

TOV

Notes

$$\begin{aligned}\frac{dP}{dr} &= -\frac{G[P + \mathcal{E}(r)][M(r) + 4\pi r^3 P/c^2]}{c^2 r^2 [1 - 2GM(r)/(c^2 r)]} \\ \frac{dM}{dr} &= 4\pi r^2 \frac{\mathcal{E}(r)}{c^2}\end{aligned}\quad (1)$$

- Compare with Newton
- Also do white dwarfs

An interpretation of these equations can be more readily seen by multiplying the first equation by $4\pi r^2 \mathcal{E} dr/c^2 = dM$ and cancelling \mathcal{E} on both sides:

$$\begin{aligned}4\pi r^2 dP &= -\frac{GM dM}{r^2} \left(1 + \frac{P}{\mathcal{E}(r)}\right) \left(1 + \frac{4\pi r^3 P}{Mc^2}\right) \\ &\quad \times \left(1 - \frac{2GM}{c^2 r^2}\right)^{-1}\end{aligned}\quad (2)$$

The term on the left hand side is the force exerted on a infinitesimal shell at radius r . The first factor on the right hand side is the newtonian gravitational force from the interior acting on this shell.

Numerical set-up

Making the substitutions $r = R_0 x$, $P = P_0 p$, $\mathcal{E} = P_0 \varepsilon$ and $M = M_0 m$, we can write

$$\begin{aligned}\left(\frac{P_0}{R_0}\right) \frac{dp}{dx} &= \\ -\left(\frac{GP_0 M_0}{c^2 R_0^2}\right) \frac{[p + (\mathcal{E}_0/P_0)\varepsilon][m + 4\pi R_0^3 x^3 P_0 p/c^2]}{x^2 [1 - 2GM_0 m/(c^2 R_0 x)]}\end{aligned}\quad (3)$$

and thus

$$\frac{dp}{dx} = -\left(\frac{G_0 M_0}{c^2 R_0}\right) \frac{[p + (\mathcal{E}_0/P_0)\varepsilon][m + 4\pi R_0^3 x^3 P_0 p/c^2]}{x^2 [1 - 2GM_0 m/(c^2 R_0 x)]}\quad (4)$$

Now, we set

$$1 = \frac{\mathcal{E}_0}{P_0} = \frac{GM_0}{c^2 R_0} = \frac{4\pi R_0^3 P_0}{c^2}\quad (5)$$

and find

$$\begin{aligned}\frac{dp}{dx} &= -\frac{(\varepsilon + p)(m + x^3 p)}{x(x - 2m)} \\ \frac{dm}{dx} &= x^2 \varepsilon\end{aligned}\quad (6)$$

The initial value for the first equation can either be $p(0) = p_0$ or $p(x_0) = 0$ where p_0 is the pressure at the center and x_0 is the radius of the star. The initial value for the second equation is $m(0) = 0$. We wish to solve this equation with a given equation of state $\varepsilon(p)$.