

Astr150 : section notes

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1 GRAVITY PROBLEMS

The notion of gravity links very closely to the concepts discussed in the lectures : mass , density, or composition. As Newton showed few hundred years ago, we can write the force of attraction between two masses as:

$$F_{grav} = \frac{GMm}{r^2} \quad (1.1)$$

Where F_{grav} stands for the force of gravitational attraction, G the gravitational constant ($G = 6.67 \cdot 10^{-8} [cm^3 g^{-1} s^{-2}]$), M and m - considered masses, and r - separation between the masses. We also have Newton's second law : that any force is proportional to mass. In other words:

$$F \propto m \quad (1.2)$$

and the constant of proportionality is acceleration \mathbf{a} :

$$F = m \mathbf{a} \quad (1.3)$$

Gravitational attraction occurs between any two bodies that have mass - thus, as much as the Earth attracts you, you also attract the Earth (if you did not, the force implies acceleration, and we ought to plunge downwards with some acceleration! Fortunately we can stay on the surface thanks to the structural forces that hold the matter together - electrostatic force - that provide the solidity of your body and feet, so that you can stand firm on the surface. The same applies to a rock, and anything else that is on the surface. That's also why ancient philosopher Aristotle claimed that the natural state of sublunar matter is rest).

One may ask - what is the distance between me and the Earth: if I am standing on the surface, isn't it 0 ? This is answered by **Gauss's theorem**, from a German mathematician Carl Friedrich

Gauss (it's one of his many theorems). This theorem states that the **gravitational attraction between the two bodies is the same as if all the mass for each body was concentrated in one point**. This means that we are attracted by the Earth in the same fashion as if all the mass of the Earth was exactly in its geometrical center (in something like a small black hole), and we were separated from it by the Earth's radius : 6478 km . Same goes for the gravitational attraction between the Moon and the Earth - where the distance between the center of the Earth and center of the Moon is roughly 364000 km.

Therefore r in Eq. 1.1, means in some cases the radius of a considered planet, and sometimes the distance between a planet and a moon, or a planet and a star.

In case of gravitational acceleration on the surface of a planet or star, we call the constant of acceleration from Eq. 1.3 g - the specific gravity.

Thus

$$g = \frac{F}{m} = \frac{GMm}{r^2} \frac{1}{m} = \frac{GM}{r^2} \quad (1.4)$$

and this expression will be the most useful one for the course :

$$g = \frac{GM}{r^2} \quad (1.5)$$

For Earth,

$$g_E = \frac{GM_E}{r_E^2} = 9.81[m s^{-2}] \quad (1.6)$$

For any other planet, we can calculate g given the mass and radius of the planet:

$$g_{planet} = \frac{GM_{planet}}{r_{planet}^2} \quad (1.7)$$

We can see that g is directly proportional to mass, and inversely proportional to the square of the radius. This means that if we increase the mass of the planet 2 times keeping radius constant, g increases 2 times - this corresponds to adding mass to the planet. However, if we increase the radius of the planet 2 times keeping mass constant, g decreases by a factor of 4. Therefore, as long as r_{planet} and M_{planet} are expressed as multiples of r_E and M_E , we can compare the g_{planet} to g_E .

Eg. if $r_{planet} = 2r_E$ and $M_{planet} = 2M_E$, we find:

$$g_{planet} = \frac{GM_{planet}}{r_{planet}^2} = \frac{G2M_E}{(2r_E)^2} = \frac{GM_E}{2r_E^2} = \frac{1}{2}g_E \quad (1.8)$$

2 LUNAR MAPPING

We can gain a plethora of information from studying an image of a surface. In this exercise we study pictures of the Moon, taken by the Lunar Orbiter in the 1960s. Those images were employed for instance by the NASA scientists to determine the best places for landing on the Moon.

Several concepts are relevant to understand the lunar surface. Firstly, layers that are on the top are younger than the ones that are on the bottom - that's the principle of stratification. Processes that were involved in shaping the lunar surface include lava flows, meteoritic impacts, and orogeny. Today only meteoritic impacts modify the surface, since by now the Moon is geologically dead. Lava flows in the past did not all happen at the same time. One episode of lava flow may flow on top of the other: filling old craters, creating lava rivers, lava tubes, and on the grand scale - lunar mare regions. Meteoritic impacts also do not happen all at once - a lot of bombardment happened in the Late Heavy Bombardment, but impacts have been happening ever since. Even now the Moon is occasionally hit by a meteorite - a treat for amateur astronomers seeing a sudden flash on the lunar surface! Orogeny includes all processes that shaped the surface through the engagement of the tectonic forces, even if they were present on a given body only temporarily (i.e. mountains are not formed any more on the Moon nor on Mars, but they were in the past, when these bodies were hotter and exhibited more geologic activity).

The stratified composition gets further disturbed by the meteoritic impacts - the material is ejected outwards, the deepest layer ejected comes from a depth equal to one tenth of the crater diameter. Deepest layers are also deposited closest to the rim of the crater. This means that the bigger the crater, the deeper the layers probed, especially those deposited about its rim. Sometimes the ejecta plows into surface, forming odd-shaped craters and trenches (which look similar to lava tubes, but we can distinguish them by remembering that the debris from impact would be deposited radially about the crater).

Craters can be further modified : filled with lava, forming in turn ghost craters (see Fig. 3). Then there can be many more impacts on the surface of lava, which implies the causality - seeing the ghost crater with some impacts on the top, we know that first there must have been a bigger impact creating the crater, which was then filled with lava, and eventually pock-marked by some smaller impacts. It may be even overlapping with some other later impact. Lava flows create, apart from mare and rather flat, uniform surface, two interesting formations : sinuous rilles, and lava tubes. Sinuous rilles are lava rivers, concave structures that indicate a flow of lava (Fig. 2). Lava tubes (see Fig. 1) are channels of underground flow of lava, which collapse because the top cools faster than the bottom.

Two tectonic processes can further modify the surface : stretching (extension) and compression. Stretching produces straight (linear) rilles - note that the name is similar to sinuous rilles, but the origin is completely different! (see Fig. 4). Compression produces wrinkle ridges (Fig. 5).

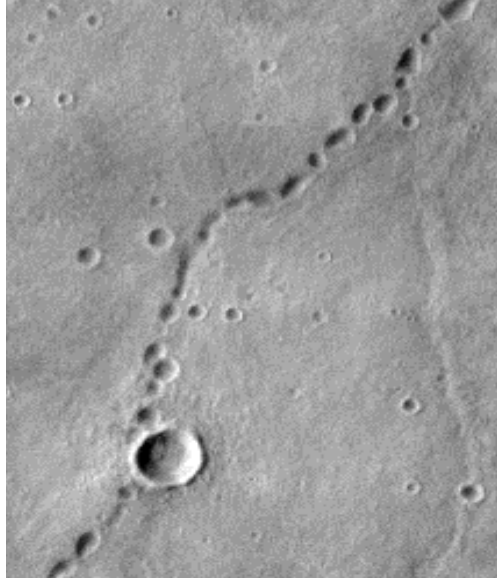


Figure 1: Lava tubes on the Moon

3 CRATER COUNTING

From lunar mapping exercise we see that the surface of a celestial body can convey a lot of useful information about its past, unless modified by surface processes. Erosion and weathering removed much of the cratering record from the surface of the Earth - indeed, most if not all craters on the Earth would be classified as ghost craters, i.e. those that have been modified. Moon, however, does not have an atmosphere, which meant that most of the impacts are left intact, and thus we can decompose its history by looking at the crater count. If we assume that meteors hit bodies in the solar system uniformly in time, i.e. at roughly the same rate now as 100 mln years ago or 1.5 bln years ago, the more craters there are, the older the body is (note that on average this is true, but in reality, the rate of impacts is not exactly constant, unless averaged over really long, >50 mln years, time). This makes sense thinking for instance about the analogy of a surface constantly undergoing bombing - if bombs are dropped at a constant rate, a surface with more bomb-craters would be older. This is the premise of this exercise, that employs crater counting as a measure of age. Even if there is some atmosphere, as in the case of Mars, there is no preferred location to be affected more by the presence of the atmosphere on the entire planet, so that wherever the meteor falls on Mars, if it is too small it will burn in the atmosphere, just like in the case of Earth. Dust devils, and other phenomena of the interaction of atmosphere with the surface on Mars have a negligible weathering effect on larger craters, certainly insignificant for craters with diameter bigger than few km. The atmosphere of Venus is so thick that only medium-sized craters would be found - large craters require big impactors, that happen very rarely, and the resurfacing that happens on Venus makes its surface very young. Small impactors get largely stopped by the atmosphere. Since craters come in a variety of sizes, it is reasonable to count their number from some

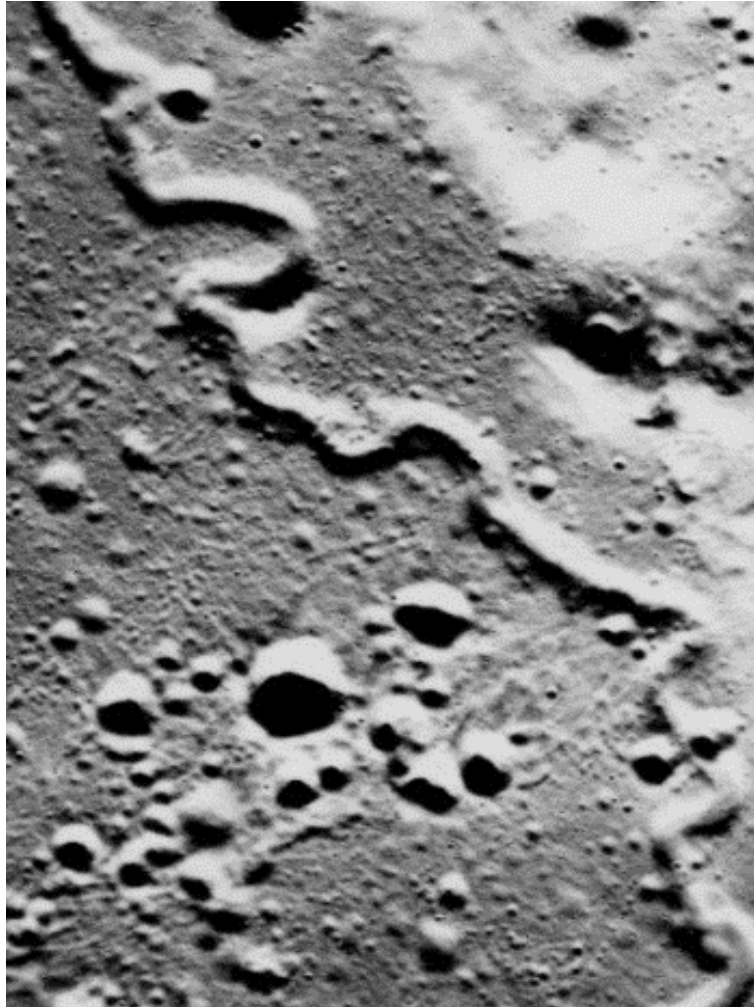


Figure 2: Sinuous Rilles on the Moon

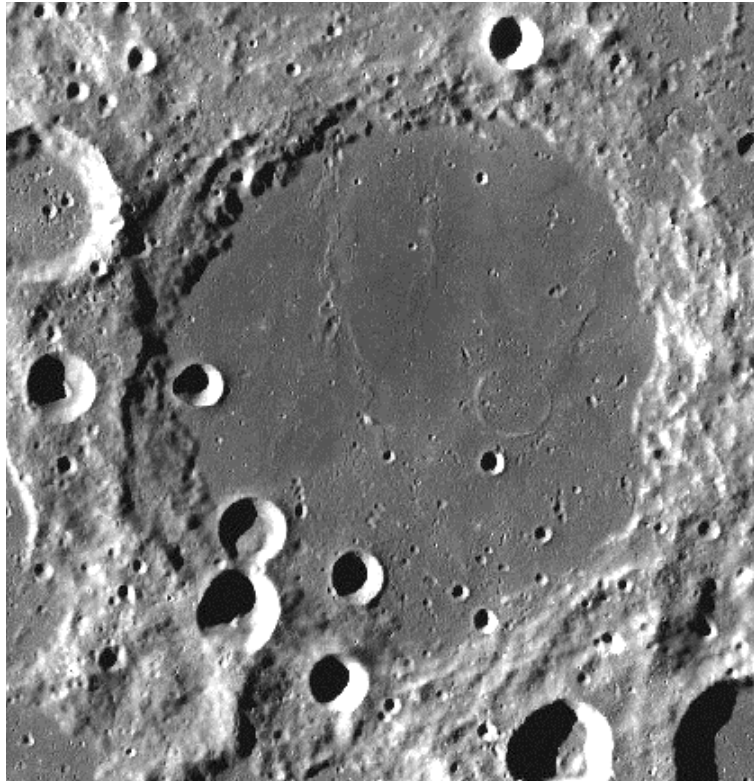


Figure 3: Ghost craters on the Moon



Figure 4: Straight rilles on the Moon

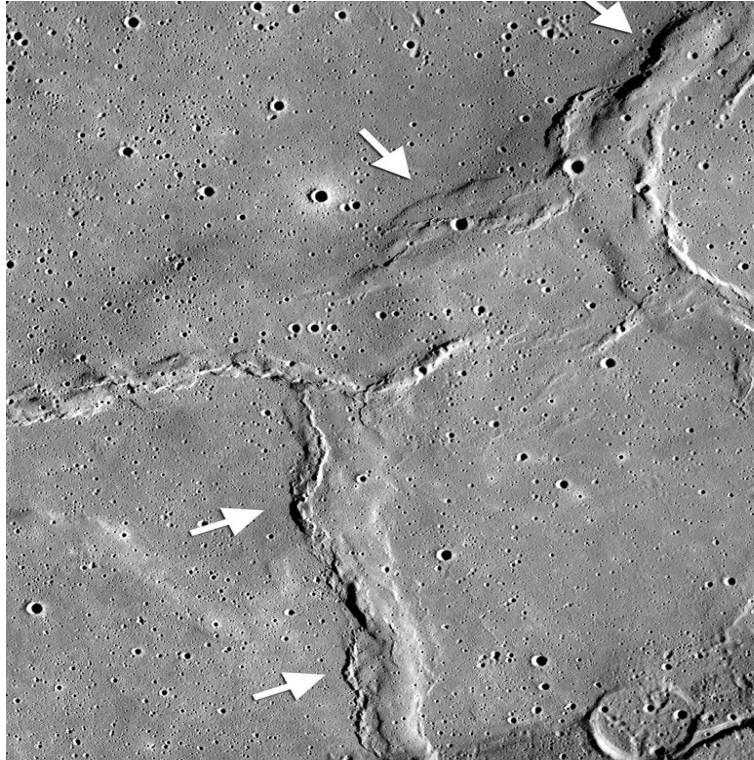


Figure 5: Wrinkle ridges on the Moon

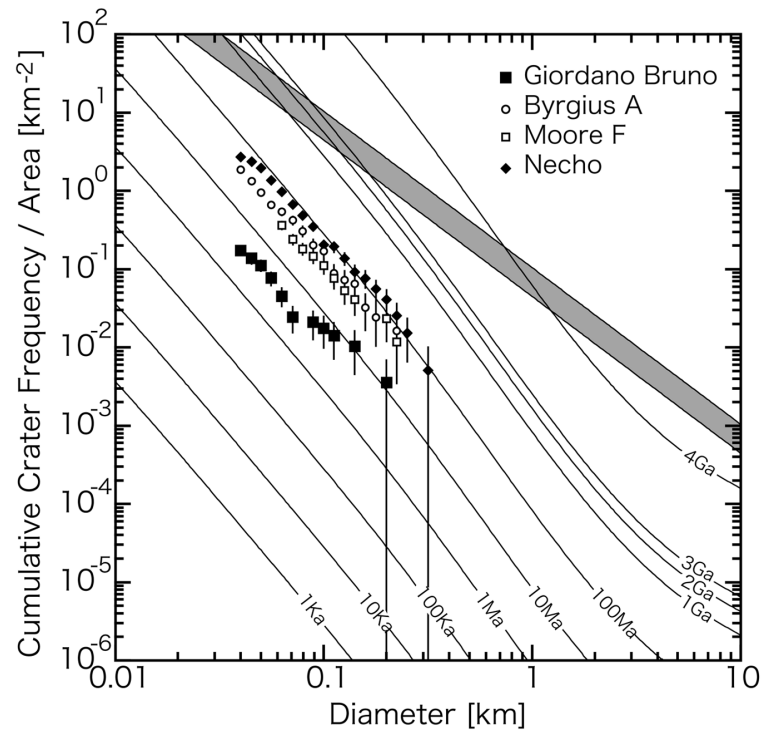


Figure 6: Cumulative crater density plot. Crater density plot is used to determine the age of a given surface. From T. Morota et al. 'Formation age of the lunar crater Giordano Bruno' *Meteoritics and Planetary Science* 44, Nr 8, 1115-1120 (2009). Also see <http://www.psrdr.hawaii.edu/Feb10/GiordanoBrunoCrater.html>

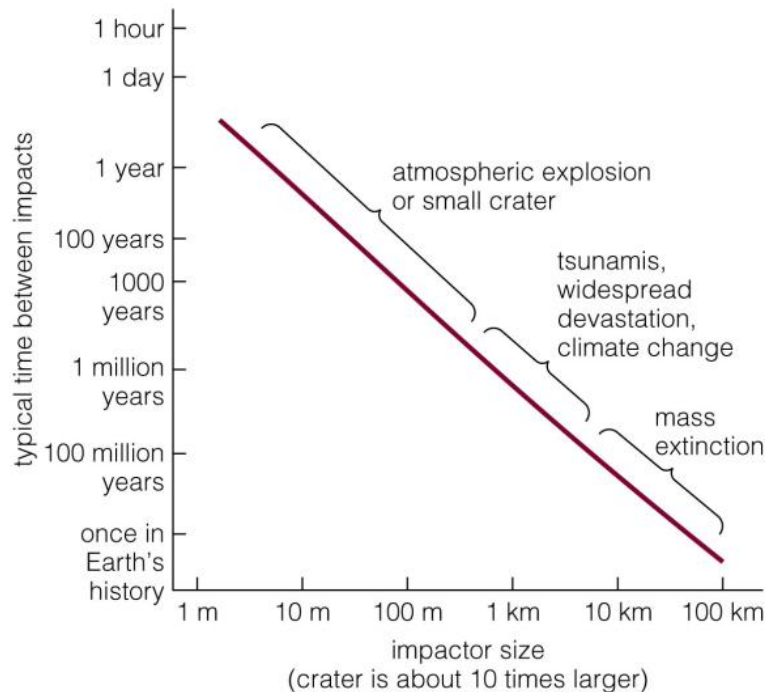


Figure 7: The frequency of impacts depends on the size of the impactor. Larger meteors hit more rarely than the smaller ones.

threshold crater diameter to be consistent with our estimate. **Cumulative Crater Density** is the number of craters N larger than a certain diameter D km, denoted as $N(D)$. The most often employed threshold is $D = 10$ km, so that by $N(10)$ we mean the number of craters with a diameter bigger than 10 km.

On a young surface, such as that which has undergone some volcanic activity, the crater count age is reset - from the moment a surface was covered by the lava, it is plain, and we call this surface a young surface. As time goes by it will be hit by more and more meteorites, creating more and more craters, so that the crater counting will give us the age from the moment of the most recent volcanic activity. Thus even though all locations on the surface of the planet (such as Mars) are being hit at constant rate (eg. one 10 km crater per 1000 years), some of them were reset by geologic or tectonic activity. Even though Mars is not geologically active now, it used to have some activity, which is witnessed by the presence of some of the largest extinct volcanoes in the Solar System, with Olympus Mons, the tallest one (27 km above the surrounding terrain). See <http://mars.jpl.nasa.gov/gallery/atlas/olympus-mons.html> and <http://hyperphysics.phy-astr.gsu.edu/hbase/solar/marsoly.html> for more stunning images of this grand volcano. Mt Rainier seems really puny compared to this 88000 ft monster (yes, 4 times the height of Mt Rainier!).

The way the surface looks changes not only with regards to the number of craters bigger than 10 km, but also with regards to their distribution. The older the surface, the more big craters there is, because big impacts happen much more rarely than small impacts. On Fig. 6 we can

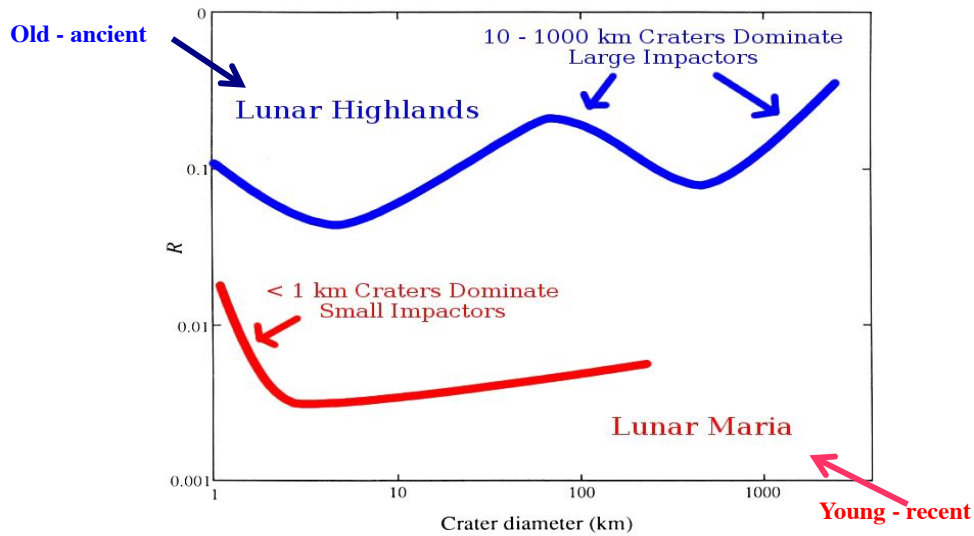


Figure 8: The R-plot shows different populations of craters on surfaces of different age

see that different regions of the Moon have different crater populations, and the older the region the more craters it has. The type of craters changes with time - the older the surface the larger of its portion is covered by the biggest craters, because there was enough time for the rarely happening big impacts to occur (see Fig. 8). Those different crater populations are illustrated by an R-plot - a plot of the fraction of the total area of a considered region covered by craters of given diameter. Some guideline ages for the Moon:

- Late Heavy Bombardment craters 4.6 Gy,
- Light-toned Terrae (highlands) , plagioclase feldspar 4.5 Gy,
- Dark-toned Mare, volcanic basalts 3.1-3.8 Gy

(Maria have 200 times fewer craters than highlands)

$$R = \frac{\text{Area covered by craters of diameter } D}{\text{Total Area}}$$

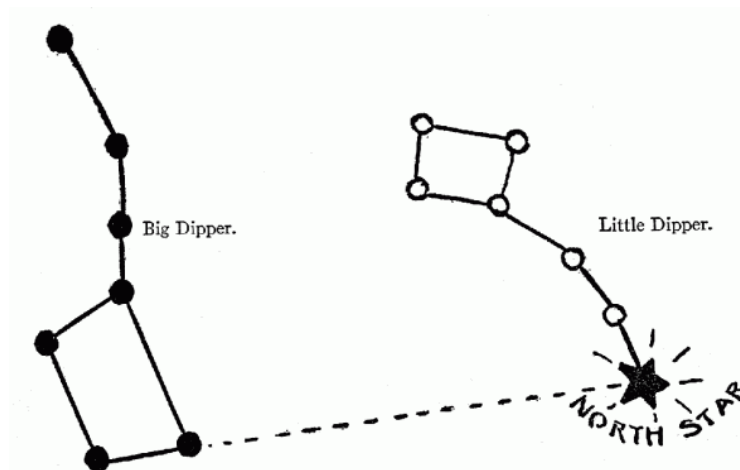


Figure 9: How to find the North Star.

4 PLANETARIUM VISIT

During our planetarium visit you have encountered several concepts. We used the World Wide Telescope software which can be used for free online at <http://www.worldwidetelescope.org/>. For looking at the night sky in terrific detail I also recommend Stellarium <http://www.stellarium.org/>, and for that as well as travel in space (eg. to look at our galaxy from outside, or travel along the space probes, visit the planets) Celestia <http://www.shatters.net/celestia/> is an excellent choice. They are all free of charge.

In the planetarium we firstly looked at the night sky as seen from Seattle (ϕ - latitude - of 47°). There you could see the constellations that we can see during the night in the summer - Ursa Major, Lyra, Cygnus, etc. You can find those and many more with Stellarium or WWT. You can also print the night sky maps for yourself at <http://skymaps.com/downloads.html>. There is also a project <http://www.sky-map.org/> worth checking. Similarly, Google came up with another way of looking at the sky: <http://www.google.com/sky/>. We particularly noted the Ursa Major - the Big Bear, or specifically a smaller part of it called the Big Dipper. It is useful because extending the end of the Dipper 5 times upward we hit the North Star, which lies in the constellation of Ursa Minor - the Small Bear (see Fig. 9).

We noted that there are two types of grids important for astronomers: Altitude/Azimuth grid, and the Equatorial grid. As the Earth rotates on its axis, stars appear to rise and set. The only star that stays still is the North Star - Polaris, which is very close to the North Celestial Pole - an extension of the rotation axis of the Earth. The Alt/Az grid is defined with respect to the observer, and defines circles parallel to the horizon as circles of the same altitude above the horizon. Hence an altitude of 90° is directly overhead - in Zenith. The Altitude and Azimuth of most stars changes during the night, hence this system is only useful for telling someone where the star is at some particular moment in time. It's easy to use, because the altitude is measured from the horizon, and the azimuth is measured from the north point eastward, and the north point is easily found by identifying the location of the North Star (see Fig. 10).

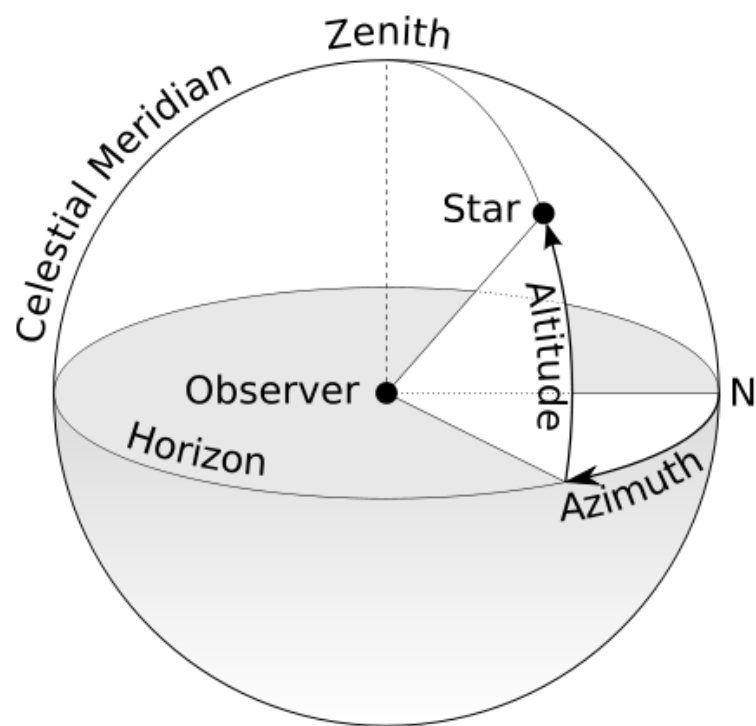


Figure 10: The horizontal - Alt/Az coordinates.

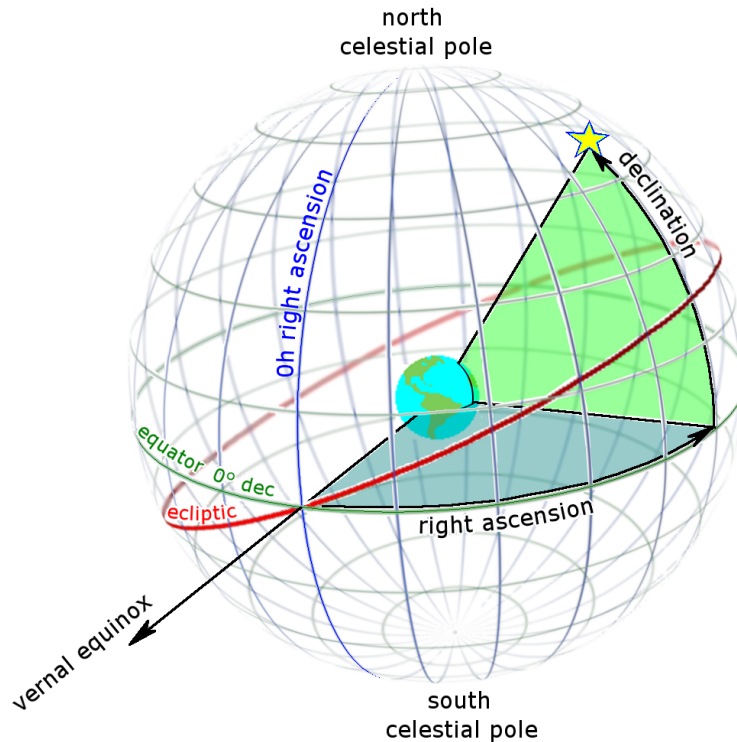


Figure 11: The equatorial coordinates.

The Equatorial grid is defined analogously to the Latitude/Longitude grid on the surface of the Earth. It is the projection of that grid onto the sky. The Latitude equivalent is called declination - δ , whereas the longitude equivalent is Right Ascension : α . δ is measured north/south from the celestial equator, $\pm 90^\circ$. α is an angular distance along the celestial equator from the vernal equinox point, which is one of the two points where the ecliptic intersects the celestial equator (Fig. 11).

We saw how the apparent movement of stars is different at different latitudes: the North Star's altitude indicates directly the Latitude. Thus in Seattle the North Star is 47° above the horizon, on the Equator : 0° , and on the North Pole - directly overhead : 90° . (see <http://cseligman.com/text/sky/motions.htm>).

We also noted the movement of the planets, in particular focusing on their retrograde motion. Planets do not move within visible timescale - but if you note the position of any planet at several consecutive nights with respect to background stars (i.e. note their δ and α), it will change. It will change faster for the planets that are closer to the Sun than to those that are further away. Some planets would appear to move backwards, because at those moments the Earth overtakes them on the orbit.

Looking at the Solar System from above, we noted the presence of smaller bodies such as asteroids and comets in very specific locations - in the Kuiper belt, the Oort Cloud, the Main Asteroid belt, and the Trojan Asteroid groups, leading and trailing the orbit of Jupiter - see

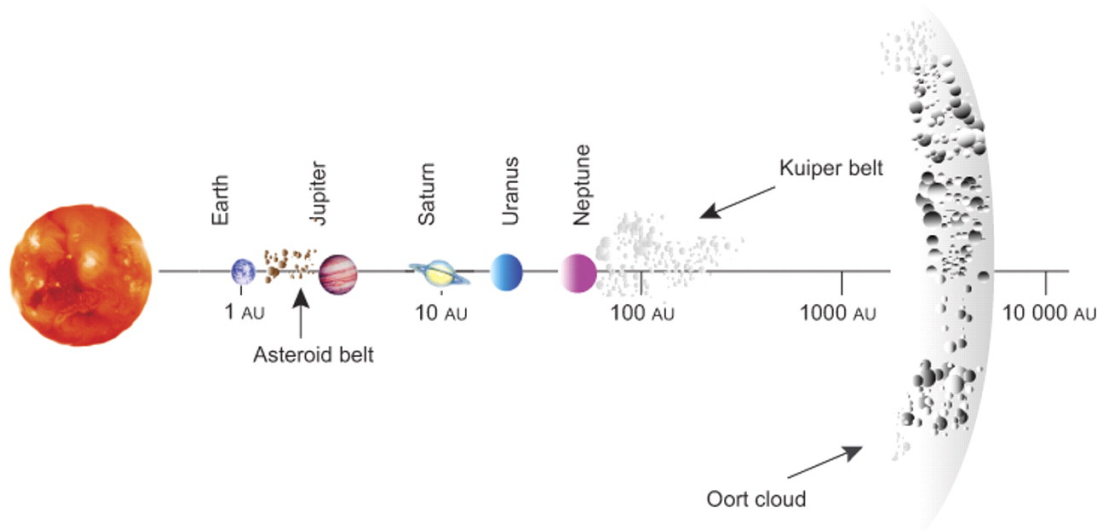


Figure 12: An overview of the Solar System, adapted from <http://elements.geoscienceworld.org/>

Fig. 12

Finally, we discussed how stars and planets are all formed in stellar nurseries, such as the Orion Nebula (Fig. 13).

5 PLATE TECTONICS

Surface of our planet or a moon is changed by a variety of processes. Apart from impacts, internal heat plays a major role. Several important concepts were discussed in the class. Firstly, one recognizes that geologic activity scales with size, so that the bigger the celestial body, the longer the duration of its geologic activity (see Fig. 14). Remember: **geologic activity scales with size**. Thus for instance Mars, which is almost two times smaller than Earth, was still geologically active until about 600 mln years ago - around that time the internal heat coming from the decay of radioactive elements became insufficient to drive its volcanism and geologic activity.

With sufficient internal heat and (as is often assumed) presence of liquid water, the crust may become divided into plates, which move around the globe. There are places where a plate is formed : a ridge, and places where a plate sinks under another one - a subduction zone. Both places would have volcanism present. But there is also a third way to form a volcano - through a hot spot. It forms a shield volcano, by forcing magma under pressure to make a channel through the crust to the surface. If the crust overhead is moving (as in the case of plate tectonics), this will form a chain of volcanoes - the plume stays in the same place, but the plate continuously moves. Several volcanoes in the chain may be active at the same time, which depends on the size of the hot spot, the thickness of the crust, and the characteristics



Figure 13: An image of the M42 nebula.

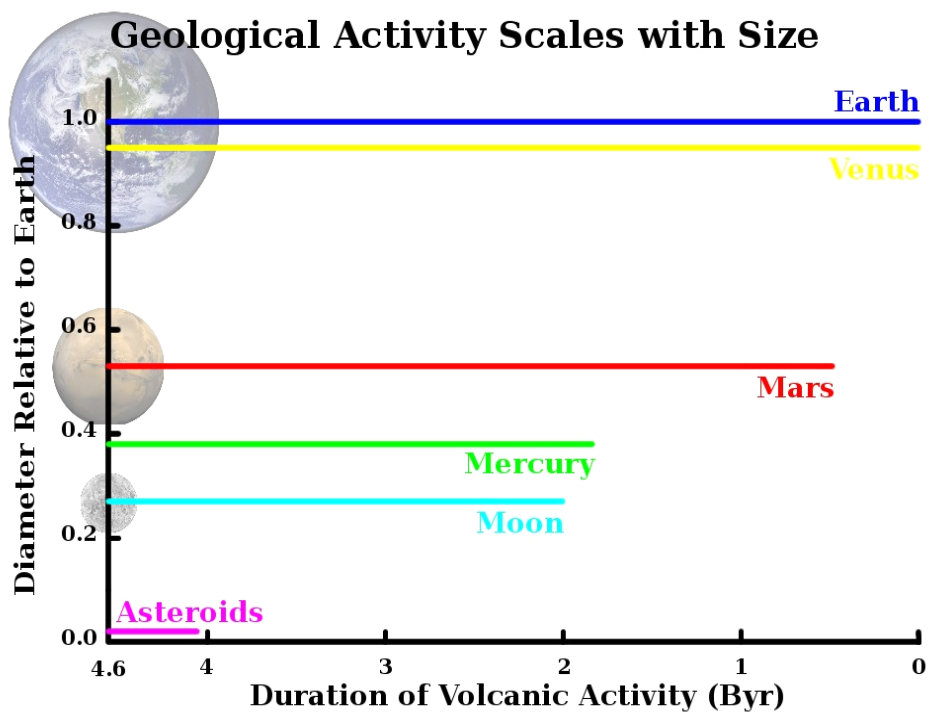


Figure 14: Geologic activity scales with size

of the crust, such as local composition, which may help to maintain or close the lava channel which connects the plume with the surface.

The hot spot volcanism together with plate tectonics implies an age relation for the volcanoes in a chain. Those furthest away from the hot spot would be the oldest. If erosion takes place (as for Hawaii islands), those would also be most eroded, weathered down, sinking under its own weight. On Mars there is no erosion currently, so the size does not imply age. If there is no plate tectonics, then the volcano grows without ceasing, just like Olympus Mons. It may be that Alba Planitia and Olympus Mons are in a chain, but there is not much evidence to support that (and explain why a tectonic plate would be moving so slowly to allow for their almost uninterrupted growth).

Given that, especially on Mars, there is no clear evidence for plate tectonics, an entirely independent way of measuring the age of any surface is crater counting. However, this is very inaccurate, because for a small number of craters bigger than 10 km ($N(10)$) it is hard to tell the age. Accuracy of crater counting is related to the $N(10)$ plot. If my count for $N(10)$ is 10 ± 2 then it corresponds to the surface of age 1 ± 0.5 bn years. However, if we count $N(10)$ to be 100 ± 2 , then it corresponds to the age 3.8 ± 0.05 bn years (age determination is read off the $N(10)$ vs age plot for the crater counting lab), which again shows that it is difficult to accurately determine the age of a young surface.

Thus a better way is to take samples and radioactively date them. We can also look for stratification of outflows, to find a relative age difference.

6 ATMOSPHERIC ESCAPE

In the previous section we discussed plate tectonics, which shows how the heat from initial contraction and the ensuing heat of the radioactive decay, provide energy necessary to drive plate tectonics and volcanism. Sometimes water is not present, which results in a lack of plate tectonics, but doesn't stop volcanism - like in a case of Venusian 'blob tectonics'.

Initially a planet is formed with many gases primordially present in the Solar System, such as Hydrogen and Helium - this is called the Primary Atmosphere. Such atmosphere may be quickly lost if the $v_{escape} > \frac{1}{6} v_{gas}$. This is true because v_{gas} - the gas velocity, is linked to temperature, as we can equate the thermal and kinetic energies of particles:

$$k_b T = \frac{mv^2}{2} \quad (6.1)$$

where k_b is the Boltzmann constant, T is the temperature at that distance from the Sun, m - mass of the gas molecules, and v - velocity of gas particles.

From that equation a rearrangement of variables yields:

$$v_{gas} = \sqrt{\frac{2k_b T}{m}} \quad (6.2)$$

and if we express m in atomic mass units : $1 \text{ amu} = 1.66 \cdot 10^{-27} [kg]$, and insert $k_b = 1.38 \cdot 10^{-23} [m^2 kg s^{-2} K^{-1}]$ then the constants simplify :

$$v_{gas} = 157 \sqrt{\frac{T}{m}} \quad (6.3)$$

which is precisely the equation used in the lab.

This shows us that not all bodies can hold on to all types of atmospheres: the escape velocity of the Earth is $v_{Earth} = 11200 \text{ [m/s]}$, and so we need $v_{gas} < (1/6) \cdot 11200$, thus $v_{gas} < 1866 \text{ [m/s]}$. At the distance of Earth to Sun : 1 [AU], the temperature is 400[K]. Thus we can calculate:

$v_{Hydrogen} = 157 \sqrt{400/2} = 2220 \text{ [m/s]}$, $v_{Water} = 157 \sqrt{400/18} = 740 \text{ [m/s]}$, so that we see that Earth could not have an atmosphere of Hydrogen, but can certainly hold on to water (vapour). Apart from primary atmosphere, due to volcanism, and other outgassing, we form a secondary atmosphere, composed of heavier elements. It would contain carbon dioxide (CO_2), nitrogen (N_2), etc. Even the secondary atmosphere changes its composition over time due to various atmospheric sinks : one of them we already described in detail: if gas molecules move too fast, they will escape from the gravitational influence of a planet or moon. This is called thermal escape. Another way for the atmosphere molecules to be lost to outer space is by interaction with the solar wind, which is the stream of energetic charged particles flowing outwards from the Sun. We are protected from a lot of this interaction by our magnetosphere, which deflects a lot of the charged particles, but some of them make it all the way to the atmosphere, and being directed by the magnetic field lines of the Earth, they are guided towards the poles, creating aurorae. There are two other atmospheric sinks. One is condensation on the surface, for instance when the temperature is too low the gas freezes - like dry ice caps on Mars. Another is the chemical interaction with the surface materials, which traps certain elements in rocks. Similar to this is the locking of carbon dioxide in skeletons of all the sea creatures, which after they sediment on the bottom of the sea, form layers of limestone. Thus on the global scale of hundreds of millions of years, carbon dioxide level in the Earth's atmosphere will steadily decrease by a combination of all those sinks. Global warming and the current increase of carbon dioxide level is only a short-term aberration from the overall trend of the decrease of its level.

7 BIG QUIZ REVISION

Revision might be made most efficiently with good use of materials and time. Lectures contain the lecture highlights on the first few slides - those are a good guideline for what to revise. Below are revision questions that may help you prepare for your exam.

- What ages should we expect when describing the Solar System? When was it formed? What are the densities of iron, rock, water ? What is the moment of inertia K ? What values would you expect for a differentiated and non-differentiated bodies?
- What is $N(10)$? How do we determine absolute ages? Why is it best to take samples? What do we measure in the samples when trying to estimate their age ?

- What is differentiation ? What K do we expect for the Moon? What's the original lunar crust made of ? What is the age of highlands / mare ? What are the 5 stages of planetary evolution? In which stage is currently Mars, in which stage is Earth ?
- What is accretion? What are dynamic and chemical constraints for the composition of the Moon? What are the various schemes of the formation of the Moon? What does the stability of lunar orbit imply for the formation of life on Earth?
- Explain why geologic activity scales with size. Why does Venus not have plate tectonics? What is blob tectonics? How do we form volcanoes?
- What is a Secondary Atmosphere ? What does it contain ? What are atmospheric sinks ?

Other questions to consider:

- what is an escape velocity? - what is a primary atmosphere? - what is the evidence for the past presence of water on Mars? - what is the age of the Martian surface? - what is the age of Venus surface? - what does the presence of Olivine on some places on the surface of Mars indicate concerning the duration of the presence of water ? - why do we often land close to craters? - why did Phoenix probe land close to the Martian pole? - what is an impact breccia? - what is the evidence for the K-T event ? - what caused the extinction of dinosaurs? what was the role of the atmosphere ?

8 FINAL REVISION I

- (Wk 6) What are chondrites? Are they differentiated or undifferentiated ? What is the rarest type of chondrite?
- What are achondrites? What types of bodies do they come from ?
- How likely are you to find a stony-iron meteorite on Earth ? Where do they come from in their parent body ?
- what are Iron meteorites? Where in their parent body do they come from ? How old would you expect one to be ?
- Asteroids - What are their characteristics? sizes - shapes - compositions surface characteristics
- What is the origin of the Asteroid Belt? What is the density of the Asteroid Belt? What are Trojan Asteroids? What are the Kirkwood Gaps?
- Physical Structure - Where are the asteroids? Why are Resonances important?
- Chemical Structure - Composition vs. Distance.
- (wk 7) Outer Moons I - Rings and Dead Worlds : What are the types/characteristics of giant planet satellites? What's characteristic about regular vs irregular satellites?

- What is the Roche Limit of a planet?
- Dead Worlds : Callisto, etc - How are their surfaces similar/different from the Earth's Moon?
- Outer Worlds II - Recently and Currently Active Worlds : Recently Active Worlds - Ganymede, etc. Why are outer worlds geologically active longer than similarly sized inner worlds ?
- Titan: Why is Titan's secondary atmosphere different from the secondary atmospheres of the inner (terrestrial) worlds?
- Why must Titan's atmosphere be constantly replenished?
- (wk8) Outer Worlds III - Active Worlds: Io (J), Europa (J), Enceladus (S) , Triton (N) - Why are they so geologically active?
- How are Io's volcanoes similar/different from volcanoes elsewhere in the Solar system?
- Why might Europa be an interesting place to look for life?
- What processes have modified the surface of Triton?
- Outer Worlds IV - The Icy Dwarfs Pluto - The Kuiper Belt - Comets : Pluto - What is its orbit like? Why all the controversy about its status?
- How did the Kuiper Belt form? What is its relation to Pluto and Triton? What is its relation to comets?

Other questions to consider:

- what is reflectance? how could it help us learn something about the surface measured ? -
- what does accretion time depend on ?

9 FINAL REVISION II

- Crater Density : how can we link it to an absolute age? In other words, how do you calibrate your $N(10)$ plot ?
- What is a crater population plot? What is R-plot ? Are crater population plots for Moon, Mars, Mercury similar? Why ?
- Are crater populations of Callisto and Moon similar? Could you use $N(10)$ to find age of surface of satellites beyond the snow line? Why or why not?
- Does geologic activity always scale with size of the planetary/ lunar size object ? What is required? What are counterexamples amongst moons of Gas Giants? What is their primary source of heat?
- (wk9) The Giant Planets - Giant Planet Atmospheres: How are they different from terrestrial atmospheres? Why they all look different? What causes the colors of the clouds?

- The Giant Planets - Magnetic Fields: What are their interiors like and how do we know this? Why does Earth's magnetic field make it a safe place? Why does Jupiter's magnetic field make it a dangerous place?
- Solar System Formation: How do we get from a homogeneous mixture of materials to the solar system we see today ?
- Why does the density of terrestrial worlds decrease as you move further from the Sun? Why is Jupiter so large?
- (wk10) Exoplanets : Why would you expect to find planets around other stars? How would you detect planets around other stars? Why is this difficult?
- How do the other planetary systems discovered differ from ours? How could you determine if there was life on some of these planets?