

DELFT UNIVERSITY OF TECHNOLOGY

# LASER SWARM

FINAL REPORT

DESIGN SYNTHESIS EXERCISE

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# Preface

This final document contains our evaluation of the Laser Swarm concept. Ten weeks have passed and we are finally ready to present our conclusions.

We did not, however, do everything by ourselves, and this is a good place to thank those who have been so kind as to give us some of their precious time.

First I want to mention professor Edoardo Charbon, who kindly explained and showed to us the workings of his **SPAD!** (**SPAD!**)<sup>1</sup>: without his help we would be hopelessly in the dark on these fascinating devices. From this place I also want to thank professor Kourosh Khoshelham for helping us to some valuable material on remote sensing. I also want to thank Jasper Bouwmeester, who provided us with some helpful information on solar panel pricing.

And finally I thank our tutor, professor Ben Gorte, who was very generous to us with his time and has been very supportive of the entire project.

I hope you will have as much fun reading this report, as we had in making it!

Yours sincerely,

Co Florijn,  
Chairman

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<sup>1</sup>See [?, 98-105], [?] and [?]

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## List of Symbols

Symbol	Description	SI units
A	accuracy	rad
A	area	$m^2$
a	semimajor axis	km
$C_D$	drag coefficient	—
D	air drag	N
e	eccentricity	—
$\dot{e}$	time rate of change of the eccentricity	$s^{-1}$
g	gravitational acceleration	$m/s^2$
$I_{rt}$	round-trip interactivity loss	Watt
$I_{sp}$	specific impulse	s
i	inclination	rad
$i_R$	relative inclination	rad
$\dot{i}$	time rate of change of the inclination	$rad/s$
$J_{ii}$	inertia tensor for i	$m^4$
$J_2$	second zonal coefficient	0.001082645 —
$J_3$	third zonal coefficient	0.000002546 —
$M_C$	torque about c	N/m
m	mass	kg
$m_f$	fuel mass	kg
$m_p$	payload mass	kg
$m_0$	dry mass	kg
N	angular separation	rad
n	nr of energy levels	—
n	mean motion	$7.292 \cdot 10^{-5} \text{ rad/s}$
$n_e$	refractive index	—
P	power	Watt
$p_{p,th}$	threshold power	Watt
R	mass ratio	—
$R_C$	radius of the orbit of the center of mass	km
$R_E$	earth equitorial radius	6378 km
s	variable in the Laplace equation	—
T	temperature	K
$T_i$	period of i	s
V	velocity	$m/s$
$V_a$	voltage at a	V
$V_B$	voltage at B	V
W	weight	N
$\eta$	pump efficiency	—
$\theta_i$	angle w.r.t. rotation axis i	rad
$\theta_{man}$	manoeuvre angle	rad
$\lambda$	wavelength	m
$\mu$	gravitational constant of the Earth	$398600 \cdot 10^9 \text{ m}^3/\text{s}^2$
$\nu$	true anomaly	rad
$\rho$	density	$kg/m^3$
s	stimulated emission cross-section	$m^2$
t	upper level lifetime	s
$\phi$	phase angle	rad
$\phi_R$	relative phase angle	

$\Omega$	right ascension of the ascending node	rad
$\dot{\Omega}_{J_2}$	rate of precession	rad/s
$\omega$	argument of perigee	rad
$\dot{\omega}$	time rate of change of the argument of perigee	rad/s

## List of Acronyms

**SPAD** Single Photon Avalanche Diode

## Abstract

In February 2010 the ICESat mission ended after 7 years of measuring ice sheet mass balance, cloud and aerosol heights, as well as land topography and vegetation characteristics using a space based **LiDAR!** (**LiDAR!**) system. ICESat used only one of the possible approaches for **LiDAR!**, namely the use of a high energy laser and a large receiver telescope. The other approach is using a high frequency, low energy laser and a single photon detector. The advantage of the latter approach is that it has a much lower mass, but it is uncertain if even a single photon per pulse reaches the receiver. One possible solution could be to use a swarm of satellites around the emitter, each equipped with a single photon detector. However, the technical feasibility of this concept has not yet been proven.

This report briefly highlights different ways in which such a mission could be accomplished, and then decides what way will be the most likely to succeed.

# Chapter 1

## Introduction

In February 2010 the ICESat mission ended after 7 years of measuring ice sheet mass balance, cloud and aerosol heights, as well as land topography and vegetation characteristics. To do all this, ICESat had only one instrument on board: a space based **LiDAR!** system (**GLAS!** (**GLAS!**)), allowing for an unprecedented 3D view of the Earth's surface and atmosphere. The laser lifetimes, however, were severely limited because of manufacturing errors in one of the laser components.

ICESat followed only one of the possible approaches for **LiDAR!**, namely the use of a high energy laser and a large receiver telescope. The other approach is using a high frequency, low energy laser and a single photon detector. The advantage of the latter approach is that it has a much lower mass, but it is uncertain if even a single photon per pulse reaches the receiver. One possible solution could be the use of a swarm of satellites around the emitter, each equipped with a single photon detector. However the technical feasibility of this concept has not yet been proven.

In this mid term review some options for this concept are proposed and traded off so as to find the best fit for the mission. This tradeoff is what the bulk of this report is about. Also described are the workings and results from the simulator, as well as the project and risk management. Chapters on operation and logistics, stability and control and sustainability are included too.

# Chapter 2

# Project Management

## 2.1 Human Resource Allocation

The first part of the project will involve a team of five specialists designing the five critical subsystems and one person designing the orbits of the swarm. The other four members will concentrate on the development of the software that will be used to assist the trade-off and verify the design. At later stages, some of the the software engineering personnel is heavily involved in designing proper algorithms for the processing of mission data, while others will be brought in to assist with detail design.

A schematic representation of the resource allocation can be found in figure ?? on page ??.

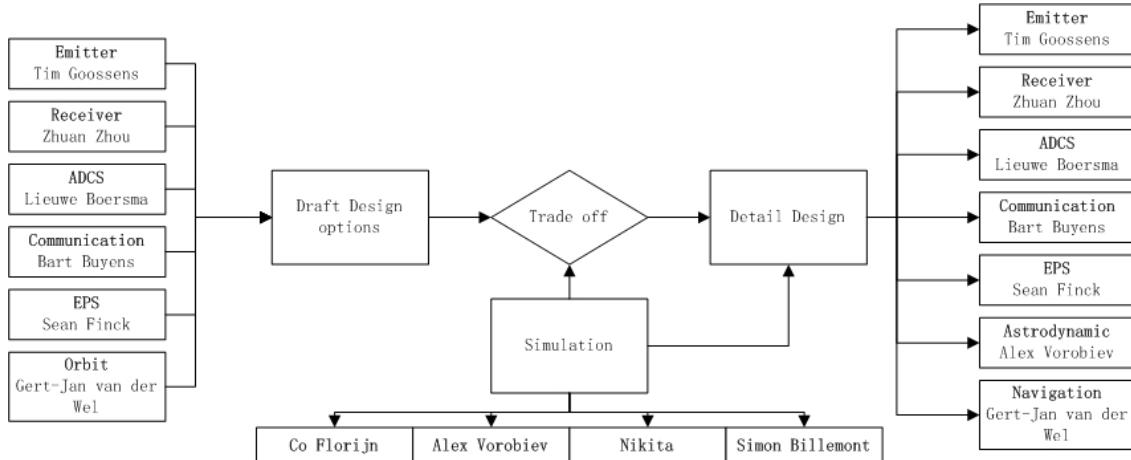


Figure 2.1: Human resource allocation chart.

The table ?? on page ?? gives the documentation distribution of each person on each chapter.

## 2.2 Mass Budget Breakdown

The table ?? on page ?? indicates mass budget breakdown for emitter and receiver satellites. The table is mainly divided into two parts. The first part gives each subsystem mass of emitter and receiver satellites in both kilograms and percentage of total dry mass. Meanwhile, the second part includes the mass of propellants and then also gives the total mass when the satellites are in their orbit. The deviation of all numbers in the table can be found in each corresponding chapter of the subsystem.

## 2.3 Cost Budget Breakdown

Figure ?? on page ?? shows a typical life-cycle for a space mission. The **RDTE!** (**RDTE!**) stage includes the planning, development and testing of all prototypes and qualification units, but does not include the technology development for different subsystems. In the case of the laser swarm this is largely dependent on the single emitter and one receiver unit. This stage is also mostly consistent of non-recurring costs. The production stage consists of actual manufacture of the physical satellites. The cost estimation in this stage is based on the **TFU!** (**TFU!**). This is done because it is assumed that the first unit (in the case of the Laser Swarm, that would be the emitter and one receiver) would be the most expensive to produce. The rest of the swarm constellation satellite costs are calculated by taking a theoretical learning curve [?].

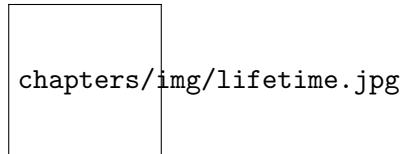


Figure 2.2: Satellite Life-Cycle

The table ?? on page ?? gives the cost breakdown of both theoretical first unit cost for either emitter or receiver satellite and also the swarm of 9 satellites including all the wraps.

Chapter	Documentation	Author	Checked by
-	Title Page	Alex	
-	Preface	Co, Team	
-	List of Symbol		
-	List of Acronyms		
-	Abstract	Co, Bart	
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2	Project Management		
2.1	-Resource Allocation	Zhuan	
2.2	-Budget Breakdown	Zhuan	
2.3	-Operations and Logistic Concept Description		
2.4	-Project Design and Development Logic	GJ	
2.5	-Project Gantt Chart	Lieuwe	
3	Mission Approach		
3.1	-Function Flow Diagram	Lieuwe, Zhuan	
3.2	-Function Breakdown Structure	Lieuwe, Zhuan	
3.3	-H/W Block Diagram	Lieuwe,	
3.4	-S/W Block Diagram		
3.5	-Electrical Block Diagram		
3.6	-Data Handling Block Diagram		
4	Risk Management	Tim, Zhuan	
5	Launch and Astrodynamic Characteristics	Alex	
5.1	-Launch Segment	Alex	
5.2	-Space Segment	Alex	
5.3	-Space Environment and Shielding	Alex	
6	Emitter		
6.1	-OEP	Tim	
6.1.1	-Principle of Diode Laser	Tim	
6.1.2	-Diode Pumped Laser Configuration	Tim	
6.1.3	-Optical Characteristics	Tim	
6.1.4	-Diffraction	Tim	
6.1.5	-Thermal Control	Tim, Lieuwe	
6.1.6	-Laser Life Time Expectancy	Tim	
6.1.7	-Laser Focus Calculation	Zhuan	
6.2	-Navigation	GJ	
6.3	-Communication	Bart	
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7.1	-ORP	Zhuan	
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7.1.5	-Payload Cost Estimation	Tim, Zhuan	
7.1.6	-Navigation	GJ	
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7.4	-EPS	Sean	
7.5	-Summary		
8	Data Validation	Simon, Co, Nikita	
8.1	-Software Tool Internals	Simon, Co, Nikita	
8.2	-Validation Results	Simon, Co, Nikita	
9	Sustainable Development Strategy		
10	Compliance Matrix		
11	Conclusion	Recommendations	
-	Others		
-	-Catia Drawing	Lieuwe	
-	-Latex Compile	Alex	

Table 2.1: Report writing distribution.

Subsystem	Emitter	Receiver		
	$M[kg]$	$\%M_{dry}$	$M[kg]$	$\%M_{dry}$
Communication	10.66	21	3	22.2
Navigation	0.25	0.5	0.25	1.85
OEP	15	29.8	-	-
ORP	0.22	0.4	0.22	1.63
EPS	5.8	11.5	3.6	26.6
ADCS	2	4	2	14.8
Thermal	1.48	3	0.3	2.22
Structure	12.35	24.5	2.45	18.12
Propulsion(tank)	1	2	0.75	5.55
Thruster	0.65	1.3	0.15	1.11
Shielding	1	2	0.8	5.92
$M_{dry}$	50.41	100	13.52	100
$M_{propellant}$	4	-	1.5	-
$M_{Loaded}$	54.41	-	15.02	-
$M_{Orbit}$	53.1	-	14.92	-

Table 2.2: Mass Budget Breakdown of the emitter and a representative receiver satellites.

Subsystem	Unit Cost				Swarm (9 satellites)	
	Emitter satellite		receiver satellite			
	Cost [k\$]	%Total	Cost [k\$]	%Total	Cost [k\$]	%Total
Payload	8215.96	21.1	4048.74	21.1	8660	21.1
<b>Bus Total</b>						
Structure	?	9.7	?	9.7	?	9.7
Thermal	?	1.1	?	1.1	?	1.1
EPS	?	?	?	?	?	?
Navigation	25	?	25	?	191.235	?
Communication	2940	?	612.5	?	7141.14	?
<b>ADCS!</b>	199.093	?	175.914	?	1405.687	?
Tank	0.713	?	0.428	?	3.649	?
Thruster	570.64	?	356.65	?	3016.9	?
<b>Wraps</b>						
<b>IAT!</b>	?	7.3	?	7.3	?	7.3
Program Level	?	12.1	?	12.1	?	12.1
<b>GSE!</b>	?	3.5	?	3.5	?	3.5
Launch	-	-	-	-	18534.25	-
Total	?	100	?	100	?	100
Total (FY00)	-	-	-	-	?	-

Table 2.3: Cost Budget Breakdown.

# Chapter 3

## Mission Approach

### 3.1 Functional Flow Diagram

The **FFD!** (**FFD!**) shows the functions the system needs to perform during its mission life. The schematic representation is divided into 2 parts. Part 1 detailed top level functions F1 and F2 (Figure ?? on page ??) meanwhile part 2 detailed top level functions F3, F4 and F5 (Figure ?? on page ??).

The first thing that needs to happen, after having been built, is that the satellites are put into their orbits and pointed towards earth. After that the measurements can start: the emitter sends down laser pulses and notifies the receivers that the signals are sent. The receivers can adjust their attitude, pick up reflected photons, turn them into a digital signal and inform the computer, which puts the data in a buffer. The data of the receivers will be sent to the emitter satellite continually and then it will transfer the data package to ground when the emitter is passing the ground station. The data on the ground can be split up into data packages, which can be distributed to research institutes and other interested parties. With those data sets, a terrain model and **BRDF!** (**BRDF!**) can be recreated. At the **EOL!** (**EOL!**) of the mission the satellites are decommissioned to make room for other satellites.

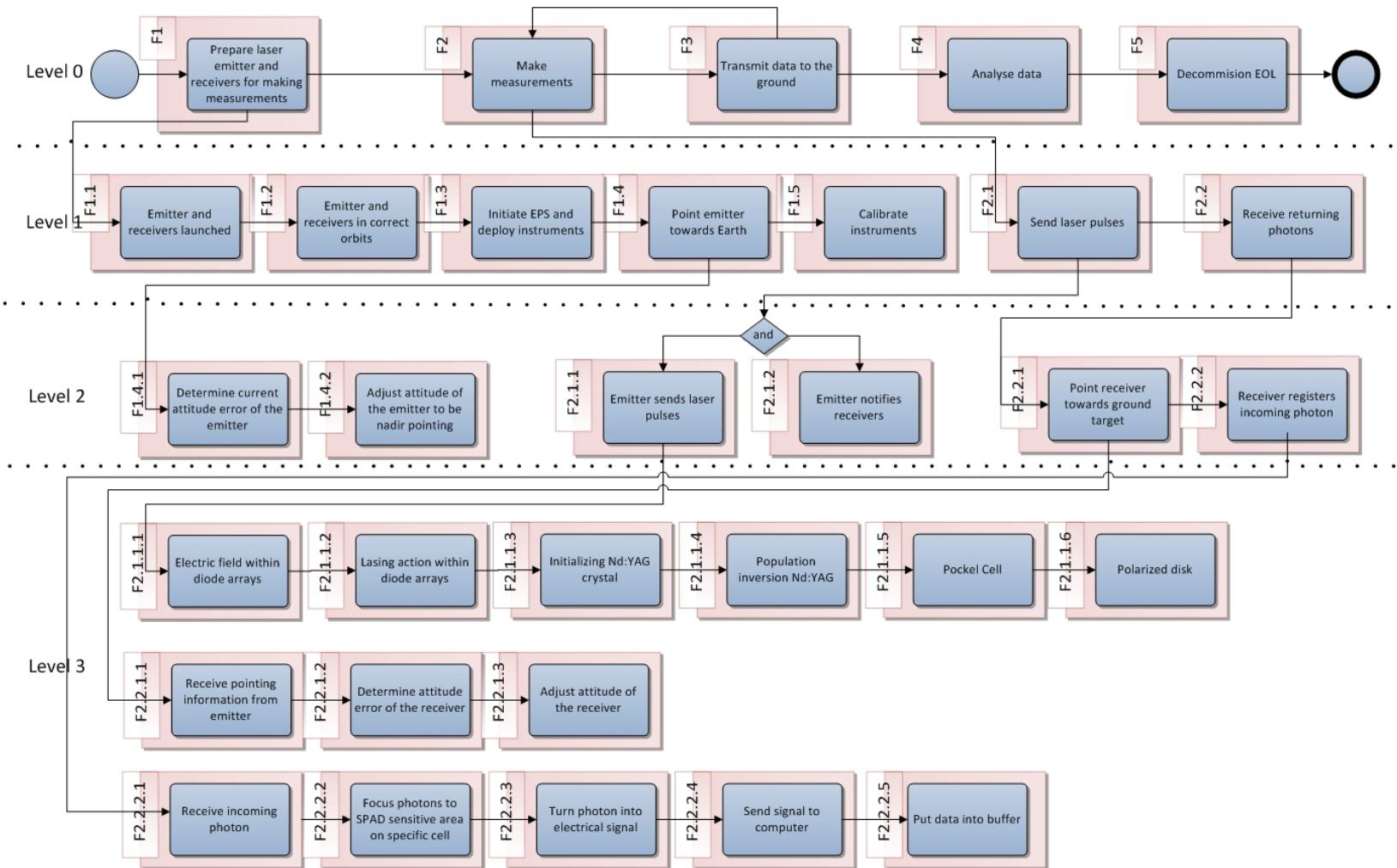


Figure 3.1: Functional Flow Diagram part 1(F1 and F2 detail) of the Laser Swarm mission

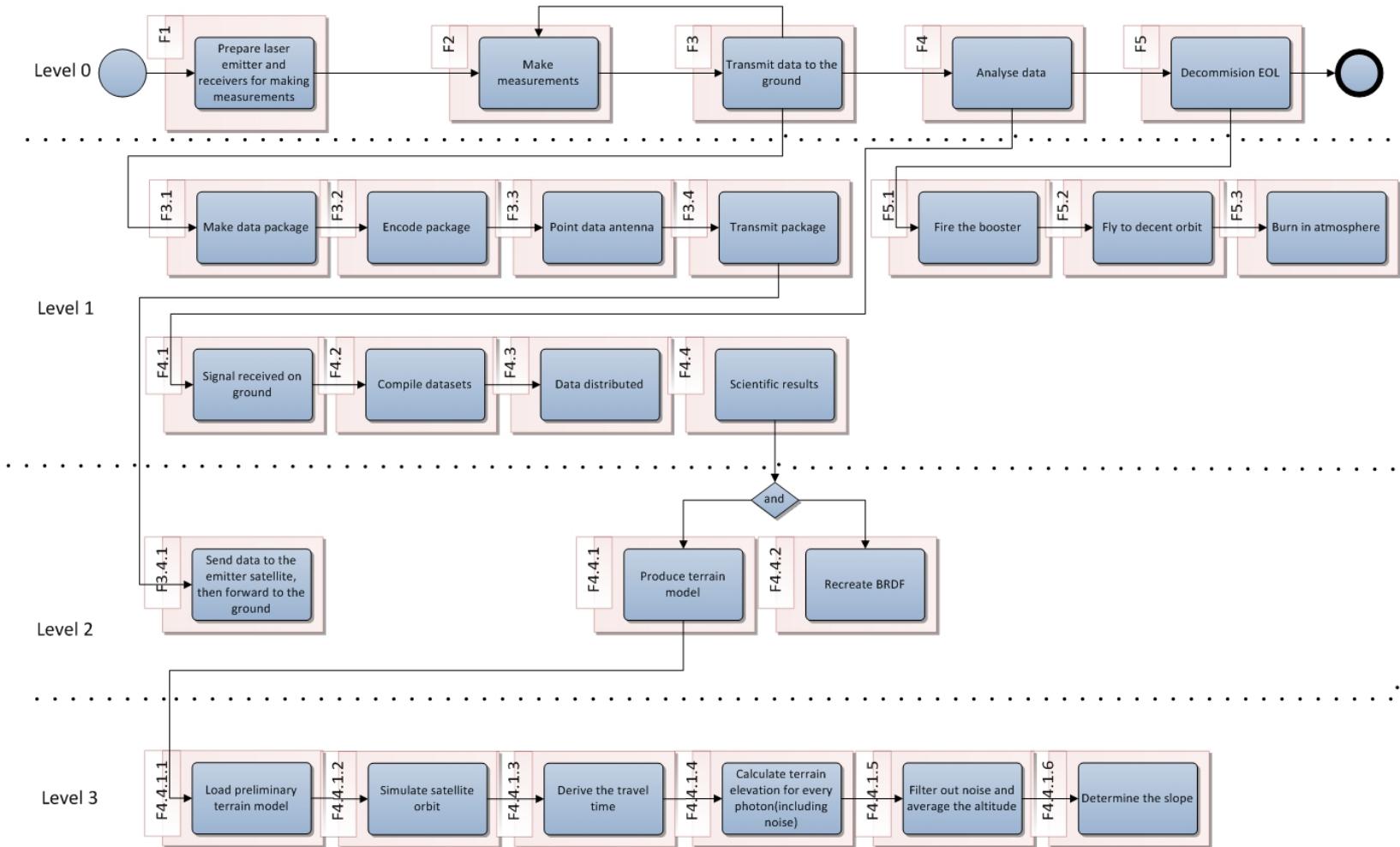


Figure 3.2: Functional Flow Diagram part 2(F3, F4 and F5 detail) of the Laser Swarm mission

## 3.2 Functional Breakdown Structure

The **FBS!** (**FBS!**) shows the functions the system needs to perform, broken up in subtasks from other functions. The schematic representation can be found in figure ?? on page ??.

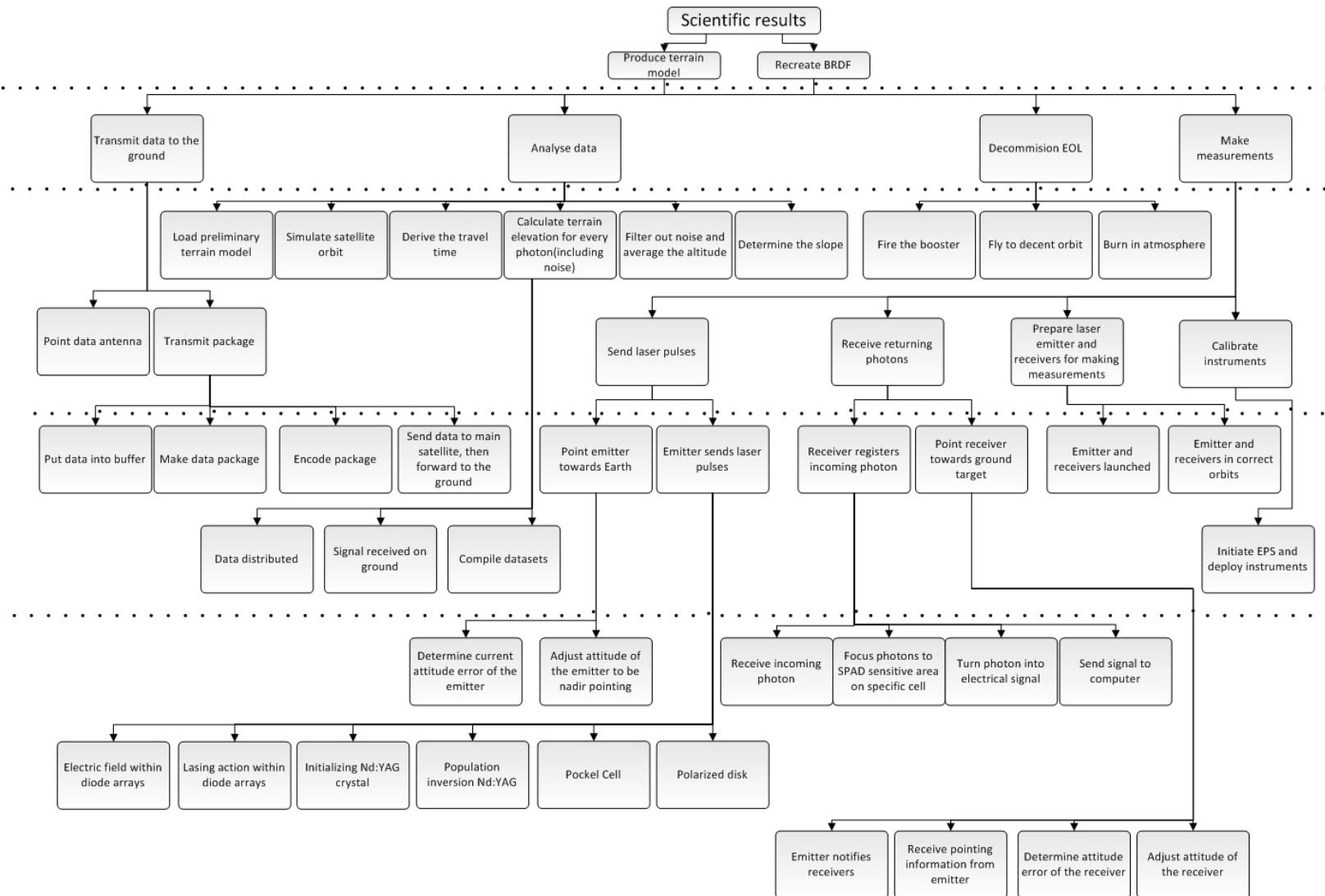


Figure 3.3: Functional breakdown structure of the Laser Swarm mission

The main function for the system is to be able to get scientific results as producing terrain model and recreating **BRDF!**, which is the reason for making the measurements. To be able to produce the terrain model the measurements need to be made, the data needs to be transferred to the ground and the data needs to be analyzed. Because the mission has to be sustainable, the satellites are decommissioned at the end of their life, so they will not pose a threat to other satellites in the same orbit.

The measurements depend on laser pulses to be sent, returning photons to be detected and for the emitter and receivers to be in orbit with the instruments calibrated. For the data to be transmitted to Earth, the antenna needs to be pointed to the ground station and the data package has to be transmitted. Data analysis depends on the data to be received on the ground and the distribution between the research institutes.

To have the satellite send out laser pulses the satellite needs to point down (nadir-pointing) and the emitter needs to send the pulses. Receiving photons depends on the receiver being pointed at the ground target and the receiver is able to register incoming photons. For the data package to be transmitted, data is put into a buffer, a data package is made, the package is coded and the receiver sends the data to the emitter satellite, which in turn forwards it to Earth. It is important to have the **EPS!** (**EPS!**) properly functioning and instruments deployed before they can be calibrated.

Determining the attitude error of the emitter and adjusting the it accordingly is required to point the emitter towards Earth. When an incoming photon is received, it is transformed into a digital signal and the signal is sent to the on-board computer so the photon is registered. For the receiver to be pointed at the ground target, the emitter needs to notify the receiver it has sent pulses. The receiver needs to receive the message, interpret it, determine attitude error and adapt accordingly to make sure that data gathering will be carried out properly.

# Chapter 4

## Risk Management

### 4.1 Technical Risk Assessment

The main objective of the technical risk assessment is to determine the reliability compared to the possible (functional or financial) consequences per specific event. To be able to determine any of these reliabilities, a definition of reliability should be stated. In this case, reliability is formulated as:

*The probability that a specific (part of a) subsystem will function without endangering the top level requirements over the expected lifetime.*

Next to formulating the definition of reliability, it should be noted that the determined reliabilities in this section are relative reliabilities, i.e. the probability that a particular subsystem outperforms another subsystem with the same core function in terms of reliability. Hence, no absolute values of reliability are determined in this section. The relative reliabilities allow for comparison material during the trade-off between multiple design options. The risk assessment analysis is divided into four main sections:

- I Ground segment (before vehicle leaves Earth's atmosphere)
- II During mission
- III Measurement protocol
- IV Post-mission

The possible events, with their respective reliability, are outlined in these sections and after that the expected consequences are shortly explained.

#### 4.1.1 Ground Segment

##### A *Financial*

###### *A1. Insufficient funds or low market-demand*

The approximate costs are determined in the cost budget. The mission data and the final results can be very interesting for a vast number of commercial parties and research or educational facilities. Every space mission is created for at least one specific (user-demanded) requirement set by a user. This third party is responsible for covering the cost. Since the space mission is developed after this request is set,

the probability that there will be insufficient funds is low (especially when more than one company can be considered as the user). However, the consequence can be severe if the funds are not enough to start or continue the development.

## B *Technological readiness*

### *B1. Technology for level zero requirements are not available*

If the technology for measuring, detecting or processing the level zero requirements is not available at present, the requirements cannot be met and alternatives should be devised, or the mission should be terminated. In our case, the technical readiness level of the payload is relatively high and hence has a high reliability. If, however, the specific payload would have a low technical readiness level, the mission should be terminated or delayed. Therefore, it has important consequences to the mission.

## C *Launch*

### *C1. Total launch failure*

Total failure indicates complete failure of the launch vehicle and all laser swarm constellation components. Needless to say, the reliability is relatively low; however, the consequences of this event are catastrophic.

### *C2. Partial launch failure*

Partial launch failure indicates non-complete failure of laser swarm constellation components, i.e. some of the satellites (more receivers and one emitter) can still perform core tasks. Considering historical launch data, the reliability that no partial failure will occur during launch is relatively high. The consequence can be very different, depending on which part (or what fraction) of the constellation cannot perform its core task. If one of the receivers will be destroyed, the level zero requirements might still be achievable. However, if the emitter is (partially) destroyed, the mission will surely be endangered.

### *C3. Delayed vehicle launch*

Delaying the vehicle launch is not particularly a problem from the technological side of the mission; however, it will affect the financial situation. Next to the fact that the data and results are delayed, extra costs will be imposed due to an increase in launch vehicle pad costs, extra personnel costs and others. The reliability of this event is actual not that high, since it is dependent on a lot of criteria like third-party companies, the weather, and atmospheric properties. The consequences are mainly financial.

### 4.1.2 During Mission

## D *Orbit accuracy*

### *D1. One or more satellites are in a wrong orbit.*

After launch and orbit initializing, it is possible one or more of the satellites are in a wrong orbit. If this deviation from the desired orbit is relatively small, the ADCS subsystem should be able to cope with this minor error and adjust the orbit. If the altitude error is large, major altitude changes should be imposed. Assuming a low to moderate error, the consequence is not really severe if the ADCS system is working properly. The chance of actually putting a satellite in the wrong orbit is also relatively small.

In the next section, the reliability of the ADCS subsystems is compared. Assuming a non-hybrid spacecraft, i.e. a spacecraft which uses one of the ADCS subsystems considered in the design option tree, the consequences of failure are equal for all subsystems and thus shall not be inspected individually. The consequences are severe considering not only the loss of pointing accuracy, but also a decrease in vehicle stability and the total failure of controllable altitude control.

## E *Altitude and control determination*

### *E1. Passive systems*

*E1a. Gravity-gradient.*

The gravity gradient technique is only dependent on gravity fluctuation in nadir direction, which makes it relatively reliable.

*E1b. Passive magnetic.*

The passive magnetic technique is only dependent on magnetic fluctuation near a celestial body. Since this is the only dependency, the technique is really reliable.

*E1c. Zero momentum.*

The zero-momentum technique uses a momentum-bias wheel, initially with no angular velocity. Like with all mechanical systems, the presence of (angular) motion will decrease the reliability (due to possible mechanical failure like static failure, fatigue etc.). The reliability however is pretty high, but lower relative to passive magnetic and gravity gradient.

*E1d. Momentum-bias wheel.*

The momentum-bias wheel technique uses, like the name already predicts, momentum wheels to dump and correct torques. In that sense it has the same reliability as the zero-momentum subsystem. However, since these momentum wheels are constantly spinning, the reliability is slightly lower than the previous mentioned subsystem.

*E1e. Spin stabilization.*

Spin stabilization can be achieved using rotation about one principal axis (single-spin) or two principal axes (dual-spin). Next to the fact that due to external torques (debris collision, aerodynamic drag) the spacecraft can become unstable, i.e. this subsystem has more dependencies, making it relatively unreliable.

Figure ?? demonstrates relative reliabilities of different passive **ADCS!** (**ADCS!**) systems.

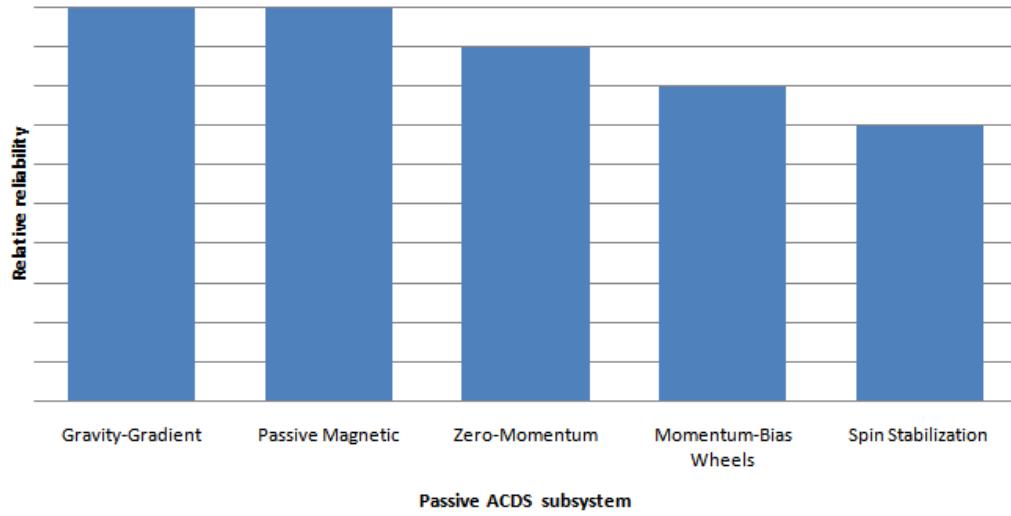


Figure 4.1: Relative reliability of passive ADCS subsystems.

## F Active systems

### F1. Actuator

#### *F1a. Thrusters (hot and cold gas).*

Multiple-axes thruster systems are very efficient ways for controlling attitude and stability. The system

is dependent on fuel consumption, combustion and mechanical properties. Each of these dependencies decreases the reliability.

#### *F1b. Reaction and momentum wheels.*

Mechanical reliability is an import aspect for using active reaction and momentum wheels.

#### *F1c. Control Moment gyros.*

A control-moment-gyro system consists of a spinning rotor and one or more motorized gimbals that tilt the rotor's angular momentum. Mechanical reliability is an important aspect for using this. Since it is also dependent on the motorized gimbals, the reliability is slightly lower than the reaction and momentum wheels.

#### *F1d. Magnetic torquers*

The magnetic torquers interact with the Earth's magnetic field, creating compensating torques to induce stability. Reliability is high due to the fact that the magnetic field is known and the system is dependent on a low number of parameters.

### *F2. Sensors*

Assuming a high technical readiness level of the sensors, the reliability is considered high. Also, the consequences of failure are high as the continuation of the mission may be impaired.

Figure ?? demonstrates relative reliabilities of different active **ADCS!** systems.

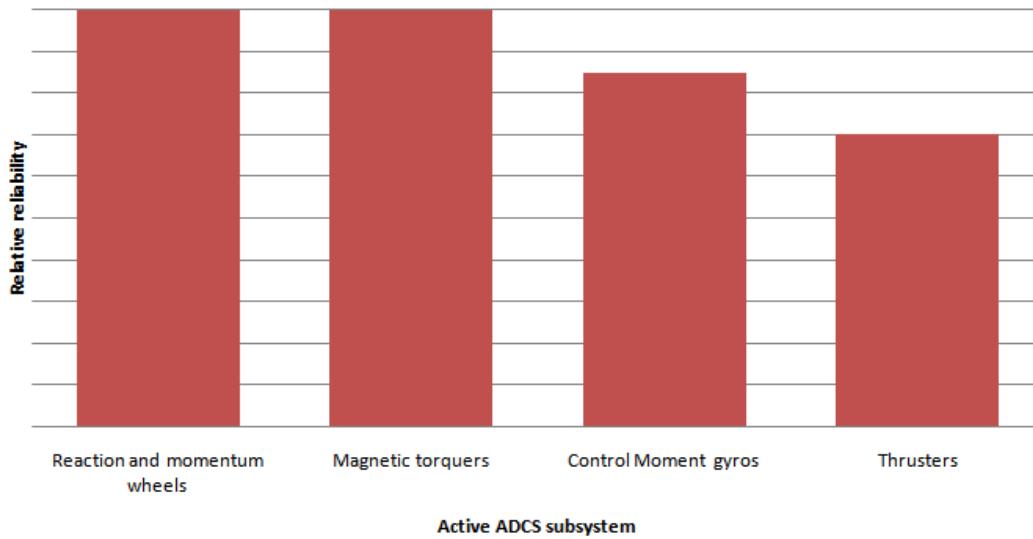


Figure 4.2: Relative reliability of active ADCS subsystems.

## G **Electric Power System (EPS)**

### *G1. Solar Panels*

#### *G1a. Solar panel deflection error or mechanical failure.*

During launch, the solar panels are retracted to achieve the lowest volume possible. During the initializing of the mission (assuming the spacecraft is in the right orbit), the solar panels need to be deployed. Errors can occur due to mechanical reasons or external disturbances. The probability of this is pretty low. The consequence can be however that the effective solar panel is decreased and hence a decrease in available electrical power will occur. This makes the consequences pretty severe. Any other mechanical failure

(broken joints, internal PN-junction failure, or a loss of an entire solar panel) will have severe consequences as well.

*G1b. Solar panel characteristics reliability (degradation).*

Degradation of solar panels should always be considered during mission development. Since this (should be) known upfront, the consequences are relatively low. The probability of this actually happening is nearly 100%.

*G1c. Severe degradation (due to external phenomenon)*

Atomic oxygen, hazardous radiation, debris collision and other external factors can influence the performance of the solar panels. Since these are not known from the start, it is difficult to cope with them. The probability of this happening is pretty small, but will have pretty severe consequences.

*G2. Batteries*

*G2a. Initial internal failure*

Considering a high level of technical readiness level, the internal reliability is high. The consequences do alter the functional capacity of the mission, since no energy can be stored if the energy capacity system would completely fail, meaning that during eclipse no energy can be used.

*G2b. Decrease in capacity*

Considering a high level of technical readiness level, the reliability is high. Consequences are low, because they are known and should be part of the mission analysis.

#### 4.1.3 Measurement Protocol

Since actual measurements are an important level zero requirement, the consequence of the items in the measurement protocol are all really severe. Unless stated otherwise, the consequences in the following section can thus be stated in this way.

**Measurement**

**H Emitter**

*H1. Laser pulses cannot be sent/ no photon generation*

Considering a high level of technical readiness level, the reliability is high.

*H2. Pointing towards nadir*

This is dependent on ADCS risks.

*H3. Laser notifies receiver (time adjustment)*

Considering a high level of technical readiness level, the reliability is high.

*H4. Laser degradation*

Laser degradation is dependent on multiple parameters: thermal properties, input power interval, external factors and internal mechanical errors (manufacturing or design errors). However, due to extensive research and development concerning laser technology, the probability of severe laser degradation within the lifetime is relatively low.

**I Receiver**

*I1. Point towards target*

This is dependent on ADCS risks.

*I2. Receive and detect photons*

Considering multiple satellite receivers, the probability of total failure to receive and detect photons using advanced single photon receiving devices (like SILAT, **GLAS!** or photon-receiving modules) is negligible.

*I3. Turn photon into electrical signal*

Considering a high level of technical readiness level, the reliability is high.

**Communication****J Inter satellite communication***J1. Determine relative position receiver and emitter*

Considering a high technical readiness level, the reliability is high.

*J2. Time differences*

Considering a high technical readiness level, the reliability is high.

**K Data handling***K1. Store data/ make data package*

Considering a high level of technical readiness level, the reliability is high.

*K2. Transmit package*

Considering a high level of technical readiness level and relative low-tech technology, the reliability is high.

*K3. Interpreted results*

Historical data comparisons for the interpretation of altimetry missions are sufficient, but not elaborate. However, the reliability is still pretty high.

*K4. Reproduce terrain model*

Historical data comparisons for the interpretation of altimetry missions are sufficient, but not elaborate. However, the reliability is still pretty high.

**L Housekeeping/ Ground communication***L1. Housekeeping data from ground station to satellite*

Considering a high level of technical readiness level, the reliability is high.

*L2. Adjusting space segment characteristics*

Considering a high level of technical readiness level, the reliability is high.

**M Structural***M1. Joints*

Considering a high level of technical readiness level, the reliability is high.

*M2. Connection points*

Considering a high level of technical readiness level, the reliability is high.

*M3. Thermal limits*

Thermal limits will alter the characteristics of pretty much all subsystems. However, thermal will be excluded in this analysis, because this is not a part of the project.

*M4. Fatigue*

High-cycle loading is usually not present (except for momentum wheels) and should therefore only play a minor role. The probability is low. The consequences are medium to high if high-cycle loading will lead to fatigue and hence partial failure.

*M5. Electrical overlay failure*

This event is dependent on the reliability of the EPS.

*M6. Launch loads*

Due to large forces and vibrations during launch the structure of the satellite can fail. Since the launch loads are well known the reliability is low, the consequences can be severe.

**N External***N1. Debris collision*

Unknown external phenomenon. The probability of this event causing total failure, which is a function of multiple parameters like orbit and celestial position, is low to medium.

*N2. Dangerous radiation*

Unknown external phenomenon. The probability of this event causing total failure, which is a function of multiple parameters like orbit and celestial position, is low to medium.

*N3. Charged particles collision*

Unknown external phenomenon. The probability of this event causing total failure, which is a function of multiple parameters like orbit and celestial position, is low to medium.

*N5. Politics or international influence*

Political decisions or international influences can alter the space mission considerably. With altimetry missions, the probability of these external influences causing a delayed or complete stop of the mission is negligible. The consequences, however, could be very severe.

*N6. Classified information (military)*

Some parts of the measurement are considered classified information, for example military ground stations or governmental classified areas. The government and/ or military can pressure the vehicle engineers to keep certain information classified. However, if this is the case, most of the measurements still can be taken and analyzed. So where the probability is medium, the consequences are very low.

**4.1.4 Post-Mission****O Satellite decommission***O1. Decommission LEO*

At the end of life the satellites have to be decommissioned to allow new mission to take their place. To decommission satellites in **LEO!** (**LEO!**) one could just wait a couple of years and air drag will cause the satellites orbit to degrade to the point when they can burn up in the atmosphere, so the consequences are low. However, it is desirable to have the satellites burn up faster, so as to remove the risks of satellite collision. The probability to no longer be able to eject the satellite from orbit depends on whether or not its propulsion system is still working; as such this probability is low.

*O2. Decommission GEO*

Satellites in GEO can not be placed in an orbit that will cause them to burn up in the atmosphere because they will cross paths with too many other satellites. Because risk is so high these satellites are instead decommissioned by ejecting them from orbit further into space. This way, new GEO satellites can take the place of the swarm. If this is not done, then the dead satellites will continue orbiting the Earth, wasting space that can be used by other satellites, as a result the consequence of failure is high. Being able to reposition a satellite depends on the ADCS systems; as such the probability of this event is low.

**4.1.5 Risk Control**

Sometimes it is possible to decrease the failure probability. For example, A1, F1a and H4 could be the top 3 risk segment. Failure probability of A1 can be pulled down by doing detailed market analysis. In case of F1a, safe combustion performance as well as fuel consumption can be tested and modified in a laboratory environment to increase the thrusters' reliability. Laser degradation (H4) is crucial in the system, and reliable energy source should be used to prevent failure. Meanwhile, thermal control can be performed to achieve high reliability of the laser emitter.

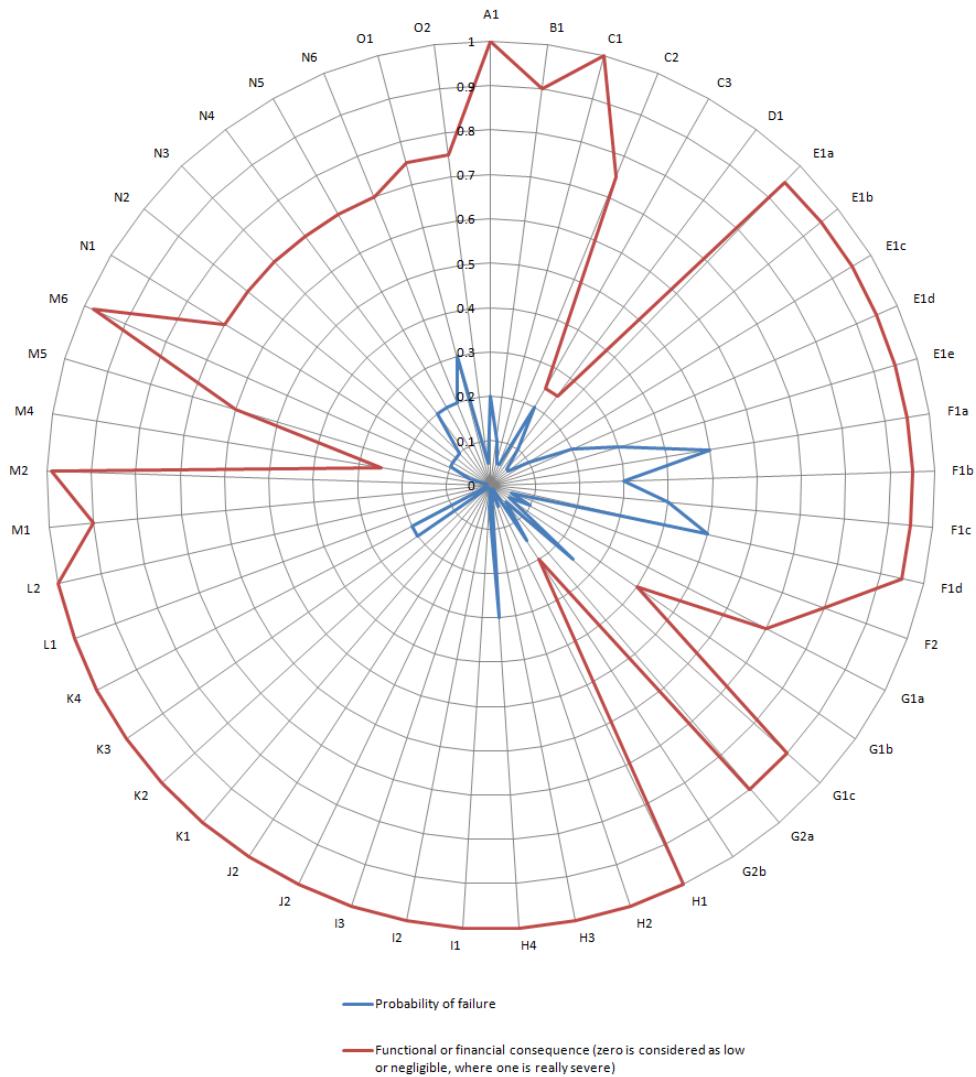


Figure 4.3: Technical Risk Map

# Chapter 5

## Launch and Astrodynamical Characteristics

This chapter explains the characteristics of the launch and space segments of the mission. Section ?? describes the three main characteristics of bringing the swarm into the orbit: launch vehicle, launch location and orbit insertion. Section ?? discusses all the aspects of the constellation and its configuration, and delves into estimations of  $\Delta V$  and needed propellant. Finally, section ?? deals with the hazardous space environment and ways of protecting the satellites against it.

### 5.1 Launch Segment

The launch segment of the mission includes the selection of the launch vehicle, the launch location and the procedure of inserting the satellites into their respective orbits. The following subsections discuss all of these aspects.

#### 5.1.1 Launch Vehicle

The costs of launch vary greatly between different vehicles. Table ?? on page ?? lists approximate total launch costs of respective platforms. The prices are given in Fiscal Year 2000 dollars for consistency.

Platform	Operator	Price (FY00\$m)
Ariane V	ESA	97.96
Soyuz	Starsem or Arianespace	8.16 - 22.1
Vega	European and Italian SA	15.1
Falcon 1E	SPACEX	8.89
PSLV	ISRO	13.88 - 16.33
Rokot	Eurokot	9.8 - 11.4

Table 5.1: Estimated price comparison of different launch vehicles. *Source: various.*

Based on the information in the above table it is possible to single out a few platforms which will be

affordable for the purpose of this feasibility study. The values represent the total launch segment costs. The Ariane V launcher is the most expensive option by far and would push the budget quite heavily, with an estimated cost of launch almost half of the total budget. However the payload capabilities of the Ariane V launcher far outweigh that of all other platforms, thus making it possible for a combined launch with other satellites, leading to shared costs. This however could jeopardize the mission in the sense that it becomes secondary priority. If that happens, the constellation would have higher requirements for orbit acquisition: an extra booster stage or higher onboard fuel capacity for the altitude and/or plane shift, which is not feasible. The Ariane V is therefore an unsuitable platform for this project.

The rest of the launchers can be analyzed with respect to reliability. All of the launch vehicles have been tested, with the exception of the Vega system, which is yet to make its maiden flight. The Vega is therefore not suitable for the analysis at this time. It is for this reason that the project will no longer consider this system at all. However, better data should be available in the near future which should allow for the Vega platform to be reevaluated.

The same goes for the Falcon - 1e system. The launcher is still under development. Furthermore, the predecessor of the 1e system is the Falcon - 1, which out of total of five launches only had two successful, fails to make a favorable impression.

Table ?? on page ?? shows some reliability statistics for the remaining 3 vehicles.

Platform	Total No. of Launches	Total Failures	Reliability	No. of Successful Launches Since Last Failure
Soyuz	1754	88	95%	57
PSLV	16	1	94%	15
Rokot	17	2	88%	6

Table 5.2: Reliability figures for several launch vehicles. *Source: various.*

The Soyuz launch vehicle presents itself as the most reliable platform, with a track record that far surpassed all other options. This launch vehicle will be the one considered for this project. In the next section, its payload capabilities are discussed.

### 5.1.1.1 Soyuz LV Payload Capability Analysis

There is a large selection of Soyuz vehicles available for consideration. For the purpose of this project, the newest modification - Soyuz-ST will be used. This vehicle is part of the Soyuz-2 family, which are technologically superior to the older Soyuz-U and U2 launchers. An illustration and technical parameters of the Soyuz-ST can be found in Appendix ?? on page ?? [?].

An estimation of mass performance for the launch vehicle can be seen in figure ?? on page ??.

The payload mass data provided in [?] is estimated, yet is good enough to have a reasonable idea about the maximum mass. The Soyuz-ST is able to launch roughly a maximum of 5000 kg into a 500km orbit. This is well above the design mass of the formation thus will allow for further considerations of joint launches (as long as the swarm mission is considered to be the primary payload) and thus spread launch costs.

The available volume in the Soyuz-ST Fairing can be seen in figure ?? on page ??.

The dimensions of the fairing are visibly too large and there is no possibility of using a different one, however that leaves a lot of possibilities for different designs of release adapters to adapt to the unique sequence of separation upon orbit injection. Again, the possibility of taking other small satellites along on

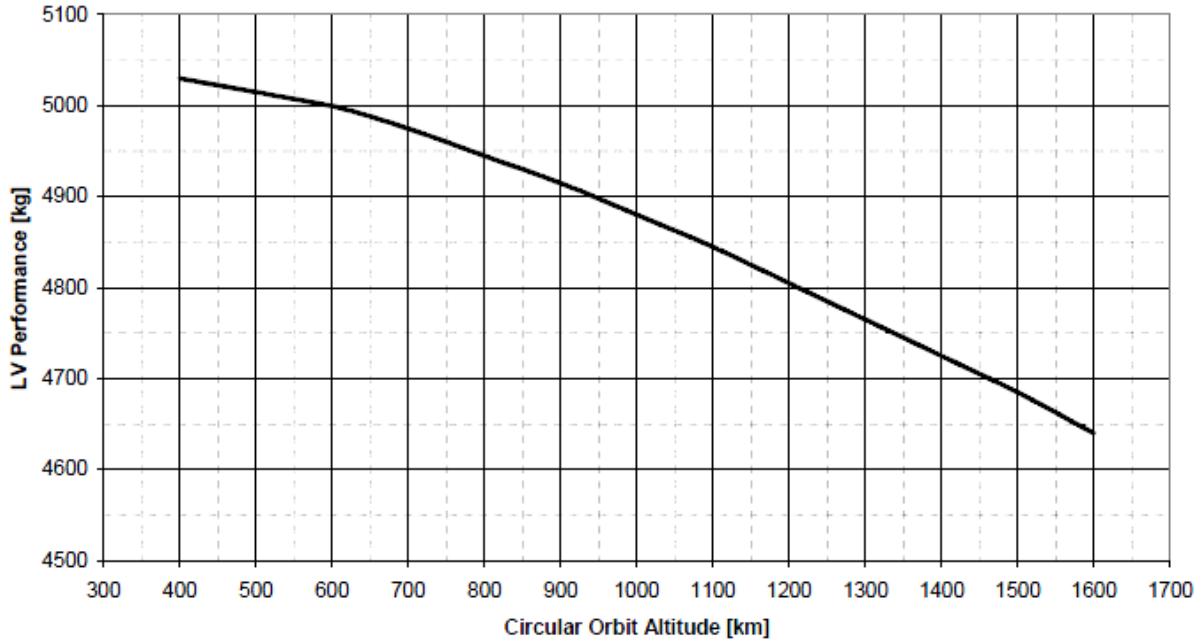


Figure 5.1: Mass performance of the Soyuz-ST for circular Orbits. *Source: [?].*

the same launch arises. The dual launch configuration shown in the previous figure is the one being considered for the launch.

The choice of the Soyuz brings forth one more advantage: the use of a unique orbit insertion booster, the Fregat. This final stage will allow for minimization of fuel on the satellite as all orbit insertion maneuvers can be done using the Fregat. The Fregat stage has been designed to handle 20 individual burns.

### 5.1.1.2 Vibrational Analysis

HERE ADD VIBRATION ANALYSIS

### 5.1.2 Launch Site

The selection of the launch site relies on several factors:

- Availability of attainable inclinations from launch.
- Compatibility with the launch vehicle.
- Accessibility and cost.
- Security and political reasons.

The first factor is crucial. It is paramount that the satellites are injected into their final inclinations at launch and do not have to perform any inclination change maneuvers, which require a substantial  $\Delta V$ . With this in mind choosing a launch site closer to the equator is necessary. Launch sites at higher latitudes

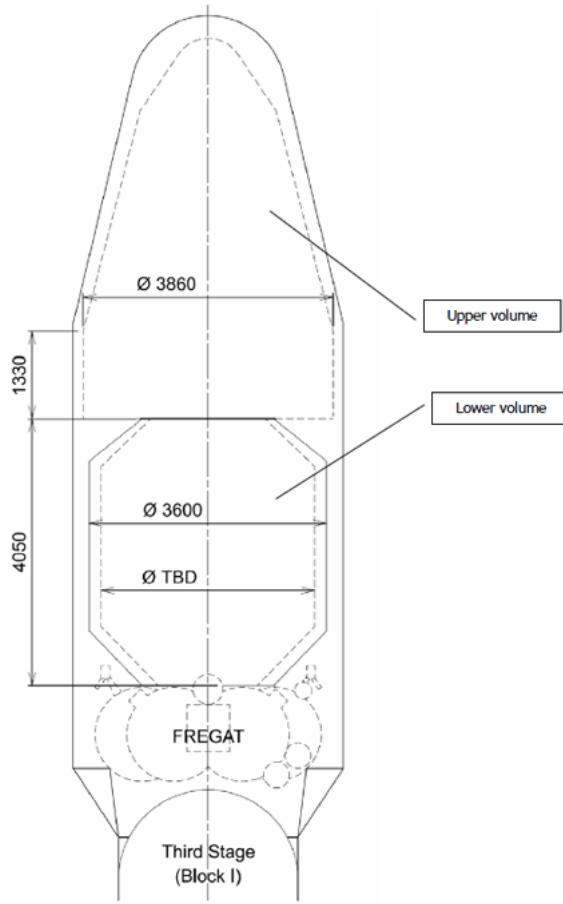


Figure 5.2: Fairing volume of the Soyuz-ST launch vehicle. Here shown in dual launch carrying configuration. *Source: [?].*

would need to sacrifice velocity and thus payload mass because of their location. Table ?? on page ?? shows a number of possible launch sites and their locations [?]. The list contains only the sites compatible with the Soyuz-ST launch vehicle. Furthermore, in figure ?? on page ??, the same sites are indicated with their respective authorized inclination ranges.

Launch Site	Operator	Latitude (deg min)	Longitude (deg min)
Baikonur LC-31/6	Russia (Starsem)	45 54 N	63 18 E
Plesetsk LC-43	Russia (Starsem)	62 48 N	40 24 E
Guiana Space Centre ELS	CNES/Arianespace	5 18 N	52 50 W

Table 5.3: Available launch sites for the Soyuz-ST. *Source: [?].*

Even though all the above sites allow inclinations of 85 degrees, it is still preferable to select a site closer to the equator in order to utilize the full effect of Earth's rotation and lower the launch costs. With this in mind, the Guiana Space Centre is selected to be the preferable location for launch.

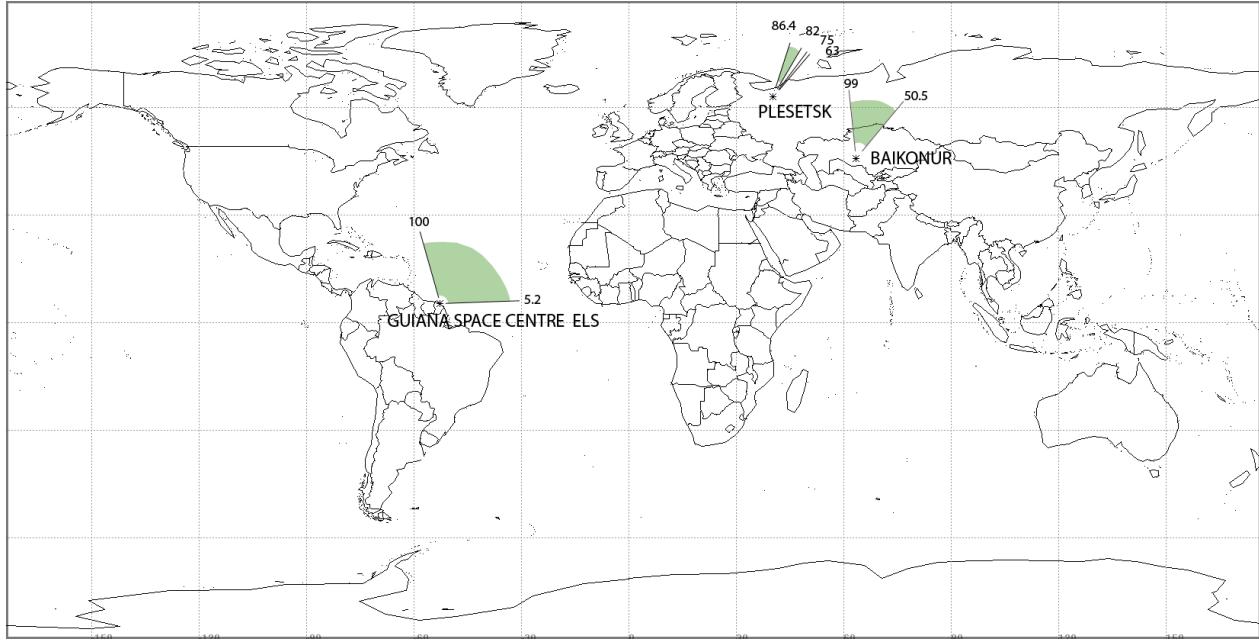


Figure 5.3: Launch site locations and allowable inclinations for the Soyuz-ST. Sources: [?] and [?].

The **ELS!** (**ELS!**) in Kourou, Guiana is currently being finalized and should accommodate its first Soyuz launch in 2010 [?].

A typical launch profile of the Soyuz launch vehicle from Kourou is shown in figure ?? on page ?? . The first three stages of the vehicle are used to propel the payload and the Frigate booster to a circular orbit at around 200 km. After separation (stage 6 on the figure), the Fregat initiates the first orbit injection burn to bring the satellites to the appropriate orbits.

With this information it is possible to further calculate some important launch parameters:

- The inertial velocity of the launch site is given by:

$$V_L = (464.5)\cos L \quad (5.1)$$

where  $L$  is the site latitude. For the case of Kourou the inertial velocity is 462.51 m/s.

- The launch azimuth in inertial frame of reference is given by:

$$A_{Z_I} = \arcsin\left(\frac{\cos i}{\cos L}\right) \quad (5.2)$$

and is equal to 5.022 degrees for this launch.

- The launch azimuth corrected for the Earth's rotation, is given by:

$$A_Z = \arctan\left(\frac{V_0 \sin A_{Z_I} - V_{eq} \cos L}{V_0 \cos A_{Z_I}}\right) \quad (5.3)$$

where  $V_0$  is the orbital velocity reached by the launcher before separation and the first burn of Fregat upper stage (see above, around 7.784 km/s),  $V_{eq}$  is the velocity of the Earth's rotation at the equator - 464.5 m/s. Thus the corrected launch azimuth becomes 1.61 degrees.

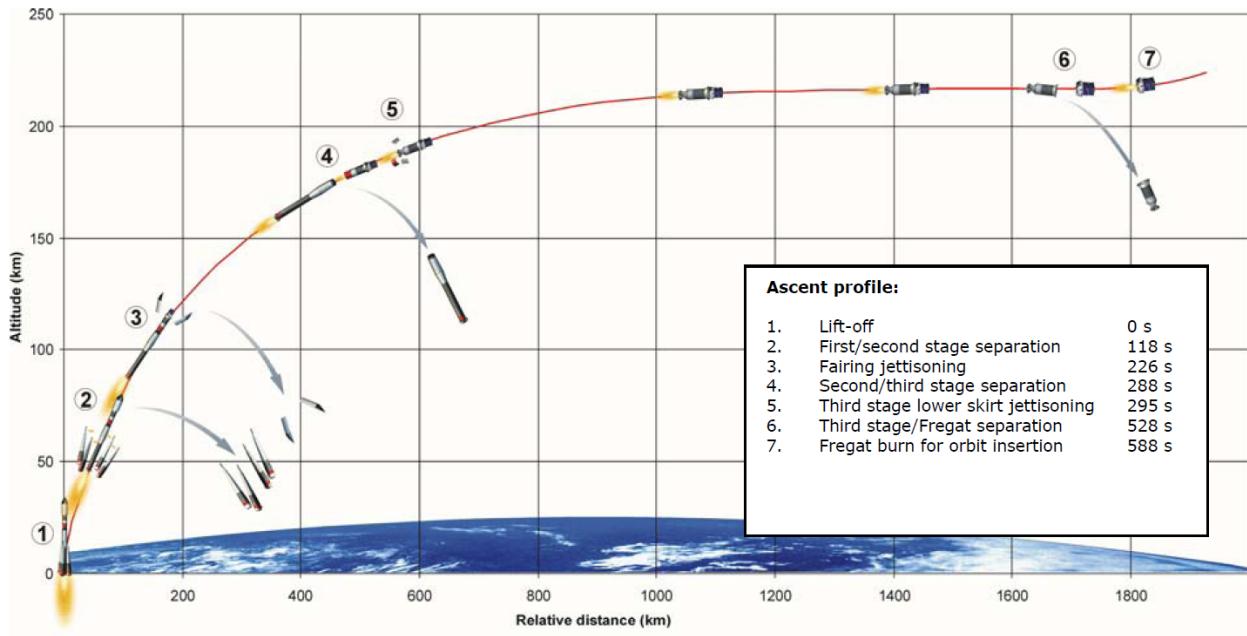


Figure 5.4: Launch profile of the Soyuz LV from Kourou. Sources: [?].

- The required launch velocity is calculated to be 7.76 km/s which means that due to Earth's rotation 26.8 m/s are saved.

Furthermore the region is politically stable and the launch site is under the protection from the French government and other security forces. All safety procedures are kept to the highest of standards and the spaceport is easily accessible by air or sea [?].

### 5.1.3 Orbit Insertion

The orbit insertion is separated into two distinct stages: primary orbits and secondary orbits. The primary orbits are located at an altitude of 500 km and contain the emitter and four initial receivers. The secondary orbits are located at a slightly different altitude of 525 km and contain the auxiliary receivers intended for replenishment.

The configuration of the primary orbits is discussed in section ?? and an image of the formation can be seen in figure ?? on page ?? . The release sequence and the important parameters are as follows (for satellite and orbit numbers please refer to figure ??):

1. The Fregat is injected in Orbit 1.
2. The ascending intersection of Orbit 1 and Orbit 2 is reached at a latitude of  $85^\circ$  and a longitude offset (from the ascending node of Orbit 1) of  $88.91^\circ$ . At this point the Fregat should be orientated in the direction of Orbit 2 and separate Rec 2. The separation  $\Delta V$  that the adapter should produce is calculated using:

$$\Delta V = 2V_i \sin \frac{\alpha}{2} \quad (5.4)$$

where  $V_i$  is the orbital velocity (7.612 km/s) and  $\alpha$  is the relative inclination ( $2.17^\circ$ , see section ?? [?]). The  $\Delta V$  is calculated to be 289.63 m/s.

3. As the Fregat crosses the descending node and approaches the second plane intersection of the orbit, it does not need to change orientation (Orbit 3 intersects in the same direction on the descent phase as Orbit 2 does on the ascent phase). At the intersection Rec 1 is separated with a  $\Delta V$  of 289.63 m/s. This is a very large value thus the method of deployment should be thoroughly investigated in further design.
4. After this the Fregat again aligns with the velocity vector of Orbit 1.
5. The Frigate should inject itself into a drift orbit with a negative drift rate of no less than 9.19 deg/orbit and re-injected back to Orbit 1 after traveling  $90 + \Delta\phi/2 = 90.09505$  degrees. This will bring the launcher 2.3 degrees behind the final emitter position, or 0.12 degrees behind Rec 4.
6. At this point the remaining three satellites: Rec 3, Base and Rec 4 should be put into drift orbits by the attachment mechanism in order to acquire the 2.18 degree orbital separation. The exact order and timing can be designed and adjusted accordingly.

After this, the satellites can be considered to be in their orbits. The Fregat can start the burn to insert into the secondary orbits. Once the altitude of 525 km is reached the Fregat performs a plane shift to change its **RAAN!** (**RAAN!**) to approximately 10 degrees behind that of the primary orbit. The exact angle has to be carefully acquired through careful modeling of differential node precessions as the satellites decay. The general idea is that the secondary satellites should have the correct RAANs by the time they decay to arrive in the right position with respect to the emitter. The formation in the secondary orbit is the same as in the primary orbits minus the emitter, thus similar maneuvers have to be performed.

#### 5.1.4 Launch Date

The launch date is dominated by development times and lifetime considerations. The satellite orbit decay is and thus lifetime is a function of atmospheric density. The density is, in turn, a function of the solar activity. The number of sun spots on the surface of the sun rises and falls every eleven years. The measurements are commonly represented in 10.7cm radio flux intensities. Figure ?? on page ?? presents a projection for the next 10 years.

In order to provide the longest possible lifetime for the mission it is essential that the launch is timed in such a way that the satellites are in orbit most of the time during solar minimum. For this reason a launch on March 1st 2017 is planned. This also gives enough time for development and production. For further information on the timeline, please refer to section ??.

## 5.2 Space Segment

This section covers the astrodynamical characteristics of the mission. Emitter orbit is covered first, then in section ??, all receiver orbits are examined. The formation and its properties are discussed in section ?? . Collision avoidance is described in section ?? and, finally, section ?? covers the orbital environment and whether the mission is going to be heavily affected by it.

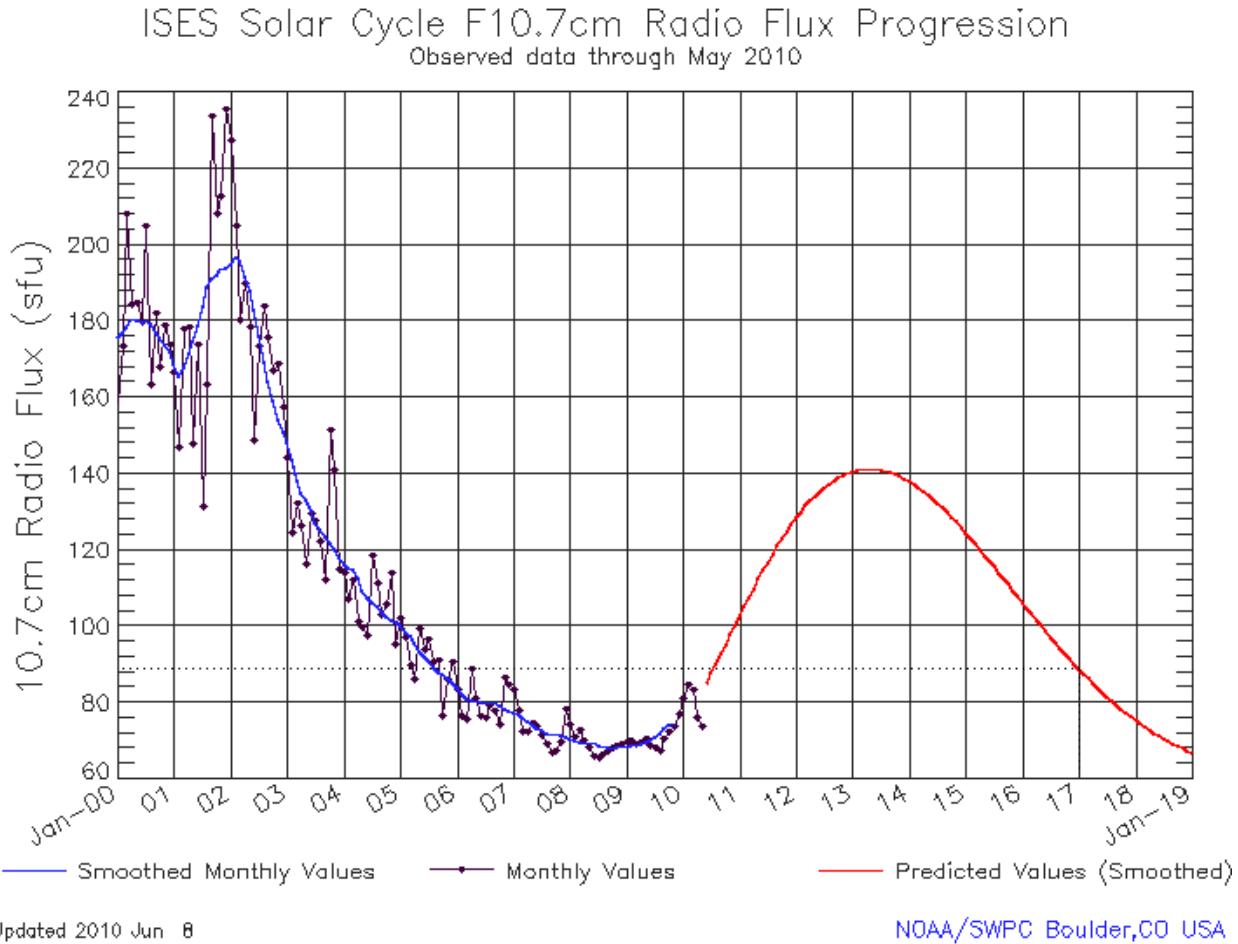


Figure 5.5: Solar activity projection up to 2019. *Source: NOAA/Space Weather Prediction Center.*

## 5.2.1 Emitter Orbit

### 5.2.1.1 Orbital Parameters

The emitter satellite is injected into a circular orbit at an altitude of 500 km. The eccentricity is frozen at 0. Inclination is chosen to be 85 degrees, as this allows access to polar areas, which are places of interest for the mission. Furthermore, the inclination provides an inherent relative phase for the crossing orbits. The **RAAN!** can be chosen arbitrary for the start of the mission, as no specific target is assumed for the mission. For consistency, in all following discussions it is assumed to be 0 degrees. The same goes for the Argument of Perigee and the True Anomaly. As the orbit is circular, these values are not really relevant. The emitter satellite acts as the reference for all the receivers.

The orbit is non-sunsynchronous and does not have a specific repeat track. This allows for larger coverage of the Earth.

Figure ?? on page ?? shows the emitter satellite (labeled as Base) in Orbit 1.

A few basic properties of this orbit can be derived and are shown in table ?? on page ???. All the values were generated with the help of [?], [?] and [?].

Parameter	Symbol	Value	Unit
Altitude	h	500	[km]
Semi-major Axis	a	6871	[km]
Eccentricity	e	0	[‐]
<b>RAAN!</b>	$\Omega$	0	[deg]
Period (mins)	P	94.6135	[mins]
Revolutions per day		15.2198	[revs/day]
Angular Velocity	n	3.805	[deg/min]
Circular Velocity	V	7.6127	[km/s]
Max. Eclipse	$T_e$	35.75	[mins]
Node Spacing		23.72	[deg]
Node Precession	$\dot{\Omega}$	-0.6667	[deg/day]

Table 5.4: Orbital properties of the emitter satellite.

### 5.2.1.2 Orbit Decay

Atmospheric drag is by far the most relevant perturbation for **LEO!** satellites that cause loss of altitude and thus decays the orbit. It directly relates to mass as it influences the amount of fuel required to maintain the orbit, where as the mass influences the rate at which the orbit decays. Altitude selection relies heavily on estimation and analysis of drag data as for longer mission times, higher altitudes are preferred, while optical instruments prefer lower altitudes for increased accuracy.

The drag that the satellite experiences due to atmospheric density is described by the following formula:

$$D = -\frac{1}{2}C_D \rho V^2 A \quad (5.5)$$

It follows that orbital parameter changes (semi-major axis, period and velocity respectively) per orbit are calculated using the following equations (assuming negligible eccentricity):

$$\Delta a = -2\pi \left( C_D \frac{A}{m} \right) \rho a^2 \quad (5.6)$$

$$\Delta P = -6\pi^2 \left( C_D \frac{A}{m} \right) \rho \frac{a^2}{V} \quad (5.7)$$

$$\Delta V = \pi \left( C_D \frac{A}{m} \right) \rho a V \quad (5.8)$$

The fundamental problem with accurately predicting effects due to atmospheric drag is twofold: firstly it is very hard to predict the satellite's ballistic coefficient:

$$\frac{m}{AC_D} \quad (5.9)$$

Even with a well known mass to area ratio, the coefficient of drag can be highly variable, highly dependent on the shape of the satellite and its orientation with respect to the velocity vector. It is usually determined in laboratory conditions. For the orbit decay analysis the following equation for the coefficient of drag was used:

$$C_D = \alpha C_{DS} + \beta C_{DD} \quad (5.10)$$

where  $C_{DS}$  is the specular coefficient of drag,  $C_{DD}$  is the diffuse coefficient of drag,  $\alpha$  and  $\beta$  are component fractions which are determined experimentally. The specular coefficient of drag is usually predominant. In reality this drag coefficient changes. The cross-sectional area normal to the velocity vector can also vary for the swarm satellites if the whole platform is reoriented for instrument pointing. These parameters were adjusted in such a way that a final average  $C_D$  of 2.22 for the emitter as well as the receiver was achieved. This is a value which is very commonly used in space mission design.

The cross-sectional area of the satellite perpendicular to the velocity vector for a sample period of 6 hours can be seen in figure ?? on page ???. The graph was generated using the correct geometry of the satellite simulated in the Satellite Tool Kit™ software, issued by AGI, Inc. The software also computes the mean area - 1.045948 m<sup>2</sup>. This will be the area used for lifetime analysis. The mass used is final estimated satellite loaded orbit mass from section ???. It is computed to be 22.87 kg/m<sup>2</sup>. This is well within the normal range for regular satellites [?].

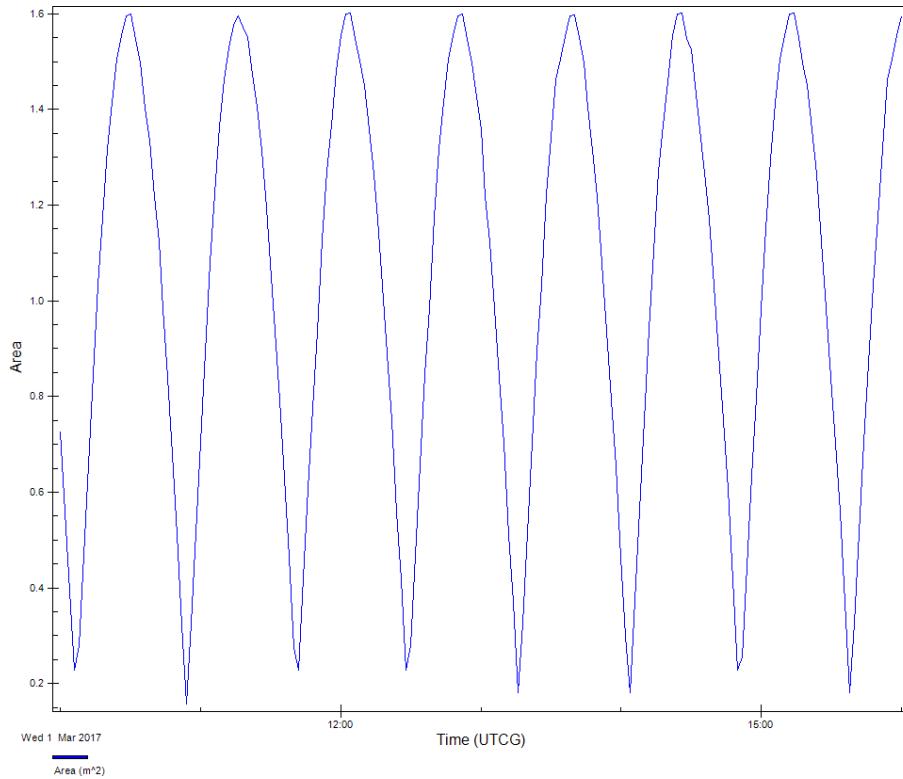


Figure 5.6: Simulated cross-sectional surface area of the emitter satellite perpendicularly to velocity vector.

The second reason drag calculations are so unreliable, is because air density at any altitude is highly variable. Raising air density is primarily connected with solar activity. As solar activity increases every 11 years the atmosphere heats up. Contrary to conventional gas laws that would dictate a fall in density as the gas expands, the atmosphere simply rises, increasing density at higher altitudes. This was previously discussed in section ??.

The density difference during maxima and minima for different altitudes is shown in figure ?? on page

???. Depending on the altitude, the density could vary for up to a whole order of magnitude between the minimum and maximum.

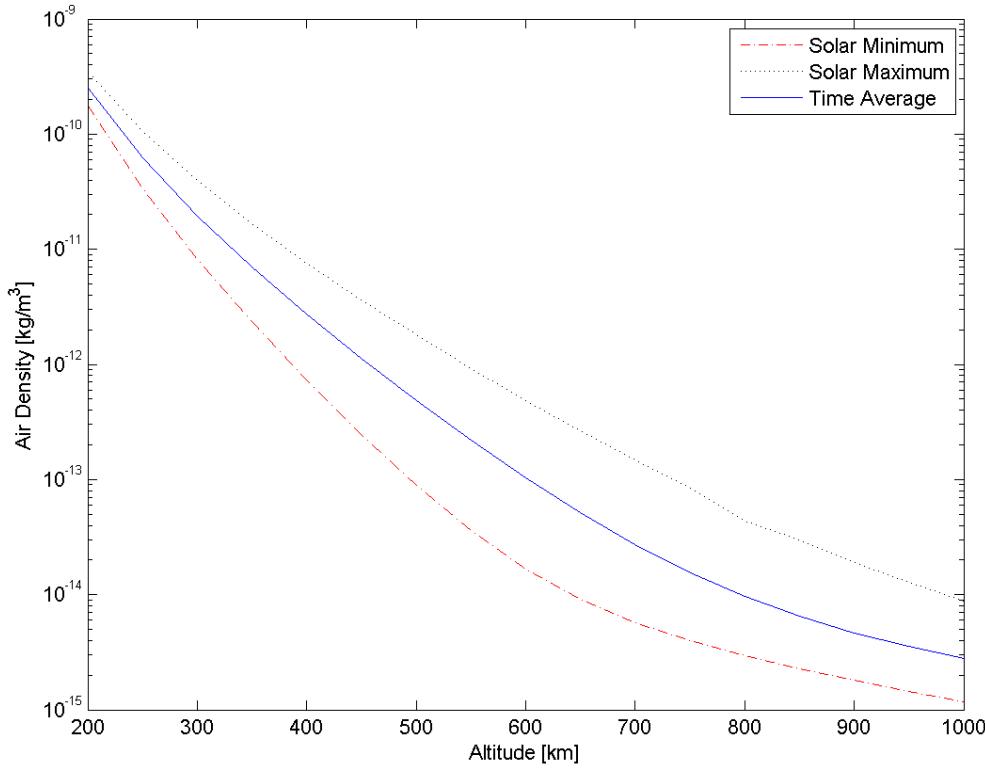


Figure 5.7: Air density vs. orbit altitude for different solar cycle stages.

The orbit lifetime simulation was produced using the software package Satellite Tool Kit™ and can be seen in figure ?? on page ??.

The simulation performs orbit propagation up to the J<sub>4</sub> factor and uses the MSIS-86 Thermospheric Model for atmospheric density calculation. In addition it takes into account solar radiation pressure with an average sun exposed area of 0.84 (this area is also estimated from the satellite model simulations).

It is obvious from the previous figure that, unmaintained, the orbit would totally decay in 4.8 years. Another notable time is the point at which the emitter passes the 450km altitude. This occurs at approximately 3.4 years after mission start. This is the altitude that will be considered as the absolute floor and will be used to approximate the altitude of the secondary satellite orbits.

### 5.2.1.3 Estimation of $\Delta V$

In order to properly estimate the mass of the propellant required for the mission all maneuvers have to be taken into account. Every maneuver requires a certain velocity change in a specified direction. During its mission, the emitter satellite will perform the following maneuvers:

- A phase change slowdown. As the satellite was released from the booster vehicle, it was injected into a

drift orbit in order to separate into a certain phase shift. Once this phase shift is achieved, the satellite needs to perform a boost to bring itself back into the circular orbit. This  $\Delta V$  is equal and opposite to the  $\Delta V$  induced by the release mechanism. This  $\Delta V$  is calculated using the following equation:

$$\omega_{drift} = (1080) \frac{\Delta V}{V} \quad (5.11)$$

where  $\omega_{drift}$  is the drift rate given in deg/orbit. For the emitter to return out of its drift orbit in 90 degrees, this maneuver requires a  $\Delta V$  of 64.8 m/s.

- A boost to change the altitude from 450 km to 500 km after approximately 3.4 years in orbit. This maneuver is performed using a simple Hohmann transfer orbit and will require two burns. A simple transfer calculation is completed using the following method [?]:

$$V_{450} = \sqrt{\frac{GM}{a_{450}}} = 7.6405 \text{ km/s} \quad (5.12)$$

$$V_{500} = \sqrt{\frac{GM}{a_{500}}} = 7.6127 \text{ km/s} \quad (5.13)$$

$$a_{ellipse} = \frac{a_{500} + a_{450}}{2} = 6853 \text{ km} \quad (5.14)$$

$$V_p = \sqrt{\frac{2GM}{a_{450}} - \frac{GM}{a_{ellipse}}} = 7.6544 \text{ km/s} \quad (5.15)$$

$$V_a = \sqrt{\frac{2GM}{a_{500}} - \frac{GM}{a_{ellipse}}} = 7.5988 \quad (5.16)$$

$$\Delta V_1 = 1000(V_p - V_{450}) = 13.9237 \text{ m/s} \quad (5.17)$$

$$\Delta V_2 = 1000(V_{500} - V_a) = 13.8984 \text{ m/s} \quad (5.18)$$

$$\Delta V = \Delta V_1 + \Delta V_2 = 27.8221 \text{ m/s} \quad (5.19)$$

Thus the required  $\Delta V$  for the Hohmann transfer is 27.8 m/s.

- A deorbit burn to take the satellite down to comply with sustainability requirements. This  $\Delta V$  is approximated by the following equation [?]:

$$\Delta V_{deorbit} \approx V \left[ \frac{0.5(H_i - H_e)}{2R_E + H_i + H_e} \right] \quad (5.20)$$

where  $H_i$  is the initial altitude,  $H_e$  is the reentry altitude and  $R_E$  is the Earth's radius. Using the initial altitude of 500 km and a reentry altitude of 50 km, the  $\Delta V$  is calculated to be 128.73 m/s. This is a very large value and in reality can be scaled down by considering a lower initial altitude (500 km is assumed in this analysis due to the consideration that a satellite might become dysfunctional at the beginning of the mission and would need to be brought down) and a higher final reentry altitude (which is arbitrary).

- The final consideration is the orbit injection accuracy of the Soyuz launch vehicle. The manufacturer states an inclination accuracy of 0.033 degrees for a 1000 km altitude orbit [?]. For the purposes of this analysis, an accuracy of 0.03 degrees is used. The following equation is used to estimate the  $\Delta V$  needed to correct this error:

$$\Delta V_{planechange} = 2V_i \sin \frac{\alpha}{2} \quad (5.21)$$

where  $V_i$  is the initial orbital velocity. The final  $\Delta V$  is 3.99 m/s.

No  $\Delta V$  is required for stationkeeping as the emitter will be allowed to decay naturally while the receivers will have to perform relative stationkeeping. Furthermore, the Soyuz launch vehicle can have a initial delivery altitude error. This is not considered, since it is assumed that as long as the constellation is delivered on the same altitude, no correction will be required and the lifetime of the mission will not be jeopardized.

The total  $\Delta V$  required for the emitter is calculated to be approximately 225.38 m/s.

## 5.2.2 Receiver Orbits

### 5.2.2.1 Orbital Parameters

Unlike the emitter, the receiver satellites are placed in six different orbits and has to be designed to be able to handle all orbits and yet all the satellites have to have the same mass configuration at all times.

The orbits are divided into primary and secondary. The primary orbits are shown in figure ?? on page ???. Receiver 1 is located in Orbit 3 which has all the same orbital characteristics of Orbit 1 as described in section ?? with the exception of its **RAAN!** having a value of +2.18 degrees. This is the maximum angle of separation and is derived from the requirement of a laser reflection angle of 30 degrees. Receiver 2 in Orbit 2 on the other hand has a **RAAN!** of -2.18 degrees. Both of these receiver satellites pass their ascending nodes at the same time as the emitter satellite. The reason for this is that the inclination of the orbits induces a relative phase between the satellites which is favorable for collision avoidance.

Receivers 3 and 4 are placed on the same orbit as the emitter. Receiver 3 travels with a positive phase offset of 2.18 degrees in front of the emitter, while the latter is trailing the emitter at the same offset.

All orbital velocities and other parameters are identical to those described in section ??.

The secondary orbits are designated 1S,2S and 3S, and are arranged in a similar manner while having an orbital altitude of 525 km. Receivers 1S, 2S, 3S and 4S also correspond in their arrangement to their primary counterparts.

Parameters of the secondary orbits that are different to those of the primary are shown in table ?? on page ??.

Parameter	Symbol	Value	Unit
Altitude	h	525	[km]
Semi-major Axis	a	6896	[km]
Period (mins)	P	95.1298	[mins]
Revolutions per day		15.1372	[revs/day]
Angular Velocity	n	3.7843	[deg/min]
Circular Velocity	V	7.6127	[km/s]
Node Spacing		23.85	[deg]
Node Precession	$\dot{\Omega}$	-0.6584	[deg/day]

Table 5.5: Orbital properties of the secondary orbits (only the parameters different to those of the primary orbits are shown).

The initial RAANs are yet to be determined and a very complex model is required to account for differential precessions between the primary and secondary orbits. This is crucial in order for the secondary constellation to arrive in the right place with respect to the emitter. The calculations though fall out of the scope of this investigation.

### 5.2.2.2 Orbit Decay

The orbit decay analysis is done in a manner identical to that described in section ???. The final loaded mass of all the primary receiver satellites should be the same as the final orbits are established in order to decay at the same rate. The mean drag area estimated with the STK<sup>TM</sup> software is  $0.3 \text{ m}^2$ . The loaded mass is taken from section ???. The area exposed to sun (for solar pressure) is calculated to be  $0.25 \text{ m}^2$ . The ballistic coefficient then becomes  $22.40 \text{ kg/m}^2$ , which is extremely close to that calculated for the emitter satellite, certainly within the margins at this point in the design. This is key to maintaining the constellation. The final design of all the satellites should ensure they decay in absolutely the same manner.

The results of the decay simulation can be seen in figure ?? on page ???. They are virtually the same as the decay shown in figure ??.

The decay rate of secondary satellites is shown in figure ?? on page ???. The actual altitude was iterated in order to see that the receivers would decay to 500 km in mid August 2020. In reality the secondary satellite will have to take 0.3 kg of fuel less than the primary (at the expense of less fuel available for the deorbit maneuver) as the emitter satellite would have lost 27.8 m/s worth of propellant after the Hohmann transfer back to 500 km. This will ensure the same decay rate of the secondary leg of the mission.

### 5.2.2.3 Estimation of $\Delta V$

During the mission different receiver satellites will require different  $\Delta V$ . The main concern is to equalize them as the final orbits are established. The differences are only different due to different phase shifts needed and different orbital velocities.

The satellites will need to perform the following maneuvers:

- A phase shift maneuver similar to the one described in section ???. For this maneuver different drift rates are required for the satellites. Receivers 1 and 2 will need a  $\Delta V$  of 32.42 m/s (drift rate assumed to be half an orbit). Receiver 3 would require 31.58 m/s, but would require a whole orbit to establish its position. Receiver 4 would require a mere 0.85 m/s.

The secondary orbit satellites have a different circular velocity thus also have different  $\Delta V$ 's. Receivers 1S and 2S require 32.37 m/s. Receiver 3S needs 31.52 m/s, and finally, Receiver 4S requires 0.84 m/s.

- The deorbit maneuver will require the same  $\Delta V$  for all primary satellites and is also equivalent to the that of the emitter - 128.73 m/s. The secondary satellites will require 27.8 m/s less (100.93 m/s in total) as explained earlier.
- The correction for orbit insertion is identical for all primary satellites - 3.99 m/s (see section ??). Secondary satellites need a  $\Delta V$  of 3.98 m/s. This difference is neglible.

The total  $\Delta V$  for each receiver satellite is computed in table ?? on page ??.

	R1/R2	R3	R4	R1S/R2S	R3S	R4S
Phase shift	32.42	31.58	0.85	32.37	31.52	0.84
Deorbit	128.73	128.73	128.73	100.93	100.93	100.93
Inclination Correction	3.99	3.99	3.99	3.98	3.98	3.98
<b>Total</b>	165.14	164.3	133.57	137.28	136.43	105.75

Table 5.6: Total  $\Delta V$  for individual receiver satellites. Values are given in m/s.

The largest  $\Delta V$  (i.e 165.14 m/s) was used for propulsion system design.

Again no numbers are specified for stationkeeping. This topic is elaborated on in section ??.

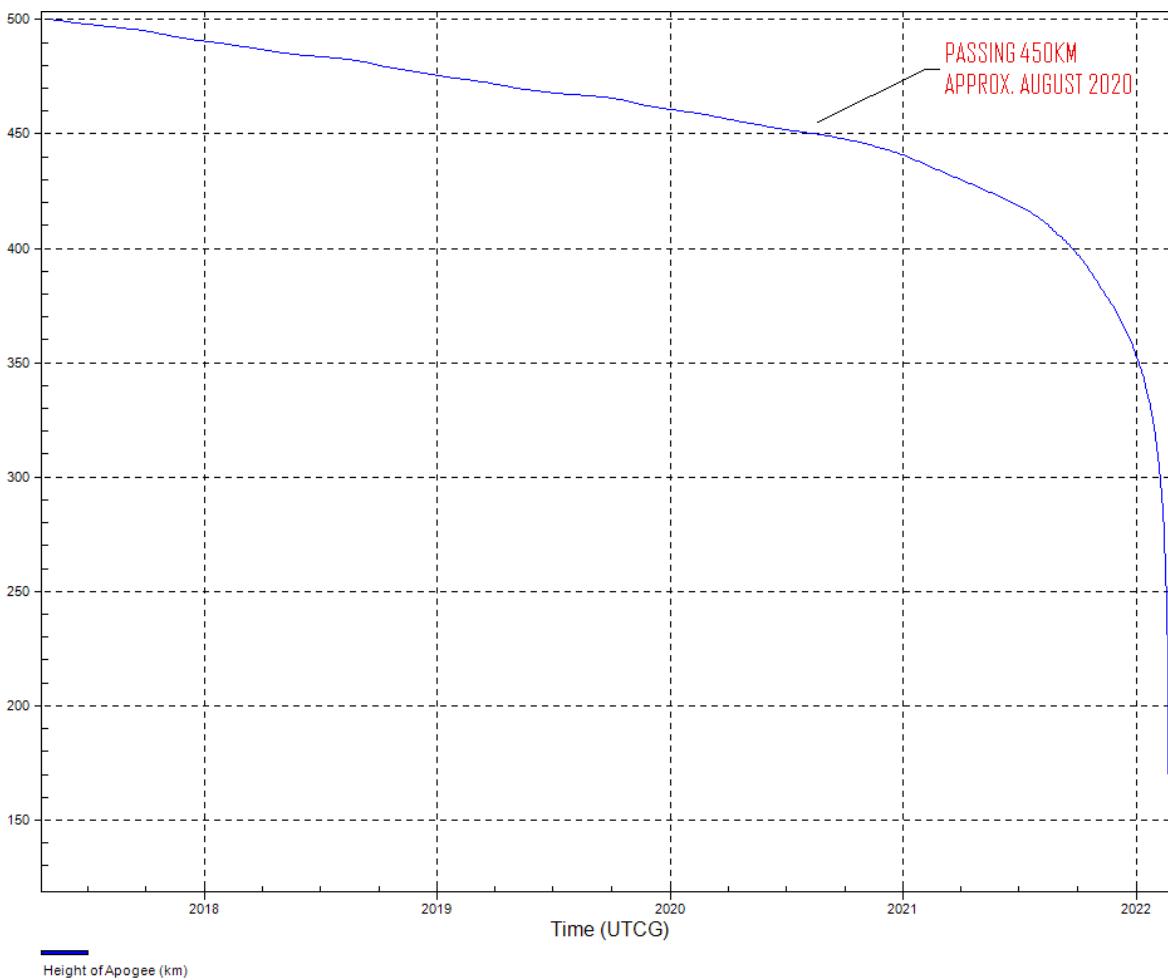


Figure 5.8: Emitter satellite orbit decay with an assumed mission start in March 2017. Values used:  $C_D = 2.22$ , mass = 53.1 kg, Area = 1.046 m<sup>2</sup>.

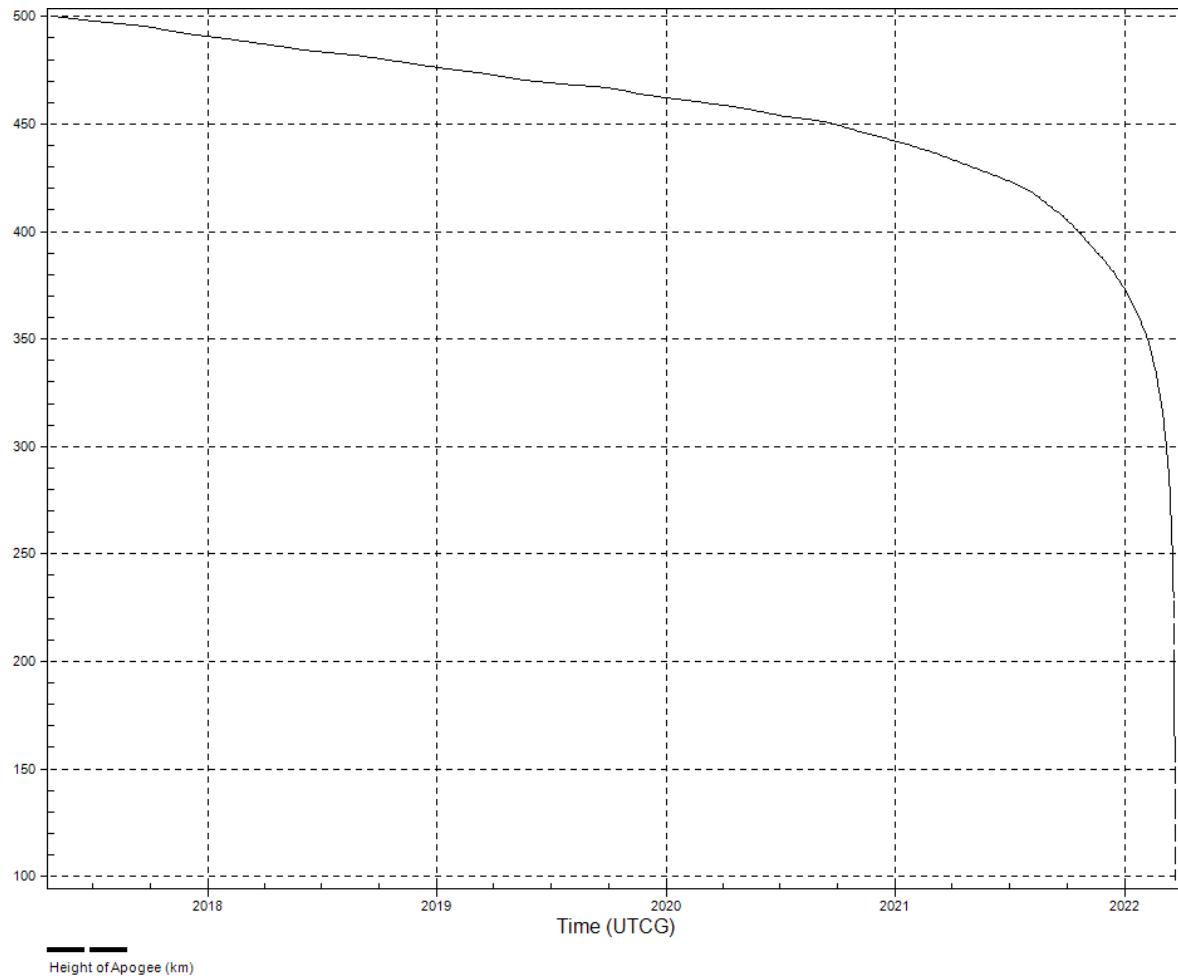


Figure 5.9: Receiver satellite orbit decay from a primary orbit with an assumed mission start in March 2017. Values used:  $C_D = 2.22$ , mass = 14.92 kg, Area =  $0.3 \text{ m}^2$ .

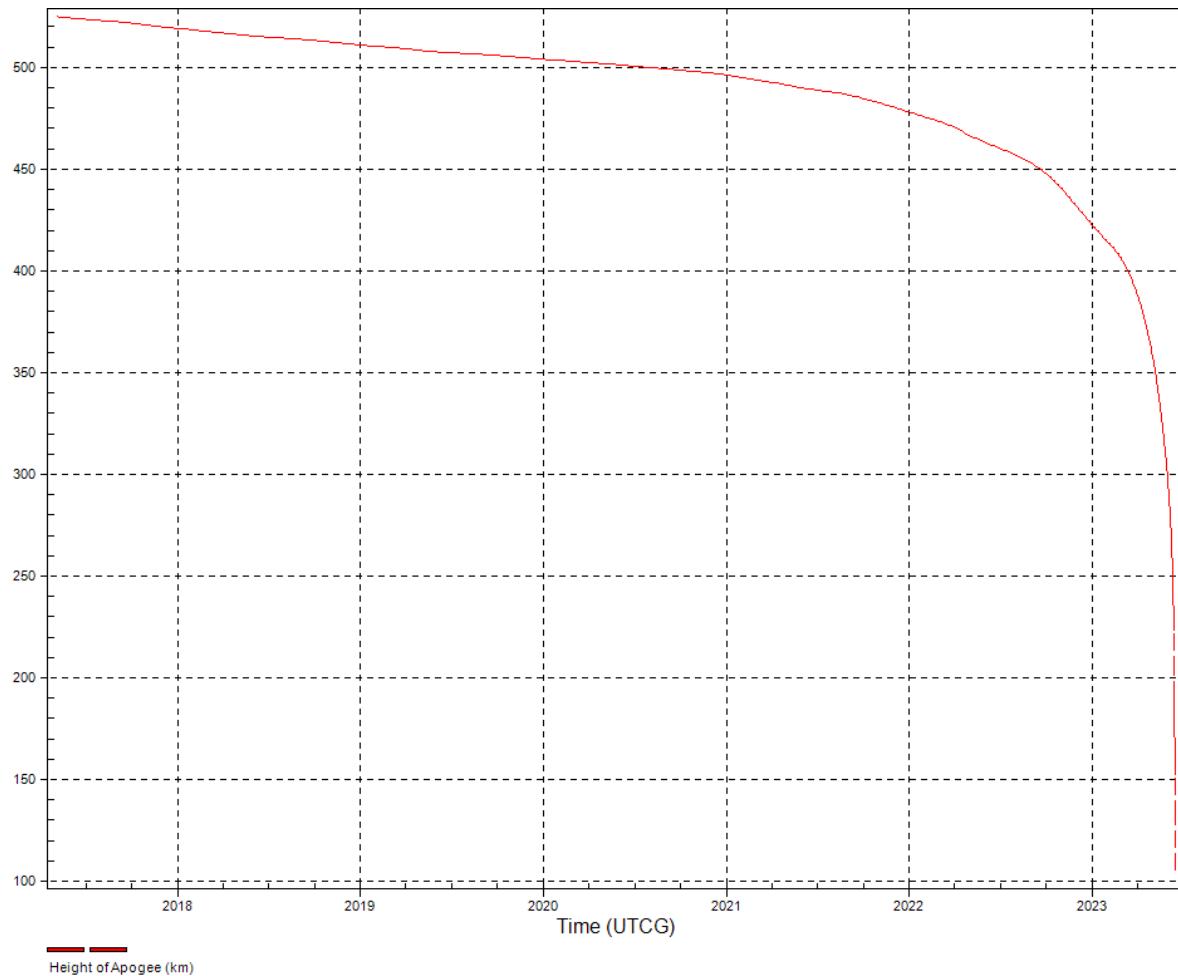


Figure 5.10: Receiver satellite orbit decay from a secondary orbit with an assumed mission start in March 2017. Values used:  $C_D = 2.22$ , mass = 14.92 kg, Area = 0.3 m<sup>2</sup>.

### 5.2.3 Swarm Configuration

The general configuration of all the satellites has been presented in the previous section and can be reviewed in figures ?? and ???. As explained, the general rule of separation is 2.18 degrees between the nodes of the different orbital planes and a 2.18 degree in-track phase offset for the satellites occupying the same orbit.

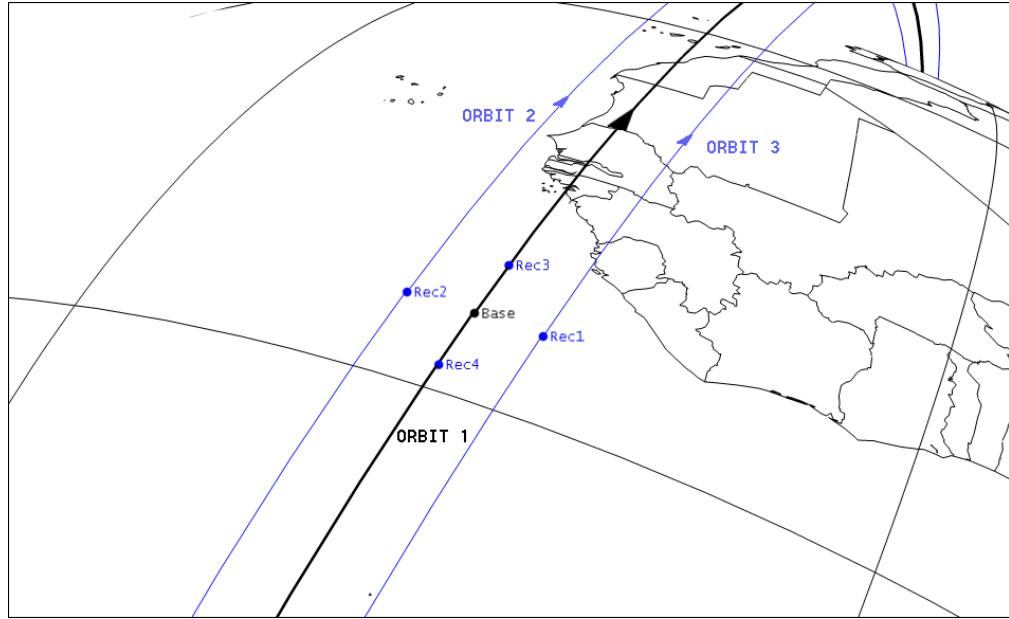


Figure 5.11: Swarm configuration as seen when the emitter (labeled here as Base) crosses its ascending node. Orbit numbers represent the different orbital planes.

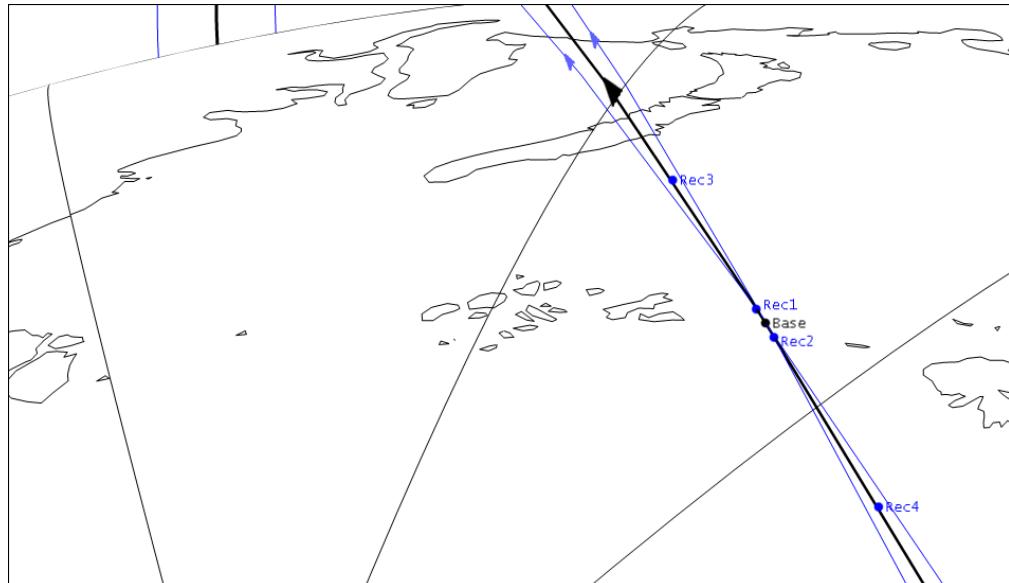


Figure 5.12: Swarm configuration as seen when the orbit planes intersect. In this figure the intersection in ascent is pictured.

A representative ground track can be seen in figure ?? on page ??.

From this general configuration several interesting parameters about the formation can be acquired. Please refer to figure ?? for a general configuration of geometry for constellations with the same inclinations.

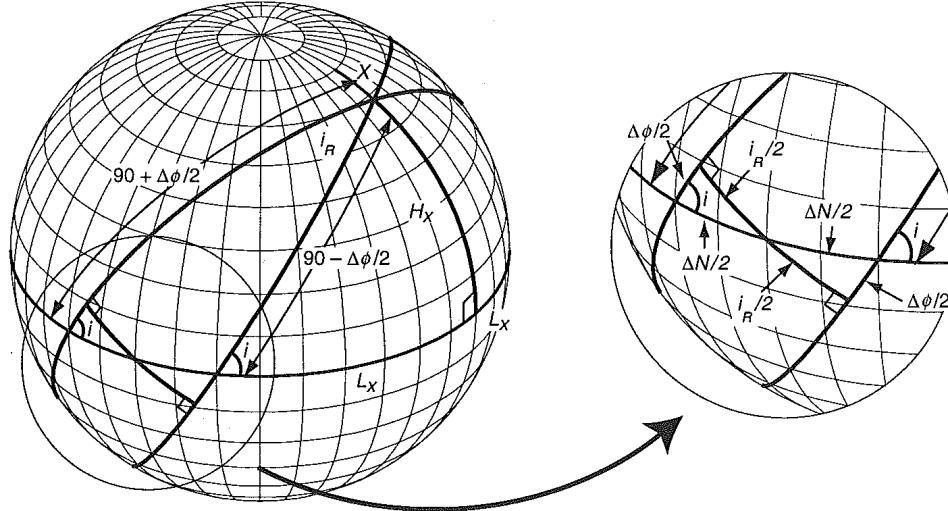


Figure 5.13: Relative geometry between two orbital planes with the same inclination. *Source: [?].*

The large-scale relative motion of two satellites in different orbits is governed by only two key variables: the relative inclination,  $i_R$  and relative phase,  $\phi_R$ . The relative inclination is the angle at which the two orbit planes intersect. The relative phase is the angle between the satellites when one intersects the other's orbit. The relative phase is the angular separation of the satellites at the time they intersect each other's orbit plane. This happens four times per orbit. The values of these angles are calculated using the following relations:

$$\cos i_R = \cos^2 i + \sin^2 i \cos \Delta N \quad (5.22)$$

$$\phi_R = (T_2 - T_1)n + \Delta\phi \quad (5.23)$$

where

$$\Delta\phi = 180 - 2\phi \quad (5.24)$$

$$\tan\phi = \frac{\tan(90 - \Delta N/2)}{\cos i} \quad (5.25)$$

$\Delta N$  is the angular separation at the ascending nodes. Using these equations the relative inclination between Orbit 1 and 2, and 3 and 1 is 2.17 degrees, this corresponds to a separation at the nodes of 261.78 km. The relative inclination between Orbit 2 and 3 is twice this value. The relative phase between the Base and Receiver 1 is 0.1901 (or 22.82 km) as Receiver 1 always crosses Orbit 1 before the emitter satellite. Receiver 2 has the same relative phase with the emitter and trails behind on intersections. The relative phase between satellites in Orbit 2 and 3 is 0.3803 degrees or around 44.6 km.

It is also possible to calculate the maximum and minimum separation angles,  $\lambda$ , using the following relations:

$$\sin\left(\frac{\lambda_{min}}{2}\right) = \sin\left(\frac{\phi_R}{2}\right)\cos\left(\frac{i_R}{2}\right) \quad (5.26)$$

$$\cos\left(\frac{\lambda_{max}}{2}\right) = \cos\left(\frac{\phi_R}{2}\right)\cos\left(\frac{i_R}{2}\right) \quad (5.27)$$

The maximum and minimum angular separation between any two closest orbital planes then becomes 2.1809 and 0.1901 degrees respectively. The nadir angle can also be calculated. This is the angle between the vector pointing to any other satellite and the Earth center vector. The relations are as following:

$$\sin(\eta_{min}) = \cos(\lambda_{max}/2) = \cos(\phi_R/2)\cos(i_R/2) \quad (5.28)$$

$$\cos(\eta_{max}) = \sin(\lambda_{min}/2) = \sin(\phi_R/2)\cos(i_R/2) \quad (5.29)$$

The actual values between cross-track satellites are then 89.91 and 88.91 degrees for maximum and minimum respectively. Along-track satellites have a constant nadir angle of 88.91 degrees.

All these values are enough to simulate along and cross-track relative motion between the various satellites. Figure ?? demonstrates the simulation of large scale relative motion between the Base, Receiver 1 and Receiver 2. Small scale motion is not simulated in this report and requires a more accurate analysis at a later point in time.

Finally, to confirm that all orbits intersect at different points, the azimuth and elevation of the intersection points are computed using the following equations:

$$\cos(\eta_{max}) = \sin(\lambda_{min}/2) = \sin(\phi_R/2)\cos(i_R/2) \quad (5.30)$$

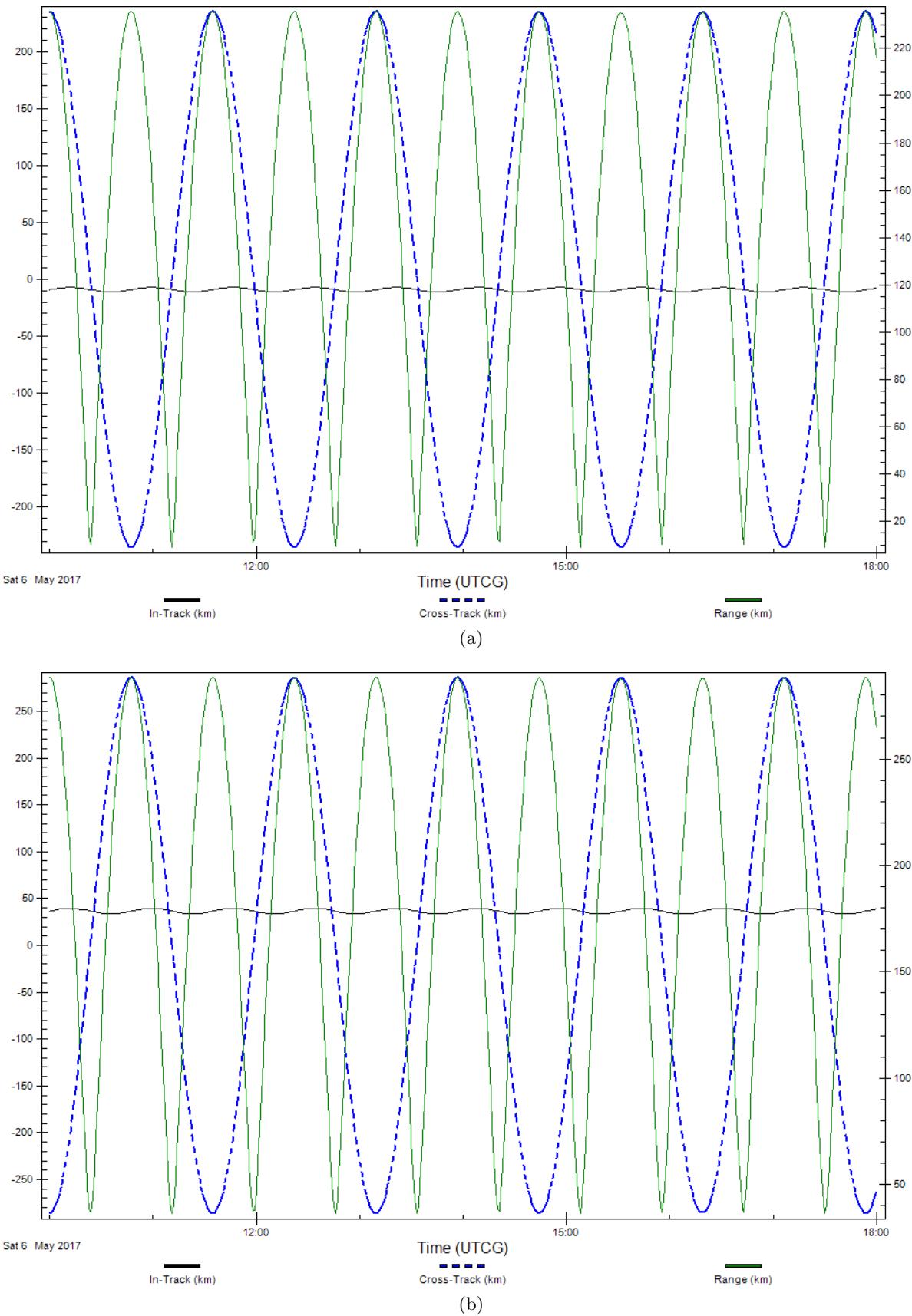


Figure 5.14: Large scale relative motion (a) of Receiver 1 w.r.t. Base and (b) of Receiver 2 w.r.t. Base.

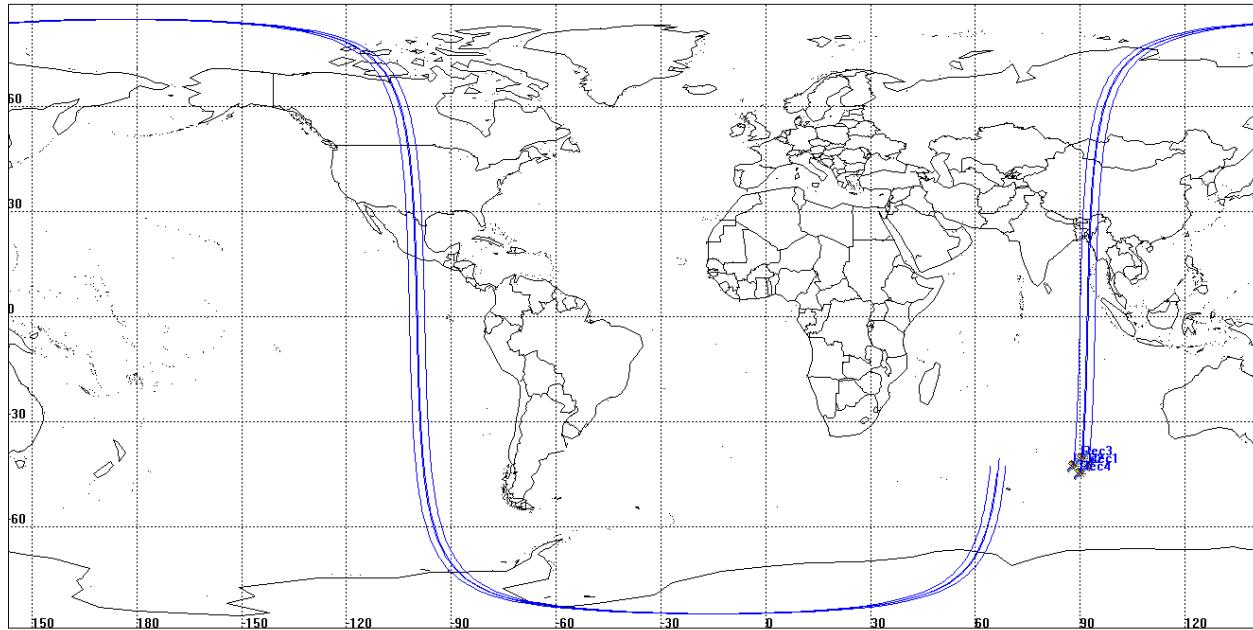


Figure 5.15: A representative ground track of the formation.

#### 5.2.4 Stationkeeping

As mentioned in section ?? the satellites will not require any  $\Delta V$  for stationkeeping maneuvers. This is because the concept of differential drag will be used to minimize propellant and to make the ballistic coefficients of the satellites more consistent.

The principle works in the following way: the satellites are maintained by following the slowest decaying one. This is known as relative stationkeeping. All the satellites are allowed to decay naturally in sync with that single reference. Since the orbits of every satellite are always known due to accurate navigation, it is possible to analyze decay rates. The satellites are then ordered to adjust the orientation of the solar arrays during eclipse in such a way that differential drag reinstates them into the right positions. This is further made possible due to the fact that not position acquisition but rather position knowledge (and thus correct instrument pointing) is of primary importance to data collection.

This concept has been successfully demonstrated by the ORBCOMM satellites [?]. It is understandable that the process of relative stationkeeping is more difficult to manage, however automated systems and prediction models can be designed to deal with this issue. Such systems however are not discussed here.

#### 5.2.5 Collision Avoidance

bla

### 5.3 Space Environment and Shielding

Bla

# Chapter 6

## Emitter Satellite

### 6.1 Detailed Design Optical Emitting Payload

**LiDAR!** is a remote sensing system comprising an optical emitting device, used to acquire topographic data, e.g. surface elevation gradients or ground composition by evaluating the **BRDF!**, considering multi-angular measurement are taken. For the generation of optical pulses, a highly efficient, diode-pumped, solid-state Nd-YAG **laser!** is considered **DPSSL! (DPSSL!)**. Solid-State **laser!**s (**laser!**s) have a high **TRL!** with relatively good properties in terms of beam quality (Q-factor), efficiency and pulse manipulation. Data products for topographical missions require that the **laser!** wave form be nearly pure Gaussian (known as transverse resonator mode  $TEM(00)$ ), both temporally and spatially, with a uniform wave front. The digitized time of flight waveform returning provides the topographic structure [?].

#### 6.1.1 Principles of *AlGaAs* Laser Diodes

**laser!** diodes are electrically pumped semiconductor **laser!**s, in which the gain is generated by an electrical current flowing through a *p-n junction* or (more frequently) a *p-i-n structure*[?]. In such a heterostructure, excitons dynamics can occur (electrons and holes can recombine), releasing the energy portions as photons. This process can be spontaneous, but can also be stimulated by incident photons, in effect leading to optical amplification. Most higher-power **laser!** diodes, however, exhibit a relatively poor beam quality, combined with other non-favorable properties, such as a large beam divergence, high asymmetry of beam radius and beam quality between two perpendicular directions, and astigmatism (property of rays to exhibit different foci in different symmetrical planes). Especially considering the long distances used in **LiDAR!** missions, these properties degrade the potential data quantity, as well as quality.

A quantum well is a thin layer which can confine (quasi-)particles (typically electrons or holes) in the dimension perpendicular to the layer surface, whereas the movement in the other dimensions is not restricted. A quantum well is often realized with a thin layer of a semiconductor medium, embedded between other semiconductor layers of wider bandgap. The thickness of such a quantum well is typically  $\sim 5 - 20$  [nm].

A major challenge is to reach the laser threshold, because the optical gain for the intracavity laser beam occurs only on a very small distance (in one or several quantum wells). It is therefore necessary to realize a laser resonator with very low losses, i.e., *Bragg mirrors* with high reflectivity. A Bragg mirror (also called *distributed Bragg reflector*) is a structure which consists of an alternating sequence of layers of two different optical materials. The principle of operation can be understood as follows. Each interface between the two materials contributes a Fresnel reflection. For the design wavelength, the optical path length difference

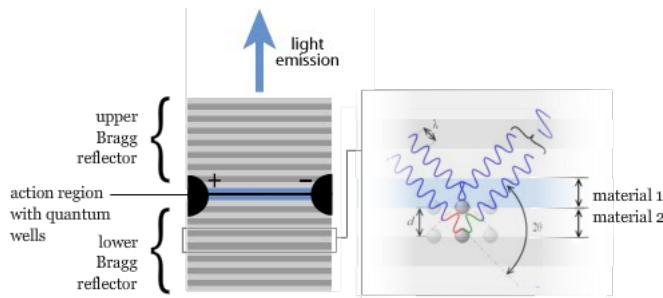


Figure 6.1: Basic diode laser configuration [?]

between reflections from subsequent interfaces is half the wavelength; in addition, the reflection coefficients for the interfaces have alternating signs. Therefore, all reflected components from the interfaces interfere constructively, which results in a strong reflection. The reflectivity achieved is determined by the number of layer pairs and by the refractive index contrast between the layer materials.

Individual **laser!** diodes normally generate quasi-continuous waves with powers  $\sim 1 - 10$  [mW]. To be able to generate higher power ( $\sim 1 - 10$  [W]) *laser diodes arrays* or *laser diode stacks* can be created, simply be combining multiple individual **laser!** diodes. High-power laser diode arrays (! (!)LDAs) are used for a variety of space-based remote sensor laser programs as an energy source for **DPSSL!**s. **LDA!**s (**LDA!**s) have been flown on NASA missions including MOLA, GLAS and MLA and have continued to be viewed as an important part of the **laser!**-based instrument component suite [?].

Laser diode bars have many single emitters arranged side-by-side and spaced approximately 0.5 [mm] apart, on a single slab of semiconductor material measuring approximately 0.5 [mm] x 10 [mm] in size. The individual emitters are connected in parallel which keeps the required voltage low at 2V, but increases the required current to 50 A/bar to 100 A/bar. Stacking these laser diode bars 2 to 20+ slabs high yields high power **LDA!**s capable of emitting several hundreds of Watts. Electrically, the bars are wired in series increasing the voltage by 2 V/bar while maintaining the total current at 50 A to 100 A. These arrays are one of the enabling technologies for efficient, high power solid-state lasers.

Traditionally these arrays are operated in QCW (Quasi Continuous Wave) mode with pulse widths of  $\sim 50$  s to  $200$  s and repetition rates of  $\sim 10$  Hz to  $200$  Hz. In QCW mode, the wavelength and the output power of the laser reaches steady-state, but the temperature does not. The advantage is a substantially higher output power than in CW mode, where the output power would be limited by the internal heating and the heat sinking properties of the device. The disadvantage is a much higher thermally induced mechanical stress caused by the constant heating and cooling cycle of the QCW operational mode, considering non-conductive cooling configurations.

Considering the fact that Nd:YAG is considered as gain medium (??), the existence of strong  $Nd^{3+}$  absorption near 808 [nm] permits efficient pumping with a GaAlAs (Gallium-Aluminium-Arsenide) diode **laser!**s for the  $F_{3/2} \rightarrow I_{9/2}$  transition. The direct band gap crystal AlGaAs is often used for laser diodes with wavelengths between 750nm and 880nm.  $Al_xGa(1-x)As$ , through changing the x, the ratio of the aluminum to gallium can be adjusted to vary the band gap and thereby control the wavelength. In the double heterostructure, stimulated emission occurs only within a thin active layer of GaAs, which is sandwiched between p- and n- doped AlGaAs layers that have a wider band gap. Laser diodes use heterojunctions to achieve simultaneous carrier and photon confinement in the active region.

A high laser efficiency demands that the light and injected charge carriers be confined as closely as possible to the same volume. The AlGaAs laser diode consists of a double heterojunction formed by an undoped (or lightly p-doped) active region surrounded by high bandgap p and n  $Al_xGa(1-x)As$  cladding

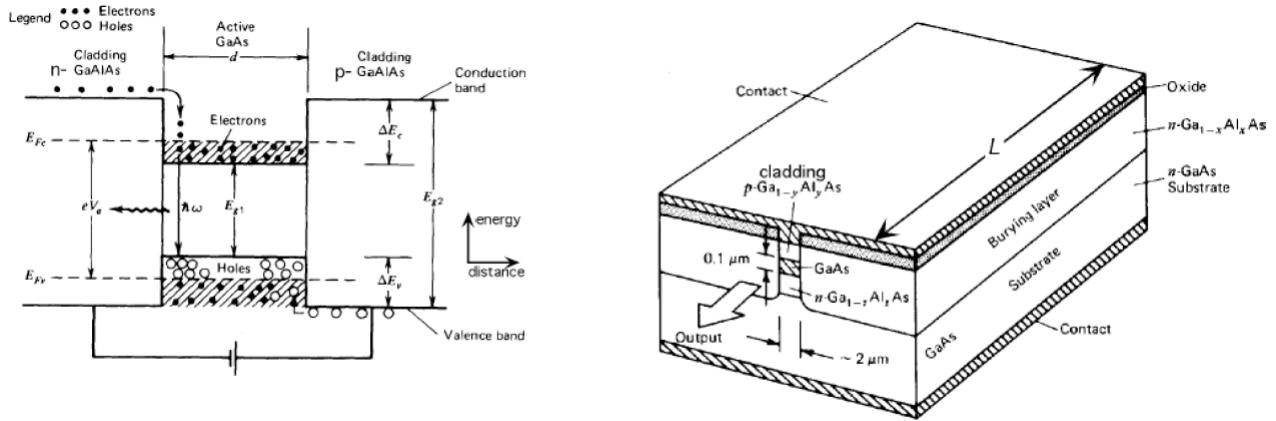


Figure 6.2: [?]

layers [?]. The surrounding cladding layers provide an energy barrier to confine carriers to the active region. The actual operation wavelengths may range from 750–880 [nm] due to the effects of dopants, the size of the active region, and the compositions of the active and cladding layers. When a certain parameter is fixed, the wavelength can vary in several (sub)nanometers due to other variables. For example, when the active layer has an energy gap  $E_g = 1.424[\text{eV}]$ , the nominal emission wavelength is  $\lambda = hc/E_g = 871[\text{nm}]$ . When a bias voltage is applied in the forward direction, electrons and holes are injected into the active layer. Since the band gap energy is larger in the cladding layers than in the active layer, the injected electrons and holes are prevented from diffusing across the junction by the potential barriers formed between the active layer and cladding layers. The electrons and holes confined to the active layer create a state of population inversion, allowing the amplification of light by stimulated emission.

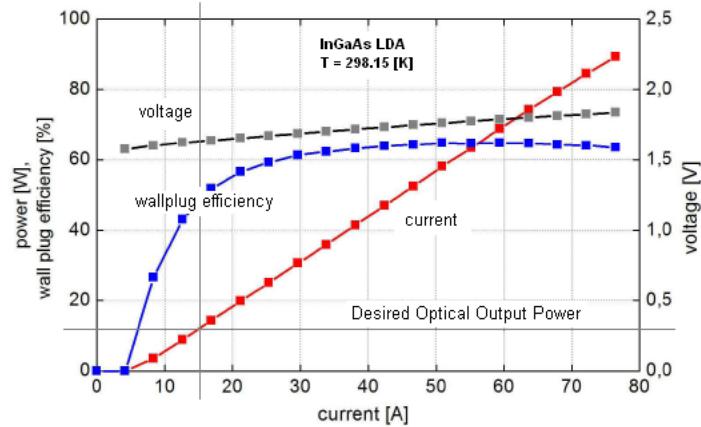


Figure 6.3: laser power [?]

Nd-YAG ( Neodymium Yttrium Aluminum Garnet $\text{Y}_3\text{Al}_5\text{O}_{12}$ )	
Nd Concentration	0.2 – 1.4 %
Diameter	0.5 – 15.0 mm
Length	1.0 – 220.0 mm
Damage Threshold	> 20 J/cm <sup>2</sup>
Refractive Index (n)	1.8169 @ 1,064 nm
Thermal	0.129 W/cm.K
Conductivity	
Specific Heat	0.59 J/g.K
Density	4.55 gm/cm <sup>3</sup>
Tensile Strength	280 MPa
Young's Modulus	282 GPa
dn/dT	+8.9 · 10 <sup>-6</sup> K <sup>-1</sup>

Figure 6.4: laser configuration

## 6.1.2 Diode Pumped Solid-State Laser Configuration

### 6.1.2.1 Nd-YAG Laser Characteristics

*Yttrium Aluminum Garnet* has emerged as the most widely produced laser gain host and has enjoyed recent popularity as a substrate material for optical components. The YAG host is a stable compound, mechanically robust, physically hard, optically isotropic, and transparent from below 300 to beyond 4,000 [nm]. YAG single crystals are able to accept trivalent laser activator ions from both the rare Earth and transition metal groups, and can be grown with very low strain.

For applications where  $TEM_{00}$  single mode operation is required, it is necessary to reduce or eliminate the variations in the bulk material and in the absorption of the pumping radiation throughout the component. In addition, wavefront distortions due to geometric imperfections and thermal gradient effects such as thermal lensing must be minimized. In this case, Neodymium concentration in the 0.4 to 0.8% range is typically specified.

### 6.1.2.2 Second Harmonic Generation

Since Nd-YAG has no principle absorption peak at the desired wavelength for the **LiDAR!** mission, the frequency should be altered from the original 946 [nm]. This can be done using *second harmonic generation* or *frequency doubling* in nonlinear crystals. The physical mechanism behind frequency doubling can be understood as follows. Due to the (2) nonlinearity, the fundamental (pump) wave generates a nonlinear polarization wave which oscillates with twice the fundamental frequency. According to Maxwell's equations, this nonlinear polarization wave radiates an electromagnetic field with this doubled frequency. Due to phase-matching issues, the generated second-harmonic field propagates dominantly in the direction of the nonlinear polarization wave. The latter also interacts with the fundamental wave, so that the pump wave can be attenuated (pump depletion) when the second-harmonic intensity develops: energy is transferred from the pump wave to the second-harmonic wave.  $\beta - \text{BaB}_2\text{O}_4$  is used for second, third, fourth and fifth harmonic generation of Nd doping **laser!**s. Typical dimensions of these crystal are  $\sim 0.05 - 10[\text{mm}]$ .

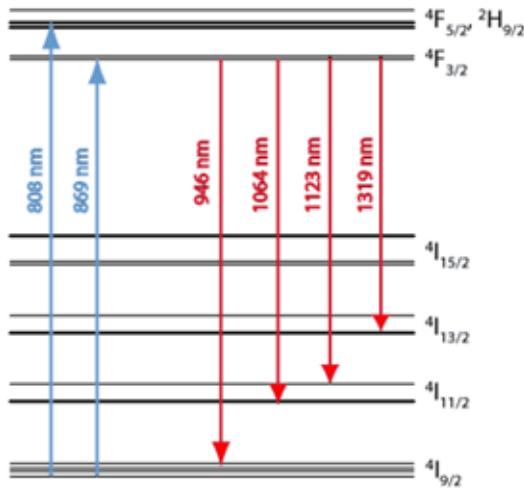


Figure 6.5: laser line

### 6.1.2.3 Pulse Generation

The generation and manipulation of pulses can highly influent the data in **LiDAR!** missions. To be able to transform the quasi-continuous wave into a pulsed wave, Q-switching is applied. Q-switching is a technique for obtaining energetic short pulses from a **laser!** by modulating the intracavity losses and thus the Q-factor (a measure of the damping of resonator modes) of the laser resonator. The technique is mainly applied for the generation of nanosecond pulses of high energy and peak power with solid-state bulk lasers. For *active Q-switching*, the losses are modulated with an active control element typically either an acousto-optic or electro-optic modulator. Both techniques rely on the fact that the optical properties within a nonlinear crystal change on the occurrence of an induced sound wave (acousto-optic) or electric field (electro-optic). There are also mechanical, less viable for space missions, Q-switches such as spinning mirrors, used as end mirrors of laser resonators. In any case, the achieved pulse energy and pulse duration depends on the energy stored in the gain medium, i.e. on the pump power and the pulse repetition rate. A Pockels cell is a device consisting of an electro-optic crystal including electrodes through which a electromagnetic beam can propagate. Dependent on the configuration, the phase delay or polarization state in the crystal (due to the *Pockels effect*) can be modulated by applying a flux electric voltage (typical second harmonic generation characteristics:  $\sim 40,000\text{V}/0.1[\text{mA}]$ ). Hence, for short periods ( $\text{dt}$ ) the polarization state of the incoming electromagnetic radiation can be altered. If a *polarizer disk* is used after the Pockel cell, the generation of pulses will begin, since the polarizer disk transmits certain polarized states only, deflecting the rest (acting like a 'polarize filter'). Pulses in the order of nanoseconds could be created this way. Care should be taken at the fact that the peak power after the Pockel cell is increased in several orders, due to the conversion from continuous to pulsed waves. Hence, the polarizer disk should be able to cope with these stresses. According to data posted after the GLAS-mission (using the same sort of **laser!**, an optical induced layer was formed at the non-linear crystal, probably induced by the high peak powers of the pulsed waves. In this configuration, the pulsed wave is created after the  $\beta - \text{BaB}_2\text{O}_4$  - crystal, reducing the risk of the creation of critical optical damage (COD).

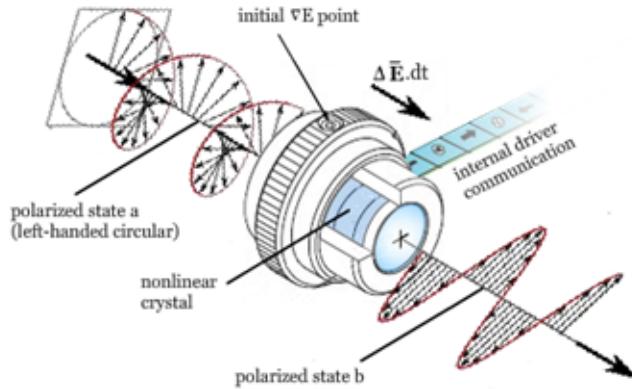


Figure 6.6: laser polarized

### 6.1.3 Optical Characteristics

Considering for a moment, the radius of the beam equals 400 [ $\mu\text{m}$ ]. Considering a  $TEM_{00}$  transverse mode of electromagnetic radiation this results in a area equal to

$$A_{beam} = \pi \cdot r^2 = \pi \cdot 400^2 = 502,654 \mu\text{m}^2 = 0.50265 \text{cm}^2 \quad (6.1)$$

The pulse energy  $E_p$  [J] (maximum optical power of a pulse) is determined using the simulator. Sufficient energy should be present within the electromagnetic radiation to ensure the optimum path from the transmitter towards the receiver. Lowering the value of  $E_p$  below this threshold energy can lead to atmospheric and surface absorption or translational mismatching due to incorrect scattering. The value of  $E_p$  of this particular mission is determined to be  $\sim 1$  [mJ] (see ??). Pulse repetition rate  $f_{rep}$  [Hz], i.e. the number of pulses emitted per second, is an important parameter for the altimetry mission. Again, using the results from the simulator, the quantity of this parameter can be determined to be  $\sim 5000$  [Hz] ( $\Delta t = 0.0002$  [s] with a pulse duration  $t_p \sim 10$  [ns]). Using the value of  $f_{rep}$ , the spatial resolution of the pulses along-track and in the nadir-direction can be calculated, considering the orbital velocity to be fixed at the determined altitude.

$$d_{along} = \Delta t \cdot v = 0.0002[\text{s}] \cdot 7,617[\text{m/s}] = 1.5234[\text{m}] \quad (6.2)$$

$$d_{nadir} = \Delta t \cdot c = 0.0002[\text{s}] \cdot 299,792,458[\text{m/s}] = 59,958.49[\text{m}] \quad (6.3)$$

Considering the value of  $E_p$  to be 1 [mJ] with a  $f_{rep}$  of 5,000 [Hz], the total power that should be induced within the electromagnetic wave can be calculated.

$$P_{output} = E_p \cdot f_{rep} = 0.001[\text{J}] \cdot 5,000[1/\text{s}] = 5.0[\text{W}] \quad (6.4)$$

The *total* electrical-to-optical power efficiency of a laser system, i.e. the *wall plug efficiency*, typically is  $\sim 10\%$ , however, linear interpolations of the current data, considering the large amount of research done on this subject, shows that the wall plug efficiency increases with one percent point every year (on average) from 2004, giving an wall plug efficiency of  $\sim 16\%$  in 2010 and  $>20\%$  in 2015 [?]. A higher wall plug efficiency reduces the electrical power consumption and also the amount of heat which has to be removed.

The pulse peak intensity equals  $E_p/t_p = 0.001[\text{J}]/10 \cdot 10^{-9}[\text{s}] = 100,000\text{W}$ . The intensity turns out to be  $\frac{E_p/t_p}{A_{beam}} = \frac{100,000[\text{W}]}{0.50265[\text{cm}^2]} = 198,950.6[\text{W/cm}^2](0.00199[\text{J}/\text{cm}^2]/10\text{ns})$ . The standard damage threshold energy

$E_{p,damage}$  for dielectric components equals  $0.5 - 10 [J/cm^2/10ns]$ . Considering the lowest value  $I_{p,damage}$ , hence,  $0.5 [J/cm^2/10ns]$ , and converting this to the appropriate dimensions, shows the intensity created within the electromagnetic pulses should do no harm to the dielectric components. Especially the polarizer disk (with the lowest  $I_{p,damage}$ ) is vulnerable for peak power caused by pulsed electromagnetic radiation.

#### 6.1.4 Gaussian Beam Propagation And Diffraction

Collimated plane wave propagation (uniform  $\mathbf{k}$ -vector distribution) in optical systems would give rise to discrete and accurate calculations. However, due to optical distortions and modifications, the  $\mathbf{k}$ -vector distribution can change, hence, altering the wave propagation.

*Gaussian Beams* For the analysis of the **laser!** beam intensity profile, a Gaussian profile (transverse resonator mode  $TEM_{00}$ ) is considered corrected by the  $M^2$  factor for optical distortion. The  $M^2$  factor is a common measure of the beam quality of a laser beam. The electric field distribution for a Gaussian beam is represented as:

$$E(r, z) = E_0 \cdot \frac{w_0}{w(z)} \cdot \exp \left[ \frac{-r^2}{w(z)} \right] \cdot \exp \left[ -i(kz - \arctan \left( \frac{z}{z_r} \right) + \frac{kr^2}{2R(z)}) \right] \quad (6.5)$$

$$E(x, y, z) = \exp \left[ \frac{-i(kz + \psi(z))}{w(z)} \right] \exp \left[ \frac{-(x^2 + y^2)}{w^2(z)} - ik \frac{(x^2 + y^2)}{2R(z)} \right] \quad (6.6)$$

The main point of this section is to determine the Gaussian beam propagations dependency on diffraction phenomenon. To characterize the Gaussian beam in more details, the following equations are used to be able to describe the propagation.

$$\theta = \frac{\lambda}{\pi \cdot w_0} \quad (6.7)$$

$$w_R(z) = w_{0R} \sqrt{\left[ 1 + \left( \frac{z(\theta + \Delta\alpha)M^2}{w_{0R}} \right)^2 \right]} \quad (6.8)$$

$$w_R(z) = w_{0R}(z) \sqrt{\left[ 1 + \left( \frac{z\lambda M^2}{\pi w_{0R}^2} \right)^2 \right]} \quad (6.9)$$

$$R_R(z) = z \left[ 1 + \left( \frac{w_{0R}}{z\theta M^2} \right)^2 \right] \quad (6.10)$$

$$R_R(z) = z \left[ 1 + \left( \frac{w_{0R}}{z(\theta + \Delta\alpha)M^2} \right)^2 \right] \quad (6.11)$$

*Fraunhofer Diffraction* Diffraction is a fundamental characteristic of all wave fields. The effect of diffraction is typically manifested when an obstacle is placed in the path of a beam. On an observation screen some distance away from the obstacle, one observes a rather complicated modulation of the time-average intensity in the vicinity of the boundary separating the illuminated region from the geometrical shadow cast by the

obstacle. With the use of high-power lasers, diffraction of radiation beams (cavity oscillating in the fundamental transverse Gaussian  $TEM_{00}$  mode) with finite transverse dimensions has significant consequences. The Fresnel number  $F = a^2/\lambda \cdot R$ , where  $a$  is the characteristic size ("radius") of the aperture,  $\lambda$  is the wavelength, and  $R$  is the distance from the aperture, determines the diffraction regime that should be considered (Fjii 1: Fraunhofer (far-field);  $F > 1$ , Fresnel). The far-field light field is the Fourier transform of the aperatured field. The far-field light field is the Fourier transform of the aperatured field.

$$E(k_x, k_y) = \mathcal{F} \left\{ \underbrace{t(x, y)}_{\text{Transmission function}} \cdot E(x, y) \right\} = \iint (\exp(-i(k_x x + k_y y)) \cdot t(x, y) \cdot E(x, y) dx dy \quad (6.12)$$

$$E(x, y, z) = \frac{\exp[-i(kz + \psi(z))]}{w(z)} \cdot \exp \left[ \frac{-(x^2 + y^2)}{w^2(z)} - \frac{ik(x^2 + y^2)}{2R(z)} \right] \quad (6.13)$$

The lens incorporates a phase delay to the outgoing electromagnetic field.

$$t_{lens} = \exp \left\{ -i \left( (n-1) \left( \frac{k}{2R} (x^2 + y^2) \right) \right) \right\} \quad (6.14)$$

Combining the above calculations, calculations for the Fraunhofer diffraction can be conducted, which shows the dependency on divergence.

$$\mathcal{F} \left\{ \left( \exp \left\{ -i \left( (n-1) \left( \frac{k}{2R(z)} (x^2 + y^2) \right) \right) \right\} \right) \otimes \left( \frac{\exp[-i(kz + \psi(z))]}{w(z)} \cdot \exp \left[ \frac{-(x^2 + y^2)}{w^2(z)} - \frac{ik(x^2 + y^2)}{2R(z)} \right] \right) \right\} \quad (6.15)$$

A different point of view, conveniently in the sense of the **LiDAR!** mission, considers the use of focal lengths to change the Gaussian beam diffraction, giving the same result as the above Fourier transform, i.e. the divergences influence the intensity profile.

$$E(x_1, y_1) = \iint \left[ \exp \left( ik \left( \frac{-2xx_1 - 2yy_1}{2z} + \frac{x^2 + y^2}{2z} \cdot t_{lens}(x, y) \cdot E(x, y) \right) \right) \right] dx dy \quad (6.16)$$

$$\frac{k}{2z} = (n-1) \frac{k}{2R_1} \quad (6.17)$$

$$\frac{1}{f} = (n-1) \left[ \frac{1}{R_1} - \frac{1}{R_2} \right] \quad (6.18)$$

$$E(x_1, y_1) = \iint \exp \left[ -i \frac{k}{f} (xx_1 + yy_1) \cdot t(x, y) \cdot E(x, y) \right] dx dy \quad (6.19)$$

### 6.1.5 Thermal Control

Basically, there are three critical parts (**LDA!**, Nd:YAG **laser!** crystal and optical components after polarizer disk) of the **laser!** configuration in terms of thermal control. All of these components shall be considered in this subsection.

**LDA!** The constituent parts and materials of a typical **LDA!** are the diode die (laser bar) and the mechanical structure. The packaging design and materials enable the array of laser bars to stay together in a stack, to be energized electrically (with a relatively high drive current), to pass the heat generated out of the unit to the mounting surface (thermal path, heat sinking), to be sufficiently rugged against mechanical insults, to provide a standard mounting interface (screws or clamps) and to be as small as possible. The active region of the **LDA!**, where heat is generated, is only about 1 micron wide, located about 3 microns from the P-side of the bar. The bars are about 0.1 [mm] wide and typically spaced about 0.5 [mm] from each other. Waste energy in the form of heat must be conductively transferred into the solder material and from there into the heat sink material (typically BeO or CuW) as rapidly as possible. The solder material of choice is a soft Indium alloy for its ductile property allowing the bar and the heat sink to expand or contract at different rate with temperature. The **LDA!** manufacturers try to use materials which possess higher thermal conductivity and a relatively comparable coefficient of thermal expansion (CTE) in order to minimize the thermal resistance of the device and the induced mechanical stresses. Additionally important to reducing mechanical stress is consideration of the use of soft solders which are highly pliable with a relatively low melting point ( $\tilde{160}$  C). Post life test analysis indicates that solder deformation caused solder roll-over, in turn creating voids, which increase thermal resistance. When coupled with built-in stress due to fabrication, such roll over, in time often obstructs emitters, leading to increased heating, or extends across the bar from anode to cathode causing bar shorts which eventually result in contaminations to the emitter face and localized hot spots, further degrading performance. Excessive heating and thermal cycling of the **LDA!** active regions plays a key role in limiting the reliability and lifetime of **LDA!**s operated in the QCW mode, particularly where pulse widths are long. To improve the assembly's heat extraction performance, advanced materials are being considered for packaging **LDA!**s, which have high thermal conductivity and a CTE (Coefficient of Thermal Expansion) that matches that of the laser bars to compensate inhomogeneous thermal strain.

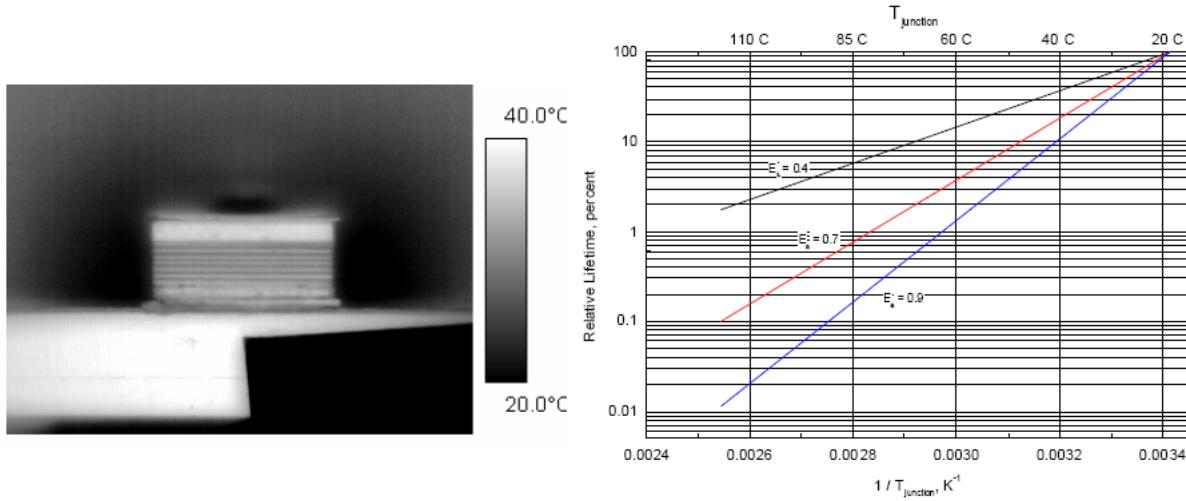


Figure 6.7: [?]

**Nd:YAG! slab.** The Nd:YAG slab plays a central role in the **laser!** configuration. To be able to cope with thermal stresses induced by the wave formation, the slab is thermally bonded to a molybdenum copper block in order to match the CTE.

*Dielectrical Component Temperature Dependency* Optical misalignment is a serious issue with the **laser!** configuration, both in the manufacturing phase, as well during the actual mission. Refractive indices of dielectric components alter the beam translation and should be considered. Given the fact that the refractive index is a material parameter with a direct dependency of temperature (temperature-dependent Sellmeier equation), beam propagation can change unwillingly during mission.

### 6.1.6 Laser Lifetime Expectance

Multiple aspects influence the expected lifetime of the optical emitter device, such as power, temperature interval, repetition rate and intracavity properties. Since most **laser!**s have a non-continuous mode of operation (i.e. the duty factor is lower than 100%), reliability data for long-term cycles are not abundant available.

For damage-free operation in a harsh, hands-off, environment such as space, a major form of damage risk reduction is the creation of a large single intracavity mode to reduce peak fluence. Since resonator efficiency depends strongly on the inversion density of the gain medium, it is advantageous to confine the desired cavity mode as close as possible. To accomplish this, the 808 [nm] light from the diode arrays should be collimated by a single plano-convex cylindrical lens (for maximal efficiency, made of undoped YAG). By doing this, the probability of the existence of thermal lensing is reduced, increasing the beam quality and the lifetime.

Considering a constant value of  $f_{rep}$  of 5,000 [Hz], the total number of pulses equals  $788.4 \cdot 10^9$  [pulses/5 years]. All optical components should be able to cope with the large amount of pulses and the peak power implied by these pulses, i.e. the energy damage threshold of the dielectric components should be higher than the incoming energy of the electromagnetic radiation. Since  $I_{p,damage}$ ) is given with a temporal resolution in the order of a single pulse width ( $\sim 10$  [ns]), individual pulses can be analyzed. Stationary calculations can be conducted with the information based on the electromagnetic radiation energy and hence, the proper optical elements could be chosen ( $I_{p,damage} > I_p$ ).

[?] shows an experimental set-up, where the lifetime of a **DPSSL!** is investigated, using approximately the same **laser!** configuration with  $f_{rep} = 242$  [Hz] and  $E_p = 0.0150$  [J]. The pulse energy is much larger than the value of  $E_p$  in the case of the **LiDAR!** mission described in this report ( $\sim 0.001$  [J]). ?? shows the results. After  $2.4 \cdot 10^9$  shots, there was no damage found in any of the cavity optics, but inspection of the diodes revealed that a single bar was lost on one array. After the first year, the pump pulse length was increased from 89 [ $\mu\text{m}$ ] to 105 [ $\mu\text{m}$ ] to restore the output energy to 15 [mJ]. This roughly simulated the procedure that would be performed in space in order to maintain an altimetry link. The final result was that after more than  $4.8 \cdot 10^9$  10 - 15 [mJ] laser pulses, there was no optical damage present in the system [?]. This clearly indicates that the **LDA!**s lifetime considerations are important for the entire **laser!** system. AlGaAs lasers can suffer from catastrophic optical damage (COD), rapid degradation, and gradual degradation. These phenomena are due to darkline defect propagation and a high surface recombination rate [?].

[?] shows an experimental set-up, where the lifetime of a **DPSSL!** is investigated, using approximately the same **laser!** configuration with  $f_{rep} = 242$  [Hz] and  $E_p = 0.0150$  [J]. The pulse energy is much larger than the value of  $E_p$  in the case of the **LiDAR!** mission described in this report ( $\sim 0.001$  [J]). ?? shows the results. After  $2.4 \cdot 10^9$  shots, there was no damage found in any of the cavity optics, but inspection of the diodes revealed that a single bar was lost on one array. After the first year, the pump pulse length was increased from 89 [ $\mu\text{m}$ ] to 105 [ $\mu\text{m}$ ] to restore the output energy to 15 [mJ]. This roughly simulated the procedure that would be performed in space in order to maintain an altimetry link. The final result was that after more than  $4.8 \cdot 10^9$  10 - 15 [mJ] laser pulses, there was no optical damage present in the system [?]. This clearly indicates that the **LDA!**s lifetime considerations are important for the entire **laser!** system.

Taken into account the fact that the total number of shots in five years exceed the number of total shots delivered by a single **LDA!** without considerable loss in power and beam quality, the obvious consequence is that multiple **LDA!**s should be implemented within the structure. Given the fact that individual **laser!** diode has dimensions  $\sim 0.01$  [m], multiple diodes could be added to form a **LDA!** matrix.

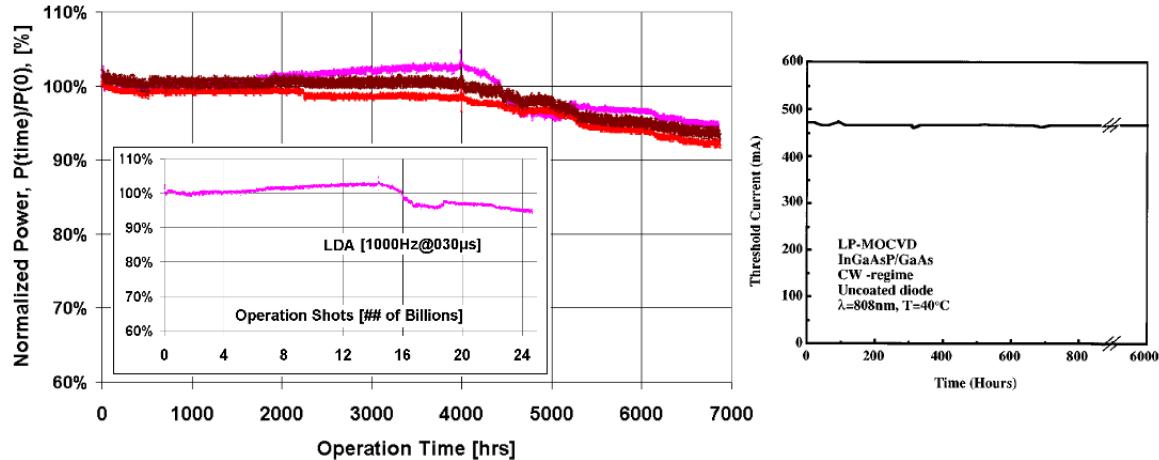


Figure 6.8

### 6.1.7 laser! Focus Calculation

The figure ?? on page ?? gives a overview of the emitter optics. In order to diverge or focus the laser beam, it is possible to move the parabolic mirror up or down from the exact focus position. In this case, the divergent angle  $\gamma$  needs to be calculated, which can be verified or optimized later on to obtain the desired footprint size. The calculation drawing is shown in figure ?? on page ??.

In the figure ??,  $p_1$  is the parabolic mirror positioned at the exact focus point  $f$ , and  $p$  is the distance between  $f$  and origin.  $p_2$  is the parabolic mirror with the exact same shape but which is moved away from focus point with distance  $\xi$ .  $L_{in}$  indicates the incoming light.  $L_{out1}$  is the outgoing light due to  $p_1$  and  $L_{out2}$  is the outgoing light due to  $p_2$ . Meanwhile,  $r_1(x_1, y_1)$ ,  $r_2(x_2, y_2)$  are the reflected points due to  $p_1$  and  $p_2$ . The purpose of this focusing calculation is to find the divergent angle  $\gamma$  with respect to the design parameters  $p$ ,  $\xi$  and reflection point  $r_1(x_1, y_1)$ . [?]Parabolic mirror  $p_1$  has the equation ??, and  $p_2$  has the equation ??.

$$y = -\frac{1}{4p}x^2 \quad (6.20)$$

$$y = -\frac{1}{4p}x^2 + \xi \quad (6.21)$$

The equation ?? for incoming light line  $L_{in}$  can be obtained since  $r_1(x_1, y_1)$  is known in this case.

$$y = \frac{y_1 + p}{x_1}x - p \quad (6.22)$$

Insert equation ?? into equation ??,  $x_2$  of  $r_2(x_2, y_2)$  can be obtained as:

$$x_2 = \frac{-\frac{y_1+p}{x_1} + \sqrt{\frac{y_1+p}{x_1}^2 - \frac{\xi-p}{p}}}{\frac{1}{2p}} \quad (6.23)$$

Next step is to find the tangent line of  $p_2$  at r2:

$$\left(\frac{dy}{dx}\right)_{x_2} = -\frac{1}{2p}x_2 = \tan(\mu_3) \Rightarrow \mu_3 = \arctan\left(-\frac{1}{2p}x_2\right) \quad (6.24)$$

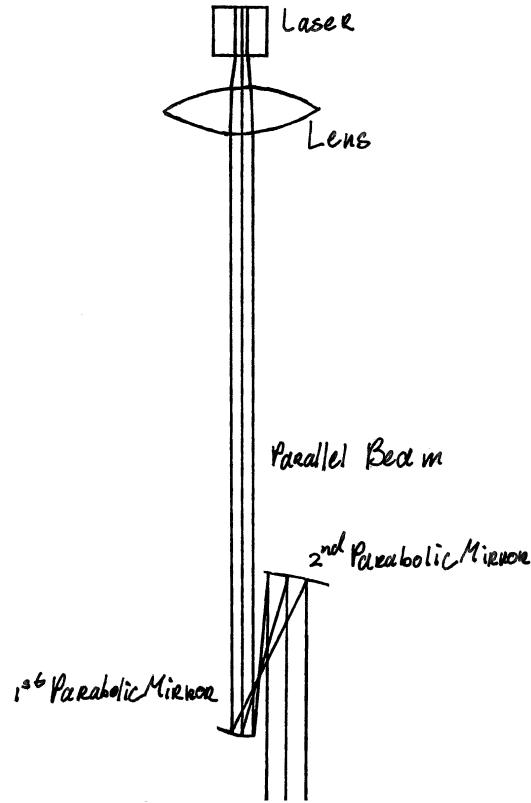


Figure 6.9: Emitter optics drawing

In the figure, 't' is the tangent line at point  $r_2$ , and 'N' is the normal line perpendicular to the tangent line. The normal line 'N' is also the angle bisect, and  $\mu_2$  is a half of the reflecting angle. From the drawing, these relations can be found:

$$\mu_1 + \mu_2 + \mu_3 = 90^\circ = \mu_1 + \mu_2 + \mu_4 \Rightarrow \gamma = \mu_2 - \mu_4 = \mu_2 - \mu_3 \quad (6.25)$$

To find  $\mu_2$ ,  $\mu_1$  need to be calculated.  $\mu_1$  is the tangent angle of  $L_{in}$  at  $r_1$  or  $r_2$ :

$$\left(\frac{dy}{dx}\right)_{x_1, y_1} = \frac{y_1 + p}{x_1} = \tan(\mu_1) \Rightarrow \mu_1 = \arctan\left(\frac{y_1 + p}{x_1}\right) \Rightarrow \mu_2 = 90^\circ - \mu_1 - \mu_3 \quad (6.26)$$

Insert value of  $\mu_2$  and  $\mu_3$  to equation ??, so the divergent angle  $\gamma = f(p, \xi, r_1(x_1, y_1))$  is obtained. Put these equations into Excel, and it is much easier to see how is  $\gamma$  verified. For instance, give values for  $p=350[\text{mm}]$ ,  $\xi = 5[\text{mm}]$  and  $x_1=5[\text{mm}]$ ,  $\gamma=0.01169[\text{deg}]$ , which will give the footprint size of 102 meters. By adjusting the  $\xi$ , the mirror has a divergence of 20.4[m/mm] for the same  $p$  and  $x_1$ .

## 6.2 Navigation

For navigation the receiver satellites are considered first to find the smallest required size of the **GPS!** receiver. However the system chosen there has an accuracy that is high enough to warrant its usage onboard the emitter satellite as well.

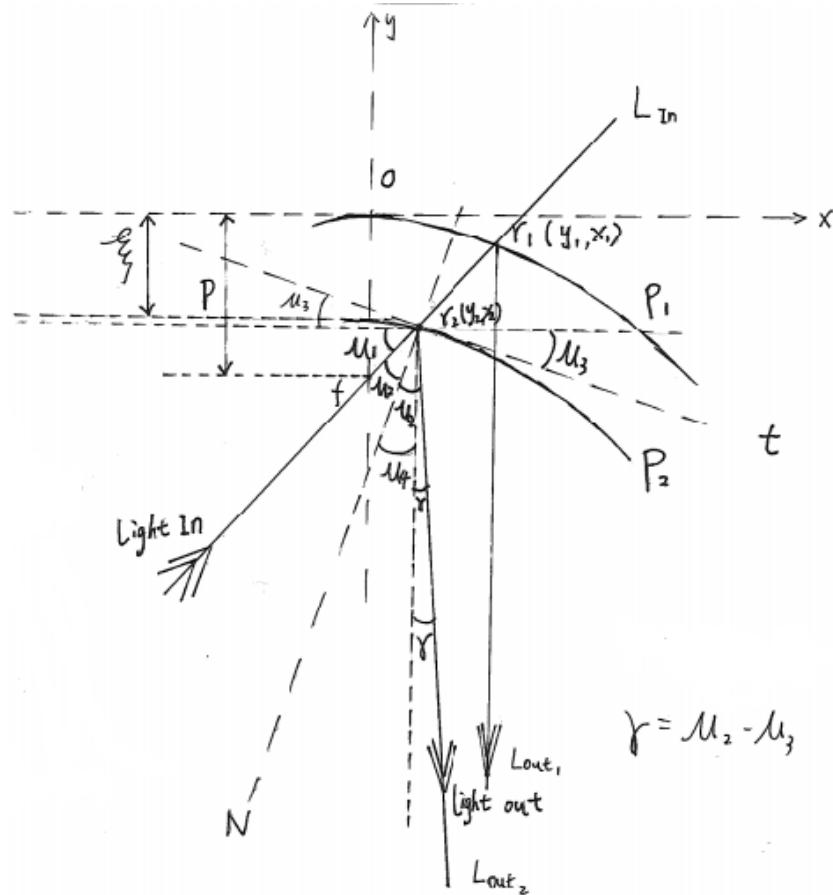


Figure 6.10: Focus calculation draft

For convenience the chosen system is repeated here. The chosen system is a **GPS!** receiver developed by SpaceQuest called the GPS-12-V1, it has a real-time accuracy of at least ten meters and a post-processing accuracy up to several millimeters. The **GPS!** antenna used is one developed by **SSTL!** (**SSTL!**) and is a SGR Patch Antenna ASY-00741-04. For further reference please go to ?? on page ??.

### 6.3 Communication subsystems

This section will discuss the three different communication links to which the emitter satellite is connected:

- Crosslink for scientific and housekeeping data
- Ground-space link for scientific data
- Ground-space link for command and housekeeping data

These links are illustrated figure XXX For each link the link budget and the communications hardware selected will be presented.

### 6.3.1 Crosslink for scientific data and housekeeping data

This link transmits the scientific data gathered by the receiver satellites to the emitter satellite. To keep the receiver satellites as small as possible and to have efficient data storage, the receiver satellites will transmit their scientific data continuously so that the emitter satellite can store all data in one data bank before it is transmitted to the ground. In addition to scientific data, command and housekeeping data is also transmitted in this link.

The link budget has the following input parameters: the frequency selected was 2 GHz, the data rate required is 1.62 Mbit/s, the maximum distance between the satellites is 261 km, the modulation used is QPSK. Atmospheric losses were not considered for obvious reasons.

In the mid term report a frequency in the Ku-band was selected, however it turned out that the data rate was low enough to use S-band or more specifically 2 GHz, which requires far less power consuming hardware.

The maximum distance is 261 km, which is longest distance that occurs between the emitter satellite and any of the receiver satellites. Distance between receiver satellites can be twice as long but were not considered after the relative tracking with these crosslinks was abandoned.

The QPSK was selected for its good balance between required  $E_b/N_0$  and spectrum utilization. Also a lot of transceivers are available for this modulation.

Another important parameter is the data rate, 1.62 Mbit/s, which was calculated in appendix YYY. The position of the satellite is registered every second (144 bits), while for each received photon the coordinates on the array, time, satellite attitude and instrument attitude are registered (295 bits), with 5000 photons expected per second. Adding an extra 10 % for overhead gives the final data rate between a single receiver satellite and the emitter satellite.

#### 6.3.1.1 Transceiver

Since the data rate is quite high for an S-band, nanosatellite S-band transceivers did not satisfy the requirements, which only allow 9.6 kb/s at maximum [?]. However microsatellite sized transceivers turned out to be sufficient. The transceiver selected was an SBTRcvr of RDLabs, which has a maximum data rate of 10 Mb/s for QPSK modulation, an input power of 12 Watt and an output power of 1 to 5 Watt [?].

#### 6.3.1.2 Antenna

For low gain applications and frequencies below 4 GHz three antennas were considered: a dipole antenna, a helix antenna and a patch antenna. The dipole antenna was not an option since its gain is too low (theoretical gain of a half wave antenna is only 2.15 dBi). A helix antenna would be too large and would decrease the ballistic coefficient too much, for example the S-Band quadrifilar helix antenna of Surrey Satellite technology [?] is 0.5 m long with a base of 100mm on 100mm, while the S Band patch antenna [?] from the same company is only 20mm thick and 82mm on 82mm large. This S Band patch antenna, which was eventually selected, has a mass of only 80 grams, a 120° beamwidth and a gain of 4 dBi. 5 antennas will be placed on the emitter satellite: 1 one nadir pointing and 4 facing the receiver satellites.

#### 6.3.1.3 Link budget

Because the selected modulation is QPSK and the maximum allowable bit error rate is  $10^{-5}$ , the required  $E_b/N_0$  is equal to 9.6 dB. With the output power of the transceiver at maximum (5 Watt or 7 dBW) it is possible to have a  $E_b/N_0$  of 10.35 dB, which leaves a margin of 2.36 dB.

### 6.3.2 Ground-space link for scientific data

As mentioned in the past subsection, all scientific data is collected on the emitter satellite which is then transmitted to the ground when it passes its ground station.

The link budget has the following input parameters: the frequency selected was 8.2 GHz, the data rate allowed is 150 Mbit/s, the maximum distance between the ground station and the satellite is 1000 km, the modulation used is SQPSK.

The frequency selected is of the X-band as was mentioned in the mid-term review, since it allows a high data rate and is commonly used for Earth observation satellites and thus a lot of ground stations have the right equipment for this frequency band.

The data rate depends on the link time between ground station and satellite, the time between the passes and the storage capacity. The exact required data rate was determined in section XXX on data storage and resulted in 110 Mb/s, nevertheless the 150 Mb/s figure was used in the calculations for the linkbudget to allow for more margin.

The distance between the station and the satellite follows from the maximum beam angle of antenna on the satellite ( $45^\circ$ ), for a detailed calculation of this distance, see appendix YYY.

The SQPSK modulation was chosen since it also allows a good balance between required  $E_b/N_0$  and spectrum utilization and also allows a high data rate than QPSK on certain transmitters.

#### 6.3.2.1 Transmitter

The transmitters considered were the XTRA-6 from TESAT Spacecom [?] and the HRT150 from General Dynamics and [?]. Since the XTRA-6 is lighter, smaller and requires less power and allows a higher data rate than the HRT150, the choice was quickly made. The XTRA-6 weighs 1.1 kg, measures 197x89x74 mm, consumes 30 Watts of power and has an output power of 6 Watts.

#### 6.3.2.2 Antenna

As was chosen in the mid term review, the antenna for the ground space link is a phased array. The phased array selected was the XPAA from Boeing's Phantom works in Seattle [?]. It has a mass of 5.5 kg, measures 330x305x74mm, has a gain of 23.03 dBi and a beamwidth of  $15^\circ$ .

#### 6.3.2.3 Ground station

The ground station selected was ESRANGE, Kiruna, which allows up to 12 passes a day as it is located above the pole circle. The ground station is operated by ESA and the SSE and is commonly used by Earth observation satellites. For X-band the base has a 13m parabolic receiver dish which has a very high gain, 58 dBi and a beamwidth of  $0.18^\circ$  [?].

#### 6.3.2.4 Link budget

Similar to the link budget of the crosslinks, the selected modulation is SQPSK and the maximum allowable bit error rate is  $10^{-5}$ , meaning the required  $E_b/N_0$  is equal to 9.6 dB. With the output power of the transceiver at 5 Watt or 7 dBW, it is possible to have a  $E_b/N_0$  of 10.35 dB, which leaves a margin of 29.3 dB.

### 6.3.3 Ground-space link for command and housekeeping data

Aside from the one way space to ground link for scientific data there is also a separate two way link for commands and housekeeping data.

The link budget has the following input parameters: the frequency selected was 2 GHz, the data rate allowed is 20 kb/s, the maximum distance between the ground station and the satellite is 1000 km, the modulation used is QPSK.

The frequency selected is 2 GHz since only a very low data rate is required, and since most ground station over the world are equipped with hardware for S-band links this allows for many contact opportunities to send commands or receive housekeeping data in addition to its main ground station at ESRANGE.

The data rate required is 20 kb/s, normally a housekeeping data is never more than a few kb/s[?] but since the housekeeping data of 5 satellites has to go through the link, it was estimated that 20 kb/s should suffice.

The maximum distance is the same as for the space to ground link for scientific data.

The modulation is again QPSK for its balance between required  $E_b/N_0$  and spectrum utilization.

#### 6.3.3.1 Transceiver

The transceiver selected was the same as the one used for the crosslinks, the SBTRcvr of RD Labs, as the data required is still too high for nanosatellite transceivers.

#### 6.3.3.2 Antenna

Also the S-band patch antenna from Surrey Satellite technologies was again selected for its small size and sufficient gain.

#### 6.3.3.3 Ground station

The ground station selected was again ESRANGE, Kiruna, for the reasons as the ground-space link for scientific data. For S-band the base has several receiver antennas from 2.4m to 13m, but the link budget showed a 2.4m parabolic reflector dish is already more than sufficient. The gain of this antenna is 53 dBi and has a beamwidth of  $30^{\text{circ}}$  [?].

#### 6.3.3.4 Link budget

For this last link budget, the selected modulation is QPSK and the maximum allowable bit error rate is  $10^{-5}$ , meaning the required  $E_b/N_0$  is equal to 9.6 dB. With the output power of the transceiver at 1 Watt or 0 dBW, it is possible to have a  $E_b/N_0$  of 39.25 dB, which leaves a margin of 27.65 dB.

### 6.3.4 Data Storage for the Emitter

In order to find an appropriate storage device it is important to know how much data will have to be stored. In order to find this number a ground station is chosen and it is determined when the satellite is

Parameter	Emitter
Max. time without contact to ground station	7:35:33
Average time without contact to ground station	1:39:00
Average duration of contact to ground station	0:08:30
Total bit rate [Mbit/s]	8.13
Max. amount of data to be stored [Gbit]	244
Required downlink rate [Mbit/s]	111
Maximum available downlink rate [Mbit/s]	150

Table 6.1: Important values used to determine the required memory on the emitter.

in view. Many European polar orbit Earth observation missions use the Kiruna station located close to the North Pole, so the laser swarm will use this station as well. Simulating a period of three weeks will give an indication of when the emitter is in view, the amount of time the satellite is in view and the time in between two passes. The first three entries in table ?? show the results for these calculations.

Next the total generated bit rate is determined, which is 8.13 [Mbit/s] as indicated in table ???. Using this number and the maximum time between two passes over Kiruna the maximum amount of storage required can be determined. The result is indicated in table ??, note that a 10% margin is included to take into account possible anomalies and housekeeping data.

Now that the required amount of storage is known, a suitable storage device can be chosen. The 64 [Gbit] flash memory module from 3D-Plus ?? is the best choice due its ability to store a large amount of data in a small, space qualified, package. One module required about 1 Watt of power and has a weight of 6.10 grams, also it's dimensions are 20.4 x 13.84 x 12.13 mm. Because 244 [Gbit] is required the emitter will be equipped with 5 of these modules. The reason 5 modules are used when 4 would do is to allow for a redundancy in data storage, because it allows more data to be stored should or if one of the other modules breaks down.

Now a suitable storage medium is chosen it is checked whether the stored data can be sent to Earth without running out of storage capacity for new measurements. For the simulated data it is observed that while most of the intervals between contact are about 1:40 hours, every 13 or more orbits the ground station is not visible for 6 or 7 hours. Taking into account the new data received during these overpasses and the average time the emitter is visible, the resulting required downlink rate is calculated to be 111 [Mbit/s] as indicated in table ???. Comparing this with the maximum possible downlink rate it is revealed that the 3D-Plus 64 [Gbit] memory is a viable option. So the emitter will carry 5 x 64 [Gbit] flash memory modules from 3D-Plus for data storage.

## 6.4 AODCS!

To be able to perform its mission the attitude and orbit of the emitter satellite need to be accurately determined and controlled. The **AODCS!** (**AODCS!**) of the satellite can be split up into a number of components, which are all described separately in this section. Section ?? covers the attitude determination, section ?? describes the attitude control, the orbit determination is threaded in section ?? and section ?? concerns the orbit control. Furthermore section ?? is about the pointing mechanism for the emitter. Section ?? gives an overview of the mass, costs and power for the complete **AODCS!** and pointing mechanism.

### 6.4.1 Attitude Determination

In the Midterm Report [?] a combination of sun sensors, for the day (Sunlit) phase, and a star tracker, for the night (eclipse) phase, was chosen for the attitude determination of the satellites. A total of five sun sensors is used, one on the top of the satellite and one on each of the side walls. Since the satellite is nadir pointing no sun sensor on the bottom of the satellite is required. The star tracker is located on the top of the satellite, so it will always be pointed towards space.

The sun sensors used are the **CoSS!** (**CoSS!**) from TNO [?], the left image in figure ?? on page ?? shows the instrument. The mass of the instrument is 24 grams and it has a volume of  $30 \times 30 \times 14.5 \text{ mm}^3$ . It has a full cone view angle of  $160^\circ$  and an accuracy of about 3 arcsec. The cost of one of these sun sensors is 11000 EUR, including all documentation and cables. Since the sensors are passive no external power is needed. It has to be noted for the processing of the data that analogue sun sensors in **LEO!** are sensitive to the influence of Earth albedo.

There are currently still few star trackers available for micro-satellites because of their complexity and consequential size. Aeroastro has developed a star tracker with a mass of 375 grams and dimensions of  $60 \times 76.2 \times 76.2 \text{ mm}^3$  (excluding a baffle) shown in the right image in figure ?? on page ?? . The instrument has a field of view of  $24 \times 30$  degrees. It uses less than 2 Watts (1 Watt nominal) of power. The attitude determination accuracy is about 90 arcsec. Up to 9 stars can be tracked at a time. **ISIS!** (**ISIS!**) [?] is working on a similar star tracker. Their dimensions are  $50 \times 50 \times 100 \text{ mm}^3$ . The instrument is planned to be ready for flight mid 2011 and will cost about 75000 EUR.

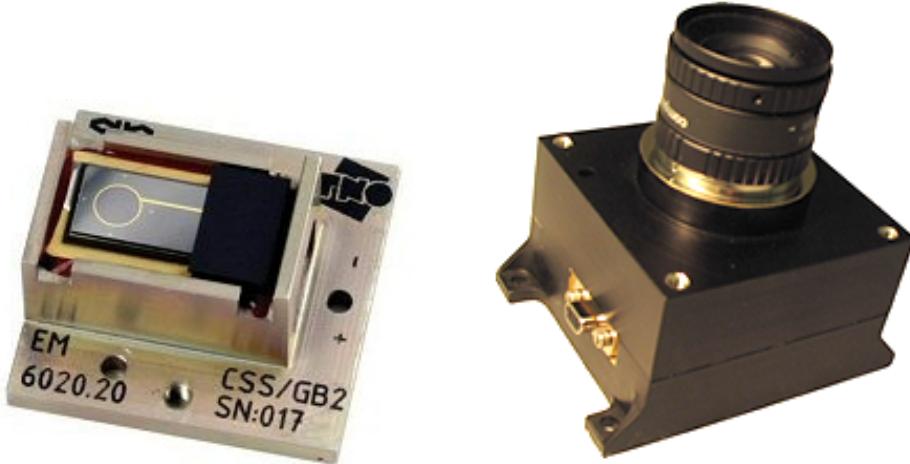


Figure 6.11: The TNO Cosine Sun Sensor [?] and the Aeroastro Miniature Star Tracker [?]

### 6.4.2 Attitude Control

The attitude control of the satellites in the Laser Swarm is done by reaction wheels for manoeuvring and magnetic torquers for desaturation (spinning down) of the wheels.

The size of the reaction wheels is such that the system is able to counter all disturbance torques that work on the satellite. The main disturbance torques are caused by the Earth's gravity gradient  $T_g$ (equation ??), Solar radiation  $T_{sp}$  (equation ??), the Earth magnetic field  $T_m$  (equation ??) and aerodynamics  $T_a$  (equation ??) [?].

$$T_g = \frac{3\mu}{2R^3} |I_z - I_y| \sin 2\theta \quad (6.27)$$

$$T_{sp} = \frac{F_s}{c} A_s (1+q) \cos i (c_{ps} - cg) \quad (6.28)$$

$$T_m = DB \quad (6.29)$$

$$T_a = \frac{1}{2} \rho C_d A V^2 (c_{pa} - cg) \quad (6.30)$$

where  $\mu$  is the Earth's gravitational constant  $398600.4418 \text{ km}^3/\text{s}^2$ ,  $R$  is the radius of the orbit,  $I$  is the satellite inertia tensor,  $\theta$  is the deviation from nadir,  $F_s$  is the Solar constant  $1367 \text{ W/m}^2$ ,  $c$  is the speed of light  $299792458 \text{ m/s}$ ,  $q$  is the reflectance factor (between 0-1, typically 0.6),  $i$  is the angle of incidence of the Sun,  $c_{ps}$  is the center of solar pressure,  $cg$  is the center of gravity,  $D$  is the residual dipole of the satellite in  $A \cdot m^2$ ,  $B$  can be approximated as  $2M/R^3$  in Tesla, where  $M$  is the magnetic moment of the Earth  $7.96 T \cdot 10^{15} \text{ m}^3$ ,  $\rho$  is the local density in  $\text{kg/m}^3$ ,  $C_d$  is the drag coefficient of the satellite,  $A$  is the surface area in  $\text{m}^2$ ,  $V$  is the satellite velocity and  $c_{pa}$  is the center of aerodynamic pressure.

Filling out equations ?? to ?? with realistic numbers for the emitter satellite results in

$$\begin{aligned} T_g &= \frac{3 \cdot 398600.4 \cdot 10^9}{2 \cdot 6878000^3} |6.542 - 2.407| \sin(2 \cdot 1^\circ) = 2.652 \cdot 10^{-7} \text{ Nm} \\ T_{sp} &= \frac{1367}{3 \cdot 10^8} 1.58 (1 + 0.6) \cos 0^\circ (0.2) = 2.304 \cdot 10^{-6} \text{ Nm} \\ T_m &= 4.893 \cdot 10^{-5} \cdot 1 = 4.893 \cdot 10^{-5} \text{ Nm} \\ T_a &= \frac{1}{2} 1.80 \cdot 10^{-12} \cdot 2.2 \cdot 1.58 \cdot 7613^2 (0.2) = 3.626 \cdot 10^{-5} \text{ Nm} \end{aligned}$$

Using the values and from SMAD [?]. The moments or inertia of the satellite are  $I_x = 2.407 \text{ kgm}^2$ ,  $I_y = 5.582 \text{ kgm}^2$  and  $I_z = 6.542 \text{ kgm}^2$  derived from the model. The maximum total disturbance torque expected on the satellite is the sum of all the above torques, i.e.  $8.776 \cdot 10^{-5} \text{ Nm}$ . For redundancy a margin of 2 is standard, so the torque the reaction wheels have to be able to produce is  $1.755 \cdot 10^{-4} \text{ Nm}$ .

The motors chosen for the reaction wheels are the Faulhauber 2209 Brushless DC-micromotors. The mass of the motor is 8.5 grams, the dimensions  $22 \times 22 \times 17.5 \text{ mm}^3$  and the maximal rotation speed is 10000 rpm [?]. They cost 176 EUR a piece. The maximum angular acceleration  $\dot{\omega}_{max}$  they can perform is  $1.03 \cdot 10^3 \text{ rad/s}^2$ .

Figure ?? on page ?? shows the general layout of a general reaction wheel suitable for the motor. Wheel basically consists of two integrated parts, a top ring and a side skirt. The top ring has a standard thickness of 1, the height  $h$  and width  $b$  of the skirt can be adapted to fit the requirements.

The torque a reaction wheel can produce is determined by

$$T_w = I_w \dot{\omega}_w \quad (6.31)$$

where  $I_w$  is the inertia tensor of the wheel and  $\dot{\omega}_w$  the angular acceleration. The inertia of the wheel depends on the dimensions. The basic build-up of the wheels is a disk with a skirt around the motor giving the mass at a distance from the rotation axis as shown in figure ?? on page ?? . The thickness  $b$  of the skirt is 8 mm for the yaw wheel and 5 mm for the other two, since the yaw wheel is also used for the instrument pointing increasing the torque requirements. Since the wheel is spinning it is required that the material of the wheel does not produce a magnetic field while rotating, therefore aluminium is chosen for the material. It has a density of  $2850 \text{ kg/m}^3$ . The mass of the wheel is

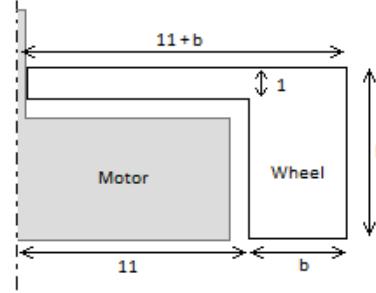


Figure 6.12: The general layout of a one of the reaction wheels. The motor is depicted in grey, the wheel in white.

$$m_w = m_{wd} + m_{ws} = \rho\pi (11 \cdot 10^{-3} + b)^2 1 \cdot 10^{-3} + \rho\pi ((11 \cdot 10^{-3} + h)^2 - (11 \cdot 10^{-3})^2) (b - 1 \cdot 10^{-3})^2 \quad (6.32)$$

where  $m_{wd}$  is the mass of the disk,  $m_{ws}$  is the mass of the skirt,  $b$  is the thickness of the skirt and  $h$  is the height. The moment of inertia of the wheel can then be calculated as

$$I_w = m_{wd} \cdot \left( \frac{11 \cdot 10^{-3} + b}{2} \right)^2 + m_{ws} \cdot \left( 11 \cdot 10^{-3} + \frac{b}{2} \right)^2 \quad (6.33)$$

By taking an angular acceleration of 100 rad/s and taking a wheel height  $h$  of 1 cm the wheel skirt thickness  $b$  of 36 mm is required. The mass is 212 grams per wheel. The inertia of the motor is not considered.

The magnetic torquer chosen for the emitter is the MT2-1, it can deliver a dipole moment of  $2.0 \text{ Am}^2$ , a mass of 0.2 kg a length of 157.5 mm and a diameter of 15 mm [?]. They use a linear power of 0.5 Watt. Three are needed in orthogonal planes to be able to desaturate all wheels.

### 6.4.3 Orbit Determination

The orbit determination will be done using the navigation subsystem, which is described in section ??.

### 6.4.4 Orbit Control

In section ?? the  $\Delta V$  required for the emitter satellite is 221 m/s. Because the main manoeuvre of the satellite is the boost to a higher orbit a bipropellant thruster, an EADS' 10 N Bipropellant Thruster Model S10 - 23 is chosen [?]. This kind of thruster has already flown on over 90 satellites. The specific impulse  $I_{sp}$  delivered by the thruster is 291s. The mass of the dual seat valve model is 650 grams, the general dimensions are  $90.3 \times 37.4 \times 178.5 \text{ mm}^3$ . The throat diameter is 2.85 mm and the diameter of the exit is 35 mm. By filling out the equation for the fuel mass over dry mass ratio

$$\frac{m_p}{m_0} = 1 - e^{-\Delta V / (I_{sp}g)} \quad (6.34)$$

the amount of fuel needed is 7.5% of the dry mass of the satellite needs to be fuel. For the 50.4 kg satellite the fuel mass is therefore 3.8 kg. The thruster works on a 1:1.5 mixture dinitrogen tetroxide  $N_2O_4$  and monomethylhydrazine  $MMH$ . This leads to fuel masses of 1.5 kg of  $N_2O_4$  and 2.3 kg of  $MMH$ , with the typical bulk densities noted in [?] this leads to volumes of 1.32 and 2.88 litres respectively.

#### 6.4.5 Pointing Mechanism

Since the **laser!** and satellite will both always be nadir pointing no additional pointing mechanism for the **laser!** is needed. The receiver on the emitter satellite also does not require a pointing mechanism since the ground target will always be on same horizontal point under the satellite. Because the satellite is moving in its orbit during the travel time of the signal the receiver instrument possibly could have to look at a small angle slightly behind the satellite to receive a reflected signal from the ground. The pulse travels two times 500 km, from orbit to ground and vice versa, in 3.34 ms. In that time the satellite covers 25.4 meters. Therefore the instrument needs to point back 0.003 degrees to always have the optimal signal in the middle of the receiver. Since the pointing accuracy required is only 0.1 degree there is no need for a back looking receiver instrument.

#### 6.4.6 Overview

In table ?? on page ?? an overview is given of all masses, costs and power requirements of the **ADCS!**, **ODCS!** (**ODCS!**) and the pointing mechanism. Since the basic structure of the satellite will not be able to carry all loads of the subsystems some extra mass is added for supporting the components ensuring that correct pointing of sensors, instruments and actuators.

Subsystem component	Number	Mass [g]	Cost [EUR]	Total mass [g]	Total cost [EUR]	Power (max) [W]
Sun sensors	5	24	11000	120	55000	0 (0)
Star trackers	1	375	75000	375	75000	1 (2)
Reaction wheel motors	3	8.5	176	25.5	528	0.15 (1.5)
Reaction wheel disk	3	36	10*	108	30*	0 (0)
Magneto torquers	3	200	3000	600	9000	1.5 (3)
Thruster	1	650	400k	650	400k	1 (2)
Fuel tanks, pumps	2	500	250*	1000	500*	0 (0)
Support structures	-	-	-	770	300*	0 (0)
Total				3650	540k 310k FY00\$	3.65 (4.4 rms)

Table 6.2: Dry mass, cost and power budget of different parts of the **ADCS!** and pointing mechanism. Starred prices are estimations based on material and machining cost.

# Chapter 7

## Receiver Satellite

### 7.1 Detailed Design Optical Receiving Payload

The figure ?? on page ?? gives an overview diagram of the receiver optics. More information about the design concept or trade off can be found in the midterm report. In this detailed design report, the wavelength filter system and the single photon detection device **SPAD!** research will be the pivot, after that the receiver payload summary and budget will be drawn.

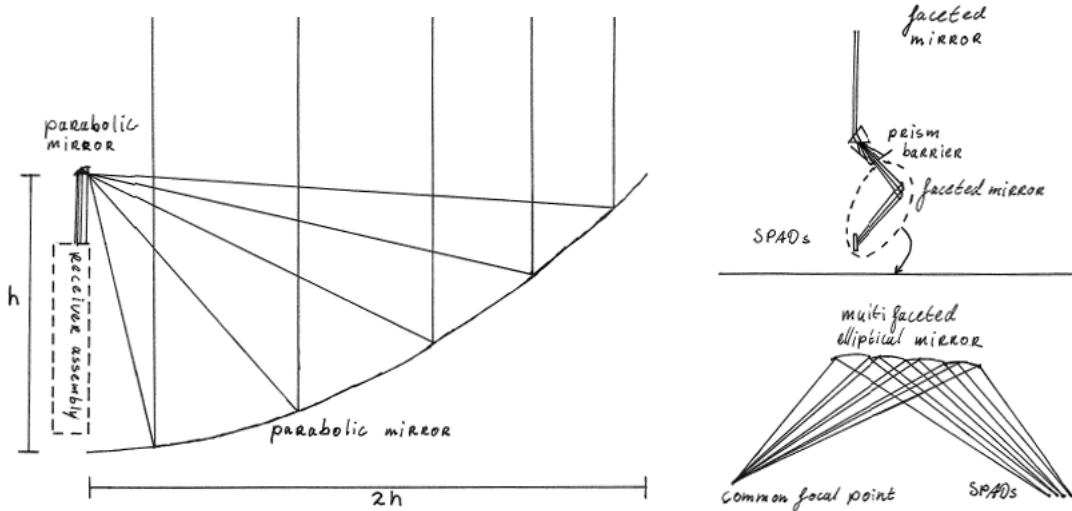


Figure 7.1: Receiver optics overview

#### 7.1.1 SPAD!

**SPAD!** identifies a class of solid-state photo detectors based on a reverse biased *p-n junction* in which a photo-generated carrier can trigger an avalanche current due to the impact ionization mechanism. This device is able to detect low intensity signals (down to the single photon) and to signal the arrival times of the photons with a jitter of a few tens of picoseconds. However, these devices are typically based on complex circuits, whose area occupation and power consumption makes their integration impossible at a pixel level.

In this case, 32 **TDC!** (**TDC!**)s have been implemented on chip serving an array of 128x128 SPAD based pixels, while reports an example of in-pixel implementation, in a linear array, of a **TAC!** (**TAC!**)<sup>[?]</sup>. The figure ?? on page ?? is how the **SPAD!** chip looks like.

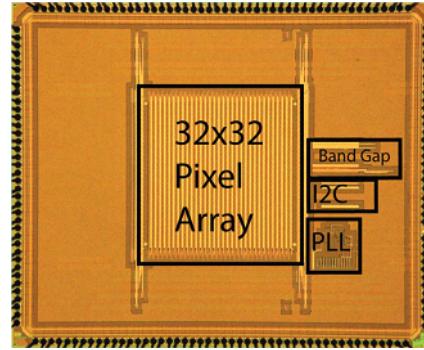


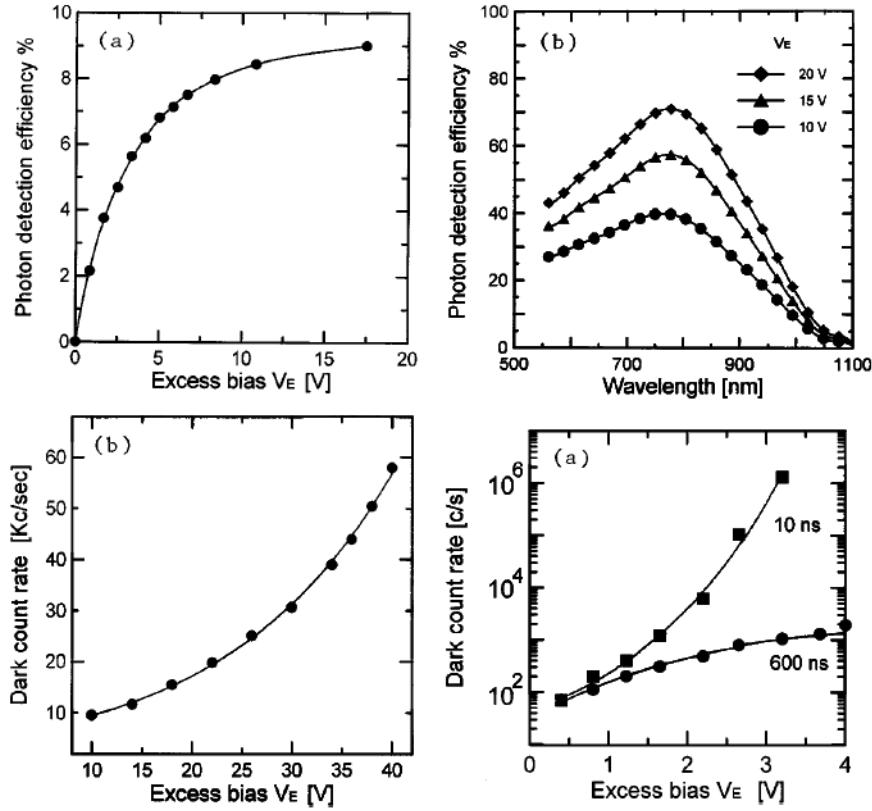
Figure 7.2: **SPAD!** chip photon-micrograph

### 7.1.1.1 Avalanche Diodes

Essentially, a **SPAD!** is a reversed biased *p-n junction* that operates at a voltage level above the breakdown voltage. At this bias, the electric field is so high that a single charge carrier injected into the depletion layer can trigger a self-sustaining avalanche. The current rises swiftly (nanoseconds or subnanosecond rise time) to a macroscopic steady level in the milliampere range. Bias supply voltage  $V_a$  exceeds breakdown voltage (non-destructive and reversible process)  $V_b$  by an amount called the *excess bias voltage*  $V_e = (V_a - V_b)$ , which has a determining influence on detector performance. Actually, the value of the ratio  $\frac{V_e}{V_b}$  is important, not the value of  $V_e$  alone, because the performance is related to the excess electric field above breakdown level.

If the primary carrier is photogenerated, the leading edge of the avalanche pulse marks the arrival time of the detected photon. The current continues to flow until the avalanche can be quenched by lowering the bias voltage to the breakdown voltage ( $V_b$ ) or below. For a photon to be detected, not only must it be absorbed in the detector active volume and generate a primary carrier (more precisely, an electron-hole pair, i.e. *an exciton*), it is also necessary that the primary carrier succeeds in triggering an avalanche. The efficiency of photon detection thus increases with excess bias voltage  $V_e$ , since higher electric field gradient enhances the triggering probability.

As it happens in **PMT!** (**PMT!**s), thermal generation effects produce current pulses even in the absence of illumination, and the Poissonian fluctuation of the dark counts represents the internal noise source of the detector. *Dark-count rate* includes primary and secondary pulses. *Primary dark pulses* are due to carriers thermally generated in the **SPAD!** junction, so that the count rate increases with the temperature as does the dark current in ordinary photodiodes. The rate also increases with  $V_e$  because of field-assisted enhancement of the emission rate from generation centers and an increase in triggering probability. *Secondary dark pulses* are due to afterpulsing effects that may strongly enhance the total dark-count rate. During avalanche some carriers are captured by deep levels in the junction depletion layer and subsequently released with a statistically fluctuating delay. Released carriers can retrigger the avalanche, generating afterpulses correlated with a previous avalanche pulse. The number of carriers captured during the avalanche pulse increases with the total number of carriers crossing the junction, that is, with the total charge of the avalanche pulse. Therefore afterpulsing increases with the delay of avalanche quenching and the current intensity, which is proportional to excess bias voltage  $V_e$ .

Figure 7.3: **SPAD!** parameters diagrams

### 7.1.1.2 Quenching Circuits

A **SPAD!** is generally implemented as a p-n junction biased above breakdown. Including quenching and recharge, there is a total 'Dead time' required by the **SPAD!** to perform detection for the next photon, which is also responsible for the upper limit of photon flux detectable by a **SPAD!**. Different quenching circuits will lead to different **SPAD!** dead time and smaller dead time is prior since during the time the **SPAD!** cannot detect any incoming photons. There exist mainly two categories of quenching circuits, named active and passive quenching.

In active quenching, the avalanche is sensed and a feedback circuit provides a mechanism to force the reverse bias of the p-n junction below breakdown. The same circuit is generally used to actively recharge the device to its initial state, above breakdown, so as to enable the next detection cycle.

In passive quenching, the avalanche current is used to directly act on the reverse bias voltage by lowering it towards breakdown voltage, which eventually quenches the current. If this is achieved, for example, using a resistance in series to the photo diode, the effective capacitance of the junction must be passively recharged through the quenching resistance.

According to the experimental result, the passive quenching circuit has the dead time in terms of hundreds nanoseconds but the active quenching circuit can reduce the dead time to a few tens of nanoseconds. That is the reason the active quenching circuit is selected in this case.

### 7.1.1.3 Space Qualification

Since the photon detection device will be operating in space, there is a big advantage of the  $32 \times 32$  pixel array **SPAD!** which did the **ESA!**'s space qualification test. According to professor Charbon(**SPAD!** developer), the **SPAD!** chip itself is survived during the test but not the motherboard connected to it, which means that a separate space application motherboard for **SPAD!** need to be designed.

### 7.1.2 Prism Design

Before the prism is taking into consideration, the optical filter is also considered as an option, which is more simple than prism and barriers. The optical filters bandwidth accuracy are in terms of several decades of nano meters(minimum 10nm[?]), which is not acceptable since the objective is to filter out all noisy light except wavelength around 473nm. Meanwhile, the transmission is another problem for normal optical filters, which have much lower value(around 50%) for smaller bandwidth accuracy (10nm to 20nm) and higher value(around 90%) for larger bandwidth accuracy (more than 30nm), but prism can has higher transmission as 97% for specified glass material.

The figure ?? on page ?? gives an overview of the receiver optics after the second parabolic mirror. the prism is used to filter out all the noisy light except wavelength 472[nm] to 474[nm]. The prism system needs to be accurate enough to perform the filtering in limited distance. The design contains the type of prism glass material, the incident angle( $\alpha$ ), prism apex angle( $A$ ) and distance between prism and barriers. They will be treated one by one.

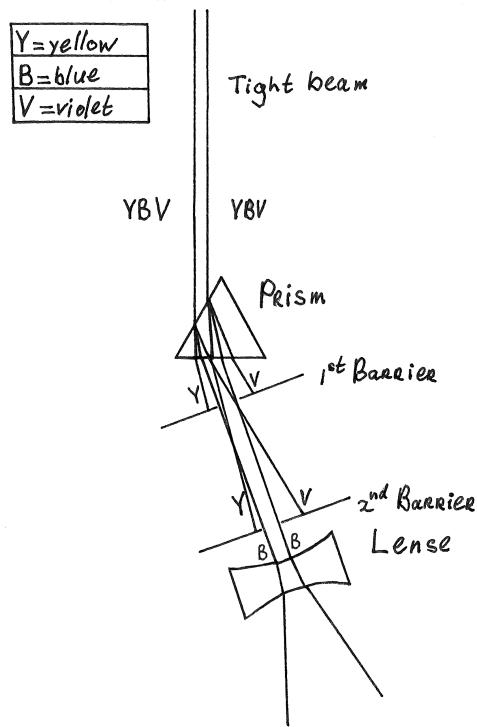


Figure 7.4: Receiver prism filter overview

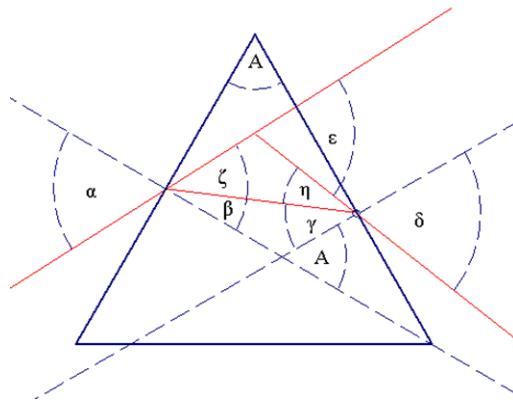


Figure 7.5: Prism angles defined in 2D

In figure ?? on page ??, the deviation angle  $\epsilon$  can be calculated using the formula[?]:

$$\epsilon = \alpha - A + \sin^{-1}(\sin(A)\sqrt{n^2 - (\sin(\alpha))^2} - \cos(A)\sin(\alpha)) \quad (7.1)$$

The key requirement to design the prism is to maximize the  $\frac{d\epsilon}{d\lambda}$ , because for wavelength 473[nm] larger  $\frac{d\epsilon}{d\lambda}$  leads to smaller distance between prism and barriers as well as the risk. The equation ?? indicates that the deviation angle( $\epsilon$ ) is a function of  $A$ ,  $\alpha$  and  $n$ . From Sellmeier formula[?], the index of refraction  $n$  can be calculated as:

$$n^2 - 1 = \frac{a_1 \lambda^2}{\lambda^2 - b_1} + \frac{a_2 \lambda^2}{\lambda^2 - b_2} + \frac{a_3 \lambda^2}{\lambda^2 - b_3} \quad (7.2)$$

In the equation ??,  $a_1, a_2, a_3$  and  $b_1, b_2, b_3$  are the dispersion coefficients, which have different values for different glasses. In this case, the most 17 common prism glasses from Schott[?][?] company are analyzed. Since it is difficult to calculate  $\frac{d\epsilon}{d\lambda}$  analytically,  $\Delta\epsilon$  due to wavelength 472[nm], 473[nm] and 474[nm] can be obtained by inserting arbitrary  $A$  and  $\alpha$ , which is actually the  $\frac{d\epsilon}{d\lambda}$  because the wavelength difference is only 1[nm] of each. During the calculation, no matter what values are given to  $A$  and  $\alpha$ , SF11 glass always has the maximum value of  $\Delta\epsilon$  which means it is the optimal glass material. Meanwhile, SF11 glass also has an internal transmittance of 97% for wavelength around 473[nm], which is acceptable. Next step is to determine the prism apex angle( $A$ ) and the incident angle( $\alpha$ ). The prism apex angle is defined as 60 degrees because which gives the most average value for  $\Delta\epsilon$ . Meanwhile, this kind of euqilateral prisms are also referred to as dispersing prisms used for wavelength separating applications(see figure ?? on page ??).

To select the correct incident angle, the figure ?? on page ?? is used. In the figure, the blue line has the asymptote about  $\alpha = 54[\text{deg}]$ . To avoid the asymptote,  $\alpha = 55[\text{deg}]$  is selected as the incident angle, which leads to  $\epsilon = 76[\text{deg}]$  and  $\Delta\epsilon = 2.1432[\text{mrad}]$ . The figure ?? is a part of the calculation in the excel sheet. There are two  $\Delta\epsilon$  in the calculation, and  $\Delta\epsilon_1 = \epsilon(472[\text{nm}]) - \epsilon(473[\text{nm}])$  meanwhile  $\Delta\epsilon_2 = \epsilon(473[\text{nm}]) - \epsilon(474[\text{nm}])$ . The driving  $\Delta\epsilon$  is the smaller one, since it is the minimum requirement. Taking  $\Delta\epsilon = 2.1432[\text{mrad}]$  into further calculation, in order to separate wavelength 472[nm], 473[nm] and 474[nm], barrier radius(or beam width) 1[mm] leads to distance 466.7[mm] between prism and barrier. To short the distance, more concentrated beam is needed from parabolic mirror and several flat mirror can also be used to redirect the beam in limited area. The figure ?? on page ?? shows the final design of all angles.

### 7.1.3 Summary

The table ?? on page ?? is the budget breakdown including mass and power consumption. To be Clear, the large parabolic mirror is used to collect the reflected photons and the small parabolic mirror is used to create

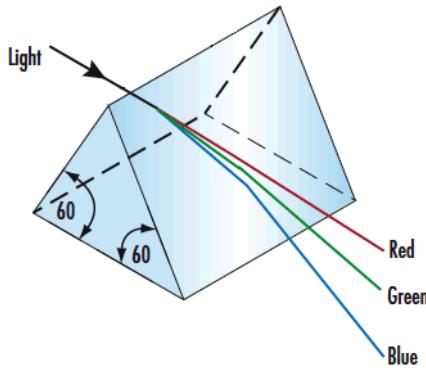
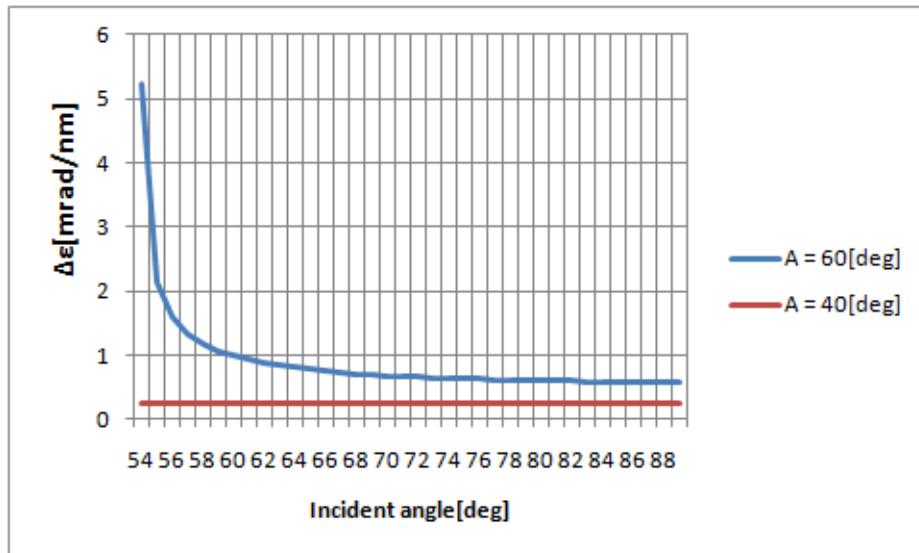


Figure 7.6: Equilateral prism for wavelength separating applications

Figure 7.7: Plot of  $\Delta\epsilon$  due to different incident angles

				$A[\text{deg}]$	$\alpha[\text{deg}]$	$A[\text{rad}]$	$\alpha[\text{rad}]$	$\epsilon[\text{rad}]$	$\epsilon[\text{deg}]$				
				60	55	1.0472	0.95993	1.33377	76.4192				
a1	a2	a3	b1	b2	b3	n(472nm)	n(473nm)	n(474nm)	$\epsilon(472\text{nm})$	$\epsilon(473\text{nm})$	$\epsilon(474\text{nm})$	$\Delta\epsilon_1[\text{mrad}]$	
SF11	1.74	3.11E-01	1.17	1.36E-02	6.16E-02	1.22E+02	1.81094	1.81061	1.81028	1.33596	1.33377	1.33162	2.191890927
												$\Delta\epsilon_2[\text{mrad}]$	
												2.14316443	

Figure 7.8: Glass 'SF11' calculations

parallel tight beam before the prism. Flat mirrors are used after the prism to focus the light on faceted mirror. A separate cost estimation can be found in the section ?? on page ??.

#### 7.1.4 Payload Cost Estimation

The table ?? on ?? is the cost breakdown of the whole emitter and receiver payload system. The table is divided into 3 main parts. The first part is the cost distribution on human resource. To design the

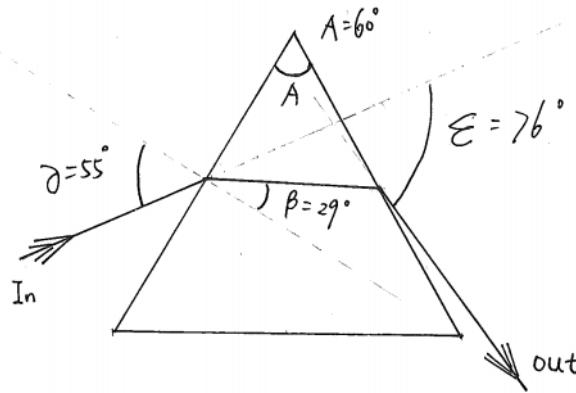


Figure 7.9: Overview of all angles

Components	Mass [g]	Quantity	Total mass [g]	Power [W]
Parabolic mirror (large)	100	1	100	0
Parabolic mirror (small)	10	1	10	0
Faceted mirror	30	1	30	0
Flat mirror	10	5	50	0
<b>SPAD!</b>	10	1	10	0.1
Prism	10	1	10	0
Diverge lens	10	1	10	0
Total			220	0.1

Table 7.1: Buget breakdown for receiver payload

payload prototype, different specific engineerings are required with a salary 122,380 \$ man work per year[?], and 'Assessment manager' gains twice salary meantime 'General manager' who is in charge of the whole project desires triple salary of the normal engineers. The second part is mainly about the manufacturing and production, meanwhile a separate space qualification test is required. The third part first calculated the theoretical first unit cost of receiver and **laser!** then obtained total cost of two types of combinations(5 receiver and 3 **laser!** or 9 receiver and 3 **laser!**) due to learning curve[?]. It is obviously that the main cost of the payload is spent on 'Brian' engineering design which is reasonable. For example, the final cost of 5 receiver and 3 **laser!**  $8.55[M\$] = a + b + c = 5.13996 + 2.92 + 0.49$  and 9 receivers 3 **laser!** leads to  $8.66[M\$] = a + b + d = 5.13996 + 2.92 + 0.6$ . Finally, the amount of cost is scaled to dollar year 2000.

## 7.2 Navigation

In this chapter the possibilities for position determination are discussed, after which the best option is chosen and developed further. The first section will discuss the options considered for navigation, the second section at page ?? contains a small trade-off, and the final section at page ?? is concerned with further development of the result of the trade-off.

Table 7.2: Cost breakdown for emitter and receiver payload

### 7.2.1 Option description

For the position determination four different ideas are considered, the first one is equipping every satellite with a **GPS!** receiver. This way all satellites can determine their position independent from each other.

Several small **GPS!** receivers have been developed specifically for small satellites.

The second option is to use lasers to determine the both the relative position of the satellites and transmit data between satellites. To determine the absolute position of the satellites the emitter is equipped with a **GPS!** receiver as well. This will allow the receivers to determine their absolute position. The **GPS!** unit will also be used to ensure the clocks on the receiver satellites are properly synchronized.

The third option is to use our intersatellite communications system to determine the relative position of the satellites, while maintaining regular communications alongside this. For this system to work the emitter is equipped with a GPS receiver in order to find its absolute position. This GPS unit is also used to ensure the clocks on board the receiver satellites are updated regularly. All satellites are designed to operate at a slightly different frequency to prevent interference from occurring. The process used to determine the position of the receivers is as follows. All satellites send out a signal to all other satellites, which send back a signal themselves after a fixed amount of time. The time between the sending and receiving of the signals and processing time can be used to determine the relative distance between satellites. When all satellites are given the distance information from the other satellites and the GPS information from the receiver, the absolute position of each satellite can be determined. Note that the satellites have to be able to accurately determine where a signal has come from.

The fourth and final option is to use the sun sensors and star trackers used for attitude determination to also determine the absolute position of the satellites.

### 7.2.2 Trade-Off

This section does not contain a trade-off table, a decision is made based on the practicality of the option. As such the practicality of each options is discussed with a final paragraph addressing the final result. The design options will be treated in no particular order.

The first option is using the sun sensors and star trackers to determine the position of the satellite, which -though possible- is a complex and tedious procedure. For example the accuracy will be different when the satellite is in and out of eclipse. It should also be noted that the resulting accuracy will not be very high.

The second option is using **GPS!**, this system is commonly used on **LEO!** satellites for several years now. As long at least four **GPS!** satellites are in view, the position of the receiver can be determined. For some years companies have been designing cheap, small GPS receiver using **COTS!** parts. Also the real-time accuracy that can be obtained is at least ten meters, and can go up to less than a meter. After post processing the satellite position can even be determined up to several [mm]. As such using **GPS!** is a viable option.

The third option considered is using the communications system to determine the position of the satellites. Obviously, the technology for this technique is available. The biggest drawback of this technique is that every receiver satellite has to be equipped with an extensive communications system in order to keep track of the other satellites very precisely, non more so than the satellites that are not in the same orbit plane as the emitter. These satellites cross the orbital plane of the emitter twice per orbit, and as such the communication system has to be able to see the entire sphere of their surroundings. Because the position determination has to be performed accurately the system will require narrow field of view, moveable antennas. As a result this option is very impractical. Added to this is the problem that at least one more **GPS!** receiver is required than has been thought necessary during the conception phase. This is another reason why this option is not considered viable.

The last option is similar to previously discussed one, as the laser is used both for communication and position determination. Also every satellite would require several laser systems to keep track of the other satellites. Added to this is the fact that it is easier to aim the laser systems if the satellites know approximately where they are, for which they can use **GPS!** receivers. Finally, the technology for this method is not yet

Parameter	SGR-05P	GPS-12-V1
Time accuracy [ns]	500	20
Position accuracy [m]	10	10
Velocity accuracy [m/s]	0.15	0.03
Time to first fix [s]	90-180	30-60
Size [mm]	100 x 65 x 12	100 x 70 x 25
Mass [g]	60	<200
Power [W]	1	1
Interface included?	no	yes

Table 7.3: Comparison of the SGR-05P and the GPS-12-V1 by **SSTL!** and SpaceQuest respectively.

compact enough for nanosatellites or small microsatellites.

The result of the investigations are that combining the communications system with the position determination is both complex and heavy. So neither of the two systems will be used, this leaves the **GPS!**-only option and using the attitude systems.

Using the attitude systems for position determination is more complex and less cost effective than using only **GPS!** ???. So the option that will be used is a **GPS!** receiver on every satellite.

### 7.2.3 Working out the best option

This section is split into two parts, the first concerning the choice of **GPS!** receiver and the second part considers the antenna(s) to use.

#### 7.2.3.1 GPS receiver

In order to determine the best **GPS!** receiver three sources have been used. These sources are the GPS World website ??, the SpaceQuest website ?? and the **SSTL!** website ???. From ??, the 2010 document has the advantage of showing whether or not a receiver can be used in space. From this list the best candidate is the **SSTL!** SGR-05P receiver due to the fact it is designed solely for space applications and has the lowest weight. Another promising candidate is the GPS-12-V1 from SpaceQuest. In order to determine the best one both are compared side by side in table ?? on page ?? for several characteristics.

From table ?? it can be seen that the SpaceQuest receiver has the best accuracy, but is heavier than the **SSTL!** receiver. However it should be noted that the GPS-12-V1 includes interface boards, for the SGR-05P there exists a receiver that is equipped with interface boards, it is the SGR-07. However that one is both heavier and larger than the GPS-12-V1. So the receiver that will be used by the swarm satellites is the GPS-12-V1.

An additional and interesting characteristic of current **GPS!** receivers is that the position accuracy can be determined up to several centimeter or millimeters. It should be noted that since the **GPS!** receiver position is known to such great accuracy the same receiver is used on the emitter. The emitter will also use the same antenna, which is chosen in the next section.

### 7.2.3.2 GPS Antenna

The main source for possible antennae is again GPS World, however in this case their receiver survey at ?? This list is checked for space qualification, just like the receiver survey. There is only one antenna specifically developed for usage in space. This is the SGR Patch Antenna ASY-00741-04, which is also the smallest antenna and also specifically developed for small satellites. The antenna has a power requirement of 24 [mA] at 2.7 to 5 [V], its dimensions are 45 x 50 x 20 mm and it uses a cable with a length of 2 [m].

## 7.3 Communications subsystem

This section will discuss the communication link to which the receiver satellite is connected:

- Crosslink for scientific and housekeeping data

These links are illustrated figure XXX For this link the link budget and the communications hardware selected will be presented. Only hardware specifications will be given, a more detailed description of the hardware selection can be found in ??.

### 7.3.1 Crosslink for scientific data and housekeeping data

This link transmits the scientific data gathered by the receiver satellites to the emitter satellite. The receiver satellites will transmit their scientific data continuously so that the emitter satellite can store all data in one data bank before it is transmitted to the ground. In addition to scientific data, command and housekeeping data is also transmitted in this link.

The link budget has the following input parameters: the frequency selected was 2 GHz, the data rate required is 1.62 Mbit/s, the maximum distance between the satellites is 261 km, the modulation used is QPSK. Atmospheric losses were not considered for obvious reasons.

A detailed description of these parameters can be found in ??.

#### 7.3.1.1 Transceiver

The transceiver selected was an SBTRcvr of RDLabs, which has a maximum data rate of 10 Mb/s for QPSK modulation, an input power of 12 Watt and an output power of 1 to 5 Watt ??.

#### 7.3.1.2 Antenna

The S Band patch antenna from Surrey Satellite technologies?? was selected as the antenna for this link, it has a mass of only 80 grams, a 120° beamwidth, a gain of 4 dBi and measures 82x82x20 mm.

#### 7.3.1.3 Link budget

Because the selected modulation is QPSK and the maximum allowable bit error rate is  $10^{-5}$ , the required  $E_b/N_0$  is equal to 9.6 dB. With the output power of the transceiver at maximum (5 Watt or 7 dBW) it is possible to have a  $E_b/N_0$  of 10.35 dB, which leaves a margin of 2.36 dB.

Parameter	Emitter
Max. time without contact to ground station	7:35:33
Total bit rate [Mbit/s]	1.63
Max. amount of data to be stored [Gbit]	49
Time required to send data to the emitter [s]	6:47:19

Table 7.4: Some values used to determine the data storage for the receiver.

## 7.4 Data Storage for the Receiver

In order for the mission to succeed it is important that all collected data can be stored somewhere. Because the emitter is the only satellite with a link to the ground, it is most logical that all data is stored on the emitter. However to minimize the chance of losing data the receivers will also be equipped with a small unit for data storage. Table ?? shows the values important for data storage on the receiver.

Assuming some problem occurs with data transmission that can only be solved after communication with a ground station, the emitters may have to remember data of up to 7.5 hours (table ??). This would require a storage capacity of 49 [Gbit]. As a result each receiver satellite will use a single 64 [Gbit] flash memory module from 3D-Plus. This also ensures the possibility to store more data should the need arise, or allow some redundancy if one of the memory banks should fail. A final calculation is made to determine how long it would take to send all the saved data to the emitter, which is about 6.75 hours. For more details on the flash memory unit please go to the emitter section on page ??.

## 7.5 AODCS!

The **AODCS!** of the receiver is basically a scaled down version of the **AODCS!** of the emitter. Section ?? will thread the attitude determination, section ?? the attitude control, the orbit determination of the receiver satellites is covered in section ?? and the orbit control in section ???. The pointing mechanism for the receiver satellites is significantly more complex compare to the emitter and is described in section ???. Section ?? will give an overview of the whole **AODCS!** and pointing mechanism of the receiver.

### 7.5.1 Attitude Determination

Just as the with the emitter the attitude determination of the receiver satellites is done with sun sensors and a star tracker. The only difference between the two kinds of satellites is that the receiver has got only four sun sensors. In stead of have a sun sensor on each side wall there is only one on the front wall and two on the back looking out at an angle of sixty degrees. This way room is saved and still the entire surrounding of the satellite is covered.

The type of sun sensor and star tracker used are the same as in the emitter satellite. Details can be found in section ??.

### 7.5.2 Attitude Control

Just as the emitter satellite the receiver uses reaction wheels and magneto torquers for its attitude control. Since the satellite itself is smaller also the wheels and torquers are scaled down. Figure ?? on page ?? shows a possible layout to have control around all axis.

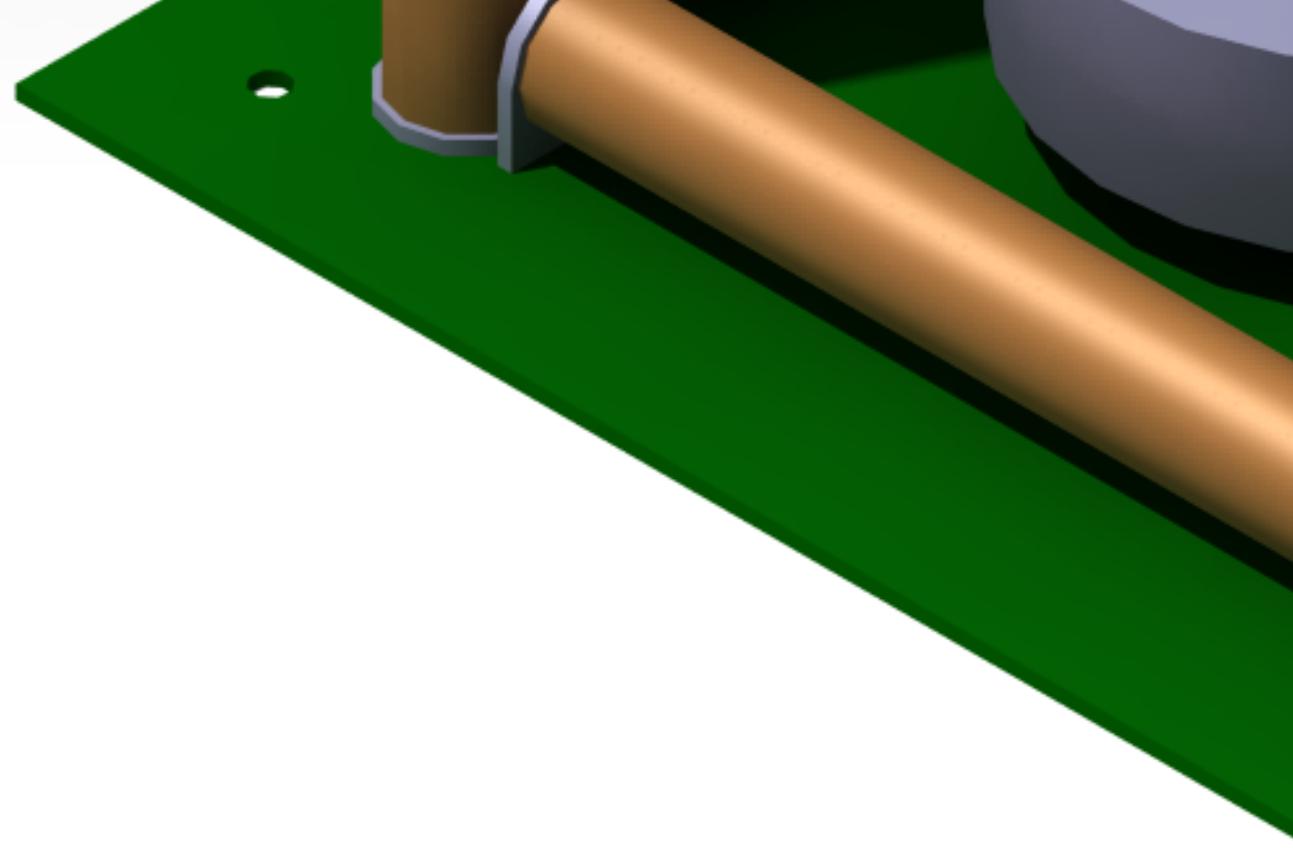


Figure 7.10: General layout of the attitude control subsystem with reaction wheels and magneto torquers on all three axis.

The sizing of the reaction wheels for the receiver satellites is done in the same way as the wheels of the emitter satellite. Filling out equations ?? to ?? for the worst case of the receiver satellites leads to

$$\begin{aligned}
 T_g &= \frac{3.398600.4 \cdot 10^9}{2.6878000^3} |1.316 - 0.101| \sin(2 \cdot 30^\circ) = 1.933 \cdot 10^{-6} \text{ Nm} \\
 T_{sp} &= \frac{1367}{3 \cdot 10^8} 0.53 (1 + 0.6) \cos 0^\circ (0.15) = 5.796 \cdot 10^{-7} \text{ Nm} \\
 T_m &= 4.893 \cdot 10^{-5} \cdot 0.1 = 4.893 \cdot 10^{-6} \text{ Nm} \\
 T_a &= \frac{1}{2} 1.80 \cdot 10^{-12} \cdot 2.2 \cdot 0.53 \cdot 7613^2 (0.15) = 9.123 \cdot 10^{-6} \text{ Nm}
 \end{aligned}$$

The moments of inertia used are derived from the drawing of the receiver satellite, i.e.  $I_x = 1.307 \text{ kgm}^2$ ,  $I_y = 0.101 \text{ kgm}^2$  and  $I_z = 1.316 \text{ kgm}^2$ . Other typical properties are derived using SMAD [?] and the master thesis of Angel Garza [?]. The total disturbance torque expected is the sum of the above values  $1.65 \cdot 10^{-5} \text{ Nm}$ . With a margin of safety of 2 applied this leads to a maximum control torque required of  $3.31 \cdot 10^{-5} \text{ Nm}$ . Taking a angular acceleration of  $100 \text{ rad/s}^2$  ( $\sim 10\%$  of maximum) this will lead to a required wheel moment of inertia of  $3.306 \cdot 10^{-7} \text{ kgm}^2$ . With a wheel height  $h$  of 1 cm this leads to a wheel thickness  $b$  of 1.1 cm. The mass of one wheel is 36 grams.

The magneto torquers chosen are CubeSat Magnetorquers, which are available off the shelf from the cubesat webshop [?]. The weight is 30 grams and the diameter is 9 mm and the length is 70 mm. They can provide a torque of  $0.2 \text{ Am}^2$ . Three are required for the 3 axis desaturation. Price is 1150 EUR a piece. Three are needed to desaturate all three wheels.

### 7.5.3 Orbit Determination

Just as with the emitter the orbit determination of the receivers is done by the navigation system. For the receiver this is described in section ??.

### 7.5.4 Orbit Control

For orbit control a  $\Delta V$  budget of 165.1383 m/s is required (derived in section ??). The thruster used for achieving this budget is the M050HP monopropellant thruster from Micro Aerospace Solutions, Inc [?]. It uses hydrogen peroxide ( $H_2O_2$ ) and has a specific impulse  $I_{sp}$  of 120 seconds. Using the propellant to dry mass ratio depicted in equation ?? and using a satellite dry mass of 13.52 kg, a propellant mass of 1.77 kg is needed. The thruster has a mass of 6.35 grams.

### 7.5.5 Pointing Mechanism

The pointing mechanism is used for pointing the receiver off-nadir towards the ground target. The right ascension difference of  $2.18^\circ$  translates into a pointing difference of about  $30^\circ$  in both directions. To be able to still make measurements when the distance between the satellites gets bigger a total design pointing angle of  $80^\circ$  is chosen. The pointing of the instrument is done by a tumbler, drive by a wheel, connected to a harmonic gear box, attached to a stepper motor. The general layout can be seen in figure ?? on page ???. The different parts of the pointing mechanism will be discussed in reverse order in this section.

The selected stepper motor is a Faulhaber ADM1220. The nominal power is 0.6 Watts. The diameter of the motor is 12 mm, the length of the motor and the axis is 17.4 mm and the motor has a mass of 9 grams. The motor is capable of making 20 discrete steps, giving it a step angle of  $18^\circ$ . Naturally this is not accurate enough, so a gearbox is needed.

For the gearbox a harmonic drive is chosen [?]. It is chosen because it has half the size, one third of the weight of a conventional gear. A Harmonic Drive CSF-8-50 is selected [?]. The gearbox has a gear ratio of 50:1, a diameter of 30 mm, a length of 22.1 mm and a mass of 26 g. With this set-up it is possible to make  $20 \times 50 = 1000$  steps for a full circular rotation of the output axis. Divided over the  $80^\circ$  of the possible pointing options this gives a pointing capability with steps of  $0.08^\circ$ .

The output of the harmonic drive is put into a wheel with the same circumference as the  $80^\circ$  sweep of the tumbler, so a complete rotation of the wheel sweeps the receiver from one maximum to the other. This translates into a wheel with a radius of 22.15 mm and a tumbler with a radius of 59.85 mm from the rotation axis of the receiver. The pointing of the receiver can be read by an array of sensors reading a 10-bit ( $2^{10} = 1024$  steps) grey code from the side of the tumbler.

### 7.5.6 Overview

In table ?? on page ?? an overview is given of all masses, costs and power requirements of the **ADCS!**, **ODCS!** and the pointing mechanism. Because it is essential for the mission that the instrument is pointed towards the ground target extra mass is added to the support structures for the pointing mechanism.

## 7.6 Receiver EPS

De EPS system of the receiver is almost the same as that of the emitter.



Figure 7.11: General layout of the receiver pointing mechanism, with the stepper motor, harmonic drive and the tumbler.

The required power generation is 42W, resulting in a required area of 0.5 square meters. Again the STK program was used to verify that the assumed cosine loss value was correct. The calculated power generation for 0.5 square meters was 52.5W, thus again it is shown that the value was correct.

Because the power requirement is much less also the amount of batteries necessary was less. For the receiver, only one battery with 2 modules was needed to be able to power the satellite during eclipse.

As with the emitter, here the architecture of the power system of the receiver is show in figure ???. The power from the solar panels is transferred to the dc-dc convertor. There part of it is sent to charge the batteries and part is sent to power the payload. During eclipse periods, the power is sent from the batteries to the payload via the convertor.

Subsystem component	Number	Mass [g]	Cost [EUR]	Total mass [g]	Total cost [EUR]	Power (max) [W]
Sun sensors	4	24	11000	96	44000	0 (0)
Star trackers	1	375	75000	375	75000	1 (2)
Reaction wheel motors	3	8.5	176	25.5	528	0.15 (1.5)
Reaction wheel disk	3	36	10*	108	30*	0 (0)
Magneto torquers	3	30	1150	90	3450	1 (3)
Stepper motor	1	9	112	9	112	0.6 (1.2)
Harmonic drive	1	26	180*	26	180*	0 (0)
Tumbler	1	22	10*	22	10*	0 (0)
Thruster	1	15	250k	15	250k	1 (2)
Fuel tanks and pumps	1	1130	300*	1130	300*	0 (0)
Support structures				1000	400*	0 (0)
Total				2500	375k	3.75 (4.5 rms)
					435500 FY00\$	

Table 7.5: Mass, cost and power budget of different parts of the **ADCS!**, **ODCS!** and pointing mechanism of the receiver. Starred prices are estimations based on material and machining cost.

Part	Dimensions [mm]			Weight [g]	Power usage [W]
	Length	Width	Height		
Driver (2 needed)	30	6	60	21.4	1
SMA Deployment (2 needed)	120	50	10	120	4*
Battery	168	102	10	1000	0
Convertor	95	60	17	80	1.5
Shunt regulator (2 needed)	2.8	2.6	1.05	0.1	0.5
Thermal knife (2 needed)	60	50	38	280	15**
Wiring	-	-	-	230	0.28
Solar panels (2 needed)	500	500	0.75	1.2	0
<b>TOTAL</b>	-	-	-	3593.5	4.875***

\* for 4 minutes

\*\* for 60 seconds

\*\*\* continuous power usage only

Table 7.6: EPS subpart details for receiver satellites

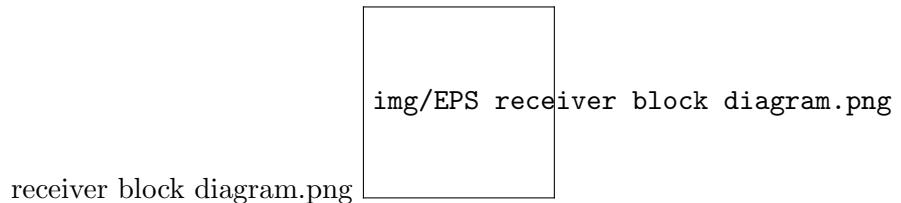


Figure 7.12: Electrical block diagram of the receiver

# Chapter 8

# Sustainable Development Strategy

In this chapter the Sustainable Development Strategy is discussed in the order of production (section ??), operations (section ??) and **EOL!** (section ??).

## 8.1 Production and Logistics

The design is aimed at a swarm of mostly identical satellites. This may allow for series production which is more efficient in terms of resources than a one-of large satellite with a lot of unique components. This also implies that the number of different spare parts could be reduced. Smaller satellites could also use smaller facilities for production and testing.

Transportation can be split up into two parts: transportation to the launch site and the launch from the surface to the final orbit in space. On both occasions the system can again profit from its small size. If the satellites are not launched all together, they can “piggyback” on another satellite’s launcher.

Spreading the swarm, i.e. piggybacking using different launchers, has several advantages. First of all, the emissions are lower than in case of a dedicated launcher. Also, if the first satellite fails before the launch of the rest of the swarm, the others can be repaired and thus less resources are wasted.

## 8.2 Operations

Once in orbit, the satellite’s influence on the Earth is very limited. The only real concern is the debris it leaves behind during launch and deployment, which can be dangerous to other satellites orbiting the Earth. However, the deployment mechanism, which is responsible for most of the debris, is not included in this technical feasibility study. Later studies developing the ideas from this feasibility study should take it into account, since more satellites could mean more deployment mechanisms and hence more waste. One aspect that can be dealt with is the efficient use of resources. The swarm can be designed in such a way that if one of the satellites fails a replacement satellite can be sent, whilst any remaining satellites can be reused.

### 8.3 End of Life

Each satellite will be at the end of its life if it cannot perform its function anymore. It is important that after the mission is over all satellites are removed from their orbit and burn up in the atmosphere so that they do not pose any danger to other satellites. Final decommissioning of the swarm will be more complex than for a regular satellite, since every individual satellite has to be decommissioned separately.

## **Appendix A**

### **Soyuz-ST Factsheet**

