

Automated rapid follow-up of *Swift* gamma-ray burst alerts at 15 GHz with the AMI Large Array

T. D. Staley,¹★ D. J. Titterton,² R. P. Fender,¹ J. D. Swinbank,³
A. J. van der Horst,³ A. Rowlinson,³ A. M. M. Scaife,¹ K. J. B. Grainge^{2,4}
and G. G. Pooley²

¹*School of Physics & Astronomy, University of Southampton, Southampton SO17 1BJ*

²*Astrophysics Group, Cavendish Laboratory, 19 J. J. Thomson Avenue, Cambridge CB3 0HE*

³*Astronomical Institute Anton Pannekoek, Science Park 904, PO Box 94249, 1090 GE Amsterdam, the Netherlands*

⁴*Kavli Institute for Cosmology Cambridge, Madingley Road, Cambridge CB3 0HA*

Accepted 2012 October 19. Received 2012 October 16; in original form 2012 September 28

ABSTRACT

We present 15-GHz follow-up radio observations of 11 *Swift* gamma-ray burst (GRB) sources, obtained with the Arcminute Microkelvin Imager Large Array (AMI-LA). The initial follow-up observation for each source was made in a fully automated fashion; as a result four observations were initiated within 5 min of the GRB alert time stamp. These observations provide the first millijansky-level constraints on prolonged radio emission from GRBs within the first hour post-burst. While no radio emission within the first six hours after the GRB is detected in this preliminary analysis, radio afterglow is detected from one of the GRBs (GRB 120326A) on a time-scale of days. The observations were made as part of an ongoing programme to use AMI-LA as a systematic follow-up tool for transients at radio frequencies. In addition to the preliminary results, we explain how we have created an easily extensible automated follow-up system, describing new software tools developed for astronomical transient alert distribution, automatic requesting of target-of-opportunity observations and robotic control of the observatory.

Key words: instrumentation: interferometers – methods: miscellaneous – methods: observational – virtual observatory tools – gamma-ray burst: general.

1 INTRODUCTION

Since the discovery of gamma-ray burst (GRB) afterglow emission at X-ray (Costa et al. 1997), optical (van Paradijs et al. 1997) and radio frequencies (Frail et al. 1997), it has been demonstrated that broad-band observations across the electromagnetic spectrum can be utilized to determine the micro- and macro-physical parameters of the GRB explosion and its immediate environment (e.g. Sari, Piran & Narayan 1998; Wijers & Galama 1999). After the initial detection in gamma-rays and X-rays, rapid follow-up observations are often essential to capture the GRB afterglow in the optical, before it fades. To this end, automated, robotic optical observations from smaller telescopes are now routine (e.g. Akerlof et al. 2003; Wozniak et al. 2005). In contrast, early follow-up observations of GRBs at radio wavelengths are rare (although some attempts have been made, see e.g. Green et al. 1995; Koranyi et al. 1995; Dessenne et al. 1996).

In this paper we explain how we have achieved fully automated rapid follow-up observations of GRB events with the Arcminute Microkelvin Imager Large Array (AMI-LA), a 15-GHz aperture synthesis radio telescope located near Cambridge, UK. A primary beamwidth of 5.5 arcmin and a sensitivity of $\sim 3 \text{ mJy s}^{-1/2}$ make this an ideal facility for follow-up of the well-localized GRB events observed by the *Swift* Burst Alert Telescope (*Swift*-BAT; Gehrels et al. 2004; Barthelmy et al. 2005).

This paper is structured as follows. First we give the scientific motivation for obtaining early-time radio follow-up observations (Section 2). A brief review of the relevant models for GRB emission is given, focusing upon predictions for radio emission at the time of the gamma-ray emission and within the first few hours thereafter. We then explain how this programme will provide novel data, complementary to existing GRB follow-up catalogues. In Sections 3–5 we describe the AMI-LA, the system we have put in place to initiate automated follow-up and our data reduction techniques. In Section 6, we report on system response times and present 15-GHz light curves. Finally, we describe our plans for further analysis of the AMI-LA data, and possibilities for follow-up with alternative facilities, in Section 7.

★ E-mail: tim.staley@soton.ac.uk

2 SCIENTIFIC MOTIVATION

2.1 GRB radio emission at short and long time-scales

The ‘relativistic fireball’ is the prevalent model to describe GRB emission across the electromagnetic spectrum. In this model a blast wave propelled by an ultrarelativistic jet ploughs through an external medium. Electrons accelerated by the shock emit synchrotron radiation in the magnetic field, which is also amplified by the shock (Rees & Meszaros 1992; Meszaros & Rees 1997). Radio observations are crucial for pinning down the evolution of the peak flux and two of the three characteristic frequencies of the broad-band synchrotron spectrum: the peak frequency and the synchrotron self-absorption frequency (with the third one being the cooling-break frequency). The radio emission from the blast wave typically peaks much later than at optical and X-ray wavelengths, on a time-scale of days to weeks depending on the observing frequency (Chandra & Frail 2012), and in some cases can be observed months or even years after the initial GRB explosion (e.g. Frail, Waxman & Kulkarni 2000a; van der Horst et al. 2008). Although radio afterglows are quite faint, with the majority having submillijansky peak fluxes (Chandra & Frail 2012), a large number of these have well-sampled light curves at time-scales from days to years (see Fig. 1).

There are a few GRBs for which the radio observations started within the first day after the gamma-ray trigger, and some resulted in the detection of an early peak (e.g. Kulkarni et al. 1999; Frail et al. 2000b). The favoured explanation for these early radio flares is emission from the ‘reverse shock’, a phenomenon arising from dynamics at the shock front. As the shell of shocked matter slows and the surrounding medium begins to influence the dynamics, the interaction region becomes bounded by two shock fronts: one propagating into the unshocked external medium (‘forward shock’) and a second, ‘reverse shock’ propagating backwards into the relativistic flow (as seen from the flow rest frame; see e.g. Kobayashi, Piran & Sari 1999). This model may be complicated further if there are pre-existing shock fronts in the external medium due to stellar winds (Pe’er & Wijers 2006). Emission from the reverse shock region may give rise to high flux levels across the frequency spectrum at much earlier times than the typical afterglow, and it has been predicted that radio flares due to the reverse shock are more likely to occur within the first hours after the GRB (Melandri et al. 2010).

Although a reverse shock should in principle be formed, the brightness may vary considerably, perhaps resulting in only a small fraction of GRBs displaying radio flares. However, this phenomenon is poorly constrained due to the scarcity of radio observations soon after the prompt gamma-ray emission. With the programme at AMI discussed in this paper we will be able to systematically probe this part of observational parameter space and look for radio flares caused by reverse shock emission. The detections of, or upper limits on, early peaks at radio frequencies are important for putting constraints on the physical parameters of the blast wave, for instance, its energy, the energy in electrons and the magnetic field. Furthermore, they can help solve an outstanding issue in GRB jet formation, namely whether the jet is Poynting flux dominated or a baryonic outflow, by constraining the energy emitted as the relativistic flow begins to interact with the external medium (Piran 2000). This would be most effective in combination with observations constraining very early optical flashes, which have been searched for systematically for some years now (see Roming et al. 2006, and references therein).

Besides the forward shock and reverse shock in GRB jets, there have been predictions of alternative sources of radio emission in

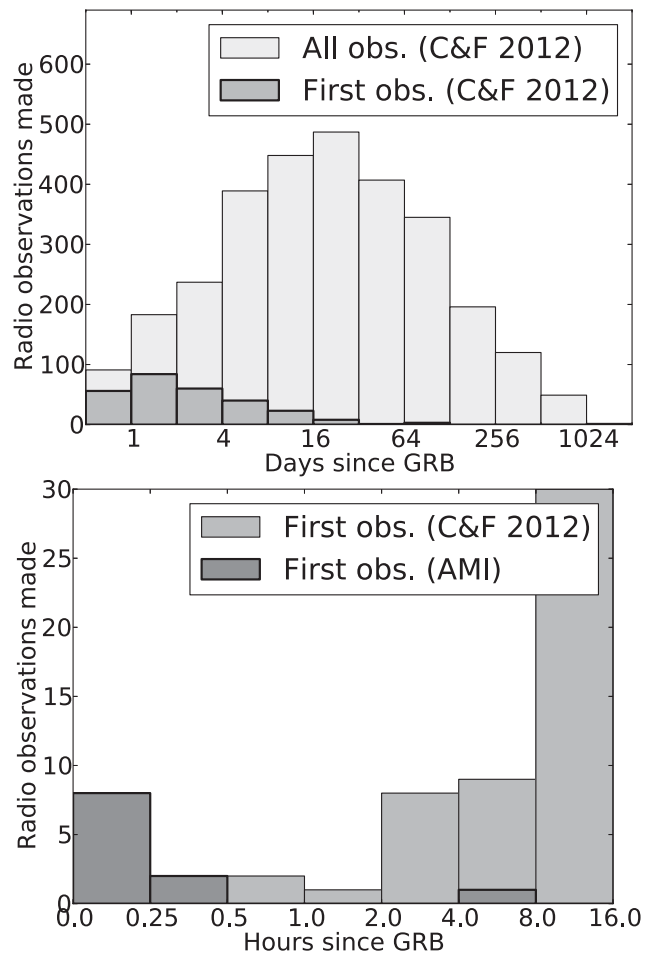


Figure 1. GRB radio follow-up delay histograms. Top: a histogram depicting radio observations of GRBs made with the VLA. Light grey bars show counts for all observations, while medium grey bars (overlaid) show counts for the first observation of any given source. Bottom: a close-up depicting the observation counts for initial radio observations of a GRB source undertaken less than 16 h after the burst. Medium grey bars depict VLA observations. Dark grey bars, overlaid, depict observations presented in this paper (there are no catalogued VLA observations within the 30-min period immediately after a burst). Data on VLA observations are reproduced from Chandra & Frail (2012).

the case of short GRBs (those with prompt gamma-ray emission durations of less than ~ 2 s; Kouveliotou et al. 1993). One favoured model for short GRBs is the merger of a binary system consisting of two neutron stars or a neutron star and a black hole. These mergers are thought to produce mildly to subrelativistic outflows, resulting in radio emission on time-scales of weeks to years (Nakar & Piran 2011; Piran, Nakar & Rosswog 2012). On shorter time-scales the predictions are more uncertain, one possible outcome being the generation of pulsar-like coherent emission by the rapidly rotating magnetar that is produced shortly after the binary merger (Pshirkov & Postnov 2010).

There may even be unpredicted mechanisms that produce early radio emission, and this follow-up programme will provide a systematic search for these at ~ 15 GHz frequencies. Intriguingly, recent observations at 1.4 GHz made with a 12 m single-dish antenna at the Parkes radio observatory have suggested the presence of short coherent bursts of radio emission within the first 20 min after the

GRB onset, hinting at the presence of just such radio phenomena occurring at early times (Bannister et al. 2012).

2.2 Sample sizes and selection bias

The sample of GRB radio frequency observations presented in Chandra & Frail (2012) – albeit large – is far from complete, being both strongly sensitivity limited and biased due to target selection effects. The observations that have been made, and are still being carried out, at different facilities [e.g. Jansky Very Large Array (JVLA), Westerbork Synthesis Radio Telescope, Australia Telescope Compact Array and Giant Metrewave Radio Telescope] are restricted to a subset of possible GRB triggers due to limited available observing time. Subsequently, only a small fraction of GRBs have been followed up, mostly those that have already shown interesting behaviour in the optical or X-ray bands. The same constraints on observing time also restrict systematic follow-up of GRB non-detections, with a radio non-detection in the first few days after the burst often resulting in no further radio follow-up, despite the fact that the peak flux of the radio emission will often occur at later times.

The recent upgrade of the JVLA allows for much deeper observations, which will help significantly in probing GRB radio afterglows at fainter flux levels. However, the number of GRBs that the JVLA and other facilities will observe is still limited, as illustrated by the fact that of the 11 GRBs observed and reported upon here, only five have been followed up by other radio telescopes (and the observations made known to the follow-up community). Although AMI is not as sensitive as the JVLA, this programme will provide a uniformly selected sample which can be used for comparison with GRB follow-up at other wavelengths. With systematic follow-up we should also detect many 15-GHz radio-afterglow peaks at later times when other observatories have stopped observing.

In summary, after the first year(s) of this project, we will have a uniquely large sample of early time radio observations of GRBs, providing important new constraints on short (minutes to hours) time-scale radio emission. We will also provide a systematically obtained and rapidly reduced data set, flagging up interesting GRBs for deeper observation both with AMI and other observatories as appropriate.

3 THE ARCMINUTE MICROKELVIN IMAGER LARGE ARRAY

The AMI-LA (Zwart et al. 2008) is a synthesis telescope composed of eight equatorially mounted 12.7 m dishes sited at the Mullard Radio Astronomy Observatory at Lord's Bridge, Cambridge. The telescope observes in the band 12.0–17.9 GHz with eight 0.75 GHz bandwidth channels. In practice, the two lowest frequency channels are not generally used due to a lower response in this frequency range and interference from geostationary satellites, which leaves an effective bandwidth of 13.5–17.9 GHz and a central frequency of 15.75 GHz using channels 3–8. After calibration, the phase is generally stable to 5° for channels 4–7 and 10° for channels 3 and 8. The telescope measures a single linear polarization ($I + Q$) and has a flux density sensitivity of $\simeq 3 \text{ mJy s}^{-1/2}$. Flux density calibration follows the Perley–Butler 2010 scale (R. Perley, private communication). AMI-LA data reduction has been described extensively in previous works (see e.g. Zwart et al. 2008) and the flux calibration is accurate to better than 5 per cent (Scaife et al. 2008; Hurley-Walker et al. 2009). During normal operation, the AMI-LA performs continuous execution of observations scheduled as entries

in a queue, with observing programmes often set up for several days in advance. Implementation of rapid response observing for this GRB study, as part of a new target-of-opportunity capability for AMI, involves automatic insertion of observing requests into the queue, if necessary displacing existing entries or an observation in progress.

4 AUTOMATED RESPONSE SYSTEM AND OBSERVING POLICY

The problem of automated transient detection and follow-up is multifaceted, with potential for detailed focus in many subfields such as detection, classification, resource allocation, etc. For the purposes of this paper we restrict our discussion to three broadly defined sub-systems: the distribution network, observation request triggering and the automated telescope response.

4.1 Transient alerts distribution via the VOEvent network

NASA uses the Gamma-ray Coordinates Network (GCN; Barthelmy et al. 1998) to rapidly distribute information about GRBs from satellites such as *INTEGRAL*, *Swift* and *Fermi*. Until recently, the most common medium was a short piece of plain text conforming to a specified format, distributed via a dedicated central server/client network. While fast and effective for this purpose, this approach will not encompass the range of hierarchical data structures that are envisioned for describing the wealth of transient events detected by a new generation of astronomical facilities such as LO-FAR, Pan-STARRS, *Gaia*, LSST, etc. A robust network of many facilities will also require a more flexible and adaptable distribution model, i.e. one that is largely decentralized.

The VOEvent standard (see footnote;¹ Williams et al. 2006) has been developed with these needs in mind, providing a flexible data format specified as an XML schema. The standard is transport independent (allowing for multiple distribution methods), allows for a distributed network of clients and servers, and has been previously employed by projects such as eStar, the Catalina Real Time Survey and SkyAlert (Allan, Naylor & Saunders 2006; Drake et al. 2009; Williams et al. 2009). GCN has now started distributing VOEvents (S. D. Barthelmy, private communication) using the VOEvent transport protocol,² which is an interoperable standard for VOEvent distribution.

We tested two freely available software packages implementing the VOEvent transfer protocol: ‘DAKOTA’³ and ‘COMET’.⁴ Both provide means of generating, distributing and subscribing to streams of VOEvents. For this work, we use COMET, primarily because it is implemented in the PYTHON programming language. The packages are dedicated to efficient and robust networking; processing of the information enclosed in the packets is delegated to an external program of the user's choice. In the current system implementation we monitor for new GRB events via a direct connection to the GCN VOEvents service, and pipe received VOEvents to a custom software package described below.

¹ ‘Sky Event Reporting Metadata Version 2.0’, Rob Seaman et al., 2011: <http://www.ivoa.net/Documents/VOEvent/index.html>

² ‘VOEvent Transport Protocol’, A. Allan and R. B. Denny: <http://www.ivoa.net/Documents/Notes/VOEventTransport/>

³ ‘DAKOTA’, R. B. Denny: <http://voevent.dc3.com>

⁴ ‘COMET’, J. D. Swinbank: <http://comet.transientskp.org>

4.2 VOEvent monitoring and trigger generation

The VOEvent identification and response routines were implemented using a PYTHON package written to facilitate such tasks: ‘PYSOVO’, which is available under an open-source licence from <https://github.com/timstaley/pysovo>. The package makes it easy for the user to extract information from the VOEvent XML packets, implement decision logic to determine whether an observation request will be triggered, determine whether the source is currently observable from a given observatory, generate requests formatted to custom observatory specifications and send notification e-mails and SMS (i.e. text to mobile phone) alerts to designated recipients. The package is designed to be easily extensible, in order to allow interaction with multiple sources of transient astronomical alerts and multiple follow-up facilities.

For this programme, we perform VOEvent monitoring and trigger generation using instances of COMET and PYSOVO running on a dedicated machine in Southampton. Each VOEvent has an ID string, referred to as an ‘International Virtual Observatory Resource Name (IVORN)’, which uniquely identifies a VOEvent in much the same way that a URL identifies a resource on the internet. We use a simple filter based on the IVORN to identify the VOEvent information packets pertaining to new *Swift* GRB alerts. Once relevant VOEvents have been identified, the coordinates are extracted, checked against our observation policy and used to generate a specially formatted observation request e-mail which is then sent to an address at the Mullard Radio Astronomy Observatory, Cambridge.

4.3 Automation of the AMI telescope

At the telescope level, incoming request e-mails to a dedicated account are picked out by an e-mail filter based on the subject line and a template contained in the message body. The request e-mail is then passed to a custom program written to respond to AMI target-of-opportunity requests, ‘RQCHECK’, running on the AMI-LA control computer. The e-mail is also forwarded to observatory members, who can act as request moderators if necessary.

The ‘rqcheck’ process validates the e-mail, checking syntax and keyword values, and then proceeds to the following.

- (i) Check whether target is currently on-sky and not obscured by the Sun or Moon. Sets start and stop times for the next possible timing within the next 24 h, according to the timing specified in the request (ASAP, next transit or specified sidereal time).
- (ii) Find the nearest phase calibrator from Very Long Baseline Array (VLBA)⁵ and JVAS (Patnaik et al. 1992) catalogues.
- (iii) Check the availability of the telescope, e.g. in the case of priority observing or engineering work.
- (iv) Construct an entry for the observing queue. This will always be a standard pointing observation with interleaved calibrator, using the default integration times for either array.
- (v) Insert this entry in the queue, displacing any previously scheduled observations if necessary.
- (vi) Append the request to a log file.
- (vii) Notify the requester of the outcome by reply e-mail.

4.4 Trigger and follow-up policy

For observations with AMI we employ a naive trigger policy, whereby all *Swift* GRB alerts with a declination above -10° cause

an observation request. This simple strategy is optimal in this case as we have access to sufficient observing time to allow systematic follow-up of all *Swift* GRB triggers in the Northern hemisphere. We are employing a logarithmic follow-up schedule for all triggers, observing promptly after the initial burst, after 1–2 d, 3–4 d, one week, two weeks and finally at one month. Follow-up of detected sources will continue to be evaluated manually. Our standard integration time for each observation is currently 1 h.

5 DATA REDUCTION AND ANALYSIS

When undertaking systematic observations such as those entailed by a GRB follow-up programme, it is desirable to have a fully automated reduction and analysis pipeline. With this goal in mind we have developed a procedure requiring minimal human intervention, which we plan to fully automate in the near future.

5.1 Data transfer and image synthesis

After observations have been made, the raw data are transferred from Cambridge to Southampton for further analysis. The data are then reduced using a procedure scripted in PYTHON which performs the following tasks.

- (i) Identifies set of observations pertaining to each GRB.
- (ii) Runs the AMI software tool ‘reduce’ on each observation to perform flagging for interference, shadowing and hardware errors, apply gain calibrations and synthesize the frequency channels to produce calibrated visibility data in *uv-FITS* format.
- (iii) Extracts certain metadata which can be used to flag observations of poor quality (e.g. system temperature estimate and RFI flagging percentages).
- (iv) Generates images using the casapy ‘clean’ algorithm, including a ‘dirty’ map and point spread function (PSF) map, for visual inspection.

5.2 Image time series analysis

Once a time series of images has been synthesized, these are visually checked for data quality, artefacts and sources coincident with the known GRB position. We then attempt to extract a light curve by running custom source-fitting software developed for LOFAR data analysis (Spreeuw 2010). This allows for source fitting at a known location, in addition to ‘blind’ source finding and extraction above a given threshold. To ensure that marginal detections are not missed, we use the source-fitting software to attempt an extraction at the location given by the enhanced *Swift* X-Ray Telescope (XRT) reduction (Burrows et al. 2005); except where this is unavailable (e.g. in the case of GRB 120403A) and we resort to the *Swift*-BAT refined position. In the case of a detection, the code attempts to fit an elliptical Gaussian profile, and errors on the fit are calculated according to the formulae of Condon (1997). When no detection is made, we assign an upper limit of three times the image RMS level.

6 RESULTS

6.1 Response times

We can estimate system response times by inspecting time stamps on the GCN notices and the AMI observation request logs. The typical response time for a target which is immediately observable with AMI is approximately 5 min from *Swift*-BAT trigger time to AMI

⁵ NRAO VLBA catalogue: <http://www.vlba.nrao.edu/astro/calib/>.

Table 1. Delay between GRB trigger time stamp and the start time of our initial observation, for each GRB followed up. No radio sources at the position of the GRB were detected at the 3σ level in our preliminary analysis of these data, and we list corresponding upper limits for the flux density averaged over the 1-h observation period. Note that GRB 120326A has a *marginal* detection at 0.337 mJy in the initial observation, ~ 7 h after the trigger time stamp.

GRB ID	Hours since burst	3σ upper limit (mJy)
GRB 120305A	0.07	0.316
GRB 120308A	0.08	0.164
GRB 120311A	0.07	0.230
GRB 120320A	14.32	0.196
GRB 120324A	0.07	0.218
GRB 120326A	7.37	0.430
GRB 120403A	7.48	0.238
GRB 120404A	11.32	0.223
GRB 120422A	6.41	0.616
GRB 120514A	50.91	0.211
GRB 120521C	0.24	0.302

taking data, and we can break this down into contributing factors. First, *Swift* GRB positions are broadcast from GCN, with a typical time delay of 7–30 s after the trigger time (approximately 0.1 s of which is due to GCN–S. Barthelmy, private communication). These are received by our VOEvent node in Southampton, processed, and an e-mail is sent to the AMI control system. Round trip time from NASA-GCN to AMI is consistently 4–6 s, including the time required to authenticate with the e-mail server. Alerts distributed solely via VOEvent transfer protocol will be relayed with an even shorter delay. If the target is on the sky, AMI then begins to slew almost immediately. The maximum slewing speed of the telescope is approximately 15° per minute, which means that response times for immediate observation will be of the order of a few minutes; currently a 4-min period is allotted to ensure that the telescope is on source before taking data.

Immediate observations with AMI are not always possible – the telescope may be undergoing maintenance, experiencing unfavourable weather conditions or the source may simply be over the horizon. However, even initial observations delayed by a few hours are faster than the large majority of previously catalogued radio follow-up, as illustrated in Fig. 1.

6.2 Radio follow-up light curves

Delays between each GRB trigger and our initial follow-up observation, and corresponding 3σ upper limits on the flux density, are given in Table 1. We detect a radio afterglow with confidence level above 5σ for GRB 120326A, as shown in Fig. 2. The remaining 10 have light curves consistent with non-detection. Specifically, in this preliminary analysis we do not detect any radio emission from the five GRBs observed within the first six hours post-GRB trigger. This places upper limits of less than 1 mJy on the 15-GHz flux density averaged over the 1-h observation period, but only weakly constrains possible emissions on much shorter time-scales. More detailed analysis of the data and upper limits is planned for future work, as discussed in Section 7. A full listing of all observations is given in Appendix A. The BAT and BAT-XRT light curves shown in Fig. 2 were obtained from the automated burst analyser page at the UK *Swift* Science Data Centre website (Evans et al. 2010) using 4σ significance bins and plotting the observed flux density at 10 keV.

7 DISCUSSION AND FUTURE WORK

In summary, this paper presents the first results from our programme of GRB follow-up observations using AMI-LA, demonstrating the success of the fully automated rapid-response follow-up system we have developed. Future observations will provide the first large sample of early-time radio observations of GRBs, placing tight constraints on theoretical models, while our monitoring strategy will result in a sample of radio afterglows unbiased by selection effects.

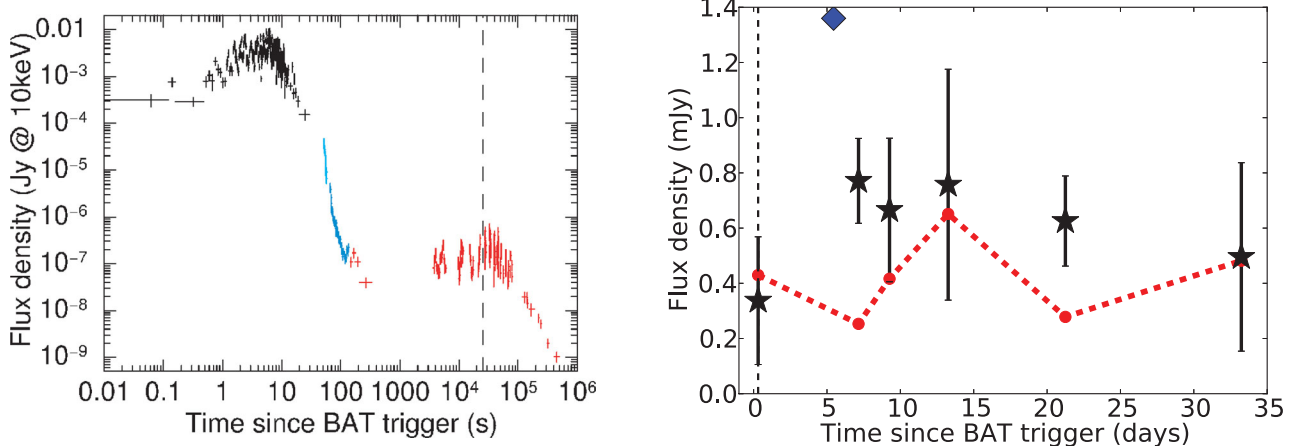


Figure 2. Light curves for GRB 120326A. Left: *Swift*-BAT and XRT measurements of the gamma- and X-ray emission. Star points with error bars represent 15-GHz observations made with AMI as part of this programme. Circle points represent 3σ blind detection thresholds – a connecting dashed line is overplotted to guide the eye. The diamond point at 5.45 d represents a JVLA detection at 22 GHz reported via GCN circulars (Laskar et al. 2012). Vertical dashed line in both plots represents the time at which AMI first started taking data on the GRB source position. Note that GRB 120326A was observed after a much longer delay than is possible with the automated response system (~ 7 h rather than the typical 5 min).

There are a number of enhancements to the AMI-LA response system we hope to implement in the near future. These include the following.

(i) Notifying the wider astronomical community when AMI-LA follow-up observations have been made, both through the NASA-GCN and the VOfEvent network.

(ii) Continued refinement of the data reduction process, including automated data transfer and initiation of the data reduction process. We will also implement time-slicing and analysis of the data on time-scales shorter than 1 h, to look for short-duration, high-intensity emission. This is a particularly exciting avenue of investigation given recent results reported in Bannister et al. (2012).

(iii) Automated discovery and user notification of GRB afterglow candidates, to enable rapid human intervention and detailed follow-up.

We note that a human-verified data set of modest size such as this provides an excellent test bed for automated transient detection and classification tools. We are working closely with the LOFAR Transients Key Project team on software to this end (Fender et al. 2006; Swinbank et al. 2011), and intend to employ it to automate much of the data curation currently done manually in the course of these follow-up observations.

This work was made possible in part through European Research Council (ERC) funding for the 4 Pi Sky project.⁶ In a wider context, this work provides a first step towards the longer term goals of that project, such as automated data reduction and distribution of *machine readable* data products (not just notifications) to the astronomical transients community, and use of such data products to enable rapid, automated decisions on allocation of telescope resources for transient follow-up.

It is hoped that the software described in this paper will be useful to the wider astronomical transients community in working towards these aims. We also plan to implement follow-up protocols in collaboration with other observatories. Notably, we are collaborating with the LOFAR-UK station (Fender 2010) which will provide extremely rapid response due to the way in which LOFAR stations are made up of phased arrays (and hence have no moving parts). We plan to start taking data with LOFAR-UK in the near future.

ACKNOWLEDGMENTS

This work makes use of GRB alerts supplied via NASA GCN, and of data supplied by the UK *Swift* Science Data Centre at the University of Leicester. We thank the staff of the Mullard Radio Astronomy Observatory for their assistance in the commissioning and operation of AMI, which is supported by Cambridge University and the STFC. We also thank Poonam Chandra for supplying the delay times of observations in the Chandra & Frail (2012) catalogue. This project was funded in part by European Research Council Advanced Grant 267697 ‘4 π sky: Extreme Astrophysics with Revolutionary Radio Telescopes’.

REFERENCES

- Akerlof C. W. et al., 2003, *PASP*, 115, 132
 Allan A., Naylor T., Saunders E. S., 2006, *Astron. Nachr.*, 327, 767
 Bannister K. W., Murphy T., Gaensler B. M., Reynolds J. E., 2012, *ApJ*, 757, 38

- Barthelmy S. D. et al., 1998, in Meegan C. A., ed., *AIP Conf. Ser.* 428, Fourth Huntsville Gamma-Ray Bursts Symposium. Am. Inst. Phys., New York, p. 99
 Barthelmy S. D. et al., 2005, *Space Sci. Rev.*, 120, 143
 Burrows D. N. et al., 2005, *Space Sci. Rev.*, 120, 165
 Chandra P., Frail D. A., 2012, *ApJ*, 746, 156
 Condon J. J., 1997, *PASP*, 109, 166
 Costa E. et al., 1997, *Nat*, 387, 783
 Dessenne C. A.-C. et al., 1996, *MNRAS*, 281, 977
 Drake A. J. et al., 2009, *ApJ*, 696, 870
 Evans P. A. et al., 2010, *A&A*, 519, 102
 Fender R. P., 2010, in Venturi T., ed., *Proceedings of Science 10th European VLBI Network Symposium and EVN Users Meeting: VLBI and the New Generation of Radio Arrays*, p. 49
 Fender R. P. et al., 2006, in Belloni T., ed., *VI Microquasar Workshop: Microquasars and Beyond*, Vol. MQW6, *Proceedings of Science*, p. 104
 Frail D. A., Kulkarni S. R., Nicastro L., Feroci M., Taylor G. B., 1997, *Nat*, 389, 261
 Frail D. A., Waxman E., Kulkarni S. R., 2000a, *ApJ*, 537, 191
 Frail D. A. et al., 2000b, *ApJ*, 538, L129
 Gehrels N. et al., 2004, *ApJ*, 611, 1005
 Green D. A. et al., 1995, *Ap&SS*, 231, 281
 Hurley-Walker N. et al., 2009, *MNRAS*, 396, 365
 Kobayashi S., Piran T., Sari R., 1999, *ApJ*, 513, 669
 Koranyi D. M., Green D. A., Warner P. J., Waldram E. M., Palmer D. M., 1995, *MNRAS*, 276, L13
 Kouveliotou C., Meegan C. A., Fishman G. J., Bhat N. P., Briggs M. S., Koshut T. M., Paciesas W. S., Pendleton G. N., 1993, *ApJ*, 413, L101
 Kulkarni S. R. et al., 1999, *ApJ*, 522, L97
 Laskar T., Zauderer A., Berger E., 2012, *GCN Circ.*, 13181
 Melandri A. et al., 2010, *ApJ*, 723, 1331
 Meszaros P., Rees M. J., 1997, *ApJ*, 476, 232
 Nakar E., Piran T., 2011, *Nat*, 478, 82
 Patnaik A. R., Browne I. W. A., Wilkinson P. N., Wrobel J. M., 1992, *MNRAS*, 254, 655
 Pe’er A., Wijers R. A. M. J., 2006, *ApJ*, 643, 1036
 Piran T., 2000, *Phys. Rep.*, 333, 529
 Piran T., Nakar E., Rosswog S., 2012, *The Electromagnetic Signals of Compact Binary Mergers*, preprint (arXiv:1204.6242P)
 Pshirkov M. S., Postnov K. A., 2010, *Ap&SS*, 330, 13
 Rees M. J., Meszaros P., 1992, *MNRAS*, 258, 41P
 Roming P. W. A. et al., 2006, *ApJ*, 652, 1416
 Sari R., Piran T., Narayan R., 1998, *ApJ*, 497, L17
 Scaife A. M. M. et al., 2008, *MNRAS*, 385, 809
 Spreeuw J., 2010, PhD thesis, Astronomical Institute Anton Pannekoek (IAP)
 Swinbank J., Accomazzi A., Mink D. J., Rots A. H., 2011, in Evans I. N., ed., *Astronomical Data Analysis Software and System XX*, Vol. 442. Astronomical Society of the Pacific, p. 313
 van der Horst A. J. et al., 2008, *A&A*, 480, 35
 van Paradijs J. et al., 1997, *Nat*, 386, 686
 Wijers R. A. M. J., Galama T. J., 1999, *ApJ*, 523, 177
 Williams R. D., Seaman R., Arviset C., Ponz D., Enrique S., 2006, in Gabriel C., ed., *Astronomical Data Analysis Software and System XV*, Vol. 351. Astronomical Society of the Pacific, p. 637
 Williams R. D., Djorgovski S. G., Drake A. J., Graham M. J., Mahabal A., Durand D., Dowler P., 2009, in Bohlender D. A., ed., *Astronomical Data Analysis Software and System XVIII*, Vol. 411. Astronomical Society of the Pacific, p. 115
 Wozniak P. R., Vestrand W. T., Wren J. A., White R. R., Evans S. M., Caspersen D., 2005, *ApJ*, 627, L13
 Zwart J. T. L. et al., 2008, *MNRAS*, 391, 1545

APPENDIX A: TABLE OF OBSERVATIONS

A full list of all AMI-LA 15-GHz observations presented in this paper is given in Table A1.

⁶ <http://www.4pisky.soton.ac.uk>

Table A1. 15-GHz measurements of the first 11 GRB sources observed with AMI-LA. For sources where radio emission was not detected (i.e. all except GRB 120326A), the flux density is given as an upper limit at the 3σ level.

Start date (UTC)	Days since burst	Flux density (mJy)	Image std. dev. (mJy)
GRB 120305A			
12/03/05 19:41:57	<0.01	<0.316	0.105
12/03/06 14:20:54	0.78	<0.158	0.053
12/03/07 15:26:46	1.83	<0.268	0.089
12/03/08 15:52:45	2.84	<0.171	0.057
12/03/09 18:08:27	3.94	<0.227	0.076
12/03/10 16:14:49	4.86	<0.193	0.064
12/03/11 17:50:36	5.93	<0.293	0.098
12/03/12 14:57:08	6.81	<0.272	0.091
12/03/13 17:17:49	7.90	<0.218	0.073
12/03/15 17:44:51	9.92	<0.226	0.075
12/03/16 18:00:52	10.93	<0.163	0.054
12/03/19 17:29:07	13.91	<0.203	0.068
12/03/22 13:08:01	16.73	<0.221	0.074
12/03/27 16:17:47	21.86	<0.197	0.066
12/04/01 16:18:04	26.86	<0.194	0.065
12/04/05 16:42:14	30.88	<0.169	0.056
GRB 120308A			
12/03/08 06:18:20	<0.01	<0.164	0.055
12/03/09 02:16:03	0.83	<0.144	0.048
12/03/13 03:20:06	4.88	<0.267	0.089
12/03/16 06:02:50	7.99	<0.183	0.061
12/03/19 05:11:09	10.96	<0.199	0.066
12/03/22 03:54:32	13.90	<0.217	0.072
12/03/25 03:27:46	16.88	<0.210	0.070
12/04/02 03:46:11	24.90	<0.169	0.056
12/04/11 02:25:55	33.84	<0.190	0.063
GRB 120311A			
12/03/11 05:37:37	<0.01	<0.230	0.077
12/04/05 04:39:13	24.96	<0.186	0.062
12/04/08 03:37:33	27.92	<0.180	0.060
GRB 120320A			
12/03/21 02:15:44	0.60	<0.196	0.065
12/04/05 01:24:45	15.56	<0.212	0.071
GRB 120324A			
12/03/24 06:03:17	<0.01	<0.218	0.073
12/04/05 05:49:01	11.99	<0.175	0.058
12/04/08 04:47:22	14.95	<0.240	0.080
12/04/16 08:25:13	23.10	<0.184	0.061
12/04/23 03:58:21	29.92	<0.329	0.110

Table A1. – *continued*

Start date (UTC)	Days since burst	Flux density (mJy)	Image std. dev. (mJy)
GRB 120326A			
12/03/26 08:42:58	0.31	0.337	0.143
12/04/02 04:55:59	7.15	0.771	0.084
12/04/04 07:37:40	9.26	0.666	0.139
12/04/08 07:31:55	13.26	0.757	0.217
12/04/16 07:15:25	21.25	0.626	0.093
12/04/28 07:03:08	33.24	0.495	0.161
GRB 120403A			
12/04/03 08:34:27	0.31	<0.238	0.079
12/04/04 16:06:16	1.63	<0.205	0.068
12/04/06 16:08:23	3.63	<0.204	0.068
12/04/10 14:52:49	7.57	<0.212	0.071
12/04/16 15:34:03	13.60	<0.254	0.085
12/05/01 14:50:02	28.57	<0.349	0.116
GRB 120404A			
12/04/05 00:09:57	0.47	<0.223	0.074
12/04/08 00:18:06	3.48	<0.237	0.079
12/04/11 00:11:17	6.47	<0.169	0.056
12/04/17 01:57:21	12.55	<0.214	0.071
GRB 120422A			
12/04/22 13:36:43	0.27	<0.616	0.205
12/04/24 16:17:23	2.38	<0.465	0.155
12/04/26 16:19:30	4.38	<0.368	0.123
12/04/29 17:17:31	7.42	<0.239	0.080
12/05/05 17:53:45	13.45	<0.228	0.076
12/05/20 21:19:03	28.59	<0.456	0.152
GRB 120514A			
12/05/16 04:07:39	2.12	<0.211	0.070
12/05/19 03:01:00	5.08	<0.460	0.153
12/05/22 01:54:22	8.03	<0.279	0.093
12/06/03 23:58:25	20.95	<0.619	0.206
12/06/11 03:20:16	28.09	<1.488	0.496
GRB 120521C			
12/05/21 23:36:44	0.01	<0.302	0.101
12/05/23 21:02:16	1.90	<0.793	0.264
12/05/25 20:54:25	3.90	<0.307	0.102
12/05/28 21:02:34	6.90	<0.328	0.109
12/06/04 18:30:23	13.80	<0.216	0.072
12/06/23 18:50:25	32.81	<0.342	0.114

This paper has been typeset from a \LaTeX file prepared by the author.