

# Impact of photometric redshifts on the galaxy power spectrum and BAO scale in the LSST survey

Reza Ansari<sup>1</sup>, Adeline Choyer<sup>2</sup>, Farhang Habibi<sup>1</sup>, Christophe Magneville<sup>3</sup>, Marc Moniez<sup>1</sup>, Stéphane Plaszczynski<sup>1</sup>, Cécile Renault<sup>2\*</sup>, Jean-Stéphane Ricol<sup>2</sup>, and Julien Souchard<sup>2</sup>

<sup>1</sup> LAL, Univ. Paris-Sud, CNRS/IN2P3, Université Paris-Saclay, Orsay, France

<sup>2</sup> Univ. Grenoble Alpes, CNRS, Grenoble INP, LPSC-IN2P3, 38000 Grenoble, France

<sup>3</sup> DSM/Irfu/SPP, CEA-Saclay, F-91191 Gif-sur-Yvette Cedex, France

Received DD MM 2018/ Accepted DD MM 2019

## ABSTRACT

**Context.** Imaging billions of galaxies every few nights during ten years, LSST should be a major contributor to precision cosmology in the 2020 decade. High precision photometric data will be available in six bands, from near-infrared to near-ultraviolet. The computation of precise, unbiased, photometric redshifts up to  $z = 2$ , at least, is one of the main LSST challenges and its performance will have major impact on all extragalactic LSST sciences.

**Aims.** We evaluate the efficiency of our photometric redshift reconstruction on mock galaxy catalogs up to  $z=2.45$  and estimate the impact of realistic photometric redshift (hereafter photo-z) reconstruction on the large-scale structures (LSS) power spectrum and the baryonic acoustic oscillation (BAO) scale determination for a LSST-like photometric survey. We study the effectiveness of the BAO scale as a cosmological probe in the LSST survey.

**Methods.** We have performed a detailed modelling of the photo-z distribution as a function of galaxy type, redshift and absolute magnitude using our photo-z reconstruction code with a quality selection cut based on a Boosted decision tree (BDT). We simulate a catalog of galaxies in the redshift range [0.2 – 2.45] using the Planck 2015  $\Lambda$ CDM cosmological parameters over 10,000 square-degrees, in the six *ugrizy* bands, assuming LSST photometric precision for a ten-year survey. The mock galaxy catalogs are produced with several redshift error models. The LSS power spectrum is then computed in several redshift ranges and for each error model. Finally we extract the BAO scale and its uncertainty using only the linear part of the LSS spectrum.

**Results.** We have computed the fractional error on the recovered power spectrum  $\sigma[P(k)]/P(k)$  which is dominated by the shot-noise at high redshift ( $z \gtrsim 1$ ), for scales  $k \gtrsim 0.1$ , due to the photo-z damping. The BAO scale can be recovered with a percent or better accuracy level from  $z = 0.5$  to  $z = 1.5$  using realistic photo-z reconstruction. Outliers (*i.e.* galaxies with catastrophic photo-z) can represent a significant fraction of galaxies at high redshift ( $z \gtrsim 2$ ), causing bias and errors on LSS power spectrum measurement.

**Conclusions.** Reaching the LSST requirements for photo-z reconstruction is crucial to exploit the LSST potential in cosmology, in particular to measure the LSS power spectrum and its evolution with redshift. Although the BAO scale is not the most powerful cosmological probe in LSST, it can be used to check the consistency of the LSS measurement. Moreover we show that the impact of photo-z smearing on the recovered isotropic BAO scale in LSST should stay limited up to  $z \approx 1.5$ , so as long as the galaxy number density balances the photo-z smoothing.

**Key words.** Galaxies: distances and redshifts – Galaxies: photometry – Cosmology: large-scale structure of Universe – Cosmology: dark energy

## 1. Introduction

The six-band (*ugrizy*) Large Synoptic Survey Telescope (LSST) survey, described in [LSST Science Collaboration et al. \(2009\)](#), will yield a sample of about ten billion galaxies over a huge volume. It will be the largest photometric galaxy sample of the next decade to study the large-scale structures (LSS) of the Universe. It aims to characterise the distribution and evolution of matter on extragalactic scales through observations of baryonic matter at a broad range of wavelengths. The LSS encode crucial information about the content of the Universe, the origin of the fluctuations and the cosmic expansion background in which the structures evolve. Of particular interest is the imprint on galaxy clustering of baryon acoustic oscillations (BAO), which reflects the acoustic waves at recombination, related to the sound horizon at that epoch. The BAO scale is sufficiently small that it is possible to measure it precisely with a large volume survey; yet,

this scale is large enough that it is not significantly altered by non-linear evolution. The BAO features can be used as a standard ruler to measure distances and constrain the dark energy equation of state, especially when used in combination with the weak-lensing.

LSST will observe every few nights half of the whole sky up to a magnitude of 24. After 10 years of survey, the depth of the stacked images should reach magnitude 27 in the i-band and LSST should provide a catalog of a few billions of galaxies with a photometry good enough to be useable for cosmology. Indeed, doing cosmology implies to accurately know the galaxy positions. While the angular coordinates will be exquisitely measured thanks to the LSST optics and camera design, it will be much more challenging to reconstruct the third coordinate, the redshift.

Results using spectroscopic data of the BOSS survey [Dawson et al. \(2013\)](#), part of the Sloan Digital Sky Survey III, allowed to measure precisely the BAO scale and lead to accurate

\*corresponding author

cosmological parameters, for instance [Anderson et al. \(2012\)](#), [Aubourg et al. \(2015\)](#) or more recently [Salazar-Albornoz et al. \(2017\)](#), [Alam et al. \(2017\)](#). Although BAO constrains with LSST might be not fully competitive with the future spectroscopic survey DESI [Levi et al. \(2013\)](#) or later Euclid [Laureijs et al. \(2011\)](#), it is crucial to perform BAO analysis to cross-correlate with the other LSST probes. The reader can refer to [Zhan & Tyson \(2018\)](#) and references therein for a recent review of the cosmology reach of LSST, including the BAO probe. The sensitivity of each probe to systematic errors and to dark energy parameters is different, so it is mandatory to check the consistency of all results and powerful to combine them [Alonso et al. \(2017\)](#). The study of BAO in LSST can in return be a way to test our photo-z reconstruction and tune the quality cuts to deal with the total number of galaxies versus the number of outliers. The current DES survey [Drinkwater et al. \(2010\)](#), with 5 filters only, has already to face similar challenges [Crocce et al. \(2017\)](#), obtaining a 4 % distance measurement of the standard ruler [Abbott et al. \(2017\)](#).

The goal of this paper is to test the impact of the use of photo-z to reconstruct the LSS power spectrum, considering several options on the required quality up to  $z=2.2$ . Compared to some previous work on BAO with photometric redshifts in LSST [Zhan et al. \(2009\)](#), [Abrahamse et al. \(2011\)](#) where simpler LSST observation and photo-z error model were used, we have done a complete analysis using realistic mock galaxy catalogs with galaxies. The cosmological impact is illustrated by comparing the BAO scales measured using spectroscopic redshifts (hereafter spectro-z) or photo-z of galaxies from a mock catalog; our BAO analysis is similar to the method described in [Blake et al. \(2011\)](#). The BAO scale extraction is simply based on the Wiggle method described in [Beutler et al. \(2016\)](#).

The paper is organised as follows. We first present the LSST survey and the goal of the paper. Section 2 is dedicated to an overview of the mock catalog of galaxies and Sect. 3 presents the photometric redshift reconstruction characteristics and performances. Section 4 presents method and results regarding the matter power spectra and the BAO scale. We have developed a Montecarlo simulation to easily link precision on the recovered BAO scale and photo-z performances presented in Sect. 5. Finally, we illustrate in Section 6 .We end by summarising our results in Sect. 6.

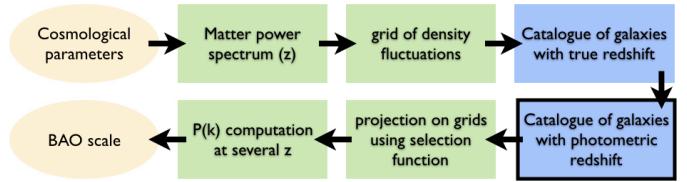
## 2. LSST mock galaxy catalog simulation and analysis procedure overview

A complete end to end simulation and analysis pipeline has been developed to measure the performance of a LSST-like photometric survey for recovering the LSS power spectrum. Starting from a set of cosmological parameters and going through mock galaxy catalogs with realistic photo-z, the pipeline is used to extract the BAO scale at several redshifts. The main blocks of this pipeline are depicted in Fig. 1.

We start from the matter power spectrum  $P(k)$  computed at a given reference redshift  $z_0$ , from which a huge grid of density fluctuations  $\delta\rho/\rho$  is calculated, through random generation of Fourier coefficients and Fourier transform (FFT).

The grid of density field is then converted into a catalog of galaxies using galaxy luminosity functions (LFs). We use a set of LFs corresponding to different galaxy classes mapped into spectral types. It should be stressed that, in this procedure, we assume that the local galaxy absolute luminosity is directly proportional to the local matter density  $(1 + \delta\rho/\rho)$ .

The key ingredient for the study presented here is the error model for assigning "observed" redshifts to the galaxies. We



**Fig. 1.** Graphic representation of the main steps of the simulation chain, from cosmological parameters up to the BAO scale measurement.

compare several cases, including a realistic photo-z reconstruction through the FastPz method discussed below.

The mock galaxy catalog is then analysed: after computation of the redshift dependent selection function, the galaxy angular positions and redshifts are converted into 3D positions using a fiducial cosmology. Galaxies are projected into 3D-grids positioned at several redshifts. The "observed" power spectrum  $P_z(k)$  for each redshift range is computed through a Fourier transform, averaging over all Fourier modes within a  $k$ -bin. Finally, the BAO scale and its uncertainty are extracted from the ratio of  $P_z(k)$  to a wiggle-less power spectrum through a dedicated fit procedure described in Sect. 4.3 and compared to the expected value from the input cosmological parameters.

This pipeline directly provides a catalog, so we ignore all systematics introduced by the catalog extraction from the images, even if effects like blending [Jones & Heavens \(2019\)](#) or dithering [Awan et al. \(2016\)](#) affect in different manner photometric and spectroscopic surveys. The simplistic way galaxy catalogs are produced do not allow to derive the precision of the effective reconstruction of the matter power spectrum by LSST. This methodology and the following analyses are however well suited to evaluate the impact of realistic photo-z redshifts, compared to spectro-z redshifts, on the matter power spectrum, and more specifically on the reconstruction of the BAO scale.

### 2.1. Computation of the density field grid

The aim is to simulate the density field on a large portion of the sky in order to get a simplified view of half of the sky that will be monitored by LSST. We define a volume, hereafter named *BigCube*, which has a comoving depth of 5,600 Mpc along the  $Oz$  axis and a transverse comoving section of  $10,000 \times 10,000$  Mpc $^2$ . The grid center is set at 3,200 Mpc from the observer. This *BigCube* is divided into cells of  $8 \times 8 \times 8$  Mpc $^3$ , this cell size being small enough compared to the BAO scale ( $\sim 150$  Mpc). Some non-linearities may affect the galaxy distribution around or below the cell scale but, for our purpose, we do not need to take into account such additional complexities.

The matter power spectrum is initially computed for the redshift of the *BigCube* center with the cosmological parameters of the fiducial model. We use the set of parameters defined in [Planck Collaboration et al. \(2016\)](#) in the conservative case, based on the TT power spectrum with lensing reconstruction, polarisation only at low multipoles and external data (column 3, Tab. 4): the Hubble constant is then  $H_0 = 67.90 \text{ km.s}^{-1}.\text{Mpc}^{-1}$ , the cold dark matter reduced density is  $\Omega_{\text{CDM}} = 0.2582$ , the baryonic matter reduced density is  $\Omega_B = 0.0483$  and the (linear) power spectrum is described by a power-law with an amplitude  $\sigma_8 = 0.8154$  at a scale of  $8 h^{-1}\text{Mpc}$  and a spectral index  $n_s = 0.9681$ . By hypothesis in the  $\Lambda$ CDM model, the neutrino mass is minimal and the geometry is flat.

Density fluctuations are generated from the power spectrum  $P(k)$  using the standard formalism, described for example in Eisenstein & Hu (1998). Fourier components  $F(\mathbf{k})$  for the matter density fluctuation field are randomly generated following the power spectrum  $\langle |F(\mathbf{k})|^2 \rangle = P(k = |\mathbf{k}|)$ . The matter density fluctuation field  $\delta\rho/\rho$  is then computed through an FFT on the array of generated Fourier components. We next apply the growth factor to each cell in the cube, according to its redshift. This method is less accurate than the one implemented in CLASS Lesgourgues (2011) but the resulting spectra differs only marginally. This simplified method is sufficient, as our aim is mainly to quantify the loss of information induced by the photo-z reconstruction on the matter power spectrum and as we do not try in this work to quantify uncertainties on cosmological parameters.

The density field grid is then clipped, setting cells with  $\delta\rho/\rho < -1$  to  $-1$  to avoid cells with negative matter density. This clipping has no effect on the results presented in this paper, as we will use the power spectrum of the clipped grid as the reference power spectrum. We have also checked that the distortions of the power spectrum due to clipping, ranging from a factor 0.5–0.6 to 0.8 depending on the k-scale, do not significantly affect its overall shape and BAO features.

## 2.2. Simulation of the catalog

The density fluctuations are afterwards converted into galaxy number fluctuations, assuming proportionality of light to matter density  $(1 + \delta\rho/\rho)$ . The key ingredient for this step is the use of observed galaxy LF to populate each grid cell with a number of galaxies with absolute magnitudes and galaxy classes or spectral types following the LFs. Details of the procedure are provided in Gorecki et al. (2014), the main difference being that we have used LF parameters derived by Zucca et al. (2009) for the study presented here, instead of the one derived by Dahlen et al. (2005) used in our previous work. Both sets of LFs provide the distribution of galaxies divided in broad types, in several redshift ranges. We have used the Schechter parameters from Zucca et al. (2009) with fixed power law index ( $\alpha$ ). LF parameters are given for galaxies divided in three broad types, labeled as Early, Late or StarBurst, corresponding roughly to Elliptical, Spiral and Irregular morphological types. These three broad categories are then mapped into a library of spectral types.

The main steps for converting the density fluctuation grid to a list of galaxies are summarised below:

- The total number of galaxies  $N$  within our survey volume is computed by integrating the LFs with a faint absolute magnitude limit well above the LSST detection threshold at all redshifts. This number  $N$  is then used to compute the average number of galaxies per *BigCube* cell.
- For each cell, the average number of galaxies in the cell  $\bar{n}(\mathbf{x})$  is computed according to the local matter density  $1 + \delta\rho/\rho$ . The number of galaxies in the cell is then drawn with a Poisson distribution with mean  $\bar{n}(\mathbf{x})$ . Galaxies positions are uniformly generated within each cell, then their angular positions and true redshifts are computed from their positions in the grid. An absolute magnitude and a type are assigned to each galaxy according to the multi-type, redshift dependent LFs. We have chosen to use the Zucca et al. (2009) parametrisation which, once extrapolated to high redshifts, provides a reasonable evolution of the number of galaxies. Only galaxies within a cone with half-opening angle of 60 degrees (defining a solid angle of  $\pi$ , so covering  $\approx 10,000$

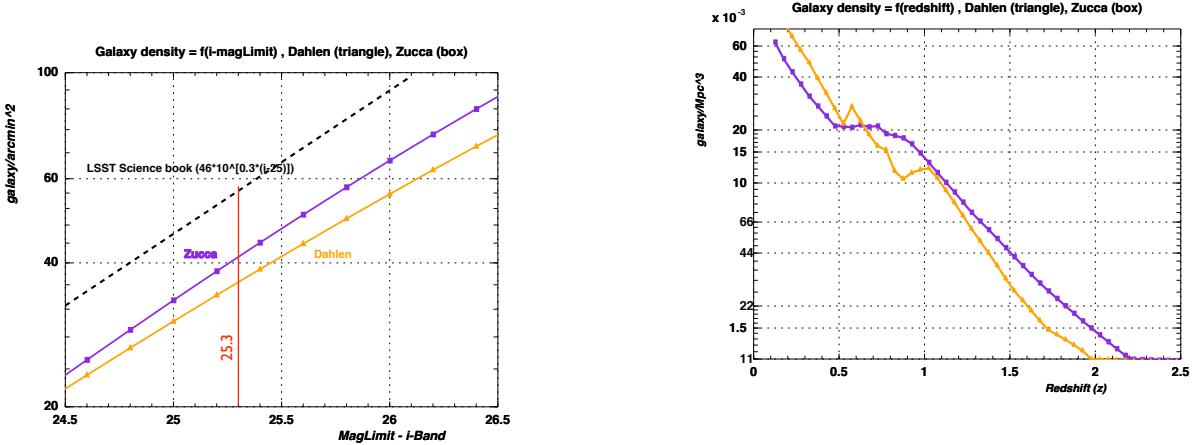
$\text{deg}^2$ ) are kept. The catalog in our simulation properly covers the redshifts from 0.2 to 2.45.

- We use a library of 51 interpolated SEDs from six main SEDs, composed of 4 SEDs of Coleman et al. (1980): Elliptical, Sbc, Scd or Im, and the 2 SEDs of Kinney et al. (1996): SB3 and SB2. Galaxies from each LF class are then distributed over a subset of these 51 interpolated SEDs. We then compute the apparent magnitudes in each of the six LSST bands (*ugrizy*) for each galaxy present in the simulation volume, taking into account the nominal LSST 10-year photometric errors for each band. The photometric errors will depend on the distribution of visits per filter and on the ability to manage the atmospheric properties during observations, as shown in Graham et al. (2018). Note that these errors have been estimated for point-like sources and are applied to extended sources. Only galaxies with  $m_i < 25.3$ , satisfying the so-called *LSST golden selection cut*, are kept for the subsequent analysis, allowing a reliable estimate of the photo-z.

As it will be shown and discussed in this paper, the effective density of galaxies observable in LSST as a function of redshift is a crucial ingredient to compute its expected performance for measuring the LSS power spectrum and the BAO scale. Many values of the Schechter parameters describing the LFs can be found in the litterature Ramos et al. (2011), Dahlen et al. (2005), Zucca et al. (2009) for instance. As expected, LF parameters derived from observations have some uncertainties and one finds significant differences when galaxy densities in the LSST survey are computed from different set of LFs.

In the left panel of Fig. 2, we show the expected galaxy surface density number as a function of the i-band magnitude limit. We compare the numbers obtained from the extrapolation used in the LSST Science Collaboration et al. (2009) and numbers derived from Dahlen et al. (2005) and Zucca et al. (2009) LF parameters. These curves are obtained by integrating the LF Schechter functions for each redshift, with the faint end absolute magnitude limit computed from the i-band magnitude limit, taking into account the K-correction and the luminosity distance. The observable galaxy density is then integrated as a function of redshift, over the cosmological volume element. As it has already been pointed out, by Chang et al. (2013) for instance, the expected galaxy density used in LSST Science Collaboration et al. (2009) is rather optimistic. Indeed, for the *LSST golden selection cut*, the Science book extrapolation yields  $\sim 55 \text{ gal/arcmin}^2$ , compared to  $\sim 36 \text{ gal/arcmin}^2$  for Dahlen et al. (2005) and  $\sim 41 \text{ gal/arcmin}^2$  for Zucca et al. (2009). The initial LSST value of  $55 \text{ gal/arcmin}^2$  seems to come from a too rough first guess. To explain the discrepancy between the two other sets of LFs, we compare their expected number of galaxies per  $\text{Mpc}^3$  as a function of redshift for the *LSST golden selection cut* in the right panel of Fig. 2. The Zucca et al. (2009) LFs lead to a higher galaxy number density at all redshifts above  $z \approx 0.6$ . So, once integrated up to  $z \geq 0.8$ , the surface galaxy number density is higher with this set of Schechter parameters. Moreover the Zucca et al. (2009) LFs provides smoother, so more realistic, galaxy density evolution along the cosmic time.

Our simulated galaxy number density should also be compared to the available observed values. The recent Hyper Suprime-Cam Survey has properties rather similar to the expected ones of the LSST survey Medezinski et al. (2018). Their galaxy number density usable for photo-z reconstruction is of almost 22 galaxies per  $\text{arc-min}^2$ , the one usable for weak-lensing is of almost 25 galaxies per  $\text{arc-min}^2$  and it exceeds 30 galaxies



**Fig. 2.** Left: expected galaxy surface density as a function of i-band magnitude limit: comparison of LSST Science Collaboration et al. (2009) extrapolation, numbers derived from the Dahlen et al. (2005) and Zucca et al. (2009) LFs parameters. Right: number of galaxies per  $\text{Mpc}^3$  as a function of redshift, for the LSST golden selection cut from Dahlen et al. (2005) and Zucca et al. (2009) LFs. Purple squares show results obtained with the Zucca et al. (2009) numbers, orange triangles show results obtained with the Dahlen et al. (2005) ones.

per arc-min<sup>2</sup> in some parts of the GAMA15H field Mandelbaum et al. (2018). So our galaxy number density sounds reasonable, in agreement with previous LSST forecast for weak-lensing too Chang et al. (2013).

From the 28 billions of simulated galaxies, about 1.4 billion satisfy the LSST golden selection cut. This number is roughly a third of the "official" volume of the LSST catalog, firstly because only 10,000 square-degrees have been simulated instead of the 20,000 square-degrees that will be scanned by LSST, secondly because the galaxy density initially foreseen was probably overestimated, as we have pointed it out previously.

Figure 3 shows the histograms of the differential number of galaxies as a function of the redshift, before and after the LSST golden selection cut, as well as the distribution per broad type. Our mock catalog extends from a redshift of 0.2 to a redshift of 2.45. About a tenth of the galaxies have a good enough photometry for cosmology purpose at  $z \approx 1$  and one thousandth at  $z \approx 2$ .

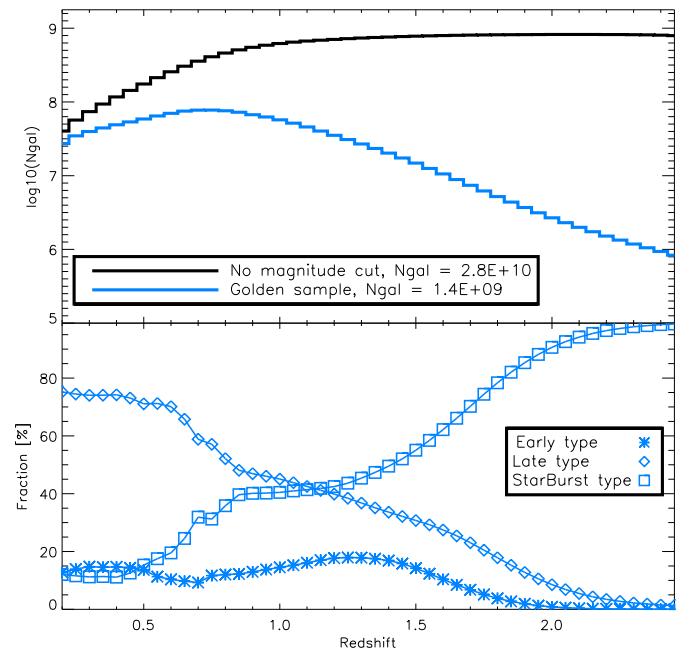
Note that in the full simulation of the mock catalog, dust extinction, photometric errors and use of 51 interpolated SEDs (instead of 6 main SEDs) are included. This comes to be small differences between the numbers from the curve labeled Zucca, derived from the full simulation and the results labeled Zucca of the simplified model presented in Fig. 2.

In order to take into account realistic uncertainties on the redshift, random errors are added to the true redshift  $z_s$ . Four error models are studied:

- a Gaussian error with  $\sigma = 0.03 * (1 + z_s)$ ,
- a more realistic photo-z reconstruction using PDF distributions described in 3.1,
- a photometric error improved by quality cuts on the BDT variable as described in 3.1.1, keeping 90 (BDT 90%) or 80 % (BDT 80%) of the galaxies.

### 3. Photometric redshift reconstruction

To guarantee the cosmological capabilities of LSST, requirements on the quality of the photo-z reconstruction are provided



**Fig. 3.** Top: Number of galaxies per redshift interval. The LSST golden selection cut is defined by  $m_i < 25.3$ . Bottom: relative distribution of the broad types of the galaxies satisfying the LSST golden selection cut as a function of the redshift.

in the LSST Science Collaboration et al. (2009). Statistical properties of the distribution of  $e_z = (z_s - z_p)/(1 + z_p)$  as a function of  $z_p$  where  $z_s$  denotes the spectro-z and  $z_p$  denotes the photo-z, have to verify the following requirements:

- the interquartile range IQR of  $e_z$ , is lower than 0.05 (the goal is lower than 0.02),
- the fraction of outliers  $f_{\text{out}}$  defined as the fraction of galaxies with  $e_z > 0.15$  is lower than 10 %,

- the bias  $b$  defined as the median value of  $e_z$  is lower than 0.003.

Note that our Gaussian error model corresponds to  $\text{IQR} \approx 0.04$  as  $\text{IQR} \approx 2 \times 0.6745 \times \sigma$ .

### 3.1. Method

The photo-z reconstruction is performed with a template fitting method presented in details in [Gorecki et al. \(2014\)](#). The SED library is made of 51 templates used for the photometry computation. Note that this step of the work is expected to be rather optimistic for LSST redshift reconstruction performance since our SED library completely represents the simulated observations. The method uses the prior proposed in [Benítez \(2000\)](#) to remove some degeneracies (mostly Lyman / Balmer breaks) based on magnitude distribution in the i-band. For each galaxy we have to compute a 3D-grid of  $\chi^2(z, T, E_{\text{B-V}})$  where the redshift  $z$ , the galaxy type  $T$  and the extinction from inner dust  $E_{\text{B-V}}$  are the three parameters determined by the photo-z reconstruction code. The extinction follows the [Cardelli et al. \(1989\)](#) or [Calzetti et al. \(1994\)](#) law, depending on the galaxy type. We marginalise the  $\chi^2$  grid over the two other parameters to get the marginalised 1D posterior probability density functions  $P(z)$ ,  $P(T)$  and  $P(E_{\text{B-V}})$ . The redshift estimator can be chosen as  $z_p^{\chi}$  (value of the triplet that minimise the  $\chi^2$  in the grid) or  $z_p^{\text{marg}}$ , value that maximises the posterior probability density function  $P(z)$ . In this work we use the latter one:  $z_p = z_p^{\text{marg}}$ . This code gives photo-z performances similar to other public codes (Ricol et al., in preparation) but allows us more freedom, in particular the possibility to compute our PDF functions.

#### 3.1.1. Outliers rejection

The originality of the method is the outliers effective rejection with a machine learning technique based on the characteristics of the PDFs.

We have improved this technique in the present work. The Likelihood Ratio method presented in [Gorecki et al. \(2014\)](#) has been indeed replaced by a Boosted Decision Tree (BDT) algorithm based on the same discriminant parameters but more robust and performant. The parameters used in the BDT are:

- the number of peaks in the marginalised 1D posterior probability density functions denoted by  $N_{pk}(\theta)$ , where  $\theta$  is either  $z$ ,  $T$  or  $E_{\text{B-V}}$ ,
- when  $N_{pk} > 1$ , the logarithm of the ratio between the height of the secondary peak over the primary peak in the 1D posterior probability density functions, denoted  $R_L(\theta)$ ,
- when  $N_{pk} > 1$ , the ratio of the probability associated with the secondary peak over the probability associated to the primary peak in the 1D posterior probability density functions, denoted by  $R_{pk}(\theta)$  (the probability is taken proportional to the integral of the PDF between two minima either side),
- the absolute difference between  $z_{pk}$  and  $z_{pk}^{\text{marg}}$ , denoted by  $D_{pk} = |z_{pk} - z_{pk}^{\text{marg}}|$ ,
- the maximum value of  $\log(\mathcal{L})$  where  $\mathcal{L}$  is the likelihood,
- the colors  $\mathbf{C} = (u - g, g - r, r - i, i - z, z - y)$ ,
- the photo-z value  $z_p = z_{pk}^{\text{marg}}$ .

The BDT method is trained on an independent sample of 100,000 galaxies. The impact of the training sample (number of galaxies, completeness) has been studied in details in [Gorecki et al. \(2014\)](#) and shows no significant bias within reasonable statistics.

After training, the BDT method provides a BDT value between -1 and 1 for each test sample galaxy, outliers having lower BDT values than good reconstructions. We can then apply a cut on this number.

The technics on photo-z reconstruction are continuously improved, see for instance [Süveges et al. \(2017\)](#), [Cavuoti et al. \(2017\)](#), [Sadeh et al. \(2016\)](#), [Gomes et al. \(2017\)](#), [Pasquet et al. \(2019\)](#). However, we do not expect significant changes in the main features of the reconstructed photo-z distributions and the impact of the LSS power spectrum reconstruction.

#### 3.1.2. Fast photo-z computation

For a sake of computing time we have developed a fast photometric redshift reconstruction (hereafter FastPZ) tool that allows to quickly compute the photo-z for each galaxy. The FastPZ tool is based on true photo-z distributions sorted in bins of true redshift, absolute magnitude and broad type. We use 121 bins in absolute magnitude  $M$  (from -24 to -12 with a step of 0.1), 3 bins in broad type  $BT$  (Early, Late or StartBurst) and 61 bins in redshift  $z$  (from 0 to 3 with a step of 0.05) leading to a total of 21,780 bins. In each bin we simulate several thousands of photometric data, compute the photo-z for galaxies satisfying the *LSST golden selection cut* and construct the  $P(z_p - z_s | M, BT, z_s)$  distributions. For very low statistic bins we increase the number of galaxies in order to get reliable distributions.

The full method uses photometric information in the six filters. It is the one applicable to real data. In a simulation framework, the fluxes are first computed from the broad type, the absolute magnitude and the true redshift of each galaxy. The FastPZ method directly uses the 3D ( $BT, z_s$ ) mapping to retrieve the reconstructed photo-z PDF from the broad type and the absolute magnitude of the galaxy. This method is able to compute 10 millions of photo-z's in 1 second while it takes several hundreds of hours with the full method, resulting in a gain of CPU of about 1 million.

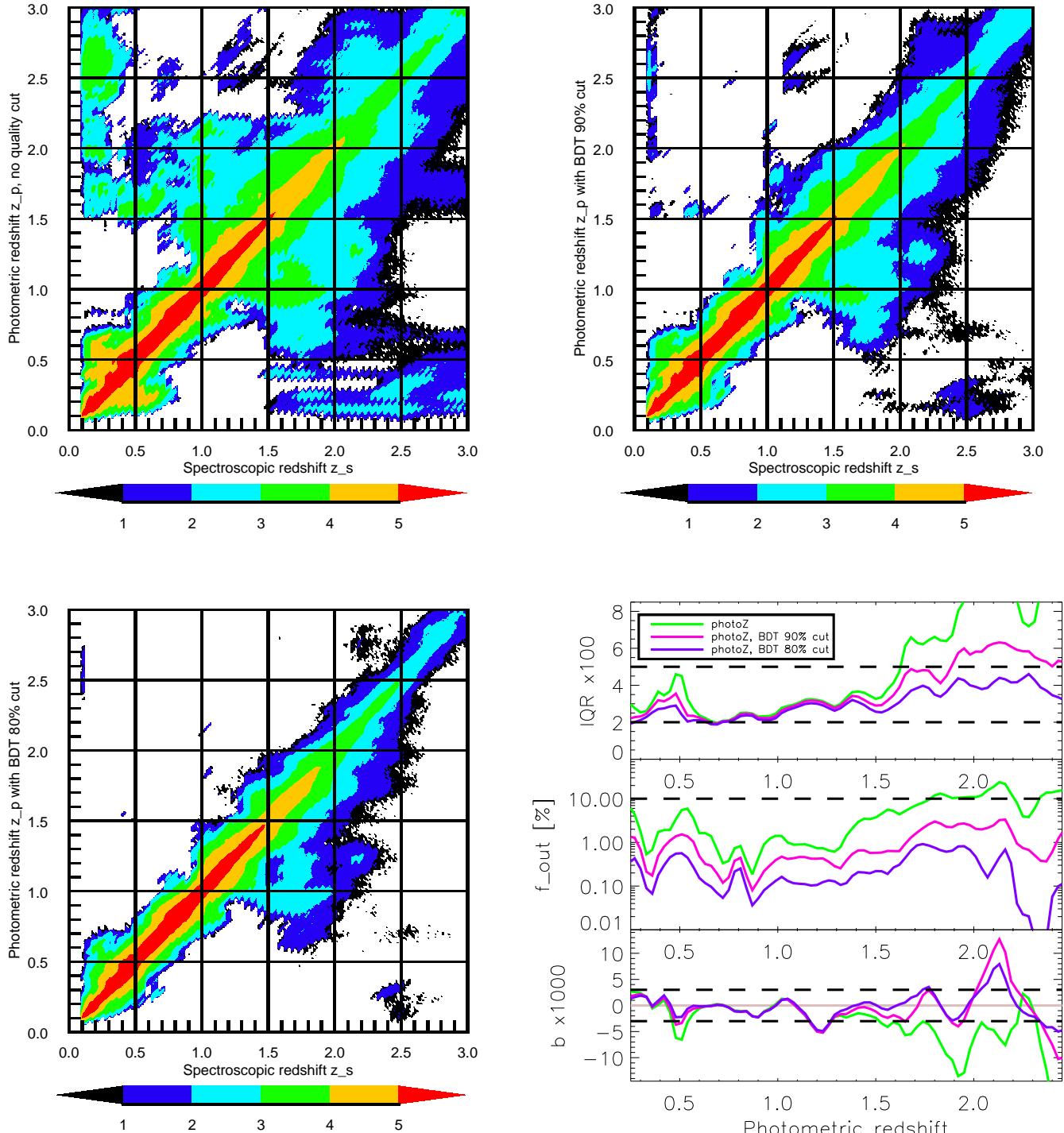
We have checked that performances of the FastPZ tool and the whole redshift reconstruction are very similar (Ricol et al., in preparation).

Figure 4 presents the results of the FastPZ and the performances of the outliers rejection with the BDT method applied on our mock catalog of  $\approx 1.4$  billion galaxies. The BDT cut value is tuned to keep a given percentage of the total number of galaxies.

Figure 4 (top left) shows the density of photo-z  $z_p$  as a function of the spectro-z  $z_s$ . While most of the reconstructed redshifts lie close to the diagonal, the plane is almost filled by some “catastrophic” reconstructions (*i.e.* galaxies labelled as outliers). Nevertheless, the galaxy density is shown with a logarithmic colour scale, so at least one thousand times more galaxies lie in the diagonal area than in the green areas.

Figure 4 (top right and bottom left) shows the same distribution as previously, but after use of the BDT 90% or BDT 80% cut to reject outliers. Galaxies lying in the diagonal area are mostly unaffected while all the space above the diagonale is now almost completely cleared out. The part below the diagonal is much cleaner too. If the BDT cut is always efficient, it is particularly powerful at high redshift ( $z_s > 1.5$ ) where the number of galaxies with “catastrophic” redshift is reduced by at least 2 orders of magnitude.

The quality of the photo-z reconstruction is evaluated with the estimators used by the LSST collaboration. These statistical properties of  $e_z = (z_s - z_p)/(1 + z_p)$  are plotted as a function

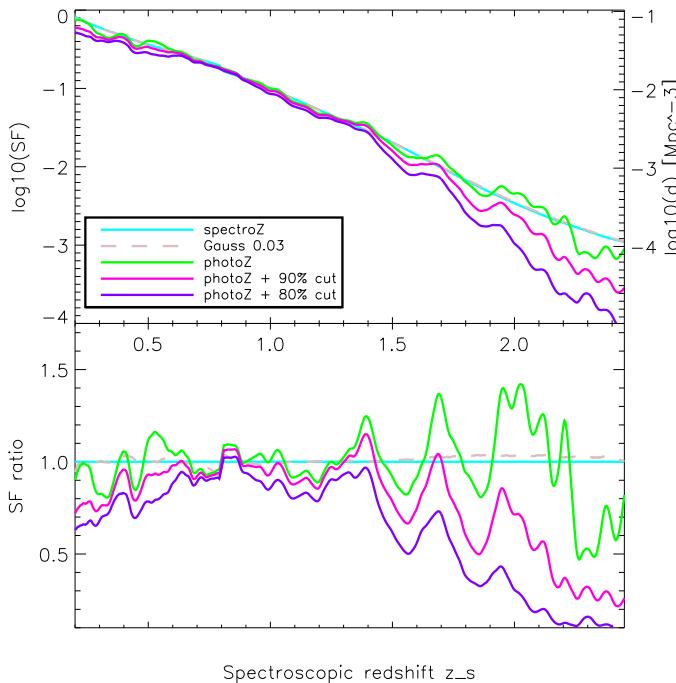


**Fig. 4.** Distribution of the photo-z  $z_p$  as a function of the spectro-z  $z_s$ ; top left: photo-z from the FastPZ method (see sect. 3.1.2), top right: with the BDT 90% cut, bottom left: with the BDT 80% cut. The colour scale for the galaxy density is logarithmic. Bottom right: statistical properties of  $e_z = (z_s - z_p)/(1 + z_p)$  as a function of  $z_p$  with (top) the interquartile range IQR of  $e_z$ , (middle) fraction of outliers  $f_{\text{out}}$ , defined by  $|e_z| > 0.15$ , (bottom) bias  $b$ , defined as the median of  $e_z$ . Green lines show the values obtained with the photo-z's without any quality cut while the pink (resp. purple) lines show the results with the photo-z satisfying BDT 90% (resp. BDT 80%) cut. The dashed black lines correspond to the LSST requirements (plus the goal value in the IQR case).

of the photo-z in Fig. 4 (bottom right). The LSST requirements are more or less satisfied by all the photo-z reconstructions up to  $z_p \approx 1.6$ . For higher redshifts, a BDT cut is mandatory, the interquartile range (IQR) remains close the requirement and the

bias varies significantly. Note that a similar degradation of the photo-z reconstruction performance is observed above  $z = 1.6 - 1.8$  in the real data of the Hyper Suprime-Cam Subaru Strategic Program Data Release 1 [Tanaka et al. \(2018\)](#).

### 3.2. Error model statistics, selection functions and quality cuts



**Fig. 5.** Selection functions SF in the redshift range [0.2-2.45], using the various estimations of the redshift. Top: selection functions defined as the ratio between the number of galaxies in the catalog (passing the *LSST golden selection cut*) and the number of simulated galaxies. The right axis gives the galaxy density  $d$  per  $\text{Mpc}^{-3}$ . Bottom: ratio between the selection functions and the selection function in the spectroscopic case.

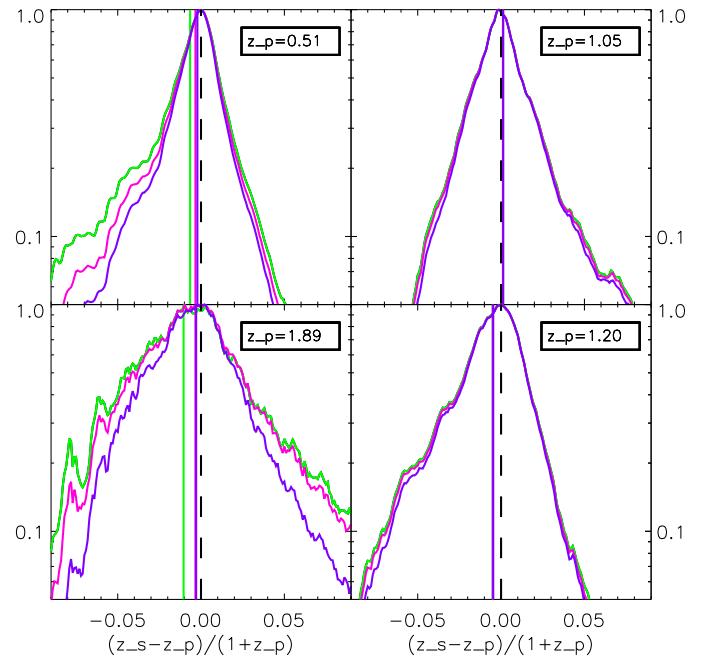
Selection functions are computed as the ratio between "observed" galaxy redshift distributions (spectroscopic, Gaussian or photometric), satisfying the *LSST golden selection cut*, and the "true" redshift distribution of all galaxies (black line of Fig. 3). Of course we have to limit the magnitude when "true" galaxies are simulated. "No magnitude cut" means in fact  $m_i < 27$ , which should be close to the depth of the stacked images after the 10-year survey completion. Galaxies are simulated with absolute magnitude in the range [-24, -13]. This range is appropriate to our simulation as, even at the lower redshifts, the selection function remains below 1 but is close to 1, so we simulate enough, but not too many very faint galaxies.

The selection functions are shown in Fig. 5. The main shape, shown by the cyan line, directly reflects the *LSST golden selection cut*. The selection function with the Gaussian error model is very close to the one using spectro-z, but the ratio of the two functions exhibits an expected slight warp as the smoothing of a power-law-like distribution leads to a flattening of this distribution. Cases with photo-z error models are very different: the ratio between the selection function of any of the photo-z cases and the spectro-z case exhibits significant structures. Asymmetric  $e_z$  distributions are responsible for these visible distortions since they induce migration of galaxies from true redshift distribution into a distorted measured redshift distribution. It is particularly visible around  $z = 0.5$  (see Fig. 6, top left). Difficulties in the photo-z reconstruction with asymmetric distributions around this

redshift are rather generic, for instance it has been deeply studied by the SDSS collaboration Beck et al. (2016).

Additionally, catastrophic redshift reconstructions significantly affect the selection function at high redshift, let's say above  $z = 1.6$ . Indeed, a small fraction (a few percents) of the large low-redshift galaxy population can significantly contaminate (up to 30 %) the much less populated regions at high redshift. The use of the BDT cut makes the selection function slightly smoother as the IQR, the bias and the number of catastrophic photo-z are significantly reduced. The cost is an additional decrease of the number of usable galaxies, from 2 % to less than 1 % at  $z = 2$  for instance.

It is interesting to pay attention on the origin of the bias. A residual bias can be corrected for only if it is attributed to a determination method that is known to systematically over or underestimate a measurement. It appears not to be the case here, as illustrated by Fig 6. The normalised histograms of  $e_z$  have been plotted for 4 bins in redshift chosen for their various bias values (visible as the gap between the dashed black and the coloured vertical lines). The  $e_z$  distributions have their maximum near zero, but when they are strongly asymmetric, the means differ significantly from the median. Here, applying global shifts to the photo-z distribution would be improper. Around  $z_p=0.51$  or 1.89, the bias in the photo-z case is far beyond the requirements but the BDT cut significantly decreases the asymmetry of the  $e_z$  distribution and the bias becomes marginally compatible with the requirements at the cost of a significant loss of galaxies. At the opposite, at  $z_p=1.05$  or 1.20, the bias is small, close to the requirement, and the IQR and the fraction of outliers are satisfactory (see Fig. 4): there is clearly a systematic effect but it remains compatible with the requirements and the BDTcut is not able to improve the symmetry of the  $e_z$  distributions.



**Fig. 6.** Normalised histogram of  $e_z = (z_s - z_p)/(1 + z_p)$  for 4 bins in photo-z. The coloured vertical lines show the value of the bias for the 3 error models on the photo-z considered (see Fig. 5 for colour legend) while the dashed black line is the reference at 0.

$z$ range	0.2-0.5	0.5-1.0	1.0-1.5	1.5-2.5	0.2-2.5
$z = z_G$	6.7	19.0	9.3	2.5	37.5
$z = z_p$	6.2	19.7	9.3	2.4	37.6
$z = z_{\text{BDT}90}$	5.5	18.6	8.8	1.7	34.6
$z = z_{\text{BDT}80}$	4.9	17.4	8.0	1.1	31.4

**Table 1.** Number of galaxies per arc-min<sup>2</sup> per redshift range. The *LSST golden selection cut* is always applied.

$f_{\text{gal}}$ with	$ e_z  < 0.02$	$ e_z  < 0.05$	$ e_z  < 0.15$
$z = z_G$	49 %	90 %	100 %
$z = z_p$	61 %	85 %	97 %
$z = z_{\text{BDT}9}$	65 %	89 %	99 %
$z = z_{\text{BDT}8}$	68 %	91 %	100 %

**Table 2.** Fraction of galaxies  $f_{\text{gal}}$  with  $|e_z| = |z_s - z_p|/(1+z_p)$  lower than several thresholds for the three photometric error models. The threshold of 0.02 is the goal, 0.05 is the requirement ; 0.15 defines the outliers.

### 3.3. Number density of galaxies

The redshift evolution of the galaxy number density (in [Mpc<sup>-3</sup>]), for galaxies satisfying the *LSST golden selection cut*, with and without BDT quality cut is shown in top panel of Fig. 5 (right vertical axis). Additionally, the average galaxy number density for few redshift ranges is given in Tab. 1. The impact of the error model on the galaxy density is driven by the BDT cut: 10 or 20 % of the galaxies are lost with the BDT 90% or BDT 80% cuts, by definition. One can notice the differences on the number of galaxies between the first two redshift bins, using Gaussian or photo-z redshifts: we retrieve the partial redistribution of the galaxies around  $z = 0.5$ , already seen in the shape of the selection functions, when photo-z are used.

Table 2 gives the fraction of galaxies with relative redshift difference  $e_z$  lower than some thresholds. The LSST requirement on the fractional photo-z error is 5 % with a goal of 2 %. Galaxies with a fractional photo-z error of 15 % are considered as outliers. The three columns of Tab. 2 correspond to these three thresholds.

## 4. Computation on the power spectra at different redshifts

Our aim is to quantify the impact of photometric redshifts on the LSS power spectrum  $P(k)$  and the BAO scale  $s_A$  determination in LSST. We have thus chosen to use a Fourier Transform based procedure, which is a rather simple, robust and well tested method to determine the LSS power spectrum. Other methods, such as direct computation of the two-point correlation function are often used when analysing observed galaxy catalogs, Szalay et al. (2003) or Anderson et al. (2012) for instance. These methods can indeed be better suited to real data, as they can handle more easily inhomogeneous sampling of space, blind spots due to bright stars for instance, which require the use of masks, but more time consuming than FFT. Nevertheless, difficulties related to Fourier-space analysis can be overcome Beutler et al. (2016).

Assuming a fiducial cosmology, galaxy angular positions and redshifts are converted into cartesian positions which are used to project the galaxies (weighted by the inverse of the selection function) into a 3D-grid  $n_{\text{obs}}(\mathbf{r})$ . There is a difficulty for this grid-FFT method. Cartesian grids are not well adapted to the spherical geometry and a clean definition of the redshift range. We have overcome this difficulty by the use of several grids, each

set paving the spherical shell in a given redshift range, as described below.

A Fast Fourier Transform (FFT) is so applied to compute the Fourier modes  $F(\mathbf{k})$  of the grid  $n_{\text{obs}}(\mathbf{r})$  and the power spectrum of each normalised, mean subtracted galaxy number density grid:

$$F(\mathbf{k}) = \text{FFT} \left[ \frac{n_{\text{obs}}(\mathbf{r})}{\bar{n}_{\text{obs}}} - 1 \right], \quad P_{\text{obs}}(k) = \langle |F(\mathbf{k})|^2 \rangle.$$

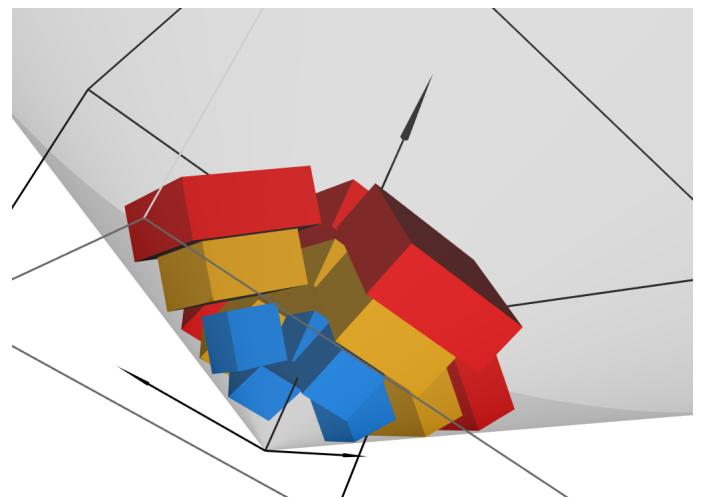
The BAO scale is then extracted from the power spectrum through a fit of a damped sinusoid to the ratio of the observed power spectrum to a wiggle less (no BAO) power spectrum (see Sect. 4.3).

### 4.1. Multiple grids covering a redshift range

First, the volume of universe covered by the simulated galaxy catalog is organised as several sets of 3D-grids, each set covering a given redshift range. Euclidian ( $x$ ,  $y$ ,  $z$ ) axes are defined with respect to the observer, with the  $Oz$  axis corresponding to the line of sight through the center of the field. A given direction  $\Omega = (\theta, \varphi)$  is defined by two angles,  $\theta$  being the angle with respect to  $Oz$  and  $\varphi$  the angle of the projection of  $\Omega$  on the  $xOy$  plane with  $Ox$  axis.

The grid thicknesses are a compromise between a reasonably small redshift range, sufficient statistics and large enough volume to be sensitive to low  $k$ -values ( $k_{\min} \sim 0.005$ ). The margin with respect to the initial *BigCube* limits must be sufficient to avoid artificial loss of galaxies, as the photo-z errors can be large, especially at high redshifts.

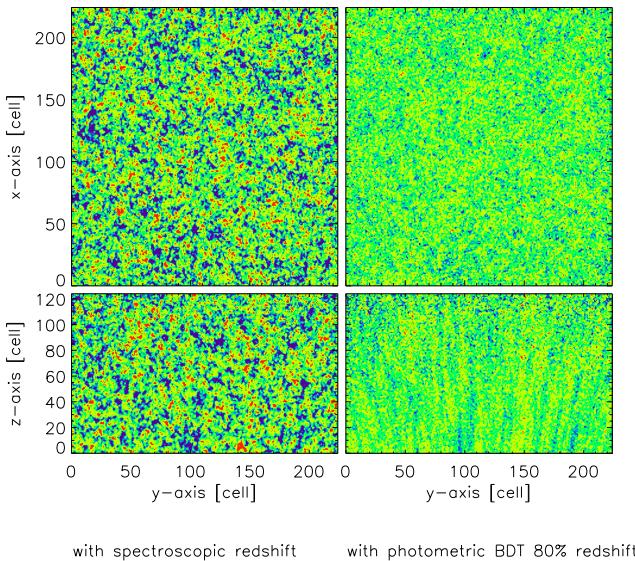
We have defined three sets of 5 grids each, centred at redshifts  $z = 0.5, 0.9, 1.3$ , each grid being subdivided into  $8 \times 8 \times 8$  Mpc<sup>3</sup> cells. One grid is centred along the Euclidian  $z$  axis while the others are rotated by 40 degrees in  $\theta$  and scattered at 0, 90, 180, 270 degrees in  $\varphi$ . A 3D-view of the grids is drawn in Fig. 7: the initial *BigCube* is represented by the white box, the field-of-view by the grey area while grids are drawn in colours according to their redshift range.



**Fig. 7.** Schematic view of the volumes used in the simulation. The black lines mark out the *BigCube*. The grey area shows the field of view of  $\pi$  sr which is filled by galaxies from  $z = 0.2$  to  $z = 2.45$ . The 5 grids used to compute each power spectrum are drawn in blue, orange or red according to their central redshift.

Table 3 summarises the main characteristics of these three grid sets. After projection of the galaxy catalog into the grids, each cell will contain the number of galaxies with a position falling in it, weighted by the inverse of the selection function corresponding to the considered error model at the appropriate redshift.

Figure 8 shows the galaxy density contours of two slices through the center of the  $z = 0.9$  central grid. The  $(x, y)$  slice corresponds to the transverse plane and the  $(z, y)$  slice contains the line-of-sight  $Oz$  axis. While density contours are isotropic and small structures are well contrasted with spectro-z (left panels), features appear along the radial (redshift) direction and small structures have faded out with photo-z (here BDT 80%, right panels).



**Fig. 8.** Galaxy density contours in slices of the grid centred at  $z = 0.9$ . Top: the transverse plane ( $y, x$ ); bottom: the radial plane ( $y, z$ ). Left : with spectro-z, right : with BDT 80% photo-z. Slices are 8 comoving Mpc thick (1 cell) and go through the grid center.

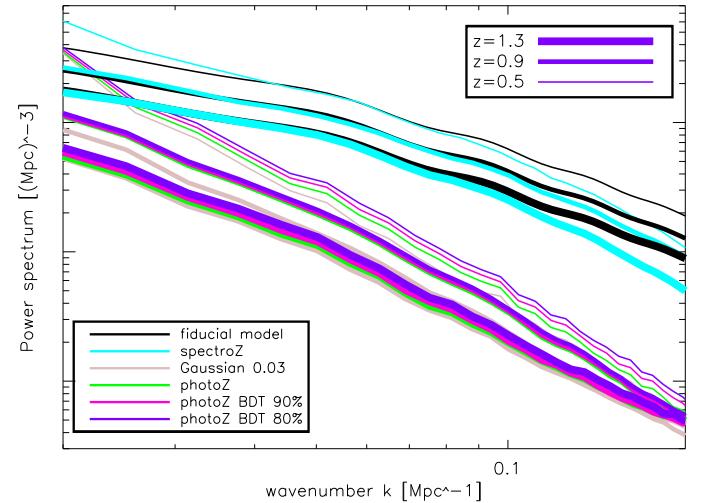
#### 4.2. Power spectra and noise estimation

We have computed power spectra  $P_{\text{obs}}(k)$  for each of the five redshift error models: spectroscopic redshifts (no-error), redshifts with gaussian error  $\sigma_z = 0.03(1 + z)$  and photo-z reconstruction without or with 90% or 80% BDT cut.

We also compute the shot-noise contribution by simulation. A separate set of grids is filled by Poisson noise using the mean galaxy density at the redshift of each cell. Further steps - application of the error model on the redshift or selection function correction for instance - are then applied as for grids filled by galaxies. The power spectra of the shot-noise grids  $P_{\text{SN}}$  are flat and the shot-noise contribution is properly approximated by a constant, as expected, which is determined with a small statistical uncertainty.

The shot-noise subtracted power spectrum is defined by  $P_{\text{D}}(k) = \langle P_{\text{obs}}(k) \rangle_{\text{3D-grids}} - P_{\text{SN}}$ , where the subscript D stands for damped. Indeed, the recovered power spectrum is damped due to redshift errors, compared to the underlying galaxy distribution power spectrum.

Theoretical (input) and recovered power spectra from simulated galaxy catalogs, at the three redshifts and for the five redshift error models, are shown in figure 9.



**Fig. 9.** Recovered power spectra  $P_{\text{D}}(k)$  computed from the 5 grids centred at each redshift bin, after subtraction of the shot-noise contribution. Black lines correspond to the theoretical (input) power spectra while other colours refer to the five redshift error models. The line thickness identifies the grids central redshift, with the thickness increasing with the redshift.

One can see the global decrease in amplitude at all scales when the redshift increases, by comparing for instance the theoretical shapes (black lines). It is related to the increasing growth factor with cosmic time. The power spectra recovered from catalogs with spectro-z for the three redshift ranges (cyan curves) follow the theoretical shapes at low  $k$ , but a moderate damping can be seen, starting around  $k = 0.1$ , which is mainly due to the sampling with  $8^3$  Mpc $^3$  cells.

The damping of the power spectrum produced by the photometric redshift smearing is clearly visible when errors on redshift are introduced. The damping factor reaches a factor  $\approx 10$  at the BAO scale, around 150 Mpc ( $k \approx 0.04$ ).

The BDT cut reduces the photo-z dispersion, so reduces the damping: the recovered power spectrum amplitude increases with more stringent BDT requirement. Note that the differences are tiny between the green, pink and purple medium thickness curves, as photo-z performance is already good around  $z = 0.9$  without any BDT cut. The amplitude of the recovered spectra from the grids centred at  $z = 0.5$  and  $z = 1.3$  are more sensitive to the BDT cut.

The power spectra for the grids centred at  $z = 1.3$  start to flatten at high  $k$  because they are not far to be shot-noise dominated with an average of only two galaxies per cell (see Tab. 3).

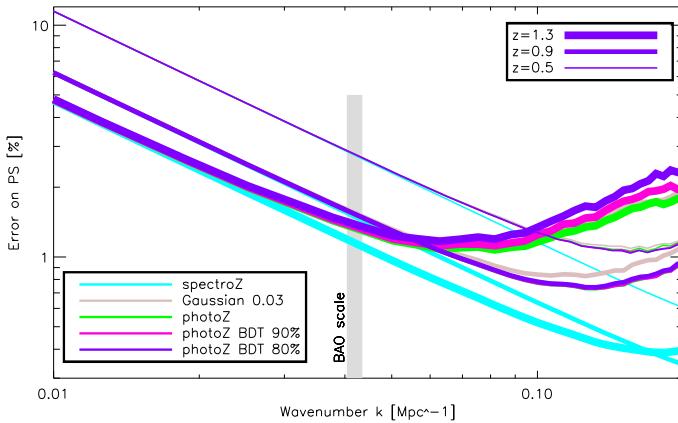
The statistical errors associated to the recovered power spectra are given by:

$$\sigma_P(k) = \frac{2\pi}{k\sqrt{V}\delta_k} \times [P_{\text{D}}(k) + P_{\text{SN}}] \quad (1)$$

where  $V$  is the total volume of the grids in a given redshift range and  $\delta_k$  is the sampling width in wavenumber. We have checked that the dispersion of the recovered power spectra from different mock catalogs do follow the above relation, although with limited number of catalogs, due to CPU and storage intensive computations needed to generate and analyse the catalogs.

grid central redshift	cells mean redshift	cells redshift range	width, thickness [ $n_{\text{cell}}$ ]	volume [Gpc $^3$ ] $1 \rightarrow 5$ grids	$N_{\text{gal}}$ with spectroZ [10 $^6$ ]	$\langle N_{\text{gal}} \rangle$ spectroZ [cell $^{-1}$ ]	$N_{\text{gal}}$ with BDT 90% [10 $^6$ ]	$\langle N_{\text{gal}} \rangle$ BDT 90% [cell $^{-1}$ ]	$N_{\text{gal}}$ with BDT 80% [10 $^6$ ]	$\langle N_{\text{gal}} \rangle$ BDT 80% [cell $^{-1}$ ]
0.5	0.51	0.36-0.68	120, 125	$0.9 \rightarrow 4.6$	26.2	14.5	24.1	13.4	21.3	11.8
0.9	0.93	0.72-1.19	225, 125	$3.2 \rightarrow 16.2$	52.3	8.7	51.1	8.1	48.4	7.6
1.3	1.36	1.08-1.73	300, 125	$5.7 \rightarrow 28.8$	26.9	2.4	25.9	2.2	22.2	2.0

**Table 3.** Geometrical description of the grids used for the power spectrum computation and mean number of galaxies falling in one of the five grids in each set, with and without BDT cut. Grids are composed of  $w \times w \times t$  cells of  $8^3$  Mpc $^3$ , with  $w$  and  $t$  the width and thickness values given in the table.



**Fig. 10.** Fractional statistical error  $\sigma_P/P(k)$  of the recovered power spectra  $P_D$  in percent. Colors identify the different redshift error models, while the grid redshifts are distinguished by different line thicknesses. The light gray area shows the wave number corresponding to the BAO scale.

The fractional statistical uncertainty on the recovered power spectrum  $\sigma_P/P(k)$  is plotted in Fig. 10. It depends naturally on the wavenumber  $k$ , but also on the redshift interval and on the redshift error model. One can distinguish two regimes on these curves:

- The fractional power spectrum error  $\sigma_P/P(k)$  is dominated by the cosmic variance at low wavenumber ( $k < 0.015$ ), for all of the five error models and all redshift intervals. It remains true with spectroscopic redshifts up to, at least,  $k \approx 0.15$ . The cosmic variance contribution is larger for low redshift as the grids are smaller and it evolves as  $1/k$ .
- At medium or high wave-numbers, the fractional statistical error  $\sigma_P/P(k)$  is dominated by the shot-noise contribution. The lower limit in wave number for this regime depends on the redshift range.

The shot-noise contribution is significantly lower when spectro-z are used. Indeed the shot-noise levels, which depend on the mean galaxy density, are very similar for the spectroscopic, Gaussian and photometric cases. However, their relative value with respect to the power spectrum  $P_D(k)$  increase significantly as  $P_D(k)$  are damped due to radial smearing (Gaussian or photo-z error models).

Note that even with spectro-z, for the 3D-grids centred at  $z = 1.3$ , the relative error contribution flattens at  $k \approx 0.2$ , as the shot-noise starts to overwhelm the LSS power spectrum.

In summary, we can expect a more accurate BAO scale determination from the grids centred at  $z = 0.9$ , compared to the grids centred at  $z = 0.5$  and  $z = 1.3$ . The precision on the recovered power spectrum is limited by the cosmic variance at low red-

shift ( $z = 0.5$ ) and by the shot-noise at high redshift ( $z = 1.3$ ). Note that non-linearities will also soften the oscillations above  $k \approx 0.1$  Mpc $^{-1}$ .

#### 4.3. Extraction of the BAO scale

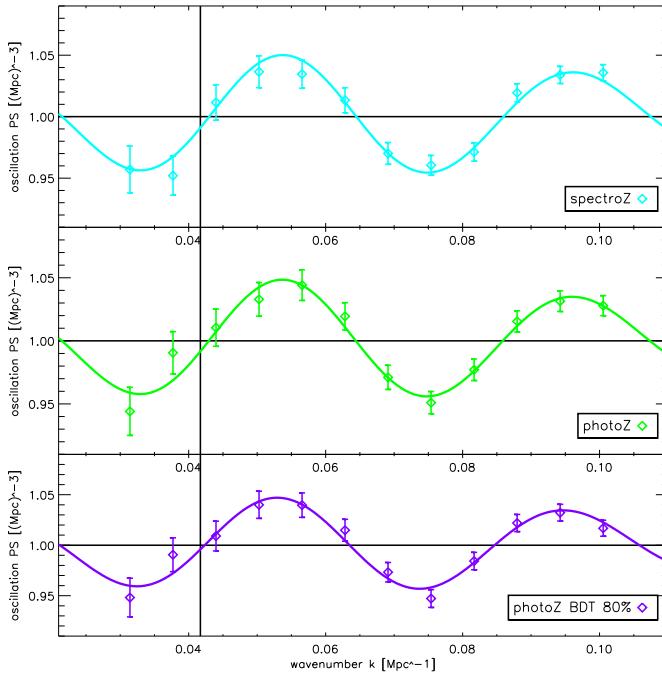
The baryon acoustic oscillations are subdominant with respect to the global matter power spectrum shape, as they are hard to see even on theoretical curves (black lines on Fig. 9). Thereby, the power spectrum, damped by any feature affecting the data or the computation method, follows a global shape with the small superimposed oscillations.

We do not want to assume any shape of the damping induced by the smearing produced by photo-z errors. So we cannot use an analytical model as it is done for Gaussian error model Glazebrook & Blake (2005). The appendix contains the description of the procedure that we have developed to estimate the smooth, wiggle-less, power spectrum from the observed one. The oscillating component in the spectrum is extracted, through the fitting of a damped sinusoid, similar to the Wiggle only method Glazebrook & Blake (2005) where the amplitude, the damping scale and the oscillation scale are left as free parameters.

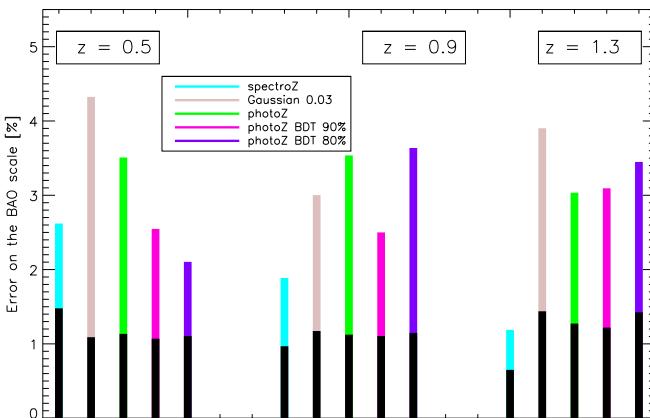
As mentioned in Sect. 2, we have used a simple model in this simulation, ignoring non-linear effects or bias on the LSS power spectrum and mock galaxy catalog generation. Indeed, non-linear clustering affects the power spectrum, leading in particular to a damping of the BAOs features at small scales Crocce & Scoccimarro (2008), Rasera et al. (2014), Obuljen et al. (2017) or Seo & Eisenstein (2007). In order to limit over-estimating LSST capability to recover the BAO scale, we have restricted the  $k$ -range used to extract the oscillating component in the power spectrum and to determine the BAO scale  $s_a$ . We have used two  $k$ -ranges, a very conservative one where only  $k \leq 0.1$  h/Mpc $^{-1}$  (i.e.  $k \leq 0.07$  Mpc $^{-1}$ ) have been kept, and a second one, using wave modes up to  $k \leq 0.15$  h/Mpc $^{-1}$  (i.e.  $k \leq 0.1$ /Mpc $^{-1}$ ). Indeed the impact of non linear clustering for low  $k$ -modes ( $k \leq 0.1$  h/Mpc $^{-1}$ ) can be safely neglected. However, as one can see for instance from Rasera et al. (2014), the damping of the BAO features due to non-linear clustering is limited up to 0.15 h/Mpc $^{-1}$ , specially at higher redshifts ( $z \geq 1$ ), which is more in focus in this work. The comparison of the reconstructed BAO scale error from these two  $k$ -ranges gives an indication of the amount of information in larger  $k$ -modes for different redshift bins.

For illustration purpose, we show the oscillating component of the LSS power spectrum in the Fourier space, as well as the fitted damped sinusoid, for the redshift  $z = 0.9$  on Fig. 11.

The estimated errors on  $s_A$  are gathered in Fig. 12 for the two tested  $k$ -ranges. In agreement with previous considerations, results obtained for 3D-grids centred at  $z = 0.9$  are more precise than results obtained from 3D-grids centred at lower or higher



**Fig. 11.** Recovered oscillation part of the power spectra in Fourier space for the grids centred at  $z = 0.9$  considering, from top to bottom, spectro-z, photo-z without and with BDT 80% cut. The coloured lines are the fitted oscillations using the simple “Wiggle only” description. The BAO scale is at the first intersection between this fitted damped sinusoid and the horizontal black line. The expected value of  $k_A$  from the fiducial Planck 2015  $\Lambda$ CDM model ( $0.0417 \text{ Mpc}^{-1}$ ) is shown by the vertical line.



**Fig. 12.** Relative errors on the BAO scale fitted from the power spectra for each redshift and each error model. The black bars are the results using the  $k$ -range  $[0.03-0.1] \text{ Mpc}^{-1}$  while the color bars are the results using the  $k$ -range  $[0.02-0.07] \text{ Mpc}^{-1}$ .

redshifts. The results are two or three times worse if the  $k$ -modes between  $0.07$  and  $0.1 \text{ Mpc}^{-1}$  are removed. Indeed the loss of half of the second oscillation impacts the ability to precisely recover the BAO scale. Nevertheless it does not suppress it as the error remains below 5 % in all cases.

The errors are probably underestimated, especially in some cases, so it prevents from proper estimation of the impact of the BDT cut on the recovery of the BAO scale. Indeed, the reduced  $\chi^2$  are reasonably close to 1 when spectro-z are used and for all

error models for grids centred at  $z = 0.9$  and  $z = 1.3$ . But the reduced  $\chi^2$  ranges from 5 to 6 when the fit is performed from grids centred at  $z = 0.5$  with Gaussian or photometric error on the redshift. We interpret this as a hint that another component in the error computation related to the error on the redshift should be included. This component is mainly hidden by the shot-noise at higher redshifts. It is due to the redshift dispersion and not to the bias or the presence of outliers as the  $\chi^2$ 's are similar with the Gaussian error model (with no bias and virtually no outlier) and with the photo-z error model.

## 5. Interpretation and discussion

We will try to explain here some of the general behaviours that we have observed using the full simulation in the previous section. We will then use a simple simulation tool implementing these effects to get BAO scale determination uncertainty in a LSST-like survey, compared to a fiducial spectroscopic survey.

The use of photo-z induces two main effects on the galaxy distribution. A first one is the smearing along the radial direction, due to reconstructed redshift uncertainties, and a second one is the presence of outliers, due to catastrophic redshifts. We can write the observed galaxy number densities  $n_{obs}(\mathbf{r})$  along the line of sight  $\mathbf{r}$  from which the underlying LSS matter power spectrum is computed as:

$$n_{obs}(\mathbf{r}) = n_g^{pz}(\mathbf{r}) + n_{out}(\mathbf{r})$$

where  $n_g^{pz}$  (resp.  $n_{out}$ ) denotes galaxy population with reasonable photo-z (resp. outliers). Introducing the outlier fraction  $f_{out}$ , we have the following relations between the different average number densities:

$$\bar{n}_g^{pz} = (1 - f_{out}) \bar{n}_{obs}; \quad \bar{n}_{out} = f_{out} \bar{n}_{obs}.$$

The observed matter distribution power spectrum is computed from the galaxy number density field renormalised by the average density :

$$P_{obs}(k) : \text{FFT}[n_{obs}(\mathbf{r})/\bar{n}_{obs} - 1].$$

Neglecting the correlation between galaxies with correctly reconstructed redshifts and outliers, the observed power spectrum can be written as:

$$P_{obs}(k) = (1 - f_{out})^2 P_g(k) \times \eta(k) + f_{out}^2 P_{out}(k) + P_{SN}.$$

The different terms of this equation correspond to:

- The  $P_g(k) \times \eta(k)$  term corresponds to the power spectrum of the galaxy number density field radially smeared due to photo-z errors, leading to a damping of the power spectrum ( $0 < \eta(k) \leq 1$ ).  $\eta(k)$  follows a scaling law:  $\eta(k) = f(k \times \sigma_R)$  with  $\sigma_R$  the standard deviation of the smearing along the radial direction as described in [Blake & Glazebrook \(2003\)](#).
- Outliers correspond to galaxies from different redshifts shuffled to a large extent, so  $P_{out}(k) \lesssim P_g(k)$ ; they contribute to the overall noise in the observed galaxy power spectrum. However, the outliers contribution would be negligible in most cases, as long as  $f_{out} < 10\%$ .
- $P_{SN}$  is the combined shot-noise power spectrum from the two populations:

$$P_{SN} = \frac{\bar{n}_g}{\bar{n}_{obs}^2} + \frac{\bar{n}_{out}}{\bar{n}_{obs}^2} = \frac{(1 - f_{out}) \bar{n}_{obs}}{\bar{n}_{obs}^2} + \frac{f_{out} \bar{n}_{obs}}{\bar{n}_{obs}^2} = \frac{1}{\bar{n}_{obs}}$$

We have computed the damping function  $\eta(k)$ , corresponding to the ratio of the recovered power spectrum to the input power spectrum, for different levels of smearing along the radial direction. The left panel of Fig. 13 shows the damping function  $\eta(k)$  for four values of the standard deviation  $\sigma_R = 50, 75, 100, 150 \text{ Mpc}$  of Gaussian smearing along the radial direction. The damping function follows very well a scaling law,  $\eta(k) = f(k/k_0)$  where  $k_0 = \frac{1}{\sigma_R}$ , as shown in Fig. 13. The damping function is well represented by the following analytic function

$$\eta(k) = \frac{1}{\sqrt{1 + (k/k_0)^2}} \quad \text{with} \quad k_0 = \frac{1}{\sigma_R}$$

We have also plotted  $\sigma_R(z) = \frac{c}{H(z)} \times \sigma_{pz}(z)$  as a function of redshift on the right panel of Fig. 13, using the redshift dependent  $\sigma_{pz}(z)$  from full photo-z reconstruction of Sect. 3 with BDT 90% selection cut.

We have carried a study of the BAO scale reconstruction uncertainty using a toy-program where the observed galaxy power spectrum and associated errors are computed using the expressions given above. The galaxy power spectrum has been damped using a redshift dependent radial smearing  $\sigma_R(z)$ , as represented by the IQR variable in Fig. 4, and a fixed value of 5% for the outlier fraction. Note that  $\sigma_R(z)$  is very irregular, as the IQR is in Fig. 4; this variability is not due to a lack of statistics but to galaxies spectra features seen by the photo-z method. We have used redshift dependent galaxy number density following the values obtained in our full simulation, slightly below the ones shown in Fig. 2. We computed the ratio of the observed power spectrum to a theoretical power spectrum without BAO oscillations  $R(k) = (P_{obs}(k) - P_{SN}) / (P_{no-osc}(k) * \eta(k))$ . The BAO scale  $s_A$  is then determined by fitting a damped sinus function  $A \exp(k/k_d) \sin(s_A k)$  to the ratio  $R(k)$ , with three free parameters:  $A, k_d, s_A$ , over the restricted  $k$ -range  $0.03 \leq k \leq 0.1 \text{ Mpc}^{-1}$ , equivalent to  $0.04 \leq k \leq 0.15 \text{ hMpc}^{-1}$ . A constant quadratic error of 0.4% has been included to compensate partially for the simplicity of this fitting method. The increase of the uncertainty due to photo-z smearing is however marginally sensitive to the fitting procedure. The error is determined from the distribution of reconstructed BAO scales  $s_A$ .

The survey volumes given in Tab. 3 correspond to a survey of slightly less than 4,000 deg<sup>2</sup>. Using our toy-program and for  $\Omega_{surv} = 4,000 \text{ deg}^2$ , we obtain a BAO scale uncertainty  $\sigma_{s_A}$  of 1.9 Mpc ( $\approx 1.3\%$ ) and 2.2 Mpc ( $\approx 1.5\%$ ) for spectro-z and photo-z, with  $k_{max} = 0.1 \text{ hMpc}^{-1}$  for the redshift bin centred at  $z = 0.9$ , in reasonable agreement with the error bars represented in Fig. 12. It should be stressed that the uncertainties given by the toy-program should be considered as a lower limit of the true error bars, as the galaxy power spectrum and the one without BAO oscillations, as well as the damping function  $\eta(k)$  are supposed to be perfectly known. However, the comparison between spectro-z and photo-z remains meaningful.

We have used this toy-program to compute  $\sigma_{s_A}$  for an LSST-like survey and an effective survey area  $\Omega_{surv} = 10,000 \text{ deg}^2$ , in six redshift bins in the range [0.4 - 2.2]. The results are summarised in Tab. 4, where one can see the increase of the error on the BAO scale due to photo-z. Despite high galaxy number density, the relatively large value of  $\sigma_{s_A}$  around  $z \sim 0.5$  can be explained by the more limited survey volume and larger photo-z dispersions, which has an even larger impact due to smaller Hubble parameter  $H(z)$ . The BAO scale would be rather well constrained by the LSST survey in the redshift range  $0.7 \leq z \leq 1.6$ , thanks to good photo-z reconstruction and large galaxy number densities. The degradation of BAO scale determination is

visible for redshifts above  $z \sim 1.6$  and even more clearly above  $z \sim 2$ , where  $\sigma_{s_A}$  is increased by a factor  $\approx 4$  in the case of photo-z, compared to a full spectro-z.

We have also included forecasts of  $\sigma_{s_A}$  for a fiducial spectroscopic survey with realistic galaxy number densities in Tab. 4. Spectroscopic surveys naturally have a significantly lower galaxy number densities than photometric surveys which limit their statistical power to map the LSS. To determine the galaxy number density for our fiducial spectroscopic survey, we have used the foreseen number of objects in the DESI survey Aghamousa et al. (2016). They expect to reach  $\sim 1625$  objects per deg<sup>2</sup> up to redshift  $\sim 2.1$ , by considering LRG's up to redshift 1 ( $0.4 < z < 1.0$ ), ELG's in the redshift range  $0.6 < z < 1.6$  and QSO's up to redshift  $z \sim 2.1$ . This galaxy density is about 75 times less than the one expected for LSST ( $\gtrsim 125000$  per deg<sup>2</sup>). However, to be conservative, we have applied a factor 50 to the LSST galaxy number densities for the first 4 redshift bins, up to  $z < 1.6$ , and a factor 30 for the last two redshift bins  $1.6 < z < 2.2$  to obtain the galaxy densities  $\bar{n}_{gal}$  for computing forecasts for our fiducial spectroscopic survey.

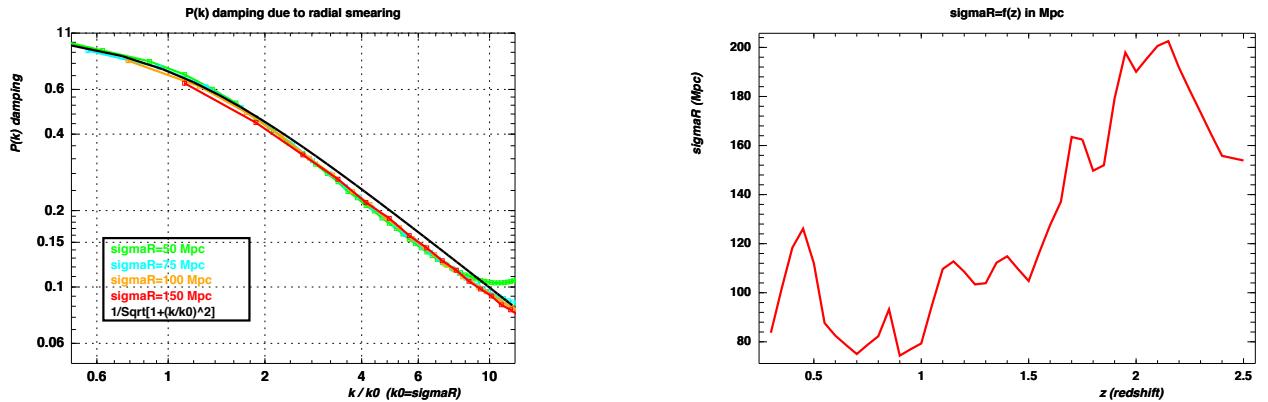
The amount of information available in the high  $k$ -modes can be evaluated by comparing the gain in the BAO scale precision by extending the  $k$ -range in the fit, *i.e.* by fixing  $k_{max}$  to 0.1 instead of  $0.07 \text{ Mpc}^{-1}$ . Such an extension would lead to a decrease of the relative uncertainty  $\sigma_{s_A}$  by almost a factor of 2. It is probably slightly less if the natural smoothing produced by non-linear effects is taken into account. The gain is more limited at higher redshifts as the shot-noise becomes dominant. The ability to properly model the non-linear part of the LSS power spectrum may give access to a still larger  $k$ -range. A hint of the improvement is given by the last column, for a fiducial spectroscopic survey, with oscillations fitted up to  $k = 0.2 \text{ Mpc}^{-1}$ . Note that such small scales are mainly whited out, hidden by the shot-noise, if photo-z are used (see Fig. 10).

Comparing the columns of Tab. 4 shows the robustness of the BAO scale as a standard ruler, even if a relatively small fraction of the galaxies is used or if the redshift suffers from some dispersion. This parameter is statistically mostly immune to these defects below  $z \approx 1.5$ . Nevertheless, the LSS power spectrum is significantly damped if photo-z are used. So using the whole shape, and not only the BAO scale, would require a very precise knowledge of the photo-z properties at the origin of the damping function  $\eta(k)$ , while it is more directly provided by spectro-z surveys.

Here, only the isotropic BAO scale has been considered. However, it should be kept in mind that the BAO probe is more powerful in spectroscopic surveys, as they will be able to determine the BAO scale in the transverse and radial directions independently, and to measure the redshift space distortions, while LSST will be, mostly, sensitive to the transverse BAO scale due to photo-z smearing.

## 6. Conclusion and perspectives

The determination of accurate photometric redshifts of galaxies for redshifts up to 2, at least, is one the main challenges of the LSST survey. We have evaluated the impact of realistic photo-z uncertainties on the LSS power spectrum and the BAO scale uncertainty. The number density of usable galaxies and its evolution with redshift is one of the main parameters with major impact on cosmology with LSST. We have determined the expected galaxy number density in the LSST survey, which decreases from  $\sim 0.015 \text{ Mpc}^{-3}$  at redshift  $z \sim 1$  to  $\sim 0.0015 \text{ Mpc}^{-3}$  at redshift  $z \sim 1$  and associated photo-z errors. We have shown



**Fig. 13.** Left: damping function  $\eta(k)$  for four values of  $\sigma_R = 50, 75, 100, 150$  Mpc (colour curves) and its analytical approximate shape  $(\sqrt{1 + (k/k_0)^2})^{-1}$  in black. Right: value of the radial smearing  $\sigma_R$  computed from photo-z reconstruction dispersion (with BDT 90%) as a function of redshift.

redshift [ $z_{min}, z_{max}$ ]	ideal survey			photometric survey			spectroscopic survey		
	$\bar{n}_{gal}$	$\sigma_{s_A} [\%]$	$k_{max1}$	$\bar{n}_{gal}$	$\sigma_{s_A} [\%]$	$k_{max1}$	$\bar{n}_{gal}$	$\sigma_{s_A} [\%]$	$k_{max1}$
[0.4 – 0.7]	1500	1.1	0.7	1500	1.2	0.7	30	1.3	0.9
[0.7 – 1.0]	1000	1.0	0.6	1000	1.2	0.7	20	1.4	0.9
[1.0 – 1.3]	600	0.9	0.6	600	1.0	0.6	12	1.6	1.1
[1.3 – 1.6]	300	0.9	0.5	300	1.1	0.8	6	2.9	2.0
[1.6 – 1.9]	130	0.9	0.6	130	1.6	1.3	4	4.9	4.0
[1.9 – 2.2]	60	1.1	0.7	60	4.2	3.4	2	11.5	10.0

**Table 4.** BAO scale determination relative uncertainty  $\sigma_{s_A}$  for different redshift ranges, for an ideal survey (left), for an LSST-like photometric survey (middle), and for a fiducial spectroscopic survey (right) with  $\Omega_{surv} = 10,000 \text{ deg}^2$ .  $\bar{n}_{gal}$  is the galaxy number density in units of  $10^{-5} \text{ Mpc}^{-3}$ . The spectro-z redshifts are used for the ideal and the spectroscopic survey, and photo-z with errors similar to the reconstructed photo-z with the BDT 90% quality cut are used for the photometric survey. The columns giving  $\sigma_{s_A}$  differ by the  $k$ -range used:  $k_{max1}$  means  $k \leq 0.07 \text{ Mpc}^{-1}$ ,  $k_{max2}$  means  $k \leq 0.1 \text{ Mpc}^{-1}$  and  $k_{max3}$  means  $k \leq 0.2 \text{ Mpc}^{-1}$ .

that LSST should be able to recover the BAO scale with few percents uncertainty up to redshift  $z \lesssim 1.5$ .

Around  $z = 1$ , photo-z are accurate, with low bias and there is a small fraction of outliers. Therefore, LSST should be able to map accurately the galaxy distribution at this crucial stage of the history of the Universe, when the dark energy starts to dominate over matter. The redshift of closer or more distant galaxies is less accurately measured in LSST, due to degeneracies between redshift and galaxy type. At low redshifts, the use of additional information like a BDT cut is helpful, leaving enough galaxies to properly compute the matter power spectrum and extract the BAO scale. At high redshifts ( $z \gtrsim 1.5 - 2$ ), the observed galaxy number density is small and cannot compensate any more the photo-z smearing and the BAO signal is mainly washed out. Our studies indicate that LSST will be able to recover with precision the BAO scale up to  $z \approx 1.5$ , with a photo-z reconstruction performances within requirements, confirming previous studies *e.g.* Zhan & Tyson (2018). However, our study also shows that previous conclusions were rather optimistic regarding higher redshifts, as the recovered LSS power spectrum suffers larger uncertainties, due to the combination of a lower galaxy number density and poorer photo-z reconstruction.

Using spherical shells and spherical harmonic transform instead of Euclidian grids and Fourier transform is better suited to the analysis of galaxy surveys with large sky coverage. In addition, the theoretical framework for LSS analysis with spherical shells has been extensively developed in recent years, for instance Bonvin & Durrer (2011), Lanusse et al. (2015), Campagne et al. (2017). We plan to carry a study of the impact of photo-z with separate determination of the radial and transverse BAO scales using spherical shells. At low redshift, specifically around  $z \sim 0.5$ , the LSS power spectrum uncertainty is not limited by shot-noise in LSST, but suffers from photo-z spread and outliers. Identifying a population of galaxies with reliable photo-z at low redshifts is another useful follow-up study.

*Acknowledgements.* Authors thanks the DESC Large-Scale Structures working group, especially D. Alonso and A. Slosar, for their useful comments. Authors thanks M. Moneuse for fruitful discussions and uses of CAMEL, even if the cosmological parameters part exceeds the scope of this paper.

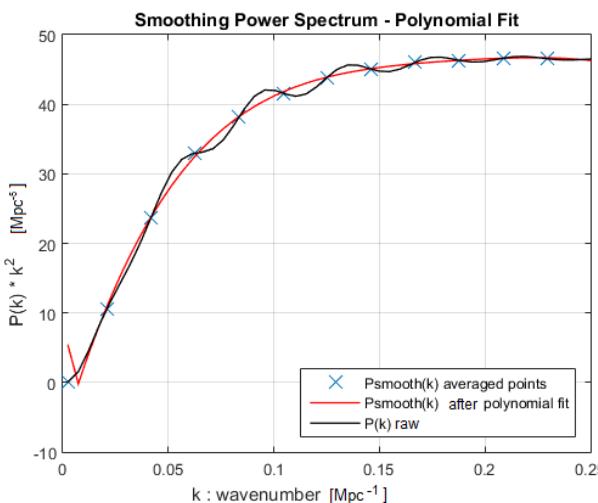
## Appendix A: Fit of the global shape of a power spectrum

The oscillating spectrum  $P_{\text{wiggle}}$  is fitted by the usual “Wiggle only” method, which aims at recovering the BAO scale  $s_A$ .

The polynomial description of the global shape  $B$  is determined as follows:

- computing abscissas of the inflection points where no contribution of baryonic oscillation is expected - they depend on the assumed value of the BAO scale,
- deriving from data the ordinates of these points by averaging the power spectrum around these abscissas,
- fitting this set of inflection points by a polynomial of fifth degree.

This method is illustrated by Fig. A.1. We have used power spectra weighted by the square of the spatial wave number ( $P(k) \times k^2$ ) to enhance the readability. The baseline derived by this method is smooth by construction and adapt itself to include all kinds of damping.



**Fig. A.1.** Example of a power spectrum and its fitted baseline. The initial power spectrum is drawn in black, the inflection points are indicated by the blue crosses and the red curve is the result of the fit, providing the global shape of the input power spectrum.

This method has been validated with results obtained from grids filled with galaxies derived from matter distribution unaffected by baryonic oscillations. Such a power spectrum, with baryons but without baryonic oscillation, are available in the Eisenstein approximation [Eisenstein & Hu \(1998\)](#).

## References

- Abbott, T. M. C. et al. 2017, Submitted to: Mon. Not. Roy. Astron. Soc., arXiv:1712.06209
- Abrahamse, A., Knox, L., Schmidt, S., et al. 2011, ApJ, 734, 36
- Aghamousa, A. et al. 2016, arXiv:1611.00036
- Alam, S., Ho, S., Satpathy, S., et al. 2017, Monthly Notices of the Royal Astronomical Society, 470, 2617
- Alonso, D., Bellini, E., Ferreira, P. G., & Zumalacárregui, M. 2017, Phys. Rev. D, 95, 063502
- Anderson, L., Aubourg, E., Bailey, S., et al. 2012, MNRAS, 427, 3435
- Aubourg, É., Bailey, S., Bautista, J. E., et al. 2015, Phys. Rev. D, 92, 123516
- Awan, H., Gawiser, E., Kurczynski, P., et al. 2016, ApJ, 829, 50
- Beck, R., Dobos, L., Budavári, T., Szalay, A. S., & Csabai, I. 2016, MNRAS, 460, 1371
- Benítez, N. 2000, ApJ, 536, 571
- Beutler, F., Blake, C., Koda, J., et al. 2016, MNRAS, 455, 3230
- Beutler, F., Seo, H.-J., Saito, S., et al. 2016, Monthly Notices of the Royal Astronomical Society, 466, 2242
- Blake, C., Kazin, E. A., Beutler, F., et al. 2011, MNRAS, 418, 1707
- Bonvin, C. & Durrer, R. 2011, Phys. Rev. D, 84, 063505
- Calzetti, D., Kinney, A. L., & Storchi-Bergmann, T. 1994, ApJ, 429, 582
- Campagne, J. E., Neveu, J., & Plaszczynski, S. 2017, A&A, 602, A72
- Cardelli, J. A., Clayton, G. C., & Mathis, J. S. 1989, ApJ, 345, 245
- Cavuoti, S., Tortora, C., Brescia, M., et al. 2017, MNRAS, 466, 2039
- Chang, C., Jarvis, M., Jain, B., et al. 2013, MNRAS, 434, 2121
- Coleman, G. D., Wu, C. C., & Weedman, D. W. 1980, The Astrophysical Journal Supplement Series, 43, 393
- Crocce, M. & Scoccimarro, R. 2008, Phys. Rev. D, 77, 023533
- Crocce, M. et al. 2017, arXiv:1712.06211
- Dahlen, T., Mobasher, B., Somerville, R. S., et al. 2005, ApJ, 631, 126
- Dawson, K. S., Schlegel, D. J., Ahn, C. P., et al. 2013, AJ, 145, 10
- Drinkwater, M. J., Jurek, R. J., Blake, C., et al. 2010, MNRAS, 401, 1429
- Eisenstein, D. J. & Hu, W. 1998, ApJ, 496, 605
- Glazebrook, K. & Blake, C. 2005, ApJ, 631, 1
- Gomes, Z., Jarvis, M. J., Almosallam, I. A., & Roberts, S. J. 2017, Monthly Notices of the Royal Astronomical Society, 475, 331
- Gorecki, A., Abate, A., Ansari, R., et al. 2014, A&A, 561, A128
- Graham, M. L., Connolly, A. J., Ivezić, Ž., et al. 2018, The Astronomical Journal, 155, 1
- Jones, D. M. & Heavens, A. F. 2019, Mon. Not. Roy. Astron. Soc., 483, 2487
- Kinney, A. L., Calzetti, D., Bohlin, R. C., et al. 1996, ApJ, 467, 38
- Lanusse, F., Rassat, A., & Starck, J. L. 2015, A&A, 578, A10
- Laureijs, R., Amiaux, J., Arduini, S., et al. 2011, arXiv e-prints, arXiv:1110.3193
- Lesgourges, J. 2011, arXiv e-prints, arXiv:1104.2932
- Levi, M., Bebek, C., Beers, T., et al. 2013, arXiv e-prints, arXiv:1308.0847
- LSST Science Collaboration, Abell, P. A., Allison, J., et al. 2009, arXiv e-prints, arXiv:0912.0201
- Mandelbaum, R., Miyatake, H., Hamana, T., et al. 2018, Publications of the Astronomical Society of Japan, 70, S25
- Medezinski, E., Oguri, M., Nishizawa, A. J., et al. 2018, Publications of the Astronomical Society of Japan, 70, 30
- Obuljen, A., Villaescusa-Navarro, F., Castorina, E., & Viel, M. 2017, Journal of Cosmology and Astro-Particle Physics, 2017, 012
- Pasquet, J., Bertin, E., Treyer, M., Arnouts, S., & Fouchez, D. 2019, A&A, 621, A26
- Planck Collaboration, Ade, P. A. R., Aghanim, N., et al. 2016, A&A, 594, A13
- Ramos, B. H. F., Pellegrini, P. S., Benoit, C., et al. 2011, AJ, 142, 41
- Rasera, Y., Corasaniti, P. S., Alimi, J. M., et al. 2014, MNRAS, 440, 1420
- Sadeh, I., Abdalla, F. B., & Lahav, O. 2016, Publications of the Astronomical Society of the Pacific, 128, 104502
- Salazar-Albornoz, S., Sánchez, A. G., Grieb, J. N., et al. 2017, MNRAS, 468, 2938
- Seo, H.-J. & Eisenstein, D. J. 2007, ApJ, 665, 14
- Süveges, M., Fotopoulou, S., Coupon, J., et al. 2017, in IAU Symposium, Vol. 325, Astroinformatics, ed. M. Brescia, S. G. Djorgovski, E. D. Feigelson, G. Longo, & S. Cavuoti, 39–45
- Szalay, A. S., Jain, B., Matsubara, T., et al. 2003, ApJ, 591, 1
- Tanaka, M., Coupon, J., Hsieh, B.-C., et al. 2018, Publications of the Astronomical Society of Japan, 70, S9
- Zhan, H., Knox, L., & Tyson, J. A. 2009, ApJ, 690, 923
- Zhan, H. & Tyson, J. A. 2018, Reports on Progress in Physics, 81, 066901
- Zucca, E., Bardelli, S., Bolzonella, M., et al. 2009, A&A, 508, 1217