

DUST AND GAS IN NGC3627

DUST AND GAS IN NGC3627 USING OBSERVATIONS FROM
SCUBA-2

By

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Abstract

Saw some dust and wanted to do something about it!

To my family and Poly.

Acknowledgements

When life looks like easy street, there is danger at your door... -Robert Hunter

Thank Chris and group members of course. Don't forget Christian!

Table of Contents

Descriptive Notes	ii
Abstract	iii
Acknowledgements	v
List of Figures	viii
List of Tables	x
 Chapter 1 Introduction	 1
1.1 The Physical Processes of Star Formation	1
1.1.1 GMC Formation	2
1.1.2 Molecular Hydrogen Formation	3
1.1.3 Dissociation of Molecular Hydrogen	6
1.2 Determining H ₂ Abundance	8
1.2.1 Methods for Determining CO-to-H ₂ Conversion Factor	9
1.3 Determining Dust Mass	12
1.4 NGC 3627	15
 Chapter 2 Observations and Data Preparation	 19
2.1 SCUBA-2	19
2.2 Image Creation and Properties	20
2.2.1 Beam Shape of the 450 μ m and 850 μ m Data	28

2.3	Ancillary Data	28
2.3.1	Key Insights on Nearby Galaxies: a Far-Infrared Survey with Herschel (KINGFISH)	30
2.3.2	Nearby Galaxies Legacy Survey (NGLS)	31
2.3.3	Nobeyama 45-m	38
2.3.4	Hetrodyne Reciever Array CO-Line Extragalactic Sur- vey (HERACLES)	38
2.3.5	The HI Nearby Galaxy Survey (THINGS)	40
2.4	Data Preparation for Analysis	42
2.4.1	Accounting for the $450\mu\text{m}$ Error Beam	44
2.4.2	Extended Structure Removal via MAKEMAP	45
Chapter 3 Spectral Energy Distribution Analysis		50
3.1	SED Fitting Method	50
3.2	Fitting the Spectral Energy Distribution	53
3.2.1	Pixel SED Fits	53
3.2.2	Total Region Flux SED Fits	63
3.3	Discussion	69
3.3.1	Reliability of Individual Pixel Fits	69
3.3.2	Comparison with Previous Work	69
3.3.3	Effects of $850\mu\text{m}$ Emission	70

List of Figures

1.1	Optical composite image of NGC3627 courtesy of the SINGS 5 th enhanced data release (Kennicutt et al., 2003). Blue areas represent the B band, green areas represent V band and red areas show the H α band.	17
2.1	Flux Values vs High-Pass Filter Sizes	23
2.2	450 μ m High-Pass Filter Images	24
2.3	850 μ m High-Pass Filter Images	25
2.4	NGC3627 450 μ m Observations	26
2.5	NGC3627 850 μ m Observations	27
2.6	SCUBA-2 Calibration and Beams	29
2.7	NGC3627 100 μ m Observations	32
2.8	NGC3627 160 μ m Observations	33
2.9	NGC3627 250 μ m Observations	34
2.10	NGC3627 350 μ m Observations	35
2.11	NGC3627 500 μ m Observations	36
2.12	NGC3627 CO J=3-2 Observations	37
2.13	NGC3627 CO J=1-0 Observations	39
2.14	NGC3627 CO J=2-1 Observations	41
2.15	NGC3627 HI Observations	43

2.16	100 μ m Filtering Steps	48
2.17	HI surface density filtering Steps	49
3.1	NGC3627 Regions	54
3.2	Li and Draine Model SED Fit Quality Using 450 μ m Data	56
3.3	SED Parameter Maps	58
3.4	Planck Model SED Fit Quality Using 500 μ m Data	60
3.5	Li and Draine Model SED Fit Quality Using 500 μ m Data	61
3.6	Emissivity as a Free Parameter SED Fit Quality using the 500 μ m Data	62
3.7	Initial Mass Dependence and Convergence of SED Fits for Region Fluxes and Variable Emissivity Index	65
3.8	Initial Mass Dependence and Convergence of SED Fits for Region Fluxes and Fixed Emissivity Index	66
3.9	Region Flux Best Emissivity Index Selection	67
3.10	Region Flux SED Fits	68

List of Tables

2.1	Properties of NGC3627 SCUBA-2 Observations	26
2.2	Properties of NGC3627 KINGFISH Observations	31
2.3	Properties of NGC3627 NGLS Observations	32
2.4	Properties of NGC3627 Nobeyama 45-m Observations	38
2.5	Properties of NGC3627 HERACLES Observations	40
2.6	Properties of NGC3627 THINGS Observations	42
3.1	Systematic Calibration Uncertainties for SCUBA-2 and KINGFISH Observations	51
3.2	Best Fit Parameters for Planck Model Using $500\mu\text{m}$ Emission for Pixel Fits	57
3.3	Best Fit Parameters for Li and Draine Model Using $500\mu\text{m}$ Emis- sion for Pixel Fits	59
3.4	Best Fit Parameters for β As A Free Parameter Using $500\mu\text{m}$ Emis- sion for Pixel Fits	59
3.5	Total Distance to 1 to 1 Line	63
3.6	Best Fit Parameters for Planck Opacity Using Region Fluxes . . .	64
3.7	Best Fit Parameters Excluding $850\mu\text{m}$ Emission	71

Chapter 1

Introduction

1.1 The Physical Processes of Star Formation

Star formation is one of the fundamental processes in astrophysics that affects not just stellar and planetary formation but also dictates the behavior in galaxy formation and evolution (Kennicutt & Evans, 2012). The study of star formation itself can be broken into several areas such as: the processes that trigger collapse and how the collapse behaves, dictating how and what type of stars can form from a given region or cloud, or how the inflow of new material can alter current star forming environments. These different sub-branches of star formation are conveniently sorted into hierarchal schemes that range from Mpc scales seen in gas accretion from the intergalactic medium, to properties of accretion disks on scales of the order R_{\odot} or AU (Kennicutt & Evans, 2012). In this body of work we focus on the physical properties of the dust and abundance of gas in the giant molecular cloud (GMC) phase of star formation prior to stellar collapse or fragmentation. However based on the resolution of

our target galaxy, we will not be able to examine individual GMCs instead we will need to focus our work on giant molecular associations (GMAs).

1.1.1 GMC Formation

After intergalactic gas has accreted into a host galaxy, it condense and lead to the formatoin of molecular clouds (Kennicutt & Evans, 2012). The dominant process of GMC formation is divided between two camps, either a “bottom-up” or “top-down” formation scenario (McKee & Ostriker, 2007). The bottom-up scenario consists of small clouds of cold HI coagulating to eventually form a GMC (Field & Saslaw, 1965; Kwan, 1979). The major concern with the bottom-up scenario is if we include heating mechanisms in the cloud, coagulation will cease before the observed masses are reached (McKee & Ostriker, 2007). In fact, the time scales required to form a cloud typical of those we observe would take around 10^8 years to accumulate a minimum of $10^5 M_{\odot}$, which is much greater than the expected GMC lifetime (McKee & Ostriker, 2007).

The alternative formation scenario, top-down, postulates the GMC formation comes from instabilities in the diffuse ISM causing the clouds to collapse from their surrounding medium (McKee & Ostriker, 2007). Two main instabilities are thought to be responsible for the collapse. The first type of instability is a Parker instability, which involves distortion of magnetic field lines in the mid-plane of the galaxy, and at these distortions gas will begin to accumulate (Parker, 1966; Dobbs et al., 2013). The second instability responsible for collapse comes in the style of a more complex Jeans instability which is deter-

mined further by the amount of rotational shear present (McKee & Ostriker, 2007). If a strong rotational shear is present, such as in the inter arm region of a spiral galaxy, then the collapse will be due to the accumulation of gas due to differential rotation in a process similar to the formation of spiral arms known as swing amplification will occur (McKee & Ostriker, 2007; ?). If no rotational shear is present, such as the inner regions of a galaxy or spiral arms, then the collapse can be attributed to a magneto-Jeans instability which is a Jeans instability where the collapse is aided by magnetic fields (Elmegreen, 1987; Kim & Ostriker, 2001).

1.1.2 Molecular Hydrogen Formation

Regardless of either collapse or coagulation, molecular hydrogen is being formed inside the cloud. The formation of molecular hydrogen can be summarized using either the two body reaction, three body reaction, formation using a free electron or proton, or surface formation (Krumholz, 2014). The two body formation scenario is the simplest reaction using two hydrogen atoms to produce molecular hydrogen, reaction 1.1.



However the two body formation is not the major mechanism in the formation of H_2 due to the requirement of a forbidden photon that arises from the combination of hydrogen atoms in the ground state (Gould & Salpeter, 1963). If one of the hydrogen atoms is excited, the transition can occur and

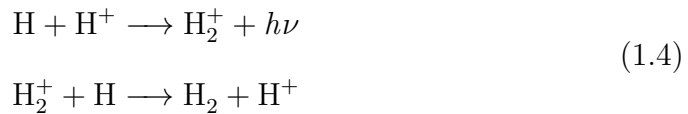
molecular hydrogen is formed, but the amount of excited hydrogen atoms in the temperature ranges typical of the cold and warm phases of the ISM are expected to be nearly nonexistent (Krumholz, 2014).

The second formation scenario listed, the three body formation, involves three hydrogen atoms coming together to form molecular hydrogen with a spare hydrogen atom shown in reaction 1.2.



The required density for the three body formation to occur is on the order of 10^8 cm^{-3} (Palla et al., 1983; Abel et al., 1997), while the typical GMC density is on the order of 300 cm^{-3} . The disparity between these densities eliminates any possibility of this being the primary mechanism to form molecular hydrogen.

An alternative to the two or three body reactions uses either an electron or proton to ionize the hydrogen forming either H^- or H_2^+ (Krumholz, 2014). The chemical reaction involving an electron is shown in reaction 1.3 and reaction utilizing a proton is shown in reaction 1.4



The main limitation of the free electron/proton formation mechanism is an undersupply of free electrons and protons. Typical Milky Way conditions

show that regions $>1 \text{ cm}^{-3}$ have free electron and proton densities $<10^{-4} \text{ cm}^{-3}$ (Wolfire et al., 2003). Secondly, the H^- ion in reaction 1.3 is more likely to interact with a stray photon than an electron resulting in the ion returning to the atomic state (Glover, 2003). While we observe singly ionized hydrogen regions, HII regions, only a small fraction of the region will successfully produce molecular hydrogen.

While the free electron/proton formation method is suspected to be the primary H_2 catalyst in the early universe (Herbst et al., 2005), at low redshifts the formation of molecular hydrogen on the surface of dust grains is the dominant mechanism of formation (Krumholz, 2014). The surface formation of molecular hydrogen will occur when a hydrogen atom strikes a dust grain and successfully sticks to the grain. The hydrogen then will then interact with another hydrogen atom to form H_2 and be ejected from the dust particle after the reaction has occurred (Pirronello et al., 1997). The same process occurs on the dust grain as the two body formation, but the dust grain will act as a medium to absorb the energy that would create the forbidden photon (Krumholz, 2014).

With a dominant mechanism for molecular hydrogen formation, a reaction rate can be defined based on the cross section of the grain, Σ_{gr} , a striking probability dependent on the temperature of the observed dust, $S(T)$, the probability of molecular hydrogen forming on the grain, ϵ_{H_2} , the density of free hydrogen in the GMC, n_H , and the density of hydrogen attached to the grain surface, $n_{\text{H}0}$ (Krumholz, 2014). Scaling this with the integrated collisional probability of a Maxwellian gas, a reaction rate is determined to be equation

1.5 where k_b is the Boltzmann constant, T is the temperature, and m_H is the mass of the hydrogen atom.

$$\frac{dn_{H_2}}{dt} = \frac{1}{2} \left(\frac{8k_b T}{\pi m_H} \right)^{\frac{1}{2}} \Sigma_{gr} S(T) \epsilon_{H_2} n_H n_{H_0} \quad (1.5)$$

Equation 1.5 is often simplified to equation 1.6 by introducing a variable known as the formation rate, \mathcal{R}_{gr} , that is constrained using the column densities CI, CII, HI, H₂. Typical reaction rates for the Milky Way have been found to be $3 \times 10^{-17} \text{cm}^3 \text{s}^{-1}$ (Jura, 1975; Gry et al., 2002; Wolfire et al., 2008).

$$\frac{dn_{H_2}}{dt} = \mathcal{R}_{gr} n_H n_{H_0} \quad (1.6)$$

1.1.3 Dissociation of Molecular Hydrogen

When the molecular hydrogen has formed, it is still susceptible to photodissociation from far ultra-violet (FUV) photons. The energy required for a single photon to break the bonds of molecular hydrogen is 14.5 eV (Krumholz, 2014). Conveniently, this is also enough energy to excite atomic hydrogen, so photodissociation using a single photon is highly unlikely due to the abundance of HI in the ISM (Krumholz, 2014). However, a photon with an energy of 11-13.6 eV will not be able to excite atomic hydrogen, but will be able to excite the molecular hydrogen to its first and second excitation levels, the Lyman and Werner bands respectively. The excited H₂ will eventually settle via photon emission to its ground state with a finite probability of returning

to two hydrogen atoms rather than maintaining its molecular state (Krumholz, 2014).

A dissociation rate, ζ_{diss} , can be obtained by scaling the total excitation rate, ζ_{exc} , by the fraction of excited hydrogen molecules that will settle to an atomic ground state (Krumholz, 2014). The total excitation rate is found by summing each of individual excitations from the ground state e.g. $\zeta_{exc,0-1}$, $\zeta_{exc,0-2}$. In the Milky Way's diffuse ISM, the interstellar radiation field is $6-9 \times 10^{-14} \text{ erg cm}^{-3}$ over the range of 6-13.6 eV resulting in $\zeta_{exc} \approx 3 \times 10^{-10} \text{ s}^{-1}$ (Draine, 2011). The expected fraction of H_2 to disassociate is between 0.11 and 0.13 resulting in $\zeta_{diss} \approx 4 \times 10^{-11} \text{ s}^{-1}$ (Draine, 2011). Equating the formation and dissociation rates and then solving for a molecular hydrogen vs atomic hydrogen ratio gives in equation 1.7 (Krumholz, 2014). From equation 1.7 we can see, in the diffuse medium, atomic hydrogen is far more abundant than molecular hydrogen, as expected.

$$\begin{aligned} \zeta_{diss} n_{H_2} &= \mathcal{R} n_{H_0} n_H \\ \frac{n_{H_2}}{n_H} &= \frac{\mathcal{R} n_{H_0}}{\zeta_{diss}} \\ &= 8 \times 10^{-6} \left(\frac{4 \times 10^{-11} \text{ s}^{-1}}{\zeta_{diss}} \right) \left(\frac{n_{H_0}}{10 \text{ cm}^{-3}} \right) \end{aligned} \tag{1.7}$$

However, as the density of the gas increases via collapse or coagulation, the optical depth will also increase limiting the amount of FUV photons able to penetrate into the core of the cloud, a process known as shielding (Draine, 2011). This shielding will aid in the decrease of ζ_{diss} which will increase the value in equation 1.7. The increase in molecular to atomic hydrogen ratio signifies the accumulation of molecular gas reservoirs. The molecular gas build

up will lead to fragmentation within the cloud, and will eventually lead to forming stars.

1.2 Determining H₂ Abundance

From the previous section, the formation of molecular hydrogen will result in the energy released during creation to be absorbed by a dust grain rather than through emission resulting in H₂ creation being a dark process, §1.1.2. Furthermore, since the molecule is made of two hydrogen atoms, an extreme symmetry and low mass will negate any permanent dipole effects making low energy rotational transitions nonexistent (Bolatto et al., 2013; Kennicutt & Evans, 2012). The high symmetry and low mass does not mean H₂ is unexcitable, but the temperatures required to excite molecular hydrogen are ($T \gtrsim 100\text{K}$) above the temperature of a typical GMC (Bolatto et al., 2013), and the settling of H₂ will more often than not result in the molecule separating into two hydrogen atoms §1.1.3. For our purposes, we can consider molecular hydrogen a dark molecule requiring special treatment to determine the amount present.

Calculating the amount of molecular hydrogen in a system can be performed in several ways, and for the purpose of extragalactic sources they all involve using the a molecular tracer to determine the amount of H₂ present. The molecule most commonly used in extragalactic studies is CO due to its plentiful amount and ability to be easily observed (Bolatto et al., 2013). Other tracers that have been used are CO's photodissociated counterpart, CII, in order to trace regions that little to no CO emission (Madden et al., 1997), and

molecules such as OH that have been used in the past but have been limited to Milky Way targets (Barrett, 1964). The amount of CO determined can then be converted to H₂ using a conversion factor (either X_{CO} or α_{CO} depending on whether the final solution is in terms of column density or surface density). Typical values for X_{CO} of normal spiral galaxies tend to be around $1-4 \times 10^{20} \text{ cm}^{-2} (\text{K km s}^{-1})^{-1}$, and a commonly used value for the Milky Way is $2 \times 10^{20} \text{ cm}^{-2} (\text{K km s}^{-1})^{-1}$ (Bolatto et al., 2013). Values of α_{CO} for nearby extragalactic sources have shown a mean value of $3.1 \text{ M}_{\odot} \text{ pc}^{-2} (\text{K km s}^{-1})^{-1}$ which is slightly lower than the assumed Milky Way value of $4.4 \text{ M}_{\odot} \text{ pc}^{-2} (\text{K km s}^{-1})^{-1}$ (Sandstrom et al., 2013).

1.2.1 Methods for Determining CO-to-H₂ Conversion Factor

Several methods are present to determine the conversion factor for extragalactic sources, and each have their own caveats. One common method used to determine a conversion factor is to utilize the virial nature of GMCs (Bolatto et al., 2013). This method works best for well defined clouds, however in the case of more distant nearby galaxies, the issue of whether or not GMAs display the same virialization as their constituent GMCs can limit concrete results (Bolatto et al., 2013). The second caveat of using the virial mass to determine a conversion factor is the method will only trace CO bright regions which lead to opposing results obtained using other methods in particular any variations in the conversion factor with the objects metallicity (Bolatto et al., 2013). The mechanism behind this discrepancy is believed to be due to weak dust shielding present in the system and corresponds to a low metallicity. The

poor shielding will allow the CO to be dissociated resulting in large amounts of molecular hydrogen traced by ionized and atomic carbon gas instead of CO (Bolatto et al., 2013). While the virial method is suitable for determining conversion factors within CO bright regions, excluding any low metallicity regions will underestimate the total amount of H_2 present in the system as well as skew any possible relations of the conversion factor with metallicity (Bolatto et al., 2013).

A second method to ascertain a conversion factor is to incorporate observations of isotopologues of CO that are optically thin, commonly ^{13}CO (Bolatto et al., 2013). The temperature, density, and column or surface density of the ^{13}CO can be used to restrict the possible outcomes of the physical conditions of the GMC or GMA being examined. In addition to degeneracies in the radiative transfer models, this method shares the same problem as the virial technique in that it only probes CO bright regions of the target leaving any CO-faint H_2 gas untraced (Bolatto et al., 2013). The issue of CO-faint molecular hydrogen has been examined using the photodissociated tracer of CO, [CII], using the radiative transfer method. The CII ion was used as a tracer of molecular hydrogen in dwarf galaxies as was able to reveal large reservoirs of self-shielding H_2 (Madden et al., 1997).

The third way to determine a CO-to- H_2 conversion factor is by incorporating the emission from dust to determine the amount of molecular gas present. This assumes that the gas is well mixed and has a constant ratio of dust mass to gas mass present in the galaxy (Leroy et al., 2011), which has been shown

to be true for the Milky Way (Boulanger et al., 1996). A suitable conversion factor is found by solving equation 1.8.

$$\begin{aligned}\delta_{GDR}\Sigma_{dust} &= \Sigma_{H_2} + \Sigma_{HI} \\ &= \alpha_{CO}I_{CO} + \Sigma_{HI}\end{aligned}\tag{1.8}$$

Where Σ_{dust} , Σ_{H_2} , and Σ_{HI} are the respective surface densities, I_{CO} is the CO line intensity, α_{CO} is the conversion factor in $M_\odot \text{ pc}^{-2} \text{ K}^{-1} \text{ km}^{-1} \text{ s}$, and δ_{GDR} is the total mass of the gas divided by the total mass of the dust known as the gas-to-dust ratio (Leroy et al., 2011; Sandstrom et al., 2013). In equation 1.8, we can measure Σ_{dust} , Σ_{HI} , and I_{CO} leaving only the conversion factor and gas-to-dust ratio free to vary. An appropriate α_{CO} value will generate a molecular gas mass that produces a constant gas-to-dust ratio over the galaxy being studied, and has been carried out extensively by Sandstrom et al. (2013) on kpc^2 scales on nearby galaxies as well as on both of the Magellenic clouds by Leroy et al. (2011).

Determining a conversion factor in this manner, provides the capability to trace the CO faint regions of target unlike the previous methods (Israel et al., 1996). Despite this advantage over the other two methods, using equation 1.8 leaves any gas not associated with atomic hydrogen to be left as H_2 increasing the conversion factor and overall amount of molecular hydrogen (Bolatto et al., 2013). Another caution of this method rests with the assumption of a constant dust-to-gas ratio (Bolatto et al., 2013). Issues such as the gas-to-dust ratio's dependence with metallicity (Draine et al., 2007) can render this assumption null if not treated properly. Nevertheless, given the agreement of conversion

factors between the dust based method and other methods used for more local targets, suggests that any added bit of gas, incorrectly assumed to be H_2 , and local fluctuations in metallicity of the target galaxy have very little effect on the final results, (Bolatto et al., 2013).

1.3 Determining Dust Mass

Approximating a conversion factor using dust emission requires a knowledge of the amount of dust present in our system to determine a gas-to-dust ratio. Given that a significant portion of the dust mass lies within the cold phase we can calculate a dust mass using a modified blackbody fit over the cold portion of the dust's spectral energy distribution (SED) (Galametz et al., 2012). A slight modification needs to be introduced to our blackbody resulting in what is known as a greybody or modified blackbody (MBB). The modification is necessary given that the dust does not absorb and re-emit all of the light incident on its surface. If we assume an isolated optically thin medium ($\tau_\nu \ll 1$) with a blackbody source function, $B_\nu(T)$, the radiative transfer equation can be written as

$$I_\nu = (1 - e^{-\tau}) B_\nu(T), \quad (1.9)$$

and simplified using a first order Taylor expansion allowed because of the small optical depth, τ_ν , to get

$$I_\nu = \tau_\nu B_\nu(T) \quad (1.10)$$

The optical depth can be expanded to incorporate the surface density of the dust, Σ_{dust} ,

$$\tau_\nu = \kappa_\nu \Sigma_{dust} \quad (1.11)$$

The dust emissivity cross-section per unit mass is represented by κ_ν and is often referred to as the opacity. It is important to note that the opacity will reflect the chemical makeup and grain structure of the dust, but will not indicate the grain size (Abergel et al., 2013). The behavior of the opacity has been well fit with a power law such that opacity shows a dependence with the frequency of the observed emission (Hildebrand, 1983). The opacity is typically written as

$$\kappa_\nu = \kappa_{\nu,0} \left(\frac{\nu}{\nu_0} \right)^\beta \quad (1.12)$$

where ν_0 , is a reference frequency, κ_ν is a reference opacity based on the reference frequency, and β the dust emissivity index. Values for κ_ν have been calculated for several models of varying dust compositions and can range from 0.2 to 2 m² kg⁻¹ over the wavelengths used to model the cold component (100 μ m -850 μ m)(Li & Draine, 2001). The emissivity index has commonly been fit between ranges of 1.0 to 2.0 suggested from laboratory experiments (Walcher et al., 2011). However, with the influx of Herschel data the ability to fill in the sub-millimeter portion of the SED between 100 μ m and 500 μ m has allowed the emissivity index to be determined based on observations and have shown a range from 1 to 2.5 (Galametz et al., 2012). It should be noted that the lower limit of $\beta=1$ is a hard cutoff due to limits determined by the

Kramers-Kronig relationship in the optically thin case (Li, 2005), but in extreme instances when the dust becomes optically thick β will approach 0.

With the opacity well defined, substituting equation 1.12 into equation 1.11 and then into the original equation for the specific intensity of our MBB, equation ??, gives the formula for the specific intensity of a greybody as

$$I_\nu = \Sigma_{dust} \kappa_{\nu,0} \left(\frac{\nu}{\nu_0} \right)^\beta B_\nu(T) \quad (1.13)$$

Since we measure the flux of our target, it is useful to convert from the specific intensity to the flux using

$$F_\nu = \pi I_\nu \left(\frac{r}{D} \right)^2 \quad (1.14)$$

given the target is a uniform sphere with radius, r , and distance, D (?). Substituting equation ?? into the greybody equation, 1.13, gives

$$S_\nu = \frac{\Sigma_{dust} \pi r^2}{D^2} \kappa_\nu \left(\frac{\nu}{\nu_0} \right)^\beta B_\nu(T). \quad (1.15)$$

Converting from surface density to overall mass is done using the relationship between mass and surface density of $M = \Sigma * \text{Area}$ where the targets projected area πr^2 . This gives a final modified blackbody equation of

$$S_\nu = \frac{M_{dust}}{D^2} \kappa_\nu \left(\frac{\nu}{\nu_0} \right)^\beta B_\nu(T). \quad (1.16)$$

The dust mass can be solved from equation 1.16 by either fitting the mass, temperature, and possibly the emissivity index. The fitted mass will act as a

normalizing factor for the SED and correspond with the mass associated at the peak flux. Using the fitted dust mass should be used cautiously due to a sensitivity in mass with temperature fluctuations (Draine et al., 2007). An alternative to using the fitted mass is to isolate the mass from equation 1.16 to get

$$M_{dust} = \frac{S_\nu D^2}{\kappa_\nu B_\nu(T)} \left(\frac{\nu}{\nu_0} \right)^{-\beta} \quad (1.17)$$

and use the parameters returned from fitting. This equation is best used with a longer wavelength observation that will explore the Rayleigh-Jeans tail of the modified blackbody. This portion of the SED will see less of a dependence in the mass and temperature fluctuations (Draine et al., 2007). The main concern with both methods to ascertain the mass arises from fitting both the temperature and dust emissivity index and their strong anti-correlation (Galametz et al., 2012; Tabatabaei et al., 2014). A common method to help break the degeneracy is to fix emissivity index to a reasonable value based upon the opacity model (Tabatabaei et al., 2014).

1.4 NGC 3627

NGC 3627 is the most prominent member of the interacting trio galaxies called the Leo triplet (NGC 3623, NGC 3627, and NGC2628). An optical image is shown in figure 1.1 while the data we used can be seen in sections §2.2 and §2.3. The interaction between NGC 3627 and NGC 3628 has been well documented via HI emission in the form of a large HI debris tail extending

from NGC 3628 in the direction of NGC 3627 (Rots, 1978; Haynes et al., 1979). The interaction between these galaxies has resulted in a high star formation activity, making NGC 3627 a prime candidate to study the current star forming conditions. The morphology of NGC 3627 is that of an SABb with an inclination of 60° (Reuter et al., 1996) at a distance of 9.4 Mpc determined by cepheid variable observations (Kennicutt et al., 2011). Furthermore, the strong spiral arm features are unique to this galaxy in the fact that speculation exists they are not contained within the plain of the galaxy, which has been supported by magnetic field lines (Soida et al., 2001) and CO emission (Dumke et al., 2011).

NGC 3627 is a member of the Spitzer Infrared Nearby Galaxies Survey (SINGS) (Kennicutt et al., 2003) and has been observed throughout the electromagnetic spectrum. ^{13}CO observations by Watanabe et al. (2011) have shown that a majority of the star formation is occurring in the bar ends of the galaxy. This information corresponds with Warren et al. (2010) showing the star formation efficiency being highest in the bar ends. Warren et al. (2010) also shows the galaxy to be molecular gas dominant, with dense warm gas dominating the emission at the bar ends, nucleus and bright region located beneath the southwestern spiral arm, and more diffuse and cooler molecular gas outside of these regions. The atomic gas is located primarily in the spiral arms with little to no emission in the nucleus of the galaxy. A mean conversion factor of $\alpha_{\text{CO}} = 1.2 \text{ M}_\odot \text{ pc}^{-2} (\text{K km s}^{-1})^{-1}$ was found by Sandstrom et al. (2013).

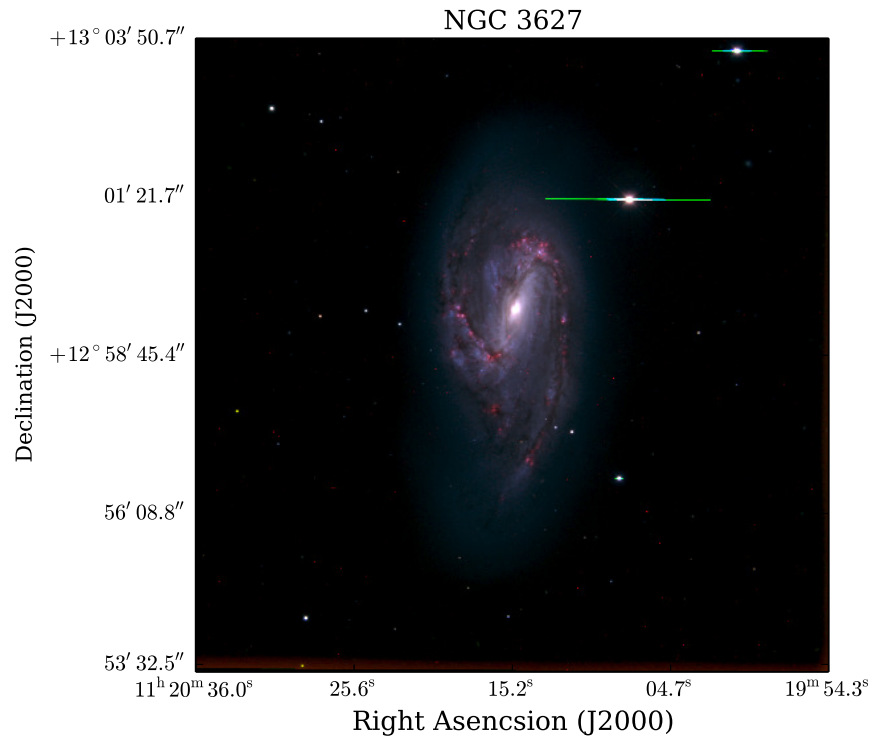


Figure 1.1: Optical composite image of NGC3627 courtesy of the SINGS 5th enhanced data release (Kennicutt et al., 2003). Blue areas represent the B band, green areas represent V band and red areas show the H α band.

The dust emission of NGC 3627 follows the same trend as the CO emission with the brightest regions at the bar ends, nucleus, and a bright region below the southwest bar end. The southeastern spiral arm shows some knotted features in the dust emission. Global values for the cold and warm components of the SED were fit by Galametz et al. (2012) and reveal dust temperatures of $T_W=55.8\pm5.6$ K and $T_C=20.2\pm1.4$ K for the warm and cold components respectively with a best fit emissivity index of $\beta=2.3\pm0.2$ with a total dust mass of $782\times10^5 M_\odot$.

Chapter 2

Observations and Data Preparation

2.1 SCUBA-2

The Submillimetre Common-User Bolometer Array 2 (SCUBA-2) was designed to decrease the observing time of the sub-millimeter sky relative to its predecessor SCUBA (Holland et al., 2013). This would benefit the community by allowing for rapid data acquisition in the submillimeter regime of the electromagnetic spectrum, at the $450\mu\text{m}$ and $850\mu\text{m}$ bands in particular. Prior to SCUBA-2, other bolometer cameras such as LABOCA, BOLOCAM and SHARC-II were limited to less than 100 pixels, while the new SCUBA-2 has been able to incorporate over 10,000 pixels in its design and effectively reduce the required observing time (Chapin et al., 2013). Increasing the amount of pixels by a factor of 100 was possible by the advent of new technology such as high precision micromachining, superconducting transition edge sensors, and superconducting quantum interference device amplifiers (SQUIDs) (Holland et al., 2013).

The observations of NGC3627 were taken from the Nearby Galaxies Legacy Survey’s (NGLS) initial science images using SCUBA-2 from December 29, 2011 to January 21, 2012, and consist of 24 18' by 18' scans taken in grade 3 weather or better ($0.08 < \tau < 0.12$) with observations centered at $450\mu\text{m}$ and $850\mu\text{m}$ emission with a $32\mu\text{m}$ and $85\mu\text{m}$ bandpass respectively. 16 of the 24 scans were deemed useable, and whether or not an observation was deemed worthwhile was determined by factors such as the behavior of the image background or whether the image was flagged during observing to be unusable. The observations of NGC3627 were taken using a daisy scanning pattern to help remove any random noise by introducing crossing points. The scanning speed of the JCMT was $150''/\text{second}$ in order to negate any drifting effects seen from the instrument or sky (Chapin et al., 2013).

2.2 Image Creation and Properties

For any imaging process to be successful, the final image needs to have a limited white noise (Chapin et al., 2013). White noise in the sense of our bolometer observations arises from thermal variations in the instrument and atmosphere during data acquisition. The random noise is minimized through scanning methods and during image processing (Chapin et al., 2013). To create the final SCUBA-2 data products, we use the Submillimetre User Reduction Facility (SMURF) procedure MAKEMAP. This procedure reduces the noise of the observations while maintaining the source’s emission by incorporating a combination of principal component analysis and a maximum likelihood analysis (Chapin et al., 2013). Both of these methods have proven useful in

reducing bolometer data on their own, but due to the size of raw SCUBA-2 data, either method on its own would result in extreme run times or the process becoming resource intensive (Chapin et al., 2013).

MAKEMAP breaks down the image creation into several steps performed in iteration in order to successfully reduce any background noise (Chapin et al., 2013). The raw data is modeled by

$$b_i(t) = f [e_i(t) * a_i(t) + g_i * n_c(t) + n_f(t) + n_r(t)] \quad (2.1)$$

such that $b_i(t)$ is the i^{th} bolometer output at time t , f is a scaling factor determined from flat field calibrations, $e_i(t)$ is the extinction at time t for the i^{th} bolometer, $a_i(t)$ is the astronomical signal for i^{th} bolometer at time t , $g_i * n_c(t)$ is the noise due to the product of the common mode and gain, $n_f(t)$ is the low-frequency noise, and $n_r(t)$ is the random noise (Chapin et al., 2013).

The steps used in MAKEMAP are designed so they can estimate each noise component of the bolometer signal. The COM and GAI steps estimate a common mode signal by fitting the average time-series of each bolometer observation. The next step is EXT, which is used to apply the measured extinction corrections. Following EXT is FLT to apply high- and low-pass filters to remove any noise features not removed in the COM and GAI filtering. After the high- and low-pass filters, the AST step is performed. AST converts the data from a time series to an image and detects sources to be removed from reduction. The sources removed during this portion of MAKEMAP are intended to be astronomical in nature. The final step is NOI which determines the noise in the gridded map after each step has been performed and is calcu-

lated by isolating the white noise component in equation 2.1. A convergence check is then issued based on the magnitude of change in each pixel from the previous map and the current version. If the check fails, the COM, GAI, EXT, and FLT values are recalculated using the AST information obtained from the previous iteration, and the process is repeated until the convergence values are met or the maximum number of iterations have been carried out (Chapin et al., 2013).

In our production of maps, we used the configuration file `dimmmconfig_bright_compact.lis` and altered the AST and FLT sections of the image creation by introducing a mask made from Herschel’s $250\mu\text{m}$ map. The purpose of the mask was to exclude the target from interfering with the noise minimization as well as prohibit any emission from the galaxy to be significantly altered during image production. The filter size of the high-pass filter was also modified, and an appropriate value was determined to be $175''$. We determined an appropriate high-pass filter size by running a large range of filters from $100''$ to $300''$ and inspecting the total recovered flux of NGC3627. A plot of the returned flux values can be seen in Figure 2.1. A good filter would not show any significant decrease in flux compared to the $350\mu\text{m}$ or $500\mu\text{m}$. This requirement removes any filters greater than $200''$. Given the flux was nearly constant with filters less than $200''$, we chose an appropriate filter size by examining the structure that was returned, in particular how well the spiral arms were recovered for the $850\mu\text{m}$, and how well the disk was preserved in the $450\mu\text{m}$. The spatial effect of the filters can be seen in Figure 2.2 for the $450\mu\text{m}$ and Figure 2.3 for the $850\mu\text{m}$. The maps were returned from MAKEMAP in units of pW with a pixel size of $2''$ by $2''$ for both the $450\mu\text{m}$ and $850\mu\text{m}$.

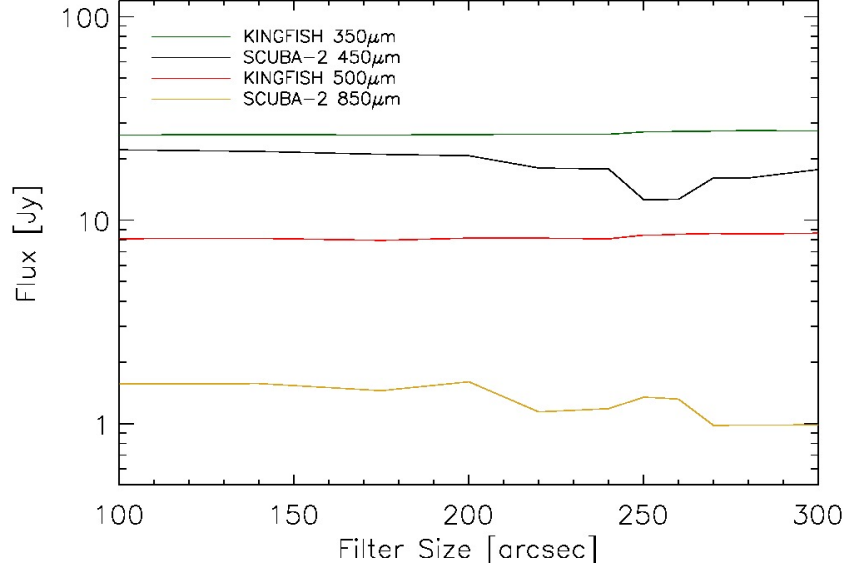


Figure 2.1: Returned flux values for NGC3627 with varying high-pass filter sizes. The KINGFISH fluxes have been processed using the fakesource process, §2.4.2

The finalized $450\mu\text{m}$ image was then re-gridded down to a $4''$ by $4''$ pixel grid, and flux calibration values of 491000 and 4710 were applied to convert from pW to mJy/beam and mJy/square arcsecond, respectively. The $850\mu\text{m}$ maps were re-gridded to an $8''$ by $8''$ pixel size and used flux calibration values of 537000 and 2340 for mJy/beam and mJy/square arcsecond. The $4''$ and $8''$ pixels corresponded to a 180pc and 360pc size scale for our target, NGC3627. To simplify the analysis, the images are also converted to Jy/pixel. The $450\mu\text{m}$ and $850\mu\text{m}$ images are shown in Figures 2.4 and 2.5. The calibration values used are the default flux calibration factors that are determined from our calibrator source, Uranus. The overall noise in the final image can be seen in Table 2.1.

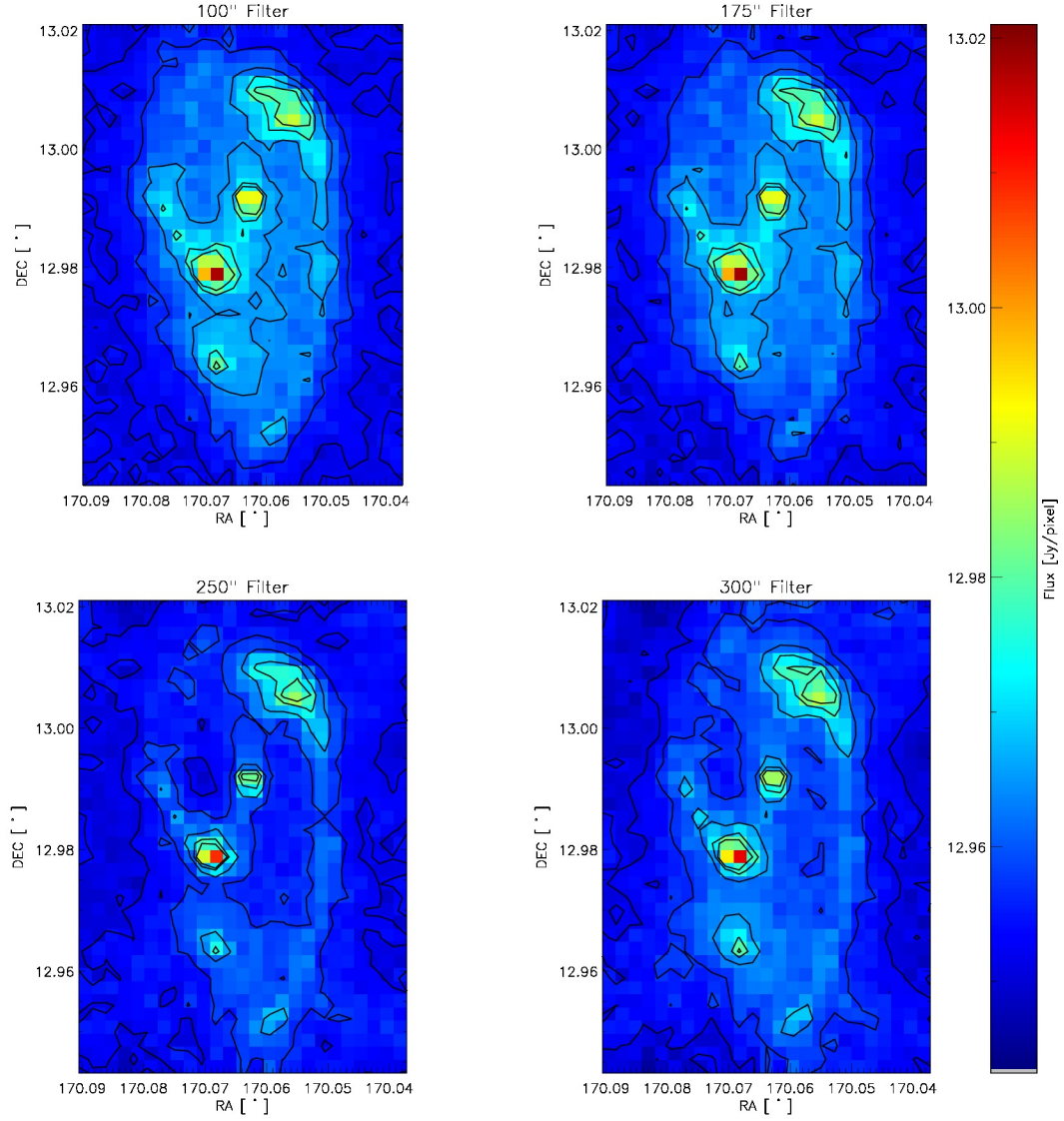


Figure 2.2: Four 450 μm maps of NGC3627 using varying high-pass filter sizes. The contours shown are for 0.0, 0.02, 0.05, 0.08 and 0.1 Jy/pixel for each image.

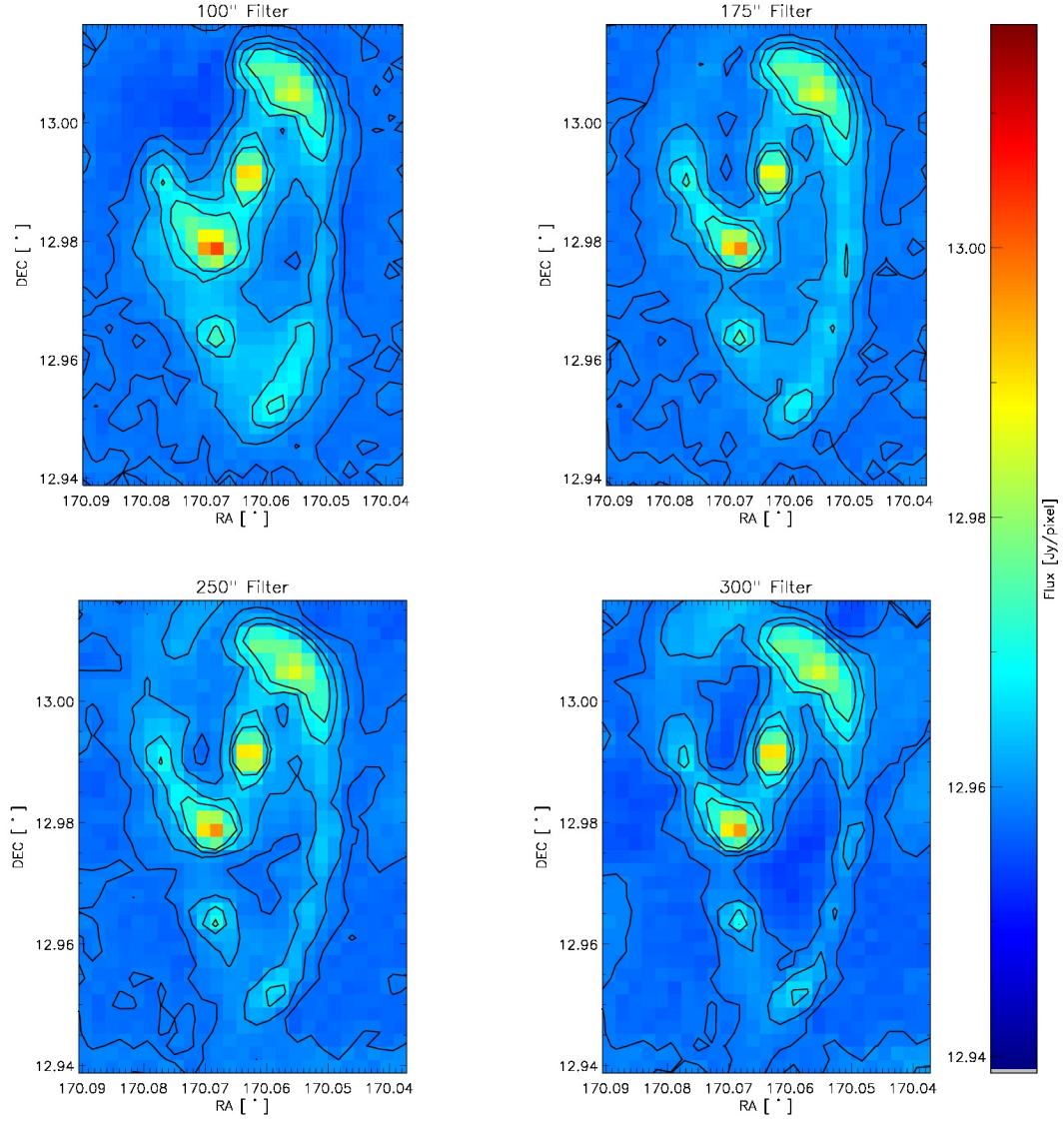


Figure 2.3: Four 850 μm maps of NGC3627 using varying high-pass filter sizes. The contours shown are for 0.0, 0.002, 0.005, and 0.008 Jy/pixel for each image.

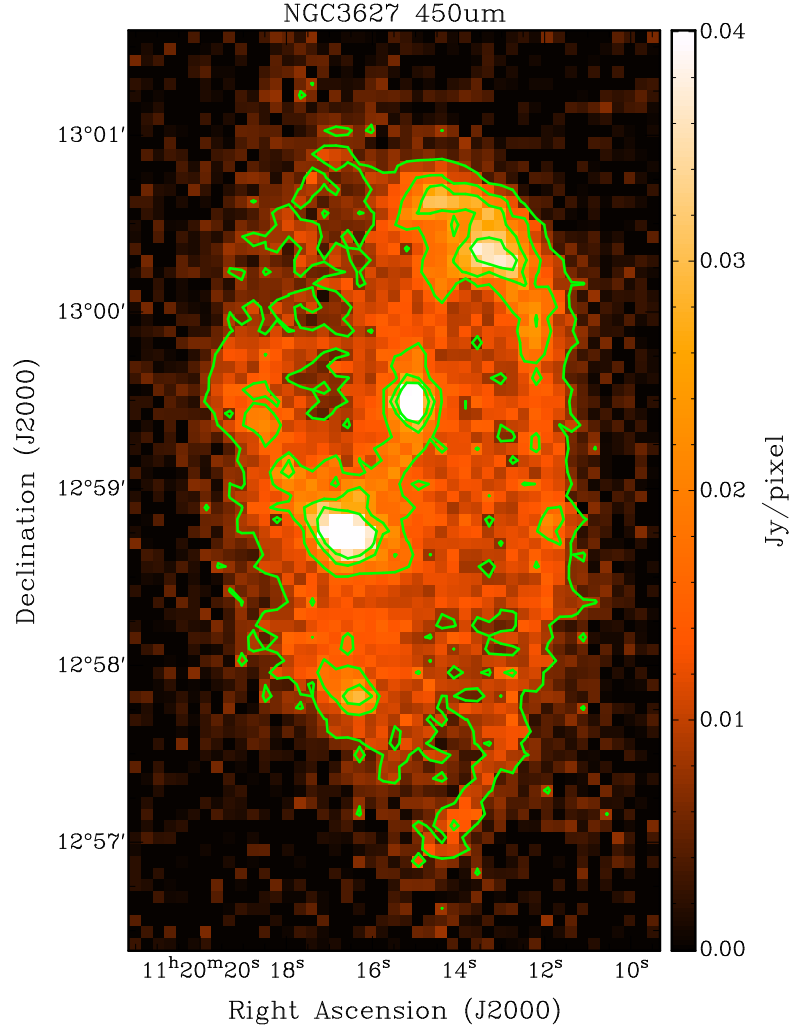


Figure 2.4: 450 μ m observation produced at the end of the image production with 20%, 40%, 60%, and 80% contours.

Table 2.1. Properties of NGC3627 SCUBA-2 Observations

Observation	Beam Properties				RMS [mJy / Pixel]
	α	θ_α	β	θ_β	
450 μ m	0.854 ± 0.002	$7.48'' \pm 0.03''$	0.146 ± 0.003	$23.1'' \pm 0.2''$	3.42
850 μ m	0.9624 ± 0.0002	$12.8'' \pm 0.004''$	0.0376 ± 0.0002	$44.5'' \pm 0.09''$	0.476

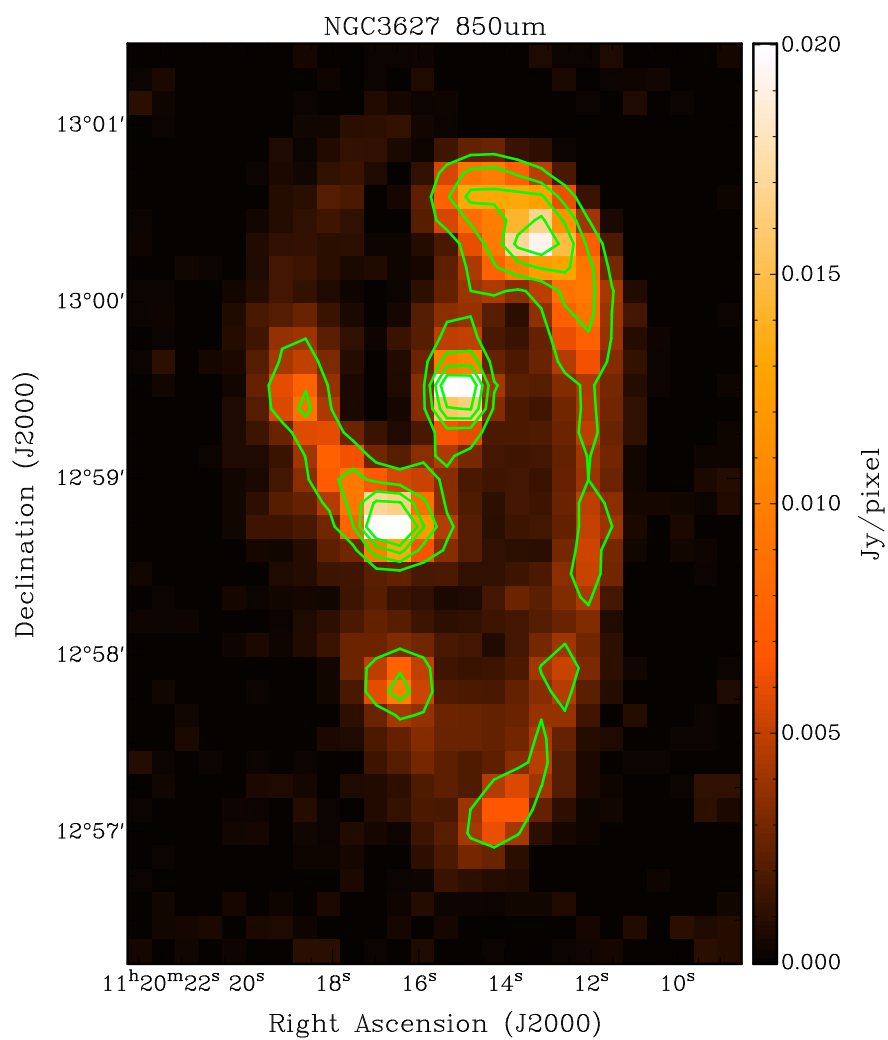


Figure 2.5: 850 μ m observation produced at the end of the image production with 20%, 40%, 60%, and 80% contours.

2.2.1 Beam Shape of the $450\mu\text{m}$ and $850\mu\text{m}$ Data

The Uranus calibration images were used to determine the shape of the beam for the $450\mu\text{m}$ and $850\mu\text{m}$ observations. The beam shape of both the $450\mu\text{m}$ and $850\mu\text{m}$ maps deviates from a single gaussian due to the second maximum of the airy diffraction pattern in the response function of the telescope and minor imperfections in the mirror of the JCMT due to boundaries of the panels (Dempsey et al., 2013). This abnormality is best represented by a sum of two gaussians whose amplitude totals to unity (Dempsey et al., 2013). The average beam resolution for the $450\mu\text{m}$ and $850\mu\text{m}$ are reported in Table 2.1 and match the values within uncertainty found in Dempsey et al. (2013). The calibration images, fitted beams, and the residual of the fits can be seen in Figure 2.6. The contribution of the error beam in the $850\mu\text{m}$ emission is negligible, and allows the beam to be approximated by a single gaussian. However, the contribution of the error beam in the $450\mu\text{m}$ images was large enough to require special treatment in order to properly match the beams for analysis.

2.3 Ancillary Data

The scientific goals of this thesis required data outside the capabilities of SCUBA-2. For instance, accurately determining the dust mass involved fitting the spectral energy distribution (SED) for NGC3627. To successfully fit an SED, we needed shorter wavelength data to fully probe the cold component of this galaxy. We used data ranging from $100\mu\text{m}$ to $500\mu\text{m}$ from the KINGFISH survey (Kennicutt et al., 2011) to gain a large enough wavelength range for fitting the cold component. Secondly, the bandpass of the $850\mu\text{m}$ emission

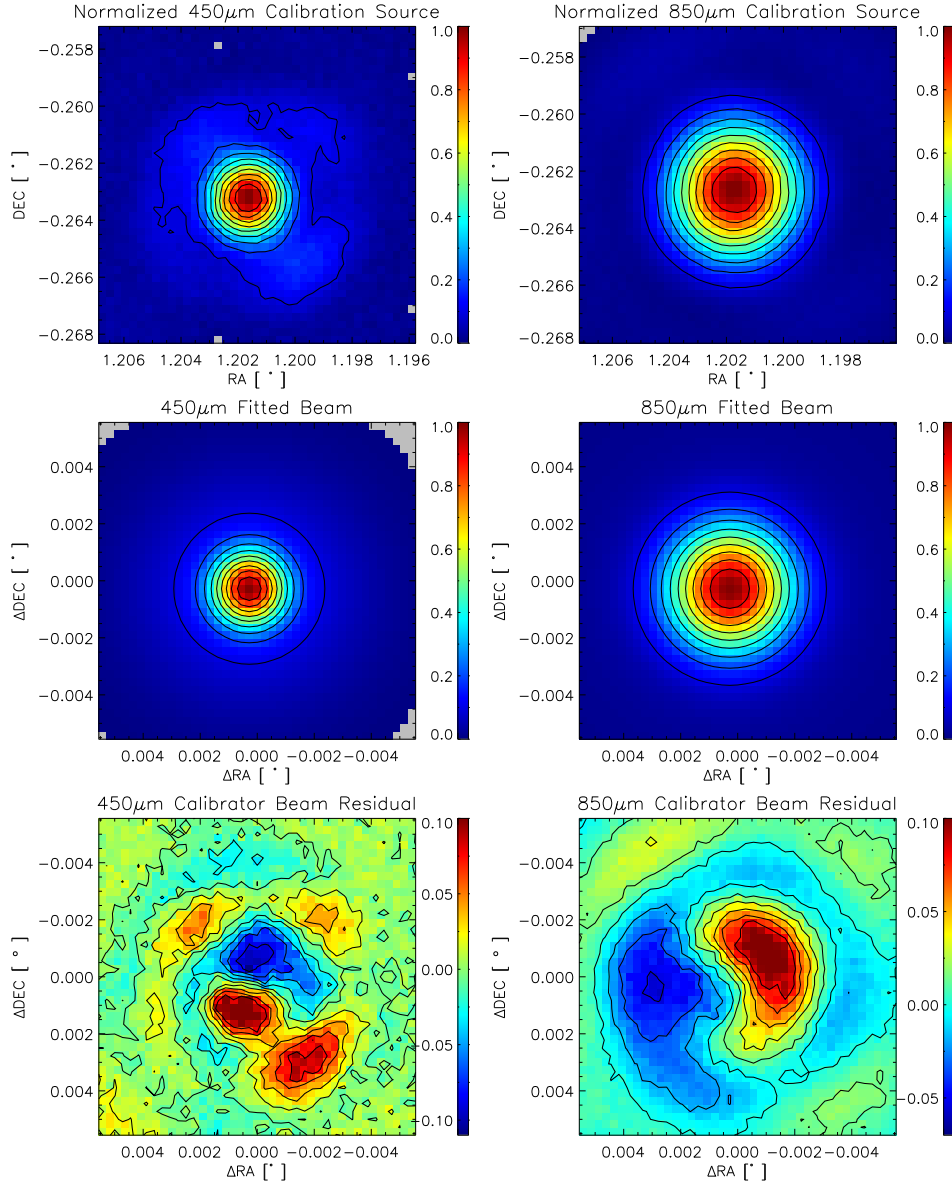


Figure 2.6: The top row shows the Uranus images taken on January, 8th 2012 for the 450 μm on the left and the 850 μm on the right. The middle row shows the fitted beams for the 450 μm on the right and 850 μm on the left using the double gaussian beam shape. The bottom row shows the difference between the original image and the fitted beam. The contours in the image are from 10% to 90% in intervals of 10%.

contains the CO J=3-2 transmission line. In order to get a valid measurement of the dust mass, this contribution had to be removed. We used emission data from the NGLS using the HARP instrument on the JCMT (Wilson et al., 2012) to fulfill this role. When a dust mass was obtained, we used CO J=1-0 from the Nobeyama 45-m telescope (Kuno et al., 2007), CO J=2-1 from HERACLES (Leroy et al., 2009), and *HI* observations from THINGS (Walter et al., 2008) to determine a reasonable molecular hydrogen mass to calculate a dust-to-gas ratio.

2.3.1 Key Insights on Nearby Galaxies: a Far-Infrared Survey with Herschel (KINGFISH)

The Key Insights on Nearby Galaxies: a Far-Infrared Survey with Herschel (KINGSFISH) was designed to be a follow up to the Spitzer Infrared Nearby Galaxies Survey (SINGS) (Kennicutt et al., 2003) with observations of the warm and cold component of dust emission using the increased resolution from Herschel (Kennicutt et al., 2011). The main science goals of the KINGFISH survey were to better understand the star formation processes that were shielded by dust, resolved studies of heating and cooling of the interstellar medium (ISM), and to build an inventory of how cold dust emission relates to other dust components in the ISM (Kennicutt et al., 2011). The survey consisted of studying 61 nearby galaxies ($d < 30 \text{ Mpc}$) that cover a range of environments. Each target was observed at $70 \mu\text{m}$, $100 \mu\text{m}$, $160 \mu\text{m}$, $250 \mu\text{m}$, $350 \mu\text{m}$, and $500 \mu\text{m}$. Our analysis focuses on fitting the cold component of NGC3627's SED, so we omitted the $70 \mu\text{m}$ emission from the fitting, and pro-

Table 2.2. Properties of NGC3627 KINGFISH Observations

Observation	Beam Properties θ_{beam}	RMS $[mJy/Pixel]$	Percentage of Emission Removed
100 μ m	6.8''	2.24	11%
160 μ m	11.6''	3.95	17%
250 μ m	18.0''	2.47	20%
350 μ m	24.9''	1.08	21%
500 μ m	36.0''	0.387	28%

cessed the data through MAKEMAP as described in § 2.4. The rms and beam size after the large scale structure has been removed can be seen in Table 2.2, while the preconvolved maps are shown in Figures 2.7 to 2.11.

2.3.2 Nearby Galaxies Legacy Survey (NGLS)

The Nearby Galaxies Legacy Survey is an HI-selected set of 155 galaxies contained in the annulus of $2Mpc \leq r \leq 25Mpc$ observed using the instrumentation on the JCMT (Wilson et al., 2012). The NGLS consists of data observed in several wavelengths that include the 450 μ m and 850 μ m data used for this thesis. As mentioned previously, the bandpass for SCUBA-2's 850 μ m emission contains the CO J=3-2 line which is contained in the NGLS data set. We used the zeroth moment CO J=3-2 maps from the NGLS to determine the percentage of CO J=3-2 emission present in the 850 μ m band as well as removing it for an accurate SED analysis. The rms and resolution of the CO J=3-2 emission are shown in Table 2.3, and the scan prior to convolution is shown in Figure 2.12.

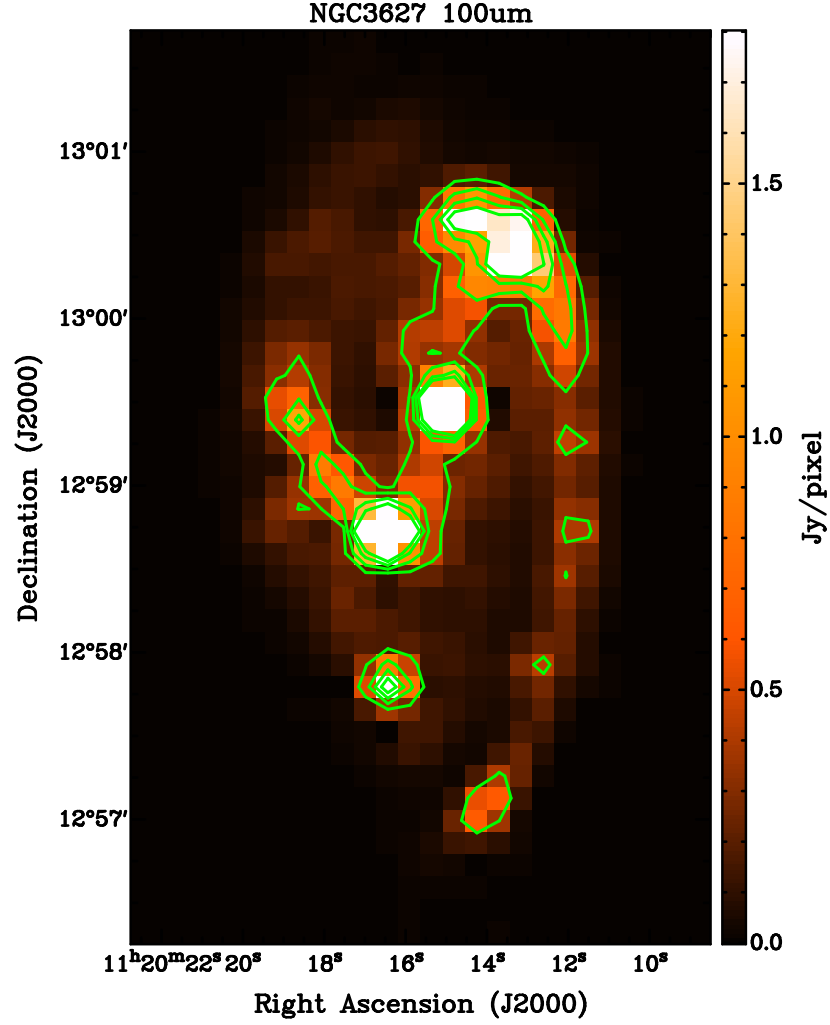


Figure 2.7: Image after the MAKEMAP filtering of $100\mu\text{m}$ observations with 20%, 40%, 60%, and 80% contours.

Table 2.3. Properties of NGC3627 NGLS Observations

Observation	Beam Properties θ_{beam}	RMS $[mJy/Pixel]$	Percentage of Emission Removed
CO J=3-2	14.5''	1.28e-2	29.8%

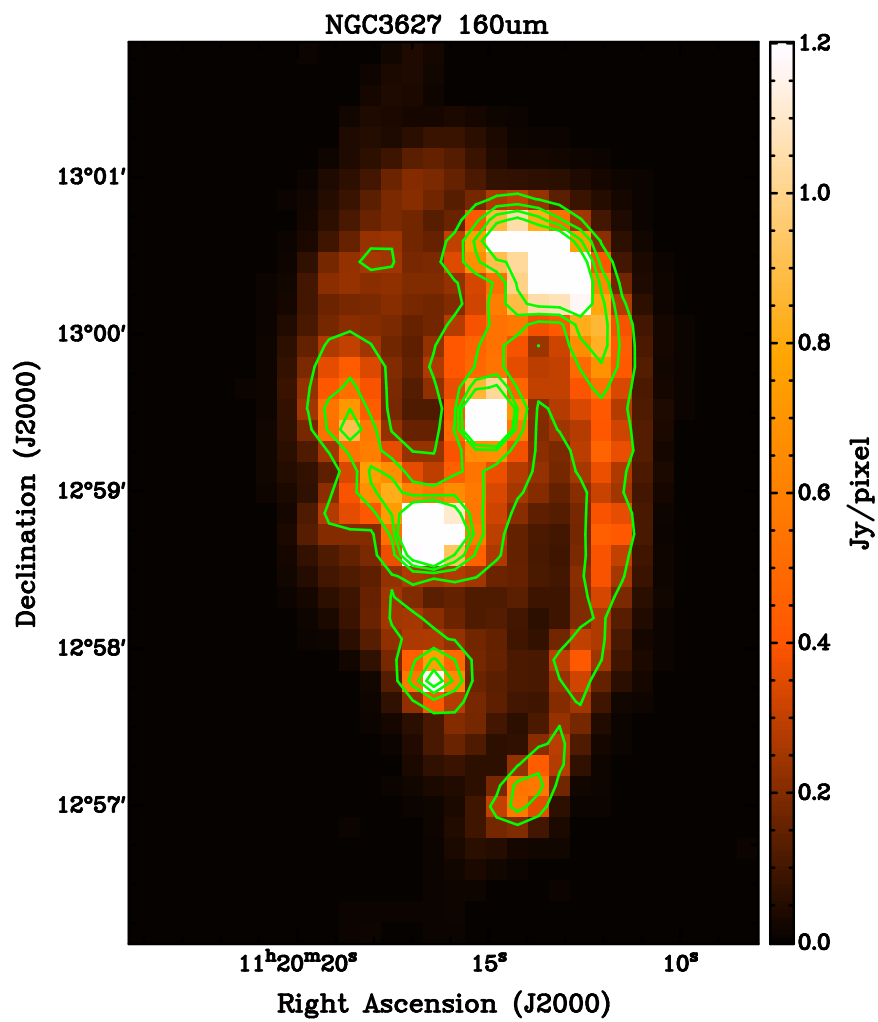


Figure 2.8: Image after the MAKEMAP filtering of $160\mu\text{m}$ observations with 20%, 40%, 60%, and 80% contours.

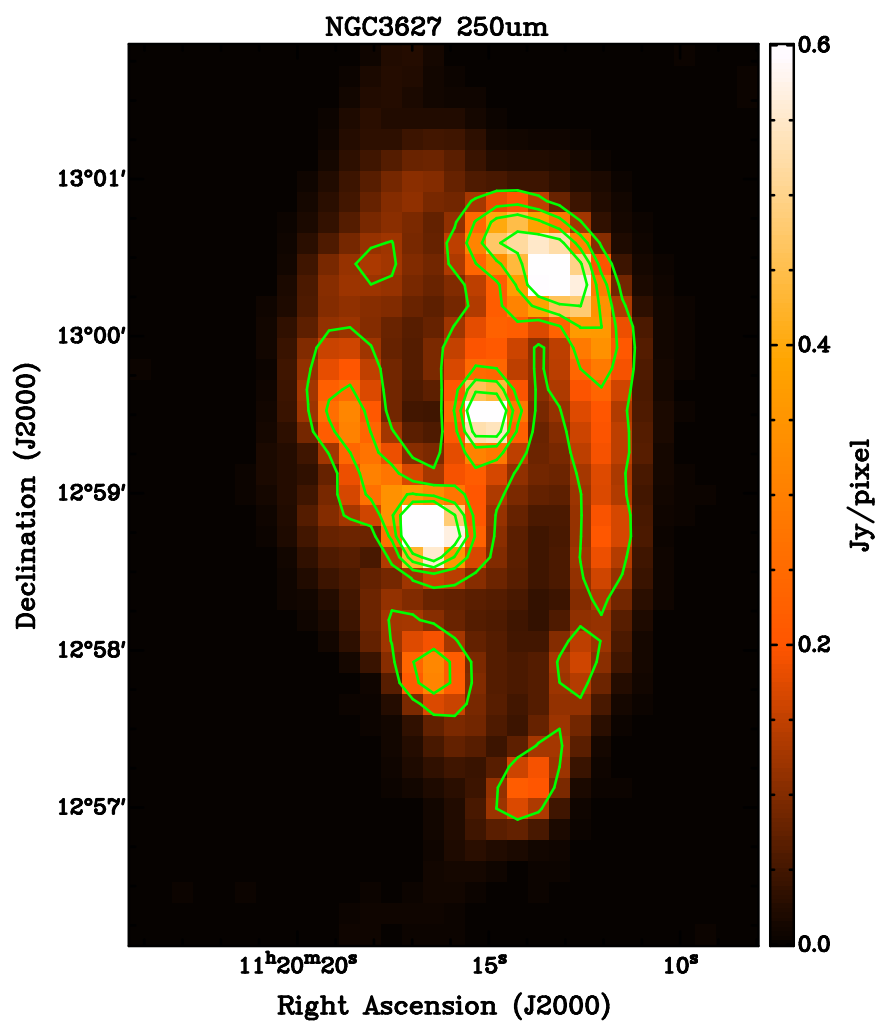


Figure 2.9: Image after the MAKEMAP filtering of 250μm observations with 20%, 40%, 60%, and 80% contours.

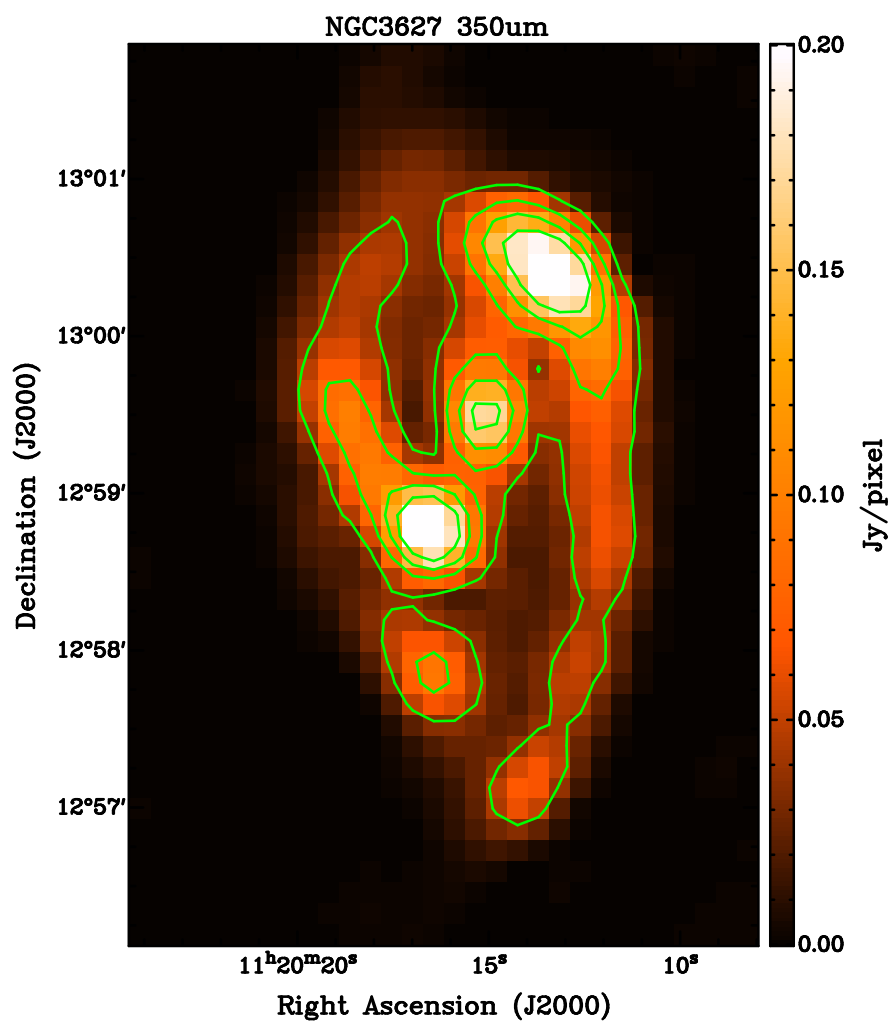


Figure 2.10: Image after the MAKEMAP filtering of the $350\mu\text{m}$ observations with 20%, 40%, 60%, and 80% contours.

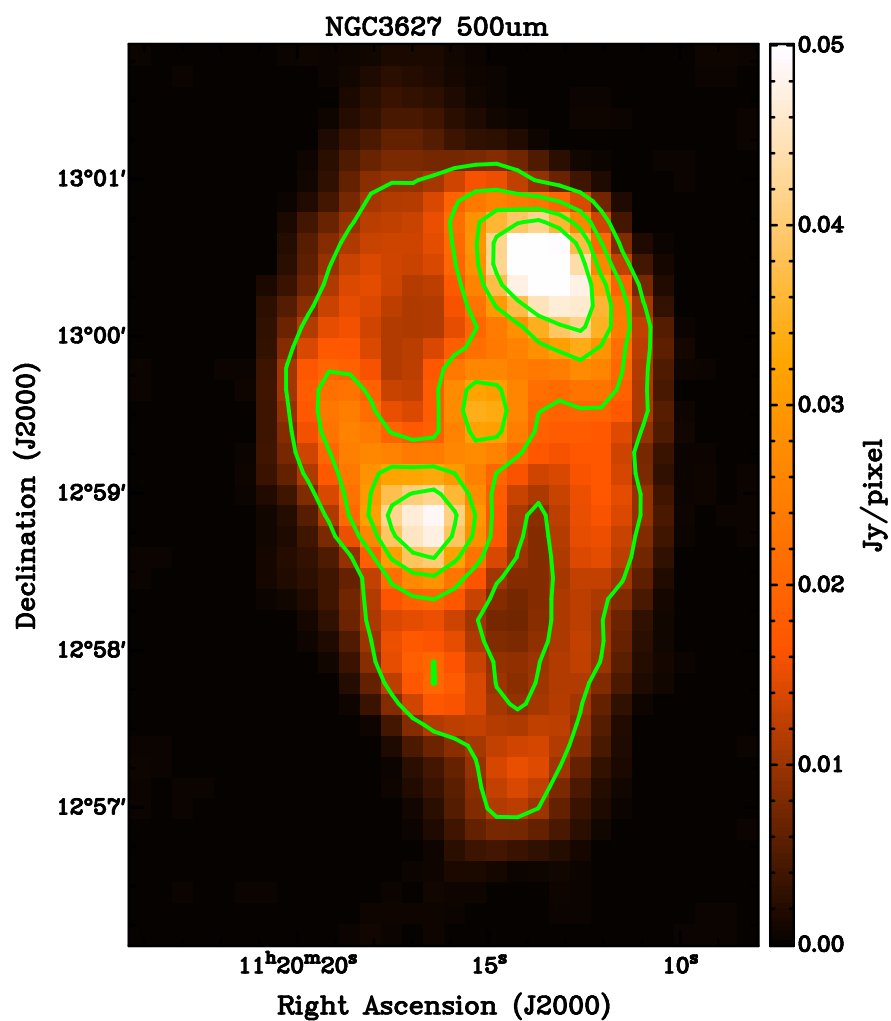


Figure 2.11: Image after the MAKEMAP filtering of the $500\mu\text{m}$ observations with 20%, 40%, 60%, and 80% contours.

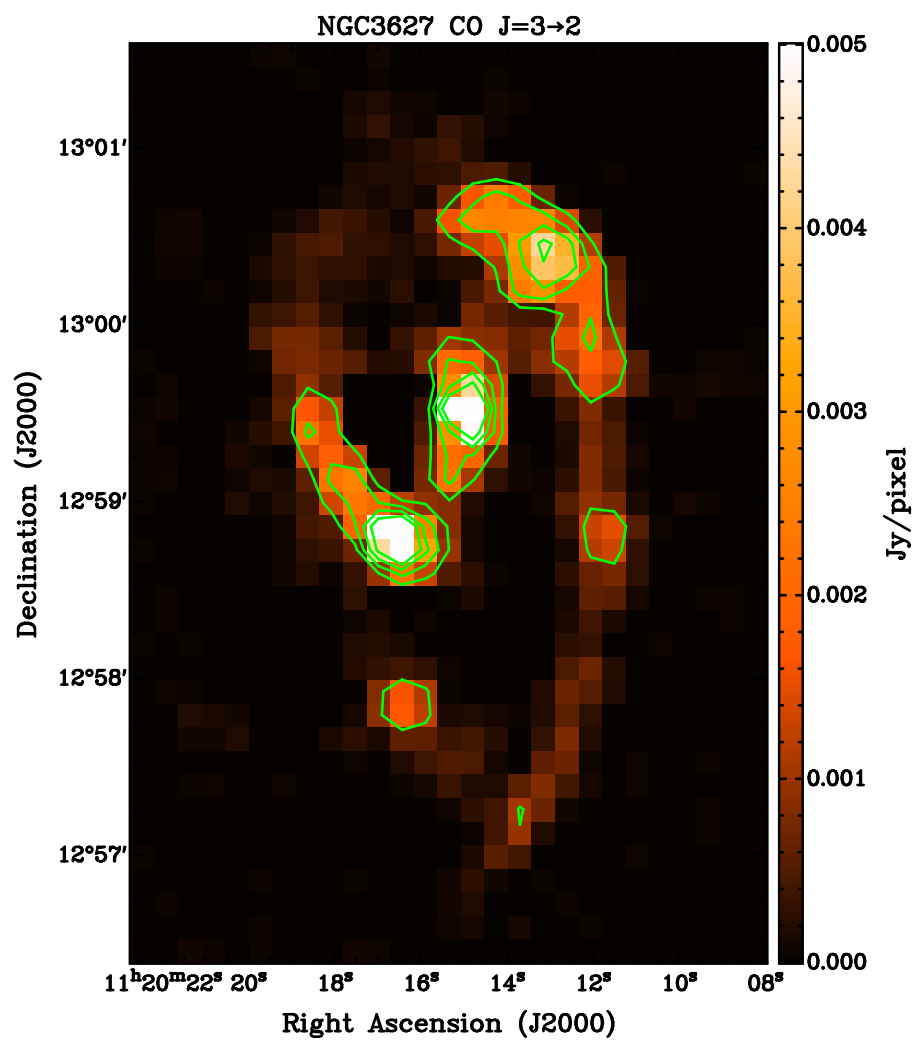


Figure 2.12: Residual of the MAKEMAP filtering of CO J=3-2 observations used to subtract line contamination from 850μm SCUBA-2 map with 20%, 40%, 60%, and 80% contours.

Table 2.4. Properties of NGC3627 Nobeyama 45-m Observations

Observation	Beam Properties θ_{beam}	RMS [K km/s]	Percentage of Emission Removed
CO J=1-0	15.0''	0.681	20%

2.3.3 Nobeyama 45-m

Determining a dust-to-gas ratio requires a molecular tracer to estimate the amount of molecular hydrogen present. The most frequently used tracer is CO J=1-0 due to its abundance in the ISM. The CO J=1-0 we used was taken from the Nobeyama 45-m CO Atlas of Nearby Spiral Galaxies observed to better understand the role of bars relating to molecular gas (Kuno et al., 2007). The Nobeyama 45-m CO Atlas consists of galaxies with morphologies ranging from Sa to Scd, located less than 25Mpc from the Milky Way, inclination values less than 79° , $100\mu\text{m}$ flux greater than 10Jy, and spiral structure that has not been compromised through interactions. Any galaxies that met these criteria were then observed with the Nobeyama 45-m telescope (Kuno et al., 2007). The beam sizes and rms of the filtered CO J=1-0 map are displayed in Table 2.4 and the final image product can be seen in Figure 2.13.

2.3.4 Heterodyne Receiver Array CO-Line Extragalactic Survey (HERACLES)

The CO J=2-1 line was used to determine a CO $2-1/1-0$ line ratio which can be used to trace a gradient in α_{CO} and hint towards regions of

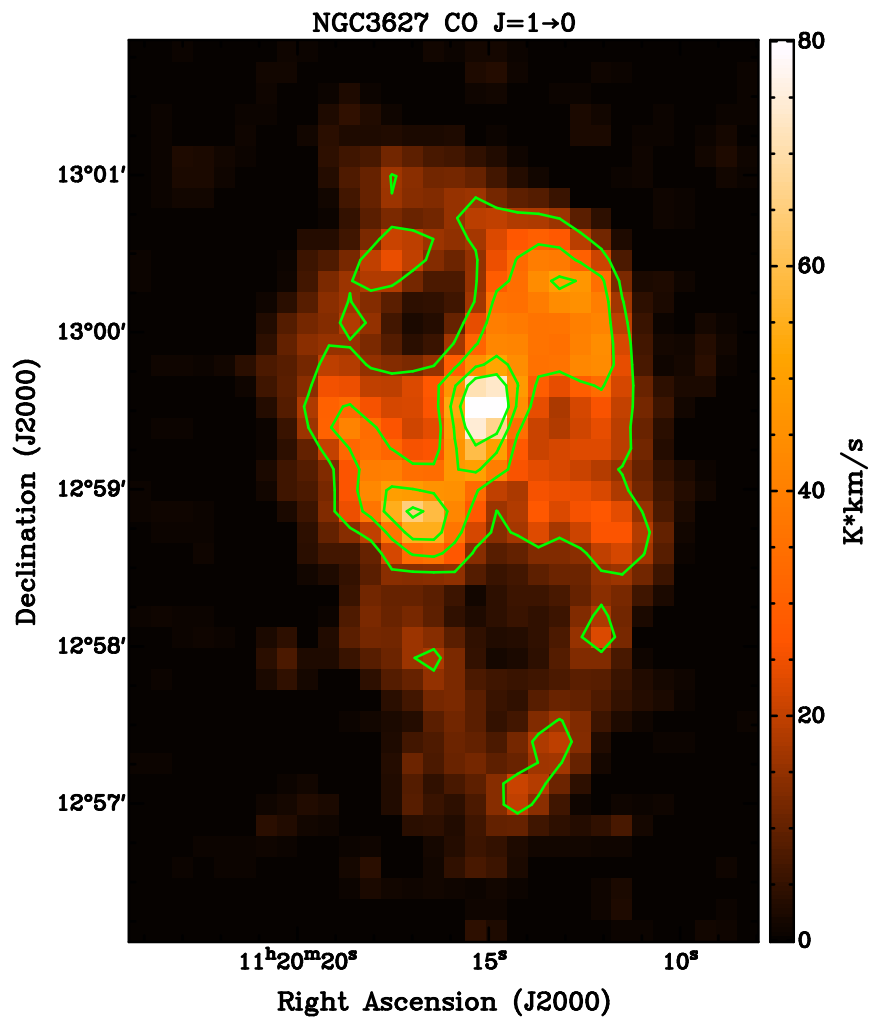


Figure 2.13: Residual of the MAKEMAP filtering of CO J=1-0 observations with 20%, 40%, 60%, and 80% contours.

Table 2.5. Properties of NGC3627 HERACLES Observations

Observation	Beam Properties θ_{beam}	RMS [K km/s]	Percentage of Emission Removed
CO J=2-1	13.0''	0.305	7%

high star-formation (Reuter et al., 1996). We used the CO J=1-0 data from the Nobeyama 45-m telescope (§ 2.3.3), and CO J=2-1 from the Hetrodyne Reciever Array CO-Line Extragalactic Survey (HERACLES) using the IRAM 30-m telescope. The main goal of HERACLES was to quantify the relationship between atomic and molecular gas and star formation using a large sample of galaxies (Leroy et al., 2009). The sample of galaxies chosen were targets contained in THINGS that were within the observing limits of the IRAM 30-m telescope. The final image can be seen in Figure 2.14 and the image properties can be seen in Table 2.5.

2.3.5 The HI Nearby Galaxy Survey (THINGS)

To determine the gas to dust ratio we had to determine the total amount of gas present which includes both atomic and molecular hydrogen. We approximated the amount of molecular hydrogen by using CO J=1-0, and measured the amount of atomic hydrogen (HI) present. The source of our atomic hydrogen came from The HI Nearby Galaxy Survey (THINGS) designed to observe HI emission in nearby galaxies with the extreme spatial resolution of the Very Large Array (VLA). Targets in THINGS included many of the SINGS targets

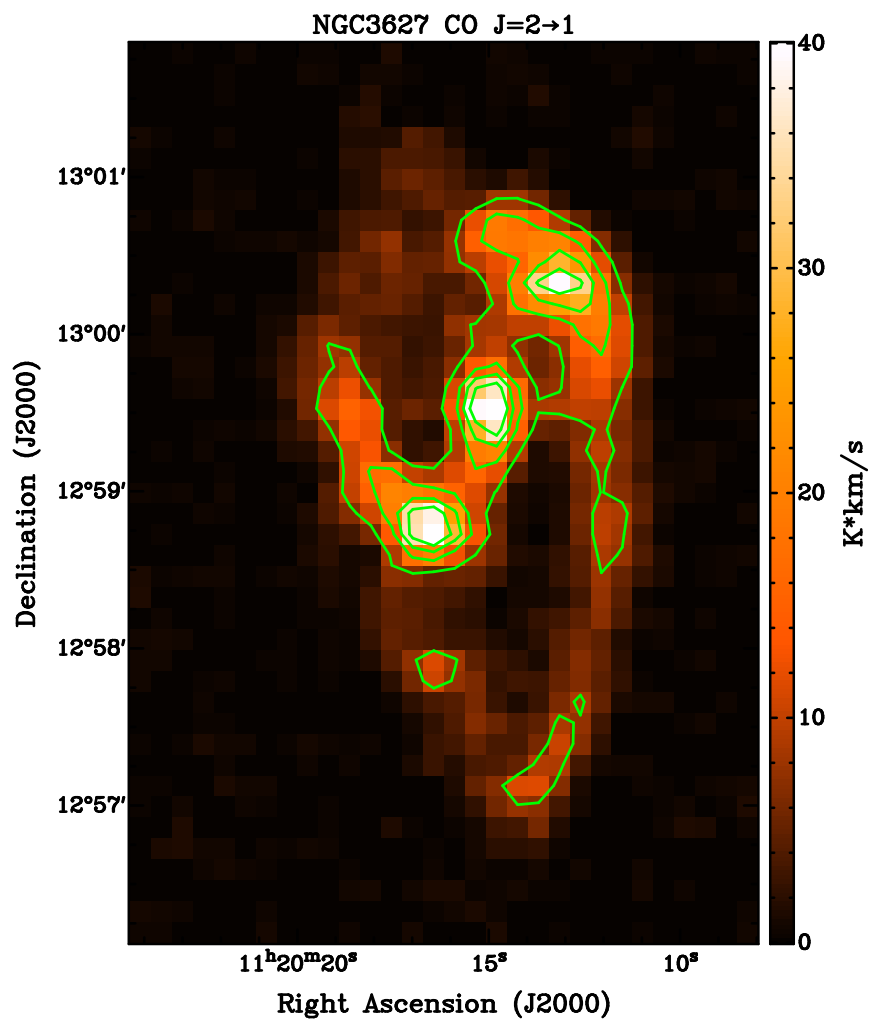


Figure 2.14: Residual of the MAKEMAP filtering of CO J=2-1 observations with 20%, 40%, 60%, and 80% contours.

Table 2.6. Properties of NGC3627 THINGS Observations

Observation	Beam Properties			RMS [M_{\odot}/pc^2]	Percentage of Emission Removed
	θ_{maj}	θ_{min}	θ_{PA}		
HI	10.6''	8.85''	-48.0°	0.760	>99%

with the exception of known HI poor sources (E/S0 type galaxies), dynamically complex systems (edge-on spirals), and large extended galaxies found in the Local Group (Walter et al., 2008). The resolution and rms of the filtered image are shown in Table 2.6. The final data product is shown in Figure 2.15 which has been converted to $M_{\odot} pc^{-2}$ using

$$M_{HI} [M_{\odot}] = 2.36 \times 10^5 D^2 \times \sum_i S_i \Delta v \quad (2.2)$$

Walter et al. (2008) and then dividing by the pixel area in pc^2 where D is the distance, and $\sum_i S_i \Delta v$ is the first moment of the flux.

2.4 Data Preparation for Analysis

Initially the data do not agree on several levels. The differences consist of the presence of the large scale/extended structure and different resolutions including the beam shape of the $450\mu m$. We have taken several steps to correct these disagreements and maximize the compatibility of the data. In order to account for the varying beam resolutions, we use a gaussian convolution to increase the resolution of our maps to the largest beam size in our dataset, $\sigma_{max}=36''$. An appropriate convolution beam width is determined by

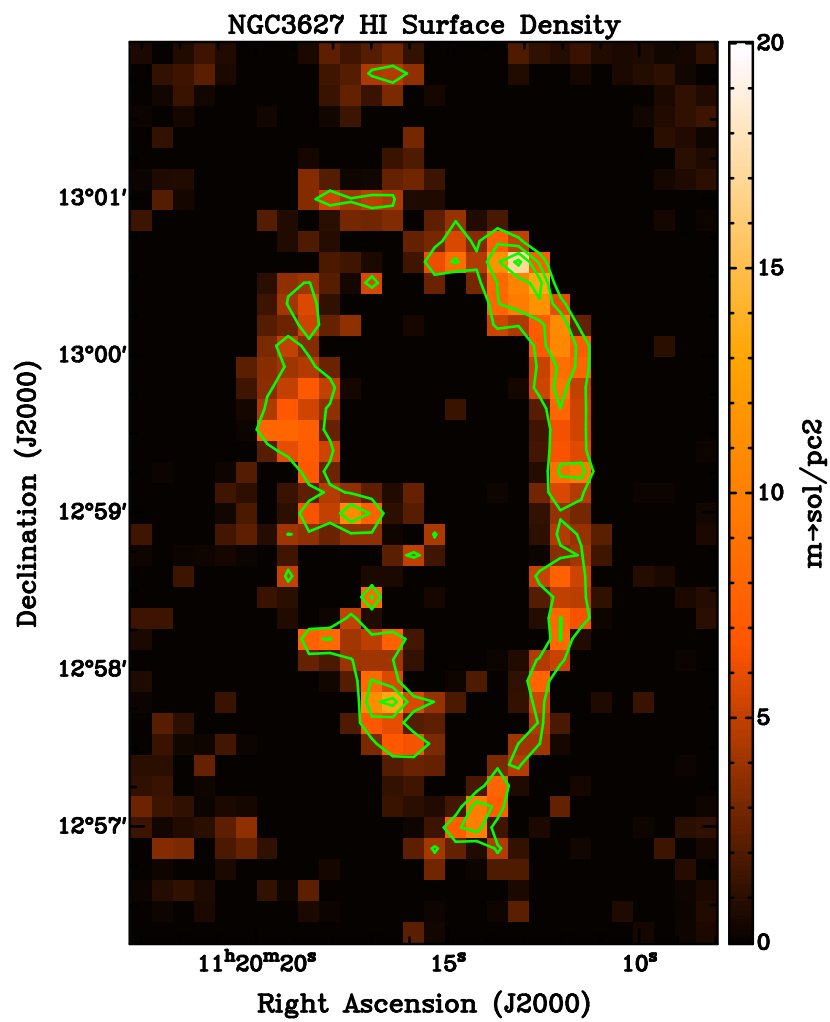


Figure 2.15: Residual of the MAKEMAP filtering of HI observations with 20%, 40%, 60%, and 80% contours.

$$\sigma_{desire} = \sqrt{\sigma_{max}^2 - \sigma_{given}^2} \quad (2.3)$$

such that σ_{desire} is the desired beam width, σ_{max} is the beam size we are convolving to, and σ_{given} is the beam size we are convolving from.

However, this only works when the beams are well approximated by a gaussian which is not the case for the $450\mu\text{m}$ beam. The steps taken to match the resolution of the $450\mu\text{m}$ beam with the rest of the data set are given in §2.4.1. Removing any large scale structure from our ancillary data is implemented using a feature built into MAKEMAP that allows us to add fake sources into the data during production. The fake source implementation allows us to remove the same amount of large scale structure from our ancillary data as was removed from the SCUBA-2 data. The steps taken to prepare the ancillary data are described in §2.4.2.

2.4.1 Accounting for the $450\mu\text{m}$ Error Beam

Taking the $450\mu\text{m}$ error beam into consideration is different than a normal beam convolution in the sense that we are not convolving the higher resolution map to the lowest resolution. Instead we are adding in an error beam similar to the error beam found in the $450\mu\text{m}$ observations. We have to take these steps because convolving a double gaussian beam with a single gaussian kernel will not sufficiently remove the error component of our beam resulting in a poor approximation to the wings of our beam shape.

In order to accommodate the $450\mu\text{m}$ map's error beam, we used a method employed by another SCUBA-2 survey, the Gould Belt Survey team. This method used the distributive nature of the Fourier transform to create similar error components in the beams we were convolving to and from. Carrying out this convolution involved breaking the $450\mu\text{m}$ beam into its two components, X_α and X_β to represent the main beam and error beam. The values of the main beam amplitude, X_α , and error beam amplitude, X_β , will sum to one so the height of the total beam is normalized to one. The poorest resolution beam, X_{max} , is convolved with the two components X_α and X_β and the resulting beams are added together. This process produces an equivalent two component beam as the $450\mu\text{m}$ beam convolved with X_{max} . The relationship can be expressed as

$$\begin{aligned} X_{max} * X_\alpha + X_{max} * X_\beta &= (X_\alpha + X_\beta) * X_{max} \\ &= X_{450\mu\text{m}} * X_{max} \end{aligned} \tag{2.4}$$

Where X_{max} is the poorest resolution beam, X_α and X_β are the main and error beam of the $450\mu\text{m}$ observations, and $X_{450\mu\text{m}}$ is the double gaussian beam shape of the $450\mu\text{m}$ observations.

2.4.2 Extended Structure Removal via MAKEMAP

Due to the combination of methods used in MAKEMAP, large scale/extended structure is removed from the final SCUBA-2 images. However, in all of our ancillary data the large scale emission was present in the initial maps. The removal of the extended features from our support data was carried out by

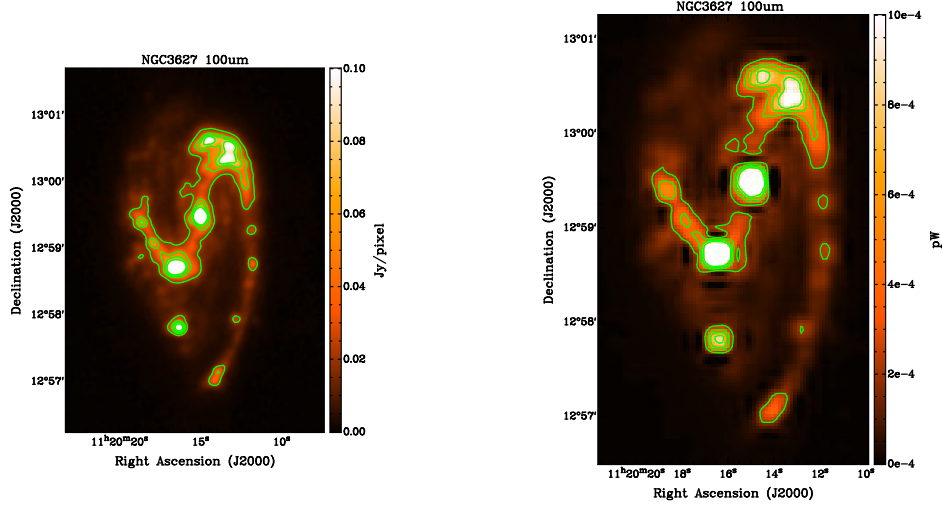
passing the data through MAKEMAP using a special function called fakemap. Implementing fakemap allows us to pass an image through processing and have it added into the image being processed. We use the $850\mu\text{m}$ map as our base image for the filtering process and add the ancillary data to the image. The output image then consists of the sum of the ancillary data and the $850\mu\text{m}$ map. The ancillary data were then isolated by subtracting the $850\mu\text{m}$ image from the fakemap image.

Preparing the data to be implemented consisted of either converting the images from their native units into pW using the $850\mu\text{m}$ flux calibration factor and scaling down to match the observed signal or by just applying a scaling factor. Which method is used is based on the desired units of the final map, and can be separated by what purpose the data had in our analysis.

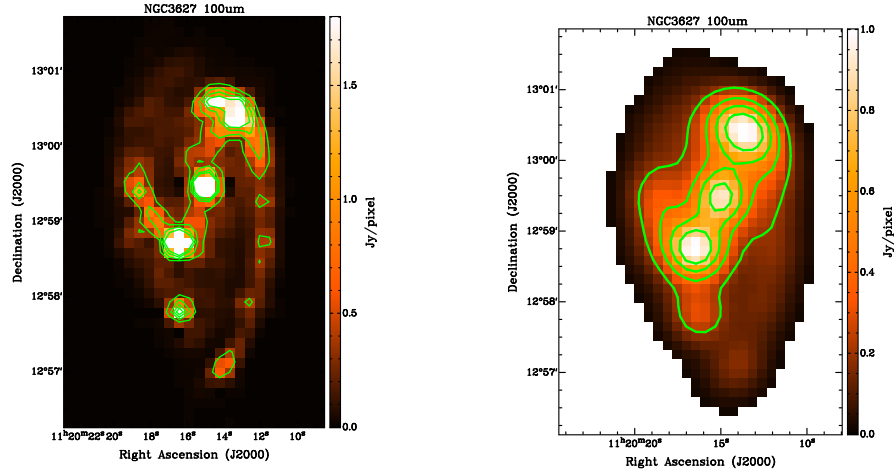
The data used in the SED fitting (KINGFISH and NGLS) were scaled to pW so they could have a similar reduction and calibration process as the SCUBA-2 maps. The KINGFISH data is regridded to a $2''$ by $2''$ pixel size and has the appropriate calibration factor from Dempsey et al. (2013) applied to convert from either MJy/sr to pW in the case of the $250\mu\text{m}$, $350\mu\text{m}$, and $500\mu\text{m}$ or from Jy/pixel to pW for the $100\mu\text{m}$ and $160\mu\text{m}$. Converting the CO J=3-2 requires the final product to be in the same units as the $850\mu\text{m}$ observations in order to properly remove its molecular gas contribution. Converting from K km/s to mJy/beam involved applying a scaling constant of $0.70 [\text{mJy/beam}][\text{K km/s}]^{-1}$ (Drabek et al., 2012) prior to applying the $850\mu\text{m}$ flux calibration factor to convert to pW. After the data has been converted to pW, the image used as a fakemap is then scaled down to the same order of magnitude as the

base image using a scaling factor specified prior to the map production. After the fakemap image has been processed and the $850\mu\text{m}$ map subtracted, the maps are scaled back up using the scaling factor used to scale them down and calibrated using the same flux calibration factors as the $850\mu\text{m}$ map and scaled to an $8''$ by $8''$ grid. The amount of extended flux lost in the KINGFISH and NGLS is shown in Tables 2.2 and 2.3. This process is shown in Figure 2.16 from the original Herschel $100\mu\text{m}$ map, to the final image used in SED fitting convolved to the $500\mu\text{m}$ beam resolution, $36.0''$.

The rest of the ancillary data is used in calculating a dust to gas ratio, and follows nearly the same process as the SED data filtering. The major difference is the CO J=1-0, CO J=2-1 and HI maps need to remain in their original units of K km/s and $\text{M}_{\odot}/\text{pc}^2$. This necessity simplified the process by only requiring a scaling factor of 0.001 to be applied to the original maps. After the data is scaled, it is filtered using the fakesource option in MAKEMAP with the $850\mu\text{m}$ observation as the base image. The atomic and molecular gas maps were then isolated in the same fashion as the KINGFISH and NGLS, however they were rescaled back to their original values and fit to an $8''$ by $8''$ grid then finally convolved to a $36.0''$ resolution. The process is shown in Figure 2.17. The amount of emission lost is shown in tables 2.4, 2.5, and 2.6.

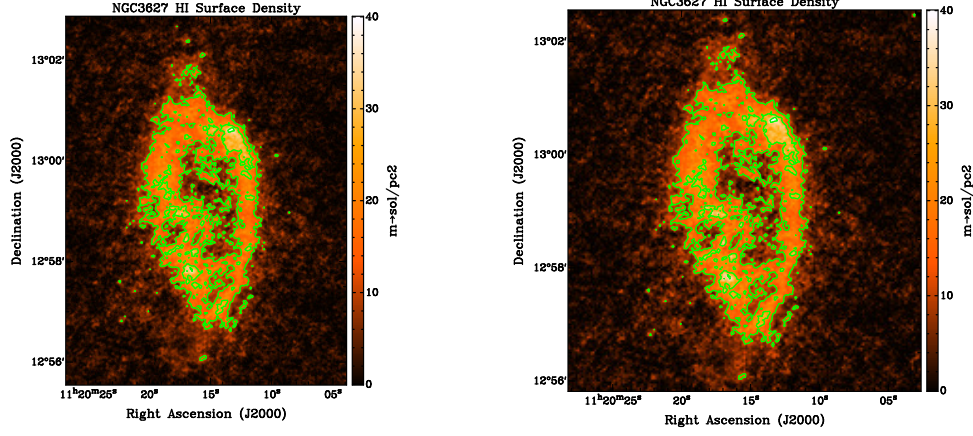


(a) Herschel 100 μ m image of NGC3627. (b) Herschel 100 μ m image converted to pW and rescaled to a 2'' by 2'' pixel grid.

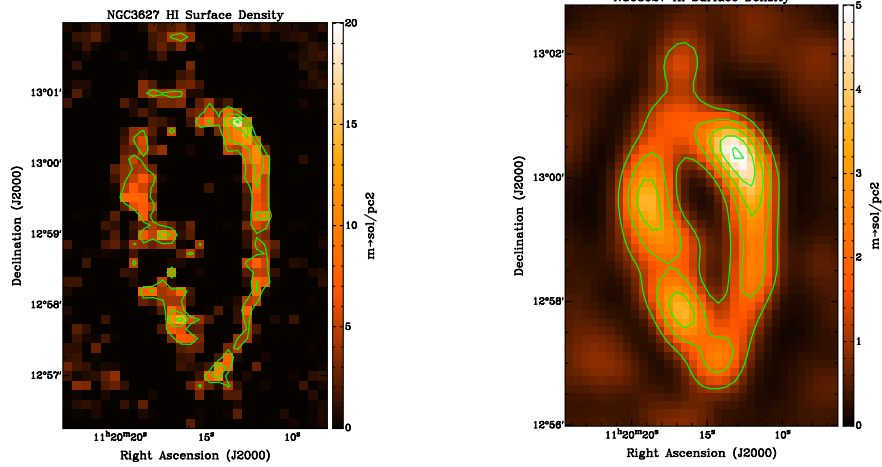


(c) 100 μ m map of NGC3627 after large scale structure has been removed and rescaled to an 8'' by 8'' grid. (d) 100 μ m map with extended structure removed and convolved to final resolution of 36.0'' with a 5 σ cut applied.

Figure 2.16: The 100 μ m image from the beginning of processing to the end of processing. All of the contours shown are 20%, 40%, 60%, and 80%.



(a) HI surface density map of NGC3627. (b) HI surface density map rescaled to a $2''$ by $2''$ pixel grid.



(c) HI surface density map of NGC3627 after large scale structure has been removed and rescaled to an $8''$ by $8''$ grid. (d) HI surface density map with extended structure removed and convolved to final resolution of $36.0''$.

Figure 2.17: The HI surface density maps from the beginning of processing to the end of processing. The contours in the top row are 30%, 60% and 90%, and the contours in the bottom row are 20%, 40%, 60%, and 80%.

Chapter 3

Spectral Energy Distribution Analysis

3.1 SED Fitting Method

In order to determine a dust mass, we use the IDL package MPFIT (Markwardt, 2009) to fit equation

$$S_{\nu}(T) = \frac{M \kappa_{\nu,0}}{D^2} \left(\frac{\nu}{\nu_0} \right)^{\beta+3} B_{\nu}(T) \quad (3.1)$$

for the temperature, T , mass, M , and the emissivity index, β while the opacity, κ_{ν} , distance, D , and reference wavelength, ν_0 are held fixed. The routine MPFIT utilizes the Levenberg-Marquardt algorithm. This algorithm uses a combination of two minimization techniques (the steepest descent method and the Newton-Raphson Method) to determine the parameter combination that corresponds to a minimum in the χ^2 space while maximizing the efficiency of the step sizes in each iteration (Burden & Faires, 2001). The algorithm begins by implementing the steepest descent method. The way this technique works is it will follow the direction opposite of the largest gradient in order to traverse the χ^2 space to locate a minimum. As the set of solutions approaches

Table 3.1. Systematic Calibration Uncertainties for SCUBA-2 and KINGFISH Observations

Observation	Scaling Factor
100 μm	3%
160 μm	5%
250 μm	7%
350 μm	7%
450 μm	12%
500 μm	7%
850 μm	8%

a minimum, it will switch to the Newton-Raphson method to locate the best set of parameters by finding where the derivative at that point is closest to zero (Gavin, 2013).

In order for MPFIT to provide the most accurate fit, we establish a reasonable error for each of our data points, and determine a realistic set of starting points for the fitting. The variance for our SCUBA-2 data is determined using

$$\sigma^2 = \sigma_{obs}^2 + \sigma_{rms,sky}^2 + \sigma_{calib}^2 \quad (3.2)$$

such that σ_{obs} is the noise determined by MAKEMAP, $\sigma_{rms,sky}$ is the RMS of the sky, and σ_{calib} is the product of calibration uncertainty and observed flux for each observed wavelength. The systemic calibration scaling factors are shown in Table 3.1.

The variances for the KINGFISH data are determined in a similar fashion using

$$\sigma^2 = \sigma_{rms,sky}^2 + \sigma_{calib}^2. \quad (3.3)$$

The observation error is excluded since the reported variance in the filtered images are reflective of the fake source image used and not of the KINGFISH data set.

The nature of the Levinberg-Marquardt method leaves the solution vulnerable to converge at a local minimum rather than converging at the global minimum. This is remedied by selecting reasonable initial conditions. The initial conditions we used were a modified blackbody with a temperature of 20K, a dust emissivity index of 2, and a mass determined by

$$\begin{aligned} M &= \frac{D^2 S_{250}}{\kappa_{\nu,0} B_{250}(T)} \left(\frac{\nu}{\nu_0} \right)^{-(\beta+3)} \\ &= 2.24 \times 10^5 S_{250} [M_{\odot}] \end{aligned} \quad (3.4)$$

using the flux from the $250\mu\text{m}$ emission and our initial temperature and dust emissivity index values with a reference opacity of $0.2665 \text{ m}^2 \text{ kg}^{-1}$ at $300\mu\text{m}$. When the best fit values have been calculated we use equation 3.4 to determine the dust mass associated with the $250\mu\text{m}$ emission instead of using the fitted peak mass in order to increase the signal to noise of our data and avoid any temperature dependences associated with emission further from the Rayleigh-Jeans tail of the SED. The error in the dust mass is then determined by taking the total derivative of the mass function with respect to each of the free variables.

3.2 Fitting the Spectral Energy Distribution

The fitting procedure was carried out in two different ways on a modified blackbody equation. One of the two methods is fitting an SED to each individual pixel in order to generate a set of parameter maps. The second method sums the flux of each of the selected regions shown in Figure 3.1 to maximize the signal to noise ratio in order to generate a more precise set of parameters. For both of the fitting methods the mass, M , and temperature, T , are set as free parameters, but the emissivity index, β , requires special treatment for each of the methods. The distance, D , has been set to 9.4 Mpc (Walter et al., 2008), and the reference opacity, $\kappa_{\nu,0}$, was tested using $0.2665 \text{ m}^2 \text{ kg}^{-1}$ (Li & Draine, 2001) and $1.0 \text{ m}^2 \text{ kg}^{-1}$ (Planck Collaboration et al., 2011).

3.2.1 Pixel SED Fits

In order to generate a parameter map of NGC3627, each individual pixel has its own SED determined from the available wavelengths described in the observations chapter. The fits are performed over $100\mu\text{m}$, $160\mu\text{m}$, $250\mu\text{m}$, $350\mu\text{m}$, $450\mu\text{m}$ or $500\mu\text{m}$, and $850\mu\text{m}$ observations. Initially, we excluded the $500\mu\text{m}$ emission in order to increase the resolution of our final maps, but an over abundance of emission in the $450\mu\text{m}$ provided unreliable results, so we exchanged the $450\mu\text{m}$ emission with $500\mu\text{m}$ emission. The treatment of the dust emissivity index is performed by fixing it at $\beta = 1.8$ for the Planck opacity model such that the opacity is $\kappa_{300\mu\text{m},0} = 1.0 \text{ m}^2 \text{ kg}^{-1}$ (Planck Collaboration et al., 2011) and $\beta = 2.0$ for the Li and Draine opacity model where the opacity is $\kappa_{300\mu\text{m},0} = 0.2665 \text{ m}^2 \text{ kg}^{-1}$ (Li & Draine, 2001). A third fit is performed

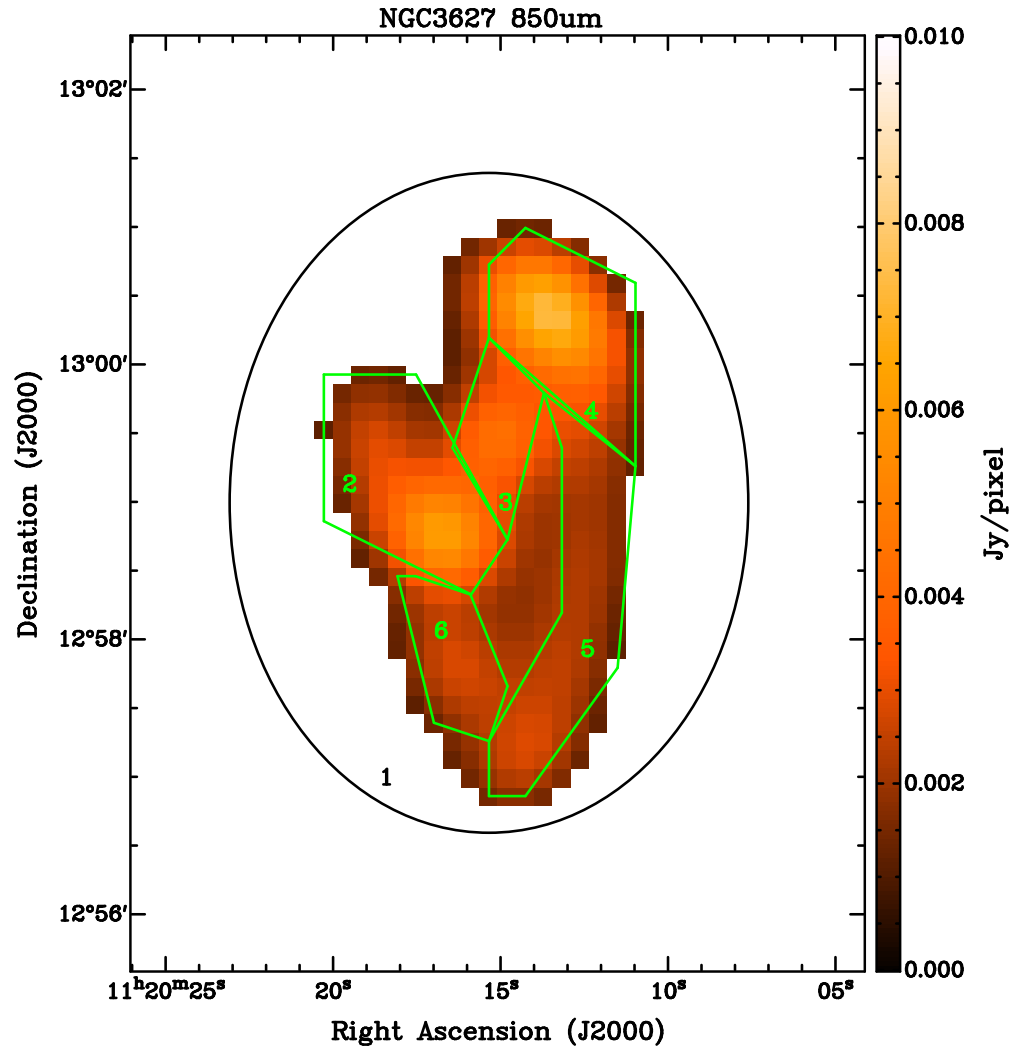


Figure 3.1: 850 μ m emission convolved to the 500 μ m beam size overlaid with the selected regions of NGC3627 labeled 1 through 6 such that region 1 includes the entire galaxy.

where the emissivity index is allowed to vary as a free parameter. While the emissivity index is allowed to vary, the opacity being used is the Planck model. Using either an opacity of $1.0 \text{ m}^2 \text{ kg}^{-1}$ or $0.2665 \text{ m}^2 \text{ kg}^{-1}$ will not effect the shape of the SED, only its normalizing value. In the case of our fits, it is the mass that acts as the normalizing value, so increasing or decreasing the opacity will yield the opposite effect in the mass. The inverse proportionality can be seen in equation 3.4.

The quality of the fit using the $450\mu\text{m}$ observations for the Li and Draine model of $\beta=2$ and $\kappa_{\nu,300}=0.2665$ is shown in Figure 3.2. The quality of the fit is determined by how well the expected flux from the SED fitting matches the observed flux. If the fits are able to recreate the observed emission, then all of the points will lie on the line $y = x$ shown in the plots as a black line. If the SED is underestimating the flux, the points will appear below the 1 to 1 line, and an over estimation from the SED will result in points above the line.

Using Figure 3.2, the fitting procedure is able to handle the $100\mu\text{m}$, $250\mu\text{m}$, and $850\mu\text{m}$ with some slight under estimation in the $160\mu\text{m}$ and over estimation in the $350\mu\text{m}$. The most compelling piece of information in Figure 3.2 is the SED's inability to recreate the observed $450\mu\text{m}$ flux despite the care taken while accounting for the odd beam shape of the $450\mu\text{m}$ data. The discrepancy between the observed and fitted fluxes was not a result of a poor fit, but rather an intrinsic error in the $450\mu\text{m}$ observations. The over abundance of flux present in SCUBA-2 $450\mu\text{m}$ data was determined to be an issue with the data itself rather than a signature of a physical process. Determining if the $450\mu\text{m}$ data was overestimating the flux was tested by extrapolating a

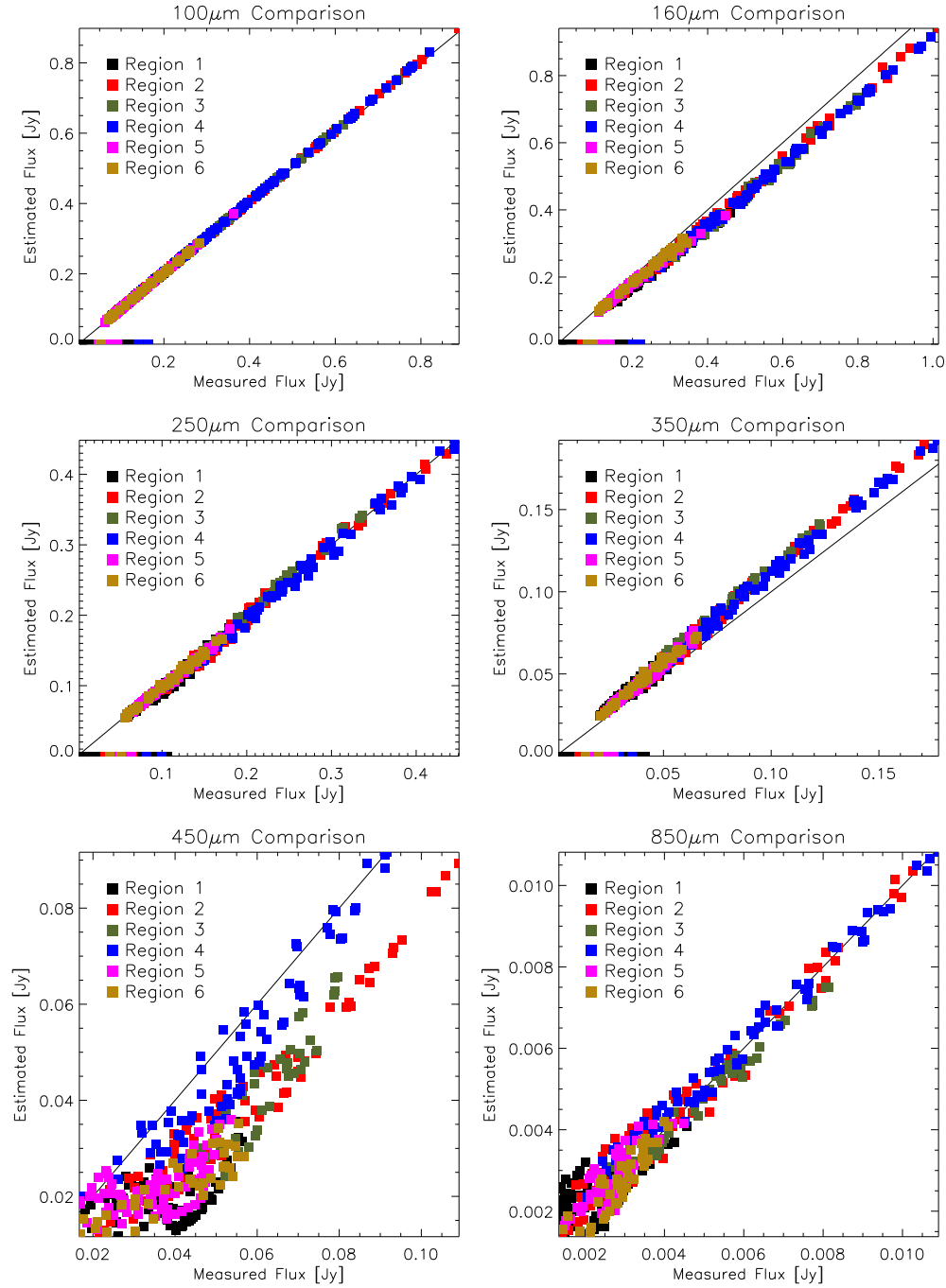


Figure 3.2: Quality of the SED fits to the Li and Draine model using the 450 μ m emission. The regions shown are the regions in figure 3.1.

Table 3.2. Total Distance to 1 to 1 Line Using the $500\mu\text{m}$ Emission

Observation	Planck Model	Li and Draine Model	Variable Emissivity Index
$100\mu\text{m}$	0.000	0.01245	0.1277
$160\mu\text{m}$	15.55	8.979	2.159
$250\mu\text{m}$	5.808	3.071	0.4465
$350\mu\text{m}$	0.8045	0.3541	0.1161
$500\mu\text{m}$	0.08192	0.1147	0.1866
$850\mu\text{m}$	0.02825	0.05218	0.09528
Total	22.23	12.583	3.131

$450\mu\text{m}$ flux using a linear fit between the filtered $350\mu\text{m}$ flux of 27.1 Jy and the filtered flux of the $500\mu\text{m}$ flux of 8.56 Jy. The flux expected from these two data points was 14.7 Jy, however our $450\mu\text{m}$ flux measured to be 20.6 Jy.

In order to avoid any errors in our final parameter maps, we substitute the $450\mu\text{m}$ emission with KINGFISH $500\mu\text{m}$ emission. The quality of fitted SEDs using the $500\mu\text{m}$ emission are shown in Figures 3.4, 3.5, and 3.6 for the Planck model, Li and Draine model and the emissivity as a free parameter respectively. To assign a numerical quantity to the quality of the fit the vertical distance was calculated for each point to the 1 to 1 line and then summed. The summed distances are shown in Table 3.5 and suggest the best fit comes from allowing the emissivity index to vary. The resulting parameter maps of the fits using the $500\mu\text{m}$ observations are shown in Figure 3.3. The numerical values from fitting each model are shown in Tables 3.2, 3.3, and 3.4 where each region corresponds to the regions labeled in Figure 3.1.

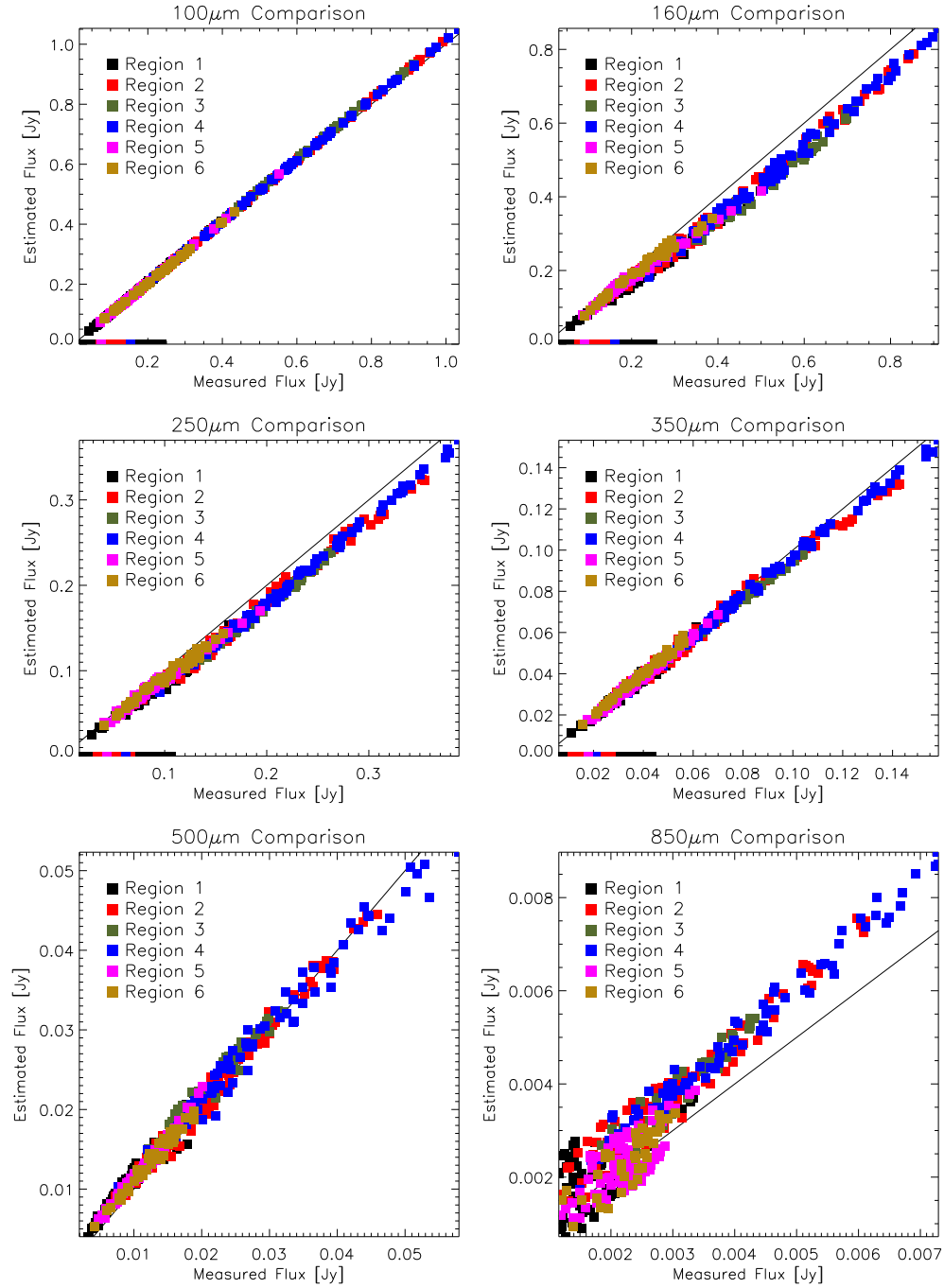


Figure 3.3: Quality of the SED fits using to the Planck model with the 500 μ m emission. The regions shown are the regions in figure 3.1.

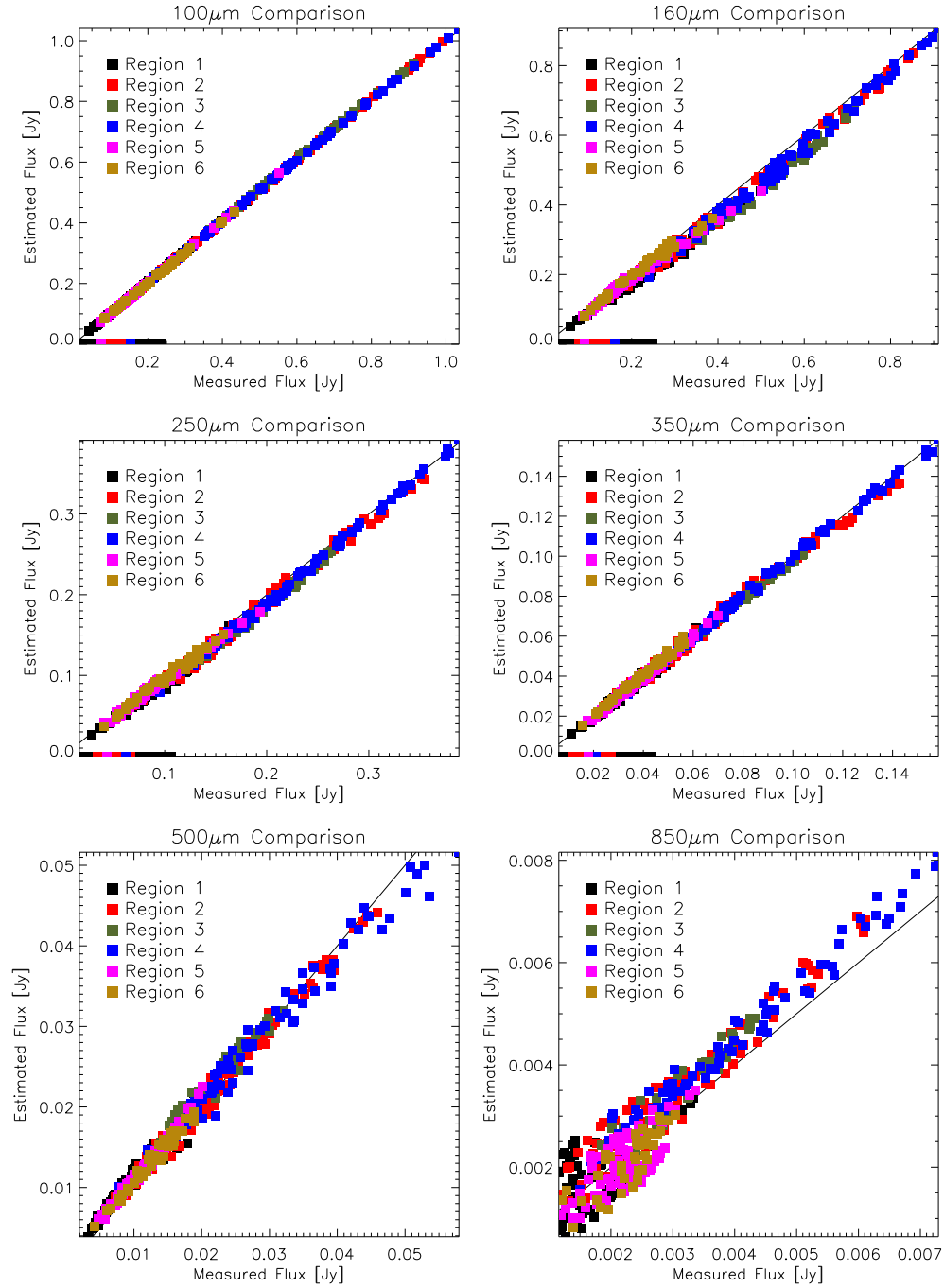


Figure 3.4: Quality of the SED fits using the Li and Draine model with the 500 μ m emission. The regions shown are the regions in figure 3.1.

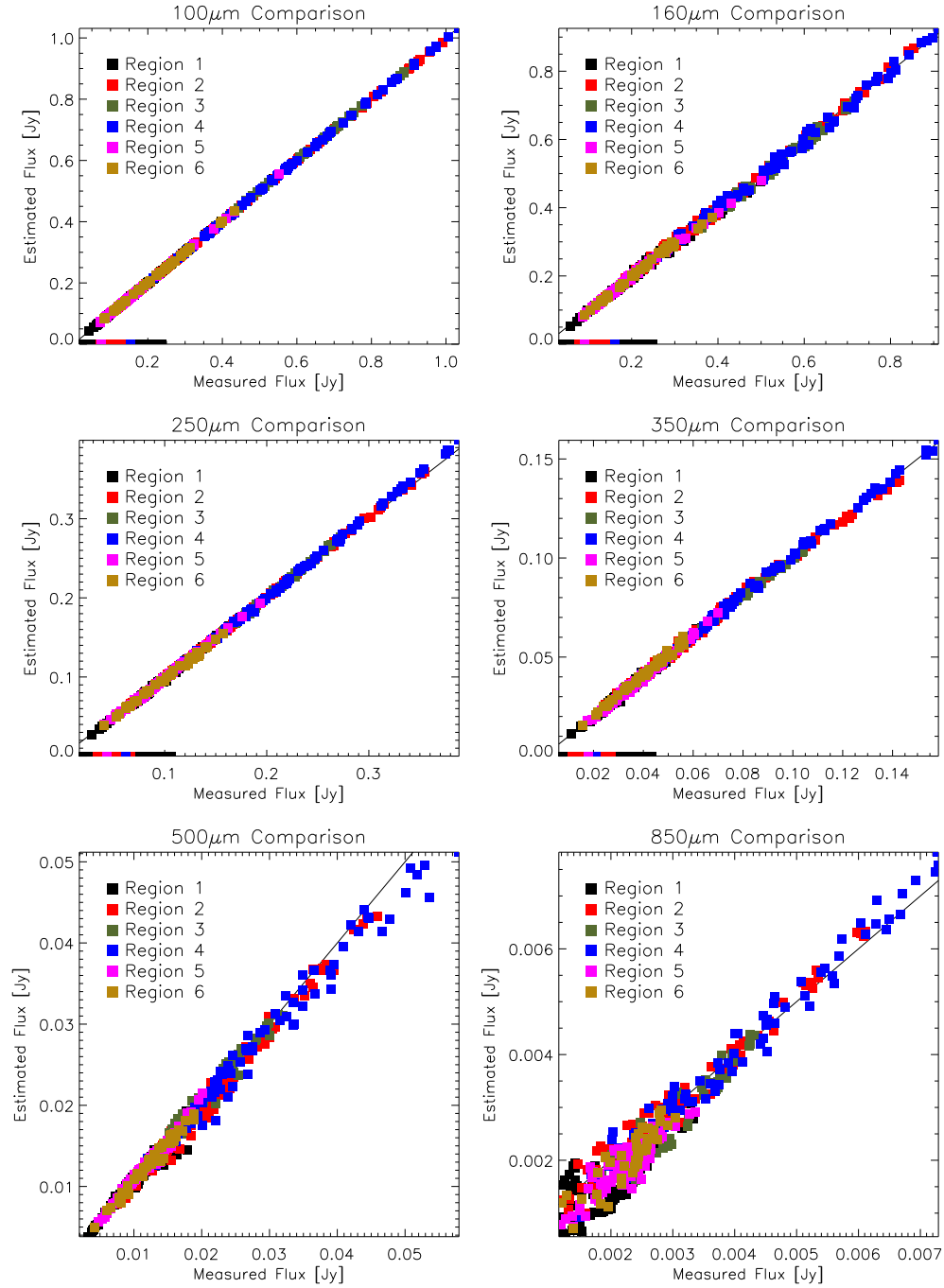


Figure 3.5: Quality of the SED fits with the emissivity index as a free parameter with the 500 μ m emission. The regions shown are the regions in figure 3.1.

Table 3.3. Best Fit Parameters for Planck Model Using $500\mu\text{m}$ Emission for Pixel Fits

Region	Average β	Total Mass [$10^5 M_\odot$]	Surface Density [$M_\odot \text{pc}^{-2}$]	Average Temperature [K]
1	1.8	66 ± 30	0.13 ± 0.06	26 ± 2
2	1.8	16 ± 6	0.16 ± 0.05	27 ± 1
3	1.8	8 ± 1	0.15 ± 0.03	28.8 ± 0.6
4	1.8	20 ± 6	0.19 ± 0.06	26 ± 1
5	1.8	12 ± 3	0.10 ± 0.02	24 ± 1
6	1.8	5 ± 1	0.09 ± 0.02	25 ± 1

Table 3.4. Best Fit Parameters for Li and Draine Model Using $500\mu\text{m}$ Emission for Pixel Fits

Region	Average β	Total Mass [$10^5 M_\odot$]	Surface Density [$M_\odot \text{pc}^{-2}$]	Average Temperature [K]
1	2.0	278 ± 128	0.5 ± 0.3	24 ± 2
2	2.0	70 ± 23	0.7 ± 0.2	25 ± 1
3	2.0	35 ± 6	0.7 ± 0.2	26.4 ± 0.5
4	2.0	84 ± 27	0.8 ± 0.3	24.4 ± 0.8
5	2.0	49 ± 11	0.40 ± 0.09	22.4 ± 0.8
6	2.0	21 ± 5	0.40 ± 0.10	23.5 ± 0.9

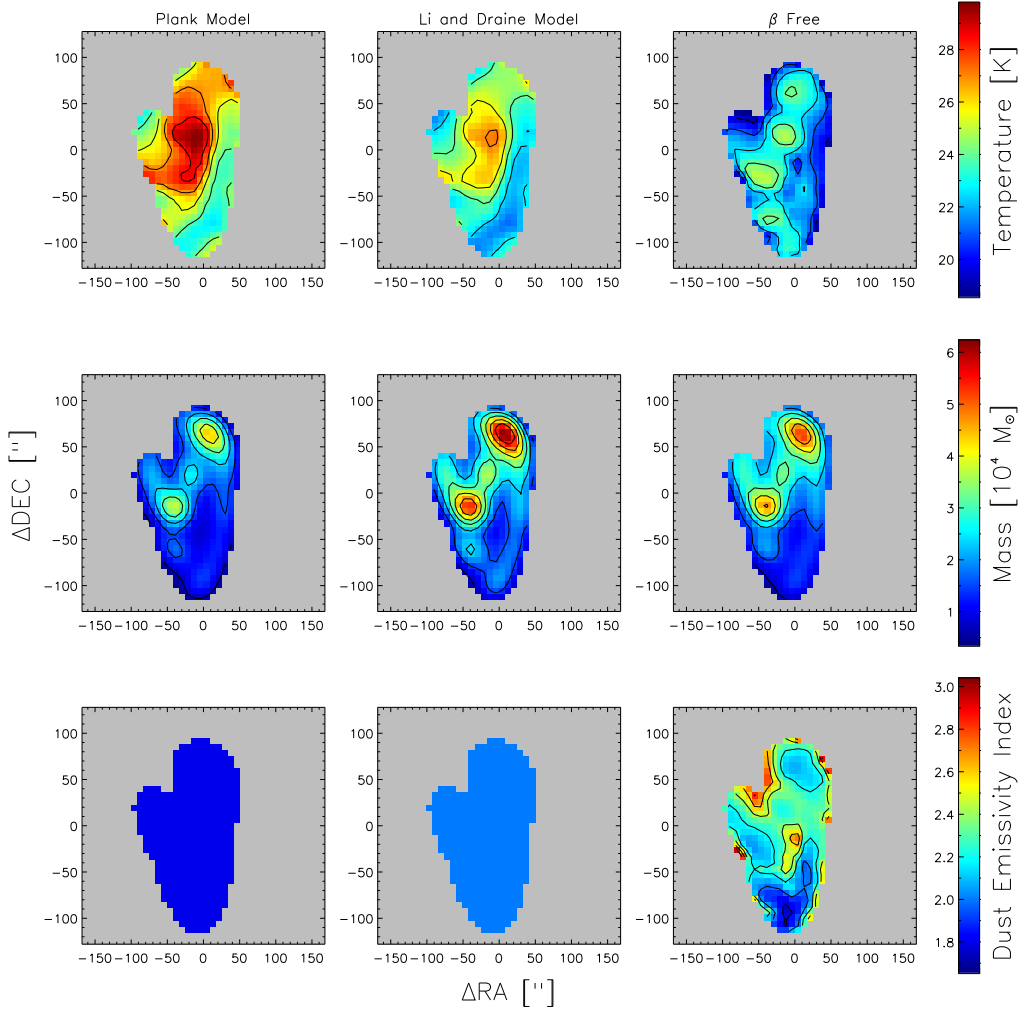


Figure 3.6: Returned value for the SED fits using the 500 μm observations with the Planck model in the left column, the Li and Draine model in the middle column, and β as a free variable in the right column. The top row shows the temperature with contours from 1.5K to 28.5K in 1.5K increments. The second row show the returned masses with contours from 1.9 M_\odot to 13.3 M_\odot in 1.9 M_\odot increments. The Li and Draine mass fits have been divided by three to better show the features. The bottom row shows the returned dust emissivity index values with contours from 1.8 to 2.8 with 0.2 increments.

Table 3.5. Best Fit Parameters for β As A Free Parameter Using $500\mu\text{m}$ Emission for Pixel Fits

Region	Average β	Total Mass [$10^5 M_\odot$]	Surface Density [$10^5 M_\odot pc^{-2}$]	Average Temperature [K]
1	2.2 ± 0.2	85 ± 38	0.17 ± 0.07	22 ± 1
2	2.3 ± 0.2	22 ± 7	0.21 ± 0.06	22 ± 2
3	2.34 ± 0.09	11 ± 1	0.21 ± 0.03	23.0 ± 0.9
4	2.2 ± 0.1	26 ± 7	0.25 ± 0.06	22 ± 1
5	2.1 ± 0.2	14 ± 4	0.12 ± 0.03	22 ± 1
6	2.1 ± 0.1	6 ± 2	0.11 ± 0.03	23.0 ± 0.9

3.2.2 Total Region Flux SED Fits

The second method used to determine the dust mass, was to fit the SED to the flux of each region in Figure 3.1. Performing the fit in this manner is beneficial because it increases the signal to noise of the region and produces a more precise set of parameters. Initially, fitting the total flux of each region was carried out in the same fashion as the individual pixel fits by fixing the emissivity index to 1.8 and 2.0, and allowing the emissivity index to vary. This method ran into issues when determining an initial mass to use in the fitting. If we used the same procedure as we did for the pixel fitting, equation 3.4, our initial masses would either converge on a local minimum or to not be able to produce a useable fit. This was due to the Levinberg-Marquardt method not being able to efficiently step through the parameter space to find a useable minimum, or the method would converge to a local minimum instead of the global minimum. The behavior of the fitted parameters with respect to the initial mass is shown in Figure 3.7 for a variable emissivity index.

Table 3.6. Best Fit Parameters for Planck Opacity Using Region Fluxes

Region	β	Mass [$10^5 M_\odot$]	Surface Density [$M_\odot pc^{-2}$]	Temperature [K]
1	2.20 ± 0.03	82 ± 6	0.15 ± 0.01	22.8 ± 0.5
2	2.30 ± 0.03	22 ± 1	0.21 ± 0.01	22.5 ± 0.3
3	2.35 ± 0.03	11.3 ± 0.6	0.21 ± 0.01	23.3 ± 0.3
4	2.30 ± 0.03	26 ± 1	0.25 ± 0.01	22.2 ± 0.2
5	2.00 ± 0.03	13.3 ± 0.8	0.019 ± 0.007	22.3 ± 0.4
6	1.95 ± 0.03	5.5 ± 0.3	0.104 ± 0.006	23.9 ± 0.4

In order to prevent any false convergences, we establish a range of emissivity index values from 1.6 to 2.9 with increments of 0.05 and found the best fit mass and temperature for each of these indices. This method ended up being significantly more robust than allowing the emissivity index to vary. The relationship between the initial mass and the fitted parameters with a fixed emissivity is shown in Figure 3.8, and shows the results are independent of the initial mass up to a cutoff point where the Levinberg-Marquardt method is unable to produce reliable fits. Determining the best emissivity index was decided by which value returned the lowest reduced χ^2 value. The reduced χ^2 and emissivity index plots are shown in Figure 3.9 for each of the regions of NGC3627, where the minimum value is marked by the red line. The best fit results using this method are shown in Table 3.6 using the opacity suggested by the Planck model and the fitted SEDs are shown in Figure 3.10.

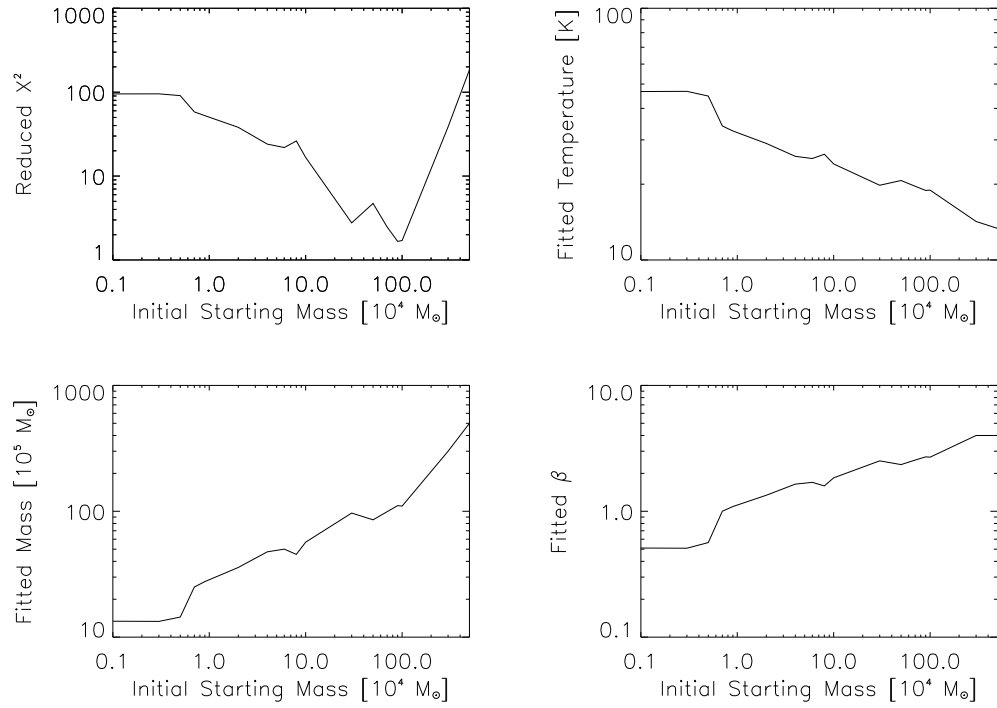


Figure 3.7: Returned SED fitting output with Planck opacity for the χ^2 , temperature, mass, and emissivity index with varying initial conditions. The top left panels shows the χ^2 values for each starting mass. The top right, bottom left, and bottom right panels show the returned temperatures, mass, and emissivity index.

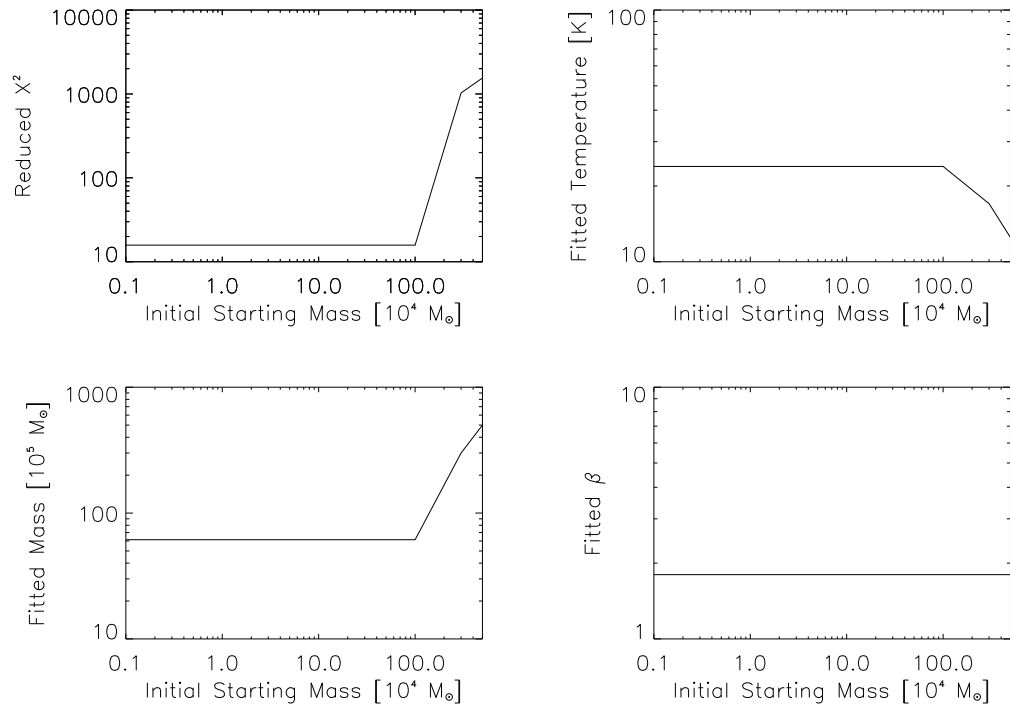


Figure 3.8: The same as figure 3.7, however the emissivity index has been fixed to 1.8.

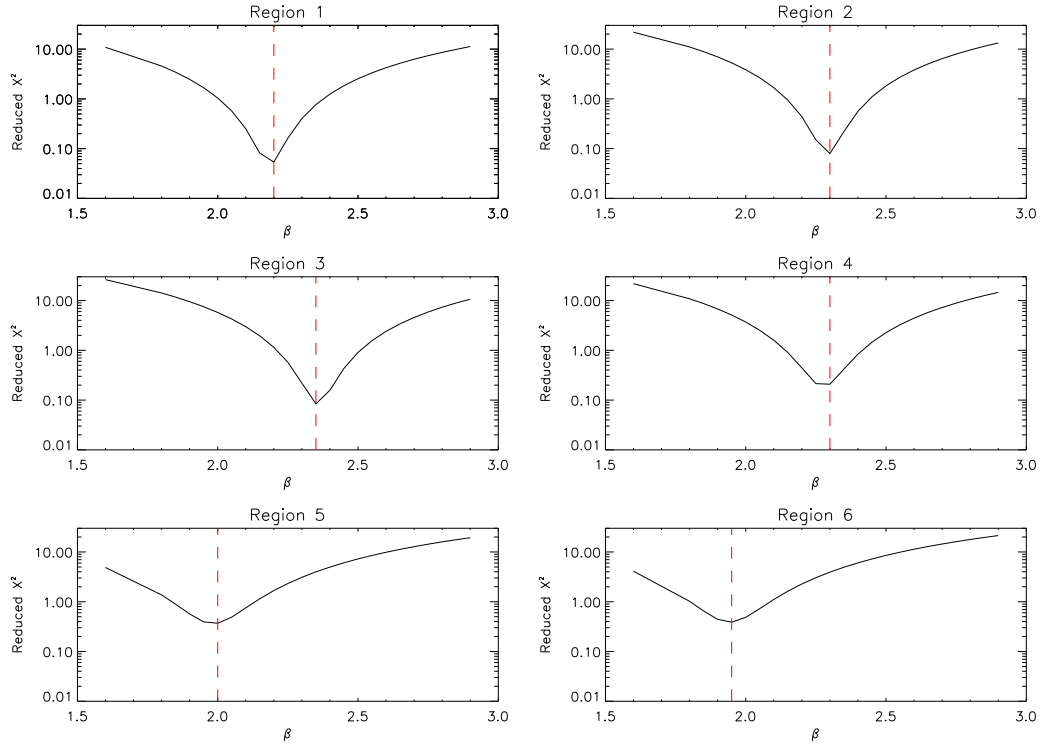
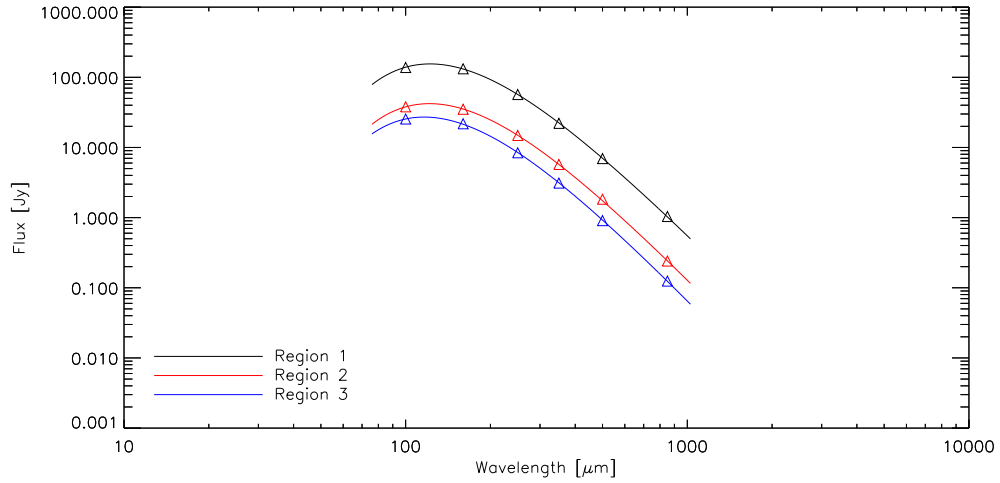
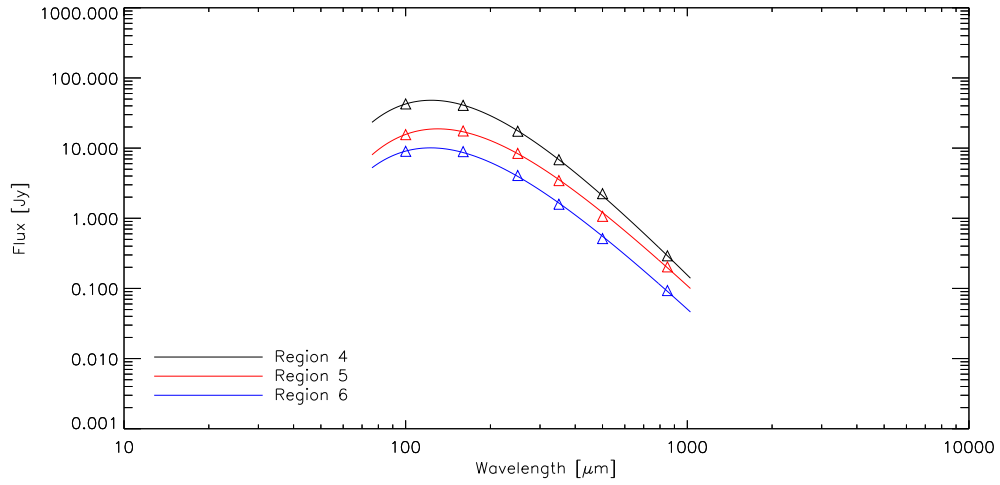


Figure 3.9: Plots for each region from figure 3.1 showing the fixed emissivity index value and the resulting reduced χ^2 value.



(a) Regions 1, 2, and 3 SEDs



(b) Regions 4, 5, and 6 SEDs

Figure 3.10: SED fits for the flux of each region in Figure 3.1 using the Planck Opacity.

3.3 Discussion

3.3.1 Reliability of Individual Pixel Fits

The issue of whether or not the individual pixel fits are reliable for a variable emissivity index is raised after seeing a strong dependence with the initial mass of the region SED fits shown in Figure 3.7. The argument can be made that since both fits have returned the same temperature ($T_{pixel}=22\pm1$ K and $T_{region}=22.8\pm0.5$ K, emissivity index ($\beta_{pixel}=2.2\pm0.2$ and $\beta_{region}=2.20\pm0.03$), and mass ($M_{pixel}=85\pm38$ M_{\odot} and $M_{region}=82\pm6$ M_{\odot}), and the removal of any initial mass dependence for the region fits, the initial values for the pixel fits are able to locate the global minimum of the χ^2 space regardless of any dependence. The inability of the region fits to successfully converge was due to an overestimate of the initial mass. Essentially, the mass associated with the $250\mu\text{m}$ flux using the Li and Draine model was much larger than the the peak of the SED. While the $250\mu\text{m}$ flux was able to better approximate the SED's peak in the pixel fits.

3.3.2 Comparison with Previous Work

We can check the validity of our results by comparing them to previous SED fits using the KINGFISH data (Galametz et al., 2012). The best fit parameters from their fitting were $T=20.2\text{K}\pm4\text{K}$ and $\beta=2.3\pm0.2$ which agree within error to our results of $T=22\text{K}\pm1\text{K}$ and $\beta=2.2\pm0.2$ for a variable emissivity index. To test agreement with our mass we need to take into account several factors.

The first is a lessened flux from the extended structure removal, second is the selection of the opacity value, and the last issue is which wavelength the mass was calculated for. If we use the same opacity model as Galametz et al. (2012) who used Li & Draine (2001), and run our fits for the mass at the same wavelength, $500\mu\text{m}$, then the only issue that remains is the spatial filtering. The amount of emission removed can be scaled to directly to the reported mass given by Galametz et al. (2012), where the amount of emission lost in the $500\mu\text{m}$ processing is 28% (Table 2.2). Removing the extended structure from the fits of Galametz et al. (2012) decreases the value from $571_{-30}^{+24} \times 10^5 \text{M}_{\odot}$ to $411_{-21}^{+17} \times 10^5 \text{M}_{\odot}$ which agrees with our mass calculation at $500\mu\text{m}$ of $396 \pm 240 \times 10^5 \text{M}_{\odot}$.

3.3.3 Effects of $850\mu\text{m}$ Emission

The major difference between our work on NGC3627 and the work done in Galametz et al. (2012) is the presence of the $850\mu\text{m}$ data point. In our analysis the presence of the $850\mu\text{m}$ has shown to have little effect on the returned parameters. The returned values for a variable emissivity index are shown in table 3.7.

The absence of a significant change in the returned parameters is indicative of a lack of any excess submillimeter emission suggestive of an abundance of a very cold ($T \lesssim 10\text{K}$) dust component (Dale et al., 2012). This excess in emission is typically seen in low-metallicity systems such as dwarf galaxies (Madden et al., 2011). The dwarf galaxies showing an excess typically have a metallicity of $\log(\text{O}/\text{H}) + 12 \lesssim 8.3$ (?). An $850\mu\text{m}$ excess has also been seen in

Table 3.7. Best Fit Parameters Excluding $850\mu\text{m}$ Emission MASS NEEDS TO BE RECALC

Region	Average β	Total Mass [$10^5 M_\odot$]	Surface Density [$M_\odot pc^{-2}$]	Average Temperature [K]
1	2.3 ± 0.3	75 ± 29	0.15 ± 0.06	22 ± 2
2	2.3 ± 0.3	18 ± 5	0.17 ± 0.05	23 ± 2
3	2.4 ± 0.2	10 ± 1	0.19 ± 0.02	23 ± 1
4	2.2 ± 0.2	22 ± 5	0.21 ± 0.05	23 ± 2
5	2.3 ± 0.2	13 ± 3	0.11 ± 0.03	21 ± 1
6	2.2 ± 0.2	5 ± 1	0.10 ± 0.02	22 ± 1

NGC3627 (Galametz et al., 2014), but our results show no excess of emission at $850\mu\text{m}$ (evident in Figures 3.4, 3.5, and 3.6). Furthermore, we should not expect to see any submillimeter excess given the $\log(\text{O}/\text{H}) + 12$ is 8.99 for NGC3627 (Moustakas et al., 2010).

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