

Assessing the Scientific Value of Amateur Photometric Observations of Binary Star Transits

Thomas Maurant

Polytech Montpellier

thomas.maurant@etu.umontpellier.fr

January 24, 2025

Abstract

Amateur astronomers have increasingly contributed to astronomical research through advancements in affordable equipment and accessible software. This paper investigates the precision and reliability of photometric observations of binary star transits conducted by amateur astronomers compared to data collected by professional observatories. We analyze case studies, compare methodologies, and assess the implications for collaborative scientific research.

Keywords: photometry, binary star systems, amateur astronomy, professional astronomy, data comparison

Contents

1	Introduction	3
2	What are transits?	3
3	On the U Cephei system	4
3.1	Why U Cephei?	4
3.1.1	Period	4
3.1.2	Eclipse duration	4
3.1.3	Magnitude variation	4
3.1.4	Data availability	4
4	Methodology	5
4.1	Observation Setup	5
4.2	Observation Procedure	5
4.3	Cleaning the amateur data	5
4.4	Professional Data Acquisition	8
4.5	Standardization of Data Formats, light curves and comparative analysis .	8
5	Results	8
5.1	Light curve	8
5.2	Comparative Analysis	9
6	Discussion	10
6.1	Figure 1	10
6.2	Figure 2	11
6.3	Figure 3	11
7	Conclusion	11

1 Introduction

In recent years, the field of astronomy has witnessed a growing collaboration between amateur and professional astronomers. The advent of high-quality, low-cost equipment has enabled amateur astronomers to perform sophisticated observations, including photometric studies of binary star transits. This paper focuses on photometric observations of the U Cephei (U Cephei) binary star system conducted by the author using personal equipment from a Bortle 7 location. The study aims to examine the precision and reliability of these amateur observations in comparison to professional datasets, exploring the potential for collaboration and the challenges inherent in combining data from diverse sources.

2 What are transits?

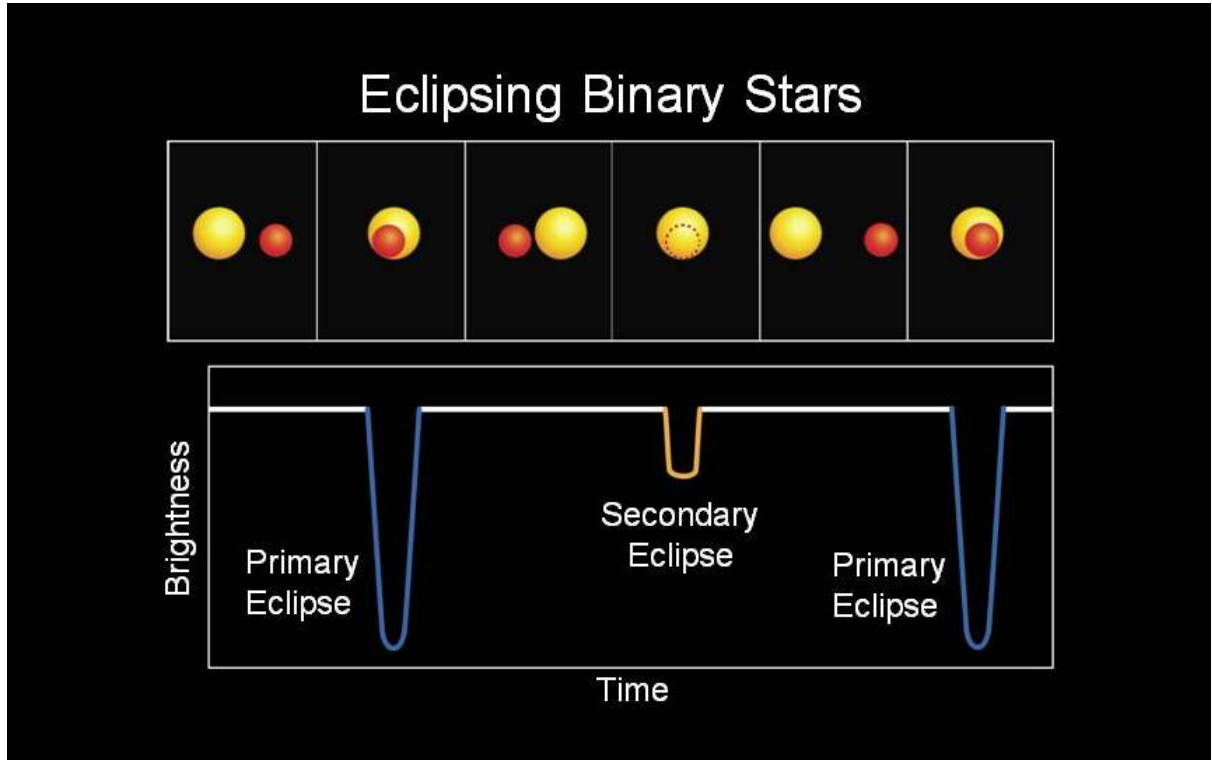


Figure 1: Transit of binary star system schema and associated light curves.

Transits are a phenomenon that occurs when a celestial body passes in front of another, partially or totally blocking its light. Those events are very interesting for astronomers as they can provide a lot of information about the objects involved, such as their size, mass, and even atmosphere.

The first exoplanet was discovered using the transit method, and it's still one of the most used methods to detect exoplanets. One of the way to know if a star is part of a binary system is to observe it and see if it's light intensity varies over time, which can be a sign of a transit event.

3 On the U Cephei system

U Cephei is a very interesting binary star system and a great pick to try catching transit eclipses.

3.1 Why U Cephei?

There are multiple reasons why U Cephei is a great candidate for me in this study.

3.1.1 Period

First of all U Cephei period is 2.493 days, which is quite short and allows for multiple chances to observe it on a short period of time. Even though this period seems short since it's not exactly 3 days but 2 and a half it means that one observation out of 2 will be at night and the other during the day. Making observation nights possible every 5 days. Considering the fact weather conditions are critical for this field, managing to have a night with good seeing condition can also be a challenge, so having multiple chances to observe the same system is a great advantage.

3.1.2 Eclipse duration

Secondly, this system eclipse is short, lasting only a couple of hours it's possible in one night to observe the complete eclipse, which is not the case for all binary systems.

3.1.3 Magnitude variation

Stars magnitude define how bright they are. This scale of magnitude is logarithmic, meaning that a difference of 1 magnitude is a difference of 2.512 times in brightness and a difference of 2 magnitudes is a difference of **2.512²** times in brightness. the scale is inverted, meaning that the lower the magnitude, the brighter the star is.

The magnitude of U Cephei varies from 6.7 to 9.2. Which is actually a huge variation in brightness. If we wanna have a comparison in percentage we can do the following calculation:

We can right the magnitude variation as a ratio of brightness variation:

$$2.512^{m_2-m_1} = \frac{F_1}{F_2} \quad (1)$$

$$2.512^{9.2-6.7} \approx 10 \quad (2)$$

In other words, during the transit, U Cephei will be 10 times dimmer than when it's not in transit. This is a huge variation and makes it easier to detect the transit.

3.1.4 Data availability

Finally, U Cephei is a well-known system, with a lot of data available from professional observatories, which makes it a great candidate for me as an amateur astronomer to try to catch some transits and compare my data to the professional ones.

4 Methodology

To address the research question, this study employs a comparative analysis of photometric data. Observations of the U Cephei binary star system were conducted using personal equipment under urban light pollution conditions (Bortle 7). The following steps were undertaken:

4.1 Observation Setup

Specifying the equipment used in the study is essential to understanding the context and limitations of the observations. I employed the following equipment for the photometric observations of U Cephei:

- Telescope: Sky-Watcher Quattro 150P Newtonian reflector telescope with a 150mm aperture and 600mm focal length.
- Reducer: Sky-Watcher 0.85x Coma Corrector lowering the focal length to 510mm.
- Camera: ZWO ASI 585MC pro monochrome CMOS camera with a 8.29 MP sensor.
- Filters: UV and IR cut filter for reducing light pollution and enhancing contrast.
- Mount: Celestron Celestrong cg-5 goto equatorial mount for trackin.
- Autoguider: ZWO ASI 120MM mini monochrome camera and 50mm guide scope for guiding.
- Software: N.I.N.A for image acquisition and PHD2 guiding software for tracking. AstroImageJ for data analysis and Siril for aligning the images.
- Light Pollution: Montpellier, Bortle 7 (suburban/urban transition) location with moderate light pollution.

4.2 Observation Procedure

4.3 Cleaning the amateur data

The first step in analyzing the data is to clean it, this means removing any outliers or bad data points that could affect the results but also in our case making sure our star is aligned on each frame to make the process easier.

Using Siril which is a free and opensource software for astrophotography, it is very easy to align all the images of the observation to make sure the star is always at the same place in the frame. As a comparison here is the same data after the cleaning process:

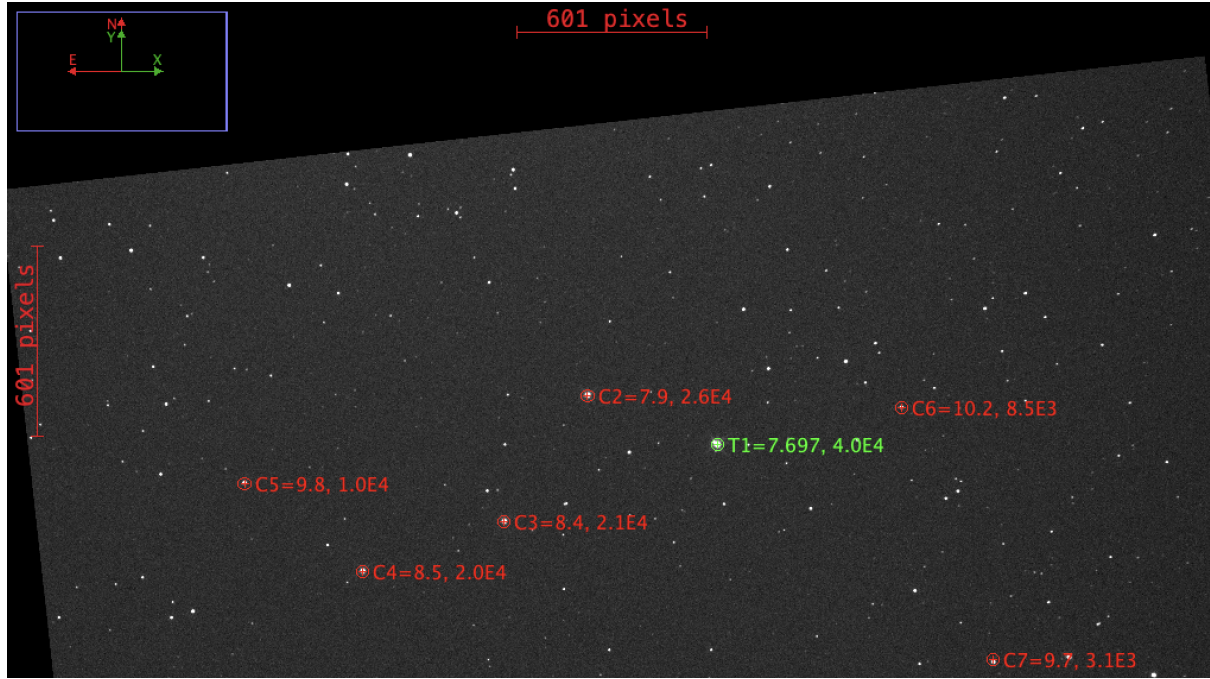


Figure 2: Cleaned data of U Cephei and it's comparison stars used for the study.

This screenshot has been taken in AstroImageJ, which is a software used to analyze astronomical images and easily create light curves. Here we can see in green my star of interest U Cephei, and in red the comparison stars used to normalize the data. You can see that next to each star there is a number, this is the magnitude of the star, which is a measure of its brightness.

I chose manually the stars and took them because they have a known magnitude, allowing the software to guess the magnitude of U Cephei. To know the magnitude of those reference stars I used the AAVSO website which is a great tool to get information about stars and their magnitude.

By searching for U Cephei you basically get a chart of the star and the stars around it, with their magnitude, which is very useful to pick the right stars for the study.

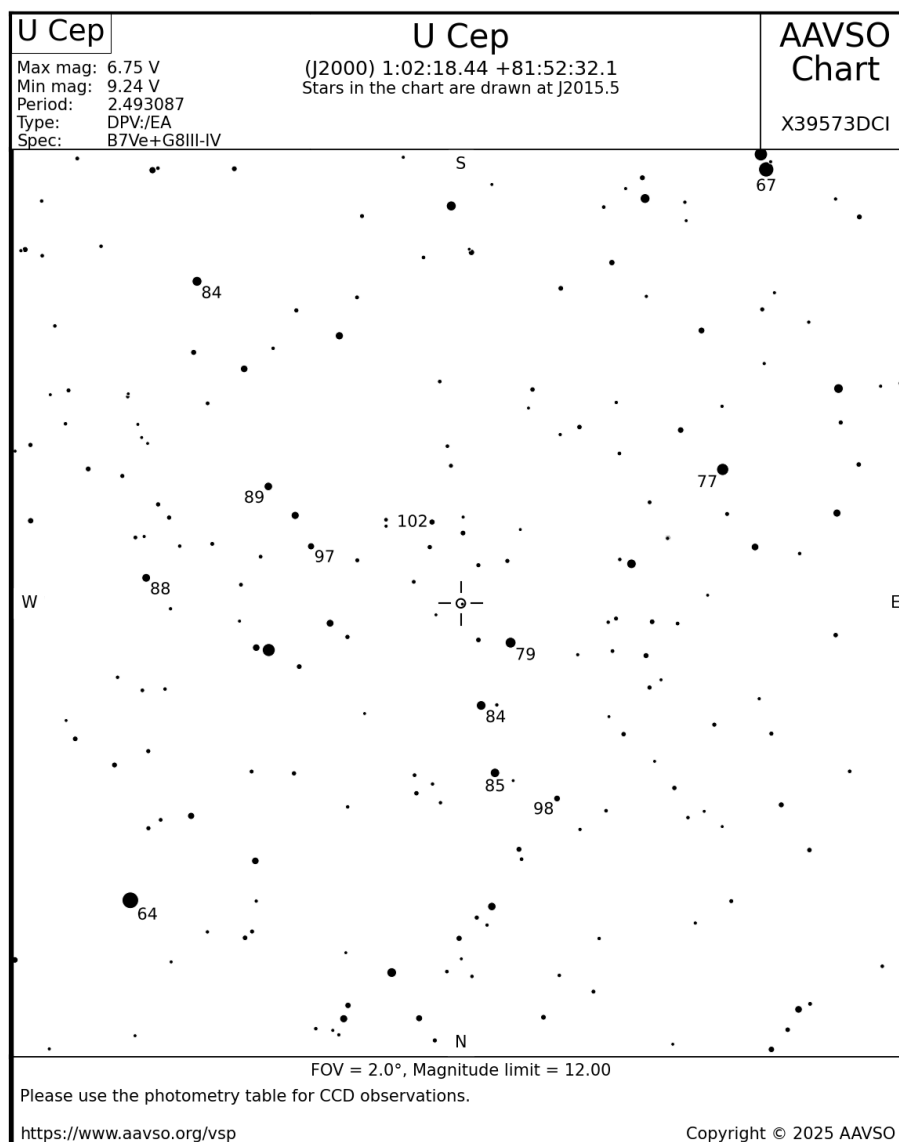


Figure 3: AAVSO Star chart of U Cephei and neighbors star magnitudes.

It's important to note something else in the 2 image, we see that the frame has been shifted, this is what Siril did to make sure each frame is aligned.

After cleaning the data I wanted to see if it was possible to detect the transit event just by creating a video of the frames. Since this eclipse variation of magnitude is so important I felt like it could be visible enough.

Using a python script I converted each of the fit file into a grayscale png and then was able to use the ffmpeg library to create a video of the frames. The result is available [here](#).

In it we can clearly see the star getting dimmer until the frame 150 and then brighter again, this is the transit event happening.

4.4 Professional Data Acquisition

4.5 Standardization of Data Formats, light curves and comparative analysis

5 Results

5.1 Light curve

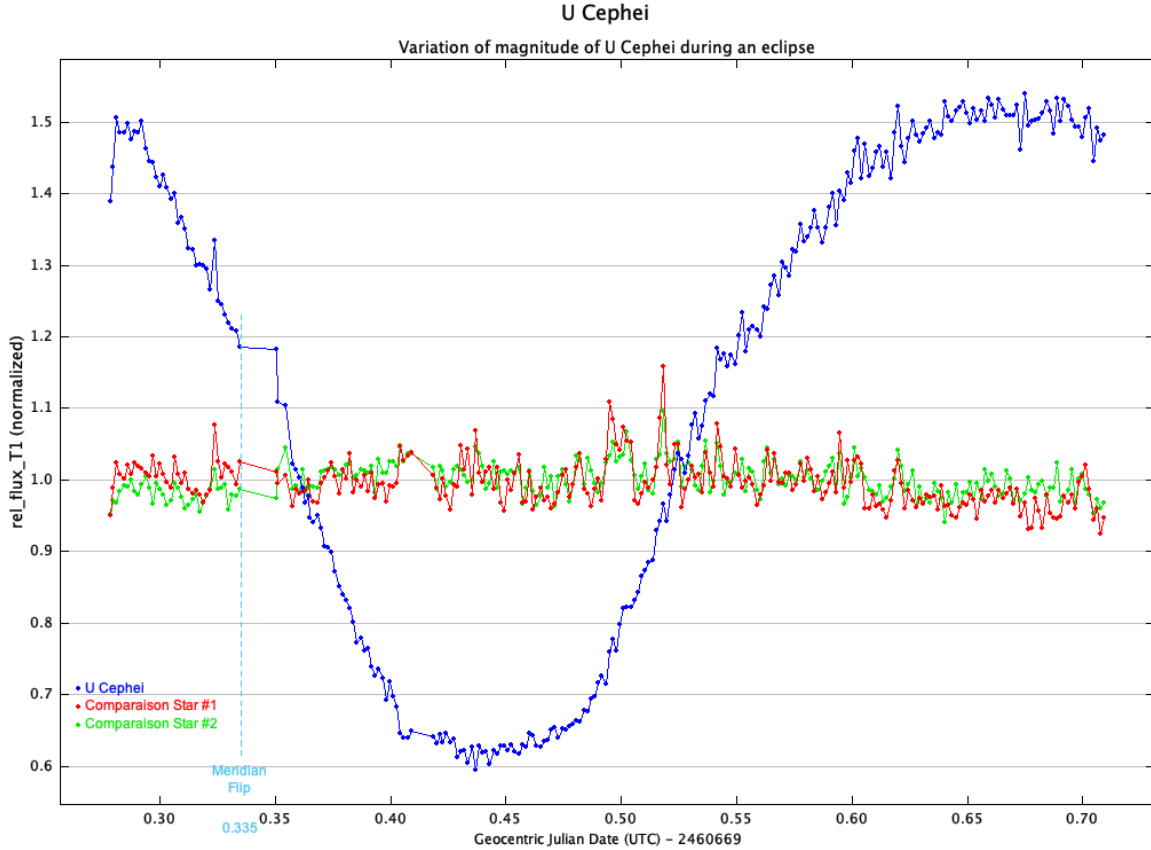


Figure 4: Light curve of U Cephei showing a clear transit event.

In this graph, we see multiple lines: the blue one represents the normalized relative flux of U Cephei, while the red and green ones show the normalized relative flux of the comparison stars. The flux of the comparison stars remains fairly linear around 1, whereas the flux of U Cephei follows a parabolic curve, starting from 1.5, decreasing to 0.6, and then returning to 1.5 over 9.6 hours. This duration is calculated as $0.7 - 0.3 = 0.4$ days, and $0.4 \text{ days} \times 24 = 9.6$ hours.

We also can notice a small gap in the data around the meridian flip since this is not fully automated yet, the telescope had to be flipped manually, which caused a small gap in the data.

5.2 Comparative Analysis

To validate the reliability of the amateur observations, I conducted a comparative analysis with professional data from the Transiting Exoplanet Survey Satellite (TESS) mission.

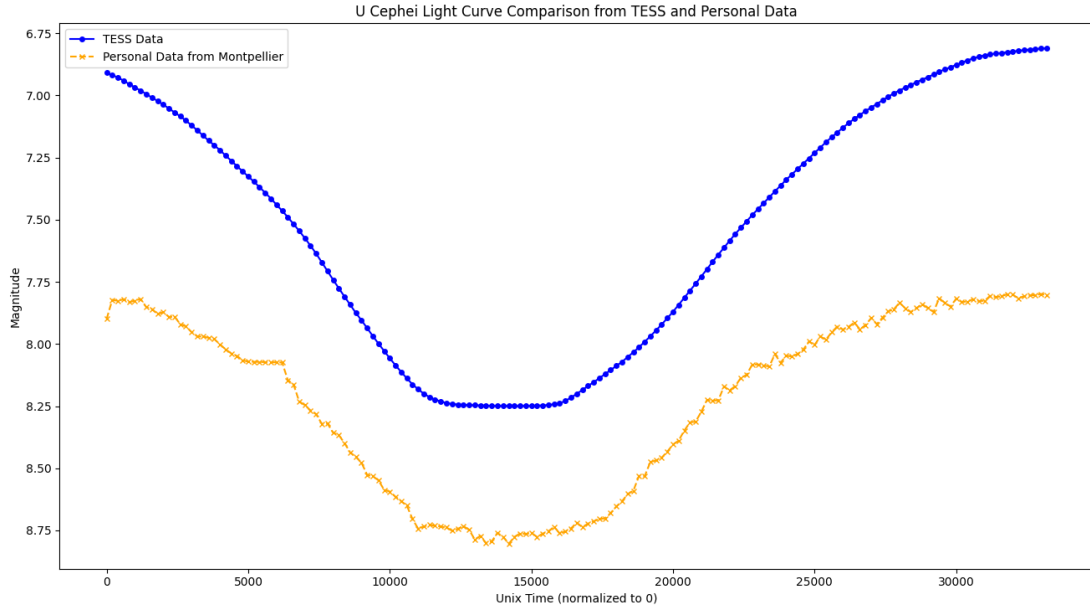


Figure 5: Light curve comparison from TESS data and amateur data.

This graphs shows the light curve of U Cephei from the TESS data in blue and the amateur data I collected in orange. Both curves display a parabolic similar shape. The TESS data light curve actually being a bit brighter, starting around 6.85 and lowering to almost 8.25 than the amateur one, starting at 7.75 and lowering to 8.75.

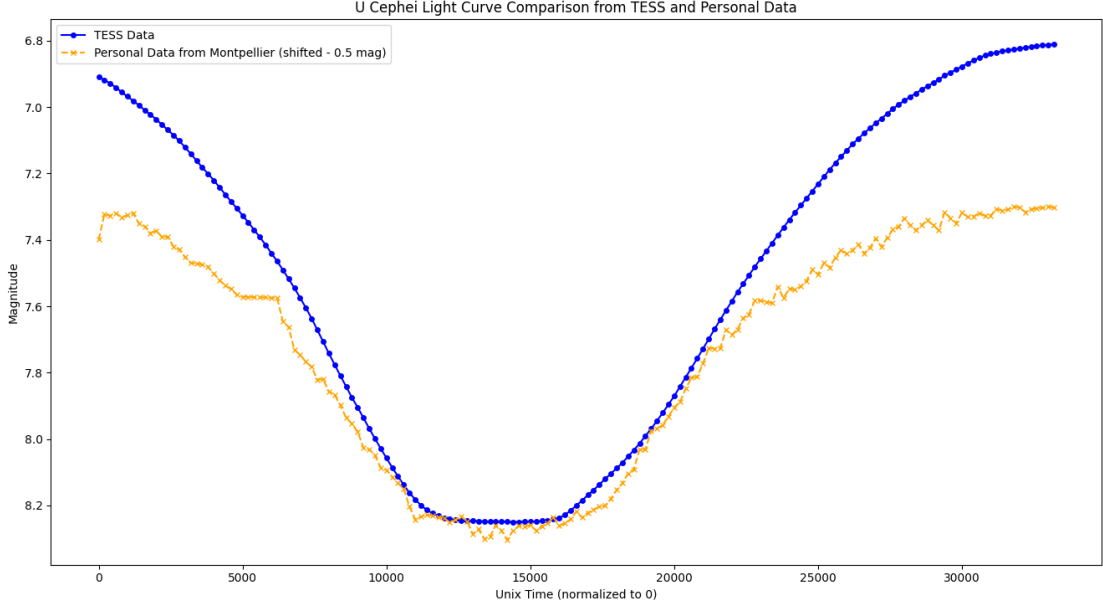


Figure 6: Light curve comparison where the amateur data was shifted by -0.5.

This is another light curve comparison very similar to the previous one, but this time the amateur data has been shifted by -0.5.

6 Discussion

6.1 Figure 1

After our data has been collected and prepared we can plot a light curve which is a graph showing the variation of the flux of the star over time. On the x axis we will have the time of the observation in Julian Date (JD), and on the y axis we have the relative flux of the stars.

To understand a bit more those scales I think it is important to take the time to explain each of them.

- **Relative flux:** In astronomy the light emitted by a star is measured in flux, which is the amount of light received per unit area. Here we are dealing with *normalized relative flux*. The relative flux is basically the flux of a star compared to other, it is very important in our case because of some perturbations that could affect the light received by the camera, such as clouds, or the moon or the atmosphere.

$$\text{Relative Flux} = \frac{\text{Flux of the Target Star}}{\text{Flux of the Comparison Star}} \quad (3)$$

We also normalize the data to have a relative flux of 1 when the star is not in transit, this way we can see the variation of the flux during the transit.

- **Julian Date:** The Julian Date is a continuous count of days since the beginning of the Julian Period on January 1, 4713 BC. It is widely used in astronomy because it is a continuous count of days, allowing for easy comparison of dates and times.

The light curve of U Cephei is showing a clear transit event happening.

This is a great result, it means that the observation was successful and that we managed to catch a transit event, now it is time to compare it to data that was collected by professional observatories.

Only 2 comparison stars are showcased on this graph, but in reality 7 comparison stars were picked to get a more accurate result. When looking at those 2 light curve we clearly see that they remain pretty linear through the night, and normalized around 1, which is a good sign that the data is reliable. The perturbations on the light curves can be explained by the fact that the stars are not at the same altitude in the sky, and that the atmosphere can affect the light received by the camera, as well as the fact this data is collected by an amateur setup.

6.2 Figure 2

One thing we see when looking at those light curves is the fact they are not overlapping. This can be explained because TESS is in space and doesn't have to deal with the atmosphere, making the star look brighter. It could also be explained by the fact TESS sensor is more sensitive than the one used in this study, making the star look brighter. This is why the data is not overlapping, but the transit event is still clearly visible on both light curves.

6.3 Figure 3

To have a better comparison I shifted the amateur data by -0.5, this way both light curves are overlapping. This shift allows us to see that the light curves are following the same path but that the amateur data is just a bit dimmer than the TESS data. The variation of magnitude is different which is something I cannot explain. Even if my data is less precise in theory the variation of magnitude should be the same, but it is not the case here.

This might be explained by a difference of sensitivity of the sensors used, or by the fact that the data was not normalized correctly, but this is something that needs to be investigated further.

7 Conclusion

The personal photometric observations of U Cephei conducted in this study illustrate the potential for amateur astronomers to achieve a high degree of precision, even in light-polluted environments. These observations, while slightly less precise than professional datasets, offer valuable insights and demonstrate the viability of citizen science in advancing binary star research. Further research should focus on determining why we have a magnitude variation difference between the amateur and professional data, therefore we could get data from less polluted areas to see if the difference is due to the light pollution or to the equipment used.

References

- [1] NASA TESS Mission, *The Transiting Exoplanet Survey Satellite (TESS)*, <https://www.nasa.gov/tess>, accessed January 2025.
- [2] American Association of Variable Star Observers, *Variable Star Charts, Data and Resources*, <https://www.aavso.org>, accessed January 2025.
- [3] Siril, *Free Astronomical Image Processing Software*, <https://siril.org/>, accessed January 2025.
- [4] Collins, K. A., Kielkopf, J. F., Stassun, K. G., Hessman, F. V. (2017). *AstroImageJ: Image Processing and Photometric Analysis Software for Astronomical Research*. The Astronomical Journal, 153(2), 77.
- [5] International Dark-Sky Association (IDA), *Understanding Light Pollution*, <https://www.darksky.org>, accessed January 2025.
- [6] Guinan, E. F., Engle, S. G., Devinney, E. J. (2014). *A Satellite Constellation for Photometry and Monitoring of Variable Stars*. Journal of the American Association of Variable Star Observers, 42, 235. <https://ui.adsabs.harvard.edu/abs/2014JAVSO...42..235G/abstract>, accessed January 2025.
- [7] Mobberley, M. (2010). *Amateur Astronomer's Handbook*. Springer, Chapter 1. https://link.springer.com/content/pdf/10.1007/978-1-4471-0055-3_1.pdf, accessed January 2025.
- [8] Mattei, J. A. (1990). *The Contribution of Amateur Astronomers to the Study of Variable Stars*. Proceedings of the International Astronomical Union, 125, 73. <https://www.cambridge.org/core/services/aop-cambridge-core/content/view/6E3175E20C284177D4F534F506B1DD3E/S0252921100092794a.pdf/the-contribution-of-amateur-astronomers-to-the-study-of-variable-stars.pdf>, accessed January 2025.
- [9] Wikipedia, *Variable Star*, https://en.wikipedia.org/wiki/Variable_star, accessed January 2025.