

PERTURBED UNIVERSES AND INFLATIONARY MODELS

MSci. THESIS

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1 Acknowledgments

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2 Abstract

The aim of this project is to investigate inflationary models and discover new models that give near scale invariance in accordance with Cosmic Microwave Background data.

3 Introduction

3.1 Cosmological Principle

3.2 Hubble Law

3.3 Cosmic Microwave Background (CMB)

3.4 Cosmological Problems

3.5 Inflation and why it is needed

3.6 Aims of the Project

4 Theory

4.1 Friedmann Equations

4.2 Geometry

4.3 Cosmological Constant

4.4 Friedmann Models

4.5 Conformal Time

4.6 Mukhanov Sasaki Equation

4.7 Power Spectrum

4.8 Spectral Index

$$P(k) = A \left(\frac{k}{k_0} \right)^{n_s-1} \quad (1)$$

$$n_s - 1 = \frac{\log(P(k))}{\log(k)} \quad (2)$$

4.9 Spectral Index

5 Early Work

5.1 Numerical Work

6 Matching Conditions

6.1 Outline

The matching condition is a simple analytical method that can be used to approximate the spectral index of an inflationary model to a good degree of accuracy so long as $a(\eta)$ is given. This can be achieved by assuming the growing modes that freeze out and exit the horizon are proportional to the scale factor and some function of k . The functional dependence on k can then be calculated by matching the modes at horizon crossing. At horizon crossing the following equation is valid.

$$\left[k^2 - \frac{a''}{a} \right] = 0 \quad (3)$$

$$\frac{1}{\sqrt{k}} \approx a(\eta)F(k) \quad (4)$$

Furthermore

$$P(k) = k^3 \zeta^2 = k^3 \frac{v^2}{a^2} = k^3 F(k)^2 \quad (5)$$

6.2 Constant w

$$a \approx \eta^{\frac{2}{3w+1}} \quad (6)$$

$$n_s(w) = 4 \pm \frac{3(1-w)}{3w+1} \quad (7)$$

6.3 Jacobi Elliptic Function Solutions

The matching condition method can be used to derive an expression for $a(\eta)$ that gives scale invariance. Equation 4 can be expressed purely in terms of the scale

factor and its derivatives with respect to conformal time using equation 3 to give the following.

$$a(\eta) \approx \frac{1}{F \left(\sqrt{\frac{a''}{a}} \right) \left(\frac{a''}{a} \right)^{\frac{1}{4}}} \quad (8)$$

From equations 2 and 5 it can be seen that $F(k) \propto k^{-\frac{3}{2}}$ is required for exact scale invariance. If this condition is imposed in equation 8 a second order differential equation can be found for the the scale factor $a(\eta)$.

$$a''(\eta) \propto a(\eta)^3 \quad (9)$$

It is immediately clear to see that $a \propto \frac{1}{|\eta|}$ is a solution.

7 Hamiltonian Jacobi

7.1 Outline

The evolution of scalar fields is described by the Klein [1] Gordon Equation

$$\ddot{\phi} + 3H\dot{\phi} + V'(\phi) = 0 \quad (10)$$

$$H^2 = \frac{8\pi}{3m_{pl}^2} \left[\frac{1}{2} \dot{\phi}^2 + V(\phi) \right] \quad (11)$$

8 Generalised Equation of State

8.1 $n_s = 1$ (Intermediate Inflation)

$$H''H - 2(H')^2 = 0 \quad (12)$$

$$H(\phi) = \frac{\alpha}{\beta + \phi} \approx \frac{1}{\phi} \quad (13)$$

8.2 $n_s \neq 1$

$$\left(\frac{H''}{2}\right) - \left(\frac{H'^2}{H}\right) + (n_s - 1) \pi H = 0 \quad (14)$$

A trial solution of $H(\phi) = \exp(C\phi)$ can be shown to satisfy equation 14 with $C = 5\sqrt{2}$. This is the same as Power Law Inflation.

8.3 Proof of Slow roll

8.4 Equation of State

$$\rho_\phi = \left(\frac{8\pi(n_s - 1)}{9A^2}\right) p_\phi^2 - \left(\frac{n_s + 2}{3}\right) p_\phi \quad (15)$$

8.5 Potential

8.6 Number of e-foldings

9 Conclusion

References

- [1] Herbert Goldstein. *Classical mechanics*. Addison-Wesley, Reading, Mass. ; Wokingham, 2nd ed. edition, 1980.