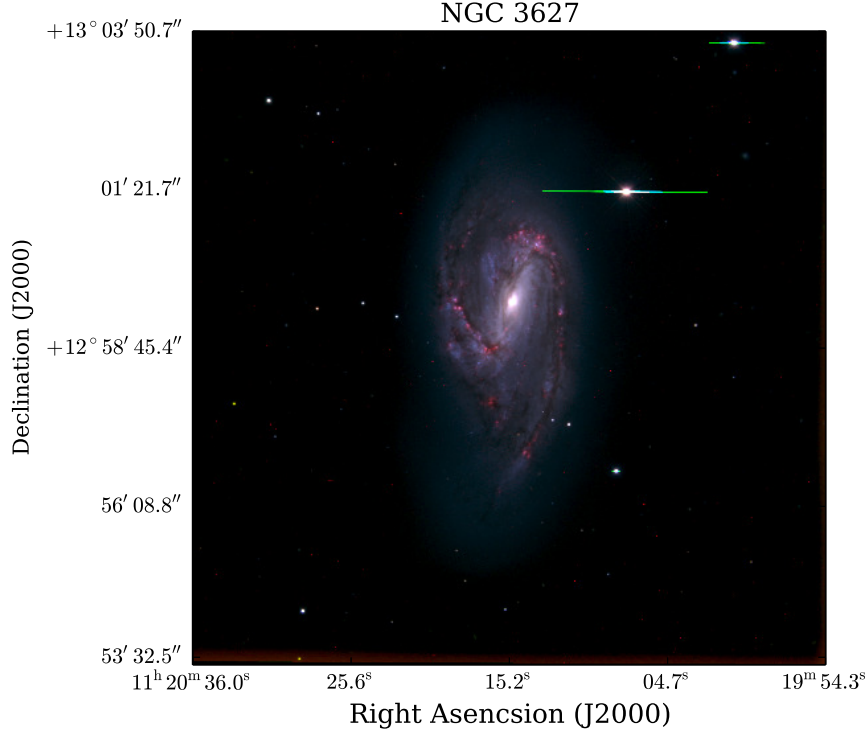


Figure 1.1: Optical composite image of NGC3627 made using the SINGS 5th enhanced data release (Kennicutt et al., 2003). Blue areas represent the B band, green areas represent V band and red areas show the H α band. Component images were retrieved from NED.

that of a barred spiral (SABb) with an inclination of 60° (Reuter et al., 1996) at a distance of 9.4 Mpc determined by Cepheid variable observations (Freedman et al., 2001). Furthermore, the strong spiral arm features are unique to this galaxy in that speculation exists they are not contained within the plane of the galaxy. This interpretation is supported by magnetic field lines traced by dust polarization that do not trace the spiral structure in the southeastern bar end (Soida et al., 2001) and CO observation with bimodal emission features in velocity space over the southeastern bar end (Dumke et al., 2011).

NGC 3627 is a member of the Spitzer Infrared Nearby Galaxies Survey (SINGS) (Kennicutt et al., 2003) and has been observed throughout the electromagnetic spectrum. ^{13}CO observations by Watanabe et al. (2011) have shown that a majority of the star formation is occurring in the bar ends. This information corresponds with Warren et al. (2010) showing the star formation efficiency being highest in the bar ends. Warren et al. (2010) also show the ISM of the galaxy to be dominated by molecular gas, with dense warm gas dominating the emission at the bar ends, the nucleus and a bright region located beneath the southwestern spiral arm, and more diffuse and cooler molecular gas outside of these regions. The atomic gas is located primarily in the spiral arms with little to no emission in the nucleus of the galaxy. A mean conversion factor of $\alpha_{CO}=1.2 \text{ M}_{\odot} \text{ pc}^{-2} (\text{K km s}^{-1})^{-1}$ was found by Sandstrom et al. (2013) using the dust method (§1.2.1). An HI mass was given as $8.18 \times 10^8 \text{ M}_{\odot}$ (Walter et al., 2008), and a corresponding H_2 mass was calculated as $5.79 \times 10^9 \text{ M}_{\odot}$ (Kennicutt et al., 2003), and a stellar mass of $2.8 \times 10^{10} \text{ M}_{\odot}$ (Skibba et al., 2011).

The dust emission of NGC 3627 follows the same trend as the CO emission with the brightest regions at the bar ends, nucleus, and a bright region below the southwest bar end. The southeastern spiral arm shows some knotted features in the dust emission. Global values for the cold and warm components of the SED were fit by Galametz et al. (2012) and reveal dust temperatures of $T_W=55.8 \pm 5.6 \text{ K}$ and $T_C=20.2 \pm 1.4 \text{ K}$ for the warm and cold components, respectively, with a best fit emissivity index of $\beta=2.3 \pm 0.2$ and a total dust mass of $7.82 \times 10^7 \text{ M}_{\odot}$.

This thesis will explore the properties of dust emission from NGC3627 by exploring its SED using new observations from SCUBA-2 in order to measure the mass and dust emissivity index at a higher resolution in order to build on the work of Galametz et al. (2012). Furthermore, we can use our SED derived dust masses to determine the overall amount of molecular gas present in the galaxy while gaining insight into the dust-to-gas ratio and conversion factor for this galaxy using the methods outlined by Leroy et al. (2009) and Sandstrom et al. (2013). The information in this thesis is arranged as follows: Chapter 2 will describe the data reduction and properties of the SCUBA-2 data and any ancillary data, Chapter 3 will describe the SED fitting method, Chapter 4 will go over the method to determine a dust-to-gas ratio, and Chapter 5 will summarize our findings.

Chapter 2

Observations and Data Preparation

2.1 SCUBA-2

The Submillimetre Common-User Bolometer Array 2 (SCUBA-2) was designed to decrease the observing time of the sub-millimeter sky relative to its predecessor SCUBA (Holland et al., 2013). This would benefit the community by allowing for rapid data acquisition in the submillimeter regime of the electromagnetic spectrum, at the $450\mu\text{m}$ and $850\mu\text{m}$ bands in particular. Prior to SCUBA-2, other bolometer cameras such as LABOCA, BOLOCAM and SHARC-II were limited to less than 100 pixels, while the new SCUBA-2 has been able to incorporate over 10,000 pixels in its design and effectively reduce the required observing time (Chapin et al., 2013). Increasing the amount of pixels by a factor of 100 was possible by the advent of new technology such as high precision micromachining, superconducting transition edge sensors, and superconducting quantum interference device amplifiers (SQUIDs) (Holland et al., 2013).

The observations of NGC3627 were taken from the Nearby Galaxies Legacy Survey’s (NGLS) initial science images using SCUBA-2 from December 29, 2011 to January 21, 2012, and consist of 24 $18'$ by $18'$ scans taken in grade 3 weather ($0.08 < \tau < 0.12$) or better with observations centered at $450\mu\text{m}$ and $850\mu\text{m}$ emission with a $32\mu\text{m}$ and $85\mu\text{m}$ bandpass respectively. 16 of the 24 scans were deemed useable, and whether or not an observation was deemed worthwhile was determined by factors such as the behavior of the image background or whether the image was flagged during observing to be unusable. The observations of NGC3627 were taken using a DAISY scanning pattern to help remove any random noise by introducing crossing points. The scanning speed of the JCMT was $150''/\text{second}$ in order to reduce any drifting effects seen from the instrument or sky (Chapin et al., 2013).

2.2 Image Creation and Properties

For any imaging process to be successful, the final image needs to have a limited white noise (Chapin et al., 2013). White noise in the sense of our bolometer observations arises from thermal variations in the instrument and atmosphere during data acquisition. The random noise is minimized through scanning methods and during image processing (Chapin et al., 2013). To create the final SCUBA-2 data products, we use the Submillimetre User Reduction Facility (SMURF) procedure MAKEMAP. This procedure reduces the noise of the observations while maintaining the source’s emission by incorporating a combination of principal component analysis and a maximum likelihood analysis (Chapin et al., 2013). Both of these methods have proven useful in

reducing bolometer data on their own, but due to the size of raw SCUBA-2 data, either method on its own would result in extreme run times or the process becoming resource intensive (Chapin et al., 2013).

MAKEMAP breaks down the image creation into several steps performed in iteration in order to successfully reduce any background noise (Chapin et al., 2013). The raw data is modeled by

$$b_i(t) = f [e_i(t) * a_i(t) + g_i * n_c(t) + n_f(t) + n_r(t)] \quad (2.1)$$

such that $b_i(t)$ is the i^{th} bolometer output at time t , f is a scaling factor determined from flat field calibrations, $e_i(t)$ is the extinction at time t for the i^{th} bolometer, $a_i(t)$ is the astronomical signal for i^{th} bolometer at time t , $g_i * n_c(t)$ is the noise due to the product of the common mode and gain, $n_f(t)$ is the low-frequency noise, and $n_r(t)$ is the random noise (Chapin et al., 2013).

The steps used in MAKEMAP are designed so they can estimate each noise component of the bolometer signal. The COM and GAI steps estimate a common mode signal by fitting the average time-series of each bolometer observation. The next step is EXT, which is used to apply the measured extinction corrections. Following EXT is FLT to apply high- and low-pass filters to remove any noise features not removed in the COM and GAI filtering. After the high- and low-pass filters, the AST step is performed. AST converts the data from a time series to an image and detects sources to be removed from reduction. The sources removed during this portion of MAKEMAP are intended to be astronomical in nature. The final step is NOI which determines the noise in the gridded map after each step has been performed and is calcu-

lated by isolating the white noise component in equation 2.1. A convergence check is then done based on the magnitude of change in each pixel from the previous map and the current version. If the check fails, the COM, GAI, EXT, and FLT values are recalculated using the AST information obtained from the previous iteration, and the process is repeated until the convergence values are met or the maximum number of iterations has been carried out (Chapin et al., 2013).

In our production of maps, we used the configuration file `dimmmconfig_bright_compact.lis` and altered the AST and FLT sections of the image creation by introducing a mask made from Herschel’s $250\mu\text{m}$ map. The purpose of the mask was to exclude the target from interfering with the noise minimization as well as to keep any emission from the galaxy from being significantly altered during image production. The filter size of the high-pass filter was also modified, and an appropriate value was determined to be $175''$. We determined an appropriate high-pass filter size by running a large range of filters from $100''$ to $300''$ and inspecting the total recovered flux of NGC3627. A plot of the returned flux values can be seen in Figure 2.1. A good filter would not show any significant decrease in flux compared to the $350\mu\text{m}$ or $500\mu\text{m}$ fluxes. This requirement removes any filters greater than $200''$. Given the flux was nearly constant with filters less than $200''$, we chose an appropriate filter size by examining the structure that was returned, in particular how well the spiral arms were recovered for the $850\mu\text{m}$ image, and how well the disk was preserved in the $450\mu\text{m}$ image. The spatial effect of the filters can be seen in Figure 2.2 for $450\mu\text{m}$ and Figure 2.3 for $850\mu\text{m}$. The maps were returned from MAKEMAP in units of pW with a pixel size of $2''$ by $2''$ for both the $450\mu\text{m}$ and $850\mu\text{m}$.

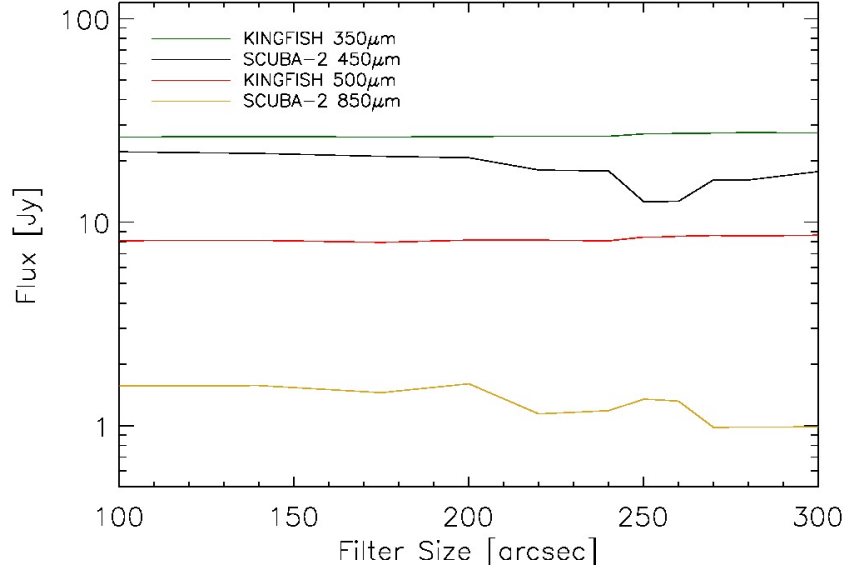


Figure 2.1: Returned flux values for NGC3627 with varying high-pass filter sizes. The KINGFISH fluxes have been processed using the fakesource process, §2.4.2

The final $450\mu\text{m}$ image was then re-gridded down to a $4''$ by $4''$ pixel grid, and flux calibration values of 491000 and 4710 (Dempsey et al., 2013) were applied to convert from pW to mJy/beam and mJy/square arcsecond, respectively. The $850\mu\text{m}$ maps were re-gridded to an $8''$ by $8''$ pixel size and used flux calibration values of 537000 and 2340 for mJy/beam and mJy/square arcsecond (Dempsey et al., 2013). The $4''$ and $8''$ pixels correspond to a 180pc and 360pc size scale for our target, NGC3627. To simplify the analysis, the images are also converted to Jy/pixel. The $450\mu\text{m}$ and $850\mu\text{m}$ images are shown in Figures 2.4 and 2.5. The calibration values used are the default flux calibration factors that are determined from our calibrator source, Uranus, in (Dempsey et al., 2013). The overall noise in the final image can be seen in Table 2.1.

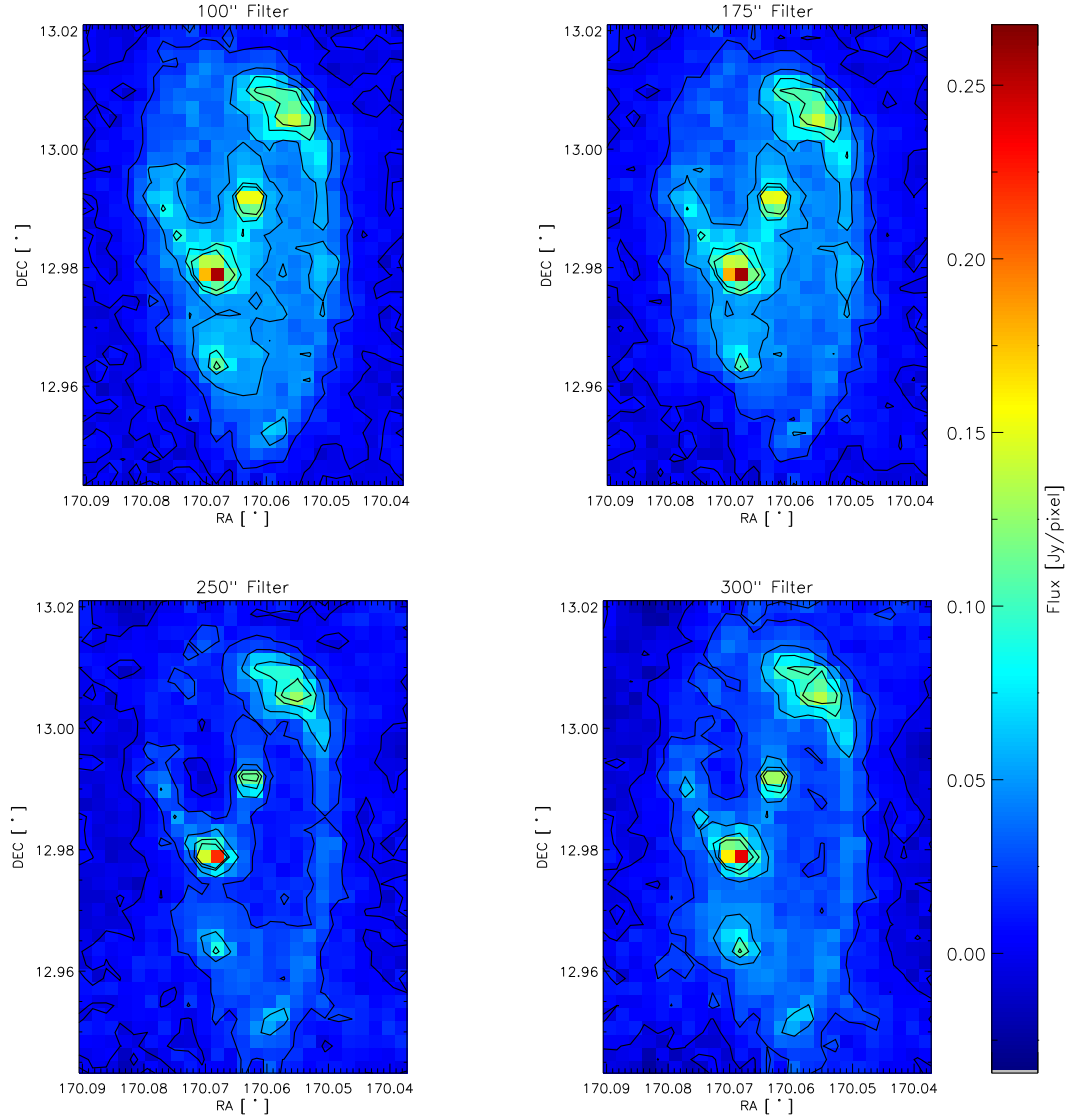


Figure 2.2: Four 450 μm maps of NGC3627 using varying high-pass filter sizes. The contours shown are for 0.0, 0.02, 0.05, 0.08 and 0.1 Jy/pixel for each image with 8'' by 8'' pixels.

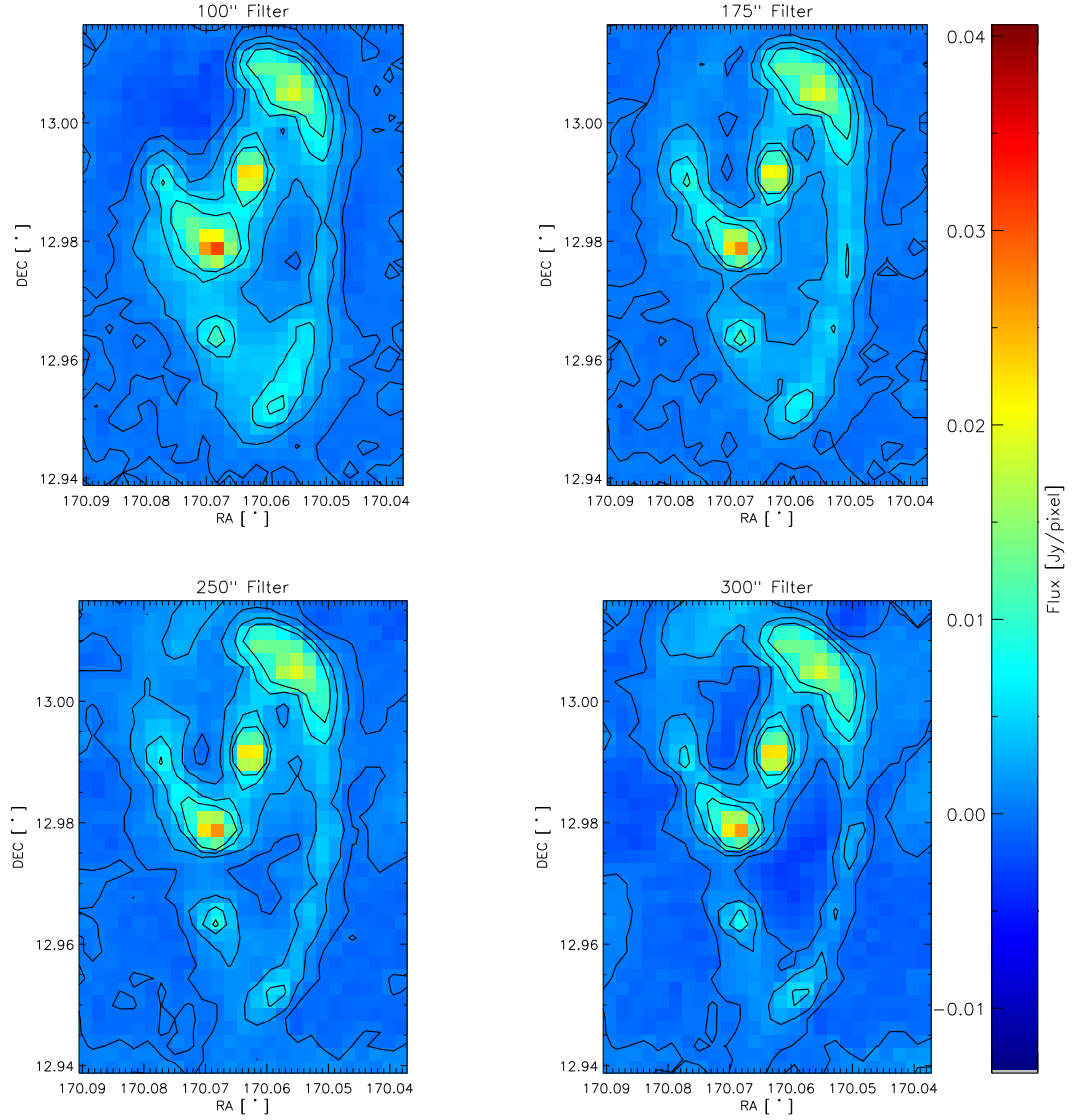


Figure 2.3: Four 850 μm maps of NGC3627 using varying high-pass filter sizes. The contours shown are for 0.0, 0.002, 0.005, and 0.008 Jy/pixel for each image with 8'' by 8'' pixels.

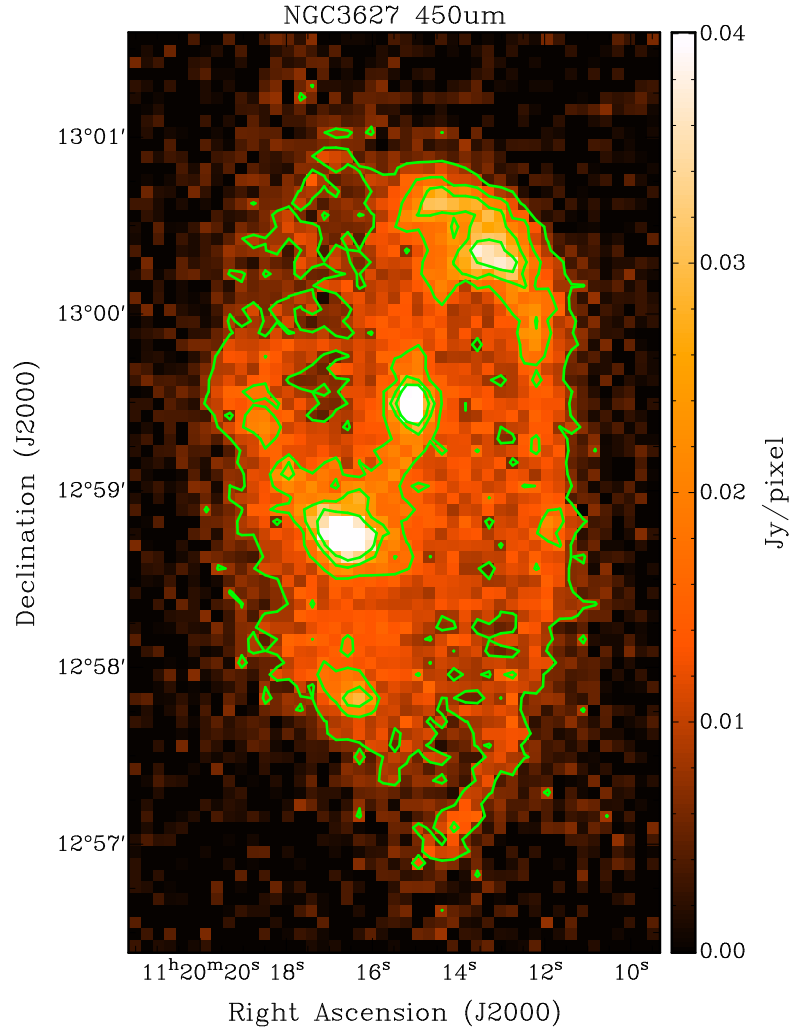


Figure 2.4: 450 μ m observation produced at the end of the image production with 20%, 40%, 60%, and 80% contours with 4'' by 4'' pixels.

Table 2.1. Properties of NGC3627 SCUBA-2 Observations

Observation	Beam Properties				RMS [mJy / Pixel]
	α	θ_α	β	θ_β	
450 μ m	0.854 ± 0.002	$7.48'' \pm 0.03''$	0.146 ± 0.003	$23.1'' \pm 0.2''$	3.42
850 μ m	0.9624 ± 0.0002	$12.8'' \pm 0.004''$	0.0376 ± 0.0002	$44.5'' \pm 0.09''$	0.476

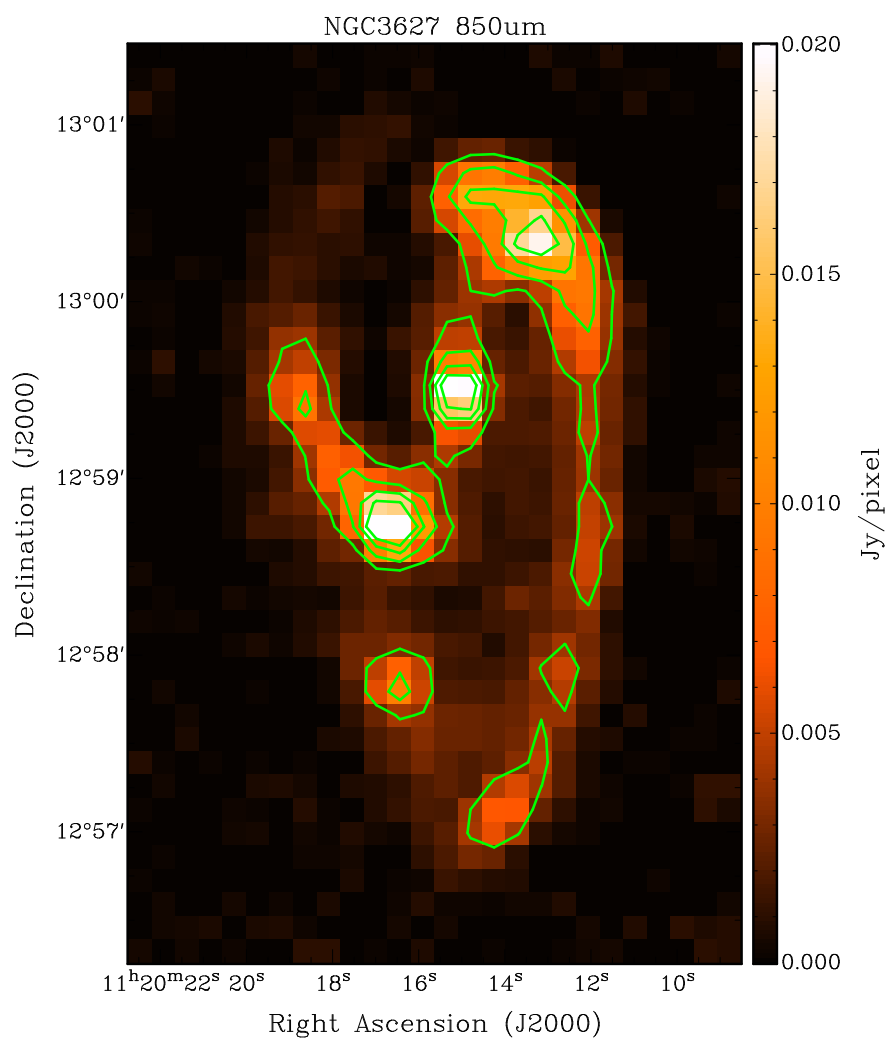


Figure 2.5: $850\mu\text{m}$ observation produced at the end of the image production with 20%, 40%, 60%, and 80% contours with $8''$ by $8''$ pixels.

2.2.1 Beam Shape of the $450\mu\text{m}$ and $850\mu\text{m}$ Data

The Uranus calibration images were used to determine the shape of the beam for the $450\mu\text{m}$ and $850\mu\text{m}$ observations. The beam shape of both the $450\mu\text{m}$ and $850\mu\text{m}$ maps deviates from a single gaussian due to the second maximum of the airy diffraction pattern in the response function of the telescope and minor imperfections in the mirror of the JCMT due to boundaries of the panels (Dempsey et al., 2013). This abnormality is best represented by a sum of two gaussians whose amplitudes sum to unity (Dempsey et al., 2013). The average beam resolution for the $450\mu\text{m}$ and $850\mu\text{m}$ data are reported in Table 2.1 and within an acceptable difference from the values found in Dempsey et al. (2013). The calibration images, fitted beams, and the residual of the fits can be seen in Figure 2.6. The residuals seen in here are due to asymmetries in the Uranus PSF. The contribution of the error beam in the $850\mu\text{m}$ emission is negligible, and allows the beam to be approximated by a single gaussian. However, the contribution of the error beam in the $450\mu\text{m}$ images was large enough to require special treatment in order to properly match the beams for analysis.

2.3 Ancillary Data

The scientific goals of this thesis require data outside the capabilities of SCUBA-2. For instance, accurately determining the dust mass involves fitting the spectral energy distribution (SED) for NGC3627. To successfully fit an SED, we need shorter wavelength data to fully probe the cold component of this galaxy. We used data ranging from $100\mu\text{m}$ to $500\mu\text{m}$ from the KINGFISH

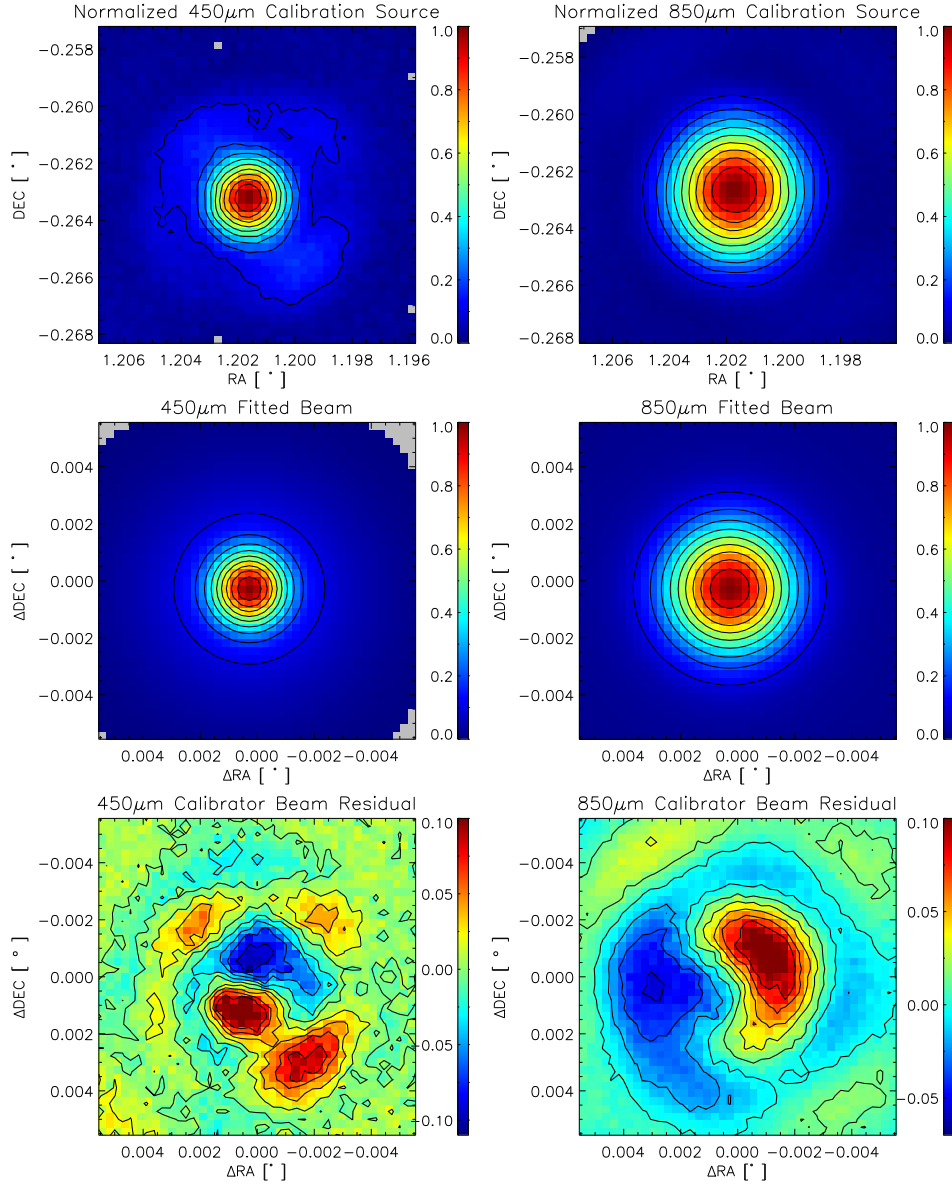


Figure 2.6: The top row shows the Uranus images taken on January, 8th 2012 for 450 μ m on the left and 850 μ m on the right. The middle row shows the fitted beams for 450 μ m on the right and 850 μ m on the left using the double gaussian beam shape. The bottom row shows the difference between the original image and the fitted beam. The contours in the image are from 10% to 90% in intervals of 10%.

survey (Kennicutt et al., 2011) to gain a large enough wavelength range for fitting the cold component. Secondly, the bandpass of the $850\mu\text{m}$ filter contains the CO J=3-2 line. In order to get a valid measurement of the dust mass, this contribution had to be removed. We used emission data from the NGLS from the HARP instrument on the JCMT (Wilson et al., 2012) to remove this contamination. When a dust mass was obtained, we used CO J=1-0 from the Nobeyama 45-m telescope (Kuno et al., 2007), CO J=2-1 from HERACLES (Leroy et al., 2009), and *HI* observations from THINGS (Walter et al., 2008) to determine a reasonable molecular hydrogen mass to calculate a dust-to-gas ratio. In the next sections we give some details for these ancillary data sets.

2.3.1 Key Insights on Nearby Galaxies: a Far-Infrared Survey with Herschel (KINGFISH)

The Key Insights on Nearby Galaxies: a Far-Infrared Survey with Herschel (KINGSFISH) was designed to be a follow up to the Spitzer Infrared Nearby Galaxies Survey (SINGS) (Kennicutt et al., 2003) with observations of the warm and cold component of dust emission using the increased resolution from Herschel (Kennicutt et al., 2011). The main science goals of the KINGFISH survey were to better understand the star formation processes that were shielded by dust, to make resolved studies of heating and cooling of the interstellar medium (ISM), and to build an inventory of how cold dust emission relates to other dust components in the ISM (Kennicutt et al., 2011). The survey consists of 61 nearby galaxies ($d < 30\text{Mpc}$) that cover a range of environments. Each target was observed at $70\mu\text{m}$, $100\mu\text{m}$, $160\mu\text{m}$, $250\mu\text{m}$, $350\mu\text{m}$,

Table 2.2. Properties of NGC3627 KINGFISH Observations

Observation	Beam Properties θ_{beam}	RMS $[mJy/Pixel]$	Percentage of Emission Removed
100 μ m	6.8''	2.24	11%
160 μ m	11.6''	3.95	17%
250 μ m	18.0''	2.47	20%
350 μ m	24.9''	1.08	21%
500 μ m	36.0''	0.387	28%

and 500 μ m. Our analysis focuses on fitting the cold component of NGC3627's SED, so we omitted the 70 μ m emission from the fitting, and processed the data through MAKEMAP as described below in § 2.4. The rms and beam size after the large scale structure has been removed can be seen in Table 2.2 given for 8'' by 8'' pixels, while the preconvolved maps are shown in Figures 2.7 to 2.11.

2.3.2 Nearby Galaxies Legacy Survey (NGLS)

The Nearby Galaxies Legacy Survey is an HI-selected set of 155 galaxies contained in the annulus of $2Mpc \leq d \leq 25Mpc$ observed using the instrumentation on the JCMT (Wilson et al., 2012). The NGLS consists of data observed in several wavelengths that include the 450 μ m and 850 μ m data used for this thesis. As mentioned previously, the bandpass for SCUBA-2's 850 μ m emission contains the CO J=3-2 line which is contained in the NGLS data set. We used the zeroth moment CO J=3-2 maps from the NGLS to determine the percentage of CO J=3-2 emission present in the 850 μ m band as well as

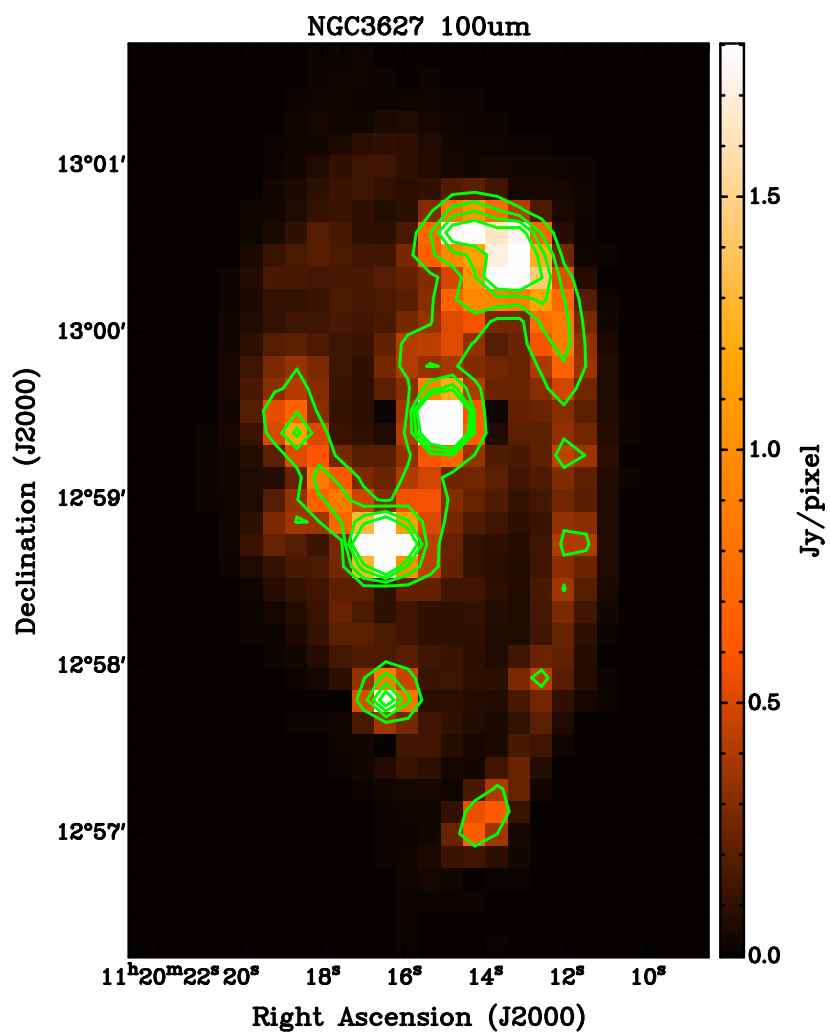


Figure 2.7: Image after the MAKEMAP filtering of $100\mu\text{m}$ observations with 20%, 40%, 60%, and 80% contours.

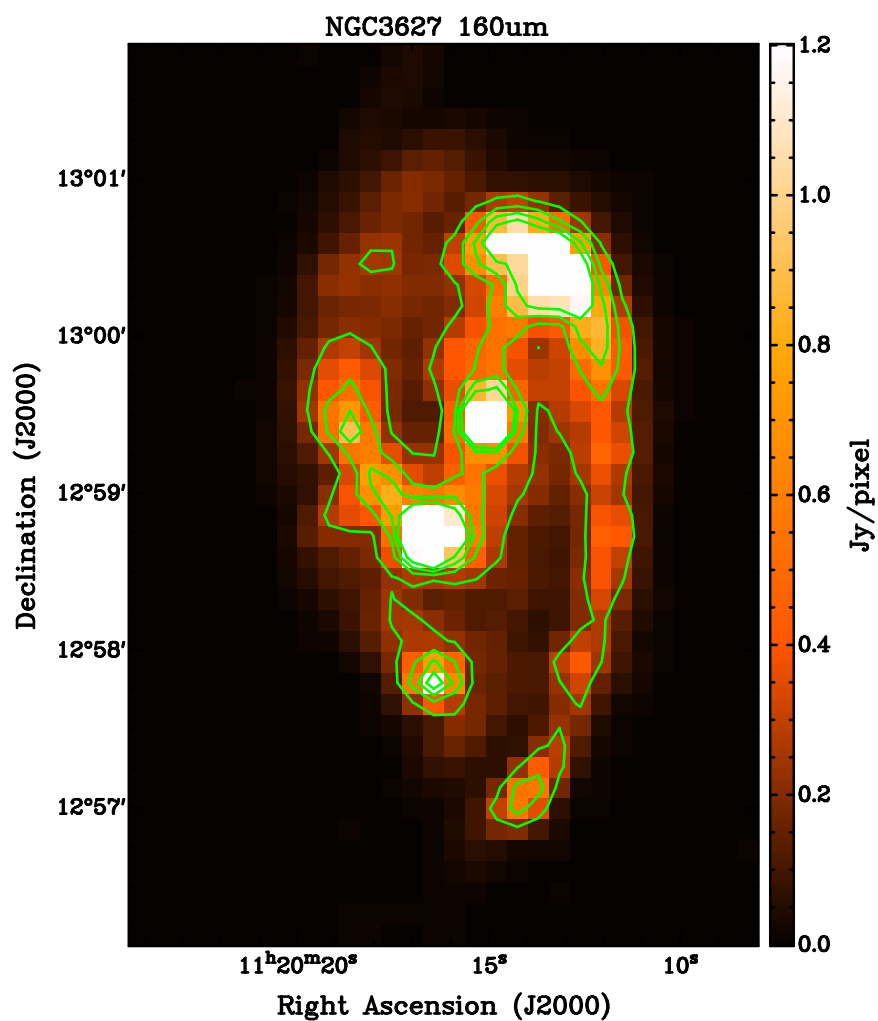


Figure 2.8: Image after the MAKEMAP filtering of $160\mu\text{m}$ observations with 20%, 40%, 60%, and 80% contours.

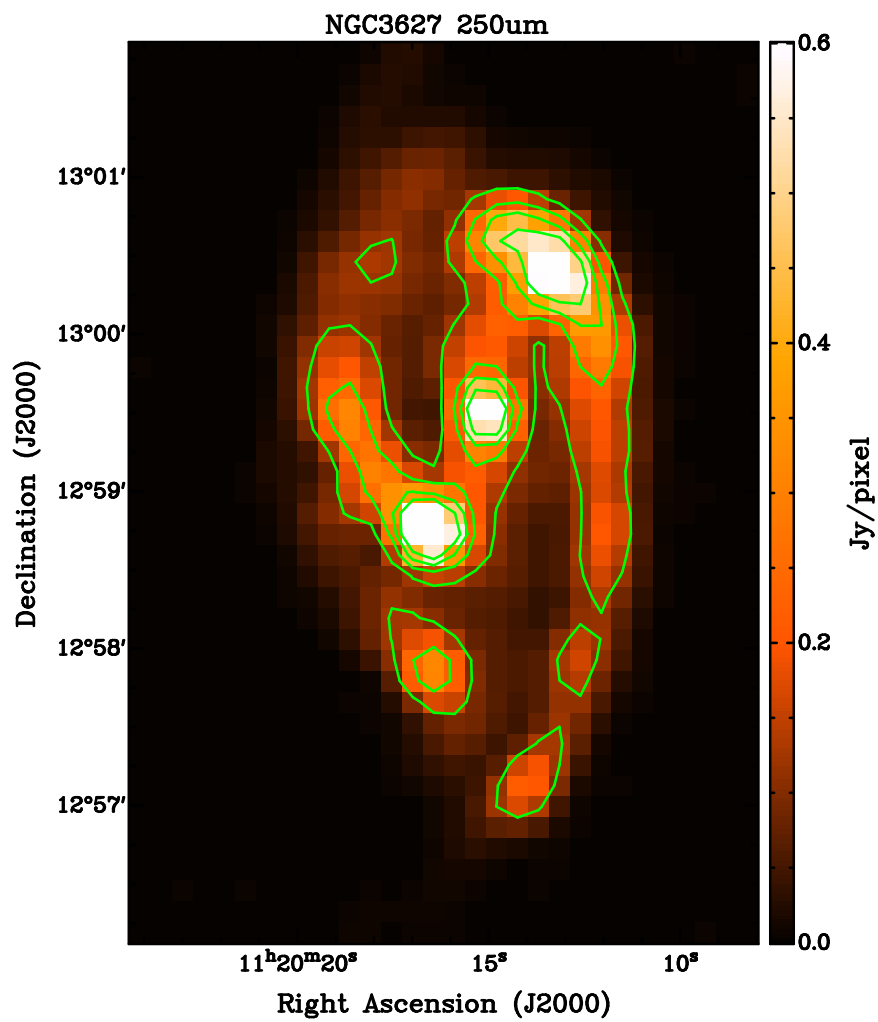


Figure 2.9: Image after the MAKEMAP filtering of 250μm observations with 20%, 40%, 60%, and 80% contours.

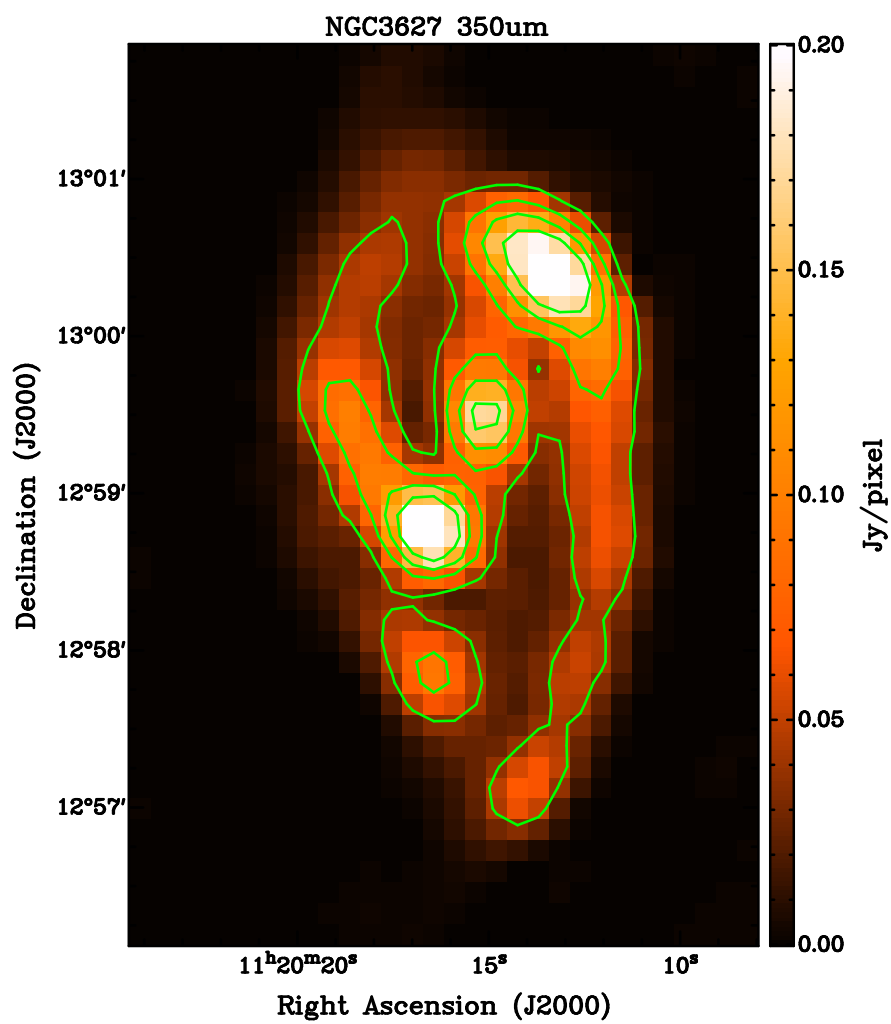


Figure 2.10: Image after the MAKEMAP filtering of the 350 μ m observations with 20%, 40%, 60%, and 80% contours.

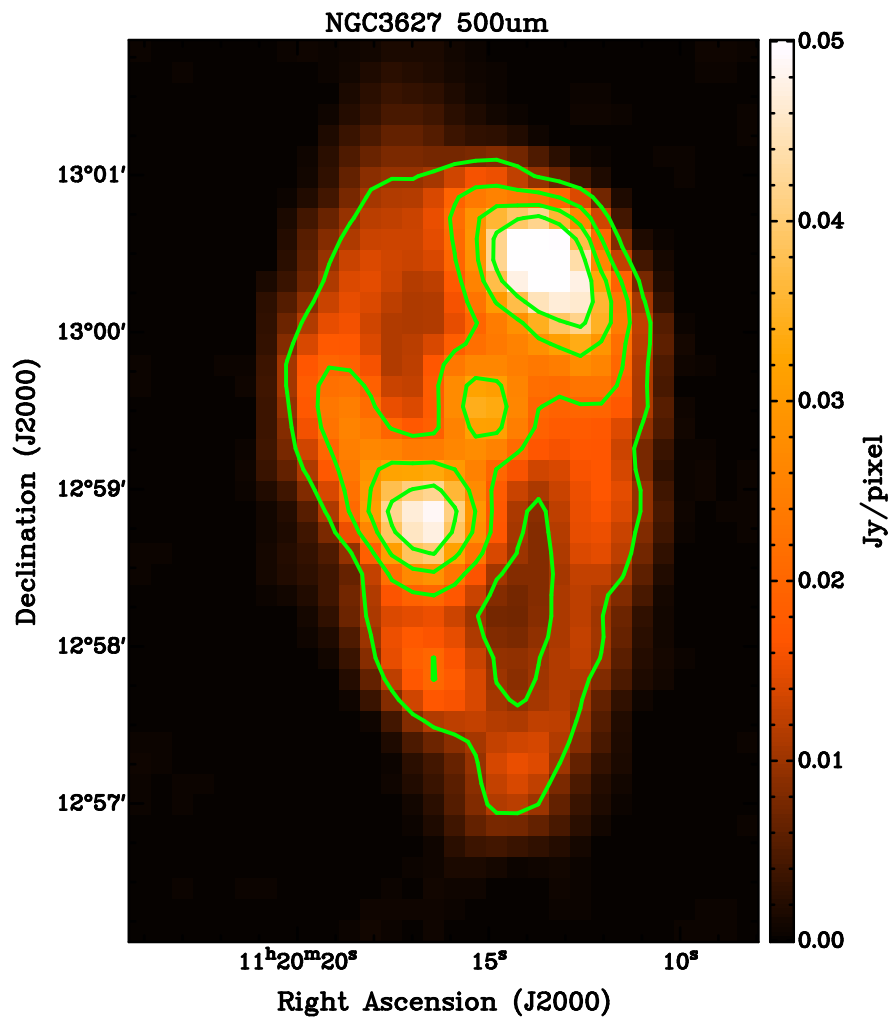


Figure 2.11: Image after the MAKEMAP filtering of the $500\mu\text{m}$ observations with 20%, 40%, 60%, and 80% contours.

Table 2.3. Properties of NGC3627 NGLS Observations

Observation	Beam Properties θ_{beam}	RMS $[mJy/Pixel]$	Percentage of Emission Removed
CO J=3-2	14.5''	1.28e-2	29.8%

removing it for an accurate SED analysis. The rms and resolution of the CO J=3-2 emission are shown in Table 2.3 for 8'' by 8'' pixels, and the scan prior to convolution is shown in Figure 2.12.

2.3.3 Nobeyama 45-m

Determining a dust-to-gas ratio requires a molecular tracer to estimate the amount of molecular hydrogen present. The most frequently used tracer is CO J=1-0 due to its abundance in the ISM. The CO J=1-0 data we used was taken from the Nobeyama 45-m CO Atlas of Nearby Spiral Galaxies observed to better understand the role of bars relating to molecular gas (Kuno et al., 2007). The Nobeyama 45-m CO Atlas consists of galaxies with morphologies ranging from Sa to Scd, located less than 25Mpc from the Milky Way, inclination values less than 79°, 100 μ m flux greater than 10Jy, and spiral structure that has not been compromised through interactions. Any galaxies that met these criteria were then observed with the Nobeyama 45-m telescope (Kuno et al., 2007). The beam sizes and rms of the filtered CO J=1-0 map are displayed in Table 2.4 for 8'' by 8'' pixels, and the final image product can be seen in Figure 2.13.

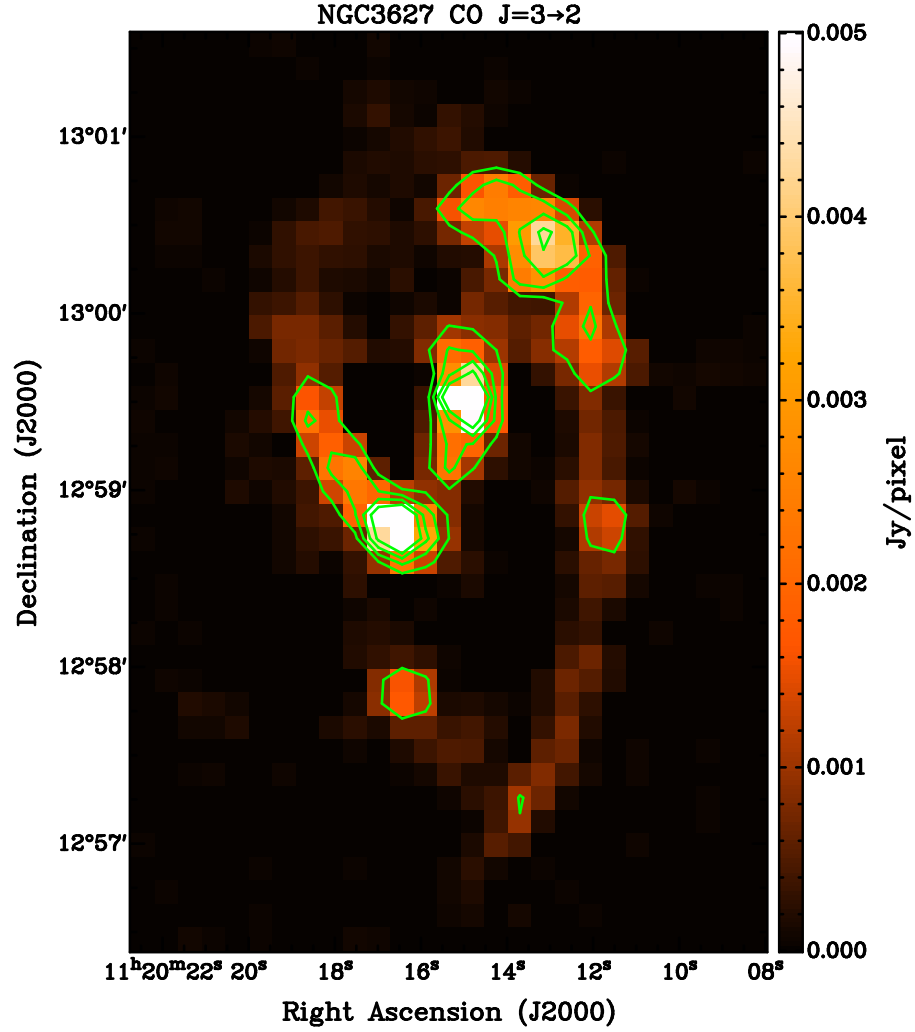


Figure 2.12: Residual of the MAKEMAP filtering of CO J=3-2 observations used to subtract line contamination from 850 μ m SCUBA-2 map with 20%, 40%, 60%, and 80% contours.

Table 2.4. Properties of NGC3627 Nobeyama 45-m Observations

Observation	Beam Properties θ_{beam}	RMS [K km/s]	Percentage of Emission Removed
CO J=1-0	15.0''	0.681	20%

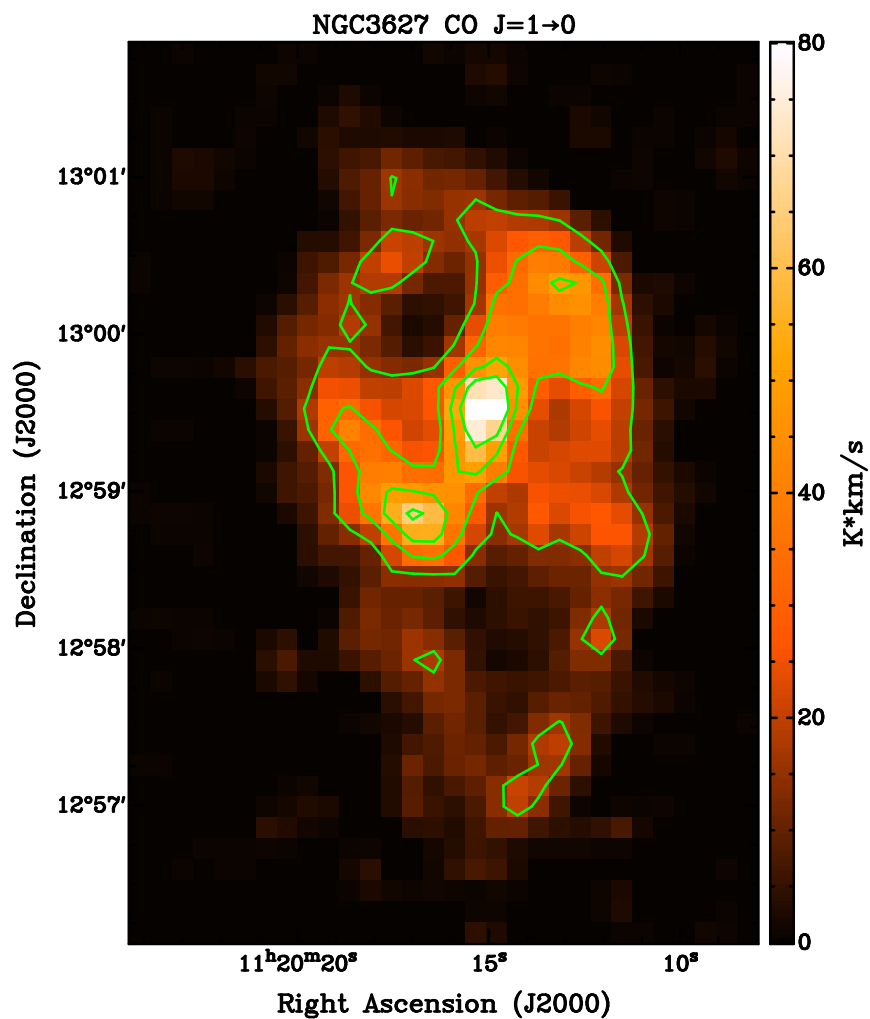


Figure 2.13: Residual of the MAKEMAP filtering of CO J=1-0 observations with 20%, 40%, 60%, and 80% contours.

Table 2.5. Properties of NGC3627 HERACLES Observations

Observation	Beam Properties θ_{beam}	RMS [K km/s]	Percentage of Emission Removed
CO J=2-1	13.0''	0.305	7%

2.3.4 Hetrodyne Reciever Array CO-Line Extragalactic Survey (HERACLES)

The CO J=2-1 line was used to determine a CO $2 - 1/1 - 0$ line ratio which can be used to trace a gradient in α_{CO} and hint towards regions of high star formation (Reuter et al., 1996). We used the CO J=1-0 data from the Nobeyama 45-m telescope (§ 2.3.3), and CO J=2-1 from the Hetrodyne Reciever Array CO-Line Extragalactic Survey (HERACLES) using the IRAM 30-m telescope. The main goal of HERACLES was to quantify the relationship between atomic and molecular gas and star formation using a large sample of galaxies (Leroy et al., 2009). The sample of galaxies chosen were targets contained in THINGS that were within the observing limits of the IRAM 30-m telescope. The final CO J=2-1 image can be seen in Figure 2.14 and the image properties can be seen in Table 2.5 with 8'' by 8'' pixels.

2.3.5 The HI Nearby Galaxy Survey (THINGS)

To determine the gas-to-dust ratio we had to determine the total amount of gas present which includes both atomic and molecular hydrogen. We approximated the amount of molecular hydrogen by using CO J=1-0, and measured

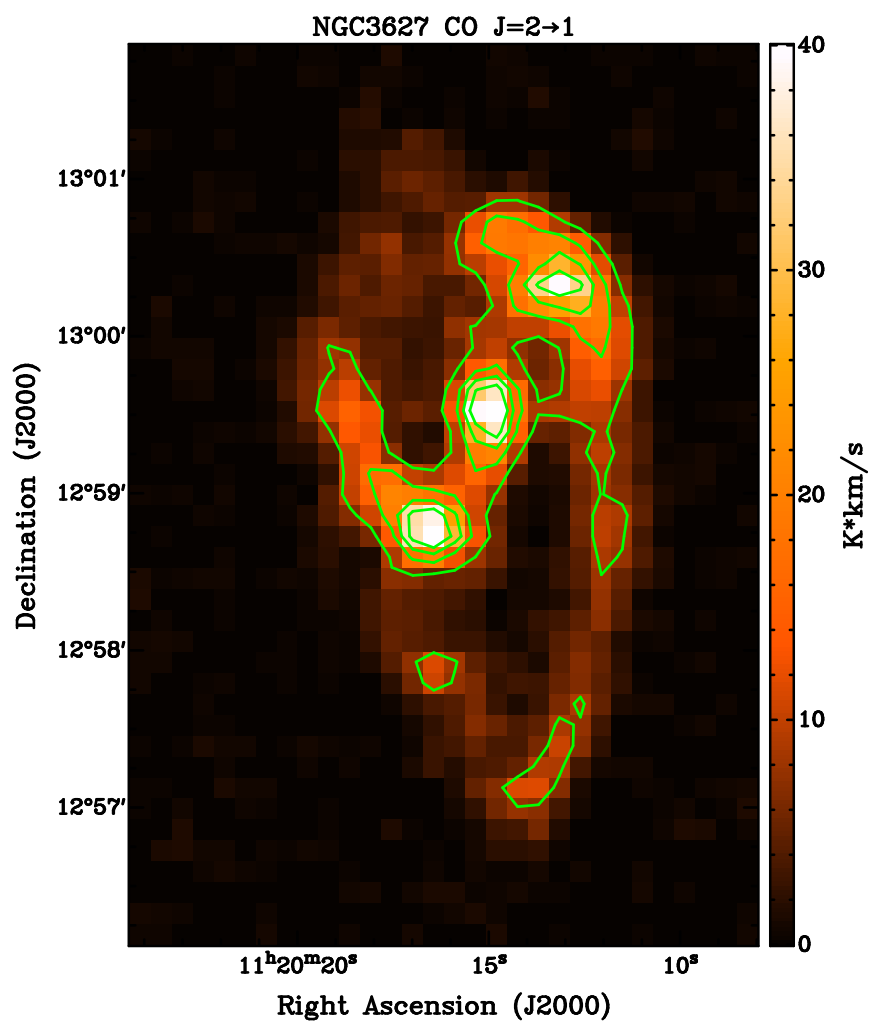


Figure 2.14: Residual of the MAKEMAP filtering of CO J=2-1 observations with 20%, 40%, 60%, and 80% contours.

Table 2.6. Properties of NGC3627 THINGS Observations

Observation	Beam Properties			RMS [M_{\odot}/pc^2]	Percentage of Surface Density Removed
	θ_{maj}	θ_{min}	θ_{PA}		
HI	10.6''	8.85''	-48.0°	0.760	>99%

the amount of atomic hydrogen (HI) present. Our atomic hydrogen map came from The HI Nearby Galaxy Survey (THINGS) designed to observe HI emission in nearby galaxies with the extreme spatial resolution of the Very Large Array (VLA). Targets in THINGS included many of the SINGS targets with the exception of known HI poor sources (E/S0 type galaxies), dynamically complex systems (edge-on spirals), and large extended galaxies found in the Local Group (Walter et al., 2008). The resolution and rms of the filtered image are shown in Table 2.6 for 8'' by 8'' pixels. The final data product is shown in Figure 2.15 and has been converted to $M_{\odot} pc^{-2}$ using

$$M_{HI} [M_{\odot}] = 2.36 \times 10^5 D^2 \times \sum_i S_i \Delta v \quad (2.2)$$

(Walter et al., 2008) and then dividing by the pixel area in pc^2 where D is the distance, and $\sum_i S_i \Delta v$ is the first moment of the flux. Since most of the HI emission is diffuse and extended, the filtering process resulted in a severe bowling lowering the surface densities, as well as removing a significant portion.

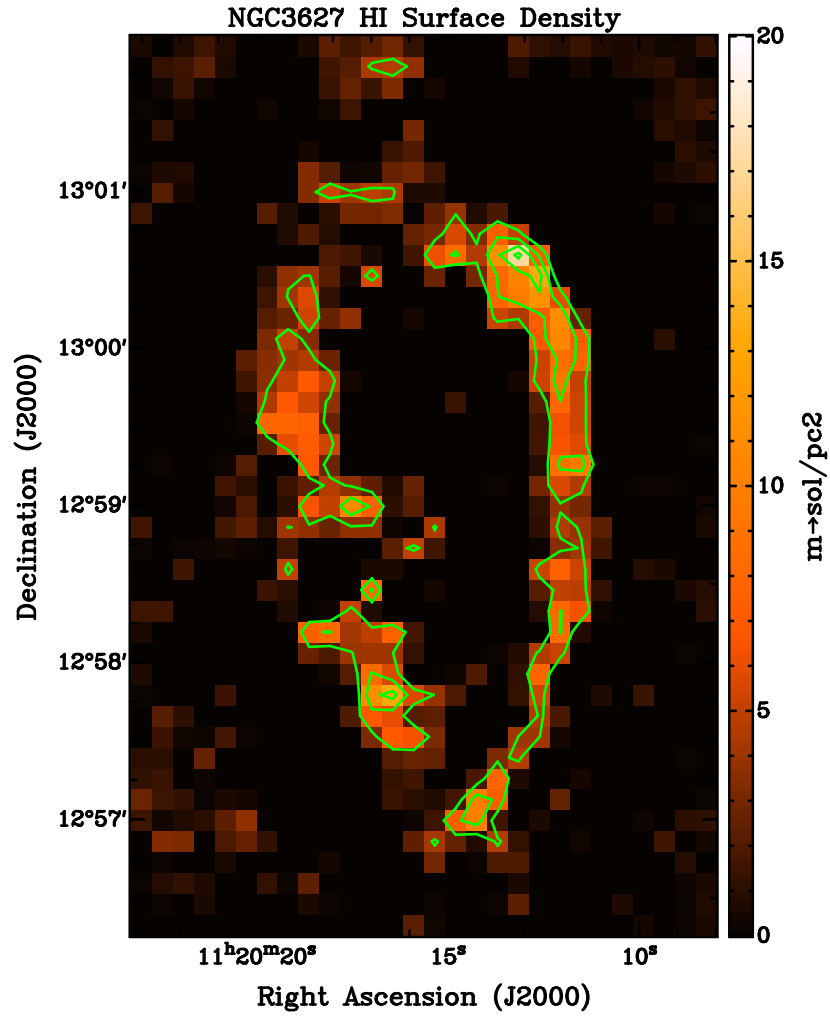


Figure 2.15: Residual of the MAKEMAP filtering of HI observations with 20%, 40%, 60%, and 80% contours.

2.4 Data Preparation for Analysis

Initially the data at the different wavelengths do not agree on several levels. The differences consist of the presence of the large scale/extended structure and the different resolutions including the beam shape of the $450\mu\text{m}$ data. We have taken several steps to correct for these disagreements and maximize the compatibility of the data. In order to account for the varying beam resolutions, we use a gaussian convolution to increase the resolution of our maps to the largest beam size in our dataset, $\sigma_{max}=36''$. An appropriate convolution beam width is determined by

$$\sigma_{desire} = \sqrt{\sigma_{max}^2 - \sigma_{given}^2} \quad (2.3)$$

where σ_{desire} is the desired beam width, σ_{max} is the beam size we are convolving to, and σ_{given} is the beam size we are convolving from. However, equation ?? only works when the beams are well approximated by a single gaussian which is not the case for the $450\mu\text{m}$ beam. The steps taken to match the resolution of the $450\mu\text{m}$ beam with the rest of the data set are given in §2.4.1.

Removing any large scale structure from our ancillary data is implemented using a feature built into MAKEMAP that allows us to add fake sources into the data during production. The fake source implementation allows us to remove the same amount of large scale structure from our ancillary data as was removed from the SCUBA-2 data. The steps taken to prepare the ancillary data are described in §2.4.2.

2.4.1 Accounting for the $450\mu\text{m}$ Error Beam

Taking the $450\mu\text{m}$ error beam into consideration is different than a normal beam convolution in the sense that we are not convolving the higher resolution map to the lowest resolution. Instead we are adding in an error beam similar to the error beam found in the $450\mu\text{m}$ observations. We have to take these steps because convolving a double gaussian beam with a single gaussian kernel will not sufficiently remove the error component of our beam resulting in a poor approximation to the wings of our beam shape.

In order to accommodate the $450\mu\text{m}$ map's error beam, we used a method employed by another SCUBA-2 legacy survey, the Gould Belt Survey. This method uses the distributive nature of the Fourier transform to create similar error components in the beams we were convolving to and from. Carrying out this convolution involves breaking the $450\mu\text{m}$ beam into its two components, X_α and X_β , to represent the main beam and error beam. The values of the main beam amplitude, X_α , and error beam amplitude, X_β , will sum to one so the height of the total beam is normalized to one. The poorest resolution beam, X_{max} , is convolved with the two components X_α and X_β and the resulting beams are added together. This process produces an equivalent two component beam to the $450\mu\text{m}$ beam convolved with X_{max} . The relationship can be expressed as

$$\begin{aligned} X_{max} * X_\alpha + X_{max} * X_\beta &= (X_\alpha + X_\beta) * X_{max} \\ &= X_{450\mu\text{m}} * X_{max} \end{aligned} \tag{2.4}$$

where X_{max} is the poorest resolution beam, X_α and X_β are the main and error beam of the $450\mu\text{m}$ observations, and $X_{450\mu\text{m}}$ is the double gaussian beam shape of the $450\mu\text{m}$ observations.

2.4.2 Extended Structure Removal via MAKEMAP

Due to the combination of methods used in MAKEMAP, large scale/extended structure is removed from the final SCUBA-2 images. However, in all of our ancillary data the large scale emission was present in the initial maps. The removal of the extended features from our ancillary data was carried out by passing the data through MAKEMAP using a special function called fakemap. Using fakemap allows us to pass an image through the SCUBA-2 processing and have it added into the image being processed. We use the $850\mu\text{m}$ map as our base image for the filtering process and add the ancillary data to the $850\mu\text{m}$ image. The output image then consists of the sum of the ancillary data and the $850\mu\text{m}$ map. Finally, the ancillary data were isolated by subtracting the $850\mu\text{m}$ image from the fakemap output image.

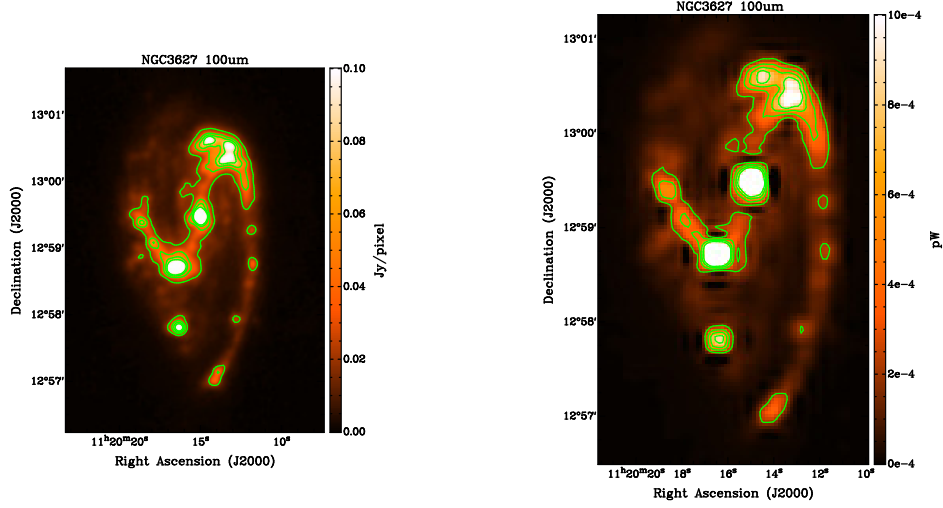
Preparing the data to be added via fakemap consists of either converting the images from their native units into pW using the $850\mu\text{m}$ flux calibration factor and scaling down to match the observed signal or by just applying a scaling factor. Which method is used is based on the desired units of the final map, and can be separated by what purpose the data had in our analysis.

The data used in the SED fitting (KINGFISH and NGLS) were scaled to pW so they could have a similar reduction and calibration process as the SCUBA-2 maps. The KINGFISH data are regridded to a $2''$ by $2''$ pixel size

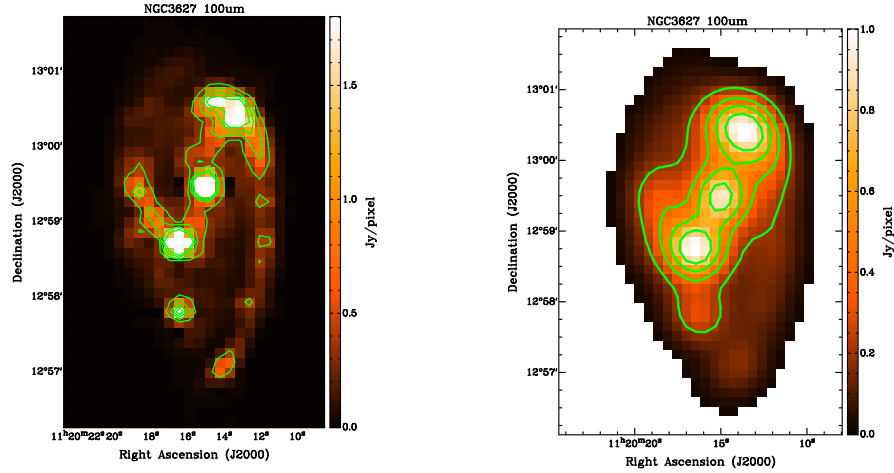
and have the appropriate calibration factor from Dempsey et al. (2013) applied to convert from either MJy/sr to pW in the case of the $250\mu\text{m}$, $350\mu\text{m}$, and $500\mu\text{m}$ or from Jy/pixel to pW for the $100\mu\text{m}$ and $160\mu\text{m}$. Converting the CO J=3-2 data requires the final product to be in the same units as the $850\mu\text{m}$ observations in order to properly remove the molecular gas contribution. Converting from K km/s to mJy/beam involves applying a scaling constant of $0.70 [\text{mJy/beam}][\text{K km/s}]^{-1}$ (Drabek et al., 2012) prior to applying the $850\mu\text{m}$ flux calibration factor to convert to pW.

After the data has been converted to pW, the image used in fakemap is then scaled down to the same order of magnitude as the base image using a scaling factor specified prior to the map production. After the fakemap image has been processed and the $850\mu\text{m}$ map subtracted, the maps are scaled back up using the same scaling factor used to scale them down and calibrated using the same flux calibration factors as the $850\mu\text{m}$ map and scaled to an $8''$ by $8''$ grid. The amount of extended flux lost from the KINGFISH and NGLS images is shown in Tables 2.2 and 2.3. This process is illustrated in Figure 2.16 from the original Herschel $100\mu\text{m}$ map, to the final image used in SED fitting convolved to the $500\mu\text{m}$ beam resolution, $36.0''$.

The rest of the ancillary data are used in calculating a dust to gas ratio, and follow nearly the same process as the SED data filtering. The major difference is the CO J=1-0, CO J=2-1 and HI maps need to remain in their original units of K km/s and $\text{M}_{\odot}/\text{pc}^2$. This requirement simplified the process by only requiring a scaling factor of 0.001 to be applied to the original maps. After the image is scaled, it is filtered using the fakesource option in MAKEMAP



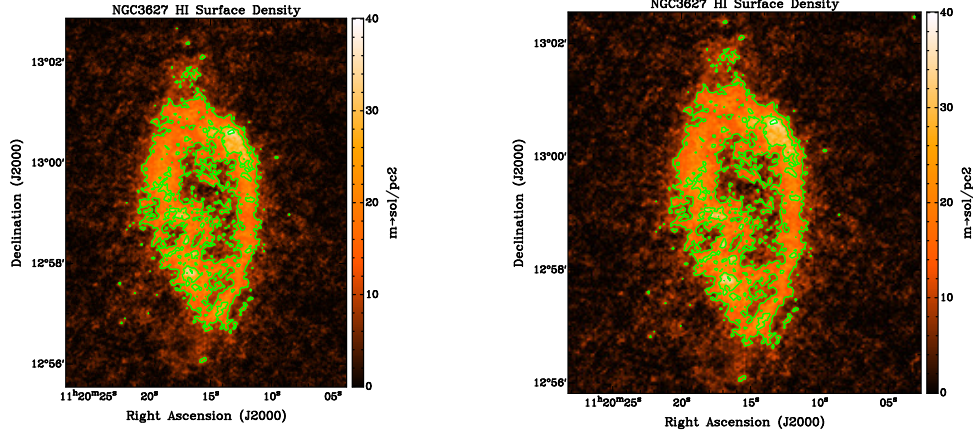
(a) Herschel 100μm image of NGC3627 with 1.7'' by 1.7'' pixels. (b) Herschel 100μm image converted to pW and rescaled to a 2'' by 2'' pixel grid.



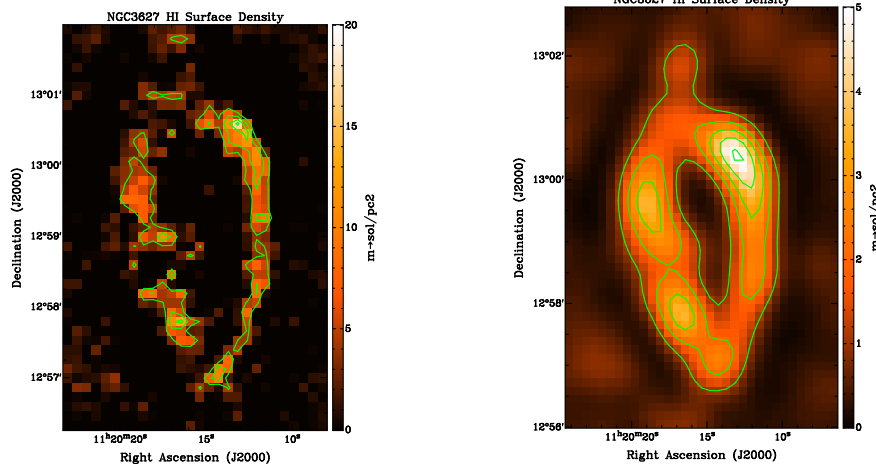
(c) 100μm map of NGC3627 after large scale structure has been removed and rescaled to an 8'' by 8'' grid. (d) 100μm map with extended structure removed and convolved to final resolution of 36.0'' with a 5σ cut applied.

Figure 2.16: The 100μm image from the beginning of processing to the end of processing. All of the contours shown are 20%, 40%, 60%, and 80%.

with the $850\mu\text{m}$ observation as the base image. The atomic and molecular gas maps were then isolated in the same fashion as the KINGFISH and NGLS by subtracting the $850\mu\text{m}$ map. Then they were rescaled back to their original values, fit to an $8''$ by $8''$ grid and finally convolved to a $36.0''$ resolution. The process is illustrated in Figure 2.17. The amount of emission lost is shown in tables 2.4, 2.5, and 2.6.



(a) HI surface density map of NGC3627 with $1.5''$ by $1.5''$ pixels. (b) HI surface density map rescaled to a $2''$ by $2''$ pixel grid.



(c) HI surface density map of NGC3627 after large scale structure has been removed and rescaled to an $8''$ by $8''$ grid. (d) HI surface density map with extended structure removed and convolved to final resolution of $36.0''$.

Figure 2.17: The HI surface density maps from the beginning of processing to the end of processing. The contours in the top row are 30%, 60% and 90%, and the contours in the bottom row are 20%, 40%, 60%, and 80%.

Chapter 3

Spectral Energy Distribution Analysis

3.1 SED Fitting Method

In order to determine a dust mass, we use the IDL package MPFIT (Markwardt, 2009) to fit equation

$$S_{\nu}(T) = \frac{M \kappa_{\nu,0}}{D^2} \left(\frac{\nu}{\nu_0} \right)^{\beta+3} B_{\nu}(T) \quad (3.1)$$

for the temperature, T , mass, M , and the emissivity index, β while the opacity, $\kappa_{\nu,0}$, distance, D , and reference wavelength, ν_0 are held fixed. The routine MPFIT utilizes the Levenberg-Marquardt algorithm. This algorithm uses a combination of two minimization techniques (the steepest descent method and the Newton-Raphson Method) to determine the parameter combination that corresponds to a minimum in the χ^2 space while maximizing the efficiency of the step sizes in each iteration (Burden & Faires, 2001). The algorithm begins by implementing the steepest descent method. The way this technique works is it will follow the direction opposite to the largest gradient in order to traverse the χ^2 space to locate a minimum. As the set of solutions approaches

Table 3.1. Systematic Calibration Uncertainties for SCUBA-2 and KINGFISH Observations

Observation	Scaling Factor
100 μm	3%
160 μm	5%
250 μm	7%
350 μm	7%
450 μm	12%
500 μm	7%
850 μm	8%

a minimum, it will switch to the Newton-Raphson method to locate the best set of parameters by finding where the derivative at that point is closest to zero (Gavin, 2013).

In order for MPFIT to provide the most accurate fit, we establish a reasonable uncertainty for each of our data points, and determine a realistic set of starting points for the fitting. The variance for our SCUBA-2 data is determined using

$$\sigma^2 = \sigma_{obs}^2 + \sigma_{rms,sky}^2 + \sigma_{calib}^2 \quad (3.2)$$

such that σ_{obs} is the noise determined by MAKEMAP, $\sigma_{rms,sky}$ is the RMS of the sky, and σ_{calib} is the product of calibration uncertainty and observed flux for each observed wavelength. The systemic calibration scaling factors are shown in Table 3.1.

The variances for the KINGFISH data are determined in a similar fashion using

$$\sigma^2 = \sigma_{rms,sky}^2 + \sigma_{calib}^2. \quad (3.3)$$

The observation error is excluded since the reported variance in the filtered images reflects the fake source image used and not the KINGFISH data set.

The nature of the Levinberg-Marquardt method leaves the solution vulnerable to converge at a local minimum rather than converging at the global minimum. This is remedied by selecting reasonable initial conditions. The initial conditions we used were a modified blackbody with a temperature of 20 K, a dust emissivity index of 2, and a mass determined by

$$\begin{aligned} M &= \frac{D^2 S_{250}}{\kappa_{\nu,0} B_{250}(T)} \left(\frac{\nu}{\nu_0} \right)^{-\beta} \\ &= 5.40 \times 10^5 S_{250} [M_{\odot}] \end{aligned} \quad (3.4)$$

using the flux from the $250\mu\text{m}$ emission and our initial temperature and dust emissivity index values with a reference opacity of $0.2665 \text{ m}^2 \text{ kg}^{-1}$ at $300\mu\text{m}$. When the best fit values have been calculated we use equation 3.4 to determine the dust mass associated with the $250\mu\text{m}$ emission instead of using the fitted peak mass in order to increase the signal to noise of our data and avoid any temperature dependences associated with emission further from the Rayleigh-Jeans tail of the SED. The error in the dust mass is then determined by taking the total derivative of the mass function with respect to each of the free variables.

3.2 Fitting the Spectral Energy Distribution

The fitting procedure was carried out in two different ways on a modified blackbody equation. One of the two methods is fitting an SED to each individual pixel in order to generate a set of parameter maps. The second method sums the flux of each of the selected regions shown in Figure 3.1 to maximize the signal to noise ratio in order to generate a more precise set of parameters. For both of the fitting methods the mass, M , and temperature, T , are set as free parameters, but the emissivity index, β , requires special treatment for each of the methods. The distance, D , has been set to 9.4 Mpc (Walter et al., 2008), and the reference opacity, $\kappa_{\nu,0}$, was tested using $0.2665 \text{ m}^2 \text{ kg}^{-1}$ (Li & Draine, 2001) and $1.0 \text{ m}^2 \text{ kg}^{-1}$ (Planck Collaboration et al., 2011).

3.2.1 Pixel SED Fits

In order to generate a parameter map of NGC3627, each individual pixel has its own SED determined from the available wavelengths described in the observations chapter. The fits are performed over $100\mu\text{m}$, $160\mu\text{m}$, $250\mu\text{m}$, $350\mu\text{m}$, $450\mu\text{m}$ or $500\mu\text{m}$, and $850\mu\text{m}$ observations. Initially, we excluded the $500\mu\text{m}$ emission in order to increase the resolution of our final maps, but an over abundance of emission in the $450\mu\text{m}$ map provided unreliable results, so we exchanged the $450\mu\text{m}$ emission with $500\mu\text{m}$ emission. The treatment of the dust emissivity index is performed by fixing it at $\beta = 1.8$ for the Planck opacity model where the opacity is $\kappa_{300\mu\text{m},0} = 1.0 \text{ m}^2 \text{ kg}^{-1}$ (Planck Collaboration et al., 2011) and $\beta = 2.0$ for the Li and Draine opacity model where the opacity is $\kappa_{300\mu\text{m},0} = 0.2665 \text{ m}^2 \text{ kg}^{-1}$ (Li & Draine, 2001). A third fit is performed

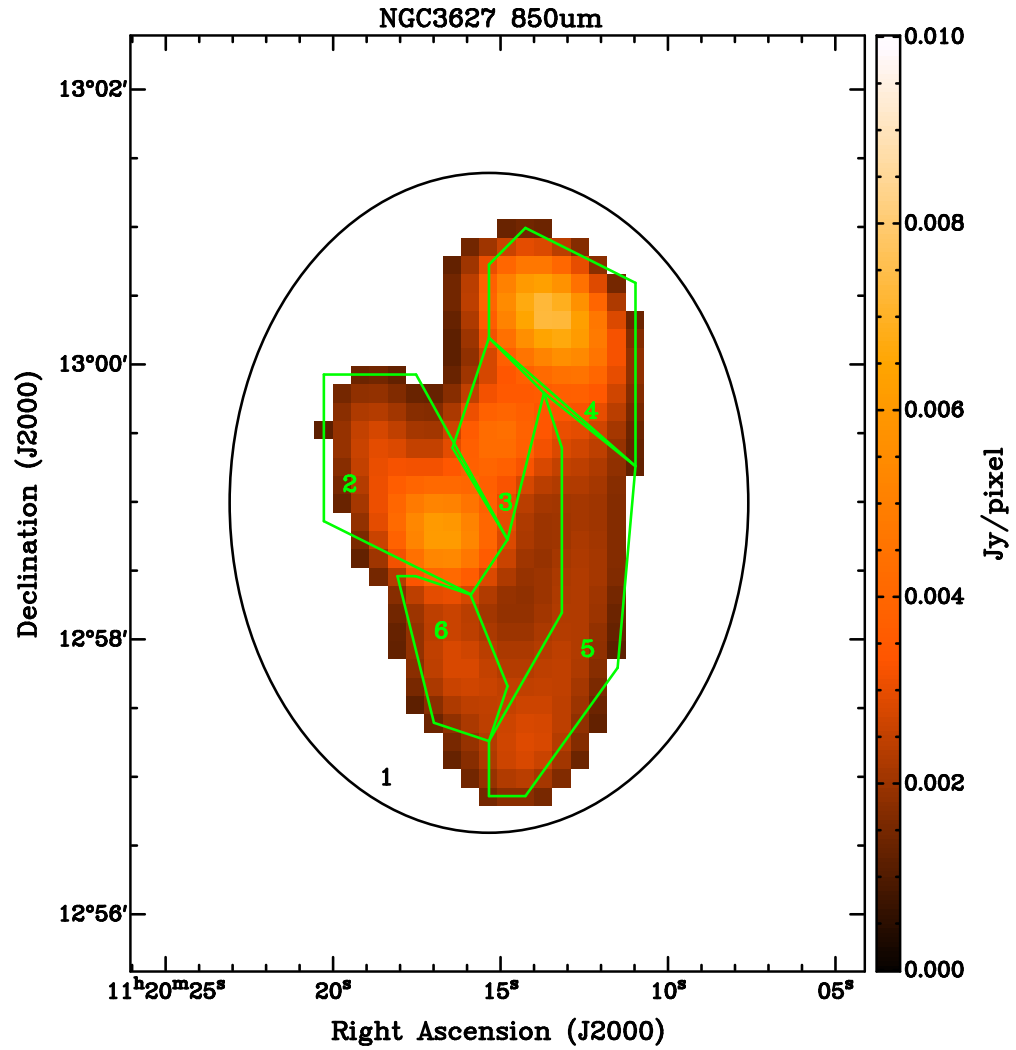


Figure 3.1: 850 μ m emission convolved to the 500 μ m beam size overlaid with the selected regions of NGC3627 labeled 1 through 6 such that region 1 includes the entire galaxy.

where the emissivity index is allowed to vary as a free parameter. While the emissivity index is allowed to vary, the opacity being used is the Planck model. Using either an opacity of $1.0 \text{ m}^2 \text{ kg}^{-1}$ or $0.2665 \text{ m}^2 \text{ kg}^{-1}$ will not affect the shape of the SED, only its normalizing value. In the case of our fits, it is the mass that acts as the normalizing value, so increasing or decreasing the opacity will yield the opposite effect in the mass. The inverse proportionality can be seen in equation 3.4.

The quality of the fit using the $450\mu\text{m}$ observations for the Li and Draine model of $\beta=2$ and $\kappa_{\nu,300}=0.2665$ is shown in Figure 3.2. The quality of the fit is determined by how well the expected flux from the SED fitting matches the observed flux. If the fits are able to recreate the observed emission, then all of the points will lie on the line $y = x$ shown in the plots as a black line. If the SED is underestimating the flux, the points will appear below the 1 to 1 line, and an over estimation from the SED will result in points above the line.

From Figure 3.2, the fitting procedure is able to handle the $100\mu\text{m}$, $250\mu\text{m}$, and $850\mu\text{m}$ with some slight under estimation in the $160\mu\text{m}$ and over estimation in the $350\mu\text{m}$. The most compelling piece of information in Figure 3.2 is the SED's inability to fit the observed $450\mu\text{m}$ flux despite the care taken to account for the odd beam shape of the $450\mu\text{m}$ data. The discrepancy between the observed and fitted fluxes is not a result of a poor fit, but rather an intrinsic error in the $450\mu\text{m}$ observations. The over abundance of flux present in the SCUBA-2 $450\mu\text{m}$ data was determined to be an issue with the data itself rather than a signature of a physical process. We determined that the $450\mu\text{m}$ data was overestimating the flux by calculating a $450\mu\text{m}$ flux using a linear fit

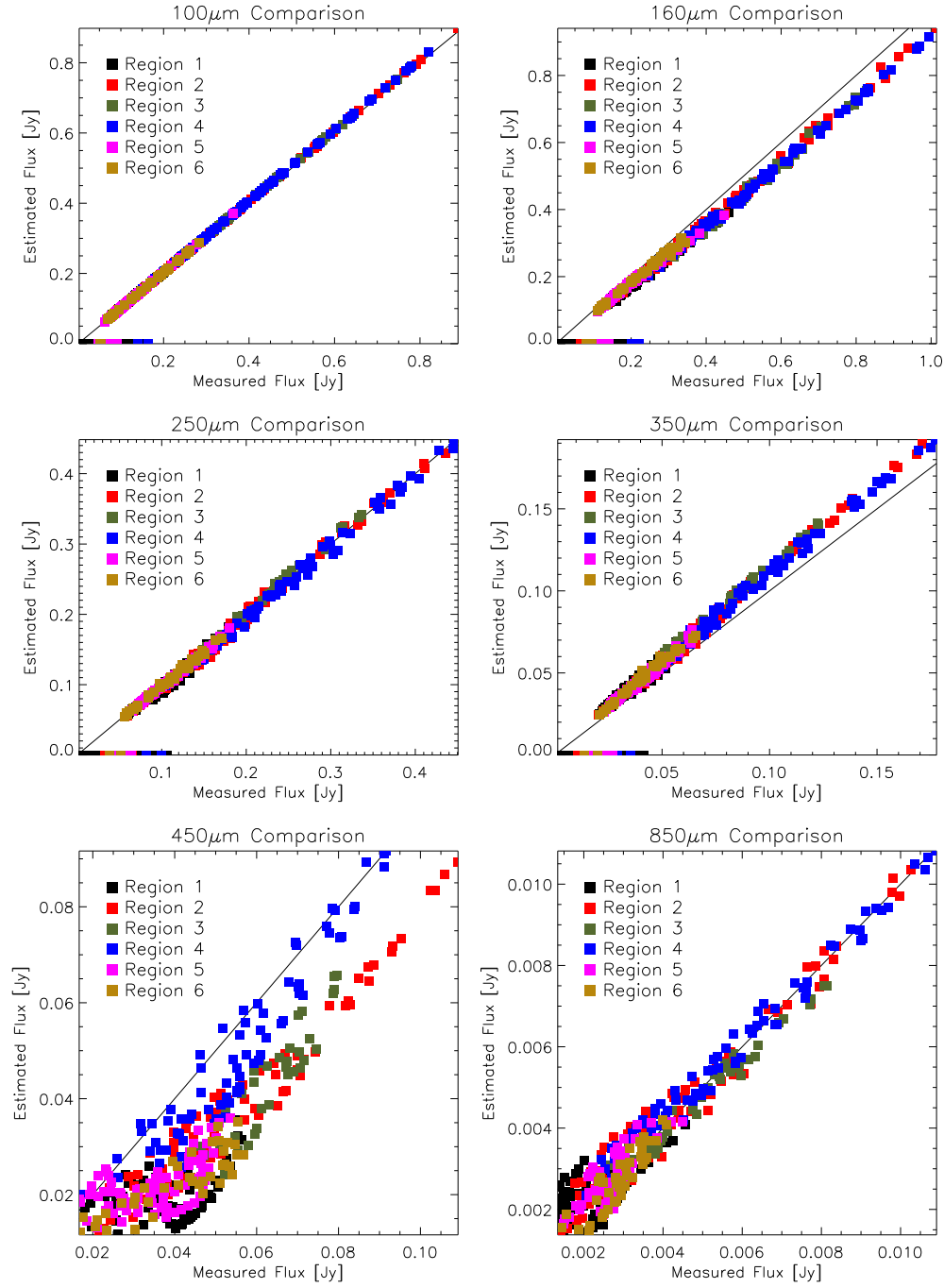


Figure 3.2: Quality of the SED fits to the Li and Draine model using the 450 μ m emission. The regions shown are the regions in Figure 3.1.

Table 3.2. Total Distance to 1 to 1 Line Using the $500\mu\text{m}$ Emission

Observation	Planck Model	Li and Draine Model	Variable Emissivity Index
$100\mu\text{m}$	0.000	0.01245	0.1277
$160\mu\text{m}$	15.55	8.979	2.159
$250\mu\text{m}$	5.808	3.071	0.4465
$350\mu\text{m}$	0.8045	0.3541	0.1161
$500\mu\text{m}$	0.08192	0.1147	0.1866
$850\mu\text{m}$	0.02825	0.05218	0.09528
Total	22.23	12.583	3.131

between the filtered $350\mu\text{m}$ flux of 27.1 Jy and the filtered $500\mu\text{m}$ flux of 8.56 Jy. The $450\mu\text{m}$ flux expected from these two data points was 14.7 Jy; however our $450\mu\text{m}$ flux was measured to be 20.6 Jy.

In order to avoid any errors in our final parameter maps, we substitute the $450\mu\text{m}$ emission with KINGFISH $500\mu\text{m}$ emission. The quality of the fitted SEDs using the $500\mu\text{m}$ emission are shown in Figures 3.3, 3.4, and 3.5 for the Planck model, Li and Draine model and the emissivity as a free parameter, respectively. To assign a numerical quantity to the quality of the fit the vertical distance was calculated for each point to the 1 to 1 line and then summed. The summed distances are shown in Table 3.2 and suggest the best fit comes from allowing the emissivity index to vary. The resulting parameter maps of the fits using the $500\mu\text{m}$ observations are shown in Figure 3.6. The numerical values from fitting each model are shown in Tables 3.3, 3.4, and 3.5 where each region corresponds to the regions labeled in Figure 3.1.

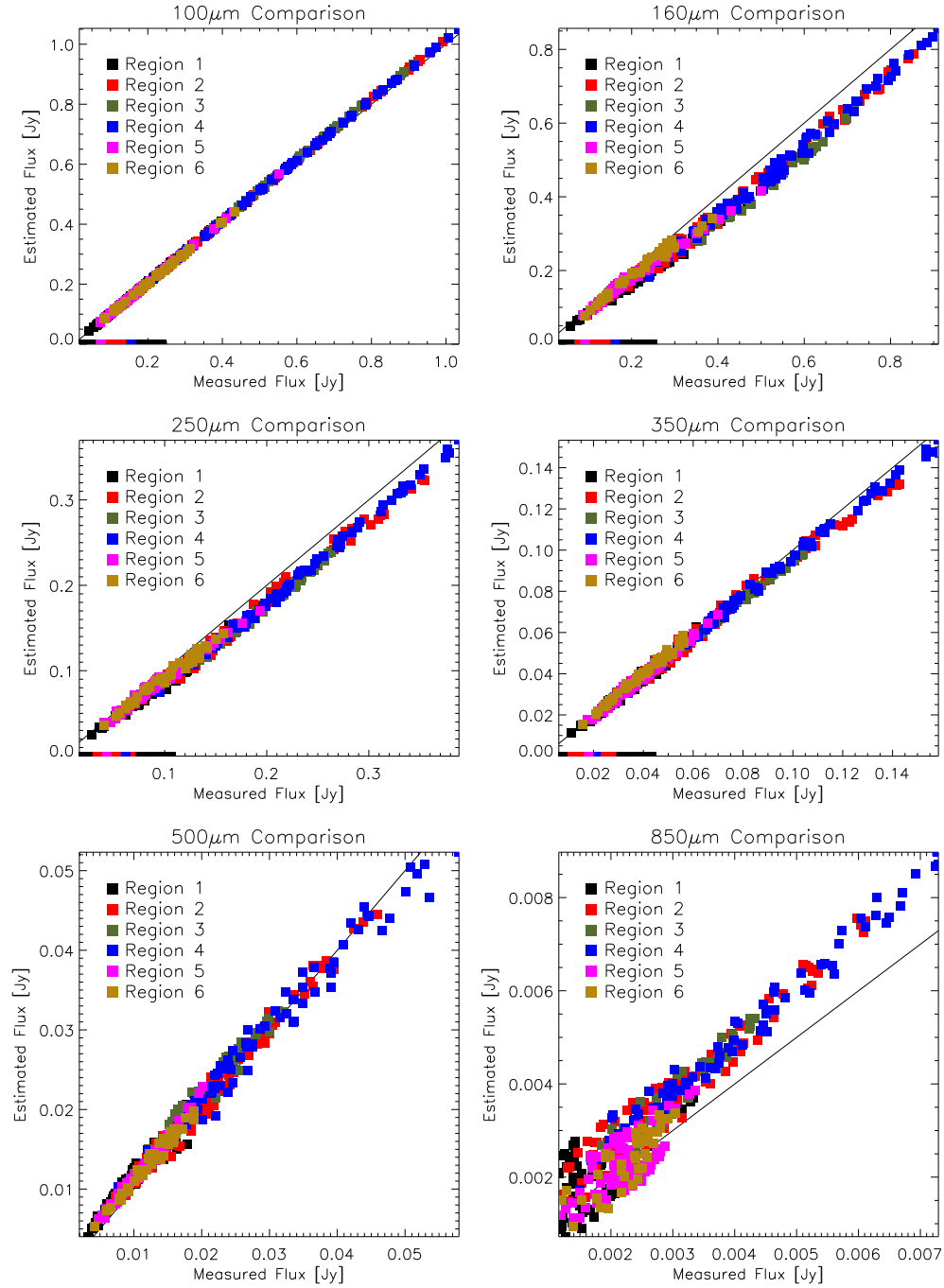


Figure 3.3: Quality of the SED fits using the Planck model with the 500 μ m emission. The regions shown are the regions in Figure 3.1.

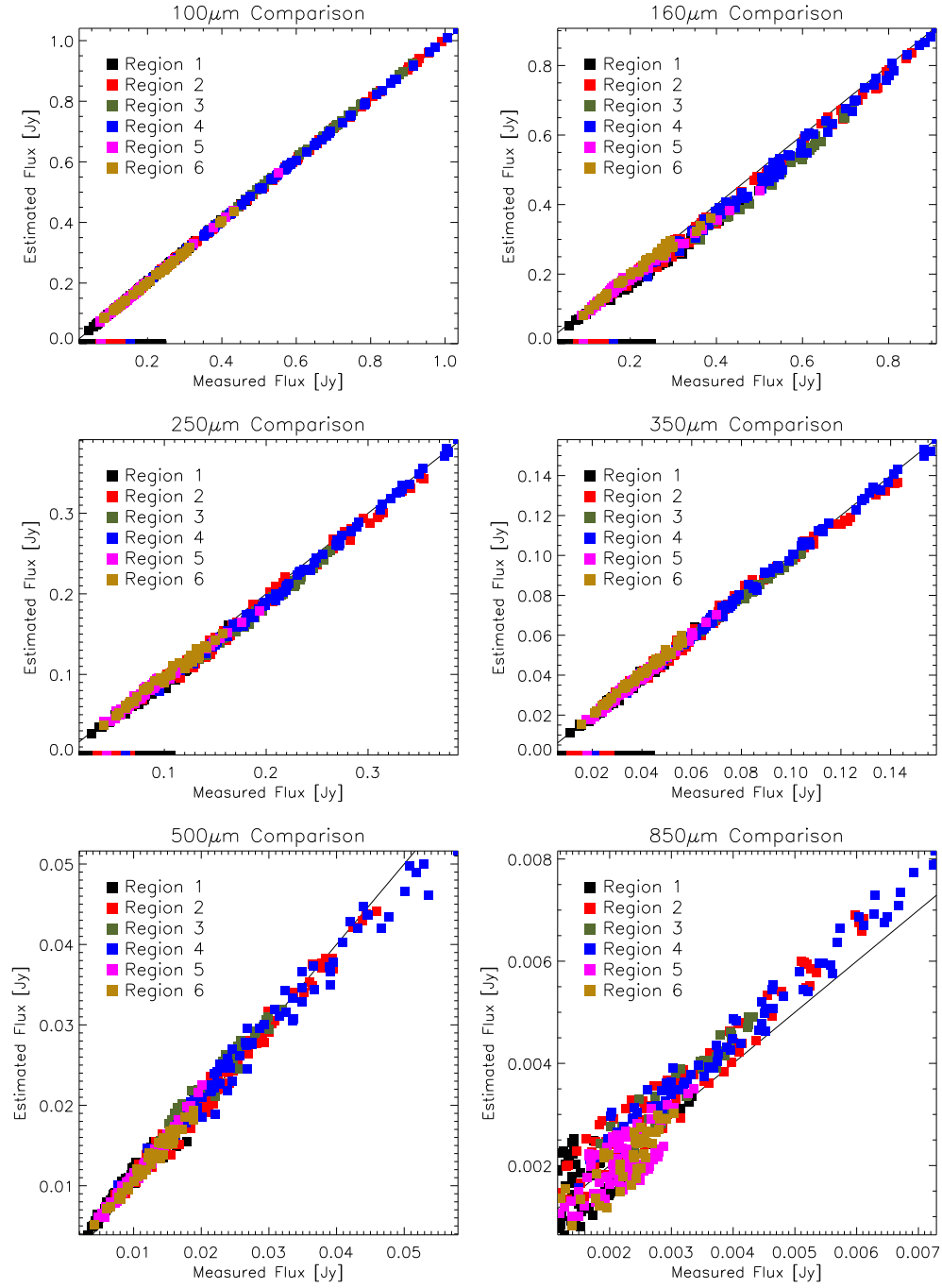


Figure 3.4: Quality of the SED fits using the Li and Draine model with the 500 μ m emission. The regions shown are the regions in Figure 3.1.

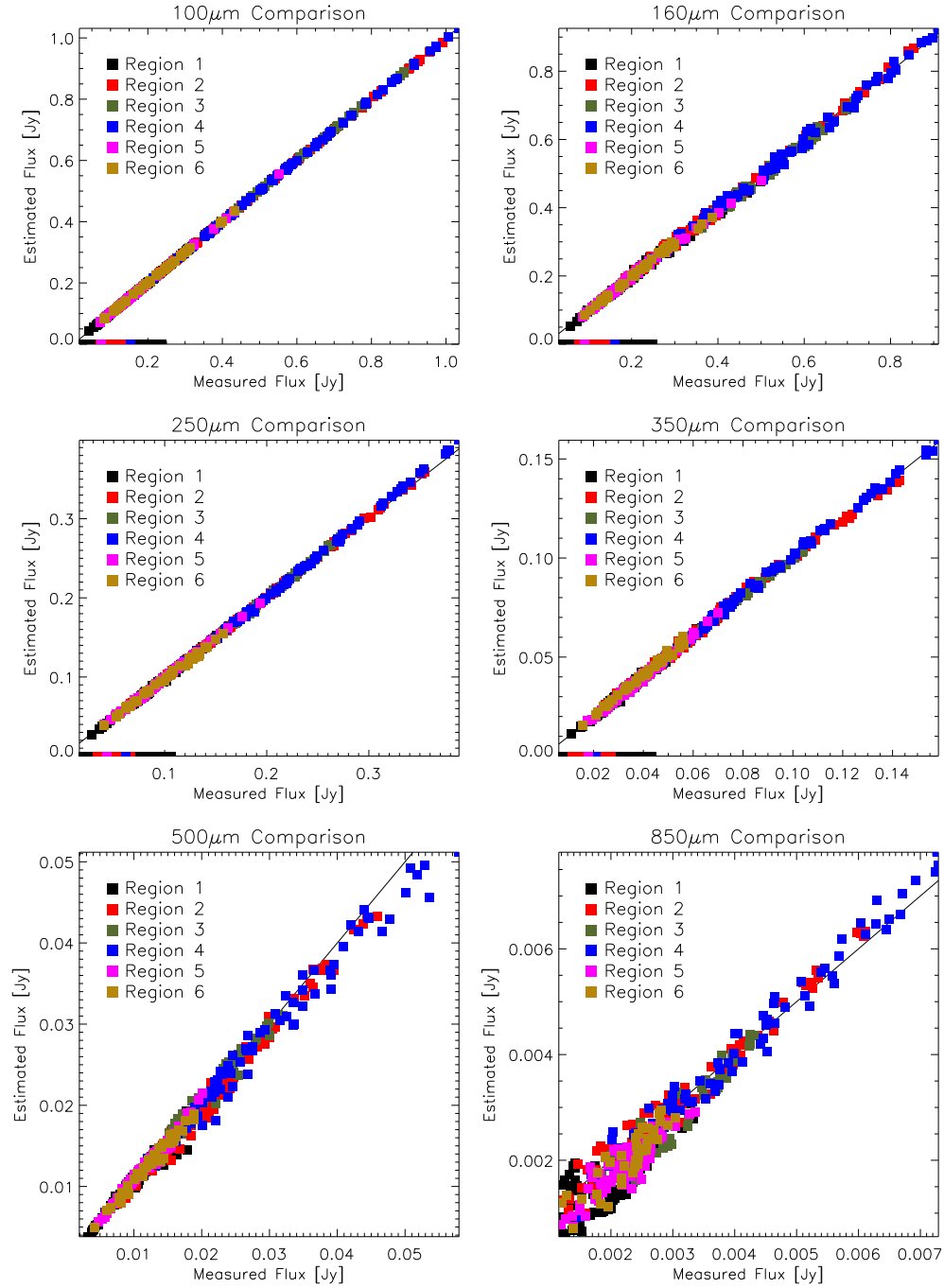


Figure 3.5: Quality of the SED fits with the emissivity index as a free parameter with the 500 μ m emission. The regions shown are the regions in Figure 3.1.

Table 3.3. Best Fit Parameters for Planck Model Using $500\mu\text{m}$ Emission for Pixel Fits

Region	Average β	Total Mass [$10^5 M_\odot$]	Average Surface Density [$M_\odot pc^{-2}$]	Average Temperature [K]
1	1.8	52 ± 23	0.10 ± 0.05	26 ± 2
2	1.8	13 ± 4	0.12 ± 0.04	27 ± 1
3	1.8	6 ± 1	0.12 ± 0.02	28.8 ± 0.6
4	1.8	16 ± 6	0.15 ± 0.05	26 ± 1
5	1.8	10 ± 2	0.08 ± 0.02	24 ± 1
6	1.8	4 ± 1	0.08 ± 0.02	25 ± 1

Table 3.4. Best Fit Parameters for Li and Draine Model Using $500\mu\text{m}$ Emission for Pixel Fits

Region	Average β	Total Mass [$10^5 M_\odot$]	Average Surface Density [$M_\odot pc^{-2}$]	Average Temperature [K]
1	2.0	230 ± 103	0.5 ± 0.2	24 ± 2
2	2.0	57 ± 19	0.6 ± 0.2	25 ± 1
3	2.0	28 ± 5	0.52 ± 0.09	26.4 ± 0.5
4	2.0	70 ± 22	0.7 ± 0.3	24.4 ± 0.8
5	2.0	42 ± 9	0.34 ± 0.08	22.4 ± 0.8
6	2.0	18 ± 4	0.33 ± 0.08	23.5 ± 0.9

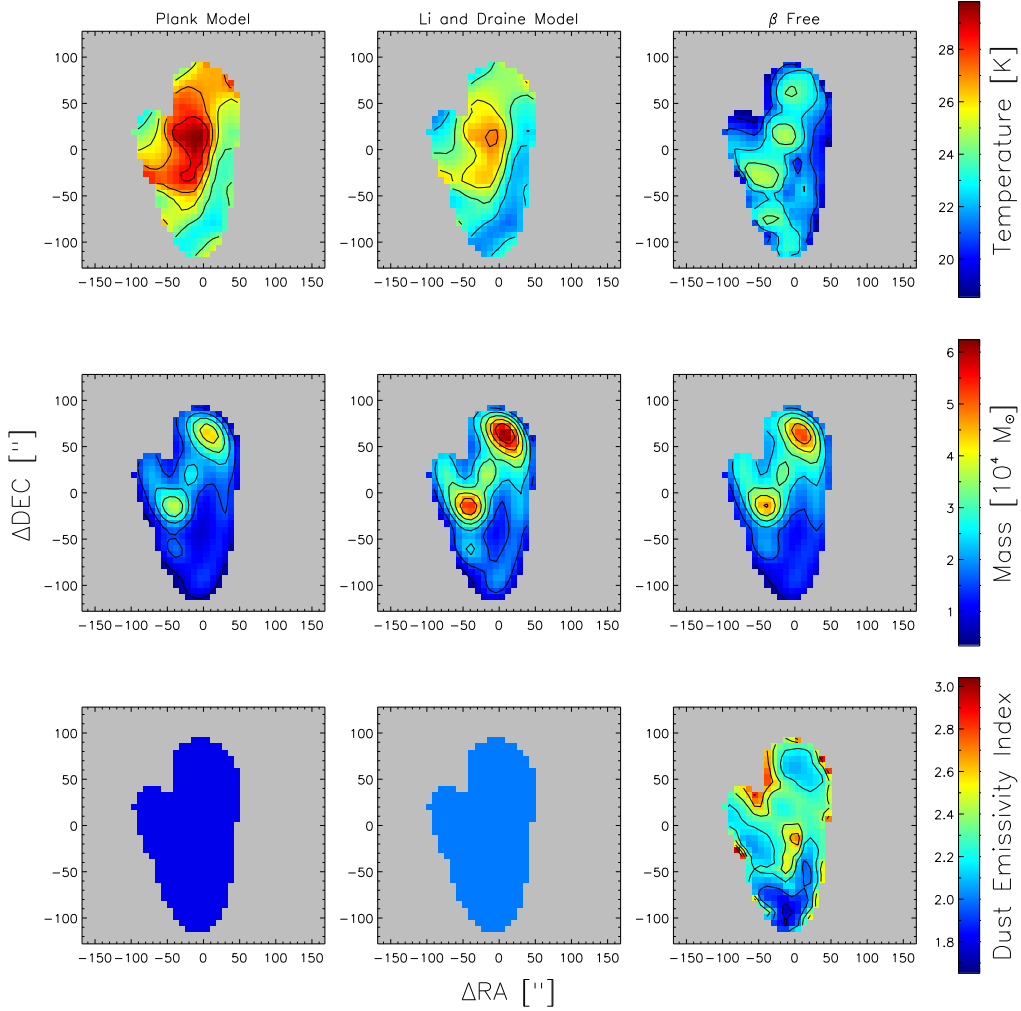


Figure 3.6: Returned value for the SED fits using the 500 μ m observations with the Planck model in the left column, the Li and Draine model in the middle column, and β as a free variable in the right column. The top row shows the temperature with contours from 19.5K to 28.5K in 1.5K increments. The second row show the returned masses with contours from 1.9 M_{\odot} to 13.3 M_{\odot} in 1.9 M_{\odot} increments. The Li and Draine mass fits have been divided by three to better show the features relative to the other two fits. The bottom row shows the returned dust emissivity index values with contours from 1.8 to 2.8 with 0.2 increments.

Table 3.5. Best Fit Parameters for β As A Free Parameter Using $500\mu\text{m}$ Emission for Pixel Fits

Region	Average β	Total Mass [$10^5 M_\odot$]	Average Surface Density [$10^5 M_\odot pc^{-2}$]	Average Temperature [K]
1	2.2 ± 0.2	73 ± 31	0.14 ± 0.06	22 ± 1
2	2.3 ± 0.2	18 ± 5	0.18 ± 0.05	22 ± 2
3	2.34 ± 0.09	9 ± 1	0.18 ± 0.02	23.0 ± 0.9
4	2.2 ± 0.1	22 ± 5	0.21 ± 0.05	22 ± 1
5	2.1 ± 0.2	12 ± 4	0.10 ± 0.03	22 ± 1
6	2.1 ± 0.1	5 ± 1	0.09 ± 0.03	23.0 ± 0.9

3.2.2 Total Region Flux SED Fits

The second method used to determine the dust mass was to fit the SED to the flux of each region in Figure 3.1. Performing the fit in this manner is beneficial because it increases the signal to noise of the region and produces a more precise set of parameters. Initially, fitting the total flux of each region was carried out in the same manner as the individual pixel fits by fixing the emissivity index to 1.8 and 2.0, and allowing the emissivity index to vary. This method ran into issues when determining an initial mass to use in the fitting. If we used the same procedure as we did for the pixel fitting, equation 3.4, our initial masses would either converge on a local minimum or not be able to produce a useable fit. This was due to the Levinberg-Marquardt method not being able to efficiently step through the parameter space to find a useable minimum, or the method would converge to a local minimum instead of the global minimum. The behavior of the fitted parameters with respect to the initial mass is shown in Figure 3.7 for a variable emissivity index.

Table 3.6. Best Fit Parameters for Planck Opacity Using Region Fluxes

Region	β	Mass [$10^5 M_\odot$]	Mean Surface Density [$M_\odot pc^{-2}$]	Temperature [K]
1	2.20 ± 0.03	70 ± 5	0.14 ± 0.01	22.8 ± 0.5
2	2.30 ± 0.03	19 ± 1	0.18 ± 0.01	22.5 ± 0.3
3	2.35 ± 0.03	9.5 ± 0.6	0.21 ± 0.01	23.3 ± 0.3
4	2.30 ± 0.03	23 ± 1	0.22 ± 0.01	22.2 ± 0.2
5	2.00 ± 0.03	11.4 ± 0.8	0.093 ± 0.007	22.3 ± 0.4
6	1.95 ± 0.03	4.6 ± 0.3	0.087 ± 0.006	23.9 ± 0.4

In order to prevent any false convergences, we establish a range of emissivity index values from 1.6 to 2.9 with increments of 0.05 and found the best fit mass and temperature for each of these indices. This method ended up being significantly more robust than allowing the emissivity index to vary. The relationship between the initial mass and the fitted parameters with a fixed emissivity is shown in Figure 3.8, and shows the results are independent of the initial mass up to a cutoff point where the Levinberg-Marquardt method is unable to produce reliable fits. Determining the best emissivity index was decided by which value returned the lowest reduced χ^2 value. The reduced χ^2 and emissivity index plots are shown in Figure 3.9 for each of the regions of NGC3627, where the minimum value is marked by the red line. The best fit results using this method are shown in Table 3.6 using the opacity suggested by the Planck model and the fitted SEDs are shown in Figure 3.10.

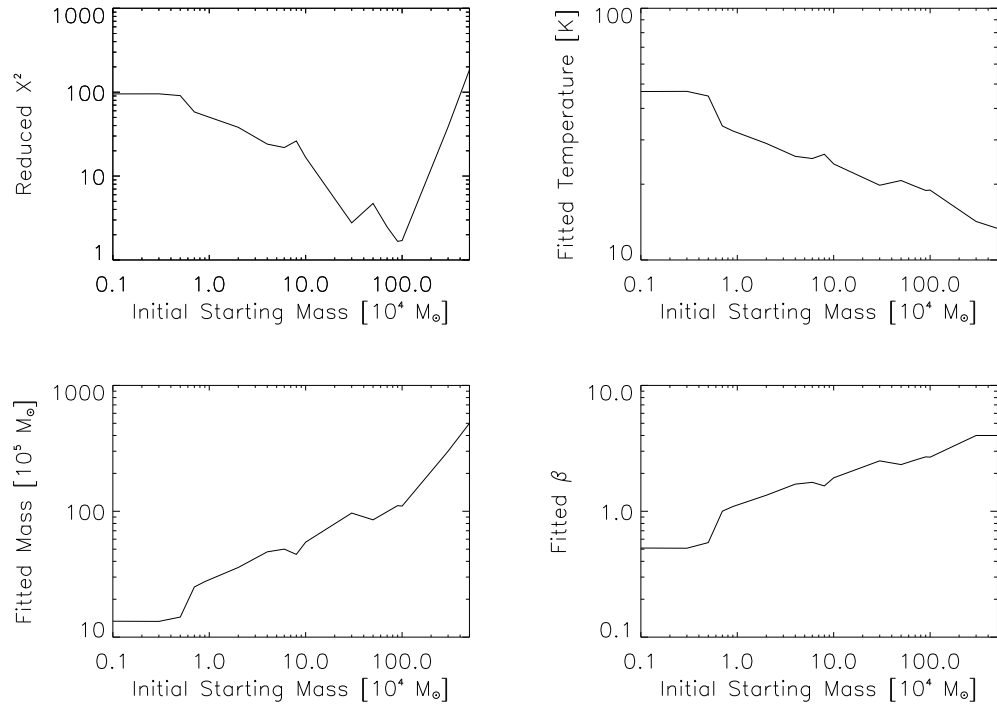


Figure 3.7: Returned SED fitting output for region 1 with Planck opacity for the χ^2 , temperature, mass, and emissivity index with varying initial mass. The top left panels shows the χ^2 values for each starting mass. The top right, bottom left, and bottom right panel show the returned temperatures, mass, and emissivity index.

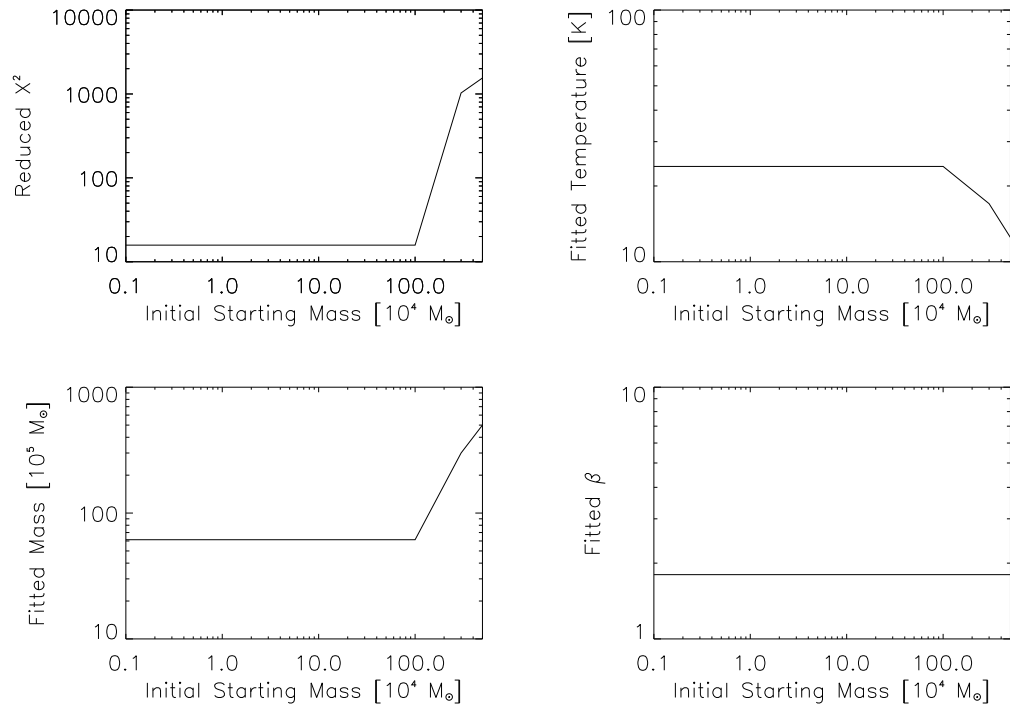


Figure 3.8: The same as Figure 3.7, however the emissivity index has been fixed to 1.8.

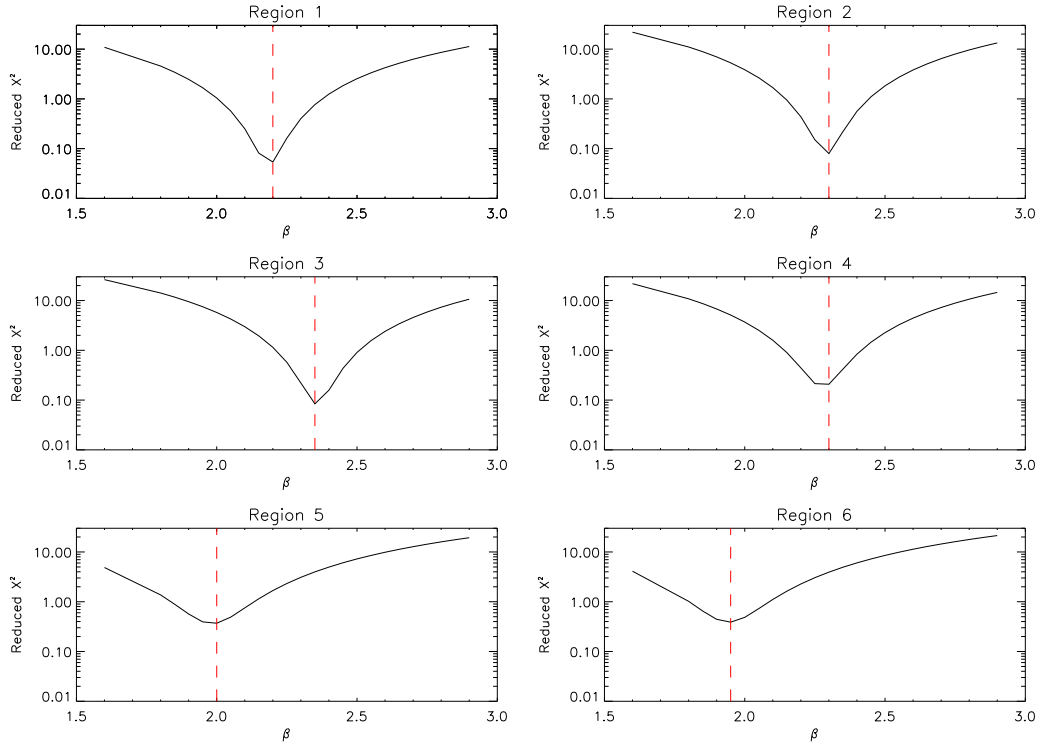
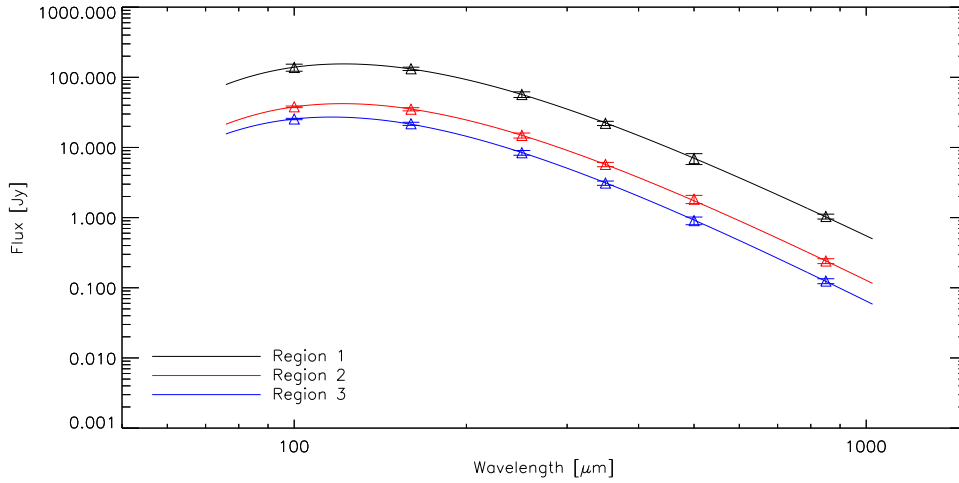
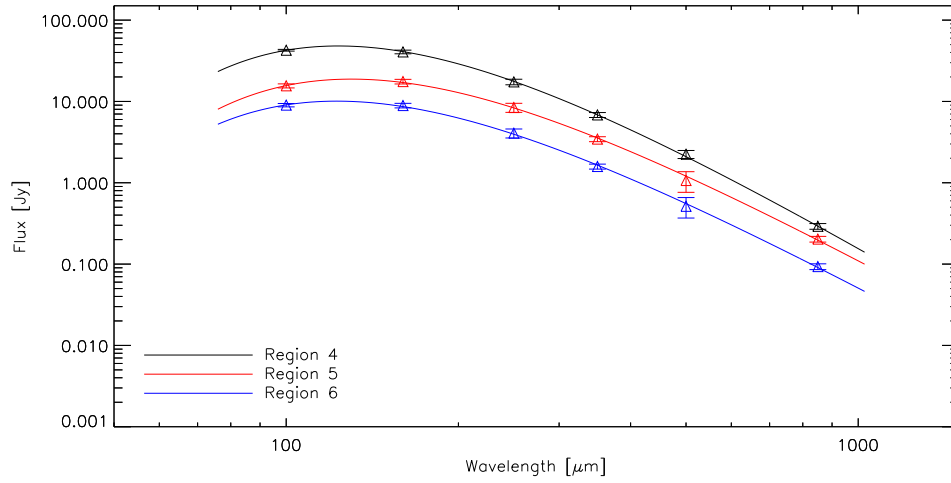


Figure 3.9: Plots for each region from figure 3.1 showing the Fixed emissivity index value and the resulting reduced χ^2 value.



(a) Regions 1, 2, and 3 SEDs



(b) Regions 4, 5, and 6 SEDs

Figure 3.10: SED fits for the flux of each region in Figure 3.1 using the Planck Opacity and allowing the emissivity index, β to vary.

3.3 Discussion

3.3.1 Reliability of Individual Pixel Fits

The issue of whether or not the individual pixel fits are reliable for a variable emissivity index is raised after seeing a strong dependence with the initial mass of the region SED fits shown in Figure 3.7. The argument can be made that since both fits have returned the same temperature ($T_{pixel}=22\pm1$ K and $T_{region}=22.8\pm0.5$ K), emissivity index ($\beta_{pixel}=2.2\pm0.2$ and $\beta_{region}=2.20\pm0.03$), and mass ($M_{pixel}=85\pm38$ M_{\odot} and $M_{region}=82\pm6$ M_{\odot}), after the removal of any initial mass dependence for the region fits, the initial values for the pixel fits are able to locate the global minimum of the χ^2 space regardless of any dependence. The inability of the region fits to successfully converge was due to an overestimate of the initial mass. The mass associated with the $250\mu\text{m}$ flux using the 20 K Li and Draine model was scaling the peak of the SED to values larger than what would be expected using the Planck model which is the opacity used for the region fits. The difference between the two initial masses was large enough for the Levinberg-Marquardt method to fail at finding a minimum. This could be remedied by retooling the initial mass to use the Planck model; however by cycling through emissivity index values, we have provided a more reliable check to our pixel maps by breaking the dependency with the initial mass.

3.3.2 Comparison with Previous Work

We can check the validity of our results by comparing them to previous SED fits using the KINGFISH data (Galametz et al., 2012). The best fit parameters from their fitting were $T=20.2\text{K}\pm 4\text{K}$ and $\beta=2.3\pm 0.2$ which agree within error to our results of $T=22\text{K}\pm 1\text{K}$ and $\beta=2.2\pm 0.2$ for a variable emissivity index. To test agreement with our mass we need to take into account several differences in the two analyses. The differences consist of a lessened flux from the extended structure removal, the opacity values used, and which wavelength was used to calculate the mass. The easiest approach to check whether our masses agree is to use the parameters found by Galametz et al. (2012) with our reduced flux. In order to lessen the amount of free parameters we can use the Li and Draine model to calculate a mass. Using our flux with their fits we get a mass of $276\pm 59\times 10^5 M_{\odot}$ which agrees with the our mass of $230\pm 103\times 10^5 M_{\odot}$.

3.3.3 Effects of $850\mu\text{m}$ Emission

The major difference between our work on NGC3627 and the work done in Galametz et al. (2012) is the presence of the $850\mu\text{m}$ data point. We tested the effect of the $850\mu\text{m}$ data by excluding it from our fits. The returned values for a variable emissivity index are shown in table 3.7. The presence of the $850\mu\text{m}$ has little effect on the returned parameters.

The absence of a significant change in the returned parameters indicates of a lack of any excess submillimeter emission, which can be indicative of an abundance of a very cold ($T \lesssim 10\text{K}$) dust component (Dale et al., 2012). This excess in emission is typically seen in low-metallicity systems such as

Table 3.7. Best Fit Parameters Excluding $850\mu\text{m}$ Emission Using Planck Opacity with Variable Emissivity Index

Region	Average β	Total Mass [$10^5 M_\odot$]	Average Surface Density [$M_\odot pc^{-2}$]	Average Temperature [K]
1	2.3 ± 0.3	75 ± 28	0.15 ± 0.06	22 ± 2
2	2.3 ± 0.3	18 ± 5	0.17 ± 0.05	23 ± 2
3	2.4 ± 0.2	10 ± 1	0.19 ± 0.02	23 ± 1
4	2.2 ± 0.2	22 ± 5	0.21 ± 0.05	23 ± 2
5	2.3 ± 0.2	14 ± 3	0.11 ± 0.03	21 ± 1
6	2.2 ± 0.2	5 ± 1	0.10 ± 0.02	22 ± 1

dwarf galaxies (Madden et al., 2011). The dwarf galaxies showing an excess typically have a metallicity of $\log(\text{O}/\text{H}) + 12 \lesssim 8.3$ (Rémy-Ruyer et al., 2013). An $870\mu\text{m}$ excess has also been seen in NGC3627 (Galametz et al., 2014), but our results show no excess of emission at $850\mu\text{m}$ (evident in Figures 3.3, 3.4, and 3.5). The discrepancy between our results and those of Galametz et al. (2014) may be attributed to their treatment of the removal of the CO J=3-2 emission from their data and possibly to how their $870\mu\text{m}$ data were processed. If their image processing method does not remove extended emission unlike to SCUBA-2, then submillimeter excess might be attributed to diffuse emission regions that are absent in our analysis. However, we would not expect to see any submillimeter excess given that $\log(\text{O}/\text{H}) + 12 = 8.99$ for NGC3627 (Moustakas et al., 2010), well above the estimated threshold where the $850\mu\text{m}$ excesses are typically found.

Chapter 4

Dust-to-Gas Ratio and α_{CO}

4.1 Minimizing Scatter Dust-to-Gas Ratio

In order to determine the amount of molecular hydrogen present, we have chosen to use the dust based method to determine a conversion factor. In our analysis we have used the conversion factor α_{CO} which will return the amount of molecular hydrogen present in terms of its surface density. Calculating the amount of molecular hydrogen present with this method involves reducing the scatter of the dust to gas ratio of each pixel. The dust-to-gas ratio is calculated using

$$\delta_{DGR} = \frac{\Sigma_{dust}}{I_{CO} * \alpha_{CO} + \Sigma_{HI}} \quad (4.1)$$

where δ_{DGR} is the ratio of the surface densities of the dust and gas present in the system, Σ_{dust} is the surface density of the dust in $M_{\odot} \text{ pc}^{-2}$, I_{CO} is the CO intensity in K km s^{-1} , α_{CO} is the conversion factor in units of $M_{\odot} \text{ pc}^{-2} (\text{K km s}^{-1})^{-1}$, and Σ_{HI} is the surface density of the atomic hydrogen. The scatter of the dust-to-gas ratio is determined by finding the standard deviation of

the pixels assuming a Gaussian distribution of dust-to-gas ratios. The atomic hydrogen surface density map had a cut applied where any points below 20% of the maximum were set to zero. Applying this cut removed all but two of the HI emission data points associated with the nucleus region. By removing most of the data, the shape of the scatter in the dust-to-gas ratio vs α_{CO} doesn't show as dramatic of an upturn at higher α_{CO} , so the reported values correspond to the lowest α_{CO} that produces the minimum scatter.

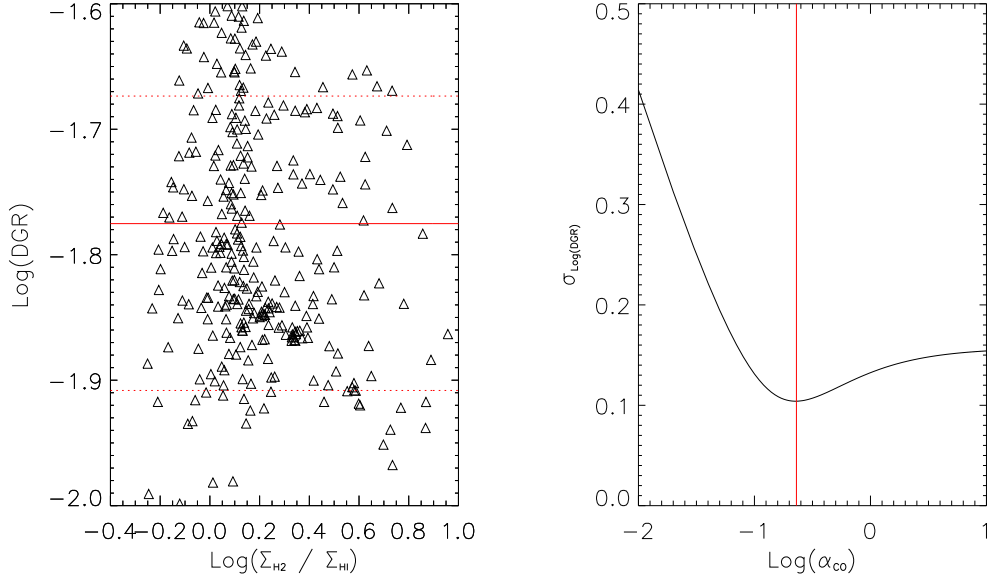
The ideal configuration was determined using the dust surface densities found in Chapter 3 for the Planck and Li and Draine models. The α_{CO} was handled by assigning it a single value over the range of 0.01 to 100 $M_{\odot} \text{ pc}^{-2} (\text{K km s}^{-1})^{-1}$, and the best approximated α_{CO} was determined by the lowest resulting dust-to-gas scatter in the pixels of each region in Figure 3.1. A 7th region was added that is the galaxy without the nucleus. The seventh region was created to accommodate the lack of HI emission in the nucleus of NGC3627 to avoid lowering the α_{CO} value. The surface density of molecular hydrogen is then determined by scaling the CO intensity by the calculated conversion factor, and the dust-to-gas ratio is then found by dividing the dust surface density by the sum of the HI and H₂ surface densities. The results using Planck model with the CO J=1-0 line are shown in Figure 4.1 where the right panel shows the ratio of H₂ to HI against the value of the dust-to-gas ratio in each pixel where the solid red line indicates the mean value for the dust-to-gas ratio, and the dotted lines show the variance associated with the configuration. The left panels of Figure 4.1 show the trend of α_{CO} and the scatter in the dust-to-gas ratio. The numerical values of the dust-to-gas ratio, α_{CO} , and H₂ surface density using both the Planck and Li and Draine dust

masses are shown in Table 4.1. The error on the dust-to-gas ratio in Table 4.1 is the same as the minimum scatter found.

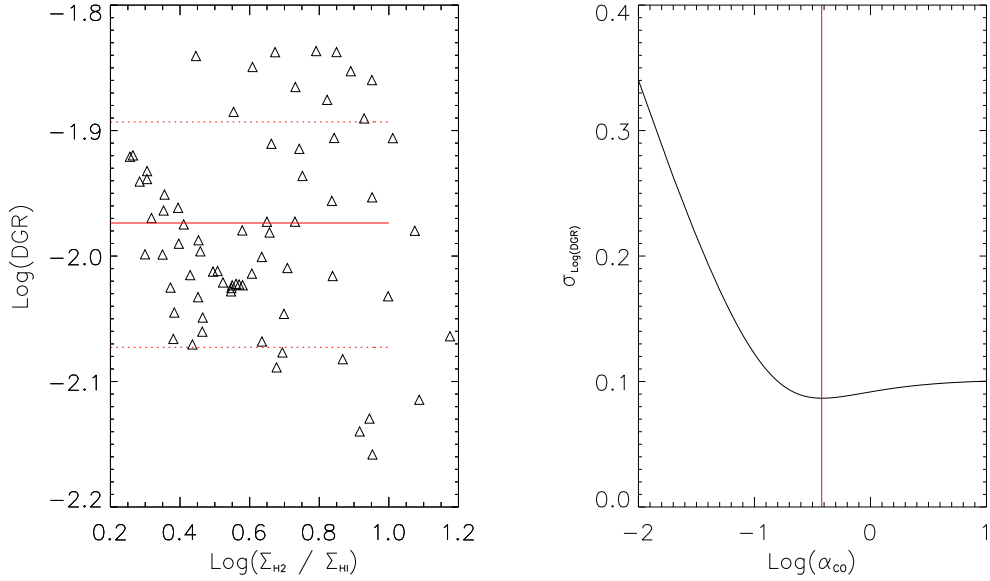
While CO J=1-0 is the standard molecular tracer of H₂ (Bolatto et al., 2013), the observations are not always available. In the situation without CO J=1-0 data an alternative excitation state of CO is used. This was done by Sandstrom et al. (2013) using the CO J=2-1 transition, and by Warren et al. (2010) using the CO J=3-2 transition, where the observations were scaled to the expected intensity of the CO J=1-0 using a ratio of the observed CO transition to the CO J=1-0 transition. We have found the line ratio of the 2-1/1-0 to be 0.39 for the galaxy on a whole, and can examine the effects of using this technique to approximate the CO J=1-0 transition in this method. The results for the dust-to-gas ratio, α_{CO} , and H₂ surface density are shown in Table 4.2.

4.2 Effects of the Dust Model and CO Treatment

The effects of the dust model are mainly seen in the dust-to-gas ratio for the CO J=1-0 emission. Since the major difference in the two models was the outputted mass, this would make sense to see the Li and Draine model produce larger dust-to-gas ratios. The conversion factors show some decrease, but the over all molecular gas surface density all agree within error for each region. The same trend of larger dust-to-gas ratios are seen when scaling the CO J=2-1 emission where the dust-to-gas ratio of the Planck model is larger than the Li and Draine model, however the galaxy as a whole and regions 5

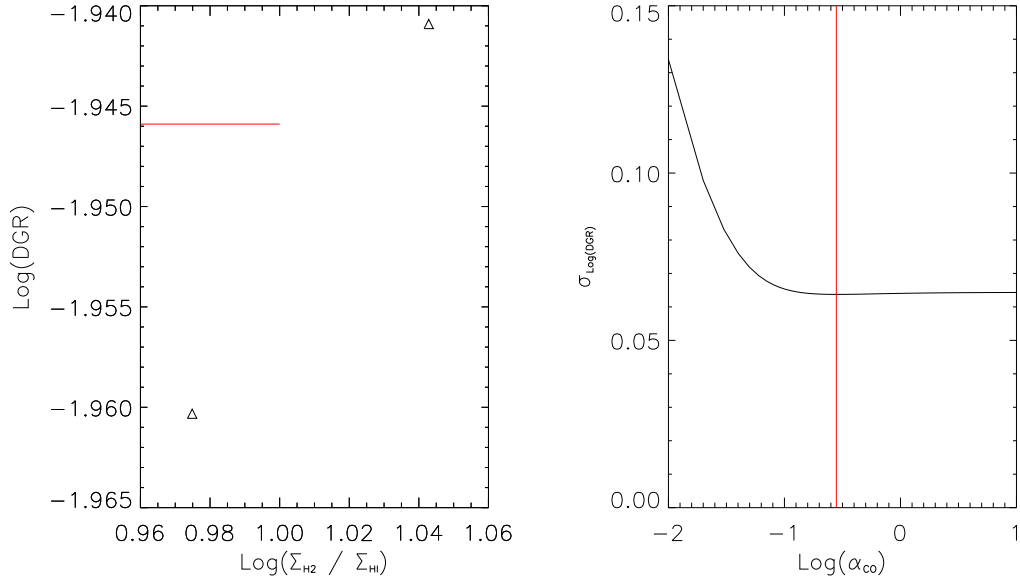


(a) Region 1

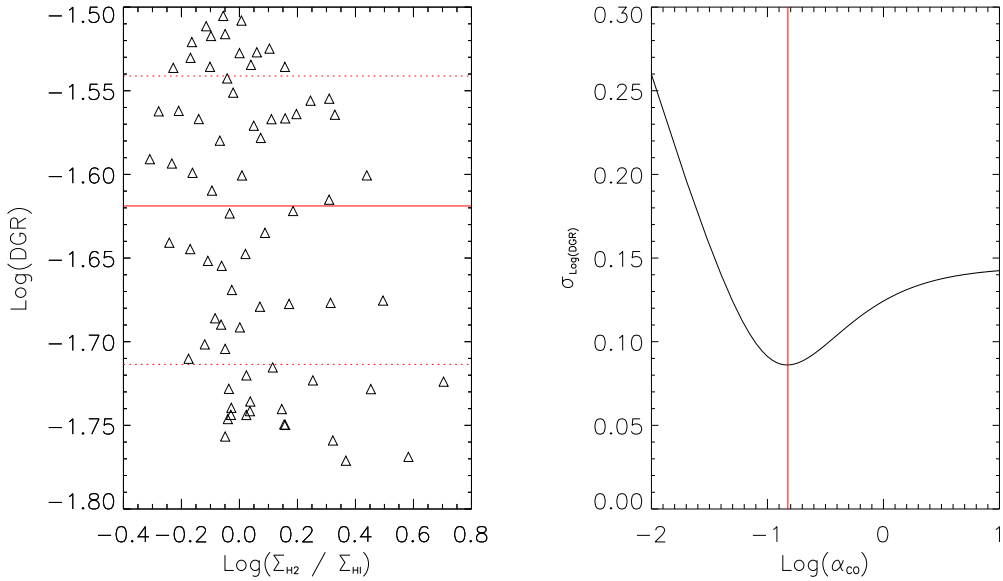


(b) Region 2

Figure 4.1: Plots of the dust-to-gas ratio vs the H_2 to HI surface densities using the calculated α_{CO} and the scatter in the dust-to-gas ratio against the α_{CO} values used. Each were calculated using the CO J=1-0 line and the Planck dust model.

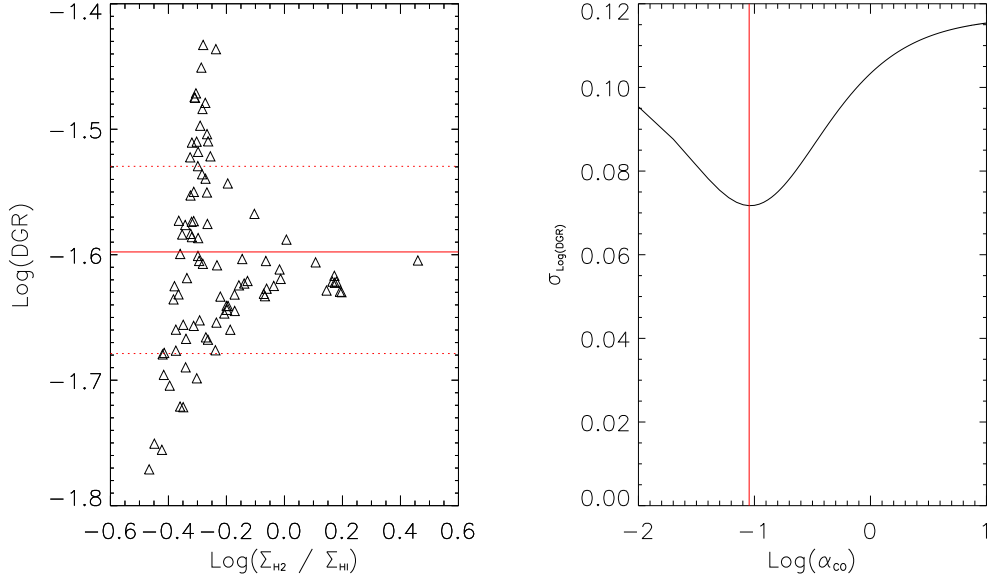


(c) Region 3

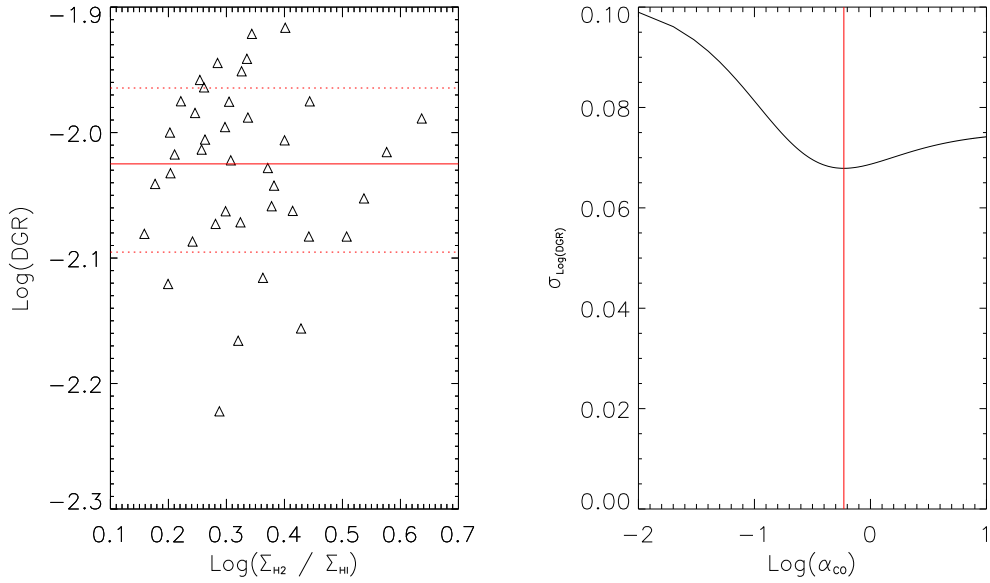


(d) Region 4

Figure 4.1: Plots of the dust-to-gas ratio vs the H_2 to HI surface densities using the calculated α_{CO} and the scatter in the dust-to-gas ratio against the α_{CO} values used. Each were calculated using the CO J=1-0 line and the Planck dust model.

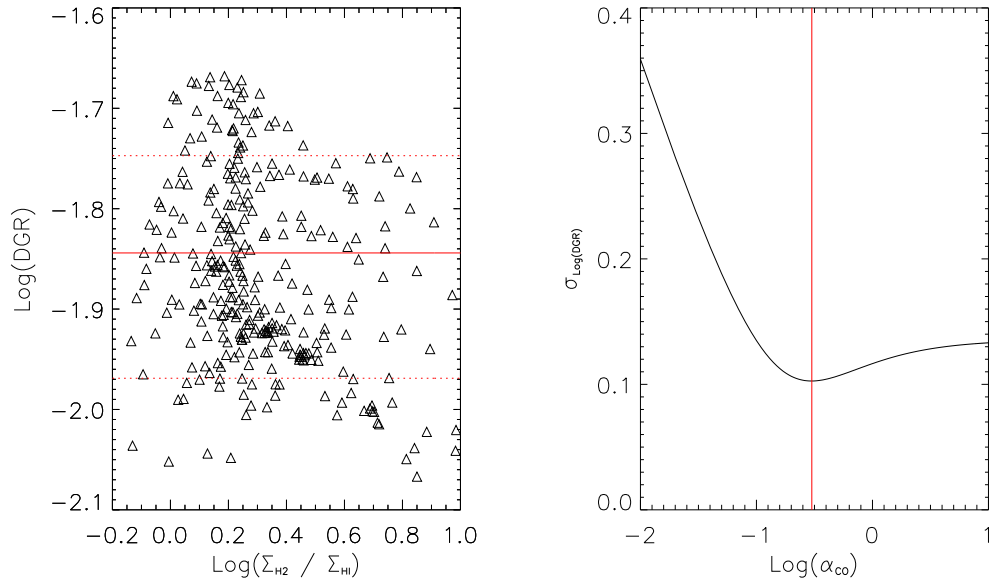


(e) Region 5



(f) Region 6

Figure 4.1: Plots of the dust-to-gas ratio vs the H_2 to HI surface densities using the calculated α_{CO} and the scatter in the dust-to-gas ratio against the α_{CO} values used. Each were calculated using the CO J=1-0 line and the Planck dust model.



(g) Region 1 Without the Nucleus

Figure 4.1: Plots of the dust-to-gas ratio vs the H_2 to HI surface densities using the calculated α_{CO} and the scatter in the dust-to-gas ratio against the α_{CO} values used. Each were calculated using the CO J=1-0 line and the Planck dust model.

Table 4.1. Dust-to-gas ratio, α_{CO} , and H_2 Surface Density using CO J=1-0 Tracer

Opacity Model	Dust-to-Gas Ratio	α_{CO} [$M_{\odot} \text{ pc}^{-2} (\text{K km s}^{-1})^{-1}$]	Average Σ_{H_2} per Pixel [$M_{\odot} \text{ pc}^{-2}$]
Region 1			
Planck	0.017 ± 0.004	0.230 ± 0.005	4 ± 3
Li and Draine	0.07 ± 0.02	0.230 ± 0.005	4 ± 3
Region 2			
Planck	0.011 ± 0.002	0.380 ± 0.005	10 ± 3
Li and Draine	0.044 ± 0.009	0.410 ± 0.005	11 ± 4
Region 3			
Planck	0.011 ± 0.002	0.280 ± 0.005	10 ± 2
Li and Draine	0.049 ± 0.008	0.290 ± 0.005	11 ± 2
Region 4			
Planck	0.024 ± 0.005	0.150 ± 0.005	4 ± 2
Li and Draine	0.11 ± 0.02	0.150 ± 0.005	4 ± 1
Region 5			
Planck	0.03 ± 0.09	0.090 ± 0.005	1.2 ± 0.6
Li and Draine	0.11 ± 0.02	0.100 ± 0.005	1.4 ± 0.6
Region 6			
Planck	0.010 ± 0.001	0.590 ± 0.005	6 ± 1
Li and Draine	0.040 ± 0.006	0.610 ± 0.005	6 ± 1
No Nucleus			
Planck	0.014 ± 0.004	0.300 ± 0.005	5 ± 3
Li and Draine	0.06 ± 0.02	0.310 ± 0.005	5 ± 3

Table 4.2. Dust-to-gas ratio, α_{CO} , and H_2 Surface Density using CO J=2-1 Tracer

Opacity Model	Dust-to-Gas Ratio	α_{CO} [$M_{\odot} \text{ pc}^{-2} (\text{K km s}^{-1})^{-1}$]	Average Σ_{H_2} per Pixel [$M_{\odot} \text{ pc}^{-2}$]
Region 1			
Planck	0.01 ± 0.01	0.390 ± 0.005	8 ± 5
Li and Draine	0.048 ± 0.01	0.410 ± 0.005	8 ± 5
Region 2			
Planck	0.0128 ± 0.004	0.310 ± 0.005	8 ± 3
Li and Draine	0.054 ± 0.003	0.330 ± 0.005	8 ± 3
Region 3			
Planck	0.014 ± 0.009	0.270 ± 0.005	9 ± 2
Li and Draine	0.037 ± 0.02	0.280 ± 0.005	9 ± 3
Region 4			
Planck	0.0085 ± 0.004	0.520 ± 0.005	15 ± 1
Li and Draine	0.037 ± 0.005	0.540 ± 0.005	16 ± 6
Region 5			
Planck	0.013 ± 0.008	0.340 ± 0.005	4 ± 1
Li and Draine	0.054 ± 0.008	0.360 ± 0.005	5 ± 2
Region 6			
Planck	0.0108 ± 0.004	0.450 ± 0.005	5 ± 1
Li and Draine	0.037 ± 0.004	0.460 ± 0.005	5 ± 2
No Nucleus			
Planck	0.009 ± 0.007	0.550 ± 0.005	10 ± 6
Li and Draine	0.034 ± 0.008	0.580 ± 0.005	10 ± 6

and 6 also show a larger α_{CO} which increases the average surface density of those regions.

The CO J=2-1 emission shows the same trend with the dust models as the CO J=1-0 which show that the Li and Draine model yields a higher dust-to-gas ratio. The difference between the two transitions is seen in the value of α_{CO} . By using the CO J=2-1 emission with a 2-1/1-0 ratio applied, the conversion factor is significantly increased for each region by a factor of 1.5 to a factor 3.5. Despite the increase in α_{CO} we still see the same masses for all of the regions except for regions 4 and 5. The conflicting surface densities are due to the behavior of the 2-1/1-0 ratio in these regions. Region 5 displays a large gradient in 2-1/1-0 ratio as you move south along the spiral arm, and the 75% percent decrease in region 4 is due to the entire region's 2-1/1-0 ratio is nearly double the mean value of the galaxy. The difference in α_{CO} between the CO J=1-0 emission and the converted CO J=2-1 emission raise the question of which value should we trust? The answer lies primarily in whether we are concerned with the dense ISM or diffuse ISM. If we are interested in the diffuse ISM, Wilson & Scoville (1990) showed in M33 that over half of the CO J=1-0 emission is due to diffuse gas, so we should use the CO J=1-0 component. Since we are concerned with determining molecular hydrogen abundances, we would focus on the more dense regions of NGC3627's ISM, therefore we should use the scaled CO J=2-1 emission as our tracer of choice.

4.3 Validity of Results

If NGC3627 were to have Milky Way like values for α_{CO} , we would expect $\alpha_{CO} \approx 4$ (Sandstrom et al., 2013). Except the α_{CO} values given for the the CO J=2-1 emission suggest NGC3627 is similar to U/LIRG type galaxies whose α_{CO} range has been given as $0.3 - 1.3 \text{ M}_{\odot} \text{ pc}^{-2} (\text{K km s}^{-1})^{-1}$ (Downes & Solomon, 1998), and the CO J=1-0 emission falls below the U/LIRG range. Even our smallest dust-to-gas ratio falls above the typical values found in late type galaxies of $0.005 - 0.01$ (Smith et al., 2012). If we compare our results to values recently calculated using the same method by Sandstrom et al. (2013), we find that our conversion factor is much lower than their average of $1.2 \text{ M}_{\odot} \text{ pc}^{-2} (\text{K km s}^{-1})^{-1}$, and our dust-to-gas ratio is much larger than their average of ≈ 0.017 . The low α_{CO} and high dust-to-gas ratios are due mainly to the filtering process applied to our data (§2.3), in particular, the amount of HI removed by the filtering. If we use the unfiltered gas data, and use the Planck dust surface densities, we can establish a minimum for our expected dust-to-gas ratios. Doing this we get the results shown in Table 4.3. Applying the filtering has increased the CO J=2-1 emission to within a reasonable value of what Sandstrom et al. (2013) reported, however our dust-to-gas ratio is much lower than their values. The lower dust-to-gas ratios are attributed to our examining only the cold component, and Sandstrom et al. (2013) using the full infrared SED to determine a dust mass, which will include the warm component and the PAH contribution as well as the cold component we use.

Table 4.3. CO Transition, α_{CO} , and H_2 Surface Density using the Planck dust model with No Emission Filtering

Region	Dust-to-Gas Ratio	α_{CO} [$M_{\odot} \text{ pc}^{-2} (\text{K km s}^{-1})^{-1}$]	Average Σ_{H_2} per Pixel [$M_{\odot} \text{ pc}^{-2}$]
Region 1			
J=1-0	0.0030 ± 0.0008	0.820 ± 0.005	20 ± 9
J=2-1	0.0023 ± 0.0003	1.410 ± 0.005	30 ± 17
Region 2			
J=1-0	0.0026 ± 0.0005	1.050 ± 0.005	31 ± 10
J=2-1	0.0025 ± 0.0002	1.270 ± 0.005	33 ± 14
Region 3			
J=1-0	0.0032 ± 0.0003	0.580 ± 0.005	25 ± 5
J=2-1	0.0040 ± 0.0002	0.520 ± 0.005	18 ± 5
Region 4			
J=1-0	0.008 ± 0.001	0.130 ± 0.005	4 ± 1
J=2-1	0.0030 ± 0.0003	1.140 ± 0.005	35 ± 12
Region 5			
J=1-0	0.0031 ± 0.0003	0.670 ± 0.005	13 ± 5
J=2-1	0.0021 ± 0.0002	1.810 ± 0.005	26 ± 8
Region 6			
J=1-0	0.0012 ± 0.0002	3.360 ± 0.005	47 ± 7
J=2-1	0.0015 ± 0.0002	2.920 ± 0.005	33 ± 9
No Nucleus			
J=1-0	0.0022 ± 0.0006	1.370 ± 0.005	30 ± 13
J=2-1	0.0018 ± 0.0002	2.100 ± 0.005	41 ± 23

Chapter 5

Summary

We have shown two new continuum maps of NGC3627 in $450\mu\text{m}$ and $850\mu\text{m}$ from SCUBA-2, and analyzed them with the goal of determining a dust-to-gas mass ratio, a suitable conversion factor, and the amount of molecular hydrogen present in the galaxy as a whole and 5 individual regions picked based upon the morphological features of the galaxy. In order to produce our results we utilized data from the KINGFISH survey (Kennicutt et al., 2011) and NGLS (Wilson et al., 2012) to calculate a dust mass using SED fitting on a pixel by pixel scale and a total of the flux. The dust-to-gas ratio, α_{CO} , and molecular gas surface density was determined using a minimization of the dust-to-gas scatter in the galaxy using data from the Nobeyama 45-m CO J=1-0 survey (Kuno et al., 2007), HERACLES (Reuter et al., 1996), and THINGS (Walter et al., 2008). The important results from our analysis are as follows:

1. We have created high quality images of NGC3627 using the Starlink program MAKEMAP by incorporating masks in the AST and FLT portions of MAKEMAP and adding a high-pass filter of $175''$ to in the FLT

portion of MAKEMAP in order to reduce the noise in the final image products.

2. Using results from the KINGFISH survey (Kennicutt et al., 2011), we were able to show an excess emission was present in the $450\mu\text{m}$ observations that was most likely due to calibration issue rather than a physical process in the ISM of NGC3627.
3. The results of the SED analysis were tested using two competing dust models calculated by the Planck observation group (Planck Collaboration et al., 2011), Li and Draine (Li & Draine, 2001), and the third model was tested using the opacity values from the Planck model with an emissivity index that was allowed to vary over 5 regions of NGC3627 and the galaxy as a whole. The results are shown in Tables 3.3, 3.4, 3.5, and 3.6. The calculated temperatures for the entire galaxy are 23 ± 2 K, 24 ± 2 K, 22 ± 1 K, and 22 ± 2 K for the Planck model, Li and Draine model, free emissivity index, and regional flux, respectively. The calculated masses using the Planck opacity are $52\pm 23 M_{\odot}$, $73\pm 31 M_{\odot}$, and $75\pm 28 M_{\odot}$ for the Planck model, free emissivity index and regional flux, respectively. The mass returned for the Li and Draine model is $230\pm 103 M_{\odot}$. When the emissivity index was allowed to vary the returned values were 2.2 ± 0.2 and 2.3 ± 0.3 for the pixel by pixel fits and regional flux, respectively. The results from the fitting agreed with previous fittings by Galametz et al. (2012).
4. There was no observed excess of emission at $850\mu\text{m}$ in our fits, which contradict the results of Galametz et al. (2014) using $870\mu\text{m}$ emission

from LABOCA. The difference in results can be attributed to their handling of the molecular contribution and whether the reduction process of the $870\mu\text{m}$ emission preserves extended features.

5. Varying dust models used in determining the dust-to-gas ratio showed that the Li and Draine model yielded higher dust-to-gas ratios than the Planck model. This is due to the Li and Draine model having a higher dust surface density.
6. Using either the CO J=1-0 or scaling the CO J=2-1 using a 2-1/1-0 ratio will not effect the dust-to-gas ratio, but will change the conversion factor. Using the scaled CO J=2-1 emission will provide higher conversion factors because the CO J=1-0 emission will trace more diffuse gas (Wilson & Scoville, 1990) where atomic hydrogen will be more dominant than its molecular counterpart.
7. The gas data with the extended emission removed yielded α_{CO} values typical of U/LIRGs. This was due to the significant portion of the HI emission that is made up of extended emission. When using unfiltered gas data, we recovered conversion factors within the vicinity of Sandstrom et al. (2013). The dust-to-gas ratios however were much smaller than the reported values of Sandstrom et al. (2013). The lower gas-to-dust ratios is attributed to a larger dust mass that contains not only the cold component we used, but their analysis also include the warm component and PAH contribution.

Future work on the NGC3627 and NGCLS SCUBA-2 catalogue would include:

1. Investigating the excess of emission at the $450\mu\text{m}$ to determine its source. If the unphysical excess $450\mu\text{m}$ could be found, the resolution of this analysis could be reduced and higher quality SED fitting could take place.
2. Implementing a MCMC method to determine an α_{CO} values for each pixel would help to decrease the overall scatter of the dust-to-gas ratio and provide dust-to-gas ratio, α_{CO} , and H_2 surface density maps to go with the SED fitting maps. These maps could be used to determine correlations in α_{CO} with parameters returned from the SED fitting.
3. Implementing the SED and dust-to-gas ratio analysis on other targets in the NGLS catalogue would allow for a large scale statistical survey over the properties of the dust in the local Universe.

Bibliography

- Abel, T., Anninos, P., Zhang, Y., & Norman, M. L. 1997, *New A*, 2, 181
- Abergel, A., Ade, P. A. R., Aghanim, N., et al. 2013, *ArXiv e-prints*, arXiv:1312.1300
- Barrett, A. H. 1964, *Nature*, 202, 475
- Bolatto, A. D., Wolfire, M., & Leroy, A. K. 2013, *ARA&A*, 51, 207
- Boulanger, F., Abergel, A., Bernard, J.-P., et al. 1996, *A&A*, 312, 256
- Burden, R. L., & Faires, J. D. 2001, *Numerical analysis*
- Calzetti, D., Wu, S.-Y., Hong, S., et al. 2010, *ApJ*, 714, 1256
- Chapin, E. L., Berry, D. S., Gibb, A. G., et al. 2013, *MNRAS*, 430, 2545
- Dale, D. A., Aniano, G., Engelbracht, C. W., et al. 2012, *ApJ*, 745, 95
- Dempsey, J. T., Friberg, P., Jenness, T., et al. 2013, *MNRAS*, 430, 2534
- Dobbs, C. L., Krumholz, M. R., Ballesteros-Paredes, J., et al. 2013, *ArXiv e-prints*, arXiv:1312.3223
- Downes, D., & Solomon, P. M. 1998, *ApJ*, 507, 615
- Drabek, E., Hatchell, J., Friberg, P., et al. 2012, *MNRAS*, 426, 23
- Draine, B. T. 2011, *Physics of the Interstellar and Intergalactic Medium*

- Draine, B. T., Dale, D. A., Bendo, G., et al. 2007, *ApJ*, 663, 866
- Dumke, M., Krause, M., Beck, R., et al. 2011, in *Astronomical Society of the Pacific Conference Series*, Vol. 446, *Galaxy Evolution: Infrared to Millimeter Wavelength Perspective*, ed. W. Wang, J. Lu, Z. Luo, Z. Yang, H. Hua, & Z. Chen, 111
- Elmegreen, B. G. 1987, *ApJ*, 312, 626
- Ferrière, K. M. 2001, *Reviews of Modern Physics*, 73, 1031
- Field, G. B., & Saslaw, W. C. 1965, *ApJ*, 142, 568
- Freedman, W. L., Madore, B. F., Gibson, B. K., et al. 2001, *ApJ*, 553, 47
- Galametz, M., Kennicutt, R. C., Albrecht, M., et al. 2012, *MNRAS*, 425, 763
- Galametz, M., Albrecht, M., Kennicutt, R., et al. 2014, *MNRAS*, 439, 2542
- Gavin, H. P. 2013, *The Levenberg-Marquardt method for nonlinear least squares curve-fitting problems*, Tech. rep., Department of Civil and Environmental Engineering, Duke University
- Glover, S. C. O. 2003, *ApJ*, 584, 331
- Gould, R. J., & Salpeter, E. E. 1963, *ApJ*, 138, 393
- Gry, C., Boulanger, F., Nehmé, C., et al. 2002, *A&A*, 391, 675
- Haynes, M. P., Giovanelli, R., & Roberts, M. S. 1979, *ApJ*, 229, 83
- Herbst, E., Chang, Q., & Cuppen, H. M. 2005, *Journal of Physics Conference Series*, 6, 18

- Hildebrand, R. H. 1983, QJRAS, 24, 267
- Holland, W. S., Bintley, D., Chapin, E. L., et al. 2013, MNRAS, 430, 2513
- Israel, F. P., Bontekoe, T. R., & Kester, D. J. M. 1996, A&A, 308, 723
- Jura, M. 1975, ApJ, 197, 575
- Kennicutt, R. C., & Evans, N. J. 2012, ARA&A, 50, 531
- Kennicutt, R. C., Calzetti, D., Aniano, G., et al. 2011, PASP, 123, 1347
- Kennicutt, Jr., R. C., Armus, L., Bendo, G., et al. 2003, PASP, 115, 928
- Kim, W.-T., & Ostriker, E. C. 2001, ApJ, 559, 70
- Krumholz, M. R. 2014, ArXiv e-prints, arXiv:1402.0867
- Kuno, N., Sato, N., Nakanishi, H., et al. 2007, PASJ, 59, 117
- Kwan, J. 1979, ApJ, 229, 567
- Leroy, A. K., Walter, F., Bigiel, F., et al. 2009, AJ, 137, 4670
- Leroy, A. K., Bolatto, A., Gordon, K., et al. 2011, ApJ, 737, 12
- Li, A. 2005, in American Institute of Physics Conference Series, Vol. 761, The Spectral Energy Distributions of Gas-Rich Galaxies: Confronting Models with Data, ed. C. C. Popescu & R. J. Tuffs, 123–133
- Li, A., & Draine, B. T. 2001, ApJ, 554, 778
- Madden, S. C., Poglitsch, A., Geis, N., Stacey, G. J., & Townes, C. H. 1997, ApJ, 483, 200

- Madden, S. C., Galametz, M., Cormier, D., et al. 2011, in EAS Publications Series, Vol. 52, EAS Publications Series, ed. M. Röllig, R. Simon, V. Ossenkopf, & J. Stutzki, 95–101
- Markwardt, C. B. 2009, in Astronomical Society of the Pacific Conference Series, Vol. 411, Astronomical Data Analysis Software and Systems XVIII, ed. D. A. Bohlender, D. Durand, & P. Dowler, 251
- McKee, C. F., & Ostriker, E. C. 2007, ARA&A, 45, 565
- Moustakas, J., Kennicutt, Jr., R. C., Tremonti, C. A., et al. 2010, ApJS, 190, 233
- Palla, F., Salpeter, E. E., & Stahler, S. W. 1983, ApJ, 271, 632
- Parker, E. N. 1966, ApJ, 145, 811
- Pirronello, V., Liu, C., Shen, L., & Vidali, G. 1997, ApJ, 475, L69
- Planck Collaboration, Abergel, A., Ade, P. A. R., et al. 2011, A&A, 536, A25
- Rémy-Ruyer, A., Madden, S. C., Galliano, F., et al. 2013, A&A, 557, A95
- Reuter, H.-P., Sievers, A. W., Pohl, M., Lesch, H., & Wielebinski, R. 1996, A&A, 306, 721
- Rosenberg, M. J. F., Kazandjian, M. V., van der Werf, P. P., et al. 2014, A&A, 564, A126
- Rots, A. H. 1978, AJ, 83, 219
- Rybicki, G. B., & Lightman, A. P. 1986, Radiative Processes in Astrophysics

- Sandstrom, K. M., Leroy, A. K., Walter, F., et al. 2013, *ApJ*, 777, 5
- Skibba, R. A., Engelbracht, C. W., Dale, D., et al. 2011, *ApJ*, 738, 89
- Smith, M. W. L., Gomez, H. L., Eales, S. A., et al. 2012, *ApJ*, 748, 123
- Soida, M., Urbanik, M., Beck, R., Wielebinski, R., & Balkowski, C. 2001, *A&A*, 378, 40
- Tabatabaei, F. S., Braine, J., Xilouris, E. M., et al. 2014, *A&A*, 561, A95
- Walcher, J., Groves, B., Budavári, T., & Dale, D. 2011, *Ap&SS*, 331, 1
- Walter, F., Brinks, E., de Blok, W. J. G., et al. 2008, *AJ*, 136, 2563
- Warren, B. E., Wilson, C. D., Israel, F. P., et al. 2010, *ApJ*, 714, 571
- Watanabe, Y., Sorai, K., Kuno, N., & Habe, A. 2011, *MNRAS*, 411, 1409
- Wilson, C. D., & Scoville, N. 1990, *ApJ*, 363, 435
- Wilson, C. D., Warren, B. E., Israel, F. P., et al. 2012, *MNRAS*, 424, 3050
- Wolfire, M. G., McKee, C. F., Hollenbach, D., & Tielens, A. G. G. M. 2003, *ApJ*, 587, 278
- Wolfire, M. G., Tielens, A. G. G. M., Hollenbach, D., & Kaufman, M. J. 2008, *ApJ*, 680, 384