

8.13 JUNIOR LAB MANUAL FOR COSMICWATCH: THE DESKTOP MUON DETECTOR (v3)

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Chapter 1

DOCUMENT OVERVIEW

This lab manual describes the physics associated with the third version (v3) of the CosmicWatch Desktop Muon Detectors. It will begin by outlining the forms of ionizing radiation that can trigger the detector. Then, we describe what effect these forms of ionizing radiation have on the detector and why they are important for our consideration.. We then move to information specific to the Desktop Muon Detectors, by giving an overview of the technology used and the physics behind them. The latter half of the document is dedicated to illustrating various physical phenomena described in the earlier chapters using measurements made with the detectors. Along the way, we'll briefly mention some interesting aspects of physics, which hopefully encourages you to dive a bit deeper than what we go into here. As students develop new ideas for studies with this detector, and new data is analyzed, we hope these final chapters will be continuously updated.

All supplementary material can be found in the GitHub repository located here:

<https://github.com/spenceraxani/CosmicWatch-Desktop-Muon-Detector-v3>

There are also a set of videos that provide an overview of the detectors here:

<https://youtu.be/Js-1OOS7sVE>

There are several documents that have been particularly useful when putting this together. A valuable reference now, and for the future, is the Particle Data Group's (PDG) summary of cosmic rays [1] and energy loss in matter [2]. MIT Prof. Bruno Rossi's *Cosmic Rays* [3] textbook from 1964 is a great entry level book that gives an early account of the initial investigation into cosmic-ray physics. It's a great read, and you can repeat many of the experiments he did for the first time using the CosmicWatch detectors. A more modern perspective that particularly focuses on the history can be found in Prof. Michael W. Friedlander's book entitled *A Thin*

Cosmic Rain [4]. A most thorough and comprehensive overview of cosmic-ray physics can be found in Dr. Peter K. F. Grieder's textbook *Cosmic Rays at Earth* [5]; we will be referencing it often. When discussing higher energy cosmic-ray physics, Prof. Thomas K. Gaisser's book *Cosmic Rays and Particle Physics* [6] is extremely useful. Although it is more heavily focused on the neutrino side of cosmic rays and particle physics, Prof. Masataka Fukagita's two textbooks *Physics of Neutrinos* and *Physics of Neutrinos and Applications to Astrophysics* [7], was found to often have interesting bits of information that we could not find in the other readings. The overall best description on energy loss in matter comes from *Techniques for Nuclear and Particle Physics* [8] by William R. Leo, which is a great resource for experimental particle physics. We would also recommend Prof. Claus Grupen's textbook, *Particle Detectors* [9], which is a great reference for describing detection methods in particle physics.

The text of this lab manual is a lightly modified version of Spencer Axani's Masters Thesis, available at <https://arxiv.org/abs/1908.00146>. The primary changes are to chapter 5, and thereafter. We now describe an upgraded detector compared to the version in the thesis. We remove the instructions on how to construct the detector, since students are sent fully assembled devices. And we have divided the ideas for studies into three chapters. The first describes experiments that we have tested and can be done at home; the second describes untested ideas that, in principle, can be done at home with some careful planning; and the third are studies at exotic but inspiring locals. This manual assumes that students have access to two fully constructed CosmicWatch detectors and a Red Pitaya oscilloscope.

The CosmicWatch detector design has been developed with co-collaborator Dr. Katarzyna Frankiewicz. The project has had funding from the MIT Physics Department, the Wisconsin Particle Astrophysics Center and the National Science Foundation.

We are excited to have the CosmicWatch detector incorporated into Junior Lab! Developing this detector has been a passion, not just a hobby. We hope you enjoy it, and even consider building your own. See Spencer's thesis for step-by-step instructions for a slightly simpler version if you would like to build one yourself.

Chapter 2

SOURCES OF IONIZING RADIATION

We'll begin by discussing the various sources of ionizing radiation that are capable of triggering the Desktop Muon Detectors. There are a few forms of ionizing radiation that are particularly important due to their abundances and their energies. In passing, we will also discuss why we are not sensitive to certain other forms of ionizing radiation, however a description of them is often important for completeness.

2.1 Cosmic radiation

The Earth is continuously being bombarded by a flux of particles called *cosmic-rays*. Approximately 74% by mass of this flux comes from ionized hydrogen (free protons), 18% from helium nuclei (two protons and two neutrons), and the remainder is trace amounts of heavier elements [1]. The majority of the cosmic-ray flux observed at Earth is relativistic, in that the individual nuclei have kinetic energies greater than their rest mass ($E_k/mc^2 > 1$). The lower energy cosmic-rays (GeV-scale) are greatly influenced by the solar wind as well as the geomagnetic field (see for example Fig. 2.1), which limits the flux interacting with the Earth. The high energy flux extends up to 10^{11} GeV at which point the cosmic-ray loses energy from interactions with the cosmic microwave background. This is often referred to as the GZK cutoff¹ (GZK stands for the initials of the three principle authors of the first theoretical prediction: K. Greisen [12], G. T. Zatsepin, and V. A. Kuzmin [13])). The energy of the cosmic-rays falls off rapidly with energy: below 10^6 GeV, the flux falls off as $E^{-2.7}$, and above this it steepens to approximately $E^{-3.1}$ [6]. For perspective, the number of 1 GeV cosmic-ray protons is 8.1 orders of magnitude higher than that at 1000 GeV (i.e. 2.7×3), or 16.2 orders of magnitude higher² than that at

¹The highest energy cosmic-ray observed was measured to have approximately 3×10^{20} eV [10] (48 joules), equivalent to a brick falling on your toe [11], all contained in a single proton, and later named the *Oh-My-God Particle*.

²The steep fall off in energy of the cosmic-ray flux is why we require large detectors to measure the rare high energy events.

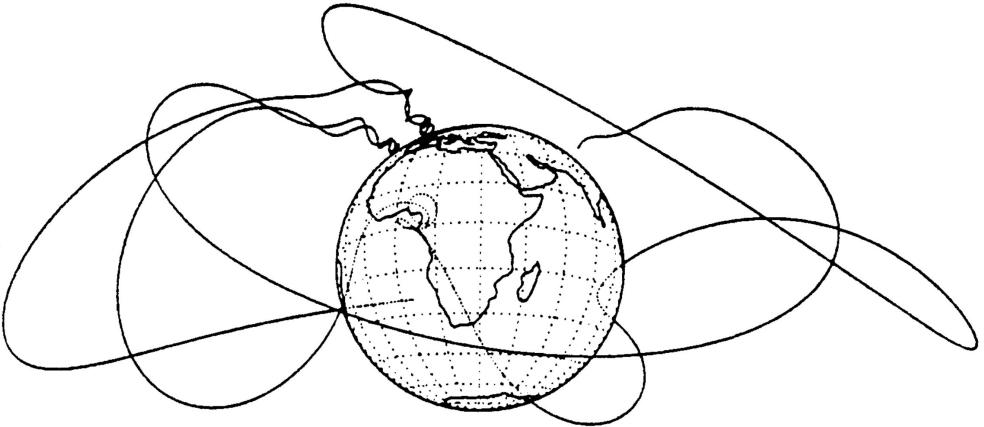


Figure 2.1: Simulated trajectories of low energy cosmic-rays interacting with the geomagnetic field. From Ref. [4]

10^6 GeV (i.e. 2.7×6).

When a primary cosmic-ray collides with a nucleus in the upper atmosphere (typically with the nucleus of an oxygen or nitrogen molecule), the energies can be large enough to break apart both or either of the primary particle or the target nucleus through a nuclear interaction. Much of the energy of the collision goes into producing short lived particles known as *mesons*³, the most common being the π -meson or pion (π^+, π^-, π^0) and then the K-meson or kaon (K^+, K^-, K^0). The charged pions (π^\pm) decay within approximately ten billionths of a second [16] producing a same charge muons and a neutrinos (the charged kaons, K^\pm , are a bit more complicated, but they also preferentially decay this ways as well, or to pions). The neutral mesons (π^0, K^0), decay approximately one billion times faster (10^{-17} s) than the charged mesons, preferentially to gamma rays. Unlike the neutral mesons, the charged mesons are able to travel far enough before decaying to interact with another molecule in the atmosphere. This interaction in turn can be in the form of another nuclear interaction (since mesons are made of quarks, they experience the strong force responsible for the nuclear interaction), much like the original cosmic-ray interaction. The interaction may then produce even more mesons, contributing to the shower of particles induced by the primary interaction. The primary cosmic-rays do not penetrate directly to Earth's surface due to the shielding provided by the atmosphere, however, a small flux of nuclear fragments (such as protons and neutrons) from these interactions can occasionally cascade down and make it to the surface. An illustrative diagram of a cosmic-ray interaction is shown in Fig. 2.2. The first interaction of vertical cosmic-rays takes place at an altitude approximately between 15 and 20 km. cosmic-rays entering at an angle, will interact at higher altitudes due to their path passing through more atmosphere [7].

³Mesons are particles that, unlike the proton and neutron, contain only two quarks: one quark and one anti-quark. The lightest meson is the pion, then the kaon. There exists many other combinations of quark/anti quark pairs, however due to their higher masses, they are not preferentially produced and will not be discussed here. More information on this side of particle physics can be found in Ref. [14, 15]

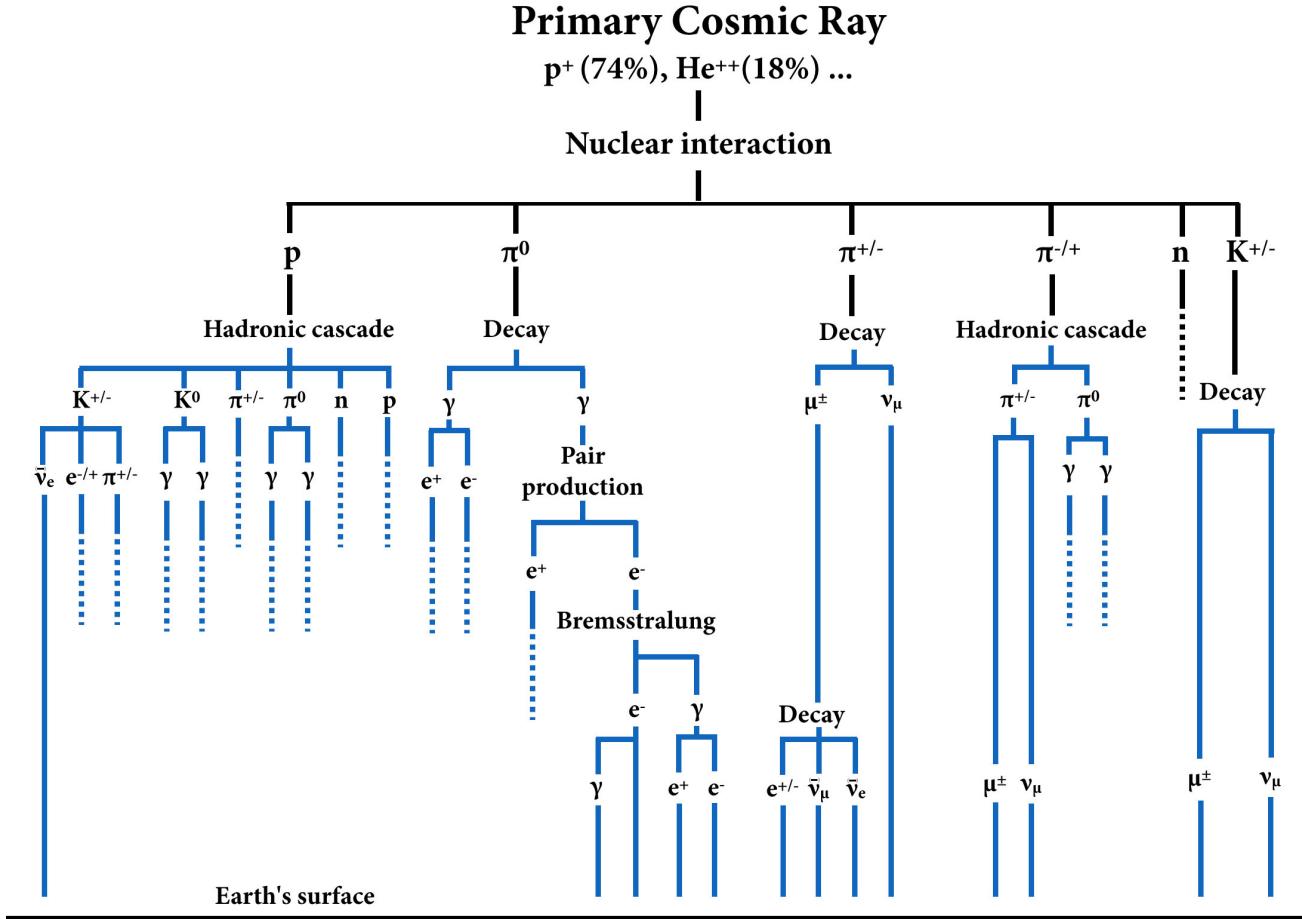


Figure 2.2: An schematic representation of the various decay and interactions chains that result from the interaction of a cosmic-ray in Earth’s atmosphere. This is a modified figure from Ref. [3].

The high energy photon from the decay of the neutral mesons quickly materialize into an electron-positron pair, also referred to as *pair production*. These electron-positron pairs then radiate high energy photons, which can again materialize into another electron-positron pair. This electromagnetic cascade process continues, dividing the original energy of the photons between the numerous electrons, positrons, and lower energy photons at the end of the cascade. Photons with energies less than 1.022 MeV cannot further pair produce and their interactions will be dominated by Compton scattering and photoelectric absorption. At lower altitudes, there isn’t a fresh supply of high energy neutral mesons due to the rapid decrease of nuclear interactions at lower altitudes.

The cosmic-ray muons (μ^\pm) originate from the decay of the charged mesons. A charged pion will decay to a same-sign muon (and muon-neutrino) with a branching fraction of 99.98%, whereas a charged kaon decays to a muon (and muon-neutrino) 63.5% of the time [17]. The neutrinos are not electrically charged and only interact through the weak force, therefore they can be fully

ignored in this discussion.

$$\begin{aligned}\pi^\pm &\rightarrow \mu^\pm + \nu_\mu(\bar{\nu}_\mu) \dots (99.98\%) \\ K^\pm &\rightarrow \mu^\pm + \nu_\mu(\bar{\nu}_\mu) \dots (63.5\%) \end{aligned}\tag{2.1}$$

Approximately 80-90% of the muon flux in the energy range of interest (GeV to TeV-scale) comes from the decay of pions, and the remainder from kaons [18]. The muons are particularly penetrating, that is, they essentially only lose energy due to ionization as they travel through the atmosphere and other matter and can make it through a large amount of material. This is contrasted with *baryons* (particle comprised of quarks), which also interact through the strong force. This makes muons the most numerous charged particle showering down on the Earth's surface. Muons have a mass of 105.65 MeV and are also unstable particles with a half-life of 2.2×10^{-6} s. They decay to an electron and two neutrinos.

$$\mu^\pm \rightarrow e^\pm + \nu_e(\bar{\nu}_e) + \bar{\nu}_\mu(\nu_\mu) \dots (100.0\%),\tag{2.2}$$

however, again we can ignore the neutrinos. A cosmic-ray muon with an energy greater than 2.4 GeV will be sufficiently relativistic that its half-life, as seen by an observer on the Earth, will be dilated enough that it has a significant chance of reaching the Earth's surface before decaying. In other words, the muon decay length becomes greater than 15 km – the approximate altitude of the cosmic-ray muon production. Further, a typical muon will lose approximately 2 GeV of energy due to ionization as it passes through the atmosphere on its way to the ground. Combining these two facts with the fact that the geomagnetic field and interstellar solar winds drive back the GeV-scale cosmic-rays, as well as the steeply falling cosmic-ray energy spectrum, we can expect the average muon energy at Earth's surface to be greater than a few GeV. We typically quote the mean cosmic-ray muon energy at Earth's surface to be approximately 4 GeV [1]. The muons that do not survive the journey to the Earth's surface decay. The resulting electrons (or positrons in the case of a μ^+), also referred to as *Michelle electrons*, contribute to the low energy electromagnetic component from cosmic-ray showers at sea level.

Let us now compare the number of particles showering down on us at sea level. We'll limit ourselves to thinking about particles that are coming from one steradian about the zenith, this can be thought of as circular disk around the vertical (zenith) part of the sky whose area represents approximately 1/6th of the total visible sky, or equivalently, a half-angle of 32° from the zenith. From this direction, we can describe the number of particles passing through a $1 \times 1 \text{ cm}^2$ horizontal surface per minute ($\text{cm}^{-2} \text{ min}^{-1} \text{ sr}^{-1}$) by following the measurements outlined in Ref. [1]. We expect approximately $0.4 \text{ cm}^{-2} \text{ min}^{-1} \text{ sr}^{-1}$ from μ^\pm with energies greater than 1 GeV; $0.2 \text{ e}^\pm \text{ cm}^{-2} \text{ min}^{-1} \text{ sr}^{-1}$ above 10 MeV, but the flux falls off fast with energy, becoming negligible above 1 GeV; and $0.0054 \text{ cm}^{-2} \text{ min}^{-1} \text{ sr}^{-1}$ from protons above 1 GeV; and a charged meson flux above 1 GeV two orders of magnitude lower than that of the proton flux. This means that the protons and charged mesons are insignificant at sea level, however, there is a significant muonic and low-energy electromagnetic component. In Ref. [5], the flux is divided into a hard component (essentially fully muons), which can penetrate 15 cm of lead,

and soft component (approximately 60-65% muons and the remainder is electrons, positrons, and photons), which cannot. As we'll soon see, there are a variety of physical phenomena that influence fluxes, but the relative contributions listed here represent a useful approximation.

As we increase in altitude, the relative contribution from the ionizing radiation fluxes change. In particular, we see a larger contribution from both the protons and electromagnetic component, whereas the charged mesons are still sub-dominant. Once we pass the primary interaction region where the primary cosmic-rays are most likely to interact (typically around 15-20 km), the secondary particles produced by the initial nuclear interaction die off and we see a decrease in the ionization radiation flux. The shape of the curve describing the ionizing radiation flux as a function of altitude is called the Pforzner curve, and where the ionizing particle production reaches a maximum is termed the Regener-Pforzner maximum [19].

2.1.1 Flux variations due to solar system properties

There are several properties associated with the interstellar medium that modulate the cosmic-ray flux, and in particular the observable cosmic-ray muon flux. These properties are primarily associated with behaviour of the Earth and Sun's magnetic field.

The latitude effect: Roughly speaking, the Earth has a magnetic field that behaves similarly to a magnetic dipole orientated from north to south. The magnetic field points parallel to the surface of the Earth near the equator, and perpendicular to the surface near the poles. Particles traveling towards the Earth will be less deflected ($F = q\vec{v} \times \vec{B}$) near the poles compared to the equator. Low energy charged particles passing through the magnetic field may even become trapped in what's known as the *Van Allen radiation belt*. This presents a low energy cutoff, where the magnetic field is able to deflect protons below approximately 10 GeV near the equator (corresponding to a rigidity of 10 GV) and near 1 GeV at higher latitudes [20, 21].

The East-West asymmetry: The cosmic-ray muon flux is larger looking towards the west compared to the east due to the Earth's magnetic field. This is an effect produced by primary cosmic-ray particles being predominately positively charged. The positively charged muons curve towards the east, meaning that the intensity from the west is stronger. This effect is more evident in the upper atmosphere [5], and obviously a larger effect at the geomagnetic equator than at the poles.

Magnetic anomalies: There are local geomagnetic field variations, which causes a change in the cosmic-ray intensity. The most prominent being the *South Atlantic Anomaly (SAA)* [22], which extends from the east coast Brazil to the west coast of southern Africa (-50.0 to 0.0 geographic latitude and from -90.0 to 40.0 longitude). This is the region where Earth's inner Van Allen radiation belt extends closest to the planet's surface and provides the smallest amount of protection from cosmic-rays. In fact, the increased level of ionizing radiation when passing through the SAA is responsible for radiation damage to electronics on-board Low-Earth Orbit

(LEO) satellites⁴.

Solar modulation: The observed cosmic-ray flux at the top of the Earth's atmosphere depends partially on solar activity, which manifests itself as an 11-year cycle⁵. Solar winds can drive back low energy cosmic-rays entering the solar sphere and the modulation effect decreases with an increase in energy. According to Ref. [7], the 1 GeV cosmic-ray proton flux is twice as small during maximum solar activity compared to minimum solar activity; similarly there is a 10% reduction in the 10 GeV cosmic-ray protons during the solar maximum.

Solar Flares: Solar-flares can eject protons with energies up to several GeV, the upper end of which is able to produce muons through nuclear interactions. These events are rare transients, and since the energy is low, it primarily has an effect on the low energy muon flux [28].

2.1.2 Flux variations due to atmospheric properties

Similar to the previous subsection, there exists terrestrial phenomena that also modulate the cosmic-ray muon flux.

The Cosine Squared Law: At greater angles from the vertical, cosmic-ray muons must travel through a much larger distance, and therefore amount of matter, to reach a ground-based observer. A cosmic-ray muon traveling vertically downwards may only travel through 15 km of atmosphere, whereas one traveling in the horizontal direction must pass through approximately 500 km of atmosphere. The larger path length means that the muon will lose more total energy due to ionization in the atmosphere and also have a higher probability of decaying before reaching the ground. As a function of zenith angle, the cosmic-ray muon intensity is expected to follow a cosine squared dependence [1].

The atmospheric attenuation: Recall that the nuclear interactions between the primary cosmic-ray and atmospheric nucleus happen in the upper atmosphere. Therefore, particles reaching sea level must have had sufficient energy to penetrate the remainder of the atmosphere. An increase in atmospheric density (perhaps due to atmospheric pressure changes) will cause secondary particles to lose more energy as they propagate to the Earth's surface. Due to this, the muon rate turns out to be anti-correlated with the pressure (i.e. if the atmospheric pressure increases, the cosmic-ray muon rate decreases). The density of the atmosphere changes with the season and therefore exhibits a time-dependence. From other measurements, this is expected to be a percent level effect [29].

⁴The International Space Station (ISS) passes intermittently through the SAA [23] and has dedicated instrumentation for measuring the increased radiation dose to astronauts [24]. Astronauts from NASA missions as early as Apollo 11 [25] have also reported seeing flashes of light while being in orbit [26]. These flashes of light are attributed to high-energy particles in the space radiation environment, however many details on the origin are still unknown [27]

⁵There is also a 22-year cycle since the solar magnetic dipole flips polarity at every solar maximum, which occurs every 22 years [5]

The positive temperature effect: To produce a muon, we require a charged meson to decay. However, recall that the charged mesons are typically relativistic and have lifetimes on the order of nanoseconds⁶. This gives the charged mesons sufficient time to potentially interact with another nucleus in the atmosphere rather than decay. As the temperature increases, the atmosphere expands and there are fewer particles to interact with, thus increasing the probability of decaying rather than interacting [30].

Rather than correlating this with the ground based pressure (as in the paragraph above), it is more commonly correlated with atmospheric temperature – taking into account the temperature profile of the atmosphere. This effect is larger at higher energies and therefore is typically measured in laboratories located deep underground where the low energy cosmic-rays have less of an influence [31–33].

The negative temperature effect: As the temperature of the atmosphere increases, the atmosphere expands, moving the muon production region further out. This means that the muon path length increases, which gives them a higher probability of decay prior to making it to the ground. During the winter when the atmosphere is colder, shallower and more dense, cosmic-ray interactions happen closer to the Earth’s surface. The charged mesons quickly begin to lose energy and have a less likely chance of decaying into muons.

2.2 Radioactive backgrounds

The previous section described the ionizing radiation that we expect from showers of particles raining down from the upper atmosphere, and the expected phenomena that can modulate this flux. This section will describe ionizing radiation that originates on the surface of Earth and can also influence our measurements; we’ll refer to these as the *radioactive backgrounds*. Radioactive backgrounds are sub-divided into primarily three main processes called alpha, beta, and gamma radiation. Radioactivity is a quantum mechanical effect, which is non-deterministic, that is, we cannot predict when a particle will decay, rather we can only assign a probability to it. The energy scale of these processes are relatively low (MeV-scale) compared to the energies associated with the cosmic-rays (GeV and above), but their natural abundance on the surface of the Earth is sufficient that these are typically the dominant source of triggers in the Desktop Muon Detector.

Alpha decay is the result of an unstable nucleus ejecting a helium nucleus (a bound state of two protons and two neutrons), $(Z, A) \rightarrow (Z - 2, A - 4) + \alpha$. This is a quantum mechanical effect, where a helium state (helium is a very tightly bound state) forms in the nucleus, then quantum tunnels through the nuclear potential barrier, exiting the nucleus. The emitted alpha particle

⁶For example, a 5-GeV π^\pm produced at 15 km will travel approximately 300 m before decaying. This distance is small compared to the interaction path length of approximately 13 km, which means that most charged pions will decay rather than interact. However, at approximately 115 GeV, the pion has an equal probability to interact or decay in the atmosphere.

is mono-energetic, and since the helium nucleus has a charge of $+2e$ and mass of approximately 4 GeV (therefore, it moves slow and has a large charge), it will lose energy rapidly in matter. A 5-MeV alpha particle will have a range of 3.5 cm in air before losing all of its energy, or equivalently, 23 micrometers in aluminum⁷ [34].

Beta radiation is described as the decay of a neutron to a proton⁸: $n \rightarrow p + e^- + \bar{\nu}_e$. The proton remains in the nucleus, while the electron and electron-neutrino are ejected. Since this is a 3-body decay, the electron is not mono-energetic. It is emitted with a continuous energy spectrum whose maximum energy is approximately at the total energy available for the decay (the Q-value). Beta decays typically have energies that can range from tens of keV to a few MeV.

Gamma radiation is simply a high-energy photon, emitted during the de-excitation of an atomic nucleus. When the nucleus is in an unstable state (for example, maybe the nucleus absorbed a neutron or was left in an excited state after a beta decay), it will de-excite into a lower energy configuration releasing a photon. This is analogous to the de-excitation of an atomic electron, emitting a characteristic mono-energetic photon. Since the energy levels in the nucleus are quantized, gamma ray are also mono-energetic (with a small spread due to nuclear motion). These energy scales are in the 100 keV to MeV-range.

⁷The high energy loss rate that alpha radiation makes it useful for cancer therapies. An alpha particle will deposit all of its kinetic energy into a very local space (order micrometers in human tissue), which is capable of destroying cancerous cells.

⁸More fundamentally, during neutron decays, a down-quark in the neutron converts to a up-quark, emitting a virtual W-Boson $u \rightarrow d + W \rightarrow d + e^- + \bar{\nu}_e$. On the macroscopic level, this appears as the transmutation of an atom converting to another atom with an extra proton and one fewer neutron: $A(Z, N) \rightarrow A'(Z + 1, N - 1) + e^- + \bar{\nu}_e$.

Chapter 3

PARTICLE INTERACTIONS WITH MATTER

In order to detect a particle, it must undergo an interaction with the material of a detector. This section will go through the main interactions that transfer energy from the particle to the absorbing material in sufficient detail to understand the data from the Desktop Muon Detector. We'll begin with the muons, protons, pions/kaons, and generically other high-energy heavy charged particles, then move to high-energy electrons (and positron), and finally onto the interactions associated with the radioactive backgrounds.

3.1 High energy heavy charged particles

The description below will be useful when thinking about any charged particles with a mass much greater than that of the electron ($m \gg m_e$). This happens to be all charged particles except for the electron and positron (e.g. the muon, the next lightest charge particle, is 206 times more massive than the electron). Unlike the trajectories of heavy charged particles, the trajectories of electrons are not straight lines in a target and special consideration is required. The description below represents an approximation to the underlying processes responsible for energy loss in matter, since the breadth of this subject is far too large to cover in a single document. More information can be found in Ref. [2, 9, 35, 36].

The energy loss rate (often called the stopping power), $-dE/dx$, is a measure of how much energy is lost per unit distance traveled. It is often expressed in units of MeV cm²/g (referred to as *mass stopping power*), where one simply has to multiply by the density (in terms of g/cm³) of the absorbed to get the energy lost per cm. We've taken care of this by expressing the energy loss rate in terms of water ($\rho = 1.0\text{ g}/\text{cm}^3$), which coincidentally has approximately the same density of plastic scintillator. As a particle travels through matter, it can be broken up into three energy ranges that are mass dependent:

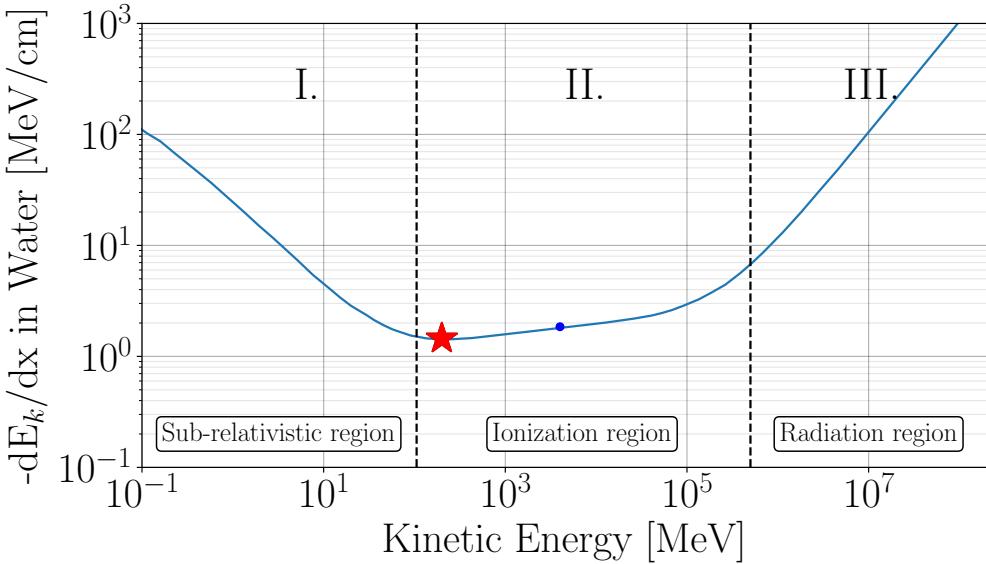


Figure 3.1: The kinetic energy loss per centimeter traveled, of a muon traveling through H_2O ($\rho = 1.0 \text{ g}/\text{cm}^3$). Modified from Ref. [16]. The blue dot represents the mean energy of the cosmic-ray muons at sea level (4 GeV), and the red star represents where the particle is minimally ionizing.

- I. The sub-relativistic region: $(E_k < mc^2)$
- II. The ionization region: $(E_k > mc^2 \text{ and } E < 400 \text{ GeV m}^2/\text{m}_{\mu}^2)$
- III. The radiation region $(E > 400 \text{ GeV m}^2/\text{m}_{\mu}^2)$.

These three regions are shown in Fig. 3.1 for a muon traveling in water, however the energy loss can be scaled to another material by simply multiplying by the density (in g/cm^3) of the material (for example, the stopping power in lead would be scaled by a factor of 11.34). The energy loss can be scaled to another particle by multiplying by the charge of the particle squared. For example, the energy loss of an alpha particle will be scaled by a factor of 4. This plot is often represented in terms of momentum, but we've scaled it in terms of the kinetic energy of the incident particle to simplify the description. The vertical-dashed lines separate the three regions and can be scaled to another charged particle using the equalities above. For example, a proton will enter the sub-relativistic region at approximately 1 GeV.

In the **sub-relativistic region** (Muon: $E_k \approx 100 \text{ MeV}$; Proton: $E_k \approx 1 \text{ GeV}$), as the particle loses energy, the amount of energy loss per unit distance travels increases. Essentially, this means that once a particle enters this region, it comes quickly to a stop. This phenomena is also known as the *Bragg peak* [8].

In the **high energy radiation region** (Muon: $E_k \approx 400 \text{ GeV}$; Proton: $E_k \approx 27 \text{ TeV}$), the energy loss is associated with *bremsstrahlung*, pair production, and nuclear interactions, and

scales linearly with energy. The radiation term begins dominating at approximately 400 GeV for muons, however, recall that the cosmic-ray flux falls off fast with energy. The muons that are in this regime represent a small percentage of the flux and lose energy fast.

Finally, the **ionization region** is where the majority of the cosmic-ray muon flux lies. Recall that the mean muon energy at sea level is 4 GeV (indicated as a blue marker in Fig. 3.1). The energy loss in this region is due to ionization (breaking electromagnetic bonds) and excitation (raise the electron to a higher-lying shell within the absorber atom) of the incident particle and described by the *Bethe Bloch formula* (great descriptions of the formula can be found in Refs. [8, 9, 16]). This region is particularly interesting since the energy loss rate is nearly constant (it actually increases logarithmically), with an average energy loss rate of 2.2 MeV/cm in a density 1.0 g/cm³ material, over many orders of magnitude. The minimum here, indicated by a red star in Fig. 3.1, is where the muon is said to be a minimum ionizing particle (MIP) and represents the energy at which the muon is most penetrating. The function is so flat in the region of this minimum (up to approximately 400 GeV), that any particle whose energy is near this red star is often called a MIP. It is also coincidental that the majority of the cosmic-ray flux falls into this region. This means that if we want to approximate the penetrating depth of a typical cosmic-ray muon, we can simply divide the energy by 2.2 MeV/cm and multiply by the density of the absorber. As an example, a 10 GeV muon will penetrate through approximately 17 m of concrete ($\rho = 2.7 \text{ g/cm}^3$). What about the other heavy charged particles discussed in this document: the protons, pions, and kaons? These will also lose energy through ionization, but since they are composed of quarks, they can also interact via the strong force. The strong force is responsible for the nuclear collisions that can greatly impact the particle and trajectory. This is what makes the muons so unique – they do not interact via the strong nuclear force and they are heavy, which allows them to penetrate through materials with minimal loses due to collisions with the electrons in the absorbers and with minimal deflection on their trajectory¹.

3.2 High energy electrons/positrons and photons

As described in Section 2.1, there is a non-negligible source of electrons/positrons with energies below 1 GeV showering down onto the Earth's surface. This section will describe the energy loss in terms of high energy electrons, but the description is valid for positrons as well. It turns out that for an electron above a few tens of MeV, the energy loss will be dominated by radiation losses, predominantly *bremsstrahlung radiation* ("braking radiation" in German). Bremsstrahlung radiation is the emission of photons produced by a particle accelerating and decelerating as it passes near the electric field of the nucleus of the material.

¹Since high energy muons are able to penetrate very large distances through material, many experiments are buried kilometers underground to shield against them. For example, the neutrino detector Super-Kamiokande, is buried underneath a 1-km mountain in Japan, in order to reduce the muon flux by a factor of 10^5 so that they are not swamped when looking for the rare, less energetic interactions from neutrino. For similar reasons, the IceCube neutrino detector is buried under 1.4 km of ice in the Antarctic glacier at the South Pole.

A bremsstrahlung photon with sufficient energy can pair produce an electron and positron, which will subsequently radiate other bremsstrahlung photons, thereby creating a cascade of electrons, positrons, and photons. This process dominates the energy loss of the electron until the energy drops below a few tens of MeV (typically referred to as the critical energy). The radiation energy loss rate actually scales with energy, so in the region above the critical energy, dE/dx is proportional to energy. This means that a 20 GeV electron will start off by losing 1000 times more energy per centimeter traveled than a 20 MeV electron in the form of bremsstrahlung radiation. This causes the high-energy particles to lose their energy rapidly as a function of distance.

A *radiation length*, X_0 , is defined as the average thickness of a material that reduces the mean energy of the electron or positron by a factor $1/e$ (Euler's number = 2.71828) due to bremsstrahlung radiation. This means that an electron will lose a factor of e^{-t} energy after traveling through t radiation lengths. For example, after four radiation lengths, a 1 GeV electron will end up with approximately 20 MeV. A simplified model known as the *Heitler Model for electromagnetic cascades*, approximates this saying that one electron-positron pair will be created per radiation length, each of which will receive half of the energy of the photon that produced them. After t radiation lengths, the cascade will contain 2^t particles (electron, positrons, and photons) each with an average energy of approximately $E = E_0/2^t$ [37]. We've illustrated this in Fig. 3.2.

A useful list of radiation lengths for various materials can be found in Table 3.1. This table includes the information regarding pertinent materials used in the Desktop Muon Detectors or during some measurements below.

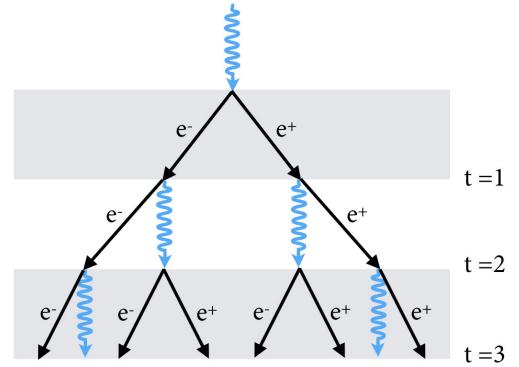


Figure 3.2: The Heitler Model for electromagnetic cascades.

Material	Density [g/cm ³]	Radiation length [cm]	Critical Energy [MeV]
Water (H ₂ O)	1.00	36.1	92
Lead (Pb)	11.35	0.56	9.51
Concrete	2.5	10.7	
Air at STP*	1.2931	30420	102
Scintillator (Polystyrene)	1.032	42.4	109
Aluminium (Al)	2.70	8.9	51.0

Table 3.1: A table of materials that are mentioned in this document and their corresponding radiation length. *STP indicates that the air is at the standard temperature and pressure: 20°C and 101.325 kPa. The data was collected from Ref. [1, 8].

3.3 Low energy electrons/positrons

Electrons and positrons are the lightest charged particles. The qualitative behavior for electron scattering is different from the high-energy particles in two ways. First, the energy loss by electrons fluctuates much more than for heavy particles. For example, the maximum transferable kinetic energy of an electron from a 4 GeV electron is the full 4 GeV (since they have the same mass), whereas a muon with the same energy has a maximum transferable energy of approximately 1 GeV. Further, because of its small mass, electrons are particularly susceptible to large angle deflections by scattering off a nucleus. This probability is so high, in fact, that multiply scattered electrons may be turned around in direction altogether, defined as *backscattering*. The backscattering probability is higher at lower energies, and if backscattered, the electrons do not deposit all their energy in the absorbing medium. A 1 MeV electron has approximately a 10% chance of backscattering off of a thick slab of aluminium, and a 50% chance of backscattering off a slab of gold [38].

The previous section described electrons and positrons with energy above the critical energy of a material (typically tens of MeV), where their energy loss is completely dominated by bremsstrahlung radiation. At lower energies, the electrons-positrons can in-elastically interact through Coulomb collisions with atomic electrons to lose energy [39]. This leads to ionization and excitation like the heavier particles. At lower energies still, MeV-scale, the electrons (positrons) also exhibit *Møller (Bhabha)* scattering.

3.4 Low energy gamma rays

Gamma-rays interact slightly differently from the charged particles due to their lack of electric charge. The three main interactions of gamma rays (and X-rays) are shown in Fig. 3.3.

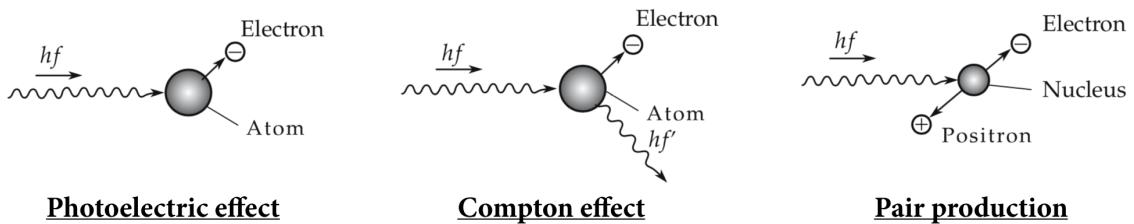


Figure 3.3: Modified from Ref. [34]

An atomic electron can fully absorb the energy of a gamma ray², resulting in an electron with the energy of the initial gamma ray (MeV-scale) minus the binding energy of the atomic electron (eV-

²In order to conserve momentum, the photoelectric effect cannot occur on a free electron, it requires a nucleus to absorb part of the recoil.

scale). This process is known as the *photoelectric effect*. As shown in Fig. 3.4, the photoelectric effect dominates for low-energy gamma-rays with a moderate-to-high density absorber.

Compton scattering results in the partial transfer of energy from the incident gamma-ray to an atomic electron, resulting in the electron being bumped into a higher energy level or being ionized. The gamma-ray may change direction, exiting the material that contained the electron, or perhaps it scatters again off another electron. Compton scattering dominates at low energy for very low-Z materials, and the probability of scattering will be proportional to the electron density and therefore proportional to the proton number of the material.

At energies above 1.022 MeV, electron-positron pair production plays a role. Pair production follows the same description as that found in the description on high-energy electrons (Sec. 3.2). The only difference being that the chain begins with a photon rather than an electron.

Beyond this, there are second order effects, such as *Rayleigh scattering* where the photon wavelength is large enough that it coherently scatters off of the entire atom, and photonuclear interactions at higher energies that break up the nucleus [16].

3.5 Neutrons

Like the photon, the neutron is not electrically charged, therefore it is not subject to Coulomb interactions with the electrons and nuclei. Rather, it interacts through the strong force with nuclei. Due to the short range nature of the strong force, these interactions are comparatively rare (the neutron needs to get close to the nucleus in order to interact). Several interactions that may occur are [8]:

1. Elastic scattering off a nucleus
2. Inelastic scattering off nucleus that leaves the nucleus in an excited state that may decay emitting a gamma ray. To excite the nucleus the neutron must transfer MeV-scale energies.
3. Neutron capture. At low energies, the neutron might be captured by the nucleus, emitting a gamma ray.

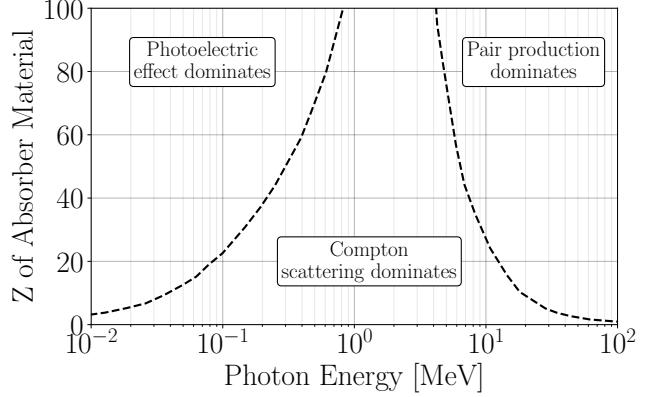


Figure 3.4: Dominant region for the three gamma ray interactions as a function of the photon energy and the target charge number Z . Modified from Ref. [9]

4. Fission
5. Hadronic shower, particularly at high energies (> 100 MeV).

High energy neutrons produced in the primary cosmic-ray interaction, will often collide with another nucleus creating a similar interaction as to the primary cosmic-ray.

Chapter 4

DETECTION METHODS

4.1 Single photon detections: photomultipliers

One of the most common instruments used by particle physicists are photomultipliers. Photomultipliers are devices capable of producing a measurable electrical signal from the interaction of a single photon. Photon detection enables us to extract information related to the incident particle by measuring the photon emission of a particle as it loses energy in a material. The most prolific single photon sensing device is a photomultiplier tube (PMT), which offers a large photo-sensitive area of coverage at modest cost. They are however bulky and require high-voltage. Other common technologies, such as the Avalanche Photodiode (APD) and P-type and N-type semiconductor photodiode (PIN photodiode), both have their benefits and drawbacks. Recent improvements in the manufacturing of silicon chips have made it possible to make a new type of photon detector called a *silicon photomultiplier*, or SiPMs (often abbreviated as SPM). SiPMs have many advantages over PMTs, such as being able to operate at low voltages (we'll operate ours at +28.8V), being insensitive to magnetic fields, being robust, and having a compact form factor. They are also single photon detectors and have a peak responsivity near the peak emission from typical scintillating materials. This modern technology is what we use in the Desktop Muon Detectors.

SiPMs are constructed out of densely arranged *microcells* (see Fig. 4.1), each of which is a separate P-type and N-type semiconductor junction (P-N junction). When a P-N junction is first formed, the free electrons from the N-type semiconductor diffuse towards the P-type semiconductor and annihilate (similarly, the holes from the P-type diffuse into the N-type). Upon annihilation, the boundary region between the P and N-type semiconductors becomes an insulator known as the *depletion region*. When a photon travels through the depletion region, and deposits sufficient energy to a bound electron, the electron can be transported to the conduction band, thereby producing an electron-hole pair. If a potential difference is applied between the P-N junction, the electron will gain energy and at a certain point collide with other electrons

also transporting them into the conduction band. When the potential difference is sufficiently high ($> 5 \times 10^5$ V/cm), an electron avalanche (or cascade) can occur (a Geiger discharge) in which a single electron turns into a current of order millions of electrons. Once the flow of electrons is initiated, the silicon becomes conductive at which point a quenching resistor lowers the potential difference across the PN junction sufficiently to stop the electron cascade. In this way, each microcell acts as a photon triggered switch, which lets a small amount of current to briefly flow if it was struck by a photon. The sum of the total current flow is proportional to how many microcells were triggered and therefore is proportional to the incident photon-flux (when the number of triggering photon ll number of microcells).

The Desktop Muon Detector employs a single 6 mm by 6 mm On Semiconductor MicroFC 60035 C-Series [40] SiPM. These are most sensitive in the 450 nm range [41], which is a deep blue to purple. If a photon wavelength is too large (> 1000 nm), the absorption length in the silicon is also too large and the necessary size of the SiPM would be too bulky. If the photon wavelength is too short, it will not penetrate into the sensitive region of the SiPM, which is required for the detection.

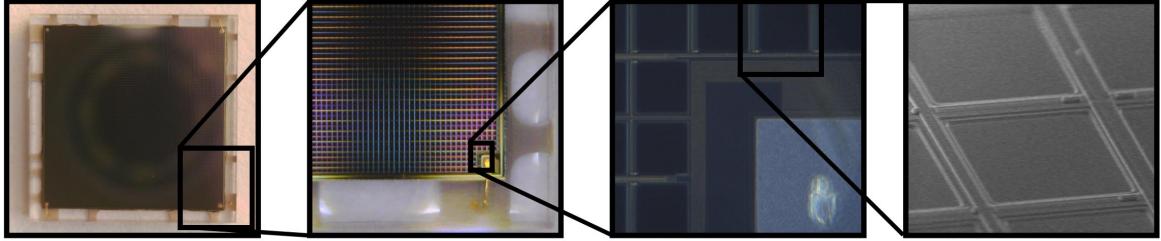


Figure 4.1: An image of a On Semiconductor MicroFC 60035 C-Series SiPM [40]. The SiPM has a length and width of 7.00 ± 0.05 mm and a thickness of 0.65 ± 0.05 mm. Each one of the tiled squares represents a single microcell, each of which operate independently in Geiger mode.

The applied potential difference is termed "bias voltage," and dictates both how large of a region is able to produce the avalanche (the depletion region), and the amount of energy gained by the electron-hole pair. The "breakdown voltage" defines the voltage at which the voltage gradient in the depletion region is large enough to create a Geiger discharge. This is typically between 24.2 and 24.7 V for the C-series SiPMs. If the bias voltage is increased beyond the breakdown voltage, the microcells will still operate in Geiger mode, however the electron cascade in the P-N junction will carry more energy, thus increasing the charge output (or gain) linearly. The difference in the bias voltage to the breakdown voltage is termed *over-voltage*. An over-voltage between 1.0 and 5.0 V is recommended. The Desktop Muon Detectors operate at an over-voltage of 4.0 V, which corresponds to a gain of roughly 4×10^6 [40].

Thermal fluctuations can produce electron-hole pairs, which mimic single photon events. For our SiPMs, this occurs at a rate of approximately 100 kHz per mm^2 , or at several MHz for the full SiPM. This can be undesirable for many applications that rely on distinguishing between small numbers of photons. The breakdown voltage required to initiate the electron cascade is temperature dependent; a lower temperature has a lower breakdown voltage.

4.2 Scintillators

Scintillators are simply a material that absorbs energy (through Coulomb interactions) and re-emits that energy in the form of electromagnetic radiation (scintillation light). Scintillators can come in different forms: for example, the scintillator could be grown as a crystal (referred to as an inorganic scintillator) with an added dopant, or the scintillator could have a fluorescing material embedded in a plastic (such as polystyrene or acrylic) or mixed into a liquid (like toluene or mineral oil) – these would be examples of organic scintillators. Inorganic scintillators are typically more expensive, however they can also have a higher density and emit more photons per unit energy deposited. This makes them more useful for calorimetry. Organic scintillators on the other hand are typically cheaper since the fluorescent material is suspended in a common, often low density material like a plastic, which enables them to be easily manufactured.

Scintillators are particularly useful since they emit light proportional to the energy deposited in the material, therefore, one will often see a metric of the quality of the scintillator expressing how many photons are emitted per absorbed MeV of energy (often called the scintillator efficiency). A common organic scintillator may have an efficiency of 10,000 photons/MeV. Another important quantity associated with scintillators photon emission profile of the scintillator, which determines what photon wavelengths are emitted after de-excitation. Scintillators must also be transparent to the scintillation light so that it can propagate to the photon detector. Plastic scintillators may have attenuation lengths on the order of 0.3 meters to 3 meters [42, 43], whereas liquid scintillators like Linear Alkyl Benzene (LAB) can have attenuation lengths of up to 25 meters [44]. Scintillators also have a very fast response and recovery (excitation and de-excitation of the fore-closing molecules), which happens on the order of nanoseconds for organic scintillators and hundreds of nano-seconds for inorganic scintillators.

Figure 4.2: A UV flash light illuminating scintillator.

The CosmicWatch Desktop Muon Detector was designed using an organic plastic scintillator, which consists of a polystyrene base (essentially just an inexpensive transparent plastic) mixed with a primary dopant of 1% by-weight of POP (2,5-diphenyloxazole) and 0.03% secondary dopant POPOP (1,4-bis[2-(5-phenyloxazolyl)]benzene) [45]¹. This plastic scintillator does not emit below 400 nm, and has a maximum emission around 420 nm (deep-purple light). This particular scintillator was developed by FermiLab for the Mi-NOS [47]/MINER ν A [48] experiments. We'll focus on a description of organic scintillators below, but a good description of inorganic scintillators can be found in Ref. [46].

¹Interestingly, PPO was one of the earliest compounds to be investigated as a scintilator solute by Hayes et. al (see Ref. [46]) from 1953-1958 and is still one of the most widely used

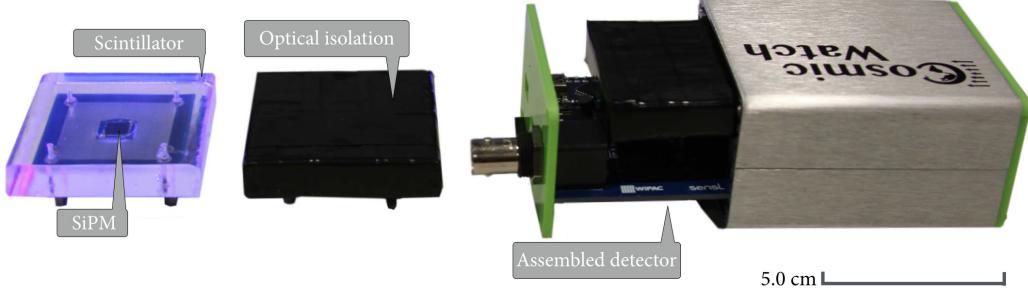


Figure 4.3: The components of the Desktop Muon Detector. Modified from Ref. [49].

The plastic scintillator requires three components:

- A transparent (in the visible spectrum) base that is used to suspend the fluorescent material. This can be some sort of plastic (like polystyrene) or transparent liquid like mineral oil.
- A primary fluorescing agent that is excited by the energy transfer from the incident charged particle. The de-excitation of the primary fluorescent material releases ultra-violet light. Ultra-violet light will not travel far in the base (order mm), before being absorbed.
- A secondary fluorescent agent that absorbs the UV light and converts it to the visible spectrum. The visible light then travels through the scintillator, internally reflecting off the walls until it is absorbed. Some of the visible light will hopefully strike the photon sensor that is coupled to the scintillator.

A simple illustration of the UV conversion is shown in Fig. 4.2. Here, we have illuminated several pieces of scintillator with a UV flash light. The UV is absorbed by the secondary fluorescent agent and re-emitted as a deep blue/purple light. Polystyrene-based scintillator has a density of roughly 1.032 g/cm^3 (similar to water) and a refractive index at standard atmosphere pressure of $n = 1.581$ [9].

The Desktop Muon Detector couples a SiPM (see Sec. 4.1) to a slab of scintillator via optical gel, which reduces the probability of a photon being reflected at the interface by matching the index of refraction from the scintillator to the housing of the SiPM, eliminating the $n = 1.0$ air gap. We also wrap the remaining surface in aluminum foil to reflect photons that escape the scintillator. It is then wrapped in 2-3 layers of black electrical tape to make the whole thing light-tight. The assembly of a detector is shown in Fig. 4.3.

Chapter 5

THE COSMICWATCH DETECTOR

5.1 Introduction to the Detector

As described in Ref. [49], the CosmicWatch Desktop Muon Detector consists of a $5 \times 5 \times 1$ cm³ slab of extruded plastic scintillator instrumented with a silicon photomultiplier (SiPM). When a charged particle passes through the scintillator and deposits energy, some of that energy is re-emitted isotropically along the particle track in the form of photons. Photons incident on the photosensitive area of the SiPM can induce a Geiger discharge in the SiPM microcells. When the microcells discharge, they produce a measurable current. A single photon can trigger a single microcell (neglecting second-order effects), whereas multiple photons may trigger multiple cells. The produced current is sent through a custom designed printed circuit board (PCB), which amplifies and shapes the signal such that it is suitable to be measured by an inexpensive micro-controller (in our case, the Teensy 4.0). The Teensy 4.0 measures the event time stamp and peak value on the analog-to-digital converter (ADC). The measured peak ADC value can then be converted back into a SiPM peak voltage, which is proportional to the number of photons incident on the SiPM. If the measured ADC value is above a software defined threshold, the micro-controller records the event data either to a microSD card or directly to a computer through a USB connection. For each event, the detector records the event number, event time and date, average measured 10-bit ADC value at two levels of pulse amplification, calculated SiPM peak voltage (which is proportional to the SiPM pulse charge), event dead-time, temperature, pressure, and whether a coincident detector also saw a pulse.

When the detector is plugged directly into a computer via a USB cable, the data can be recorded through a python-based program that replaces the date and local time stamp from the realtime clock on the Teensy 4.0 with the one from the computer. This time stamp tends to be more accurate (since the realtime clock drifts over time, like your wristwatch) with a precision of approximately 5 ms.

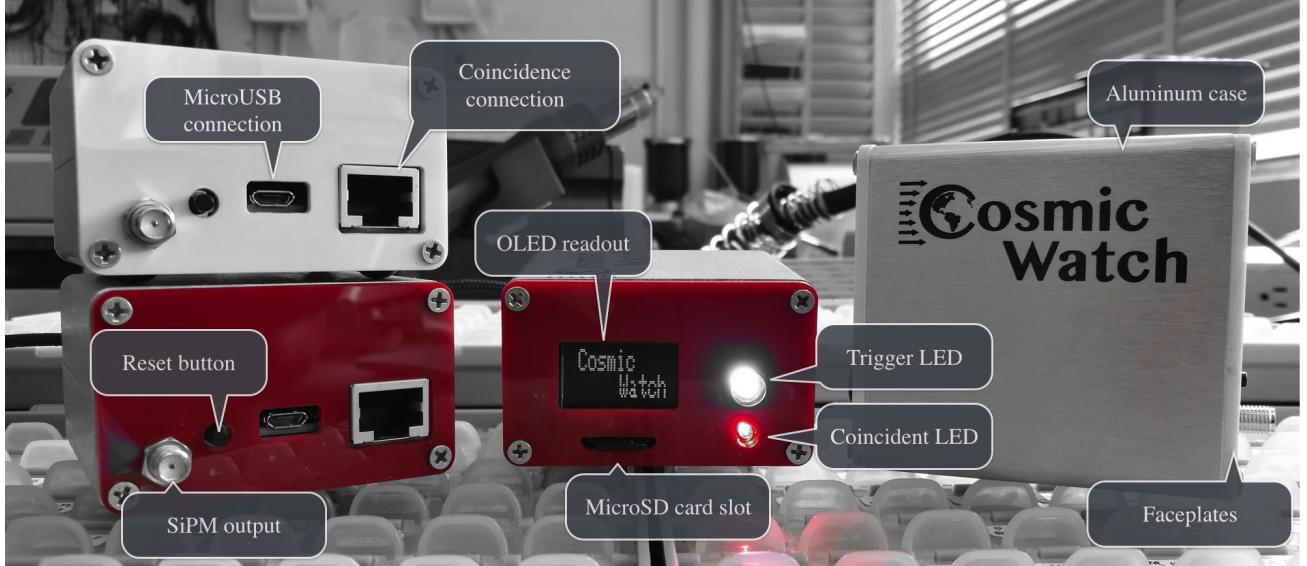


Figure 5.1: An array of Desktop Muon Detectors.

An array of complete detectors is shown in Fig. 5.1. The front of the detector includes a microSD card slot and built-in 0.96" OLED screen, which displays the count number since the last reset, total uptime, count rate (compensated for the detector dead-time), and an indication bar whose length is proportional to the calculated SiPM peak voltage of the last triggered event. Multiple detectors can also be linked together using a common ethernet cable to make a *coincidence measurement* (see Sec. 6.1). The top 5 mm LED triggers on all events, while the bottom 3 mm LED triggers only on coincident events.

The detector was measured to draw 0.6 W and can be powered through any USB connection (provided it supplies greater than 4.5 V). This includes the USB port on a computer, USB power bank, or power outlet USB. The total mass of the detector, including the aluminum case, is 200 g and the outer dimensions of the detector are 66.4 mm × 70.0 mm × 39.9 mm. Excluding the aluminum enclosure and end-plates, the detector has a mass of 100 g.

The backside of the detector includes a micro USB port for powering the detector, uploading new code to the Teensy 4.0, or recording data directly through a USB port; an ethernet connection used for connecting multiple detectors together to make a coincidence measurement; a reset button used for resetting the detector; and an SMA connection, which is connected directly to the SiPM output, and can be used for injecting signals into the PCBs for testing.

If you were to take apart the aluminum case (feel free), you can pull out the actual detector. There are a few other connections here that might be of use to you. The side SMA connector is used to output other analog information. The jumpers can be positioned to have this connection output the FAST signal from the SiPM, which has a faster rise and fall time, or any of the four test point on the circuit. There are also two test points near the microUSB connection which can be used to show the coincidence signal sent through the RJ45 connector.

5.2 Recording data

Data can be collected from the Desktop Muon Detector in many ways:

1. **(Recommended) Through a microSD card.** Each time that the detector is reset or powered on, a new file is created with a file name that counts sequentially upwards from the previous file, and an "M" or "C" indicating if the detector was in *master mode*, indicating that a coincidence detector was not used during startup, or *coincident mode*, indicating that a second detector was observed. Extracting data using the microSD card is least prone to error, and likely the simplest.
2. **Directly to a computer through a microUSB cable.** When the detector is plugged into a computer USB port, and the import_data.py is run using python 2.7, the user is prompted to supply the path and name of the file to where the data is to be stored. It will then begin recording the data in real time to the output file. Two notes: first, this method swaps the Teensy 4.0 date and local time with the one from the computer. And secondly, we could not reset the detector with the python code, so you will likely start recording data after a few events. This should not be a problem, but it means that you will have to shift the start time of the detector when you want to use the saved data.
3. **In realtime through the Teensyduino IDE.** When a detector is plugged into the computer, and the Teensyduino IDE is opened, the data can be seen accumulating in real-time in the serial monitor (top right magnifying glass). The data can be copied and pasted into a text editor for later analysis.
4. **Simply through the OLED screen.** The OLED readout accounts for the detector deadtime. Therefore, if you simply need to know the count rate, the OLED readout might be enough for your measurement. That being said, the data is still recorded to the microSD card.

The data is always saved in a simple to use .TXT file. Comments are always indicated with the "#" marker. Each row is a different event, and each column represents some different property of the measurement. The columns are tab-delimited. The detector actually starts out printing a set of diagnostics, which shouldn't be particularly relevant to this document. An example of the data is shown in Fig. 5.2.

The data shown in Fig. 5.2 is formatted into tab delimited columns, each of which is labeled in the header, and defined as:

- **Event:** The event number of the detector.
- **Time:** The time given by the realtime clock. If the data is taken directly onto the computer through a microUSB cable, then this date is replaced with the one from the

```
#####
### CosmicWatch: The Desktop Muon Detector
### Device ID: Tweety
### Launch time: 16:22:11 21/8/2020
### Questions? Email Spencer N. Axani (saxani@mit.edu)
### Event Time Date TimeStamp[ms] ADC1 ADC2 SiPM[mV] Temp[C] Pressure[Pa] DeadTime[ms] Coincident
#####
1 17:22:11 21/8/2020 697 222.27 5.99 7.92 0.00 0.00 5519 0
2 17:22:11 21/8/2020 753 131.77 5.64 5.69 0.00 0.00 4804 1
3 17:22:11 21/8/2020 978 486.97 60.39 13.94 0.00 0.00 776 0
4 17:22:12 21/8/2020 1563 717.02 262.26 23.18 0.00 0.00 6310 0
5 17:22:12 21/8/2020 1590 347.42 7.59 10.23 0.00 0.00 778 0
6 17:22:12 21/8/2020 1804 462.43 53.72 13.19 0.00 0.00 771 1
7 17:22:13 21/8/2020 2226 392.39 12.25 11.26 0.00 0.00 6311 0
8 17:22:13 21/8/2020 2260 488.40 51.75 13.99 0.00 0.00 778 0
9 17:22:13 21/8/2020 2636 219.94 5.78 7.87 0.00 0.00 776 0
```

Figure 5.2: Example data format. The header for the file is commented out with "#" and then there are nine events that follow. This data was recorded using the detector named "Tweety." The definitions of the columns are listed in the header, as well as a more descriptive description in the text below.

computer. The realtime clock time is updated everytime the software is re-uploaded to the detector.

- **Date:** The date given by the realtime clock. If the data is taken directly onto the computer through a microUSB cable, then this date is replaced with the one from the computer. The realtime clock date is updated everytime the software is re-uploaded to the detector.
- **TimeStamp [ms]:** The total elapsed time, measured in milliseconds. This is only accurate to roughly ± 1 min per day, and precise to the nearest millisecond.
- **ADC1 [0-1023]:** The high gain ADC measurement for the event. The Teensy 4.0 has a 10-bit ADC, meaning the values reported are from 0-1023 (2^{10} values). The ADC is referenced between ground and 3.3 V.
- **ADC2 [0-1023]:** The low gain ADC measurement for the event. The Teensy 4.0 has a 10-bit ADC, meaning the values reported are from 0-1023 (2^{10} values).
- **SiPM [mV]:** The calculated SiPM peak voltage. This a number calculated from the measured ADC value. Small SiPM pulses are primarily measured using the high-gain ADC channel (HGain), while large SiPM pulses primarily use the low-gain channel (LGain). It represents a number roughly proportional to the number of photons that triggered the SiPM.
- **Temp [°C]:** The measured temperature of the detector via the on-board BMP280 temperature sensor. Measured in degrees Celsius.
- **Pressure [Pa]:** The measured temperature of the detector via the on-board BMP280 pressure sensor. Measured in degrees Pascals.
- **Dead-time [ms]:** The dead-time since the previous event. This must be accounted for when making any rate measurement. Take the total time and subtract the total deadtime to get livetime.

- **Coincident:** If the detector observes a coincidence signal in the RJ45 connection (ethernet connection) within approximately 3 us, the event is considered coincident, and this value is set to "1," otherwise, it's "0." Events with a "1" in this column are overwhelmingly due to cosmic ray muon interactions.

5.3 Setting detectors in master and coincidence mode

Many measurements in this document rely on the detectors operating in *coincidence mode* (e.g. see Sec. 6.1). This mode allows us to improve the purity of the cosmic ray muon sample by rejecting events that likely came from interactions from radioactive backgrounds. Coincidence mode requires two detectors connected together using an Ethernet cable (I'll often refer to this as the *coincidence cable*). Once connected together, only one of the detectors requires power while the others are powered through one of the internal wires in the cable. To set the detectors into the coincident configuration, simply reset both detectors at the same time, while they are connected together. You could also simply unplug the power cable (microUSB) and plug it back in. They will acknowledge that a coincidence detector is seen by brightly illuminating both LEDs on the front panel for 1 second. An example of two detectors being setup in coincidence mode is shown in Fig. 5.4.

Master mode	Coincidence mode
Total counts: 12588	Total counts: 12588
Uptime: 1:10:11	Uptime: 1:10:11
M --	Rate: 3.140 +/- 0.21 Hz
Rate: 3.140 +/- 0.21 Hz	Rate: 0.170 +/- 0.01 Hz

Figure 5.3: The OLED output when set in Master (left) and Coincidence (right) mode. In Master mode, the third line displays an "M" and a bar indicating the pulse amplitude of the last event. In Coincidence mode, both the full trigger rate and coincidence rate are readout in the third and fourth line.

All events are always recorded (using the default settings in the Teensy software). However, coincident events will be designated in the data with a "1" in the coincident column. In Coincidence mode, the OLED screen will readout both the full trigger rate, and the coincidence rate.

Coincident events are likely to be due to a cosmic-ray muon, since the backgrounds and accidental coincidence are unlikely to trigger both detector simultaneously. A summary of why the purity of the cosmic-ray muon signal increases when in coincidence mode is below.

- Alpha particles will not penetrate a single detector (either the aluminum enclosure or even the black electrical tape) and therefore cannot trigger both the master and coincident detector at the same time.
- Beta particles can be significantly attenuated by the aluminum case, and have a significant chance of scattering, thus losing energy. It's unlikely that the beta particle will be able to deposit sufficient energy within the scintillator of the master, exit, then depositing sufficient energy in the coincidence detector.



Figure 5.4: Two detectors in coincidence mode, powered by a USB power bank. Everytime a detector triggers on an event, it sends a digital pulse to the other detector. If the other detector also triggers on a pulse during at the same time, the event is said to be coincident and is very likely to have been caused by a cosmic ray muon. Coincident events flash the coincidence LED (red LED on the front plate).

- Gamma rays can penetrate the aluminum enclosure and plastic scintillator, however they have a significant chance of Compton scattering, which will change the direction. If a gamma ray does interact with both detectors, this means that it likely Compton scattered off the scintillator slabs, lost sufficient energy to trigger the detector, then Compton scattered or photoelectrically absorbed in the second detector, also depositing sufficient energy to trigger the detector. This process will be rare, and represents a small part of the coincidence signal (an estimate of this rate is found in Sec. 9.5).
- Accidental coincidences from uncorrelated random events overlapping in the coincident trigger window is also a rather rare occurrence at the standard background rates. This will be elaborated on in Sec. 7.
- A typical minimum ionizing muon passing through the slab of scintillator will typically deposit more than 2 MeV of energy in the scintillator, without being deflected. If the muon passes through both scintillators, it will likely trigger both detectors simultaneously.

5.4 What Can Possibly Go Wrong?

When performing a measurement using the Desktop Muon Detectors, there are a few important things to keep in mind.

1. The Teensy 4.0 is a relatively slow device. It's quick compared to other commonly used microcontrolloers (such as the Arduino), however, there is inherant issues that arise from this.

A critical aspect of any rate measurement performed with the Desktop Muon Detector is that every command (such as adding two and two, or printing the detector information to the serial port or microSD card) takes time to perform. Accounting for the time that each command takes is important since the detector is unable to trigger on an event if the detector is busy. The term associated with the time that we are unable to make a measurement is *dead-time*. The dead-time is calculated in the detector software (Software.ino) code by measuring the time each command takes in microseconds. The dead-time between each event is readout in the output file. This must be subtracted from the up-time in order to accurately calculate the time that the detector was able to make a measurement; the result is called *live-time*. Dead-time is a common feature in all particle physics detectors, however for us it is particularly important due to the limited speed of the Teensy 4.0 and single core processor.

2. The orientation of the master and coincident detector when making a measurement will have different characteristics. When making a coincidence measurement, it is important to think about what is being measured, and how the orientation will effect the measurement. For example, if we are interested in the angular spectrum of cosmic-ray muons, we want to only accept muons coming from a small solid angle (a small area of the sky). Fig. 5.6 shows several possible configurations. The left side of Fig. 5.6, labeled as (a), shows two detectors spaced a few centimeters apart and connected with a coincidence cable. We see that only the muons that travel downwards through the blue area are able to trigger both detectors (think about what trajectories could trigger both detectors). This configuration would be useful for measuring the cosmic ray muon angular distribution. Fig. 5.6 (d) shows a configuration in which both detectors will trigger from down-going muons over a much larger solid angle. This (or (c)) configuration is typically used to extract strictly muon rate information since it will have the largest statistics since it is looking at a larger solid angle.

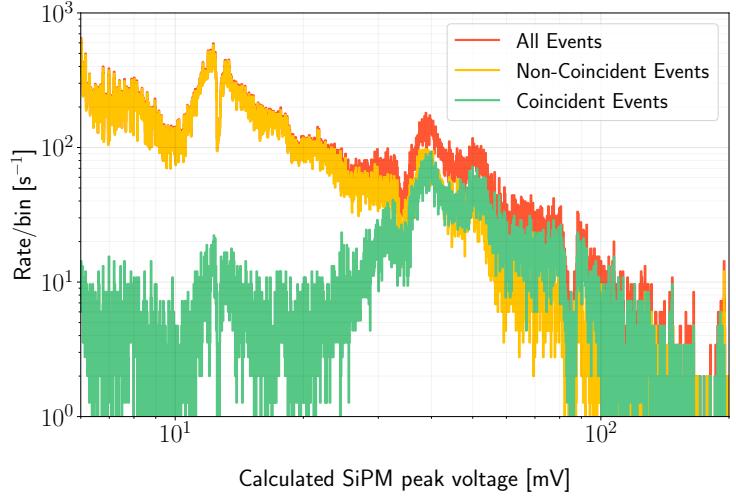


Figure 5.5: A coincidence measurement showing the calculated SiPM pulse amplitude for the coincident events and non-coincident events.

3. Two independent events could have triggered both the master and coincident detector accidentally if they happened to arrive within the coincidence time window. This becomes much more likely at higher count rates. If we assume that the individual count rates are determined by Poisson statistics and that the two detectors are independent, then the probability of observing n events when the mean number of events is expected to be μ , is given by the Poisson

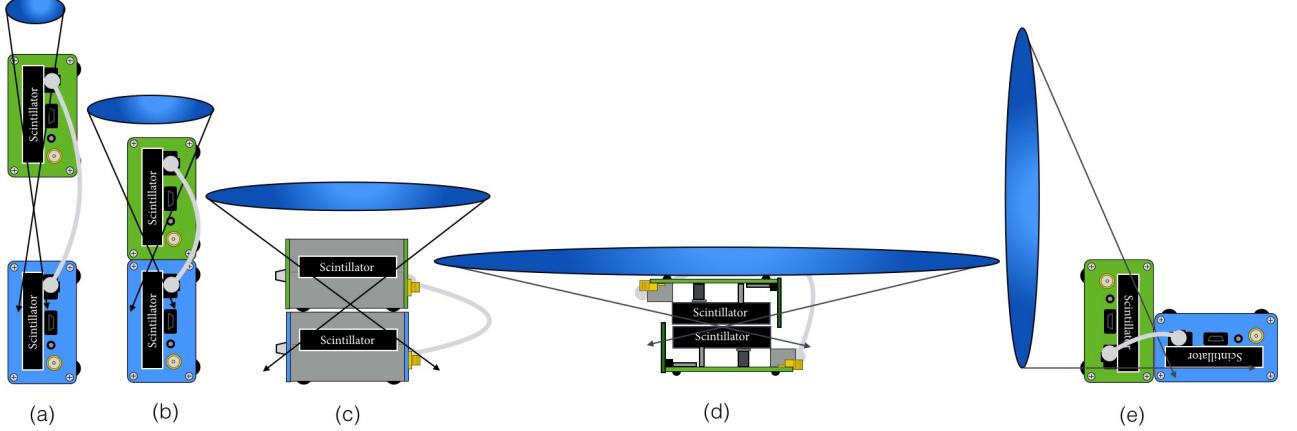


Figure 5.6: This figure shows several configurations for setting up the detectors in coincidence mode. Moving from left to right the solid angle subtended between detectors increases (illustrated as the blue ovals above the detectors) as well as the coincidence rate as measured by the coincident detector.

distribution:

$$f(n, \mu) = \frac{\mu^n e^{-\mu}}{n!} \quad (5.1)$$

Following Ref. [9], let the count rate on the two detectors be N_1 and N_2 , and the coincidence window (the time window in which an event will be labelled as a coincident event) be τ . The average count rate for both detectors at sea level is approximately 3 Hz, and given a coincident time window of roughly $3\mu\text{s}$, we expect on average 0.000009 events in the time window ($N_{1,2} \times \tau$). Therefore, the probability of observing zero event in the time window is $f(0, 9 \times 10^{-6})$ and the probability of one detector observing an event in the coincident time window is $1 - f(0, 9 \times 10^{-6})$. Since the second detector could have also triggered first, the rate of accidental triggers is: $R = 2N_1 N_2 \tau$. This corresponds to an accidental coincident rate for two detectors of 1 every 55,000 events, or roughly five events per day using two detectors, at sea level.

Chapter 6

PRE-LAB PREP WORK AND DAY 1 EXERCISES

Now that you understand the components of the detector, please do these preparatory exercises. You can find a video of the pre-lab work here:

<https://youtu.be/Js-1OOS7sVE>

6.1 Pre-Lab Prep Work: See a coincidence signal on the oscilloscope

Arrange your detectors in configuration (c) of Fig. 5.6 (also shown in Fig. 5.4) and start them up in coincidence mode. Using the two equal length SMA cables, attach the SMA output connector of one of the detectors to input 1 of the Red Pitaya oscilloscope, and the SMA output of the other detector to input 2 of the oscilloscope. Display both inputs on your laptop. Set up the oscilloscope so that both displays are on 20 mV/div, and time is on 500 ns/div. Leave both displays on zero offset so that the pulses will be aligned. Trigger on rising pulses on input 1 using NORMAL mode at a level of about 10 mV¹.

In your lab notebook, discuss some observations:

1. Input 1 will show pulses, triggering the detector at a rate of approximately 1 Hz. Sometimes input 2 will also show a “coincident” pulse. The two pulses will be “in time” (rising edges

¹The trigger rate increases as the threshold is reduced. At low enough thresholds, you may even begin triggering on electronic noise. Periodic noise in the mV-range is typically due to the digital communication on the printed circuit board, causing ground shifts observed in the analog electronics.

of the pulses line up) because the the cables are the same length. If, instead, the input 2 cable were much longer, how would the pulses on the scope shift relative to one another and the vertical trigger line (line with the blue T at the top) ? If, instead, the input 1 SMA cable were much longer, how would the pulses shift?

2. If you have set this up as suggested, then the input 1 SiPM pulseheight will range from about 1/2 of a vertical division (10 mV, or whatever is the trigger level) up to a few very high pulseheights that even go off scale. On input 2, where you sometimes see the coincident pulse, the pulseheights will tend to be 2 vertical divisions (40 mV) or larger. Why do you see many cases of small pulseheights on input 1, but not on input 2?
3. Most often, the two coincident pulses are similar in pulseheight. Why could that be? What property of the coincidence events causes the pairs of pulseheights to be substantially higher than the typical (minimum) pulse height? Why does that happen much less often than the minimum pulse height?

Change the oscilloscope trigger to SINGLE mode. The scope will now take a single trigger and stop rather than re-triggering. Hit RUN (top right corner) and the oscilloscope will take another trigger. Keep hitting RUN until you capture a nice coincidence signal. Take a photo! Copy it onto your log book as an example of a coincidence.

6.2 Day 1—Exercise 1: Pulse height distribution of your data

In this exercise, you will develop code to plot a histogram of the pulseheights. Depending on which experiment you choose, you may need this code later for your studies. The second YouTube video:

<https://youtu.be/Umbp92wuarY>

provides an example of how to use the data collected by the detector with an example iPython notebook found in the /Plotting/ directory of the GitHub repository located here:

<https://github.com/spenceraxani/CosmicWatch-Desktop-Muon-Detector-v3>

To download the repository, simply select on the green "Code" button near the top of the screen, then select "Download Zip." This repository actually contains all the information for the detector.

While you were looking at signals from the oscilloscope, the CosmicWatch detectors were automatically saving the data to the microSD card. Transfer the data from one of the detectors microSD card to your laptop. The file containing your data has a header that tells you what information was read out (see, the example data format in Fig. 5.2). For this exercise, you are interested in the column labeled SiPM voltage. That is a measure of the pulseheight. You are also interested in the last column, which has either a 0 or a 1. If it has a "0" then it means that only that detector fired. If it has a "1" then the other detector fired at the same time, within the set coincident time window.

Make a histogram of the number of counts of each pulseheight. Then, overlay a plot of the pulseheight for only those events that had a coincidence. What do you observe? Be sure to label axes and make a legend. Copy this plot into your log book and comment on how the information in this plot compares to your observations in the Pre-lab Prep.

6.3 Day 1–Exercise 2: Distribution of Rate in Hz

In this exercise, you set up code to measure the rate of pulses in Hertz, which is needed for nearly all experiments with CosmicWatch.

Loop over your events and extract the timestamp, which is the 4th column (python numbers from 0, so this is column 3). Write code that counts the number of pulses that occur in each 10 second interval. Make two plots of the full dataset:

1. the rate (in Hz) as a function of time,
2. the distribution of the rates that you measured.

Put these in your log book.

Now it is time to account for the detector deadtime. The second to last column in the data lists the deadtime, in microseconds, in which the detector was not actively taking data since the previous measurement. While your binned data contains 10 seconds of total time, the detector livetime was the total time minus the deadtime. The actual rate of the incoming pulses is defined as the number of counts per unit livetime. How does including deadtime change the above two plots?

It is very important to account for the detector deadtime in all rate measurements. While the effect is rather small for typical background rates, it becomes extremely significant when considering high count rates. For example, if you would like to see how the deadtime affects the measurement made using the FiestaWare plate described in the pre-lab video, try using your code with the example data found in /Plotting/Example_data/Fiestaware_plate.txt.

Repeat the rate measurements above, but require a coincidence. In this case, you may want to count the number of pulses that occur in a larger interval, say 60 s. Remember to convert the rate to Hz.

6.4 Before Day 2: Send us the cosmic-ray muon rate at various locations

Using the coincidence data from Exercise 2, find the mean rate in Hz and the error on the mean (accounting for the deadtime). Send the mean, the error, your location, and the altitude (look it up on Google) to your TA. Before Day 2, we will add the data from the class to a set we collected in 2018, shown in Fig. 6.1. Yeah, Spencer was in Antarctica! How fun is that? Spencer was not on the high-altitude balloon; we strapped the detectors on and sent them up (see Fig. 9.2).

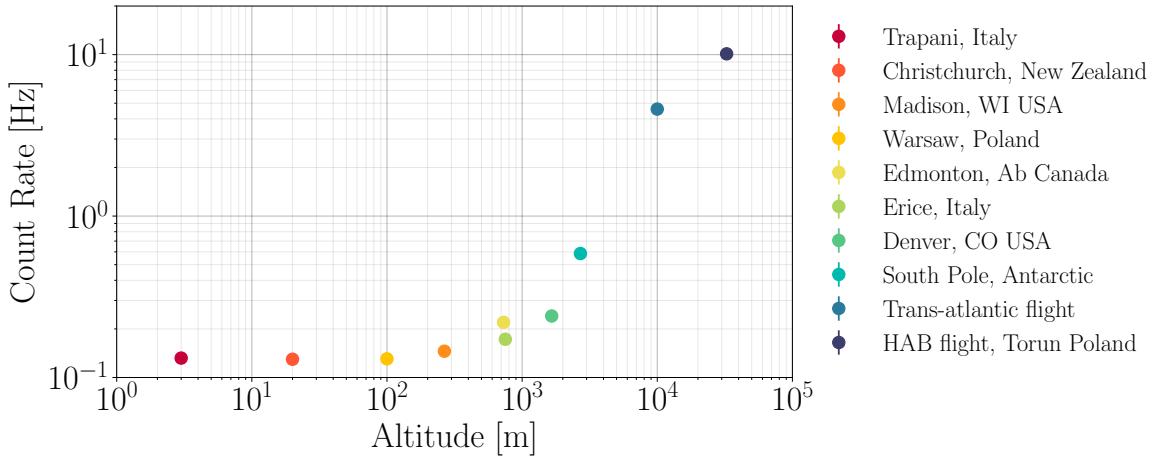


Figure 6.1: The coincident detector count rate at various locations throughout the world. The trans-Atlantic flight rate was measured at 30,000ft, while the high-altitude balloon data was simply the count-rate at the Pfotzer maximum described in Section 9.2. The statistical error bars of each point are smaller than the marker size.

Once the updated version for Fig. 6.1 is available, add it into your log book and discuss it. You can see some interesting effects. For example, the measurement in Sherwood Park, Alberta Canada, was made in a wooden house, while the measurement in Erice, Sicily Italy, was made in a large stone monastery. Although both measurements are performed at approximately the same the same altitude, there is a 27% measured rate difference between the two. What might cause this?

Final take away from this exercise: always travel with a CosmicWatch!

Chapter 7

TESTED MEASUREMENTS YOU CAN CHOOSE FROM

You will design your own experiment for this laboratory. The list below are some measurements that can be done at home. Feel free to choose from this list or develop your own idea. Inspiration may come from measurements described in Sec. 9.

There may be cases where you would like to compare your own data to data taken by others in the class. This may not be done without permission, so please contact your instructor.

7.1 Cosmic-ray muon angular distribution

This measurement illustrates the cosmic-ray muon angular dependence measured near sea level (it was performed in Madison WI, at 266 m above sea level) and was previously described in Ref. [49]. Here, two detectors were set to coincidence mode and placed side-by-side as in configuration (a) of Fig. 5.6, spaced 52 mm apart, inside their aluminium enclosures. The distance was chosen such that we gain sufficient statistics throughout a single day. If the detectors are placed too far apart, the count rate drops significantly and accidental coincidences can dominate the signal.

The angle of the detectors was determined by securing the detectors to a 100 cm long rectangular bar and then positioning the bar against a wall at a known height (see Fig. 7.1 (right)). It is important when making this measurement that the angle of the detectors are accurately measured. Fig. 7.1 (left) shows the measured relative rate as a function of zenith angle (with zero radians representing vertical). Each data point represents approximately 10 hours of data and the rate uncertainties are statistical. The horizontal (x-axis) uncertainties represent the calculated opening angle of the two detectors when spaced 52 mm apart. The measurement at

$\theta = \pi/2$ is divided by 2, since at this angle it accepts cosmic-ray muons from both directions, whereas all the other angles only accept down going muons.

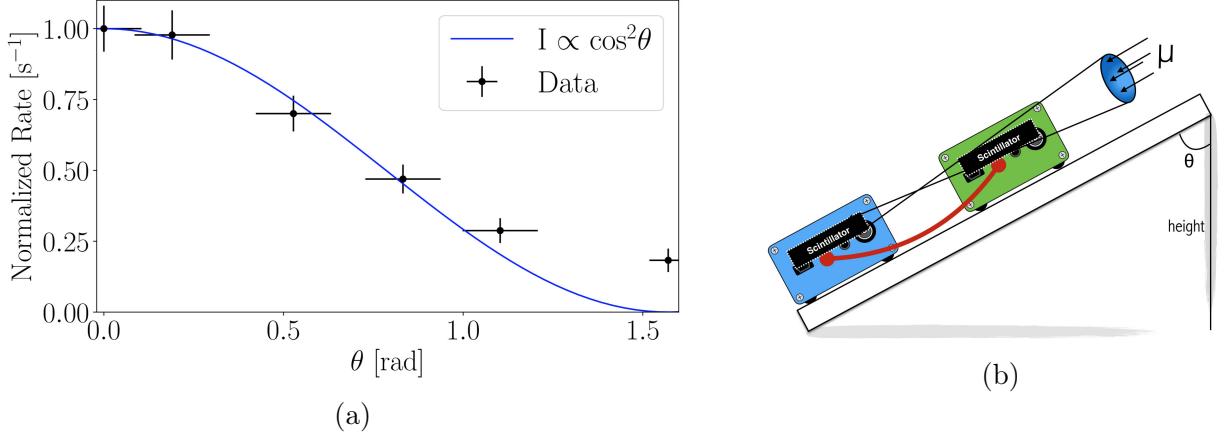


Figure 7.1: The measured cosmic-ray muon angular distribution measured by the coincident detector. The prediction by the PDG is shown in solid blue. From Ref. [49].

As indicated in Sec. 2.1.2, the angular cosmic-ray muon dependence at sea level should follow a cosine squared dependence. The overall measured shape of the distribution is found to agree relatively well with the cosine squared prediction, however it is shown that the rate does not fall completely to zero in the horizontal configuration ($\theta = \pi/2$). This appears to be related to showers of particles developing in the roof above the detector, producing a spray of high energy electrons or photons that can trigger both detectors. This can be investigated further by shielding one of the detectors with a dense material (such as lead or iron) and performing the horizontal measurement a second time. Doing the measurement near a basement wall can also add shielding.

7.2 Cosmic-ray muon attenuation in a Subway station

While visiting the Research Center for Neutrino Science (RCNS) at Tohoku University, I noticed that the Aobayama subway station was extremely deep underground and decided to see if there was a noticeable change in the muon rate compared to the office where I was working. After spending approximately 30 mins in the station, I decided to travel to another station (Aoba-Dori Ichibancho station) a bit closer to my hotel to see if the rate changed. And then took a final measurement at ground level.

The muon attenuation was found to be surprisingly pronounced, as shown in Fig. 7.2. I would anticipate this measurement also working rather well in the deep subway stations in Boston, for example at Porter Square. We could take this a step further and determine the rate as a function of depth, enabling the detectors to calculate overburden.

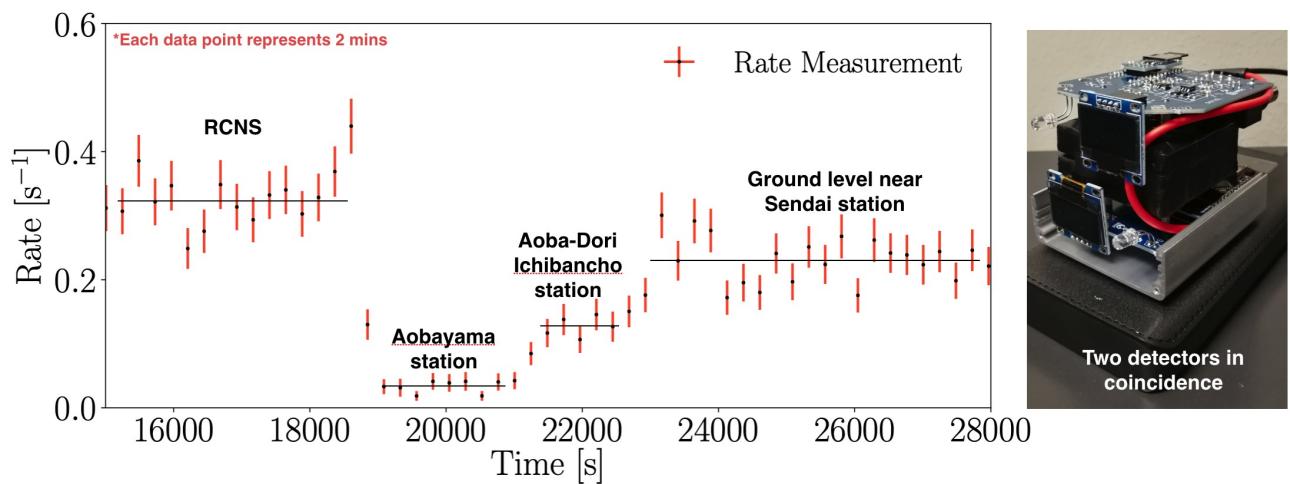


Figure 7.2: Subway in Japan.

There are likely other underground areas in your area from which this measurement can be made. While the cosmic ray muon flux per unit solid angle is from directly above, the largest flux, per unit zenith angle, is actually near 70 degrees. To get a good measurement, you must not only consider the vertical overburden, but also the overburden away from vertical.

7.3 Electromagnetic component of cosmic-ray showers

We've discussed in Sec. 2.1 that at sea level, there is a flux of electrons and positrons showering down on the Earth. The majority of this flux has energies less than 100 MeV, with a very small contribution above 1 GeV. We then discussed in Sec. 3.2 that this electromagnetic component is not particularly penetrating and can be significantly attenuated with a few tens of cm of concrete (the radiation length in concrete found in Table 3.1 is 10.7 cm). We can test this claim by measuring the coincidence rate on top of a building, then again several floors lower.

We expect a minimal attenuation of the cosmic ray muons, however a significant attenuation of the electromagnetic component. A typical building may have 15 cm of concrete between each floor. If we measure the vertical muon rate on the top floor, then again 5 floors lower, we expect the energy of the electromagnetic component to be attenuated by 7 radiation lengths (a factor of 1000).

Two detectors were placed in coincidence mode one-on-top of the other in configuration (c) of Fig. 5.6. They were placed on the roof-top of the 10-floor WiPAC building for 24-hours inside a plastic bag to protect them against the weather. After the measurement, the same two detectors were placed on the 5th floor, and the measurement was repeated. The results of the measurement are shown in Fig. 7.3.

We see that the rate on top of the building is 27% higher than that of the measurement made on the 5th floor. The attenuation of this component is thought to be mostly due to the elimination of the electromagnetic component. In Sec. 2.1, we explain that approximately 33% of the total flux from cosmic ionizing radiation at sea level comes from the electromagnetic component.

We have always thought it would be fun to do this study by lowering CosmicWatch into a lake. In that case, the level of shielding is very well defined and in physics experiments, often actually described in terms of meter water equivalent (m.w.e.). If you decide to try this, please be sure that your detectors are in a water-tight container. A 30% decrease in the coincident rate is expected within the first few feet as the electromagnetic component drops off, then a slower turn off of the muonic component as you go deeper into the water (ideally 10s of meters).

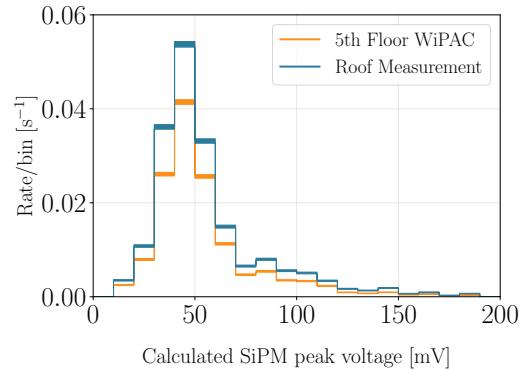


Figure 7.3: The calculated SiPM peak voltage on the roof of the WiPAC building compared to on the 5th floor.

7.4 Reducing radioactive backgrounds

There is a lot of natural radioactivity in the environment. The CosmicWatch detector can pick this up. You saw how orange FiestaWare plates are radioactive in the intro video about the detector. Potassium chloride (salt substitute) will also make a signal (you need about 11 oz for this). You may find the granite countertops in your home are radioactive too. Check for this!

It can be difficult to shield against the natural radioactivity found in the environment. Typically at sea level, the largest component of our signal comes from radioactive backgrounds and therefore it would be useful to think about how we can minimize this to improve the purity of the cosmic-ray muon component of the signal (beyond setting detectors into coincidence mode).

Dense material placed around the detector can attenuate the incoming flux from the background radiation. Ideally, we would choose a material that itself is radio-pure, however we can illustrate the effect using bricks of lead. Six lead ingots (each measuring 2"x4"x8") were positioned in such a way to provide 4π coverage around a single detector. The detector recorded data directly to the microSD card throughout a full day. It was then placed on a workbench in the same room (far away from the lead), to measure the background spectrum for another full day. The resulting calculated SiPM peak voltage for the two measurements is shown in Fig. 7.4. As expected, this is shown to significantly reduce the events that contribute to the low SiPM peak voltage region, which is dominated by the radioactive background.

You probably do not have lead bricks at home. Instead, you can work with the junior lab staff to design a lead cave to study shielding. You can also try other materials. You can dig a hole and bury your detectors. You can even try water as long as you are careful to keep the detectors in a sealed container! You can measure effectiveness of different types of materials of the same thickness, and consider why some materials are more effective than others. You can also look at the effect of different thicknesses of shielding.

You can also study the reduction in rate with distance from a radioactive gamma-ray source. If you contact the junior lab instructors, they can set up any arrangement you request for data taking, working with you on zoom. They can then send you the data. You should be able to prove the $1/r^2$ fall off associated with “ALARA” – the idea you learned in the radiation safety lecture.

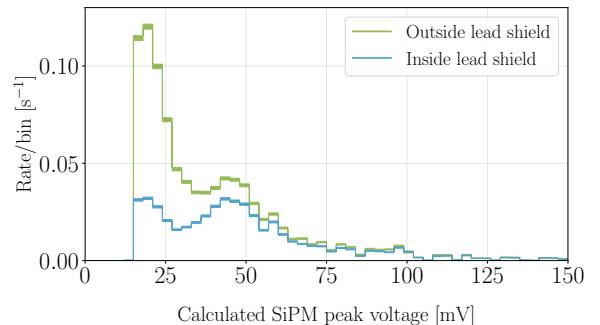


Figure 7.4: The effect of surrounding a single detector in a lead enclosure.

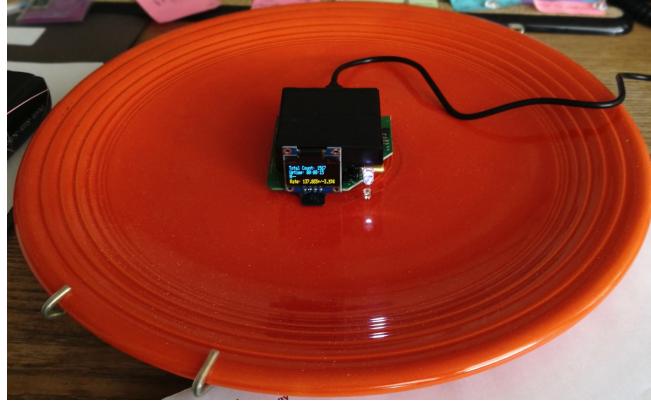


Figure 7.5: A CosmicWatch detector measuring the rate sitting on top of a radioactive Fiestaware plate. See YouTube video.... Standing far from the plate is an example of ALARA!

7.5 Correlation between the cosmic-ray muon rate and atmospheric pressure

A correlation exists between the atmospheric pressure and the cosmic-ray muon rate. It can be expressed as follows:

$$\frac{\Delta I}{\bar{I}} = \beta \Delta P, \quad (7.1)$$

where I represents cosmic-ray muon intensity, ΔP is the measured atmospheric pressure compared to the average pressure, and β is the barometric coefficient [50]. This correlation is actually the result of several processes outlined in Sec. 2.1.2. The barometric coefficient represents the percent change in detector count rate per hPa change in atmospheric pressure.

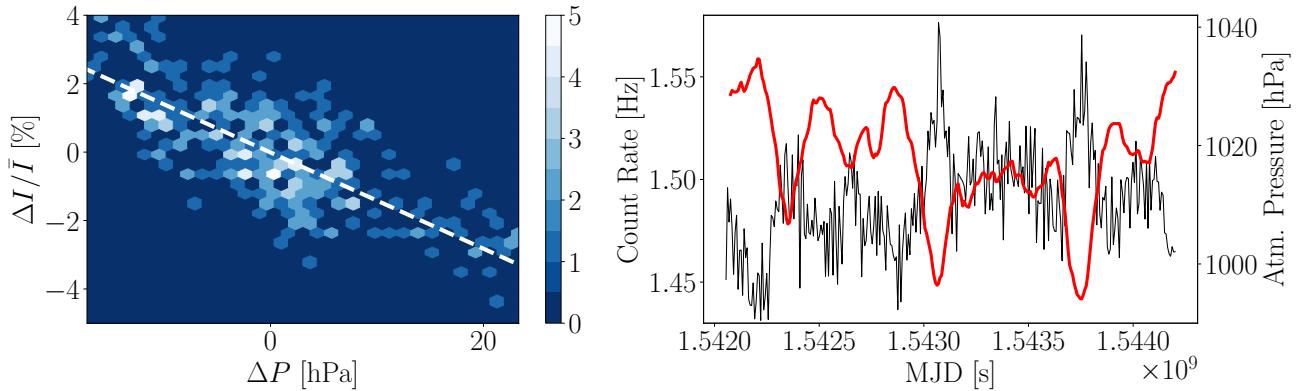


Figure 7.6: Left: The correlation between the atmospheric pressure and the detector count rate. Right: The detector count rate in blue and the atmospheric pressure in red as a function of modified Julian date (MJD). Both plots show the data binned with a bin size of two hours.

For this measurement, we decided to use an array of master-coincident detectors to improve the statistics. Five pairs of detectors were each connected in configuration (c) of Fig. 5.6, and left to record the cosmic-ray muon rate over the course of 24.8 days (from 13:35 hrs Nov. 12th to 09:00 hrs Dec. 7th 2018.). Data was recorded directly to a microSD card and powered through an 8-way USB hub powered through a wall outlet. Since the CosmicWatch does not keep accurate time over long time-scales, we assume the time drifts linearly with time, and scale the uptime of all detectors such that they are all the same. This is not ideal, and we would recommend recording the data directly to a computer through the import_data.py script in the future to get an accurate time stamp on all events. However, this is more prone to losing data, since you would require a computer that runs continuously for a significant amount of time. The array of detectors was placed in the 4th floor WiPAC lab in Madison, WI. Atmospheric pressure for Madison was found in Ref. [51]. However, the v3 version of the detector includes a built-in pressure and temperature sensor.

Fig. 7.6 (left) shows the correlation between the detector count rate and the atmospheric pressure. The calculated correlation is shown in the dotted white line. A least squares fit yielded a barometric coefficient of $-0.141 \pm 0.007\text{ \%}/\text{hPa}$, in agreement with Ref. [52].

This may be difficult to do with only two detectors since the statistics will be lower, but it will be interesting to see! You can increase the amount of data you have if you ask a friend to also take data at their location, and you add that to yours.

One of our latest adventures in HurricaneWatch. A meteorology student who is a hurricane-chaser is taking our detectors into the storms she studies. The atmospheric pressure during a hurricane can drop by over 100 hPa, which we hope to be able to see. We will be excited to get this data set!

7.6 A true random number generator?

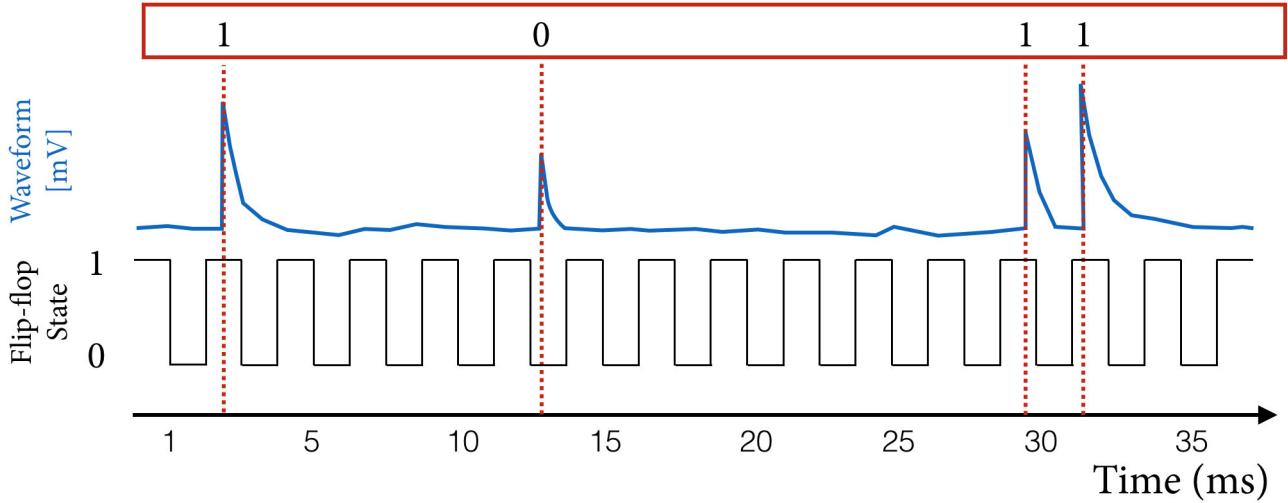


Figure 7.7: An illustrative diagram of the principle behind the random number generator. The x-axis indicates the time measure in milliseconds. Above this, we have the toggle flip-flop, which changes state every millisecond, then an illustrative waveform shown in blue. If the event triggered the detector while in the even state of the flip flop, we assign it a one, otherwise, we assign it a zero. This is shown in the red box at the top of the diagram.

Many applications require random numbers. Often, a random number may be generated through some algorithm, but this therefore becomes deterministic since if the user knew the algorithm and starting conditions, they could determine the output. A random number generator would ideally be generated from a truly random process, such as the arrival times of cosmic-ray muons or the radioactive decay of an element. The sum of two random processes will also be random, such as the signal from the radioactive backgrounds in the detector along with the cosmic-ray muon signal. For this measurement, I'm following the description found in Ref. [9].

Any number can be expressed in terms of a sequence of ones and zeros, this is known as binary. An N -length sequence is able to represent a number from zero to $2^N - 1$ (corresponding to 2^N different values). For example, the 4-bit binary sequence "1011", corresponds to $1 \times (4^2) + 0 \times (3^2) + 1 \times (2^2) + 1 \times (1^2) = 8 + 0 + 2 + 1 = 11$.

We can convert the time stamp of a radioactive decay trigger into a "1" or a "0" using a *toggle flip-flop*. The toggle flip-flop is simply a state that changes from one to zero periodically (we will use a frequency of 1 kHz). If a particle passes through the scintillator during an even time stamp (as measured in milliseconds), we assign it a "1"; if it passes through an odd time stamp, we assign it a "0". After N triggers, we can build an N -bit random number. This is schematically illustrated in Fig. 7.7.

For this measurement, we use data taken from a 20-day background lab measurement. A single

detector was used. After triggering on t total events, we can build t/N N-bit random numbers. To illustrate this, if we chose to generate 8-bit random numbers (from 0 to 255), we can plot the number of each occurrence. This is shown in Fig. 7.8.

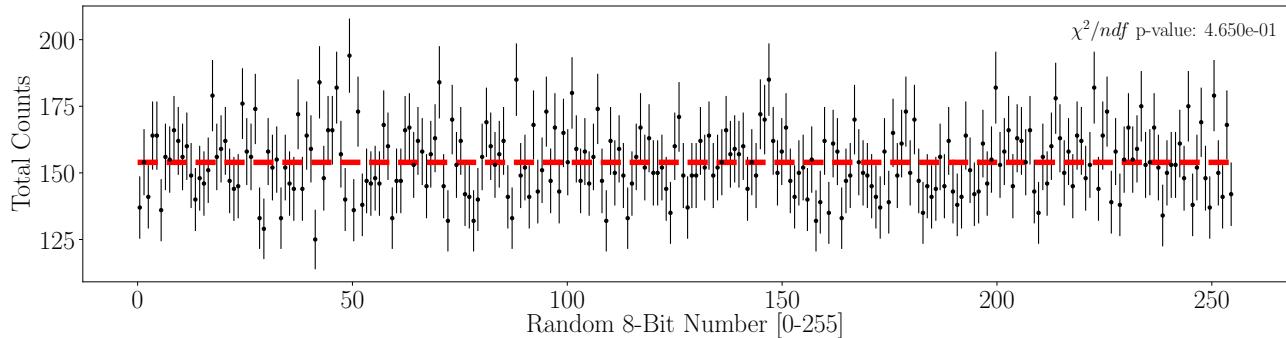


Figure 7.8: The number of occurrences of the generated numbers from 0-255. The reduced χ^2 p-value assuming that the distribution should be randomly distributed about the average number of occurrences is given at the top right of the figure.

Fig. 7.8 shows the random nature of the triggers. Each number should be equally probable to occur. The reduced χ^2 indicates a p-value for this assumption shown in the top right of the plot.

In your case, you may want to take data from a natural radioactive source instead. You can purchase an 11 oz or larger container of salt substitute (potassium chloride) for this study. It will provide a count rate that will be higher than background when near the detector.

There are several ways that the random number generator can become biased. First, let's think about extreme scenarios. Suppose that the trigger rate is roughly 1 Hz, and the flip-flop state only changes every 10 seconds. The first roughly ten triggers would give all ones, then the next roughly triggers would give all zeroes. The first numbers would be biased high, while the later numbers would be biased low. With this extreme example, we see that we need the toggle flip-flop to be changing states at a much higher rate than the trigger rate. Secondly, suppose that the microcontroller does not produce equal numbers of even and odd time stamps due to the internal configuration. We found this to be the case when using the microseconds() Teensy/Arduino function, that is, it only reported the time stamp to the nearest even microsecond. Study and quantify the randomness of data you take.



Figure 7.9: Background rate measurement with (left) and without (right) an 11 oz container of potassium chloride. The rates are measured as 7.2 ± 0.1 Hz and 3.94 ± 0.08 Hz, respectively.

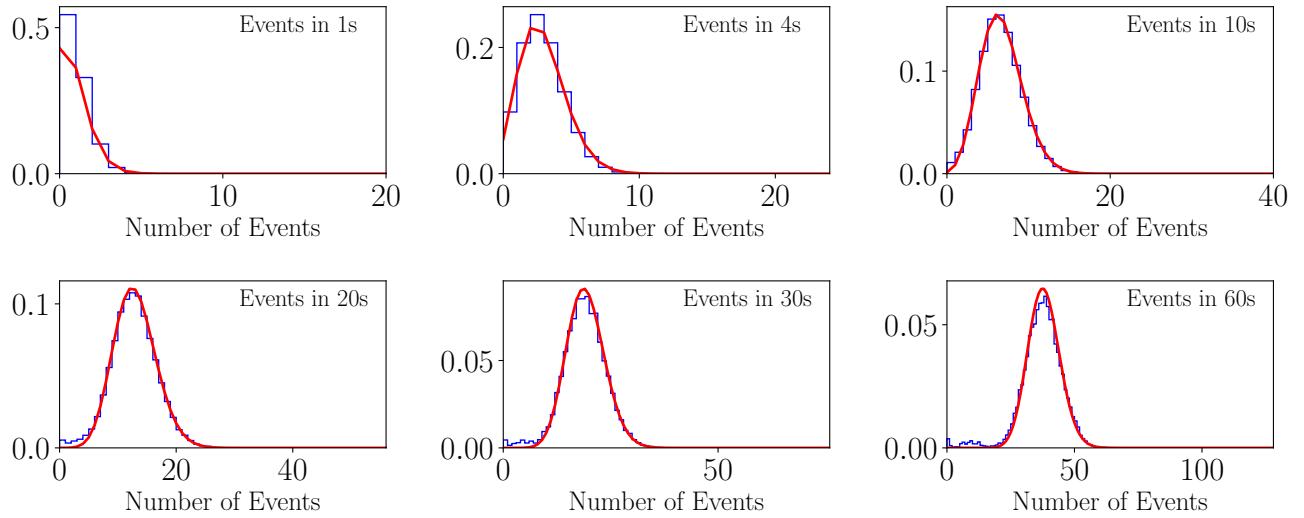


Figure 7.10: The distribution of number of events per unit time. Each cell represents a different amount of time in which the data is binned (labelled on the top right of each cell). The red line is the expected Poissonian distribution from Eq. 5.4.

Chapter 8

UNTESTED MEASUREMENTS FOR THE ADVENTUROUS

This chapter describes a few measurements that we hope to add to the list in the near future. If you choose one of these measurements, you are going into relatively uncharted territory. We point out the difficulties in the measurements (and in the case of the solar flare you have to have lucky timing!) So: Beware! However, if you are able to perform any of these measurements, they might end up in next year's curriculum! Good Luck!

8.1 Measuring the air shower size

Air showers were discovered using coincidence counters, very similar to the idea used in CosmicWatch. An airshower consists of multiple particles that hit at the same time. Showers can be produced when the muon interacted as it traverse material above the detector, like the roof of the building, or, in Fig. 8.1, a sheet of lead.

The idea of this measurement is to look for coincident events between two sets of coincidence detectors as a function of lateral distance between the two sets. The problem with this measurement is that it may require 3 to 4 detectors, as shown in Fig. 8.1. Two sets of coincidence detectors (4 total detectors) would have a lower accidental coincidence rate than using 3 detectors (i.e. emit detector (C)). The detectors aren't setup to record triple coincidences using the microSD card. However, if all detectors are plugged into a single computer (through a USB hub), the data from each detector can all be saved to a single file using the import_data.py script. The time stamp taken from the computer can then be assessed to look for coincidence events.

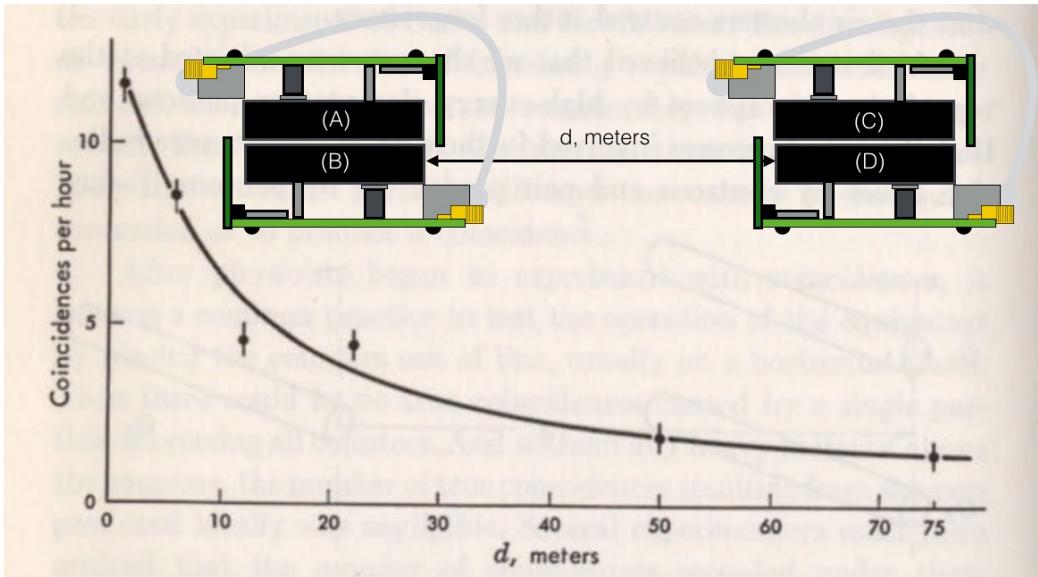


Figure 8.1: Recommended configuration for measuring air shower size. Data from Ref. [3].

8.2 The velocity of cosmic-ray muons

In order for cosmic ray muons to penetrate the atmosphere and make it to sea level, they need to be traveling at relativistic speeds. Since the speed of light is a foot per nanosecond (approximately), if a muon passes through two widely separated scintillation counters, then you can measure the time delay and extract the velocity. We have a set-up for this in the lab at MIT, and it is described here: <http://web.mit.edu/8.13/www/14.shtml>.

While the precision in the triggering electronics of the CosmicWatch detectors is order microseconds (meaning you would need 1000s of feet between two CosmicWatch detectors to be able to observe travel time differences), there potentially is a way to make this measurement using the Red Pitaya. If you open up the detector (four screws on the front, four screws and the nut on the SMA connector on the back), you can gain access to the fast output of the SiPM (you need to position the jumper near the LEDs to the FAST connection). This signal has a faster rise time than the output on the back SMA connector. Placing two detectors several feet apart, and digitizing both fast outputs using the Red Pitaya may show an observable shift.

This measurement is likely difficult, since the digitization is limited to 125 MSPS. Therefore, in order to see a shift of one ADC sample, you would require the detectors to be approximately 8 feet apart. At these distances, the coincident rate is very low and the accidental coincidence rate needs to be carefully investigated.

Another issue to consider before making this measurement is the method that you plan to use to record the data. Ideally, you would like the Red Pitaya to save the waveform of all observed coincident events so that you can analyse the data offline later. Assessing whether or not

your setup will work, you could begin by using two different length SMA cables, say 6" and 8,' the signal travel time difference between these two cables should be observable in a coincident event (with the detectors placed close together), and will give you an idea of what separation of detectors is needed to observe a difference given the limited sampling speed of the Red Pitaya. During the actual measurement, keep in mind that the SMA cable lengths need to be the same length.

8.3 Solar flares and the Forbush decrease

At higher altitudes and latitudes, where lower energy primary cosmic ray particles can be observed, the cosmic radiation can be observed to correlate with solar activity. Solar flares are also associated with coronal mass ejections can be responsible for significant temporary count rate increases. This is more noticeable at higher altitudes where the Earth's magnetic field doesn't provide the same protection to GeV cosmic ray particles as at the equator.

The Geostationary Operational Environmental Satellite (GOES) operated by NASA and the National Oceanic and Atmospheric Administration, records the incident particle flux from the sun (proton, electron, XRay...). The data is publicly available and could be used to correlate transient solar events with the cosmic ray muons. Solar flares use the letters A, B, C, M or X to classify the power per unit area of an event. A numerical suffix ranging from 0 to 9 indicates the strength of the event – that is, an X2 flare has twice the strength of an X1 flare. Each letter defines an order of magnitude different between the classes, with X being the most powerful flare classification. We may only see a few X-class solar flares per year, but there is a chance that your detector is running during an event!

The real-time clock on the CosmicWatch detectors is updated each time the software is re-uploaded to the Teensy 4.0. After that, the clock may drift resulting in an inaccurate local time stamp. However, this may not be a large issue since solar flares can be rather long. If this is thought to be an issue, then data can be recorded directly to the computer, where the

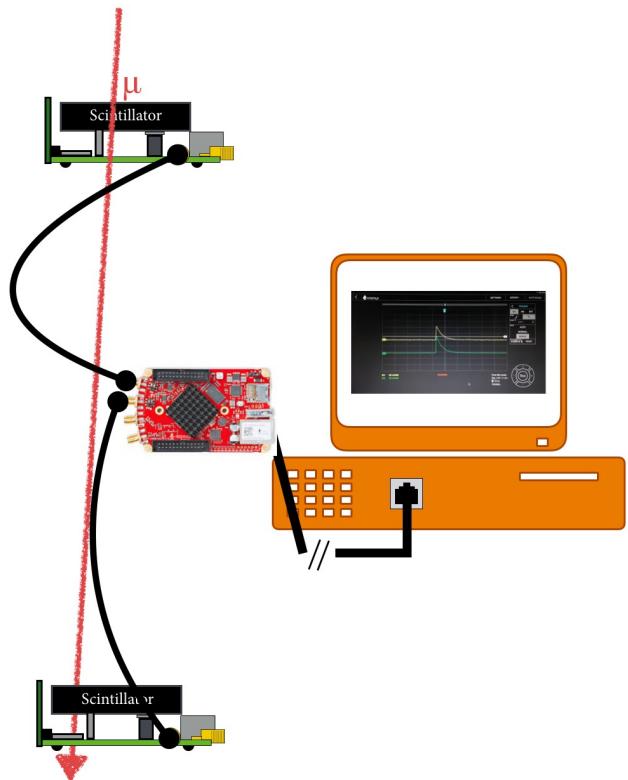


Figure 8.2: An example setup to measure the velocity of cosmic-ray muons.

computer's time stamp is replaced with the RTC time stamp – thus providing a much more accurate reading.

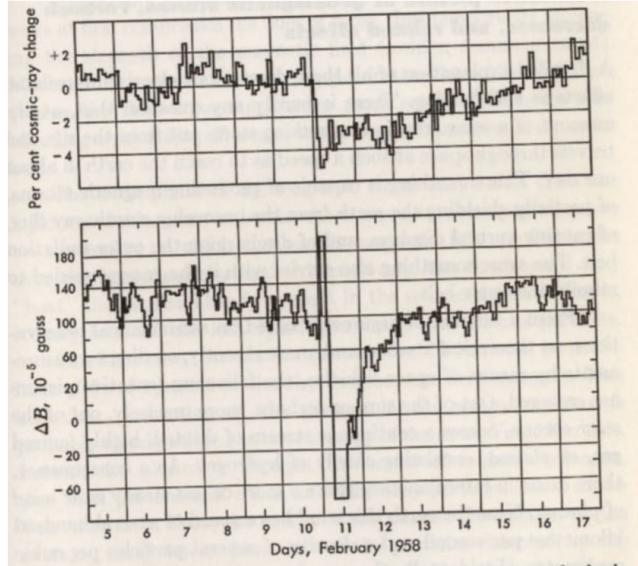


Fig. 14-2 Typical Forbush decrease. The lower curve represents the horizontal component ΔB of the magnetic field measured, from an arbitrary zero, in units of 10^{-5} gauss. The upper curve shows the corresponding variations of cosmic-ray intensity (measured by a neutron monitor in Chicago) as a percentage of the normal intensity. Shown on the horizontal axis is the universal time, in days. (From F. Bachelet, P. Balata, A. M. Conforto, and G. Marini, *Il Nuovo Cimento*, vol. 16, p. 292, 1960.)

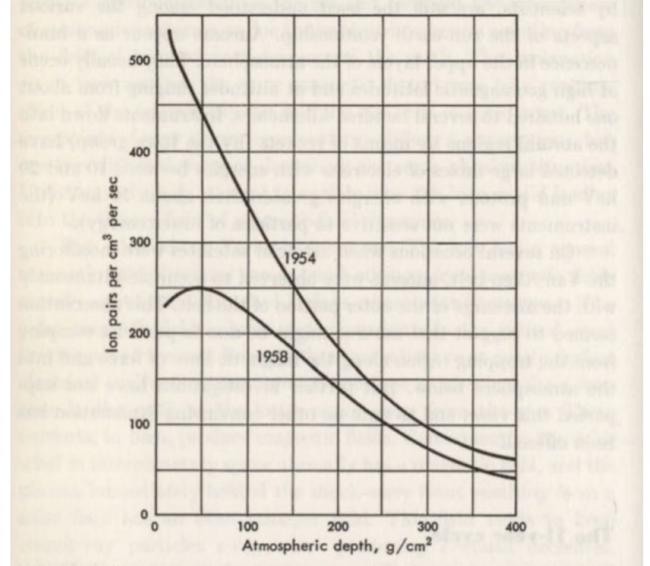


Fig. 14-3 Intensity of cosmic rays as a function of atmospheric depth, measured at a geomagnetic latitude of 88° N at the time of minimum solar activity in 1954 and at the time of maximum solar activity in 1958. The instrument used was an electroscope; the vertical scale gives the number of ion pairs per second produced by cosmic rays in 1 cm^3 of air at standard temperature and pressure. The horizontal scale gives the atmospheric depth, in grams per square centimeter. (From H. V. Neher, *Nature*, vol. 184, p. 423, 1959.)

Figure 8.3: The Forbush decrease (left) and solar maximum/minimum cosmic-ray flux (right). This figure comes from Ref. [3].

Roughly one day after a solar eruption, the earth experiences a solar storm. A short increase in the magnetic field is observed, followed by a much stronger decrease that continues for a few hours. Over the course of a few days, the magnetic field slowly recovers to its original strength. This effect is known as the “Forbush decrease.” The observed cosmic ray radiation intensity begins to first decrease

Since the Earth’s magnetic field plays such an important part in shielding us from cosmic radiation, this too should be observable. As shown in Fig. 8.3, the change in rate is percent-level. Percent level statistical changes should be observable over the course of a day, using configuration (d) of Fig. 5.6. However, recall that the atmospheric condition may also impact the measurement, so you may want to also consider the barometric pressure measurements.

Chapter 9

FUN MEASUREMENTS WE HAVE DONE

You will not be able to do the set of measurements we describe here. However, they are fun to read about. And they may inspire ideas that you can do.

9.1 Measuring the cosmic-ray muon rate in an airplane at 33,000 ft

A rate measurement was performed during a flight from the Boston International Logan Airport (BOS, latitude = 42.4°) to the Chicago O'Hare Airport (ORD, latitude = 42.0°) using a single detector. The data was recorded to a microSD card with and plugged into a 10,000 mAh USB power bank. The altitude of the airplane was collected from the flight records found at FlightAware.com [53].

Fig. 9.1 (left) shows the trigger rate of the detector in blue as a function of time, binned into 60 second intervals. The error bars shown here are purely statistical. The altitude data of the airplane was linearly interpolated between points to estimate the altitude at any given minute. The interpolated altitude data was the fit to the detector data using a simple exponential plus an offset. Since we did not know the absolute take-off time (data was recorded to the microSD card), we allow the altitude time stamps to shift during the minimization. The best-fit equation is shown at the top left of this figure, where $\text{ALT}[t]$ represents altitude measured in kilometres as a function of time. The best-fit is also plotted as a dashed red line.

Fig. 9.1 (right) shows the measured trigger rate as a function of true altitude. Here, we show the exponential fit extended beyond the measured values. The count rate uncertainties were calculated by taking the square root of the sum all the events measured at a particular altitude. One thing to note is that this is data taken with a detector in master mode, which means it is also sensitive to the background radiation from the interior of the plane.

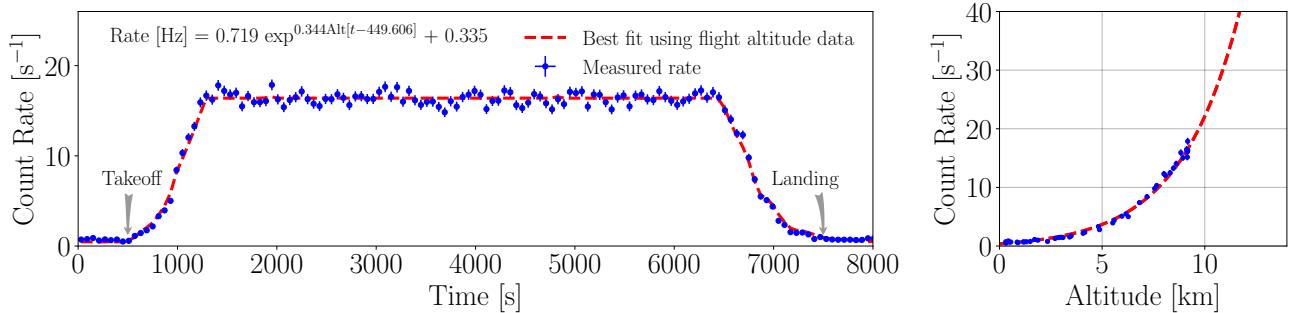


Figure 9.1: (Left) The count rate measured during a flight from Boston to Chicago as a function of flight time. The dashed red line shows the actual amplitude of the airplane [53] scaled by a fitted exponential shown at the top left of the plot. (Right) The measured count rate as a function of altitude. The dashed red line shows the fit.

The cosmic-ray muon flux is known to vary also as a function of latitude, which we have avoided by measuring the flux at a near constant latitude. We also expect the exponential fit to fail as we extend to higher altitudes due to the change in flux composition near the primary cosmic ray interaction region.

9.2 High altitude balloon measurement at 107,000 ft



Figure 9.2: An image from the high-altitude balloon flight at 107,000 ft. Photo from *Daniel Kaczmar - DNF Systems*.

As noted in Sec. 2.1, primary cosmic rays interacting in the upper atmosphere produce showers of particles, some of which decay to muons. Muons are typically produced near an altitude of

15 km. At higher altitudes, there is an increase in the contribution from other ionizing particles, primarily from electrons/positrons, and protons.

During the NearSpace2018 conference [54] in Torun Poland, we participated in a high-altitude balloon (HAB) flight to measure the ionizing radiation flux as a function of altitude. Two detectors were used for the flight so that we could measure both the total rate on the master detector and the down-going rate on the coincident detector. The detectors were placed one-on-top of another (configuration (d) in Fig. 5.6) and taped together to ensure their orientation relative to each other remained the same throughout the flight. The BNC connectors and the OLED screens were removed from the PCB to reduce the weight. An 8" 3.5 mm audio cable was used to connect them into coincidence mode and the SDCard.ino code was uploaded to both detectors. An image of the two detectors is shown in Fig. 9.3. Both detectors were powered by single cell lithium ion battery.

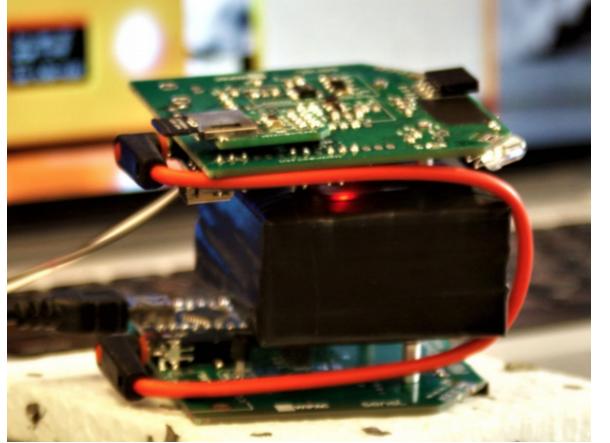


Figure 9.3: The two detectors flown in the HAB flight. The scintillators were taped together in order to preserve the detector orientation relative to each other.

The temperature during the ascent was expected to reach -60°C , therefore thermal protection was required both for the battery and in order to minimize the effect on the SiPM described in Sec. 4.1. A $10 \times 10 \times 10 \text{ cm}^3$ Styrofoam enclosure was constructed with a wall thickness of 1 cm, to house the components. This was sufficiently large that we could place the two detectors and a small heating element in the enclosure with two single-cell Lithium-ion batteries (one to power the detectors, and the other to power the heater). A micro-switch was connected to the battery and wired outside the enclosure so that we could initialize the detectors from outside the enclose, just prior to the flight.

The HAB was launched on September 22nd, 2018 at 12:53 pm. DFN System recorded the balloon altitude and location using on-board GPS, they also mounted a camera to the balloon that looked down at the payloads. An image near the maximum altitude of the flight is shown in Fig. 9.2. The master (orange) and coincident (green) detector count rate, binned in to 60 second intervals, is shown in Fig. 9.4, along with the altitude data from the GPS (black).

The shape of the measured spectrum in Fig. 9.4 corresponds to the Pforzheim curve. We find an initial maximum count rate (Regener-Pforzheim maximum) for the master detector at an altitude from approximately 16-23 km approximately 70-95 minutes into the measurement. After the balloon popped (at minute 118), the detectors fell through the Regener-Pforzheim maximum. The decrease in the trigger rate after passing the maximum occurs due to the detectors ascending beyond the primary interaction region.

The coincidence detector shows a flatter maximum at an altitude from approximately 12-25 km.

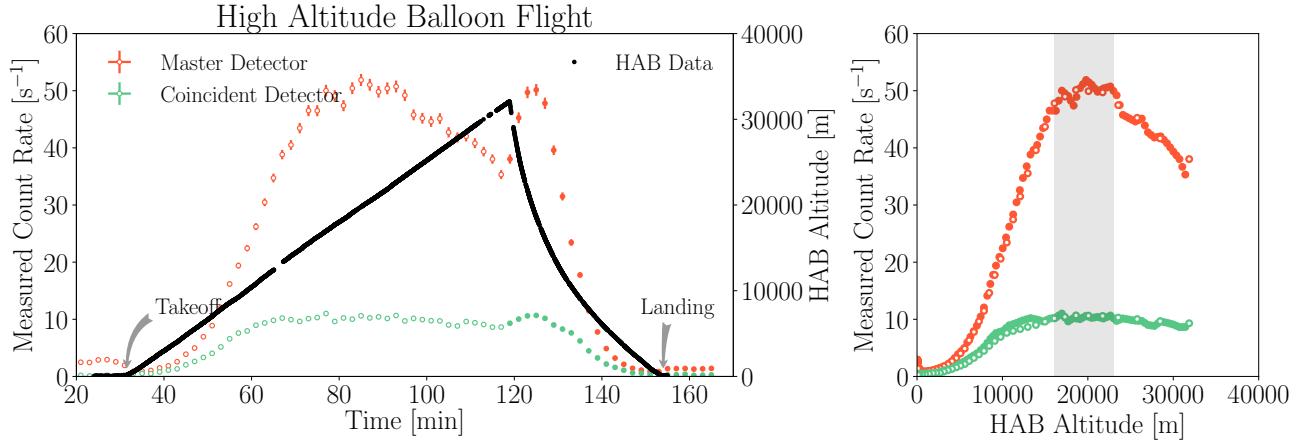


Figure 9.4: (Left) The measured trigger rate for the master (orange) and coincident (green) detector during the high-altitude balloon flight. Take-off occurred 30 minutes after powering on the detectors. The black data points correspond to the altitude as measured through the on-board GPS. (Right) The GPS altitude as a function of trigger rate. The uncertainty in both these plots are statistical.

The peak begins at lower altitudes since we are now preferentially triggering on vertically down-going particles. As described in Sec. 2.1.2, primary particles entering the Earth’s atmosphere at larger angles from the zenith will interact at higher altitudes. In agreement with the data.

9.3 Muon rate measurement while flying to the South Pole

In December 2018, I flew to the South Pole as part of a field team to perform maintenance and upgrades on the IceCube Neutrino Observatory [55]. During the flight I measured the ionizing radiation with two detectors orientated in configuration (d) of Fig. 5.6. Data was recorded to the microSD cards and powered through a single 30,000 mAh power bank.

The first four flights (from Madison, WI USA to Christchurch, New Zealand) were operated by United Airlines and New Zealand Airlines. The altitude data was publicly available for these flights on FlightAware.com [53]. The flight leaving from Christchurch New Zealand, to McMurdo Antarctica, was on a C-17 military jet operated by the US Air Force. Similarly, we flew on a C-130 Hercules, the day after to the South Pole. Since these were military flights, the altitude of this flight was not available; however, several altitude measurements on the second flight were made using GPS. We landed several days later on the 2820 m thick South Pole glacier, approximately 0.5 km from the actual Geographical South Pole.

The full master and coincident detector data are shown in Fig. 9.6, with descriptions of each flight in the text boxes. To give perspective for other measurements, the total data collected



Figure 9.5: Leaving McMurdo Station, Antarctica on a C-130 Hercules for the South Pole Station.

by the master detector was about 50 Mb, whereas the coincidence detector was approximately 15 Mb.

Fig. 9.6 illustrates several very interesting properties. First, there is a trend towards lower count rates near the equator. This is due to the latitudinal variation in the cosmic-ray flux described in Sec. 2.1.2. This effect is most obvious in the flight from SFO to ACK (labelled as flight 3 in Fig. 9.6), which travelled at nearly a constant altitude and a constant rate of latitude change from $+32^\circ$ to -32° . We see that the rate is not symmetric. This is because the magnetic latitude is offset from the geographical latitude, which in turn is because the magnetic field is not symmetric about the equator. Second, it's interesting to see that when we landed at the South Pole, there is a noticeable change in trigger rate due to the combination of the elevation and change in Earth's magnetic field.

While flying through the equator at 35,000 ft, I also performed an East-West measurement using configuration (e) of Fig. 5.6. We measure a count rate coming from the east of 0.69 ± 0.02 cps, while from the west 0.84 ± 0.03 cps. This represents a $22.2 \pm 7.4\%$ increase in the westward direction. This is due to the east-west asymmetry described in Sec. 2.1.2.

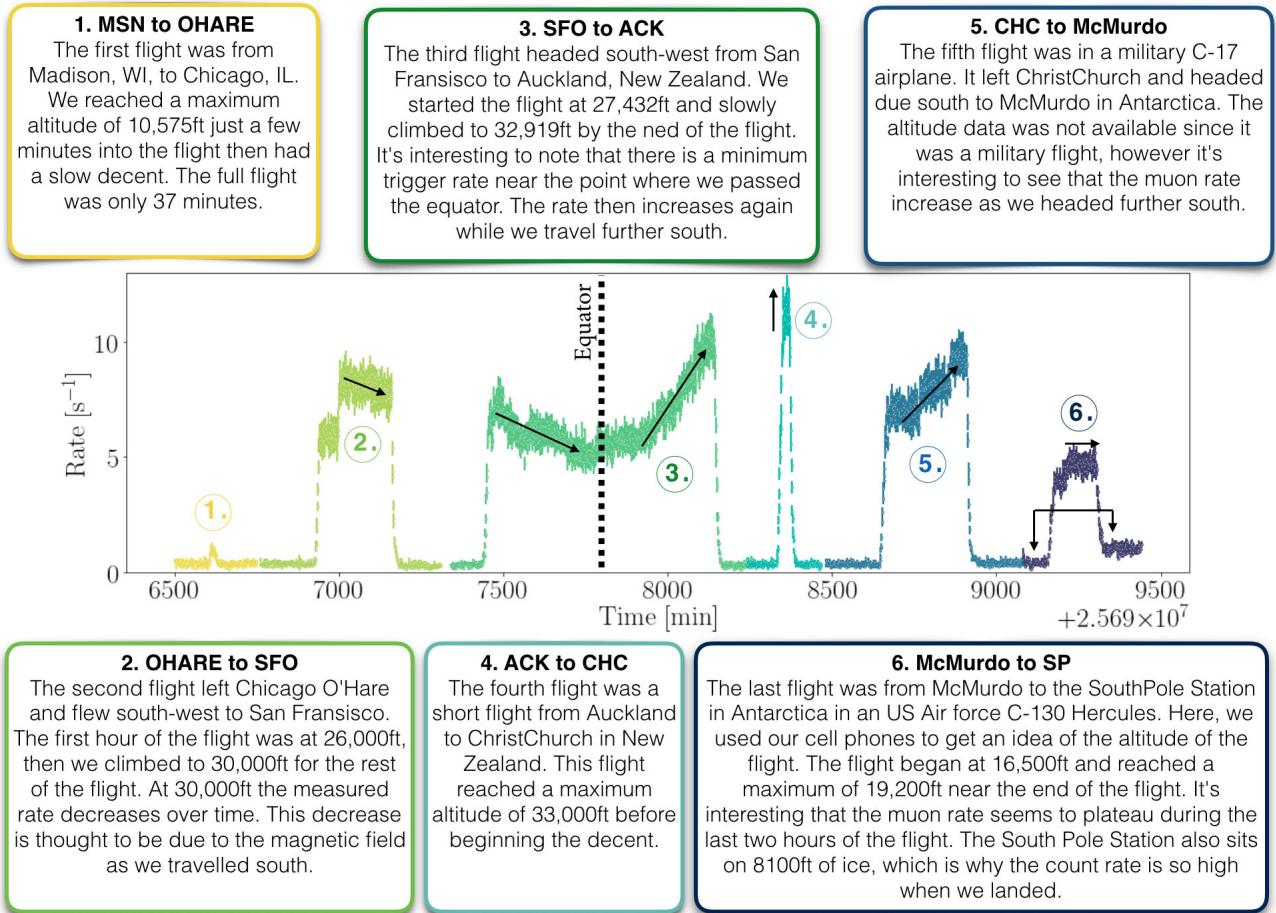


Figure 9.6

9.4 Latitude correction to the cosmic-ray muons

We observed latitudinal variation in the cosmic-ray muon rate in the previous measurements. This was an expected effect due to the change in the Earth's magnetic field as a function of latitude. Here, we will empirically attempt to account for the variation in the latitude and altitude based on the previous measurement. We will assume that the change in the rate as a function latitude follows a sine-squared form; where the minimum occurs near the equator and the maximum occurs near the poles:

$$R[Hz] = N \exp^{\alpha(\sin(LAT[t]+\theta)^2+\beta) \times ALT[t]} - b \quad (9.1)$$

Here, θ represents a phase that offsets the latitude to account for the difference between the geographical latitude and the magnetic latitude, β is a factor that permits an altitude effect at the magnetic equator (set the sine term to zero), α is a scale factor that dictates the strength of

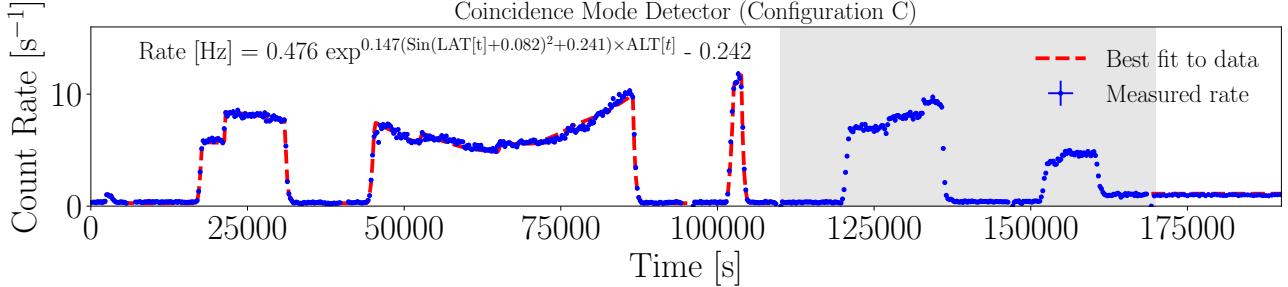


Figure 9.7: The fitted data based on the altitude and latitude data from the first four flights. The fit also includes a measurement after landing at the South Pole, where the altitude and latitude were known (2700 m and -90°). The two military flights were not included in the fit.

the latitudinal and altitude effect, N is a normalization that sets the scale for the exponential component, and b is an offset that accounts for a constant background radiation.

Fig. 9.7 uses the function in Eq. 9.1 to fit the data for the first four flights, plus the data after landing at the South Pole. The best fit values are shown at the top left of Fig. 9.7. Using this result, we can then invert Eq. 9.1 to calculate the altitude of the two military flights (assuming we flew at a constant velocity directly south). The calculated altitude is shown in Fig. 9.8 for the military flights in green, as well as the measured altitude from GPS in red.

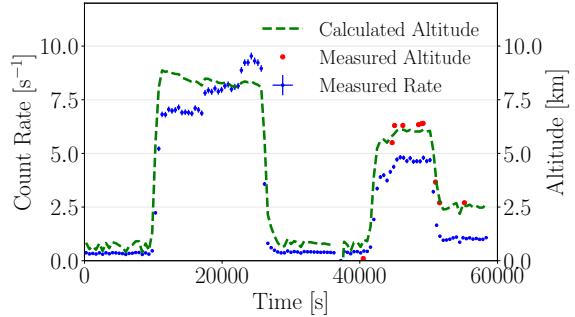


Figure 9.8: The measured muon rate in blue along with the calculated flight altitude in green. The red markers show the fly GPS altitude measurements during the flight.

9.5 Rate measurement 1 km underground at Super-Kamiokande

Two Desktop Muon Detectors were brought to the Kamioka Observatory located 1 km underground in the Mozumi Mine, Japan. This mine is home to several high profile experiments, perhaps most notably the 2015 Nobel prize winning particle physics experiment, Super-Kamiokande. Two detectors were placed in the Super-Kamiokande control room for 8 hours, and connected together via a 6-inch 3.5 mm audio cable in configuration (c) of Fig. 5.6. The data was recorded from the coincidence detector through the import_data.py script directly to a laptop. Using the same detectors and set-up, a rate measurement was also performed outside the Kamioka mine in the observatory dormitory and in the airplane at 36,000 ft when travelling between Warsaw to Tokyo. Fig. 9.10 shows the trigger rate of the coincident detector for these three measurements, as a function of calculated SiPM peak voltage.



Figure 9.9: Dr. Katarzyna Frankiewicz floating in the inner detector of Super-Kamiokande during the Gadolinium upgrade 2018.

The total number of measured coincident events measured inside the Super-Kamiokande control room was found to be 101. It was found that 96% of these events were located below the 50 mV peak described in Sec. 6.1, indicating that these are likely not minimum ionizing cosmic-ray muons.

The average rock density in the mine was measured to be 2.7 g/cm^3 , corresponding to approximately 2,700 m.w.e. (meter-water equivalent) of overburden [7]. Based on this, we expect the cosmic-ray muon rate to be attenuated by a factor of 10^5 compared to a ground level measurement. With this assumption, we only expect approximately 0.04 cosmic-ray muon events over the 8-hour measurement in the Super-Kamiokande control room.

The master detector count rate did not significantly change when it was brought into the mine, indicating that the radioactive background was still present in the control room. Given that the master detector count rate was 1 Hz, Eq. ?? suggests that we should expect 1-2 accidental coincidence events over the 8-hour period (see Sec. 7).

One unaccounted background (which was briefly mentioned in Sec. 3.3) consists of events in which a gamma-ray from a radioactive decay Compton scattered off of the scintillator of the master detector, and was then absorbed or deposited sufficient energy in the coincident detector scintillator. This is thought to be the dominant source of triggers in this dataset, and a Monte Carlo simulation is currently being developed to investigate this. If these events originate from Compton scattering, we can estimate the rate for these types of events. Given the 101 events (and 1-2 of these are assumed to be accidental coincidences and cosmic-ray muons), the calculated accidental Compton scattering coincidence rate in configuration (c) from Fig. 5.6 is found to be

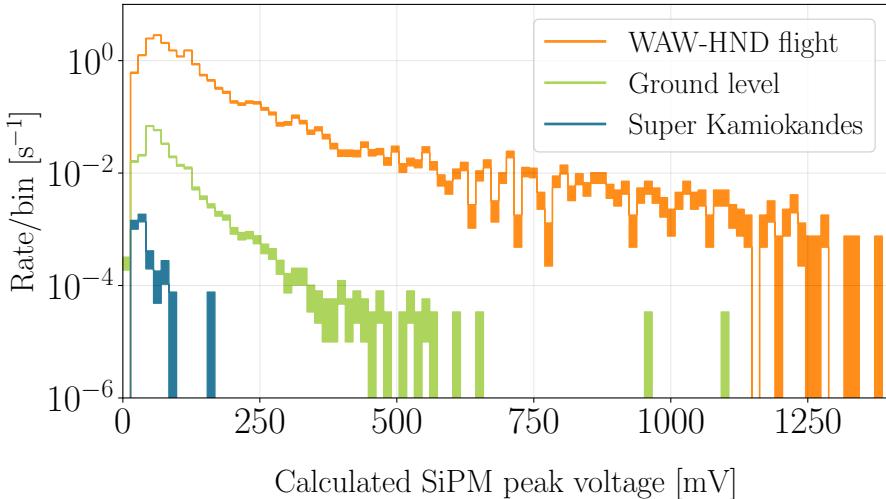


Figure 9.10: The measured coincidence rate at three locations. Data from Ref. [49].

0.0038 ± 0.0004 . A second 8-hour run was performed using the same configuration and location, which found 92 events (with a similar SiPM peak voltage spectrum) corresponding to a count rate of 0.0035 ± 0.0004 .

This result could be further investigated by repeating the measurement, however this time with a thin piece of lead between the scintillator. Lead, being a dense material, is likely to either absorb the gamma-ray or absorb some of the energy from the gamma-ray through Compton scattering. Both processes would reduce the probability of measuring the event with the coincident detector. Another potential source for these events is correlated noise. If the lead does not alter the coincident count rate, this is a potential source for this signal, however, thus far, we have not found any evidence of events due to noise.

9.6 Portable trigger system for an accelerator beamline

This measurement was previously described in Ref. [49] and represents a practical use for the Desktop Muon Detectors.

A single detector, powered by a 10,000 mAh USB power bank, was placed in the Fermilab M-Test facility to trigger on secondary particles (GeV-scale pions and electrons) from the Main Injector. The purpose of this was to trigger a downstream data acquisition system for another experiment. The BNC output at the back of the detector is the raw SiPM pulse, which has a rise time of a few nanoseconds and a decay time of roughly $0.5 \mu\text{s}$. This signal is useful for experiments that want to use a scintillators but require tens of nanosecond timing. The BNC output was connected to an 80 ft BNC cable to a NIM (Nuclear Instrument Module) rack . The signal passed through a $\times 10$ amplifier and into a discriminator. If the amplified signal was

above a certain value, a binary signal was sent to a AND gate, where it was compared against another scintillator paddle trigger that was located on the other side of the other experiment. If the AND condition was satisfied (ie. the particle passed through both the scintillator paddle and the Desktop Muon Detector) a binary signal was sent to the data acquisition system that began the recording of data of the downstream experiment. Fig. 9.11 shows the trigger rate of the detector placed in the beamline as a function of time. The beam spills occur every minute for two seconds.

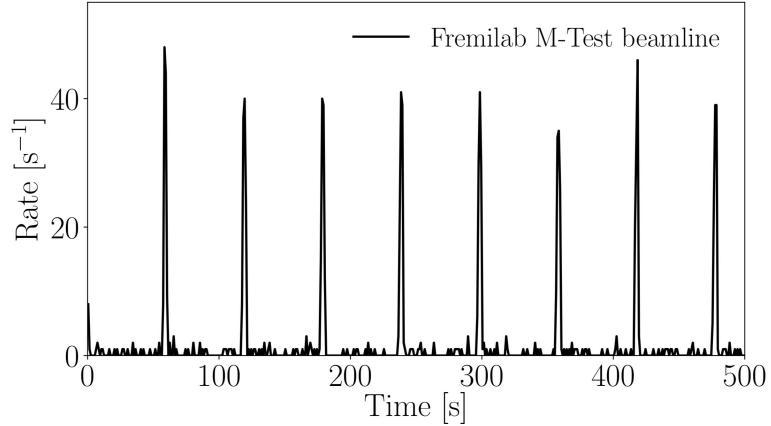


Figure 9.11: The trigger rate as a function of time of a single detector placed in the Fermilab M-Test beamline. Here, the detector is triggering primarily on GeV-scale pions and electrons from the Fermilab Main Injector.

The detector was found to be a useful beamline trigger due to its simplicity. This ability was made possible by including a BNC output connected directly to the SiPM. For this measurement, the approximate 10 ns uncertainty in the trigger was acceptable, however if users would like to use the FAST output of the SiPM, the SiPM PCB could be modified to get down to single ns precision. We plan on investigating this sometime in the future.

9.7 Gamma-ray spectroscopy

The detection material is primarily made out of polystyrene, a low density, synthetic hydrocarbon. If we look back at Fig. 3.4, we see that for carbon-materials, at MeV gamma ray energies, the dominant interaction is Compton scattering. Ideally, if you want to perform gamma ray spectroscopy, you want to capture the full energy of the gamma ray, such as with the photoelectric absorption. However, this interaction is more common with high density materials, such as germanium. What this means, is that a common interaction for us to observe is the original gamma ray Compton scattering within the scintillator, depositing some energy, and leaving the scintillator. This smears out the measured energy distribution, lowers the mono-energetic peaks, and increases the rate in the Compton valley.

Although it is difficult, some information can be extracted from gamma ray sources. An example of this is shown in Fig. 9.12. Here, we have taken data with many sources found in Junior Lab. While we cannot pull out the monoenergetic peaks, we do see some interesting features. For example, Cobalt-60 emits two rather high energy gamma rays (1.17 and 1.33 MeV), compared to Cesium-137, which has a dominant gamma ray at 0.66 MeV. Notice that the SiPM peak voltages, which is a proxy for the number of observed photons, which is related to the energy deposited by a gamma ray.

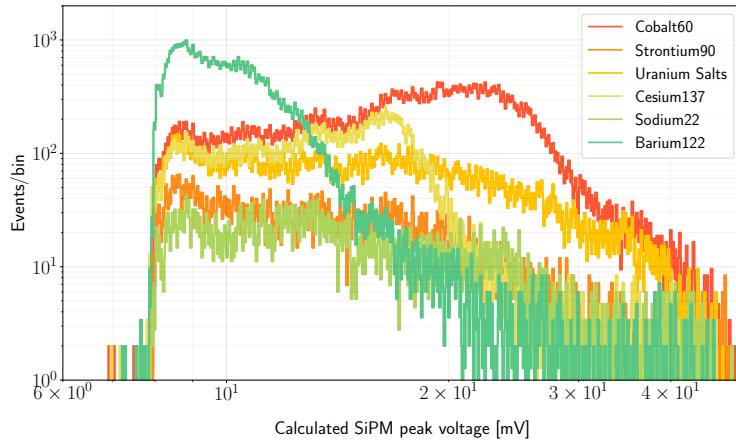


Figure 9.12: The distribution of number of events per unit time. Each cell represents a different amount of time in which the data is binned (labelled on the top right of each cell). The red line is the expected Poissonian distribution from Eq. 5.4.

Gamma ray spectroscopy with these detectors would be improved with a denser material, larger volume detectors, and larger photocathode coverage.

Chapter 10

CONCLUSION

The CosmicWatch Desktop Muon Detectors are capable of exploring various physical phenomena in nature. This document outlined in detail the physical processes that influence the detector, and how they can be extracted from data. These can be used to investigate various phenomena associated with the geomagnetic field, atmospheric conditions, cosmic-ray shower composition, attenuation of particles in matter, radioactivity, and statistical properties of a Poisson process. Students are encouraged to develop the ideas that we have already started, or design an experiment of their own. We also strongly encourage feedback so that we can improve this manual. Enjoy your CosmicWatch detectors!

Appendix A

DESCRIPTION OF THE CIRCUIT

This chapter describes the philosophy of the analog electronics used to extract information from the SiPM. The SiPM pulse amplitude is proportional to the total number of photons incident on the SiPM's photocathode. Therefore, the aim of the analog electronics is to shape the pulse such that the microcontroller is capable of measuring some information related to the pulse amplitude.

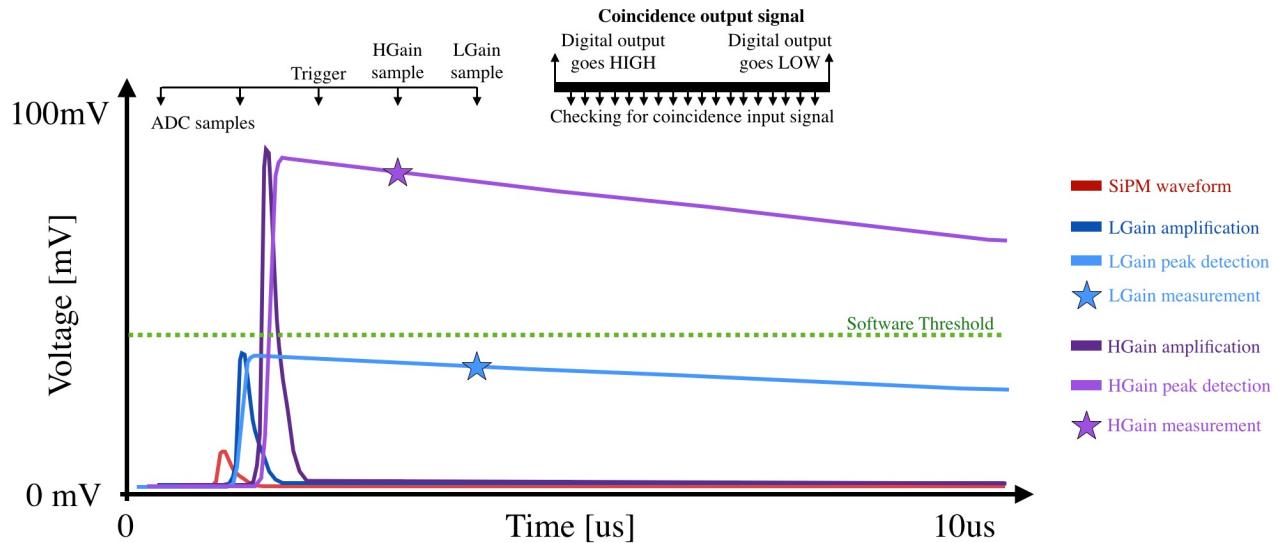


Figure A.1: A description of the analog circuitry and CPU routine. The SiPM pulse (red) is amplified (dark blue), from which we amplify for the high gain channel, and

Fig. A.1 shows a cartoon of the four stages of the analog electronics. The SiPM pulse (red) is first amplified by approximately a factor of 15 (dark blue), this is the low gain amplification section. The low gain amplification signal is then feed into two circuit: a peak detection circuit (light blue) and a second amplification stage (dark purple). The peak detector circuit holds the

value of the peak (decaying away as an exponential), for sufficient time that the microcontroller can make a measure a single measurement (light blue star). The measured ADC value of the light blue star is the Low Gain measurement. The second stage of amplification further amplifies the pulse by a factor of 7, then flows into a second peak detector circuit (dark purple). The dark purple star represents the High Gain measurement.

At the top of Fig. A.1, we show an approximation of the routine performed by the CPU. The ADC samples the high gain peak detector circuit continuously until an ADC value is found above the signal threshold (green line). This is labeled as “Trigger” in the figure. The following ADC sample is used for the High Gain measurement, and the second sample, on a separate ADC channel samples the Low Gain circuit. A single ADC sample takes approximately $1.3\mu\text{s}$. After completing the measurement, the coincidence output goes HIGH, and then we probe the digital input from the coincidence detector many times. If the digital input is observed to be HIGH, we call that a coincident event. The number of times we probe the digital input defines the coincidence time window. After, the digital coincidence output goes LOW.

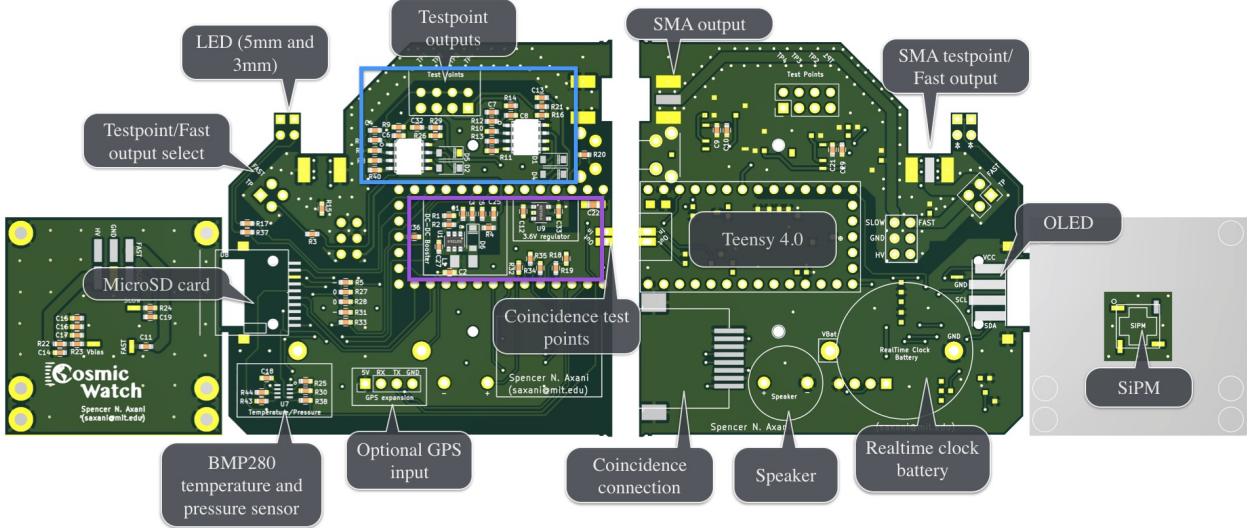


Figure A.2

Fig. A.2 shows the printed circuit boards (PCBs) used for the v3 detector. The left (right) shows the bottom (top) of the PCBs. The PCB is actually two separate pieces: the SiPM PCB and the Main PCB. Each component of the circuitry is labeled. The blue box indicates the analog circuitry, described above, while the purple box indicates the power handling. This includes a DC-DC booster that bring the 5V voltage from the USB connection to 28.9V, and a 3.6V regulator to power the op amps used in the analog circuitry.

We provide Fig. A.3 to show the internal components of the detector. Often, it is desirable to measure the full-sky muon rate. This is best done by placing the scintillators between two detectors as close as possible, as shown in Fig. 5.6 (d). If this is the case for your measurement, then feel free to remove the detectors from the cases.

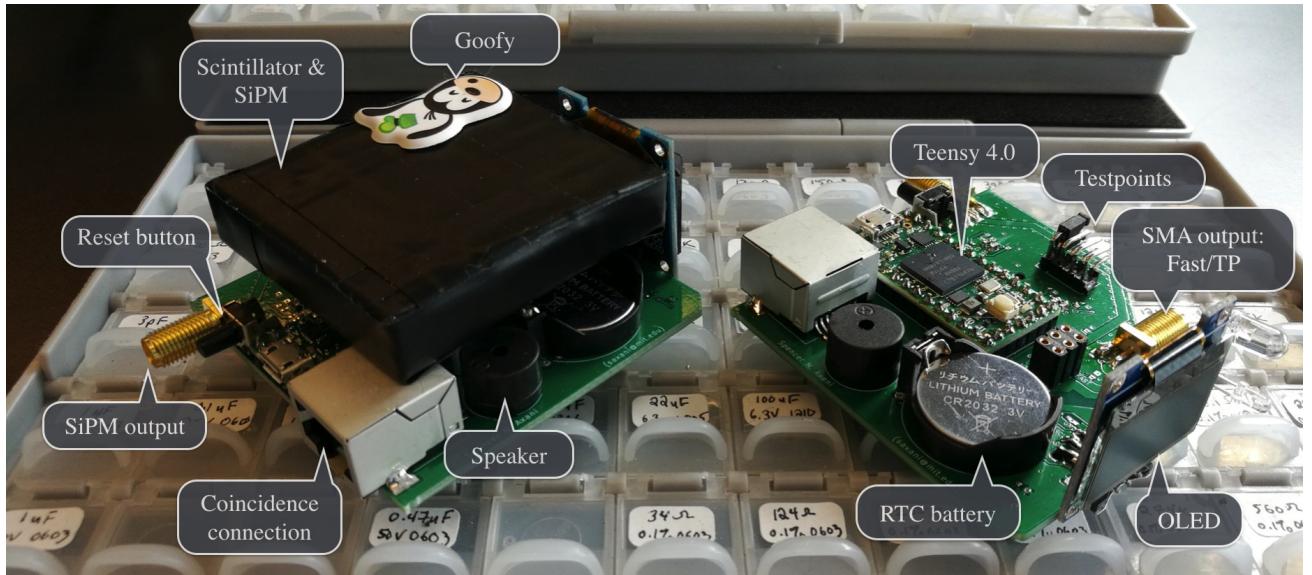


Figure A.3: Image of the detectors removed from the cases.

The final figure in this section, Fig. A.4, shows a vector image of the full circuit. Feel free to send Spencer questions (saxani@mit.edu), since we'll omit a detailed description here.

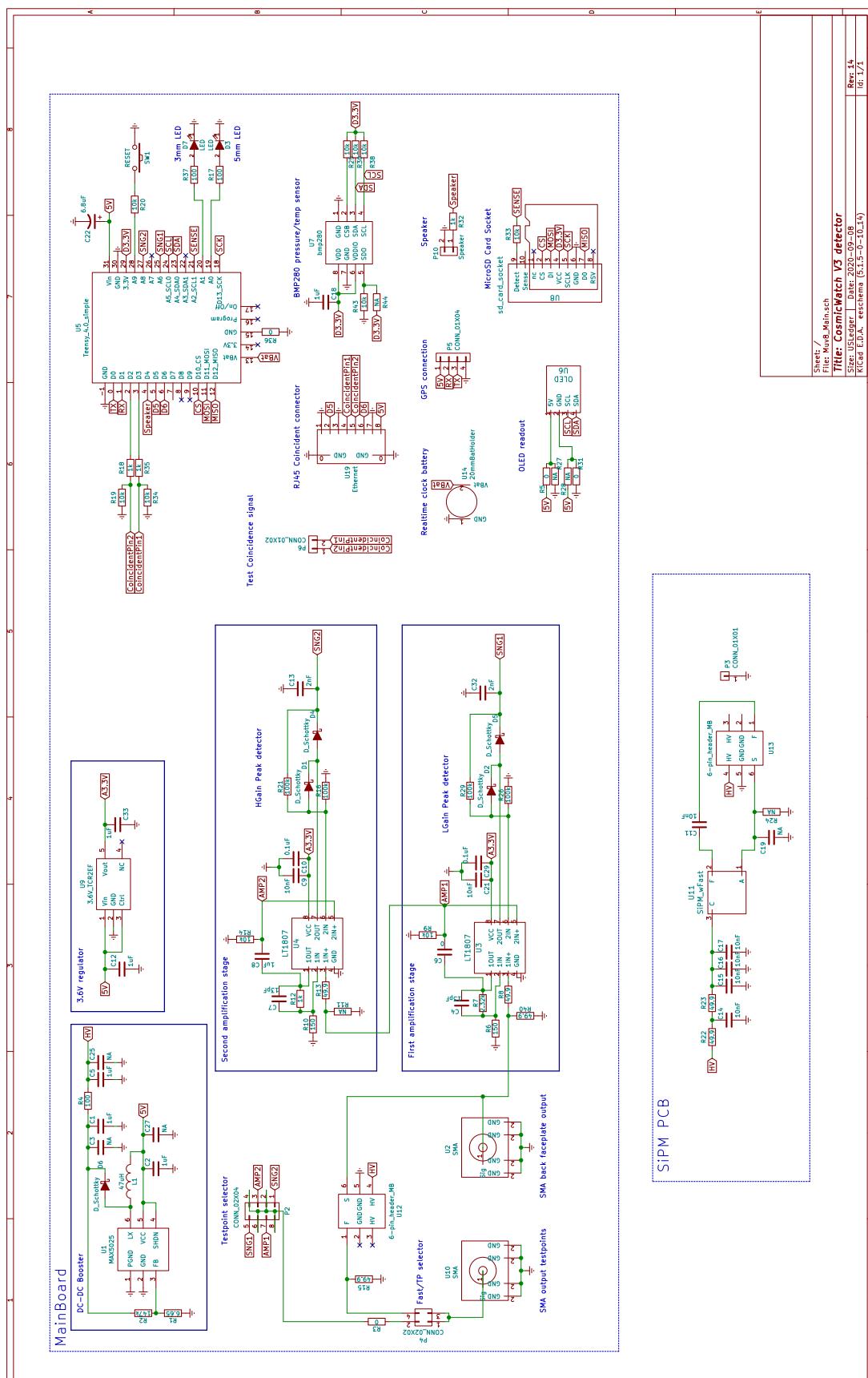


Figure A.4

Sheet 1 /
File No.8, Main sch
Title: CosmicWatch V3 detector
Status: Design
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