

# Collision rates of planetesimals near mean-motion resonances

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## ABSTRACT

In circumstellar discs, collisional grinding of planetesimals produces second-generation dust which can be observed through thermal emission. While it remains unclear when second-generation dust first becomes a major component of the total dust content, the presence of such dust and potentially the substructure within it can be used to explore a disc's physical conditions. A perturbing planet has been shown to produce nonaxisymmetric structures, as well as gaps in discs, regardless of the origin of the dust. However, small grains will have very different dynamics compared with planetesimals when in the presence of gas, and as such, the collisional evolution of planetesimals could create dusty disc structures that would not exist otherwise. Previous studies have shown that morphological differences in the collisional dust corotating with a giant planet can be used to determine whether the planet's eccentricity is large or small. In this work, we use N-body simulations that directly resolve collisions between planetesimals to investigate the collision rate profile near mean-motion resonances with the planet. We find that a distinct bump or dip feature is produced in the collision profile, the presence of which depends on the libration width of the resonance and the separation between the peri and apocenter distances of the edges of the resonance. In the collision profile, the presence of one of these two features limits possible values for the mass and eccentricity of the planet. Under the simplifying assumption that the radial profile of the dust emission in the disc traces the collision profile, we show that a bump or dip feature in the dust emission near the interior 2:1 mean-motion resonance could be used to place constraints on the orbital properties of the perturbing planet. So long as radial drift does not play a significant role in the collisional cascade that grinds planetesimals down to dust, this feature in the dust emission should be marginally observable in nearby protoplanetary disks with ALMA.

**Key words:** planet-disc interactions – planets and satellites: formation – submillimetre: planetary systems – methods: numerical

## 1 INTRODUCTION

Recent observations of circumstellar discs by ALMA have revealed a rich variety of substructure. Features such as gaps and asymmetries (ALMA Partnership et al. 2015; Pérez et al. 2016; Isella et al. 2016; Andrews et al. 2016; Cieza et al. 2016) in the emission provide diagnostics for the physical processes that drive the evolution of the discs. In many gas-rich systems, these gap features are argued to indicate the presence of a giant planet, either embedded in the disc (Dipierro et al. 2015) or orbiting adjacent to it, although the details of gas-dust interactions can have strong effects on the resulting morphologies (Dong et al. 2018). In gas-poor/debris systems, a giant planet perturber can also influence the structure of the dust continuum emission. For example, a misaligned giant planet can produce nonaxisymmetric features such as warps (Augereau et al. 2001), highly eccentric perturbers can produce structures through

secular interactions (Pearce & Wyatt 2014; Pearce et al. 2015), and mean motion resonances (MMRs) can open gaps (Nesvold & Kuchner 2015; Tabeshian & Wiegert 2016, 2018).

Collisions between planetesimals are thought to be the principal source of dust for debris discs. (see Wyatt 2008). Although some amount of primordial dust is likely still present during the early stages of debris disc evolution, ongoing collisions between small bodies will augment this and could be used to trace the underlying dynamical activity of the planetesimals. This process was directly detected when the putative object Fomalhaut b was confirmed to be an expanding debris cloud, most likely from a planetesimal collision (Gaspar & Rieke 2020).

Dobinson et al. (2013) showed that collisional dust that is generated near the gap opened by a giant planet in the inner regions of a transition disc should produce a distinct observable marker that could be used to infer the presence of the planet. In a followup study, it was found that morphological differences in the dust emission could be used to determine whether the planet had a circular or

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eccentric orbit (Dobinson et al. 2016). Moreover, substructure due to mean-motion resonances (MMRs) with a giant planet may hold additional clues that could be used to constrain the properties of the planet. The width of a resonance is set by both the mass of the perturbing planet and the unperturbed eccentricity of the planetesimals, which is in turn set by the eccentricity of the planet through secular forcing. Therefore, the structure of second generation dust produced near MMRs may encode information about both the mass and eccentricity of the planet.

The dynamics governing the motion of bodies near MMRs is extremely nonlinear, as is determining what the collision rates between planetesimals should look like in these regions. For a collection of bodies massive enough to experience the effects of gravitational focusing, a large eccentricity dispersion tends to reduce the probability of collision, while enhancements in surface density tends to increase it. Due to conservation of the Jacobi energy, MMRs simultaneously enhance the local eccentricity dispersion and also enhance the surface density adjacent to the resonance (Richardson et al. 2000; Boley 2017). Unfortunately, collision detection in an N-body simulation is extremely computationally expensive. So far, studies of planetesimal dynamics near MMRs **in which the planetesimals are directly resolved** have involved either collisionless test particles (Boley 2017; Tabeshian & Wiegert 2016, 2018) or severely limited integration times (Richardson et al. 2000).

To further elucidate this subject, we use the tree based N-body code CHANGA (Jetley et al. 2008; Menon et al. 2015) to follow the collisional evolution of a planetesimal disc under the gravitational influence of a Jupiter-sized body. Because particle positions are sorted into a tree structure, neighbor finding and collision detection can be done quickly and efficiently. This considerably relaxes the constraints on resolution and integration time. With this toolset, we explore the collision rate structure of a planetesimal disc in the vicinity of mean-motion resonances with the planet. In particular, we would like to determine (1) **What dynamics govern the collision rate profile near MMRs**, (2) whether MMRs leave a detectable signature in the collisionally-generated dust, and (3) whether these signatures can be used to determine or constrain the orbital properties of the perturbing planet.

This work is organized in the following way: In section 2, we provide an overview of the relevant dynamics that drive the evolution of a planetesimal disc under the gravitational influence of an external perturber. In section 3 we provide an overview of the N-body code used and describe the initial conditions chosen for five simulations in which a perturbing giant planet is given various masses and eccentricities. Section 4 presents the results of these simulations **and we take an in-depth look at the collision rate profiles of the planetesimals near the MMRs**. In section 5, we use the resolved collisions to generate synthetic dust emission profiles that would be detected with observing facilities like ALMA, **under the simplifying assumption that the resulting dust profile traces the collisions of the planetesimals. Under this assumption, we show** that a characteristic bump or dip feature appears in the dust emission near the interior 2:1 MMR, the presence of which depends on the mass and eccentricity of the perturbing planet. If the 2:1 MMR can be identified (presumably by identifying another prominent MMR, such as the 3:1, and measuring the spacing between the two), we discuss the potential of using this feature to place constraints on the mass and eccentricity of the planet. **In addition, we discuss the caveats of connecting the dust and planetesimal collision profiles in such a simple way**. Finally, we conclude in section 6.

## 2 OVERVIEW OF RELEVANT DYNAMICS

We begin by providing a description of the dynamical effects responsible for shaping the orbital distribution of the planetesimals. The purpose of this is twofold: (1) to motivate the initial conditions used for the simulations described in section 3.2 and (2) to justify the exclusion of certain physical effects from our simulations. Here, we focus solely on the physics relevant for full-sized ( $\sim 100$  km) planetesimals and leave a discussion of the effects on smaller bodies generated through collisions for section 4.2.

### 2.1 SECULAR FORCING

The most direct and widespread effect that a perturbing giant planet will have on a planetesimal disc is through secular forcing of the planetesimals. This will cause the complex eccentricities of the planetesimals to take on a time-independent forced value, given by (Wyatt et al. 1999) as

$$z_f = \frac{b_{3/2}^2(\alpha)}{b_{3/2}^1(\alpha)} e_g \exp i\varpi_g. \quad (1)$$

Here,  $\alpha = a_g/a$  where  $a_g$  and  $a$  are the semi-major axes of the giant planet and a planetesimal, respectively.  $e_g$  and  $\varpi_g$  are the eccentricity and longitude of pericenter of the giant planet, and  $b_s^j(\alpha)$  is a Laplace coefficient given by (Murray & Dermott 1999) as

$$b_s^j(\alpha) = \frac{1}{2\pi} \int_0^{2\pi} \frac{\cos j\theta d\theta}{(1 - 2\alpha \cos \theta + \alpha^2)^s}. \quad (2)$$

Without any nearby secular or mean motion-resonances, equation 1 will completely describe the eccentricities and longitude of pericenter orientations of the planetesimals. Additional forces due to two-body scattering between planetesimals, along with aerodynamic gas drag will add an additional free component to the complex eccentricity, which will be randomly oriented. The magnitude of the free eccentricity describes how dynamically hot the planetesimal disc is and sets the random encounter speeds of planetesimals. When the dynamical excitation of the disc is driven by gravitational stirring, the magnitude of the free eccentricity can be described by a Rayleigh distribution (Ida & Makino 1992).

**The timescale for secular forcing is given by the period at which the free eccentricity cycles in the complex plane. The imaginary and real components of the eccentricity of the  $j$ th particle are given by Murray & Dermott (1999) (see ch 7, pg 282, equation 7.51)**

$$h_j = e_{j1} \sin(g_1 t + \beta_1) + e_{j2} \sin(g_2 t + \beta_2) \quad (3)$$

$$k_j = e_{j1} \cos(g_1 t + \beta_1) + e_{j2} \cos(g_2 t + \beta_2),$$

where  $(e_{j1}, e_{j2})$  and  $(g_1, g_2)$  are the corresponding eigenvector components and eigenvalues of the matrix given by Murray & Dermott (1999) (see ch 7, pg 276, equation 7.9 and 7.10, with  $\alpha_{12} = \alpha$  and  $\alpha_{12} = 1$ )

$$A_{jj} = n_j \frac{1}{4} \frac{m_k}{m_c + m_j} \alpha b_{3/2}^{(1)}(\alpha) \quad (4)$$

$$A_{jk} = -n_j \frac{1}{4} \frac{m_k}{m_c + m_j} \alpha b_{3/2}^{(2)}(\alpha)$$

Here,  $n$  is the mean motion of a body and  $m$  is its mass. The frequency of the secular forcing is simply the difference of the two eigenvalue frequencies  $\omega_1 - \omega_2$ . For the case of a Jupiter mass planet at 5.2 au perturbing a test particle, the timescale for secular forcing is approximately 12,000 years at 3 au.

## 2.2 Mean-Motion Resonances

In regions where there are commensurabilities between frequencies, Laplace-Langrange secular theory breaks down and bodies are subject to strong perturbations. For the purposes of this study, we will ignore secular resonances, which generally occur on rather large timescales and will focus on mean-motion resonances. A MMR occurs when the orbital period ratio between two bodies is sufficiently close to

$$\frac{P}{P'} = \frac{p+q}{p}, \quad (5)$$

where  $p$  and  $q$  are integers  $> 0$  and the unprimed and primed quantities correspond to the perturber and the body being perturbed, respectively. In terms of these quantities,  $P/P'$  corresponds to a  $p+q:p$  resonance. If the perturber is much more massive than the other body and all of the bodies lie in a near-Keplerian potential, the condition for MMR is set by

$$\frac{a}{a'} = \left( \frac{p}{p+q} \right)^{2/3}. \quad (6)$$

If we further assume that the two bodies are orbiting in the same plane, the motion of the bodies near resonance is determined by the behavior of a critical angle

$$\phi = (p+q)\lambda' - p\lambda - q\varpi, \quad (7)$$

where  $\lambda = \varpi + M$  is the mean longitude of a body, with  $M$  being the mean anomaly. For bodies in resonance, the critical angle will librate around an equilibrium value, while this angle will circulate outside of resonance. For small eccentricities, this behavior is analogous to the motion of a pendulum. Furthermore, variations in the critical angle are coupled to changes in the mean motion and semimajor axis (Murray & Dermott 1999). An important point to note, which we will revisit later, is that the variation frequency of this angle approaches zero near the edge of a resonance. The width of a resonance can be defined by determining the largest variation in semimajor axis that permits librational, rather than circulatory motion of  $\phi$ . For second order and higher MMRs, the maximum libration width is given by Murray & Dermott (1999) (ch. 8, pg 338, eq. 8.58) as

$$\frac{\delta a}{a} = \pm \left( \frac{16}{3} \frac{|C_r|}{n} e^{|q|} \right)^{1/2}, \quad (8)$$

where  $n$ ,  $a$  and  $e$  are the mean motion, semimajor axis and eccentricity of the unperturbed body.  $C_r$  is a constant defined by  $m'/m_c n \alpha f_d(\alpha)$ , where  $\alpha = a/a'$  and  $f_d$  is the resonant part of the disturbing function. For an interior second-order resonance,

$$\begin{aligned} f_d(\alpha) = & \frac{1}{8} \left[ -5(p+q) + 4(p+q)^2 - 2\alpha D \right. \\ & \left. + 4(p+q)\alpha D + \alpha^2 D^2 \right] b_{1/2}^{p+q}, \end{aligned} \quad (9)$$

(see Murray & Dermott (1999) table 8.1, third row, with  $j = p+q$ ) where  $D b_s^j$  is the first derivative with respect to  $\alpha$  of the Laplace coefficient defined in equation 2. This is given by Brouwer & Clemence (1961) as

$$\frac{db_s^j}{d\alpha} = s \left( b_{s+1}^{j-1} - 2ab_{s+1}^j + b_{s+1}^{j+1} \right). \quad (10)$$

For a first-order resonance, the associated disturbing function terms are slightly simpler such that

$$f_d(\alpha) = -(p+q)b_{1/2}^{p+q} - \frac{\alpha}{2} D b_{1/2}^{p+q}, \quad (11)$$

(see Murray & Dermott (1999) table 8.1, first row) and there is an additional contribution to the motion of the critical angle by the pericenter precession rate such that

$$\frac{\delta a}{a} = \pm \left( \frac{16}{3} \frac{|C_r|}{n} e \right)^{1/2} \left( 1 + \frac{1}{27p^2e^3} \frac{|C_r|}{n} \right)^{1/2} - \frac{2}{9pe} \frac{|C_r|}{n} \quad (12)$$

(see Murray & Dermott (1999) ch 8, pg 340, equation 8.76 with  $j_2 = p$ ).

In mean-motion resonance, the orbital elements of a perturbed body, along with the critical angle described by equation 7, will exhibit sinusoidal oscillations. In the circular, restricted three-body case, these small amplitude oscillations exhibit a characteristic frequency, given by (see Murray & Dermott (1999) ch 8, pg 332, eq 8.47, with  $j_2 = -p$  and  $j_4 = -q$ )

$$\omega_0 = 3p^2 C_r n e^{|-q|}. \quad (13)$$

For the most prominent interior mean-motion resonances, including the 2:1 and 3:1, the timescale associated with these oscillations driven by a Jupiter mass planet is  $\sim 1,000$  - 2,000 years.

Changes in eccentricity are correlated with changes in semimajor axis according to the Tisserand relation

$$\frac{de}{da} = \frac{a^{3/2} - 1}{2a^{5/2}e}. \quad (14)$$

The net result of this is that MMRs will produce spikes in the semimajor axis-eccentricity distribution of a planetesimal disc. In order to preserve the Jacobi energy, equation 14 predicts that these spikes will tend to curve away from the perturbing body. This lowers the surface density of the disc near the locations of resonances, while producing a pileup of bodies on eccentric orbits adjacent to the MMR.

## 2.3 Collisions Between Planetesimals

A simple analytic model of the collision rate of a planetesimal population is given by Safronov (1969) as

$$n\sigma v = n\pi s^2 \left( 1 + 2Gm/sv^2 \right) v, \quad (15)$$

where  $n$  is the number density of the population,  $s$  and  $m$  are the radii and masses of the bodies and  $v$  is their typical encounter velocity. The encounter velocity is often described in

terms of the rms eccentricity  $\langle e^2 \rangle^{1/2}$  and inclination  $\langle i^2 \rangle^{1/2}$  of the population (e.g. Lissauer (1993)) as

$$v = \sqrt{\langle e^2 \rangle + \langle i^2 \rangle} v_k, \quad (16)$$

where  $v_k$  is the local Keplerian velocity. The second term in equation 15 can be thought of as an additional enhancement to the collision cross section due to gravitational focusing. When the typical encounter velocity is small compared to the mutual escape velocity of the planetesimals, the collision cross section greatly exceeds the geometric value. Another important feature of equation 15 is that for a fixed value of  $n$ , the collision rate has a global minimum as a function of  $v$ . For small  $v$ , gravitational focusing allows for more collisions, while for large  $v$  the number of encounters simply becomes so large that the collision rate again increases, even though gravitational focusing does not operate. This point will come up again in section 4.1 when we examine the qualitative changes in the collision rate near mean-motion resonances.

Although equation 15 works well to describe the collision rate for a homogeneous collection of planetesimals, some problems arise when localized regions of commensurabilities are introduced. Namely, the resonances cause the orbits of planetesimals to precess, and an interface between secularly aligned and randomly oriented orbits arises on each side of the resonance. Additionally, an interface between dynamically cold and hot planetesimals forms at each side. For these reasons, an N-body treatment in which collisions are directly resolved is necessary to understand how the collision rate varies near the MMRs.

## 2.4 Gas Drag on Planetesimals

Over the course of many orbits, the residual gas from the primordial nebula can damp the eccentricities and inclinations of the planetesimals. For a planetesimal-sized body, the geometric cross section is much larger than the mean free path of the gas particles and the timescale to damp its eccentricity is given by (Adachi et al. 1976)

$$t_{e,damp} = \frac{8s\rho}{3eC_D\rho_g v_{th}}, \quad (17)$$

where  $m$ ,  $e$  and  $s$  are the mass, eccentricity and radius of the planetesimal.  $C_D$  is a drag coefficient which is of order unity,  $\rho_g$  is the local density of the gas and  $v_{th}$  is its thermal speed. At 3 au, the gaseous component of the solar nebula has a density of  $4 \times 10^{-10}$  g cm $^{-3}$  and a typical thermal speed of 10 $^5$  cm s $^{-1}$  (Hayashi 1981). For a 100 km body with a density of 2 g cm $^{-3}$  and an eccentricity of 0.1, the damping timescale is on the order of 10 $^5$  years. This is much longer than the timescales related to secular forcing and libration near mean-motion resonances, as discussed above, but shorter than the typical protoplanetary disk lifetime and potentially the planetesimal formation timescale. For this reason, we do not model the effects of gas drag, although we consider its effects when constructing initial conditions for our simulations, which is discussed next.

**Table 1.** Summary of Simulations Run

Name	Mass of Planet	Eccentricity of Planet
e1m2	1.0 $M_{Jup}$	0.5 $e_{Jup}$
e2m2	1.0 $M_{Jup}$	1.0 $e_{Jup}$
e3m2	1.0 $M_{Jup}$	2.0 $e_{Jup}$
e2m1	0.5 $M_{Jup}$	1.0 $e_{Jup}$
e2m3	2.0 $M_{Jup}$	1.0 $e_{Jup}$

## 3 SIMULATIONS

### 3.1 Numerical Methods

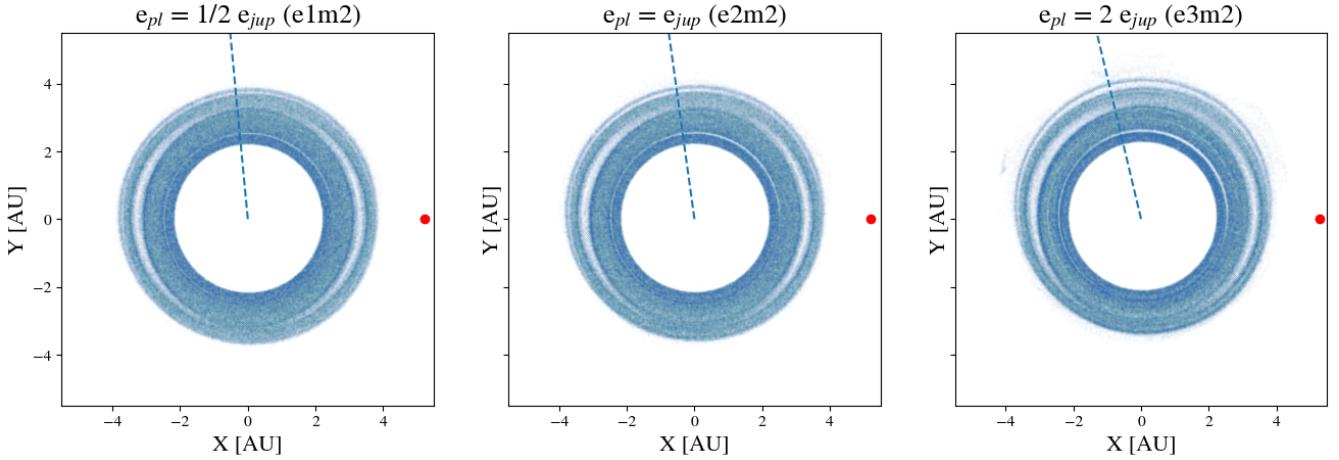
To follow the dynamical and collisional evolution of a planetesimal disc, we use the highly parallel N-body code CHANGA<sup>1</sup>. This code, which is written in the CHARM++ parallel programming language, was originally designed for cosmology simulations and has been shown to perform well on up to half a million processors (Menon et al. 2015). Using similar methods to PKDGRAV, which has been used for numerous studies of planet formation (add some references here), CHANGA calculates gravitational forces using a modified Barnes-Hut (Barnes & Hut 1986) tree with hexadecapole expansions of the moments and then integrates the equations of motion using a kick-drift-kick leapfrog scheme. All of the simulations we perform use a node opening criterion of  $\theta_{BH} = 0.7$ . More information about the code can be found in Jetley et al. (2008).

We have recently modified CHANGA to handle solid-body collisions between particles, by assigning them a fixed radius, rather than treating them as tracers of a fluid with a characteristic softening length. We provide a brief summary of the collision model here, although a full description is provided in Wallace & Quinn (2019). This work is largely based off of the solid-body collision implementation in PKDGRAV, which is described in Richardson (1994) and Richardson et al. (2000). Imminent collisions are found during the drift phase of each time step by extrapolating the positions of bodies forward using the velocity calculated during the first kick. For each body, the nearest 64 neighbors are tested for an imminent collision. If an imminent collision is detected, the two particles of mass  $m_1$  and  $m_2$  are merged together to form one single larger body with the same density and a mass of  $m_1 + m_2$ . The resulting body is then imparted with the center of mass position and velocity of its two children. Because the resolution of a collision can result in another imminent collision, these events must be resolved one at a time, with the soonest collision being resolved first. For this reason, another collision search is run each time a collision is resolved, which continues until there are no more imminent collisions during the current time step.

### 3.2 Initial Conditions

In total, five simulations are run, which are listed in table 1. The second is a “nominal” case, in which the perturbing planet’s mass and eccentricity are set to that of Jupiter’s. In the other four cases, the mass or eccentricity is altered by a factor of two from the nominal value. In all cases, the perturbing giant is placed on a 5.2 au orbit around a 1  $M_\odot$  star. The planetesimal disc extends from

<sup>1</sup> A public version of CHANGA can be downloaded from <http://www-hpcc.astro.washington.edu/tools/ChaNGa.html>



**Figure 1.** The positions of the remaining planetesimals at the end of the e1m2 (left), e2m2 (center) and e3m2 (right) simulations. The red dot indicates the position of the giant planet, and the dashed line points in the direction of the planet's longitude of perihelion. Non-axisymmetric gaps are apparent near the locations of MMRs. At higher eccentricities, more resonances become visible. At the 2:1 MMR, gap features at  $\theta = 0$  and  $\theta = \pi$  follow the giant planet in its orbit.

2.2 to 3.8 au, which covers the two most prominent mean-motion resonances with the giant planet, the 2:1 at 3.27 au and the 3:1 at 2.5 au. The planetesimal disc follows a minimum-mass solar nebula surface density profile (Hayashi 1981)

$$\Sigma = \Sigma_0 r^{-\alpha}, \quad (18)$$

with  $\alpha = 3/2$  and  $\Sigma_0 = 10 \text{ g cm}^{-2}$ . Planetesimals are given a bulk density of  $2 \text{ g cm}^{-3}$  and a diameter of 300 km. **This choice of parameters for the planetesimal disk is meant to follow the setup used in simulation B by Richardson et al. (2000).** One difference from the aforementioned study is that we use a narrower annulus, which reduces the number of particles required from  $10^6$  down to roughly 500,000. This allows us to use a finer base timestep size, which, as we will discuss in section 4, appears to be the reason why no features were seen in the collision profile near the resonances by Richardson et al. (2000). Each simulation is evolved for 5,000 years, which is about 400 orbits of the perturbing planet.

The dynamical effects of secular forcing by the giant planet, along with the effects of viscous stirring and gas drag on the planetesimals, are built into the initial conditions. **Although we do not model the effects of gas drag on the planetesimals, we construct the initial conditions such that these two effects are in balance. The viscous stirring timescale of the planetesimal disc is much longer than our chosen integration time, and so the dynamical excitation of the disc (excluding resonances) stays constant, even without modeling damping from gas.** This is done by first calculating the equilibrium eccentricity  $e_{eq}$  due to viscous stirring and gas drag as a function of semimajor axis according to equation 12 of Kokubo & Ida (2002). Although CHANGA does not account for the effects of gas drag on the planetesimals, the viscous stirring timescale is long compared to our integration time and should not cause the velocity distribution to significantly evolve. The eccentricities of the bodies are drawn randomly from a Rayleigh distribution with a mode of  $e_{eq}$ , while the inclinations are drawn from a similar distribution with a mode of  $e_{eq}/2$  (Ida et al. 1993). The arguments of perihelion  $\omega$ , longitude of ascending nodes  $\Omega$  and the mean anomalies  $M$  of the bodies are drawn uniformly  $\in [0, 2\pi]$ .

To account for the effects of secular forcing by the planet, the eccentricity vectors of the planetesimals are first decomposed into real and imaginary components:

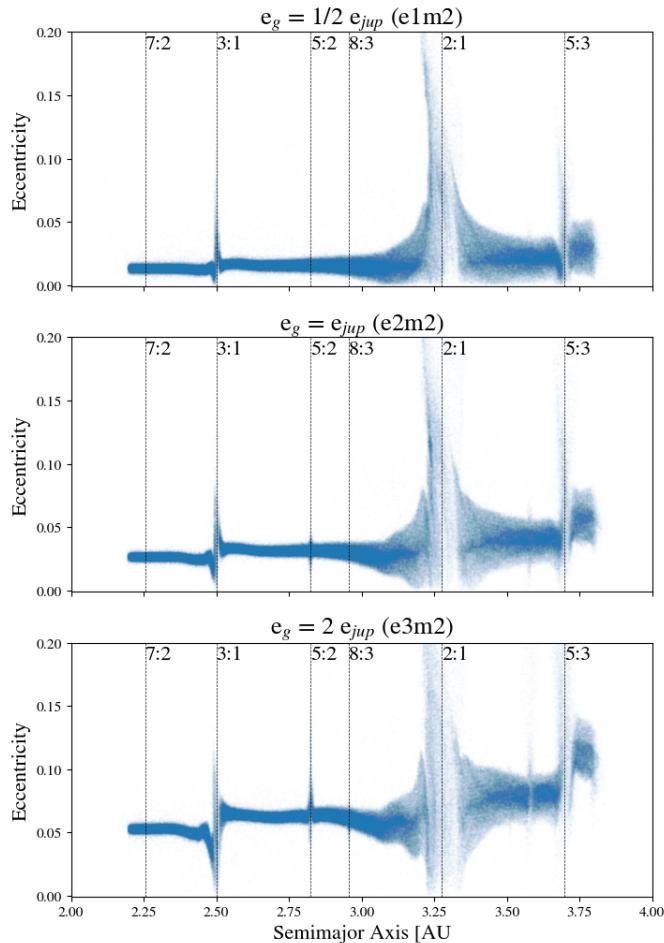
$$z = (k, i\varpi) = e \exp(i\varpi) \quad (19)$$

and a forced component is added to  $h$  according to equation 1 (where we have set  $\varpi_g = 0$ ).

### 3.3 Time Stepping Scheme

For the purposes of the integrator, there are two relevant timescales in this system. The first is the orbital dynamical time  $\sqrt{a^3/GM_\odot}$ . Experimentally, we have found that using a base timestep size with CHANGA of 3% of an orbital dynamical time at the inner edge of the disc keeps the integration symplectic. Although doing so preserves orbital frequencies, the errors associated with precession frequencies do not average out to zero. This is especially important given the effects of secular forcing by the planet. To mitigate this, we further reduce the base timestep size by a factor of 4 and use  $\Delta t = 0.0025 \text{ yr}$ . Doing so prevents the longitude of perihelia of planetesimals at the inner edge of the disc, where this artificial precession is most severe, from drifting by more than the intrinsic spread in  $\varpi$  due to the free eccentricity over the course of our integrations.

An additional timescale is set by the dynamical time of the planetesimals ( $\sim 1/\sqrt{G\rho}$ ), which is about 45 minutes. **This timescale must be resolved in order to properly follow close gravitational encounters between these bodies.** To resolve the base time step and the dynamical timescale of planetesimals simultaneously, we use a two-tiered time stepping scheme following (Leinhardt et al. 2015). To start, all bodies are placed on the base time step. A first pass of collision detection is then run in which the radii of all bodies are inflated by a factor of 2.5. Any bodies with imminent collisions predicted using the inflated radii are placed on a time step that is a factor of 16 smaller than the orbital time step. Although this is still a factor of roughly 7 larger than the dynamical time of a planetesimal, we found no difference in the collision rate



**Figure 2.** The semimajor axes and eccentricities of the remaining planetesimals are shown, with the locations of prominent resonances indicated by the vertical dashed lines. Libration of the critical angle drives large variations in eccentricity, which produce spikes in the a-e plane. Between the resonances, the nonzero eccentricity is due to secular forcing by the planet.

when using any smaller of a minimum time step size. The purpose of the two-tiered scheme is to properly resolve the gravitational interactions between any bodies that undergo a close encounter. This prevents the coarser base time step from reducing the effectiveness of gravitational focusing, while minimizing the additional computational expense.

## 4 RESULTS

All five simulations are evolved for 5,000 years, which is roughly 400 complete orbits of the perturbing planet. Because the effects of the mean motion resonances are not built into the initial conditions, the simulations must be run long enough for the distribution of orbital elements to reach equilibrium near the resonances before the collision rate is measured. **To do so, we allow each simulation to run for 2,000 years before we begin recording any collision statistics. This is comparable to the libration period (calculated in section 2.2 for the 3:1 and 2:1 resonances. Although one might worry that a single libration period may not be enough for the orbital structure due to the resonances to develop, we find that the shape of the semimajor axis-eccentricity distribution of the**

planetesimals reaches a steady state by this time. This is largely due to the fact that there are enough bodies in or near the resonances to allow an ensemble average of phase space, rather than a time average. Additionally, we find that the time evolution of the collision rate near the resonances maintains a steady state in all five simulations after  $\sim 2000$  years.

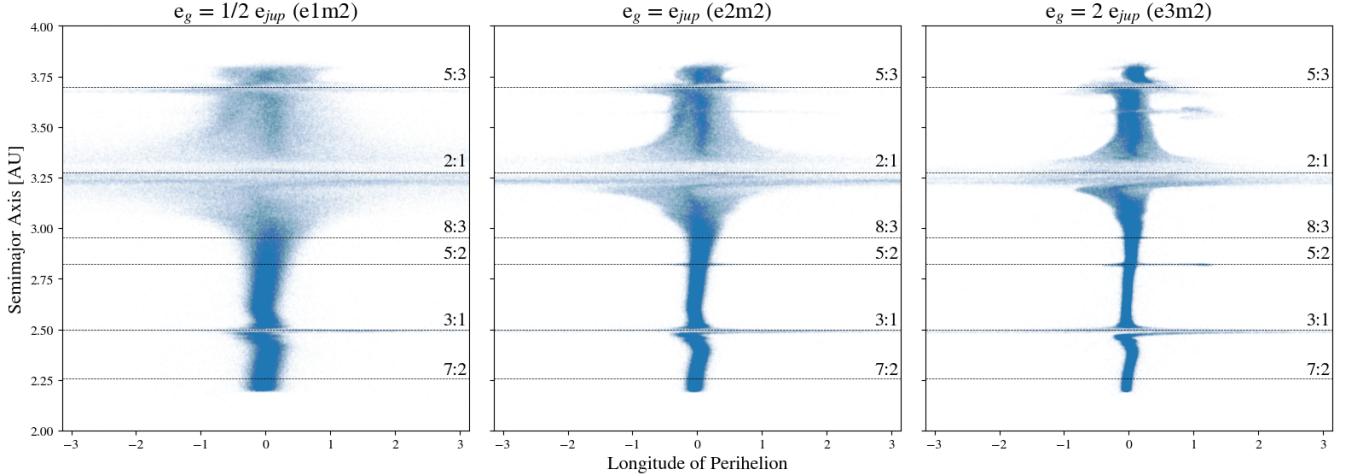
### 4.1 Varying the Eccentricity

We begin by examining simulations e1m2, e2m2 and e3m2. The positions of the planetesimals in the x-y plane after 5,000 years of integration are shown in figure 1. In all cases, the coordinate system is rotated so that the giant planet lies at  $\theta = 0$ . The longitude of perihelion of the planet is shown by the dashed line. Resonances with the perturbing planet are visible as nonaxisymmetric gaps **in this figure**. Upon close inspection of a series of simulation snapshots, these gaps appear to follow the planet in its orbit, rather than aligning themselves with the longitude of perihelion. A similar substructure, which is most noticeable near the 2:1 MMR, reveals itself in Richardson et al. (2000) (see figure 3c and 3f) and Tabeshian & Wiegert (2016) (figure 3a). This structure is also present in the Boley (2017) simulations, although it was not reported at the time (see Appendix A). It is worth noting that Richardson et al. (2000) started with a completely cold planetesimal disc and a Jupiter mass planet on a circular, coplanar orbit. Boley (2017) used a similar setup. Thus, the presence of these features seems robust to the choice of initial conditions.

The effects of the resonances become much more apparent in semimajor axis-eccentricity space, which is shown in figure 2. In all cases, the 3:1, 2:1 and 5:3 resonances are readily visible as “spikes” in the eccentricity that bend slightly inward (due to the conservation of  $E_J$ ). In the e3m2 simulation, features also appear near the 5:2, 7:3 and 5:3 resonances. The absence of these features from the e1m2 simulation can be explained by the fact that the strength of a resonance scales with  $e^q$  (Malhotra 1994).

Another important effect of the resonances is visible in figure 3, which shows the orientation of the longitude of perihelia of planetesimals in the disc. Inside of the resonances, orbits of planetesimals quickly precess, and their orientations are effectively randomized. An important point to note is that this strong precession effect quickly disappears beyond the boundaries of the resonance. This turns out to be key to explaining the nonaxisymmetric structure seen in figure 1, which will be addressed in more detail below.

Next, we examine the statistics of collisions resolved in each of the simulations. The 3D positions and velocities of the two colliding bodies are recorded to a table at the moment of impact. We derive the Keplerian orbital elements of a collision using these positions and velocities. First, we examine the semimajor axis of the first collider, which is shown in figure 4. During a collision, the “first” particle is defined as the more massive of the two. However, nearly all of the collisions happen between the initial, equal-mass planetesimals. In this case, the distinction is set by the collision search algorithm and is rather arbitrary. In all of the plots where we show collision statistics, we have verified that using the “first” or “second” collider does not qualitatively change any of the features. We find that some of the features present in figure 4 in this and subsequent figures are highly sensitive to the number of bins and the location of the bin edges. For this reason, we construct a **probability density function (PDF)** of the collisions using a Kernel Density Estimate (KDE). We use the NEIGHBORS.KERNELDENSITY function from the SKLEARN (Pedregosa et al. 2011) package to construct our KDEs, using a gaussian kernel



**Figure 3.** Shown here are the longitudes of pericenter and semimajor axes of the remaining planetesimals. The dashed lines again indicate the locations of prominent MMRs. Close to the resonances, the critical angle librates and drives fast precession of the longitudes of pericenter of the bodies. This overpowers the secular forcing by the planet and effectively randomizes the orientations of their orbits.

with a FWHM of 0.02 au. The curves shown in figure 4 are normalized such that the area underneath is normalized to 1.

Near the stronger resonances, there are noticeable suppressions or enhancements to the local collision rate. This contrasts with the findings of Richardson et al. (2000), who simulated a similar setup and found no discernible features near the MMRs. We attribute the differences to a more conservative timestepping criterion in our simulations. The most prominent features appear as a local maximum in the collision rate near the 2:1 MMR and a local minimum near the 3:1 MMR. At higher forced eccentricities, features near the 5:2, 8:3 and 5:3 resonances are also visible, due to the steep sensitivity of higher order resonant perturbations to eccentricity (Malhotra 1994).

Although the features in figure 4 near the 3:1 and 2:1 MMRs appear qualitatively different, they can be explained by one single dynamical process. As discussed in section ??, the collision rate between planetesimals depends on both the encounter rate and the strength of gravitational focusing. The encounter rate grows as the relative velocity  $v$  between the planetesimals increases, while gravitational focusing is most effective for bodies with small relative velocities. There is therefore an intermediate value of  $v$  for which the average collision rate  $\langle \sigma v \rangle$  is at a minimum, which we will refer to as  $\langle \sigma v \rangle_0$ . The dynamical excitation induced by a mean motion resonance can either increase or decrease the local average collision rate, depending on the size and mass of the planetesimals, along with their unperturbed relative velocities.

In figure 5, we show the effect that the 2:1 MMR and 3:1 MMR have on the local average collision rate. Each pair of points connected by a line represents the average collision rate before (from the initial condition file) and after (around  $T = 2,000$  yr) the effects of the mean-motion resonances develop. In this figure, the relative velocity between bodies is calculated by measuring the eccentricity and inclination dispersion near each resonance and using equation 16. Here, we assume that the eccentricity and inclination dispersions are coupled (Ida et al. 1993) and use the eccentricity as the independent variable. Because both the strength of secular forcing and the effects of the resonant perturbations are different near the 3:1 and 2:1 MMRs, the

planetesimal population at each of these locations in the disk experiences a qualitatively different change in the collision rate. Near the 3:1 MMR, the unperturbed collision rate is well below the minimum value  $\langle \sigma v \rangle_0$  and the dynamical heating introduced by the resonance acts to suppress the collision rate further. Further out in the disk, near the 2:1 MMR, the collision rate starts out much closer to the minimum value and the additional perturbations drive the collision rate upwards.

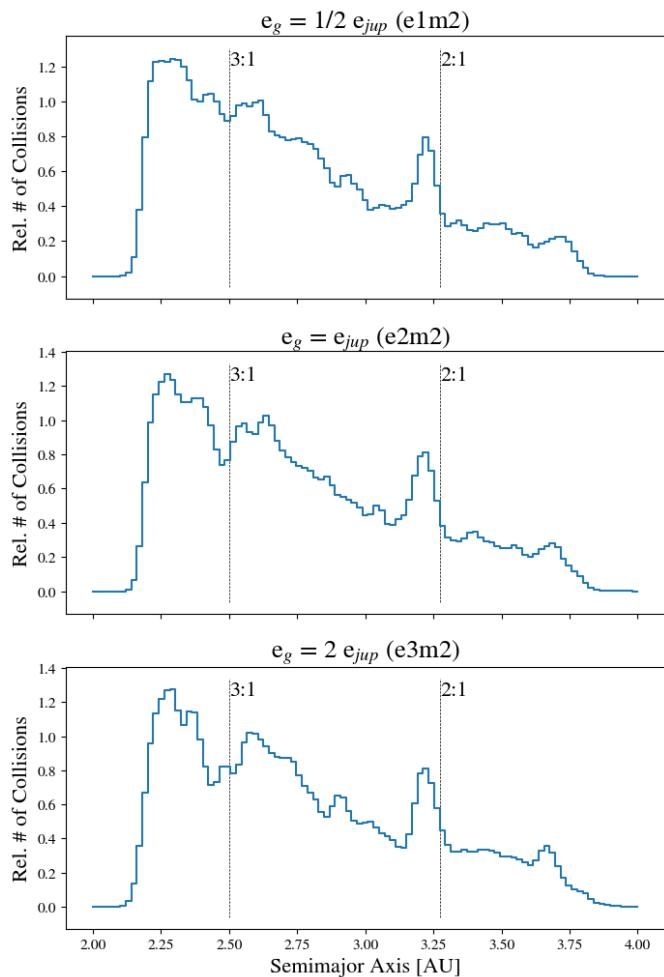
This explains the relative drop in the collision rate near the 3:1 MMR and the relative increase near the 2:1 and 5:3 resonances seen in figure 4. Although this could potentially serve as a useful diagnostic of the planetesimal size in a disc (because the size and mass of the planetesimals set the eccentricity dispersion at which  $\langle \sigma v \rangle$  is minimized), a direct measurement of the semimajor axes of the colliding planetesimals is not possible. Furthermore, as we will demonstrate next, collisions due to bodies in resonance do not have a significant effect on the final shape of the radial dust distribution.

## 4.2 Where Does the Dust End Up?

**TODO:** Introduce this a little more gently, start off by acknowledging the limitations of this model To construct a radial dust profile from the collision statistics, we begin by making the assumption that any dust generated by collisions is strongly coupled to the gas. For a dust grain larger than the mean free path of the gas particles, the timescale to damp its eccentricity is given by equation 17. At 3 au, the gaseous component of the solar nebula has a density of  $4 \times 10^{-10} \text{ g cm}^{-3}$  and a typical thermal speed of  $10^5 \text{ cm s}^{-1}$  (Hayashi 1981). For a 1 mm dust grain with a density of  $1 \text{ g cm}^{-3}$  and an eccentricity of 0.05, the damping timescale is  $4 \times 10^{-3}$  years. This is orders of magnitude smaller than the orbital timescale, therefore we conclude that collisionally generated dust grains will immediately couple to the gas.

Operating under this assumption<sup>2</sup>, a map of the relative con-

<sup>2</sup> As discussed in Boley (2017), the local generation of second-generation dust is only one consideration. The dust will also evolve spatially due to drag



**Figure 4.** A PDF of the collision rate in each disc as a function of semimajor axis, generated using a KDE with a Gaussian kernel with a FWHM of 0.02 au. In semimajor axis space, prominent features appear near the 3:1 and 2:1 MMRs. Near the 3:1, the collision rate decreases from its unperturbed value, while an enhancement appears near the 2:1.

centration of second-generation dust can be constructed from the cylindrical distance at which the collisions occur. **Here, cylindrical distance is measured as the distance between the central star and the collision in the  $(r, \theta)$  plane.** This is shown in figure 6. Similar to figure 4, we use a KDE with a width of 0.02 au to assemble the collision locations into a radial distribution. Most strikingly, the bump that was present near the 2:1 resonance is no longer visible in cylindrical distance space, although it suddenly appears again in the e3m2 simulation. The dip that is visible near the 3:1 MMR also becomes a bump in the highest eccentricity case. To determine how the resonant bodies actually contribute to the radial collision profile, we excluded collisions that fall between  $2.495 < a < 2.505$  au (near the 3:1 resonance) and between  $3.2 < a < 3.35$  au (near the 2:1 resonance), which is shown by the dashed curve. **As in figure 4, the solid curves are normalized such that the area underneath is equal to 1. For the dashed curves, the normalization factor is**

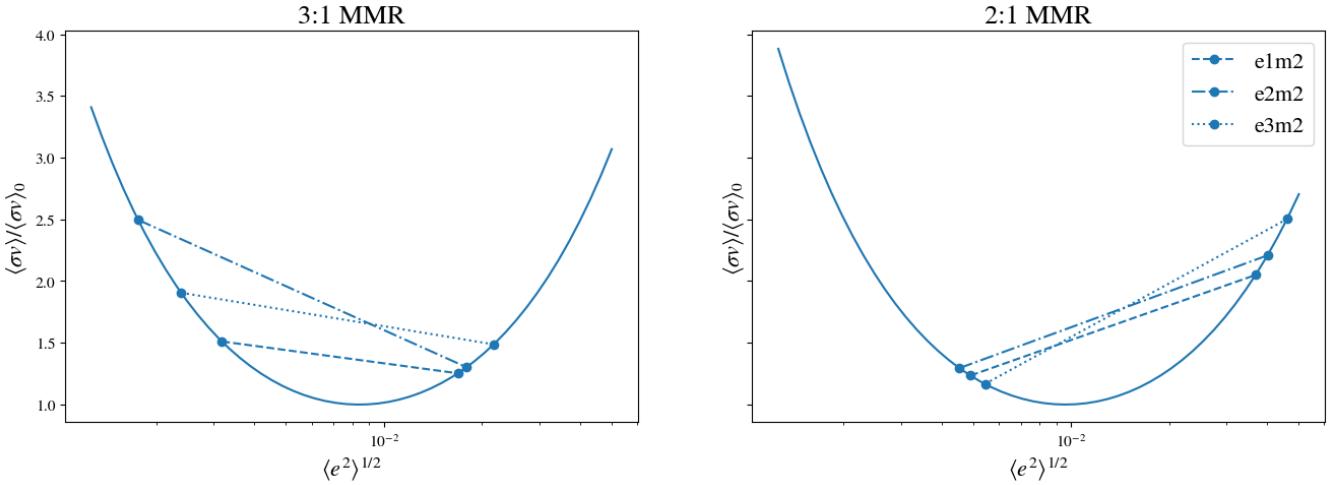
effects, but can also evolve in size or be re-accreted onto planetesimals. For simplicity, we focus on local dust sourcing only while acknowledging that subsequent evolution could affect the morphology.

scaled down according to the total number of collisions in the entire sample and the number of collisions in the subsample which excludes the collisions in resonance. Qualitatively, none of the bump or dip features present are removed by making this exclusion. This suggests that the features near resonance seen in semimajor axis space become smeared out in cylindrical distance space due to the large spread in eccentricity and orbital orientation.

This also suggests that the prominent bumps and dips seen in figure 6 are mainly produced by bodies outside of resonance, although the resonant population does make some of the features more pronounced. To further understand this, we consider the non-axisymmetric structure seen in figure 1. We compare in figure 7 the radial collision profile to the radial structure seen in the e1m2, e2m2 and e3m2 simulations. In the lowest forced eccentricity case ( $e_{\text{1m2}}$ ), the pileups near the edge of the 2:1 MMR line up with the boundaries of the dip feature seen in the radial collision profile. The locations of the edges of the resonances are calculated using equations 12 and 8. The pileups not as pronounced near the other resonances. As discussed previously, the planetesimal density enhancements and gaps seen in polar coordinates follow the position of the giant planet in its orbit, which is shown by the vertical line. For the e2m2 case, the density enhancements take up a much larger radial width, which smears out the corresponding collisional dust profile. In the highest forced eccentricity case (e3m2), the density enhancements occur well within an annulus containing the resonant region, completely altering the radial peak-valley dust morphology.

Although the nonaxisymmetric pileup of bodies near the edges of resonances has been seen previously (Richardson et al. 2000; Tabeshian & Wiegert 2016), we could not find a satisfactory explanation for why it happens, which we will try to provide here. As mentioned in section 2, the circulation frequency of the critical angle slows as one approaches the edge of the resonance from the outside. Following the pendulum analogy, this is equivalent to approaching the point where the pendulum becomes suspended in the at top dead center. When the critical angle stays relatively stationary during an encounter, the net torque will drive the longitude of pericenter towards the point of conjunction (Peale 1976). This is the basic mechanism that causes isolated resonances to be stable. Inside of resonance, the critical angle librates about an equilibrium value and this configuration is maintained. Just outside of resonance, however, the critical angle slowly drifts away, with the drift direction changing sign across the location of exact resonance. The net result of this is that bodies just outside of the resonance experience a torque from the planet that temporarily drives their pericenters toward the current orbital position of the planet. After the encounter ends, differential rotation of the disc slowly destroys the configuration over part of an orbit. As the planet passes by, adjacent planetesimals near the resonance edges temporarily align themselves in this configuration. This explains why the concentration of planetesimals appears to ‘follow’ the planet.

Tabeshian & Wiegert (2016) provide a different explanation for this phenomenon, suggesting that the nonaxisymmetric structure is the product of the path that a resonant test particle takes in a frame co-rotating with the planet (see figure 8.4 of Murray & Dermott (1999)). This explanation, however, does not hold for a large collection of planetesimals, especially when there is no forced eccentricity. In such a case, the orbits of planetesimals near resonance would be randomly aligned, and the axisymmetric structure should become washed out. When the perturbing planet is on a circular orbit, this structure is still present (see figure 3 of Tabeshian & Wiegert (2016) and Appendix A). The only connection between the structure of the gaps in the disc and the path of a resonant particle



**Figure 5.** The collision rate of bodies in the vicinity of the 3:1 (left) and 2:1 (right) MMRs, relative to the minimum value, as a function of the local eccentricity dispersion. The pairs of points connected by lines show the values of the perturbed and unperturbed collision rates in the e1m2, e2m2 and e3m2 simulations.

in the corotating frame is that both phenomena produce the same symmetries, which depend on the particular MMR.

With no secular forcing, this dynamical phenomena should produce an underdensity in the radial surface density profile near the center of the resonance and an overdensity at each edge in cylindrical distance space. When a forced eccentricity is introduced, the radial location of the resonance edges vary over the course of an orbit. The maximum and minimum radial distance that the resonance edges occupy is indicated by the solid and dashed lines in figure 7. If the the aphelion distance of the inner edge becomes close to the perihelion distance of the outer edge, the overdense regions on each side of the resonance meet and we should expect to see the central dip feature in the collision profile disappear. This is exactly what appears to be happening in the middle panel of figure 7. As the forced eccentricity is increased further, a region forms at the center of the resonance where planetesimals on both sides of the resonance spend time (although not simultaneously). When this occurs, a bump feature forms in the collision profile near the center of the resonance. This is apparent in the bottom panel of figure 7.

Although the same phenomenon appears to be happening near the 3:1 MMR, the pileups at the edges of the resonance are not as well defined. We attribute this to the more complex symmetry, compared to the 2:1 MMR. In addition, a number of other nearby resonances, including the 7:2 (at 2.25 au), the 10:3 (at 2.33 au), the 8:3 (at 2.71 au) and the 5:2 (at 2.82 au) also likely are contributing to the collision profile in cylindrical distance space near the 3:1 MMR. The 3:1 MMR is also much narrower than the 2:1, which means that any dip or bump features produced by it would require much higher resolution observations. For these reasons, we will focus on the 2:1 MMR as the main diagnostic indicator.

#### 4.3 Varying the Mass

The dynamical effects of varying the mass of the planet are somewhat simpler, in that doing so does not affect the forced eccentricity of the planetesimal disc. Instead, only the width of the resonances changes. The width of a first order resonance scales with  $m$  (see equation 8), while the leading order terms in the resonant part of the disturbing function, which set the strength of the resonance, also

scale as  $m$ . For the 2:1 MMR, the dynamics near the resonance are equally sensitive to changes in eccentricity and mass.

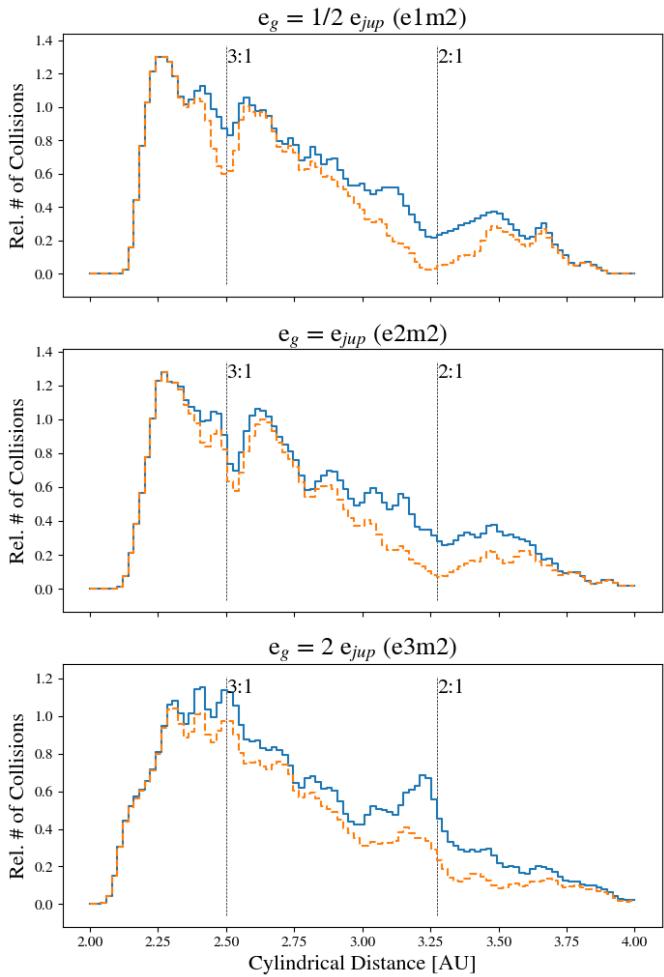
We show the polar structure of the e2m1, e2m2 and e2m3 simulations in figure 8 alongside the radial collision profile. In all three cases, the eccentricity of the perturbing planet is set to  $e_{jup}$ . Changes in the apocenter and pericenter distances of the edges of the resonances are entirely due to changes in the resonance width. For the e2m1 and e2m2 simulations, the inner apocenter and outer pericenter distances near the 2:1 MMR are quite similar and no strong features appear in the collision profile near this region. For the e2m3 case, the edges of the resonances are sufficiently separated to allow what appears to be the beginning of gap to form near the center of the 2:1 resonance in the collision profile.

#### 4.4 Observability of the Dust

While the collisional dust profiles show radial amplitude variations, we still need to assess whether those variations could be observable. To proceed, we create a sky model for simulated ALMA observations as follows. First, we use the radially-averaged collision profile from each simulation (see figures 7 and 8) as the template for an azimuthally symmetric disc. We then scale the size of each profile by a factor of ten, which places the perturbing planet at 52 au. This scaling is permitted by the dynamics (both the resonance widths and the forced eccentricities are scale-free) and makes the disc comparable in size to the many discs that have now been observed by ALMA (Huang et al. 2018). The angular size scale is then set by envisaging a face-on disc at a distance of 100 pc.

The sky model intensity (flux per pixel) is produced by interpolating the radial collision frequency onto a 2D Cartesian grid with cell widths of approximately 2 mas. In doing so, we choose to make the intensity proportional to the collision frequency (i.e., the dust) because it is the most straightforward case. Many different parameterizations are possible, depending on optical depth, grain size, disc temperature, etc., but we will avoid such complications here. We normalize the intensity by setting the total disc flux density to be 100 mJy at about 350 GHz.

The simulated observations are performed using the SIMOB-SERVE task in CASA (McMullin et al. 2007). The disc is given an



**Figure 6.** Here, collisions are instead ordered by cylindrical distance from the star. Features near the 3:1 and 2:1 MMRs are still present, but appear qualitatively different than in semimajor axis space. At low eccentricities, a dip appears around the center of the resonance. At the highest eccentricity, a bump is formed instead. The dashed lines shown the collision profile with collisions between bodies in resonance removed. The bump and dip features remain qualitatively the same, which suggests that they are produced by bodies outside of resonance.

RA =  $11^{\text{h}}01^{\text{m}}52^{\text{s}}$  and a  $\delta = -34^{\circ}42'17''$  (i.e., the J2000 coordinates for TW Hydrae) and is “observed” through transit on March 20, 2020. We observe the disc using configurations 8, 9, and 10 at 350 GHz (cycle 6 antenna file), spending six hours on-source in each configuration<sup>3</sup>. The visibilities are corrupted with thermal noise, setting the precipitable water vapor to 1.5 mm. They are then combined, imaged, and cleaned using the task `tclean`. Imaging uses Briggs weighting with a robustness parameter of -1.

The results for the nominal simulation (e2m2) are shown in figure 9. The 3:1 and 2:1 resonances can be easily identified in the cleaned image and in the corresponding radial profile. Moreover, the features are qualitatively similar to the bright rings and dark gaps seen in actual disc profiles (ALMA Partnership et al. 2015).

In figure 10, we show the azimuthally-averaged radial profiles constructed from the simulated cleaned image from all five simula-

tions. In all cases, a bump or a dip feature is clearly visible at the location of the 3:1 and 2:1 MMRs, indicated by the dashed vertical lines. As mentioned previously, the feature at the 2:1 MMR acts as the actual diagnostic indicator, while the 3:1 resonance (along with the gap that would presumably open at the location of the planet) is mainly useful for determining where the 2:1 MMR is actually located.

## 5 CONSTRAINING THE MASS AND ECCENTRICITY OF THE PLANET

The simple bump vs dip structure that we expect to reveal itself in the dust emission from colliding planetesimals in near-resonance with a giant planet could potentially be used to place constraints on the mass and eccentricity of the planet. Near the 2:1 MMR, the central dip feature is only produced when bodies at the edges of the resonance stay sufficiently separated in cylindrical distance over the course of an orbit. This is achieved when either the resonance width is large or the forced eccentricity is small. The bump feature, on the other hand, is produced when there is a sufficiently large amount of overlap between the apocenters and pericenters of the inner and outer edges of the resonance, respectively. This is achieved when the resonance is narrow and the forced eccentricity is large.

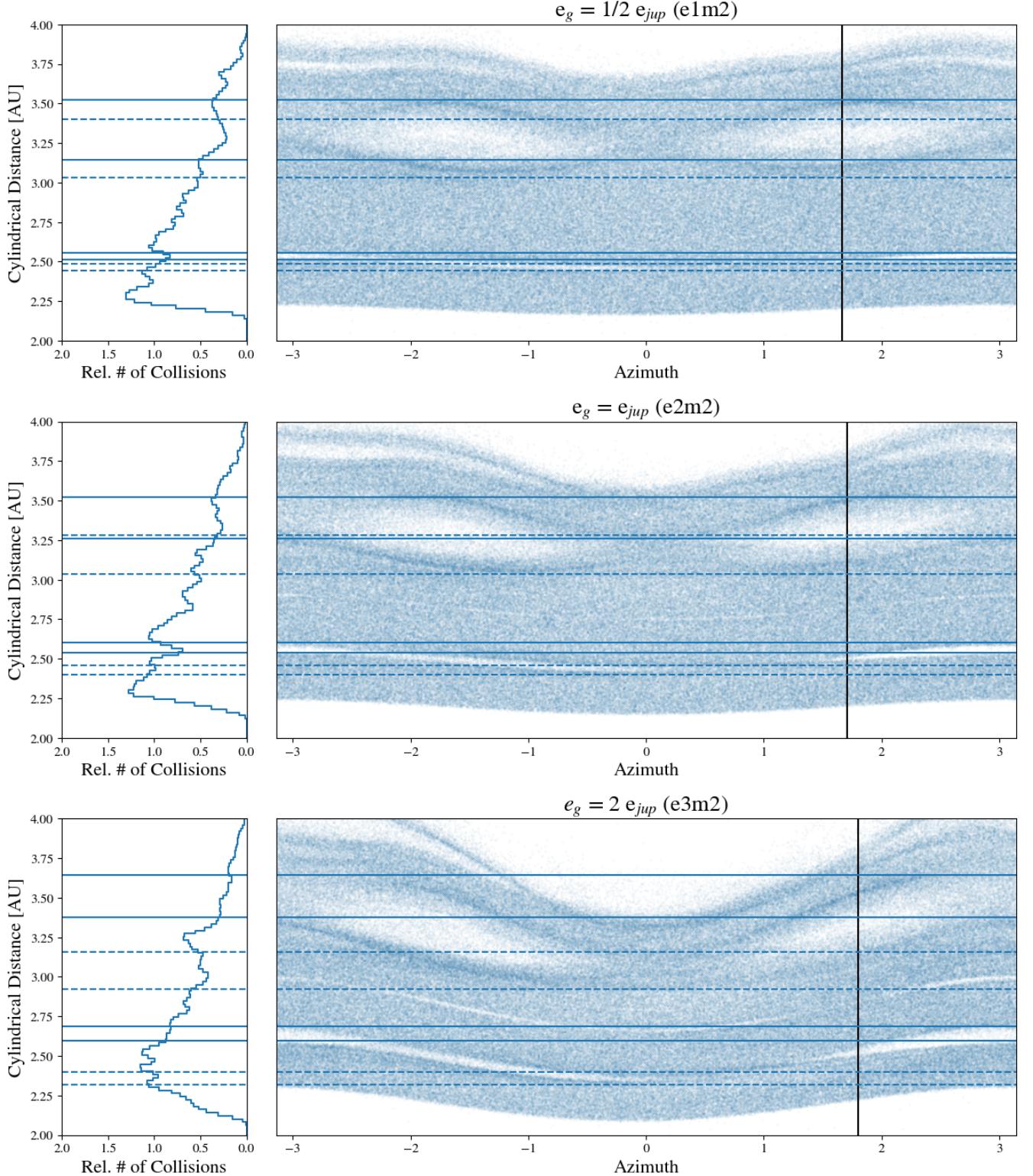
In figure 11, we show the constraints that the presence of a bump or a dip or dip at the 2:1 MMR places on possible values of  $(m_g, e_g)$ . The dashed lines indicate regions of parameter space where qualitatively different features in the collision profile are expected to form. Above the upper dashed line, the resonance is wide while the radial excursion distance of planetesimals at the edge of the resonance is small, which gives rise to a dip in the collision profile. Below the lower dashed line, the resonance is narrow, while the radial excursion distance is large, allowing the apocenter and pericenter distances of the edges of the resonance to overlap and produce a bump in the collision profile. Between the dashed lines, the absolute value of separation between the apocenter and pericenter distances of the inner and outer edges of the resonance is less than the resonance width. In this region of parameter space, no significant features in the collision profile near the 2:1 MMR are expected to form.

Because the strength<sup>4</sup> of the 2:1 resonance scales with  $e_g$  and  $m_g$ , a sufficiently low mass and low eccentricity planet will not create enough of a perturbation in the planetesimal disc to produce a detectable bump or dip feature. The colored contours in figure 11 indicate the relative strength of the resonance as these quantities are varied. Although the resonance strength is equally sensitive to changes in mass and eccentricity, the peri-apocenter overlap of the resonance edges is more sensitive to changes in eccentricity. This is because changes to eccentricity affect both the resonance width and the forced eccentricity. As we showed in figure 8, changing  $m_g$  by a factor of 4 produced very minimal changes to the resulting dust profile.

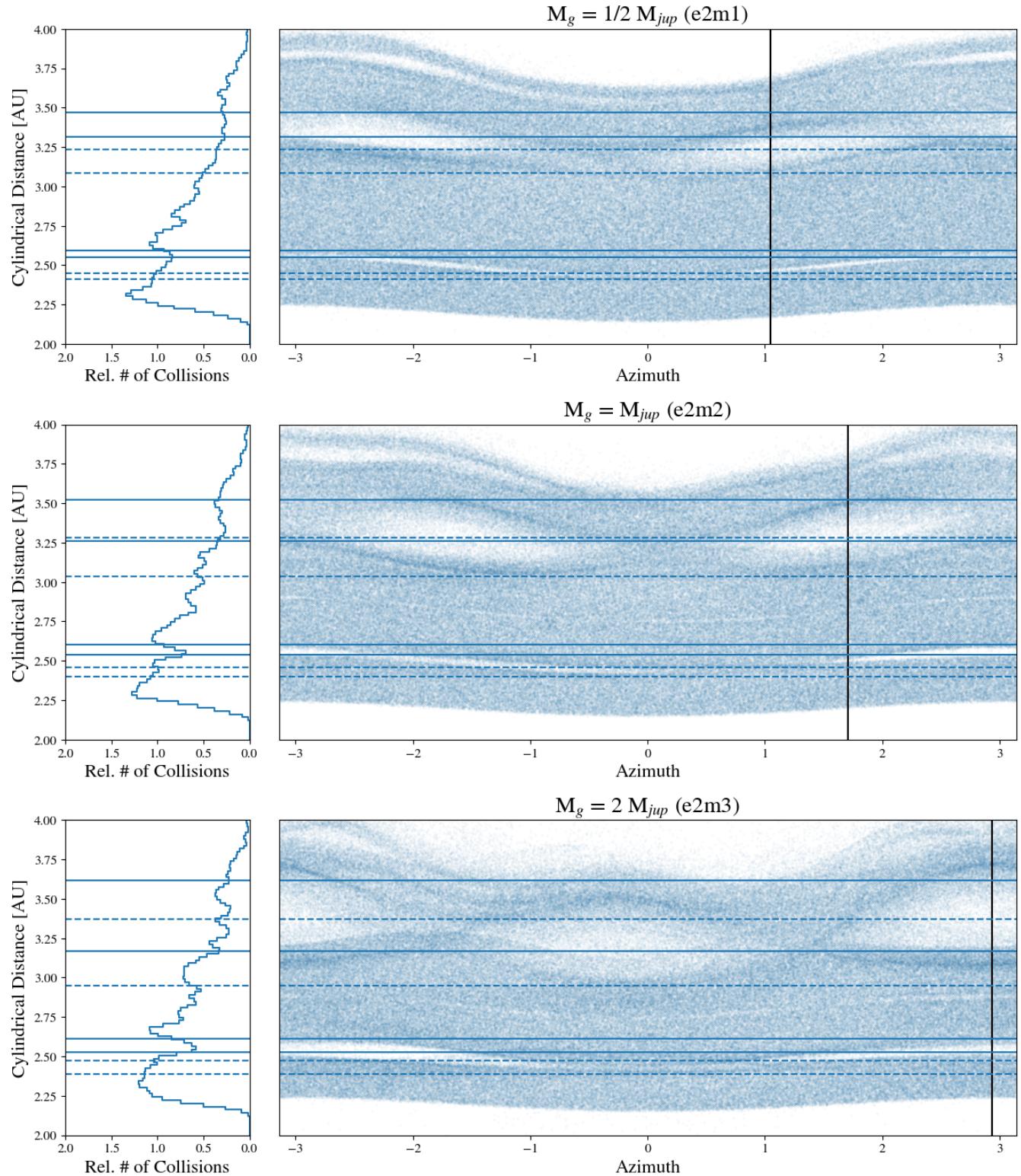
In order to actually identify the 2:1 MMR in the dust emission, at least one other resonance must be visible. Assuming the gravitational field in the disc is near-Keplerian, the distance ratios between two features can be used to determine period ratios, which can be used to confirm whether the features seen are indeed resonances. Although the 3:1 MMR does not appear to follow the simple bump

<sup>3</sup> We have imagined that the TAC really likes us.

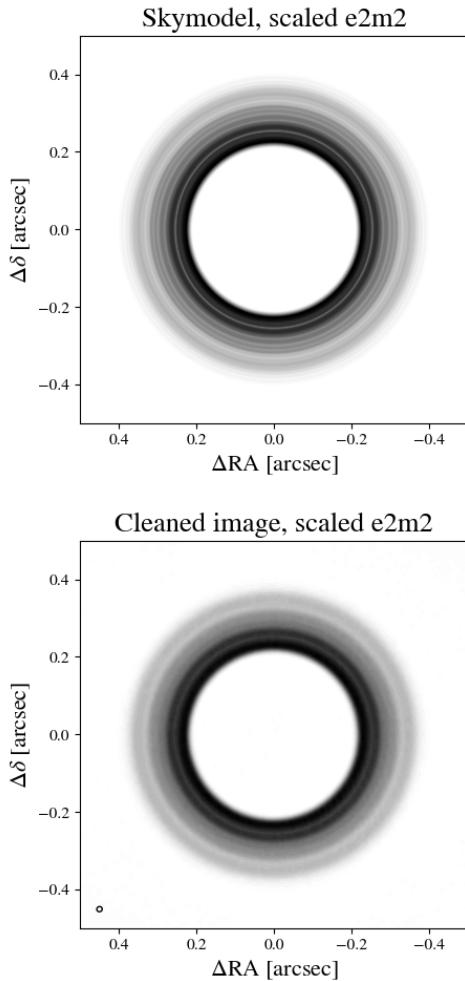
<sup>4</sup> More specifically, the strength of the resonant term in the disturbing potential



**Figure 7.** The collision profiles in figure 6 are shown alongside the positions of the planetesimals, in polar coordinates. The dashed and solid lines show the pericenter and apocenter, respectively, of bodies at each edge of the 3:1 and 2:1 MMR. In these figures, the longitude of pericenter of the planet lies at  $\theta = 0$  and the present position of the planet is indicated by the vertical line. For the 2:1 resonance, a bump, rather than a dip feature appears when the inner edge or the resonance's apocenter and outer edge's pericenter distance cross. The collision profile near the 3:1 resonance does not appear to follow this same qualitative behavior.



**Figure 8.** Similar to figure 7, except the eccentricity of the planet is kept constant and the mass is varied. This has the effect of changing the width of the resonances, without altering the relative apo or peri distances of bodies near the edges. Except for the highest mass case, the apocenter and pericenter distances of the inner and outer edges of the 2:1 resonance are too close together to produce much of a dip or a bump feature.



**Figure 9.** Simulated ALMA observations for the nominal (e2m2) case. Top: The sky model based on the radial collision distribution. The size of the disc has been scaled by a factor of 10 and placed at a distance of 100 pc. Bottom: The cleaned image using combined observations in configurations 8, 9, and 10. The circle in the bottom left indicates the simulated beam size. Gaps in the dust due to the 3:1 and 2:1 resonances are visible.

vs dip dichotomy described above, its presence should be marginally detectable in all of the cases shown above.

## 6 SUMMARY AND CONCLUSIONS

In this work, we have shown that mean-motion resonances with a perturbing planet produce significant local variations in the collision rate of a planetesimal disc. In contrast to Richardson et al. (2000), we find that the more prominent interior MMRs, including the 2:1, 3:1 and 5:3, all produce structure in the collision profile as a function of semimajor axis. Furthermore, we find that a series of distinctly different features appear when collisions are ordered by cylindrical distance from the central star. If enough gas from the protoplanetary disc is still present, dust generated from planetesimal-planetesimal collisions will immediately fall onto a circular orbit. In this case, the radial dust emission profile, which can be observed at sub-mm wavelengths, directly traces the radial collision rate and therefore the dynamical activity of the planetesimals.

Near the interior 2:1 MMR, we find that a distinct bump or dip feature is generated in the dust profile, depending on the properties of the perturbing planet. The presence of one of these two structures can be used to constrain the mass and eccentricity of the planet. For a high mass, low eccentricity planet, a dip will form in the dust profile because the edges of the resonance, where many collisions occur, stay sufficiently separated. If the planet has a low enough mass (which shrinks the size of the resonance) or is sufficiently eccentric (which decreases the separation between the apocenter and pericenter distances of the edges of the resonance), a bump will instead form. We tested this hypothesis for five different combinations of the planet's mass and eccentricity and found that a distinct bump or dip signature is produced as long as the planet properties are sufficiently far from the dividing line in parameter space. This diagnostic is more useful for massive, eccentric planets, because the strength of the resonant perturbations scale linearly with both of these quantities.

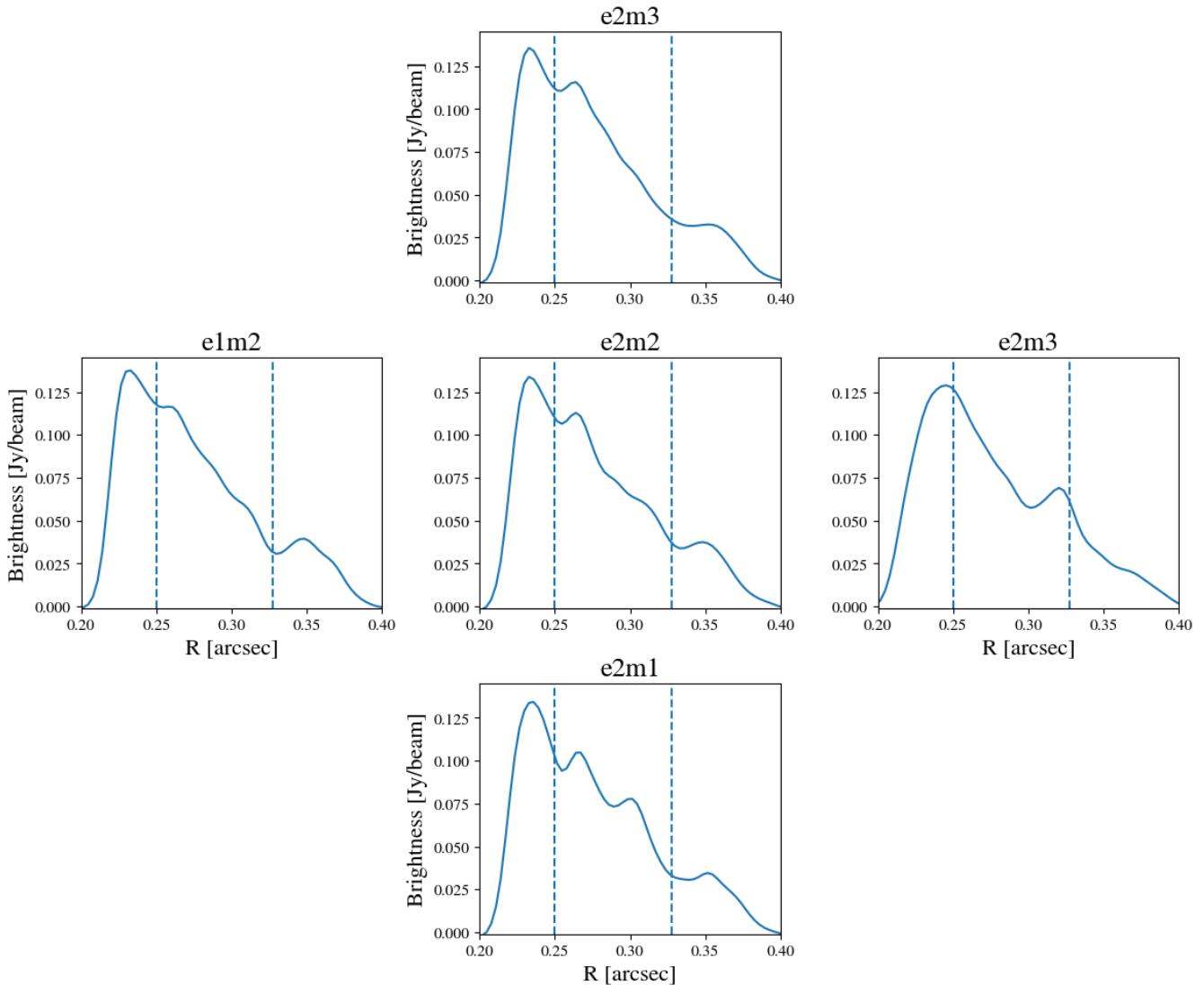
In line with Boley (2017), the results of this work suggest that, in planet forming discs, the observed dust structure should not be interpreted as simply the result of primordial dust that is perturbed by the planets and gas. Instead, some morphological features could be the product of collisional dynamics between larger bodies. Within the confines of the simulations presented here, we have shown that this second-generation dust can give rise to significant features in the dust emission profile. For this reason, we emphasize that planet-forming discs should be thought of as existing along an evolutionary continuum, with the youngest discs being dominated by primordial dust and gas and the most evolved discs being comprised of mainly collisionally generated dust.

The results presented here appear to be broadly consistent with the planetesimal dynamics seen by Tabeshian & Wiegert (2016). In their case, a dip feature can be seen in the azimuthally averaged surface density profile, although they did not test a high enough eccentricity planet in any of their simulations to produce a bump. Another thing to note is that the gap morphology seen by Tabeshian & Wiegert (2016) was markedly different for the 2:1 exterior MMR. Whether this would alter the radial collision profile for the exterior, rather than the interior 2:1 resonance, is not immediately clear. Given that the solid surface density diminishes with distance ( $\sim r^{-3/2}$  for the MMSN (Hayashi 1981)), exterior resonances may not be quite as useful for this type of measurement, however.

Compared to using the corotating gap opened by the planet as a diagnostic (Dobinson et al. 2013, 2016), measuring variations in the dust emission near MMRs is much more subtle and requires much higher spatial resolution and sensitivity. As we have shown, the bump vs dip feature near the 2:1 MMR is marginally detectable with ALMA for the nearest protoplanetary discs. Another complicating factor is that the inner  $\sim 10$  au of most protoplanetary discs are optically thick in the sub-mm. The solution to this problem is to instead observe at longer wavelengths where the inner disc is optically thin. This comes at the expense of much poorer resolution. However, future radio facilities like the NG-VLA are expected to achieve sub-AU resolution for nearby planet-forming discs (Ricci et al. 2018). In the more near-term, techniques like Gaussian process fitting present a promising way to recover substructure at sub-beam resolution (Jennings et al. 2020).

## ACKNOWLEDGEMENTS

This work was facilitated through the use of advanced computational, storage and networking infrastructure provided by the Hyak



**Figure 10.** The azimuthally-averaged radial profile based on the simulated cleaned images from all five simulations. The eccentricity of the perturbing planet increases from left to right, while the mass of the planet increases from bottom to top. The vertical dashed lines indicate the locations of the 3:1 and 2:1 resonances. As in figure 9, the size scale of the disc has been expanded by a factor of 10 and then placed at a distance of 100 pc.

supercomputer system at the University of Washington. We acknowledge the people of the Dkhw'Duw'Absh, the Duwamish Tribe, the Muckleshoot Tribe, and other tribes on whose traditional lands we have performed this work.

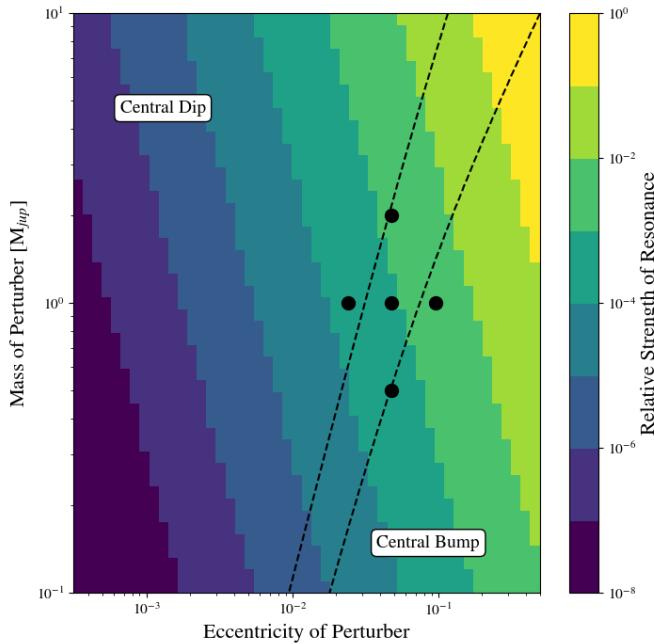
**Software:** Astropy (Astropy Collaboration et al. 2013), CASA (McMullin et al. 2007), CHANGA (Jetley et al. 2008; Menon et al. 2015), Matplotlib (Hunter 2007), NumPy (van der Walt et al. 2011), Pandas (Wes McKinney 2010), PYNBODY (Pontzen et al. 2013), SKLEARN (Pedregosa et al. 2011)

## DATA AVAILABILITY

The data presented in this article are available on our GitHub repository, at [https://github.com/spencerw/planetesimal\\_coll\\_paper](https://github.com/spencerw/planetesimal_coll_paper).

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**Figure 11.** The presence of a dip or bump near the 2:1 MMR can be used to constrain the mass and eccentricity of the perturbing planet. The color scale indicates the relative strength of the features, while the dashed lines indicate boundaries in parameter space where a dip or a bump will be produced at the resonance. Combinations of mass and eccentricity that fall between the dashed lines correspond to an inner apocenter and outer pericenter separation at the edges of the resonance that is smaller than the resonance width and therefore will not create any significant feature in the collision profile.

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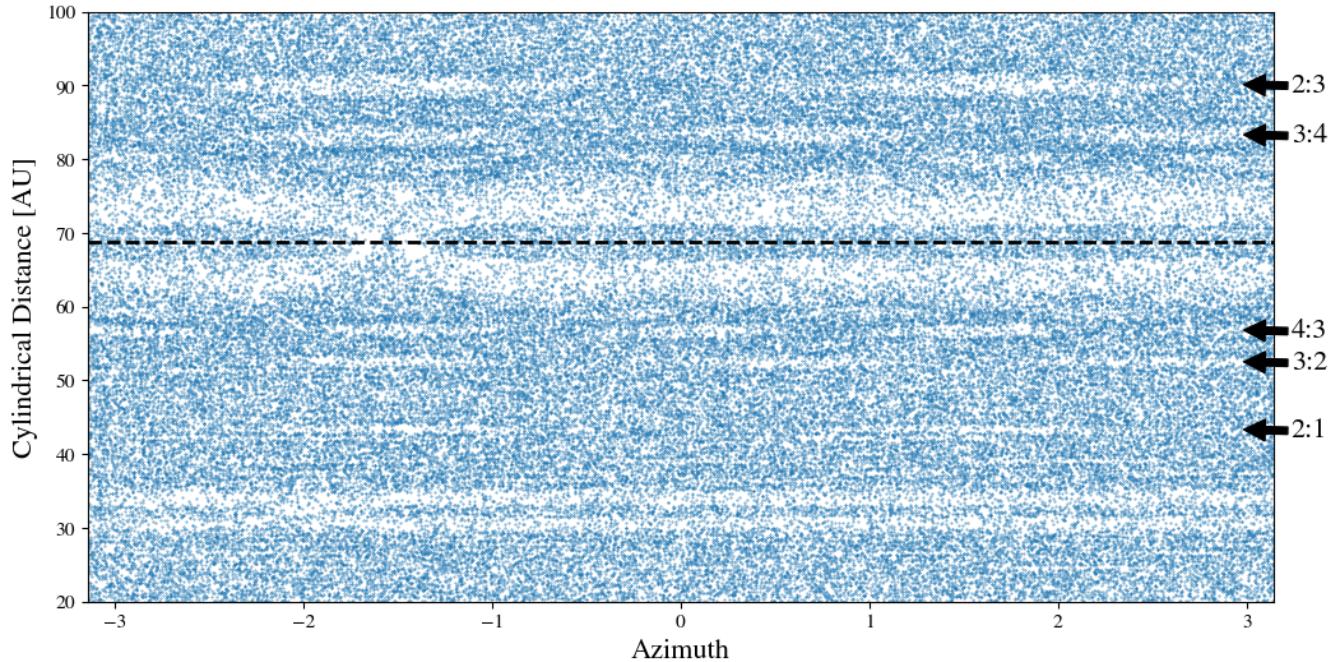
## APPENDIX A: NONAXISYMMETRIC STRUCTURE IN BOLEY 2017

As mentioned in section 4.1, the nonaxisymmetric gap structure seen in figure 1 is present in a number of other studies, including Richardson et al. (2000) and Tabeshian & Wiegert (2016). One would expect the same dynamics to be at work in the simulation of Boley (2017), although the structure is not obviously visible in figure 3 of their work.

Here, we present a closer look at the positions of the planetesimals at the end of the Boley (2017) simulation. In figure A1, we show the planetesimal positions in polar, rather than Cartesian coordinates as originally presented in the paper. Two planets on circular orbits are present here. The first has mass of  $17 M_{\oplus}$  and lies at 32.3 au. The second planet has a mass of  $50 M_{\oplus}$  and lies at 68.8 au. A horseshoe shaped collection of planetesimals is visible to be corotating with each planet. Nonaxisymmetric gaps are also visible at the locations of first order MMRs with the more massive planet, which are indicated with arrows.

Similar to figure 7, a gap is present at the same orbital phase as the planet at the 2:1 MMR, along with a second gap on the opposite side of the disc. Gaps near the 4:3 and 3:2 MMRs are also present in phase with the planet, although the number and spacing between additional gaps appears to be different for each resonance. In addition, the gap structure associated with the exterior MMRs appears markedly different than that of each corresponding interior MMR, which suggests that the structure is related to the symmetry of the path that a planetesimal in a frame corotating with the planet takes.

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**Figure A1.** Positions of planetesimals at the end of the simulation presented in Boley (2017) in polar coordinates. The dashed line marks the location of a  $50 M_{\oplus}$  planet embedded in the disc. First order MMRs with the planet are indicated with arrows. A second planet is also present at 32.3 au, although it is not massive enough to induce any significant structure in the disc near associated resonances.