

# Near-infrared observations of Type Ia supernovae: the best known standard candle for cosmology

R. L. Barone-Nugent,<sup>1★</sup> C. Lidman,<sup>2,3★</sup> J. S. B. Wyithe,<sup>1★</sup> J. Mould,<sup>4★</sup> D. A. Howell,<sup>5,6</sup>  
I. M. Hook,<sup>7,8</sup> M. Sullivan,<sup>7★</sup> P. E. Nugent,<sup>9</sup> I. Arcavi,<sup>10</sup> S. B. Cenko,<sup>11</sup> J. Cooke,<sup>4</sup>  
A. Gal-Yam,<sup>10</sup> E. Y. Hsiao,<sup>12</sup> M. M. Kasliwal,<sup>12</sup> K. Maguire,<sup>7</sup> E. Ofek,<sup>10</sup> D.  
Poznanski<sup>13</sup> and D. Xu<sup>10</sup>

<sup>1</sup> School of Physics, University of Melbourne, Parkville, Victoria 3010, Australia

<sup>2</sup> Australian Astronomical Observatory, Epping, New South Wales 1710, Australia

<sup>3</sup> ARC Centre of Excellence for All-sky Astrophysics (CAASTRO)

<sup>4</sup> Swinburne University, Hawthorn, Victoria 3122, Australia

<sup>5</sup> Las Cumbres Observatory Global Telescope Network, 6740 Cortona Drive, Suite 102, Goleta, CA 93117, USA

<sup>6</sup> Department of Physics, University of California, Santa Barbara, Santa Barbara, CA 93106, USA

<sup>7</sup> Department of Physics (Astrophysics), University of Oxford, DWB, Keble Road, Oxford OX1 3RH

<sup>8</sup> INAF, Osservatorio Astronomico di Roma, via Frascati 33, 00040 Monteporzio (RM), Italy

<sup>9</sup> E.O. Lawrence Berkeley National Lab, 1 Cyclotron Road, Berkeley, CA 94720, USA

<sup>10</sup> Weizmann Institute of Science, Rehovot 76100, Israel

<sup>11</sup> Department of Astronomy, University of California, Berkeley, CA 94720-3411, USA

<sup>12</sup> Observatories of the Carnegie Institution for Science, 813 Santa Barbara Street, Pasadena, CA 91101, USA

<sup>13</sup> School of Physics and Astronomy, Tel Aviv University, Tel Aviv 69978, Israel

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## ABSTRACT

We present an analysis of the Hubble diagram for 12 normal Type Ia supernovae (SNe Ia) observed in the near-infrared (NIR)  $J$  and  $H$  bands. We select SNe exclusively from the redshift range  $0.03 < z < 0.09$  to reduce uncertainties coming from peculiar velocities while remaining in a cosmologically well-understood region. All of the SNe in our sample exhibit no spectral or  $B$ -band light-curve peculiarities and lie in the  $B$ -band stretch range of 0.8–1.15. Our results suggest that SNe Ia observed in the NIR are the best known standard candles. We fit previously determined NIR light-curve templates to new high-precision data to derive peak magnitudes and to determine the scatter about the Hubble line. Photometry of the 12 SNe is presented in the natural system. Using a standard cosmology of  $(H_0, \Omega_m, \Omega_\Lambda) = (70, 0.27, 0.73)$ , we find a median  $J$ -band absolute magnitude of  $M_J = -18.39$  with a scatter of  $\sigma_J = 0.116$  and a median  $H$ -band absolute magnitude of  $M_H = -18.36$  with a scatter of  $\sigma_H = 0.085$ . The scatter in the  $H$  band is the smallest yet measured. We search for correlations between residuals in the  $J$ - and  $H$ -band Hubble diagrams and SN properties, such as SN colour,  $B$ -band stretch and the projected distance from the centre of the host galaxy. The only significant correlation is between the  $J$ -band Hubble residual and the  $J - H$  pseudo-colour. We also examine how the scatter changes when fewer points in the NIR are used to constrain the light curve. With a single point in the  $H$  band taken anywhere from 10 d before to 15 d after  $B$ -band maximum light and a prior on the date of  $H$ -band maximum set from the date of  $B$ -band maximum, we find that we can measure distances to an accuracy of 6 per cent. The precision of SNe Ia in the NIR provides new opportunities for precision measurements of both the expansion history of the universe and peculiar velocities of nearby galaxies.

**Key words:** cosmology: observations – distance scale.

\*E-mail: robertbn@student.unimelb.edu.au (RLBN); clidman@aao.gov.au (CL); swyithe@unimelb.edu.au (JSBW); jmould@astro.swin.edu.au (JM); sullivan@astro.ox.ac.uk (MS)

## 1 INTRODUCTION

Type Ia supernovae (SNe Ia) are widely accepted to be excellent standardizable candles at optical wavelengths. The development of various empirical corrections, which can relate peak luminosity and light-curve shape (Phillips 1993; Hamuy et al. 1996), SN colour (Tripp 1998; Jha, Riess & Kirshner 2007), spectral information (Bronder et al. 2007; Walker et al. 2011; Silverman et al. 2012) and host-galaxy mass (Sullivan et al. 2010), allows SNe Ia peak magnitudes to serve as distance indicators with a corrected scatter<sup>1</sup> of as low as 0.13 mag (Conley et al. 2011; Silverman et al. 2012).

There has been growing evidence, though, that SNe Ia may be more accurate in the near-infrared (NIR) (for a recent review, see Phillips 2011). NIR light experiences less attenuation from dust than light at optical wavelengths, and theoretical models predict a smaller intrinsic dispersion in the NIR peak magnitude (Kasen 2006). Furthermore, there appears to be little or no relationship between light-curve shape and peak luminosity (Wood-Vasey et al. 2008; Folatelli et al. 2010; Mandel, Narayan & Kirshner 2011), meaning no empirical corrections need to be applied.

Recent studies (Meikle 2000; Krisciunas, Phillips & Suntzeff 2004; Wood-Vasey et al. 2008; Folatelli et al. 2010; Mandel et al. 2011; Kattner et al. 2012) have shown that SNe Ia in the NIR can be as reliable as corrected SNe Ia observed at optical wavelengths, but these studies have had some limitations. Foremost among these limitations is that the SNe were observed at distances that are not sufficiently large to place them in the Hubble flow (defined here as  $z > 0.03$ ) and hence are affected by peculiar velocities. Here we present a study of a sample of SNe Ia selected to lie exclusively in the redshift range  $0.03 < z < 0.09$  in order to minimize redshift uncertainty due to peculiar velocity. We use NIR images on 8-m class telescopes to observe SNe in our sample, resulting in photometry that is more precise than that achieved for the handful of SNe that have been observed beyond  $z = 0.03$ .

## 2 SNE IA SAMPLE

The SNe Ia used in our sample are summarized in Table 2. These SNe Ia were discovered by the Palomar Transient Factory (PTF; Law et al. 2009; Rau et al. 2009), spectroscopically confirmed, observed in the  $g$ ,  $r$ ,  $i$  bands (Maguire et al., in preparation) at the Liverpool Telescope and followed up in both the NIR  $J$  and  $H$  bands, each with four epochs of observation. PTF09dlc was observed using High Acuity Wide field K-band Imager (HAWK-I) (Casali et al. 2006) on the European Southern Observatory (ESO) 8.1-m Very Large Telescope (VLT), and all others were observed using NIRI,<sup>2</sup> the NIR Imager and Spectrometer (Hodapp et al. 2000) on Gemini Observatory's 8.2-m Gemini-North telescope. Reference images were taken  $\sim 1$  year after initial observations so that host-galaxy light could be removed.

The data were processed in a standard manner using IRAF<sup>3</sup> and our own scripts. Darks were used to remove the pedestal, and flats were used to remove pixel-to-pixel sensitivity variations. We used

the XDIMSUM package in IRAF to remove the sky and a modified version of ISIS2 (Alard 2000) to subtract the reference image of each SN from the images with SN light, leaving just light of the SN. Zero-points were derived from standards found in the Persson et al. (1998) catalogue. All magnitudes are reported in the natural system of each instrument.  $K$ -corrections were calculated using the revised spectral template of Hsiao et al. (2007), and corrections for galactic dust extinction were applied. We did not warp the spectral template to match the observed colour.

For the SNe PTF09dlc, PTF10hmv, PTF10mwb, PTF10nlg, PTF10tce, PTF10ufj, PTF10wnm and PTF10xyt, the first observing epoch was before NIR maximum. For the remaining SNe, the first observing epoch was near NIR maximum. When fitting light curves for the SNe, we included the date of  $B$ -band maximum as an extra prior. The  $B$ -band maximum was determined using SiFTO (Conley et al. 2008), as was the  $B$ -band stretch (Maguire et al., in preparation).

The  $B$ -band stretches range from 0.8 to 1.15, so there are no fast decliners in our sample. Nor are there any heavily reddened SNe. Furthermore, all of the SNe appear to be spectroscopically normal. The absence of fast decliners and heavily reddened SNe is a potential limitation of the current sample.

The photometry for all SNe is presented in Table 1.

## 3 TEMPLATE FITTING

To find the light-curve maxima, we fit the Flexible Light-curve InfraRed Template (FLIRT) presented in Mandel et al. (2009) to our  $K$ -corrected photometry<sup>4</sup>. As we did not observe all SNe Ia before maximum light, we fit the template to the data with an additional constraint. Concurrent observations of SNe Ia at NIR and optical wavelengths provide evidence that NIR maxima occur  $\sim 5$  d before  $B$ -band maximum (Meikle 2000; Krisciunas et al. 2004). We use this prior assumption as a constraint on the date of NIR maximum,  $t_p$ , with a  $1\sigma$  error of  $\sigma_p = 1$  d, imposed on the likelihood function. By fitting the light-curve template to the SNe Ia observed before maximum without a prior, we determine that the  $J$ -band maximum occurs 5.36 d before the  $B$ -band maximum with a standard deviation of 0.74 d, and the  $H$ -band maximum occurs 4.28 d before with a standard deviation of 0.70 d. Fig. 1 display the template-fitted light curve for PTF10tce.

The errors in the peak apparent magnitudes in both the  $J$  and  $H$  bands were calculated via 100 Monte Carlo simulations of template fitting to the data, including repeated Gaussian-distributed random samplings within the errors of each point.

To convert apparent magnitudes to absolute magnitudes, we determine the distance modulus for each SN, calculated using the standard flat cosmology of  $(H_0, \Omega_m, \Omega_\Lambda) = (70, 0.27, 0.73)$ . The redshift measurement for PTF10ufj is uncertain since the redshift is measured from the SN and not the host. In this case, we have assumed an uncertainty of  $\sigma_z = 0.005$  (Maguire et al., in preparation). For all other redshifts, the errors make negligible difference to the distance modulus.

## 4 RESULTS

The peak apparent magnitudes for each SN Ia in the  $J$  and  $H$  bands are quoted in Table 2. In the  $H$  band, the median absolute

<sup>1</sup> We use the rms about the mean as a measure of the scatter.

<sup>2</sup> <http://www.eso.org/sci/facilities/paranal/instruments/hawki/> and <http://www.gemini.edu/sciops/instruments/niri/>

<sup>3</sup> IRAF is distributed by the National Optical Astronomy Observatories which are operated by the Association of Universities for Research in Astronomy, Inc. under the cooperative agreement with the National Science Foundation.

<sup>4</sup> Photometry to be published in Barone-Nugent et al. (in preparation).

**Table 1.** The  $J$ - and  $H$ -band photometry for each epoch of each SN in our sample. Apparent magnitudes have been corrected for Galactic dust extinction but not  $K$ -corrected.

SN name	MJD	$m_J$	$m_H$
PTF09dlc	55068.14	$19.803 \pm 0.023$	$18.813 \pm 0.021$
	55073.11	$19.906 \pm 0.040$	$18.933 \pm 0.044$
	55085.16	$20.997 \pm 0.059$	$19.646 \pm 0.050$
	55093.18	$21.399 \pm 0.067$	$19.635 \pm 0.028$
PTF10hdv	55340.25	$19.084 \pm 0.044$	$18.516 \pm 0.052$
	55346.25	$19.407 \pm 0.057$	$18.766 \pm 0.052$
	55350.25	$19.787 \pm 0.085$	$18.920 \pm 0.065$
	55354.41	$20.266 \pm 0.139$	$19.048 \pm 0.085$
PTF10hmv	55341.26	–	$17.887 \pm 0.019$
	55346.28	–	$17.501 \pm 0.026$
	55351.30	–	$17.595 \pm 0.017$
	55357.25	–	$17.823 \pm 0.054$
	55358.25	–	$17.825 \pm 0.019$
PTF10mwb	55359.26	–	$17.826 \pm 0.018$
	55378.41	$19.066 \pm 0.031$	$18.429 \pm 0.030$
	55381.43	$18.677 \pm 0.027$	$17.765 \pm 0.014$
	55386.41	$17.683 \pm 0.026$	$17.484 \pm 0.012$
	55391.43	$18.202 \pm 0.022$	$17.569 \pm 0.013$
PTF10ndc	55396.39	$18.584 \pm 0.023$	$17.760 \pm 0.014$
	55386.43	$19.898 \pm 0.062$	$19.249 \pm 0.079$
	55389.38	$20.110 \pm 0.088$	$19.468 \pm 0.099$
	55393.39	$20.333 \pm 0.092$	$19.488 \pm 0.107$
	55395.35	$20.496 \pm 0.101$	$19.519 \pm 0.118$
PTF10nlg	55386.39	$19.352 \pm 0.044$	$18.694 \pm 0.039$
	55387.33	–	$18.500 \pm 0.041$
	55389.35	–	$18.571 \pm 0.035$
	55393.36	$19.679 \pm 0.073$	$18.770 \pm 0.044$
	55395.37	$19.816 \pm 0.059$	$18.791 \pm 0.057$
PTF10qyx	55422.63	$19.652 \pm 0.069$	$19.005 \pm 0.046$
	55424.56	$19.701 \pm 0.045$	$19.121 \pm 0.049$
	55428.64	$20.156 \pm 0.076$	$19.487 \pm 0.102$
	55433.57	$20.654 \pm 0.104$	$19.736 \pm 0.088$
	55435.41	$18.634 \pm 0.040$	$18.047 \pm 0.032$
PTF10tce	55438.34	$18.592 \pm 0.036$	$17.943 \pm 0.026$
	55443.50	$18.778 \pm 0.041$	$18.164 \pm 0.028$
	55448.53	$19.175 \pm 0.059$	$18.352 \pm 0.035$
	55447.55	$20.496 \pm 0.112$	$19.444 \pm 0.079$
	55450.64	$20.387 \pm 0.093$	$19.241 \pm 0.068$
PTF10ufj	55455.50	$20.312 \pm 0.100$	$19.202 \pm 0.061$
	55460.58	$20.996 \pm 0.238$	$19.544 \pm 0.097$
	55469.38	$19.513 \pm 0.055$	$18.952 \pm 0.044$
	55471.48	$19.443 \pm 0.052$	$18.822 \pm 0.040$
	55476.27	$19.482 \pm 0.053$	$18.984 \pm 0.069$
PTF10wnm	55481.28	$19.883 \pm 0.072$	$19.262 \pm 0.126$
	55470.33	$19.086 \pm 0.030$	$18.488 \pm 0.033$
	55473.29	$19.201 \pm 0.039$	$18.592 \pm 0.042$
	55478.24	$19.491 \pm 0.063$	$18.855 \pm 0.051$
	55483.29	$20.071 \pm 0.071$	$19.173 \pm 0.061$
PTF10wof	55486.26	$19.116 \pm 0.051$	$19.071 \pm 0.118$
	55492.24	$19.277 \pm 0.047$	$19.438 \pm 0.179$
	55496.21	–	$19.735 \pm 0.238$
	55498.20	–	$19.596 \pm 0.227$

magnitude of the SNe in our sample is  $M_H = -18.36 \pm 0.04$ . This compares reasonably well with the median magnitudes of SNe in Folatelli et al. (2010), Wood-Vasey et al. (2008) and Krisciunas et al. (2004) for which we derive  $M_H = -18.37 \pm 0.07$ ,  $-18.29 \pm 0.04$  and  $-18.43 \pm 0.06$ , respectively. The errors on the median are computed by jackknife resampling. In part, the difference in

magnitudes comes from the limited number of SNe in each of the samples. The differences may also be due to the differences in the photometric systems that each author adopts. These differences are generally smaller than 0.04 mag.

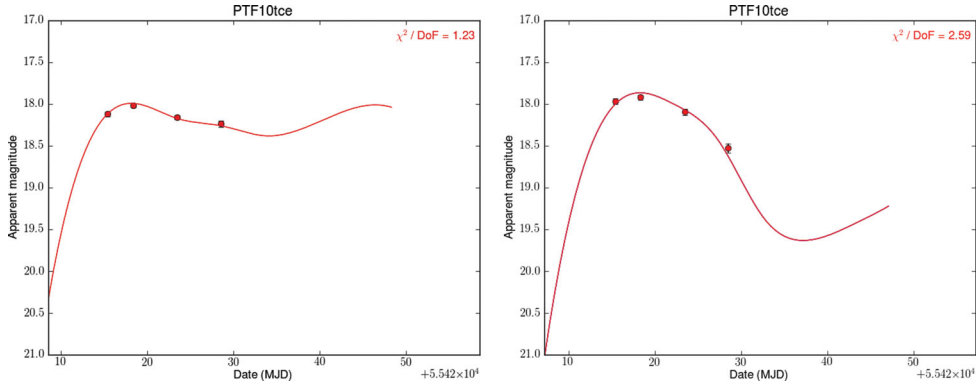
In the  $J$  band, the variations between samples are larger and more difficult to understand. We find a median magnitude of  $M_J = -18.39 \pm 0.06$ , compared to median magnitudes of  $M_J = -18.48 \pm 0.05$ ,  $-18.29 \pm 0.10$  and  $-18.73 \pm 0.05$  for the SNe in Folatelli et al. (2010), Wood-Vasey et al. (2008) and Krisciunas et al. (2004), respectively. The errors were calculated as for the  $H$  band. The largest source of uncertainty is due to sample size. Difference due to photometric systems are generally smaller. In order to constrain the dark energy equation of state beyond limits currently derived using SNe observed in the optical, the mean magnitude needs to be determined to a level of accuracy that is better than 0.01 mag.

The  $J$ - and  $H$ -band Hubble diagrams are shown in Fig. 2. Our results are compared with the surveys of Wood-Vasey et al. (2008) and Folatelli et al. (2010). Our observed scatter about the mean absolute magnitude of  $\sigma_H = 0.085 \pm 0.016$  mag and  $\sigma_J = 0.116 \pm 0.027$  mag represents a considerable improvement on previous samples, including Wood-Vasey et al. (2008) who observed  $\sigma_H = 0.28$  mag and  $\sigma_J = 0.41$  mag, Krisciunas et al. (2004) who observed  $\sigma_H = 0.17$  mag and  $\sigma_J = 0.13$  mag, Folatelli et al. (2010) who observed  $\sigma_H = 0.33$  mag and  $\sigma_J = 0.37$  mag and Mandel et al. (2009) who observed  $\sigma_H = 0.11$  mag and  $\sigma_J = 0.17$  mag. The values for  $\sigma_J$  and  $\sigma_H$  presented here are the scatter in the magnitude at  $J_{\max}$  and  $H_{\max}$ , respectively.

Photometric errors make little contribution to the scatter we measure. Peculiar velocities only make a significant contribution if they are large (i.e.  $\sim 300 \text{ km s}^{-1}$ ). The intrinsic scatters we measure, assuming peculiar velocities of  $300 \text{ km s}^{-1}$  are  $\sigma_{\text{int}} = 0.101$  in the  $J$  band and  $\sigma_{\text{int}} = 0.065$  in the  $H$  band. For peculiar velocities of  $150 \text{ km s}^{-1}$ , we obtain intrinsic scatters of  $\sigma_{\text{int}} = 0.109$  in the  $J$  band and  $\sigma_{\text{int}} = 0.077$  in the  $H$  band. Part of the scatter we observe may also be due to spectral variability affecting the  $K$ -corrections.

We note that the light curves of some of the SNe in our sample are not adequately described by the light-curve template that we have used. This is because there is a diversity in light-curve shapes, which may lead to opportunities to reduce the scatter about the Hubble line further. We investigated this by recomputing the scatter after removing SNe with poor light-curve fits. If we eliminate light curves with a  $\chi^2/\text{degrees of freedom}$  greater than 5 (a chance probability of  $Q = 0.002$ ), then we reject PTF09dlc and PTF10mwb in the  $J$  band and PTF09dlc, PTF10hmv and PTF10mwb in the  $H$  band. The resulting values of scatter in the  $J$  and  $H$  bands are 0.105 and 0.097, respectively. The scatter in the  $H$  band increases marginally but is noticeably lower for the  $J$  band. PTF09dlc, PTF10hmv and PTF10mwb exhibit no spectral or  $B$ -band light-curve peculiarities. However, applying such cuts for the purpose of reducing the scatter needs to be done with care since they may introduce biases that may affect the accuracy with which cosmological parameters can be measured.

The relationship between the peak magnitude in both filters and pseudo-colour ( $m_{J_{\max}} - m_{H_{\max}}$ ) was also considered. In the optical, there is a strong correlation between colour and luminosity (Tripp 1998; Jha et al. 2007). However, we find no obvious correlation between the  $H$ -band residual and  $J_{\max} - H_{\max}$  pseudo-colour in our sample. On the other hand, we do find a correlation between the  $J$ -band residual and pseudo-colour. From the least-squares residuals in the  $y$ -direction, we find a relation for the line of best fit, where the errors on the coefficients are  $1\sigma$  dispersions calculated via



**Figure 1.** The FLIRT template fitted to the light curves for PTF10tce in the *H* band (left) and the *J* band (right).

**Table 2.** The sample of SNe Ia showing their redshifts and *J*- and *H*-band peak apparent magnitudes. PTF10ndc, PTF10tce, PTF10ufj, PTF10wof and PTF10xyt were located by the citizen science Galaxy Zoo Supernovae project (Smith et al. 2011).

Supernova	RA (J2000)	Dec. (J2000)	$z_{\text{helio}}$	$z_{\text{CMB}}$	$m_{H_{\text{max}}}$	$m_{J_{\text{max}}}$
PTF09dlc	21:46:35.5	+06:23:23.5	0.0678	0.0666	$19.010^{+0.016}_{-0.018}$	$18.782^{+0.015}_{-0.020}$
PTF10hdv	12:07:43.4	+41:29:27.9	0.0534	0.0542	$18.521^{+0.031}_{-0.037}$	$18.453^{+0.026}_{-0.026}$
PTF10hmv	12:11:33.0	+47:16:29.8	0.0320	0.0327	$17.431^{+0.009}_{-0.008}$	—
PTF10mwb	17:17:50.0	+40:52:52.1	0.0313	0.0312	$17.371^{+0.007}_{-0.007}$	$17.213^{+0.010}_{-0.009}$
PTF10ndc	17:19:50.2	+28:31:57.5	0.0818	0.0817	$19.330^{+0.085}_{-0.133}$	$19.337^{+0.051}_{-0.043}$
PTF10nlg	16:50:34.5	+60:16:35.0	0.0560	0.0559	$18.634^{+0.018}_{-0.020}$	$18.720^{+0.032}_{-0.029}$
PTF10qyx	02:27:12.1	−04:31:04.8	0.0660	0.0652	$19.122^{+0.031}_{-0.027}$	$19.016^{+0.042}_{-0.053}$
PTF10tce	23:19:11.0	+09:11:54.2	0.0410	0.0398	$17.988^{+0.016}_{-0.011}$	$17.868^{+0.018}_{-0.021}$
PTF10ufj	02:25:39.1	+24:45:53.2	$0.077 \pm 0.005$	$0.0762 \pm 0.005$	$19.276^{+0.036}_{-0.047}$	$19.306^{+0.063}_{-0.051}$
PTF10wnm	0:13:47.26	+27:02:26.0	0.0656	0.0645	$18.941^{+0.028}_{-0.037}$	$18.764^{+0.022}_{-0.021}$
PTF10wof	23:32:41.8	+15:21:31.7	0.0526	0.0514	$18.567^{+0.025}_{-0.024}$	$18.458^{+0.022}_{-0.021}$
PTF10xyt	23:19:02.4	+13:47:26.8	0.0496	0.0484	$18.379^{+0.026}_{-0.023}$	$18.450^{+0.036}_{-0.037}$

1000 Monte Carlo simulations by randomly varying the points within their error bars:

$$J\text{-band residual} = (0.91 \pm 0.2) \times (M_{J_{\text{max}}} - M_{H_{\text{max}}}) + (0.05 \pm 0.02), \quad (1)$$

with a scatter about the line of best fit of  $\sigma = 0.071$ .

We also considered the relationship between the peak magnitude in both filters with the angular distance between the SN and the centre of the host galaxy and *B*-band stretch. However, we find no correlations between any of these quantities. It should be noted that the *B*-band stretch range covered by this sample is limited.

#### 4.1 Template fitting with a subset of points

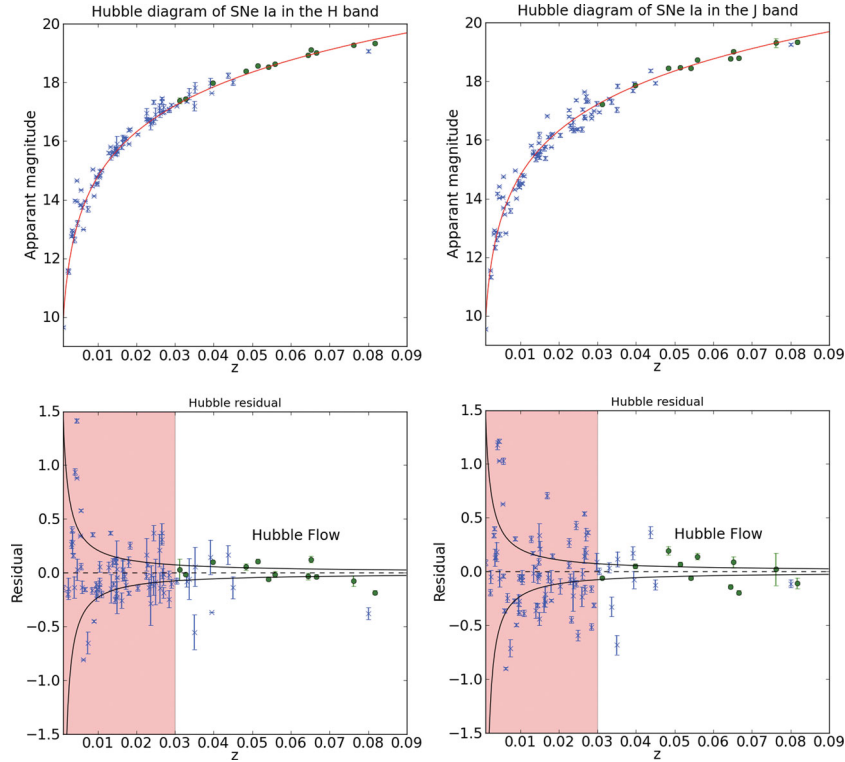
The previous discussion refers to SNe Ia light curves consisting of four data points, which we used to fit the light-curve templates. In this section we discuss the optimal number of epochs that should be observed in the NIR given finite telescope time.

We have already calculated the peak magnitude,  $M_{J,H}$ , for each SN Ia using all data available. We now calculate the peak magnitudes using a subset of the data points of size  $x$ , denoted by  $M_{x\text{-pts}}$ , and find the scatter of these magnitudes around  $M_{J,H}$ . The total scatter in the sample using  $x$  points is denoted by  $\sigma_{J,H,\text{tot}}$ . The results for these quantities based on our SN Ia sample are presented in Table 3.

We find in all cases that there is approximately 100th of a magnitude difference in the mean peak magnitude or less between different numbers of epochs used. However, there is additional scatter around the mean which decreases as more points are added. This illustrates the exchange between amount of telescope time used and dispersion measured. Table 3 suggests that the increased accuracy from three points to four is  $\sim 0.01$  mag or less in both the *J* and *H* bands. Furthermore, using just one point with a known date of maximum can be used while only sacrificing  $\sim 0.04$  mag of accuracy (i.e.  $1\sigma$  scatter) in both the *J* and *H* bands. Importantly, using just one point in the *J* band gives a dispersion that is similar to the best standardized SNe Ia observed in the optical, and using just one point in the *H* band can further improve the accuracy of SNe Ia observed in the optical. We also find that, within the range of epochs covered by our data (approximately 10 d before to 15 d after *B*-band maximum), the amount of additional scatter introduced due to using only one point has no dependence on when this point is taken with respect to *B*-band maximum. Importantly, single observations after *B*-band peak are equally as effective as points prior to the peak. We may now compare the standard error of the mean (SEM) for each subset of points. The SEM for a sample of  $n$  points is given by

$$\text{SEM} = \frac{\sigma}{\sqrt{n}}, \quad (2)$$





**Figure 2.** The upper panels show the Hubble diagrams of our sample of 12 SNe (green circles) in the  $H$  band (left) and the  $J$  band (right), including the Wood-Vasey et al. (2008) sample, the Folatelli et al. (2010) sample and the Krisciunas et al. (2004) sample (blue crosses) for comparison. The red line represents the apparent magnitude that would be observed assuming constant absolute magnitudes of  $-18.36$  ( $H$  band) and  $-18.39$  ( $J$  band). The lower panels show the Hubble residual, i.e. the deviation from the red line. The shaded red region,  $z < 0.03$ , is the region excluded in our sample due to the associated peculiar velocity errors. The solid black lines represent the change in the distance modulus due to a peculiar velocity of  $\pm 300 \text{ km s}^{-1}$ .

**Table 3.** The mean peak magnitude and dispersion in the  $J$  and  $H$  bands when using light curves with  $x$  points.

$x$	$\overline{M}_{J, x\text{-pts}} - M_J$	$\sigma_{J, \text{tot}}$	$\overline{M}_{H, x\text{-pts}} - M_H$	$\sigma_{H, \text{tot}}$
1	0.002	0.146	0.007	0.116
2	0.012	0.126	0.002	0.096
3	0.008	0.120	0.002	0.089
4	0.000	0.116	0.000	0.085

**Table 4.** The SEM for an example of 12 observing nights distributed among three, four, six and 12 SNe.

$x$	Sample size	$\sigma_H$	$\text{SEM}_H$	$\sigma_J$	$\text{SEM}_J$
1	12	0.116	0.033	0.146	0.042
2	6	0.096	0.039	0.126	0.051
3	4	0.089	0.045	0.120	0.060
4	3	0.085	0.049	0.116	0.067

where  $\sigma$  is the standard deviation of the sample. As an illustrative example, we compare observing three SNe over four epochs, four SNe over three epochs, six SNe over two epochs and 12 SNe for a single epoch. In all cases, we assume that we have well-sampled optical light curves for each SN Ia. Each of these would require equal allocation of telescope time. The relative SEM for each method is summarized in Table 4. We find that observing four times the number of SNe with only a single epoch delivers a lower SEM than observing fewer SNe over more than one epoch. The larger sample

size more than compensates for the higher uncertainty in the peak magnitude.

## 5 CONCLUSION

We have presented NIR light curves of 12 SNe Ia that are in the Hubble flow. We find that the intrinsic scatter in peak luminosities of SNe Ia in the NIR  $J$  and  $H$  bands are smaller than previously thought. This is the first sample to display an  $H$ -band rms scatter as small as  $\sigma_H = 0.085 \pm 0.016$  (with a median peak magnitude of  $M_H = -18.36$ ). Our observed  $J$ -band rms scatter of  $\sigma_J = 0.116 \pm 0.027$  (with a median peak magnitude of  $M_J = -18.39$ ) is smaller than reported elsewhere. These results provide distance errors of  $\sim 4$  per cent using  $H$ -band SNe, making them the most precise standard candles for cosmology.

We have also shown that if concurrent optical observations are made, we may use a predicted date of NIR maximum as a constraint when fitting the light curve. With this constraint, we may use as few as one NIR observation within  $\sim 5$ – $10$  d of NIR maximum per SN while still achieving scatters of  $\sigma_H = 0.116$  and  $\sigma_J = 0.146$ . As surveys improve over the coming years and more SNe Ia are discovered, single-night NIR observations undertaken concurrently with optical observations will be the most efficient and accurate way to construct samples.

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