Adaptive cancellation of power grid interference in continuous gravitational wave searches with a hidden markov model

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Continuous gravitational wave (GW) searches are commonly hampered by long-lived, narrow peaks in the instrumental frequency spectra known as 'lines'. Candidate GW signals which lie within frequency bands with known lines are typically vetoed. In this work we demonstrate a new line subtraction method based on adaptive noise cancellation (ANC). ANC is common in electrical engineering applications, including for processing audio and biomedical signals. We show how ANC can be used in conjunction with a Viterbi search to track spin-wandering continuous wave signals near the LIGO 60 Hz line.

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I. INTRODUCTION

Instrumental noise artifacts in gravitational wave (GW) searches with terrestrial, long-baseline interferometers are classified according to their duration and spectral properties. Short-lived, non-stationary recurrent noise events such as optomechanical glitches typically last for seconds and exhibit distinctive spectral signatures, e.g they can be chirp-like [1–5]. Long-lived, quasi-stationary, broadband noise sources include seismic disturbances at low frequencies, test mass thermal fluctuations at intermediate frequencies, and photon shot noise at high frequencies [6–10]. Long-lived narrowband spectral artifacts termed instrumental lines are caused by electrical subsystems (e.g. mains power, clocks, oscillators), mechanical subsystems (e.g. test mass and beam-splitter violin modes) and calibration processes [11], although often the origin of a specific feature is unknown. Instrumental lines are disruptive, especially for continuous wave (CW) searches where the target astrophysical signal is quasi-monochromatic and resembles the noise artifact spectrally. Many above-threshold candidates discovered in CW searches to date are vetoed because they coincide with known instrumental lines [12–14], for instance, in CW searches involving data from Observing Run 3 (O3) with the Laser Interferometer Gravitational Wave Observatory (LIGO), Virgo and the Kamioka Gravitational Wave Detector (KAGRA) [e.g. 15–17]

Several techniques have been implemented by the LIGO-Virgo-KAGRA (LVK) collaboration to identify, characterize and suppress instrumental noise artifacts [18, 19]. Some techniques identify and veto an artifact (or gate segments of the data) based on its time-frequency signature [20–22]. Other techniques perform offline noise subtraction with reference to auxiliary data from physical environmental monitors (PEMs) [23–26]. PEMs can be used to witness correlated noise and generate a reference signal directly or elucidate and quantify multichannel couplings [27, 28]. Finally some techniques are based on machine learning [29–32]. In CW searches specifically, the distinctive amplitude and frequency (Doppler) modulations associated with the Earth's rotation and revolution can be exploited to discriminate between terrestrial noise artifacts and astrophysical signals [33, 34].

In most of the situations above, the practical effect of an instrumental line is to excise the relevant part of the observing band from a CW search. That is, if an above threshold CW search candidate coincides with a known instrumental line, the candidate is vetoed under current practice without further analysis, such as comparing the expected strength of the noise line with the measured strength of the candidate [1]. In this paper we take a first step towards lifting the above limitation. We introduce an adaptive noise cancellation (ANC) scheme based on an adaptive recursive least squares (ARLS) method which suppresses narrowband noise proportional to a known PEM reference signal. We then apply the ANC scheme to a CW search algorithm based on a hidden Markov model (HMM) which detects and tracks quasi-monochromatic GW signals with wandering frequency and has been tested and validated thoroughly in multiple LVK searches [12–14, 35]. We demonstrate with synthetic data that the ANC scheme and HMM algorithm together can successfully detect a GW signal lying under the 60 Hz mains power line, if the signal exceeds a well-defined minimum amplitude. The approach extends naturally to other instrumental lines, a topic for future work.

The paper is organized as follows. In Section II we outline our mathematical model for the 60 Hz mains power spectral line and the PEM reference signal, and justify the assumptions of the model by reference to the strain and environmental data from the LIGO Livingston interferometer. In Section III we introduce ANC formulated as an ARLS method. We go on in Section IV to deploy the ANC method in conjunction with a HMM Viterbi

[1] A regularly updated log of narrowband instrumental lines in the LVK detector is maintained at dcc.ligo.org/LIGO-T2100200/public for public reference.

solver applied to synthetic GW strain data and demonstrate the successful recovery of a monochromatic GW signal.

II. POWER GRID INTERFERENCE

The goal of this paper is to detect a quasi-normal GW signal in a data stream contaminated by two kinds of noise: additive Gaussian noise which is fundamental and irreducible, and additive non-Gaussian interference from a long-lived narrow spectral feature which can be filtered out in principle given an accurately measured reference signal. For this work we consider the spectral line at 60 Hz that results from the North American alternating current power grid as the additive non-Gaussian interference. In Section II A we briefly review the 60 Hz LIGO interference line before proceeding in Section II B to specify the assumed mathematical forms of the interference and reference signals. In Section II C we justify the assumptions of Section II B by analysing differential arm length (DARM channel) and environmental (power grid monitoring) data from the LIGO interferometers.

A. LIGO 60 Hz Interference

LIGO data contains multiple long-duration narrow lines (see Fig. 1) in addition to the usual Gaussian noise. The provision of mains power electricity in North America via an alternating current with frequency 60 Hz leads to a line in the LIGO data at the corresponding frequency. The coupling between the mains power and the gravitational wave data channel can occur since the performance of the high-sensitivity electronic components within LIGO varies with respect to the input power voltage. Additionally, the magnetic fields that arise from the AC mains supply can couple with the magnets on the LIGO optical components. Whilst some spectral lines are static, the 60 Hz line wanders in time, due to variations in the load in the North America power grid at any one time (e.g. Fig. 2). This time variation can further impact the detector sensitivity over a broader frequency band. For a full review of LIGO spectral artifacts, including the 60 Hz line, we refer the reader to [11].

B. Statement of the problem: signal and noise models

Let $x(t)$ denote the scalar time series output by the "science" or strain channel of a LIGO-like long baseline interferometer. Suppose that $x(t)$ is sampled at discrete times t_n, with $1 \leq n \leq N$ and uniform sampling interval $\Delta t = t_n - t_{n-1}$. Let $r(t)$ denote the scalar time series output by the environmental channel relevant for filtering interference; here $r(t)$ is of one of the three phases of mains power measured at some reference point

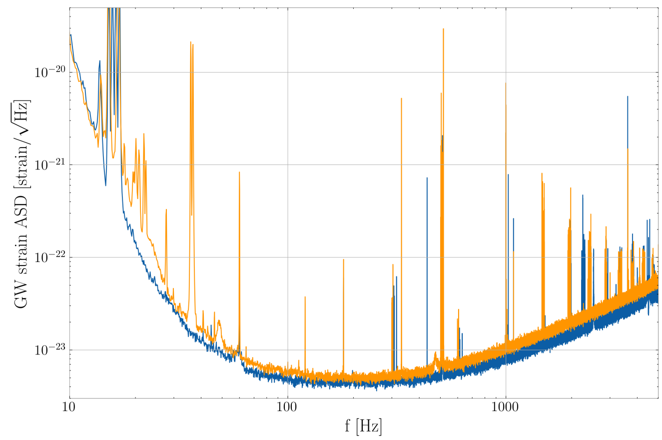

FIG. 1. Sensitivity plot for LIGO-Hanford (orange) and LIGO-Livingston (blue) for a snapshot of data (~10 minutes) from O3 (channel `*:DCS-CALIB_STRAIN_C01_AR`, see Ref. [36, 37]). The spectral line at 60 Hz is clearly visible, along with multiple other instrumental lines at other frequencies.

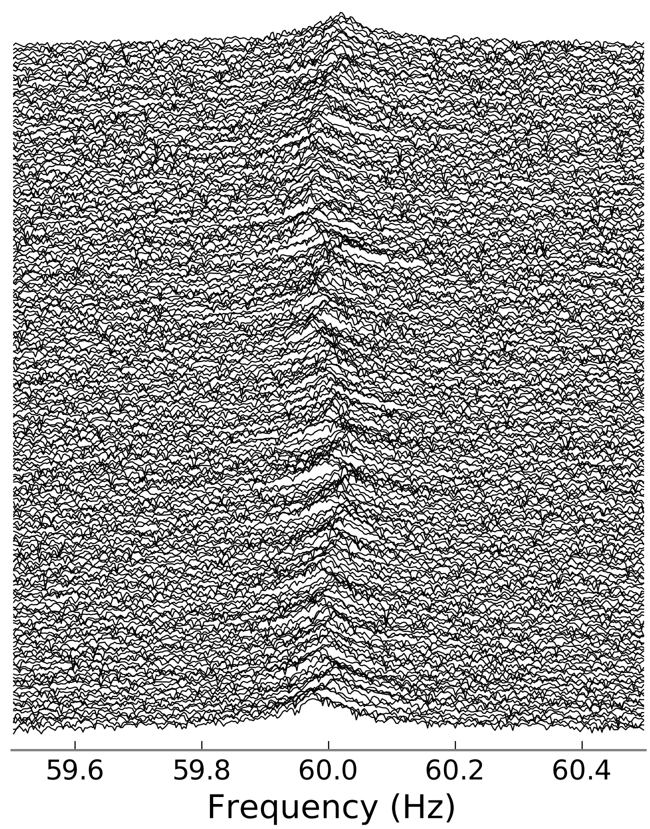

FIG. 2. Cascade plot showing the amplitude spectral density of the Hanford PEM monitor `H1:PEM-CS_MAINSMON_EBAY_1_DQ` (corner station, phase 1) over time. Each trace corresponds to 320 s (≈ 5 min) of data (210 lines plotted). The wandering of the 60Hz instrumental line about its central value can be seen.

in the detector. The reference signal is usually sampled less frequently than $x(t)$ at discrete times t_{n_k}, with $1 \le k \le K$ and $1 \le n_k \le N$. We assume for the sake of convenience that every t_{n_k} coincides with some t_n for all k, but the condition is not essential.

The strain channel is composed of a gravitational wave signal $h(t)$, non-Gaussian interference $c(t)$ (sometimes called "clutter") and Gaussian noise $n(t)$ in a linear combination:

$$x(t) = h(t) + c(t) + n(t) \ . \tag{1}$$

In this paper the gravitational wave signal takes the form predicted by Jaranowski *et al.* [38] for a biaxial rotor, e.g. a neutron star (NS) emitting continuous gravitational waves at multiples of the star spin frequency $f_\star$. The GW signal is quasi-monochromatic, amplitude-modulated by the rotation of the Earth and frequency-modulated by the Earth's orbital motion. The noise $n(t)$ is white with $\langle n(t_n) n(t_{n'}) \rangle = \sigma_n^2 \delta_{nn'}$. Noise samples $n(t_n)$ are drawn from a Gaussian distribution with zero mean and variance σ_n^2. The interference clutter $c(t)$ takes a form determined by instrumental processes but is generally a long-lived narrow spectral feature. We can relate $c(t)$ to the instrumental voltage $r(t)$ in the following.

Mains power is characterized by three properties. First, the frequency is maintained at a constant value across the grid to a good approximation by internal grid mechanisms (effectively a phase locked loop), with central frequency $f_{\rm ac} = 60$ Hz in North America. A slow periodic modulation occurs around $f_{\rm ac}$ with a small amplitude $\Delta f_{\rm ac} \lesssim 0.5$ Hz and period P which wanders randomly and uniformly in the range $0 \le P \le P_{\rm max}$. Secondly, the phase $\Theta(t)$ of the voltage $r(t)$ wanders stochastically. We assume that the phase noise is white and Gaussian, with $n_\Theta(t_n)$ drawn from a Gaussian distribution with zero mean and variance σ_Θ^2, and $\langle n_\Theta(t_n) n_\Theta(t_{n'}) \rangle = \sigma_\Theta^2 \delta_{nn'}$. Third, the voltage amplitude, $A_r(t)$, is random. We assume that samples $A_r(t)$ are distributed uniformly within $[A_{\rm ac} - \Delta A_{\rm ac}, A_{\rm ac} + \Delta A_{\rm ac}]$. We can then write the reference voltage as,

$$r(t) = A_r(t_n) \cos\left[2\pi f_{\rm ac} t + \Theta(t)\right] + n_r(t_n) \ , \tag{2}$$

with

$$\Theta(t) = 2\pi \Delta f_{\rm ac} \cos\left(\frac{2\pi t}{P(t_n)}\right) + n_\Theta(t_n) \ , \tag{3}$$

for $t_n \le t \le t_{n+1}$. That is, at time t_n, random variables $A_r(t_n)$, $P(t_n)$ and $n_\Theta(t_n)$ are drawn from the distribution $\mathcal{U}[A_{\rm ac} - \Delta A_{\rm ac}, A_{\rm ac} + \Delta A_{\rm ac}]$, $\mathcal{U}[0, P_{\rm max}]$, and $\mathcal{N}[0, \sigma_\Theta^2]$ respectively. Equation (2) then runs forward over an interval of length Δt. Hence $r(t)$ is discontinuous at each sampling time. In Equation (2), $n_r(t_n)$ is the reference signal measurement noise at t_n, assumed to be white and Gaussian with $r_r(t_n)$ drawn from a Gaussian with zero mean and variance σ_r^2. All the white measurement and process noises are assumed independent.

Mains power couples into the strain channel in various complicated ways, e.g. through electronic devices, or inductively through ambient magnetic fields. A central assumption in this work is that the interference in the strain channel is an exact, amplitude-scaled replica of the reference signal up to a delay $\tau_{\rm delay}$ which is attributed to spatial propagation effects between the reference measurement front and the interferometer mirrors or dark front. Hence we can express the interference clutter as,

$$c(t) = A_c(t_n) \cos\left[2\pi f_{\rm ac} t' + \Theta(t')\right] \ , \tag{4}$$

for $t_n \le t \le t_{n+1}$ and $t' = t_n - \tau_{\rm delay}$. The amplitude A_c is distributed as $\mathcal{U}[A_{\rm ac}, A_{\rm ac} + \Delta A_{\rm ac}]$. For this work we consider $0 \le \tau_{\rm delay} \le 10\Delta t$, but wider or narrower ranges is possible and straightforward to be implemented. The assumptions of an exact replica between the interference and the reference is tested in the following section. [2]

C. Cross-correlating the interference and reference

A key assumption of the construction in Section II B is that the 60 Hz noise that is recorded in the reference PEM channel is also present in the LIGO strain channel. That is, the noise recorded in the PEM channel is imprinted onto the strain channel. In order to test this assumption we cross-correlate the strain channel and the PEM channel. If there is a noise signal at 60 Hz present in both channels then it should be revealed by this cross-correlation. We use open sourced data for the strain and PEM channels from the first part of the third LIGO observing run, O3a [36]. This data is obtained via the Gravitational Wave Open Science Center [3] using the `GWPy` package [39]. In the auxiliary O3a data there are 9 PEM channels at LIGO-Livingston and 7 PEM channels at LIGO-Hanford [4]. For this work we consider just the LIGO-Livingston data. The strain data is obtained at a rate of 16384 Hz and downsampled to 4096 Hz to match the sampling rate of the PEM channels. In Figure II C we show the coherence between each of the 9 PEM channels and the high latency, calibrated strain channel `L1:DCS-CALIB_STRAIN_C01_AR` over a 10 minute time period. For all channels there is a clear coherence feature at 60Hz. The coherence feature is typically large, with values > 0.5 for 7 of the 9 PEM channels. Compared to the other channels, the coherence is particularly weak for the channels `L1:PEM-EY_MIC_VEA_PLUSY_DQ` and `L1:PEM-CS_MIC_LVEA_INPUTOPTICS_DQ`. These channels corresponds to microphone PEMs in the LIGO vacuum

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2 TK: Why are the amplitudes of A_r and A_c distributed differently?
3 gwosc.org
4 https://git.ligo.org/gwosc/tutorials/gwosc-aux-tutorials/-/tree/main/Channels

equipment area and so would be less sensitive to the 60Hz noise than e.g. `L1:PEM-EY_MAINSMON_EBAY_1_DQ` which directly measures the voltage. In Figure 4 we also show the coherence spectrogram i.e. the time-varying coherence between the strain channel and the PEM channel `L1:PEM-CS_MAINSMON_EBAY_1_DQ`. The coherence is calculated across 10 s windows of data. That is, each column of the spectrogram has a width of 10 s. We again observe similar results to Figure II C where there is a clear, strong spectral feature at 60 Hz. Additional coherence features can also be seen at at ~ 200 Hz and 300 Hz, corresponding to different frequency noise lines recorded by the PEM that also imprint onto the strain channel. These results showing a strong coherence between the strain and PEM channels demonstrate that our assumptions of Section II B are justified.

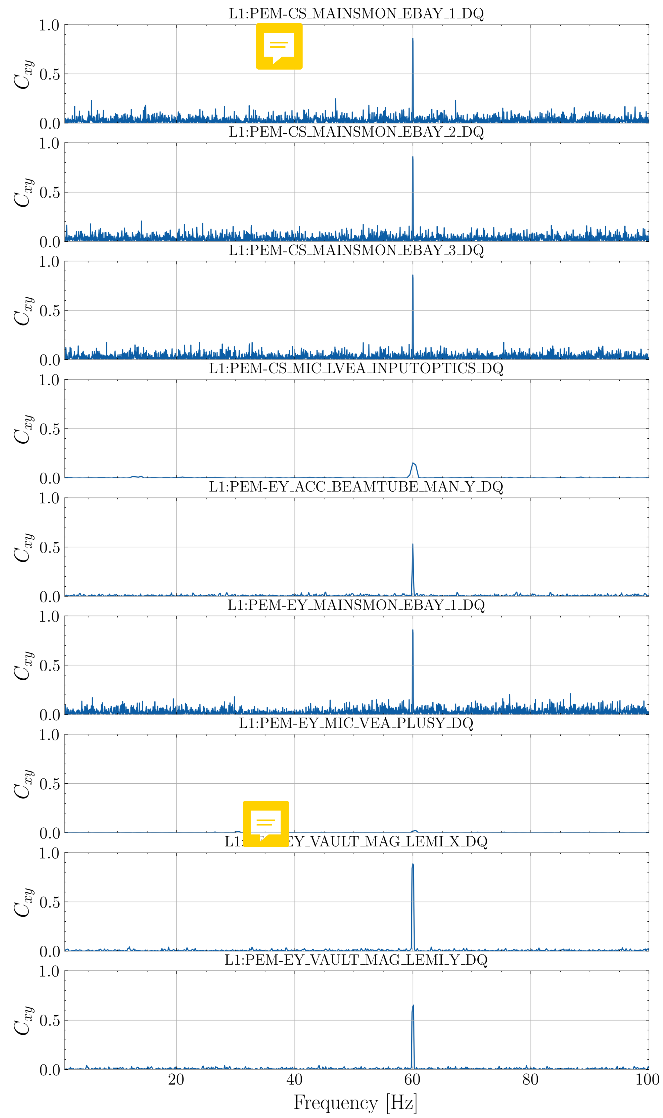

FIG. 3. Coherence C_{xy} between the LIGO-Livingston strain channel `L1:DCS-CALIB_STRAIN_C01_AR` and the PEM channels over a 10 minute observation period. Clear features at 60 Hz are present in all of the channels. The coherence is weakest for those PEM channels which do not measure voltage directly, but instead are microphones the the LIGO vacuum equipment area.

III. ADAPTIVE NOISE CANCELLATION

Adaptive Noise Cancellation (ANC) is a method for recovering an estimate of an underlying signal which has been corrupted or obscured by some additive noise interference [40]. In contrast to other common optimal filtering methods (e.g. Wiener, Kalman) ANC requires no a-priori knowledge of either the signal or the noise. Instead, ANC makes use of a reference input which is correlated in some unknown way to the noise in the primary signal. This reference can then be filtered and subtracted from the primary data series so as to recover the underlying signal. For our purposes, the primary timeseries is the gravitational-wave strain channel $x(t)$ (Equation (1)) and the reference is a PEM recording voltage data from the power grid $r(t)$ (Equation (2)). The objective of ANC is then to remove the clutter $c(t)$ of Equation (1) with the aid of $r(t)$ whilst leaving $h(t)$ i.e. the signal of interest, intact. We now describe the ANC implementation used in this work.

The reference $r(t)$ is used to construct an estimate of the clutter, $\hat{c}(t)$. The clutter estimate can then be subtracted from the primary signal in the time domain, defining a residual $e(t)$:

$$e(t) = x(t) - \hat{c}(t) \ . \tag{5}$$

As $\hat{c}(t) \to c(t)$, $e(t) \to h(t) + n(t)$ and we recover a noise cancelled timeseries. The clutter estimate is modelled by a finite duration impulse response (FIR) filter:

$$\hat{c}_k = \mathbf{w}^\intercal \mathbf{u}_k \ . \tag{6}$$

where $\hat{c}_k$ is the clutter estimate at a discrete timestep k (i.e. $c_k = \hat{c}(t_{n_k})$), $\mathbf{u}_k$ is the tap-input vector composed of M running samples of the reference signal arranged backwards in time:

$$\mathbf{u}_k \ = [r_k, r_{k-1}, \dots, r_{k-M+1}] \ , \tag{7}$$

and $\mathbf{w}$ is the tap-weight vector:

$$\mathbf{w} = [w_1, w_2, \dots, w_M] \ . \tag{8}$$

We want to determine the optimal tap weights $\mathbf{w}_{\rm opt}$, those which minimise the mean square error cost function

$$\mathbf{w}_{\rm opt} = \arg\min \sum_{t=t_1}^{t_2} |e(t)|^2 \ . \tag{9}$$

This minimization problem can be solved by way of an

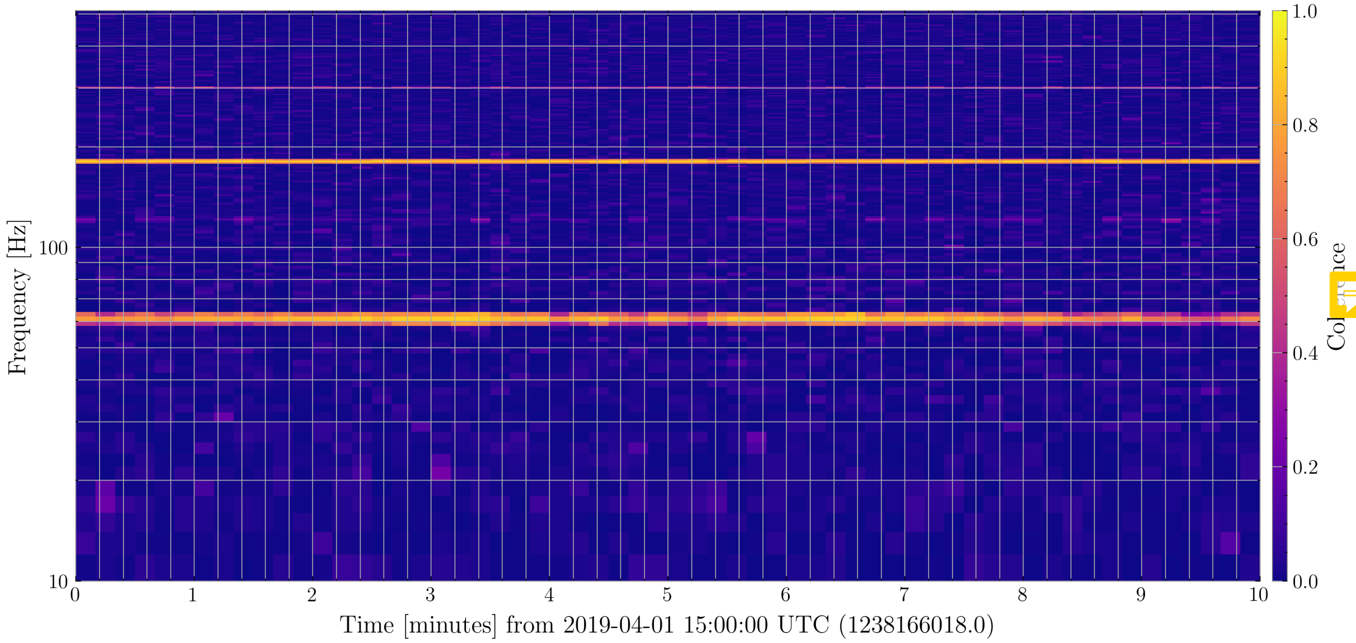

FIG. 4. Coherence spectrogram between the strain channel `L1:DCS-CALIB_STRAIN_C01_AR`and the PEM channel `L1:PEM-CS_MAINSMON_EBAY_1_DQ` over a 10 minute period of O3 data. A strong coherence is observed at 60Hz due to the main power interference. Additional coherence can also be observed at $\sim$ 200 Hz and 300 Hz due to instrumental lines of a different provenance c.f. Fig 1.

adaptive recursive least squares (ARLS) method which we now describe in Section III A

A. Adaptive Recursive Least Squares Method

We now outline the ARLS method to compute the optimal tap weights and estimate the noise-subtracted signal.

1. Initialise the tap weights $\mathbf{w} = \mathbf{0}$ and a covariance matrix $\mathbf{P} = \delta^{-1}\mathbf{I}$, for regularisation parameter $0 < \delta \ll 1$ and identity matrix $\mathbf{I}$ of rank M.
2. For $k = 1, \ldots, K$:
   (a) Estimate the clutter $\hat{c}_k$ by Equation 6
   (b) Calculate the residual e_k by Equation 5
   (c) Calculate the gain vector
   $$\mathbf{g}_k = \frac{\mathbf{P}\mathbf{u}_k}{\lambda + \mathbf{u}_k^{\intercal}\mathbf{P}\mathbf{u}_k} \tag{10}$$
   (d) Update the tap weights
   $$\mathbf{w} \mathrel{+}= e_k \mathbf{g} \tag{11}$$
   (e) Update the covariance matrix
   $$\mathbf{P} = \lambda^{-1}\mathbf{P} - \mathbf{g}\lambda^{-1}\mathbf{P}\mathbf{u}_k^{\intercal} \tag{12}$$

The algorithm is also illustrated in Figure 5 via a block diagram.

ARLS has two free parameters: the order parameter M and the "forgetting factor" λ, which is chosen so as to give exponentially less weight to older samples. The choice of M influences the latency of the FIR filter, the computational overhead and the filter accuracy. In Sec. IV we trial a selection of M values for synthetic GW data. The forgetting factor λ ranges between 0 and 1, with $\lambda = 1$ corresponding to infinite memory, causing the filter becomes an ordinary least squares method. In this work we use $\lambda = 0.9999$. We refer the reader to Chapter 9 of Ref. [41] for a full review of adaptive least squares estimation in the context of linear filtering.

IV. HMM VALIDATION TESTS

We want to test the ANC method described in the preceding section by trying to search for a continuous wave signal which has an initial frequency $f_{\mathrm{GW}}(t = 0) = 60$Hz, coincident with the instrumental line from the mains power grid. We search for the CW signal using a HMM scheme based on the Viterbi algorithm [42, 43]. Specifically, we use the method introduced by Suvorova *et al.* [35], which has been thoroughly tested through multiple LVK searches [12–14]. We do not cover any details of the HMM scheme in this work and refer the reader to

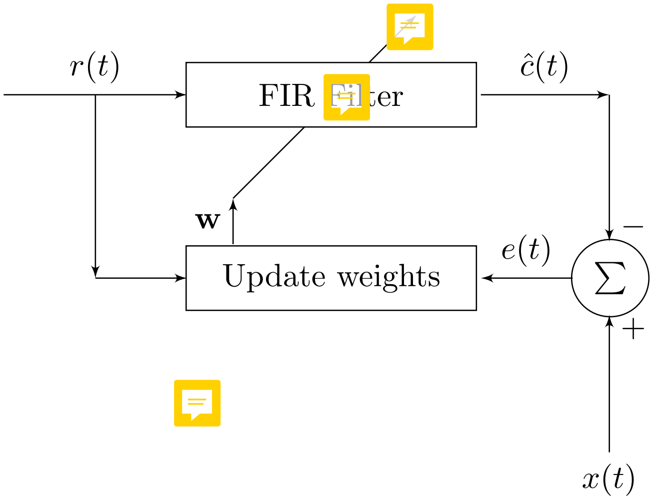

FIG. 5. Block diagram of the adaptive recursive least squares method described in Section III A. The reference signal $r(t)$ is passed to an FIR filter to construct an estimate of the signal clutter $\hat{c}(t)$. Subtracting the clutter estimate from the primary signal $x(t)$ provides a residual $e(t)$ which can then be used to update the weights $\mathbf{w}$ of the FIR filter and the method proceeds iteratively.

Suvorova *et al.* [35] for further information. The GW signal is obscured in the data by the presence of the mains power instrumental line. The challenge is to try to recover the signal by first passing the data $x(t)$ through the ANC filter before using the HMM tracker on this filtered dataset to try to recover the signal. In Section IV A we describe how we create a synthetic dataset $x(t)$. In Section IV B we introduce a representative example, illustrating the frequency tracking before and after ANC. In Section IV C we explore the performance of the ANC filter in response to different characteristics of the interference signal, e.g. variations in the central frequency $f_{\rm ac}$ relative to the initial GW frequency. In Section IV D we investigate the performance of the filter in response to different filter settings; the order of the filter M (i.e. the number of taps) and the number of reference channel inputs.

A. Creating simulated data

Throughout this paper we work with some representative synthetic data, $x(t)$, rather than directly with the LIGO data itself. We can construct synthetic data for each of the constituent parts of Equation (1) - the GW signal of interest, the Gaussian noise and the noise clutter via the formulation described in Section II B as follows.

The GW frequency evolves in general due to the intrinsic evolution of the due to the intrinsic evolution of the source. For this paper we assume that the source is isolated (i.e. it is not in a binary) and that the GW is monochromatic (i.e. all temporal derivatives of the frequency are zero). We also consider the GW source to be at a constant location with respect to the observer and neglect all contributions due to e.g. the rotation and revolution of the Earth. Under these assumptions the GW model reduces to

$$h(t) = h_0 \sin(2\pi\phi_{\rm gw}(t)) \ , \tag{13}$$

where h is the constant GW amplitude and $\phi(t)$ a random phase variable which is the integral of the underlying, piecewise linear GW frequency $f_{\rm gw}$ i.e.

$$\phi_{\rm gw}(t) = \int_0^t f_{\rm gw}(s)\, ds \ . \tag{14}$$

The GW frequency at discrete timestep m within the sampling interval Δt is labelled as $f_{\rm gw}^{(m)}$ and evolves according to

$$f_{\rm gw}^{(m+1)} = f_{\rm gw}^{(m)} + \delta_m \Delta t \ , \tag{15}$$

with δ_m a zero-mean Gaussian noise at timestep m, with variance σ_f^2,

$$\delta_m = \mathcal{N}(0, \sigma_f^2) \tag{16}$$

The synthetic GW signal $h(t)$ is then completely described by the parameters h_0, σ_f^2, and the initial GW frequency $f_{\rm gw}(t=0)$

The clutter and reference signal evolve according to Equations (2) and (4) respectively. For this initial study we take the reference voltage to have a constant amplitude $A_r(t) = a_r$, the clutter to have a corresponding constant amplitude $A_c(t) = a_c$. We also assume that the modulation in the reference voltage about $f_{\rm ac}$ has a constant amplitude $\Delta f_{\rm ac}$ with a constant period $P(t) = P$. Throughout this work we take $\Delta f_{\rm ac} = 1/2\pi$ Hz [5]. We define the variable $\gamma = P^{-1}$ and explore different values for γ in Section IV C. Under these assumptions, Equations (2) and (4) reduce to

$$\begin{aligned} r(t) = a_r \cos\left[2\pi f_{\rm ac} t + 2\pi \cos\left(2\pi\gamma t\right) + n_\Theta(t)\right] \\ + n_r(t) \ , \end{aligned} \tag{17}$$

$$c(t) = \ a_c \cos\left[2\pi f_{\rm ac} t' + 2\pi \cos\left(2\pi\gamma t'\right) + n_\Theta(t')\right] \ . \tag{18}$$

The synthetic reference and clutter data are then completely described by the amplitude parameters a_r, a_c, the central frequency $f_{\rm ac}$, the timescale γ and the noise covariances σ_Θ^2, σ_r^2. For convenience we reparametrise $f_{\rm ac}$

[5] TK: I have guessed that this is what is happening under the hood of the code, but need to verify with Sofia/Changrong.

relative to the initial GW frequency at $t = 0$, defining the new variable

$$\Delta f = |f_{\rm ac} - f_{\rm gw}(t = 0)| \tag{19}$$

The 9 free parameters of the model are summarised in Table I. We have 3 amplitude parameters h, a_r, a_c for the GW, reference and clutter respectively, with $h \ll a_c, a_r$. There are 4 noise parameters $\sigma_n, \sigma_\Theta, \sigma_r, \sigma_f$ for the Gaussian noise $n(t)$, the voltage phase noise $n_\Theta(t)$, reference signal measurement noise $n_r(t)$ and the GW frequency noise Equation 16 respectively. Additionally we have the absolute difference between the central frequency and the initial GW frequency, Δf and the timescale of the modulation in the central frequency,γ.

Throughout this work when creating synthetic data, $f_{\rm ac}$ is fixed at 60 Hz. The terrestrial noise parameters, σ_n^2,σ_r^2,σ_Θ^2 are also fixed. TK: other summary text on choice of parameters for synthetic data to go here. Waiting on input from Sofia on which parameters were actually used for the data

B. Representative example

In order to demonstrate the effectiveness of ANC in conjunction with an HMM Viterbi search in this section we consider two representative examples where $h = 0.025$ and the stochastic GW frequency wandering is either 'low', $\sigma_f^2 = 0.01$ Hz $s^{-1/2}$ or 'high', $\sigma_f^2 = 0.1$ Hz $s^{-1/2}$. All other free parameters of the model are as specified in Table I. At this stage we assume that we have just one single reference PEM channel.

Initially we verify that the ANC filter works as expected to remove the excess power from the 60Hz interference. In Figure 6 we show the Fourier amplitude of the synthetic data $x(t)$, the underlying GW signal $h(t)$, and the signal after being passed through the ANC filter, $e(t) = x(t) - \hat{c}(t)$ across the frequency range 58 - 62 Hz, for the case where $\sigma_f^2 = 0.01$ Hz. Before filtering the Fourier spectrum of $x(t)$ has multiple modes about

Parameter	Physical meaning	Injected Value
h	Strain amplitude	-
a_r	Voltage amplitude	-
a_c	Clutter amplitude	-
σ_n^2	Gaussian noise covariance	-
σ_r^2	Voltage measurement noise	-
σ_Θ^2	Voltage phase noise	-
σ_f^2	GW frequency noise	-
Δf	Central frequency shift	-
γ	Modulation frequency	-

TABLE I. Summary of parameters used to create synthetic data for testing the ANC method in Section IV, their physical meaning, and the injected values used throughout this work.

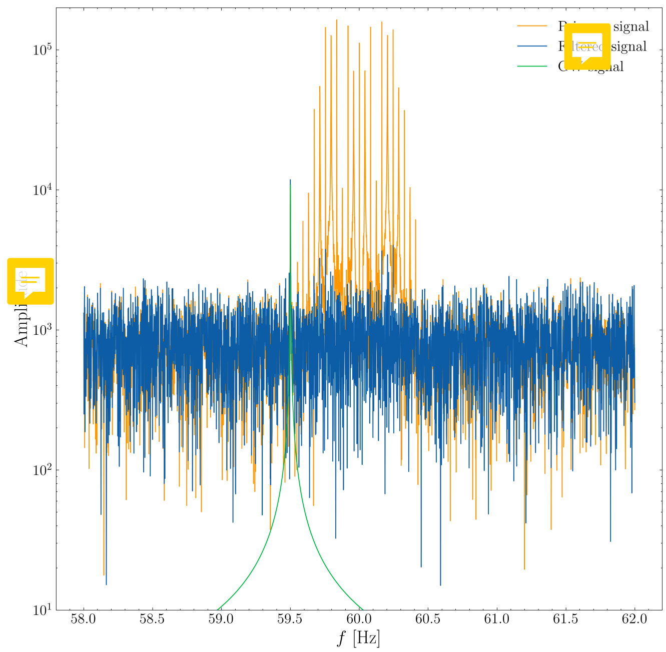

FIG. 6. Fourier response of the data $x(t)$ (orange), the GW signal $h(t)$ (green) and the ANC filtered data $x(t) - \hat{c}(t)$ (blue) for a system with parameters described in Table I. ANC filters out the excess power that results from the interference signal about the central 60 Hz frequency.

the central 60 Hz frequency as a result of the interference clutter. This clutter obscures the power from the GW signal. After filtering, this excess power is removed and the Fourier spectrum of $e(t)$ is flat with the exception of a clear feature coincident with the central frequency of the injected GW.

Having established the ability of the ANC method to filter out the interference clutter given a reference signal, we can deploy the ANC in conjunction with the Viterbi HMM. We pass the ANC filtered data to the HMM and evaluate the performance of the HMM in tracking the spin-wandering continuous wave signal. The results are shown in Figure 7 for the case of both low and high frequency wandering, for a single realisation of the noise. The figure shows the Fourier amplitude spectrogram of the data $x(t)$ before and after the application of ANC filtering. The spin wandering of the GW source (green/orange lines) and the Viterbi estimate (dashed orange/yellow lines) of the spin wandering is superimposed onto the spectrogram. In the low noise case the GW spin frequency wanders close to, but below, the 60Hz interference line. In the high noise case the GW spin frequency wanders much more strongly over a larger range of frequencies and crosses the interference line, presenting a more difficult challenge for the Viterbi tracking algorithm. Before ANC there is a clear feature in the Fourier spectrogram corresponding to the 60 Hz interference signal. In this case the Viterbi algorithm is unable to track

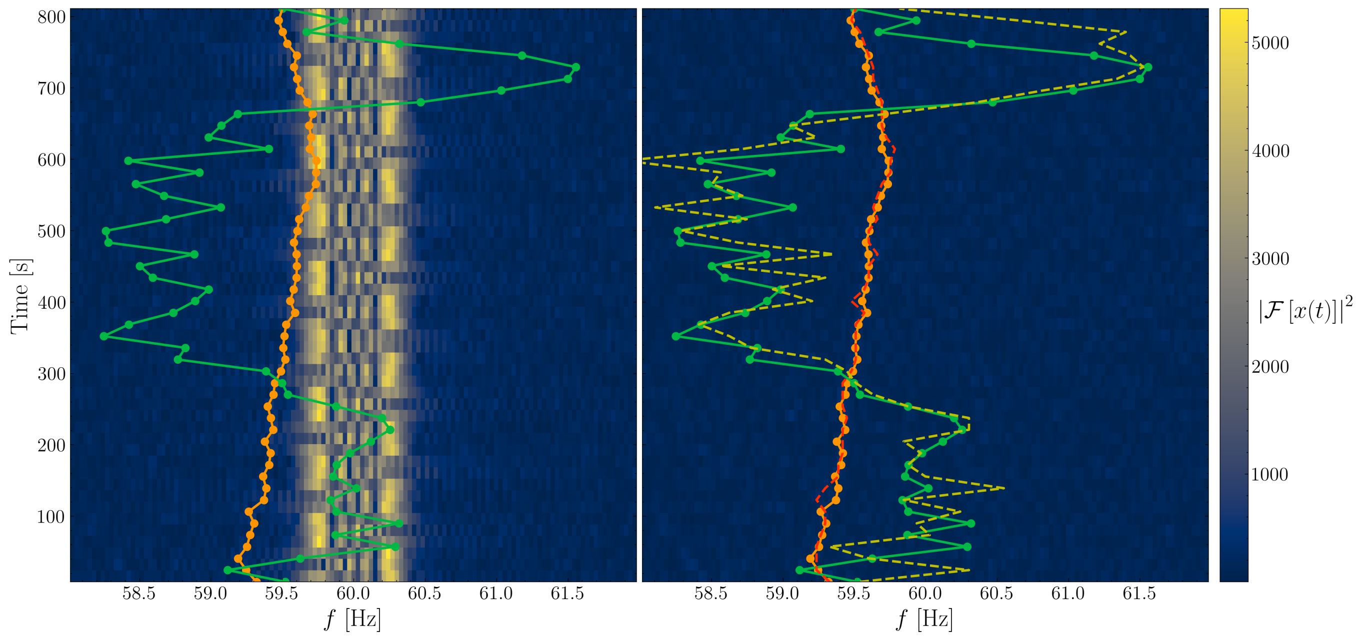

FIG. 7. Fourier amplitude spectrogram and tracking of the frequency evolution of a continuous GW with $h = 0.025$ and $\sigma_f^2 = \{0.01, 0.1\}$ Hz $s^{-1/2}$ (orange and green lines respectively) using a HMM Viterbi algorithm. *Left panel:* before applying ANC filtering to remove the interference signal centred at 60Hz, *Right panel:* after applying ANC. The Viterbi estimates of the spin wandering are noted by the dashed coloured lines. Before ANC, the Viterbi algorithm is unable to track the spin wandering. After ANC the Viterbi algorithm is able to track the GW frequency accurately for both the high and low noise cases.

the GW spin wandering frequency signal which is submerged with respect to the voltage interference at 60 Hz. Conversely, the application of the ANC enables the interference to be removed without perturbing the gravitational wave signal. In this case the Viterbi algorithm is able to track the GW frequency wandering in both the low and high noise cases with high fidelity. Specifically, the mean squared error in the frequency estimate is 1.4×10^{-3} Hz for the low noise case and 0.22 Hz for the high noise case.

C. ROC curves versus power line parameters

With the performance of the ANC and Viterbi approach established for a single example, it is of interest to explore how the algorithm performs for different power line parameters. In this section we vary Δf and γ to test how the combined ANC filter and Viterbi algorithm perform across multiple noise realisations. To this end we calculate the detection probability compared to the false alarm probability, i.e. the receiver operating characteristic (ROC), for the Viterbi search after the data has been filtered using ANC. We consider two situations. In the first situation we hold Δf constant at $\Delta f = 0.0$ and set $\gamma = \{0.001, 0.01, 0.1\}$. In the second situation we hold γ constant at $\gamma = 0.02$ and set $\Delta f = \{0.25, 0.5, 1.0\}$. The results are shown in Figure 8.

Whilst the underlying wandering GW frequency signal can generally be tracked well for a single noise realisation (c.f. Figure 7), we can see that for these power line parameters, across multiple noise realisations there is a high false alarm rate. This is a consequence of the interference not being completely removed and evidences how even a small quantity of clutter noise is sufficient to corrupt the search for continuous waves. To quantify the performance with a single scalar value we consider the Area Under the Curve (AUC), a common metric used to evaluate ROC curves. The AUC can be in the range $0.5 - -1.0$, where AUC $= 0.5$ corresponds to the performance of a random classifier (i.e. the grey dashed diagonal line in the figure) and AUC $= 1.0$ represents a perfect classifier. For the first situation with $\Delta f = 0.0$ and $\gamma = \{0.001, 0.01, 0.1\}$, AUC $= \{0.68, 0.65, 0.55\}$ respectively. For the second situation with $\gamma = 0.02$ and $\Delta f = \{0.25, 0.5, 1.0\}$, AUC $= \{0.62, 0.58, 0.58\}$ respectively. Whilst the method performs better than a random classifier for all parameters, the AUC values are generally low, especially for cases where the interference has a large amplitude with a long period. In Section IV D we explore the use of different parameters used for the ANC filter, including the inclusion of additional reference channels

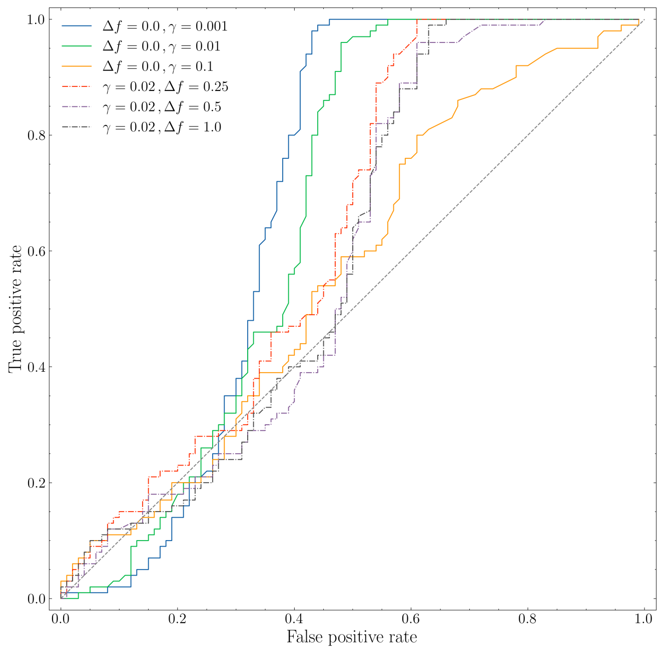

FIG. 8. ROC curve over multiple noise realisations for different power line parameters. All other parameters are as specified in Table I. There is an evident high false alarm rate for all parameters, with a low AUC value of 0.55 for $\Delta f = 0.0$, $\gamma = 0.1$, and a high AUC value of 0.68 for $\Delta f = 0.0$, $\gamma = 0.001$. The high false alarm rate is a result of the interference not being completely removed by the ANC filter.

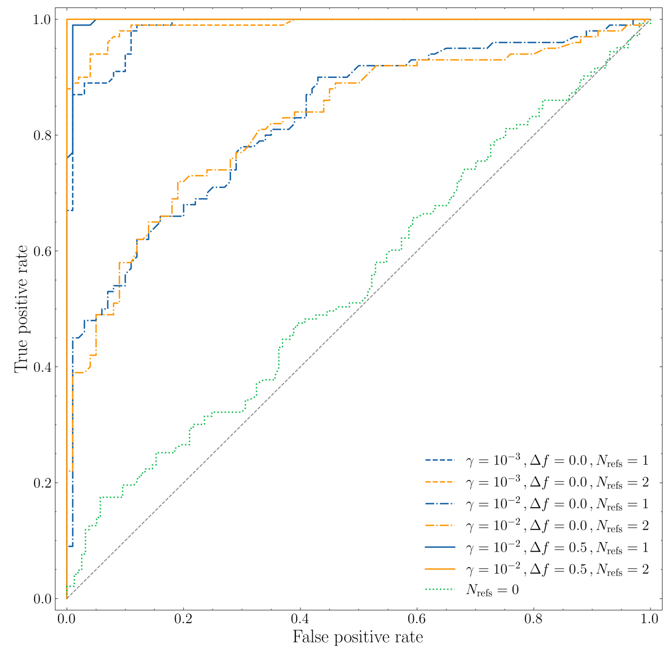

FIG. 9. ROC curve for 3 different systems: System A with $\{\gamma = 0.001, \Delta f = 0.0\}$, System B with $\{\gamma = 0.01, \Delta f = 0.0\}$ and System C with $\{\gamma = 0.01, \Delta f = 0.5\}$ (dotted, dashed, solid lines respectively). The orange lines denote the Viterbi search run with 2 reference channels, the blue lines using 1 reference channel. The green dotted line is the detection performance in the absence of ANC filtering. The grey dashed line is the performance of a random classifier. Detection using ANC filtering consistently outperforms that without ANC filtering. ANC filtering using 2 reference channels generally outperforms that using a single reference channel.

	System		
$N_{\rm refs}$	A	B	C
1	0.975	0.827	0.987
2	0.990	0.822	0.999

TABLE II. Summary of AUC values for each of the ROC curves presented in Figure 9. Adding an extra PEM reference generally improves the detection performance, with the exception of system B. The AUC value for the zero PEM reference case (i.e. no ANC filtering) is AUC = 0.55.

D. ROC curves vs. filter parameters

We have shown that the ANC filter used in conjunction with the Viterbi algorithm is effective at tracking the wandering GW spin frequency, but suffers from a high false alarm rate for the particular parameters of the ANC filter that we have been using. In this Section we investigate two important questions:

1. How does ANC benefit from multiple independent references?

2. What order ANC filter (M) is required to achieve good interference cancellation?

Regarding the first question, the preceding validation tests on synthetic data all assumed that we have a single PEM reference voltage measurement. However, as discussed in Section II C in practice there are multiple PEM channels measuring power line interference for LIGO (c.f. Figure II C). Specifically, in the open O3a data there are 9 PEM channels for LIGO-Livingston and 7 PEM channels for LIGO-Hanford. Multiple PEM channels provided additional independent measurements of the reference voltage; it seems reasonable to suspect that these additional channels may aid the performance of the ANC filter. Indeed, ANC is commonly used with multiple reference signals in other electrical engineering applications such as noise cancelling headphones [44], communication intelligibility [45, 46] and cardiac monitoring [47].

In Figure 9 we present the ROC curves for 3 different example systems:

- **System A**. $\gamma = 10^{-3}\,, \Delta f = 0.0$. Dashed lines in the figure.

- **System B**. $\gamma = 10^{-2}\,, \Delta f = 0.0$. Dash-dotted in the figure

- **System C**. $\gamma = 10^{-2}\,, \Delta f = 0.5$. Solid lines in the

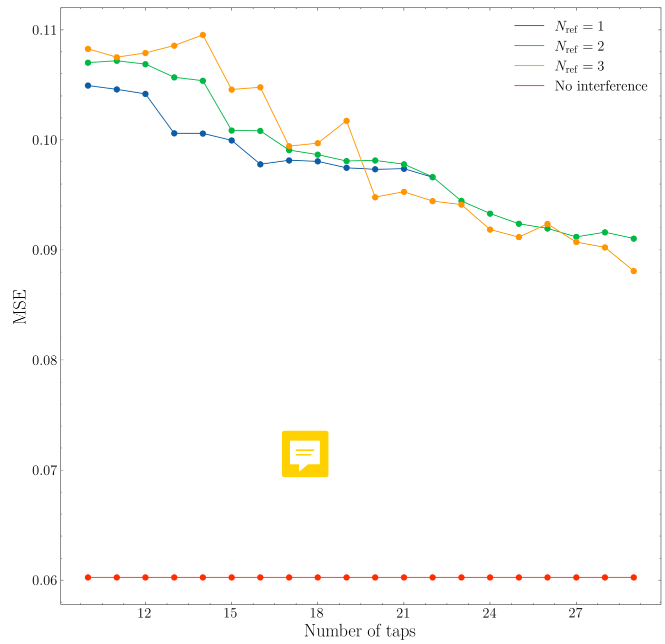

FIG. 10. Mean square error (MSE) in the Viterbi estimates of the GW wandering spin frequency for a system with $h = 0.025$ and $\Delta f = 0.0$ relative to the number of taps used in the FIR filter. Up to three PEM reference channels are used. The solution for the case with zero interference clutter is also shown.

figure

For each system we compute the ROC curve over multiple noise realisations for both one and two reference channels (blue, and orange lines respectively). This gives us six total ROC curve solutions. For comparison we also plot the case where we run the Viterbi search for one of the example systems, but with zero PEM references (dotted green line). The specific AUC values for each ROC curve are reported in Table II.

We can see that for all example systems the detection probability is high with respect to the false alarm probability. The detection probability using ANC filtering is greater than without ANC filtering for all systems (i.e. the orange and blue lines are exclusively above the green dotted line). Specifically the for the zero filtering case AUC= 0.54, whilst all cases which use ANC filtering have AUC$\geq$ 0.82 and as high as AUC= 0.99. The inclusion of an additional reference channel improves the detection probability for Systems A and C, with the AUC values rising from 0.975 to 0.990 for System A and from 0.987 to 0.999 for System B. No improvement is observed for System B, with AUC= 0.827 for $N_{\rm ref} = 1$ and AUC= 0.822 for $N_{\rm ref} = 2$. The lack of improvement for System B when using two PEM references suggests that for these parameters a single reference is sufficient to capture the dynamics of the interference clutter.

[6]

Regarding the second question, the order of the filter, i.e. the number of taps, is a parameter that can be freely chosen in ARLS. It is important to consider the filter's robustness to the choice of M. Generally an increased number of taps is expected to improve the performance of the filter due to the increased model complexity. However, this comes at an increased computational cost and also an increased latency. For real time applications tracking the wandering of the GW frequency it is important to minimize both of these variables. In Figure 10 we plot the mean squared error (MSE) in the GW frequency estimated by Viterbi compared to the true spin-wandering frequency, averaged over multiple noise realisations. We set the system to have $h = 0.025$ and $\Delta f = 0.0$ and consider up to three PEM reference channels. As a reference we also plot the error in the Viterbi estimates for the case where there is no interference clutter and so no ANC is required. We can see that generally the accuracy improves with an increased number of taps. When M is small the $N_{\rm ref} = 1$ solution generally outperforms the high $N_{\rm ref}$ solutions; in this regime the number of taps is sufficiently small to not be able to take advantage of the increased information provided by the additional reference channels. Conversely as the number of taps increases, the $N_{\rm ref} = 3$ becomes the best performing solution. We note that rather than explicitly specifying the number of taps, adaptive tap length methods that automatically update the number of taps used are also available [e.g. 48–50].

V. CONCLUSIONS

In this paper we demonstrate a new line subtraction method based on adaptive noise cancellation for use in continuous gravitational wave searches. We use an adaptive recursive least squares method in conjunction with an independent, known PEM reference signal to suppress the interference from a long-lived narrow spectral feature. We then search for the continuous wave signal using a HMM Viterbi algorithm. We test our method on synthetic data containing the 60 Hz spectral interference line due to the North American power grid. We show how the the ANC and Viterbi algorithm together are able to successfully track the spin-wandering continuous GW signal near the 60 Hz line. We test the method over multiple noise realisations and show TK: to confirm. The performance of the filter is generally improved with an increased number of reference signals and at an increased model order.

[6] TK: Why are these results so different to Figure 8? Has the strain amplitude changed? Need to check with Sofia

[1] L. Blackburn, L. Cadonati, S. Caride, S. Caudill, S. Chatterji, N. Christensen, J. Dalrymple, S. Desai, A. Di Credico, G. Ely, *et al.*, Classical and Quantum Gravity **25**, 184004 (2008), arXiv:0804.0800 [gr-qc].

[2] J. Aasi, J. Abadie, B. P. Abbott, R. Abbott, T. D. Abbott, M. Abernathy, T. Accadia, F. Acernese, C. Adams, *et al.*, Classical and Quantum Gravity **29**, 155002 (2012), arXiv:1203.5613 [gr-qc].

[3] J. Aasi, J. Abadie, B. P. Abbott, R. Abbott, T. Abbott, M. R. Abernathy, T. Accadia, F. Acernese, C. Adams, T. Adams, *et al.*, Classical and Quantum Gravity **32**, 115012 (2015), arXiv:1410.7764 [gr-qc].

[4] B. P. Abbott, R. Abbott, T. D. Abbott, M. R. Abernathy, F. Acernese, K. Ackley, M. Adamo, C. Adams, T. Adams, P. Addesso, *et al.*, Classical and Quantum Gravity **33**, 134001 (2016), arXiv:1602.03844 [gr-qc].

[5] J. Glanzer, S. Soni, J. Spoon, A. Effler, and G. González, arXiv e-prints , arXiv:2304.07477 (2023), arXiv:2304.07477 [astro-ph.IM].

[6] LIGO Scientific Collaboration, J. Aasi, B. P. Abbott, R. Abbott, T. Abbott, M. R. Abernathy, K. Ackley, C. Adams, T. Adams, P. Addesso, and et al., Classical and Quantum Gravity **32**, 074001 (2015), arXiv:1411.4547 [gr-qc].

[7] B. P. Abbott, R. Abbott, T. D. Abbott, S. Abraham, F. Acernese, K. Ackley, C. Adams, V. B. Adya, C. Affeldt, M. Agathos, *et al.*, Classical and Quantum Gravity **37**, 055002 (2020), arXiv:1908.11170 [gr-qc].

[8] .

[9] T. Akutsu, M. Ando, K. Arai, Y. Arai, S. Araki, A. Araya, N. Aritomi, H. Asada, Y. Aso, S. Bae, *et al.*, Progress of Theoretical and Experimental Physics **2021**, 05A102 (2021), arXiv:2009.09305 [gr-qc].

[10] P. Nguyen, R. M. S. Schofield, A. Effler, C. Austin, V. Adya, M. Ball, S. Banagiri, K. Banowetz, C. Billman, C. D. Blair, *et al.*, Classical and Quantum Gravity **38**, 145001 (2021), arXiv:2101.09935 [astro-ph.IM].

[11] P. B. Covas, A. Effler, E. Goetz, P. M. Meyers, A. Neunzert, M. Oliver, B. L. Pearlstone, V. J. Roma, R. M. S. Schofield, V. B. Adya, and et al., Phys. Rev. D **97**, 082002 (2018), arXiv:1801.07204 [astro-ph.IM].

[12] O. J. Piccinni, Galaxies **10**, 72 (2022), arXiv:2202.01088 [gr-qc].

[13] K. Riles, Living Reviews in Relativity **26**, 3 (2023), arXiv:2206.06447 [astro-ph.HE].

[14] K. Wette, arXiv e-prints , arXiv:2305.07106 (2023), arXiv:2305.07106 [gr-qc].

[15] R. Abbott, H. Abe, F. Acernese, K. Ackley, N. Adhikari, R. X. Adhikari, V. K. Adkins, V. B. Adya, C. Affeldt, D. Agarwal, M. Agathos, K. Agatsuma, N. Aggarwal, O. D. Aguiar, *et al.* (LIGO Scientific Collaboration, Virgo Collaboration, and KAGRA Collaboration), Phys. Rev. D **106**, 062002 (2022).

[16] R. Abbott, T. D. Abbott, S. Abraham, F. Acernese, K. Ackley, A. Adams, C. Adams, R. X. Adhikari, V. B. Adya, C. Affeldt, D. Agarwal, M. Agathos, K. Agatsuma, N. Aggarwal, O. D. Aguiar, *et al.*, Astrophys. J. **921**, 80 (2021), arXiv:2105.11641 [astro-ph.HE].

[17] R. Abbott, T. D. Abbott, F. Acernese, K. Ackley, C. Adams, N. Adhikari, R. X. Adhikari, V. B. Adya, C. Affeldt, D. Agarwal, M. Agathos, K. Agatsuma, N. Aggarwal, O. D. Aguiar, L. Aiello, *et al.*, Astrophys. J. **932**, 133 (2022), arXiv:2112.10990 [gr-qc].

[18] D. Davis, J. S. Areeda, B. K. Berger, R. Bruntz, A. Effler, R. C. Essick, R. P. Fisher, P. Godwin, E. Goetz, A. F. Helmling-Cornell, *et al.*, Classical and Quantum Gravity **38**, 135014 (2021), arXiv:2101.11673 [astro-ph.IM].

[19] D. Davis and M. Walker, Galaxies **10** (2022), 10.3390/galaxies10010012.

[20] B. P. Abbott, R. Abbott, T. D. Abbott, M. R. Abernathy, F. Acernese, K. Ackley, C. Adams, T. Adams, P. Addesso, R. X. Adhikari, *et al.*, Classical and Quantum Gravity **35**, 065010 (2018), arXiv:1710.02185 [gr-qc].

[21] J. Lee, S. H. Oh, K. Kim, G. Cho, J. J. Oh, E. J. Son, and H. M. Lee, Phys. Rev. D **103**, 123023 (2021).

[22] B. Steltner, M. A. Papa, and H.-B. Eggenstein, Phys. Rev. D **105**, 022005 (2022), arXiv:2105.09933 [gr-qc].

[23] J. C. Driggers, M. Evans, K. Pepper, and R. Adhikari, Review of Scientific Instruments **83**, 024501 (2012), arXiv:1112.2224 [gr-qc].

[24] V. Tiwari, M. Drago, V. Frolov, S. Klimenko, G. Mitselmakher, V. Necula, G. Prodi, V. Re, F. Salemi, G. Vedovato, and I. Yakushin, Classical and Quantum Gravity **32**, 165014 (2015), arXiv:1503.07476 [gr-qc].

[25] D. Davis, T. Massinger, A. Lundgren, J. C. Driggers, A. L. Urban, and L. Nuttall, Classical and Quantum Gravity **36**, 055011 (2019), arXiv:1809.05348 [astro-ph.IM].

[26] J. C. Driggers, S. Vitale, A. P. Lundgren, M. Evans, K. Kawabe, S. E. Dwyer, K. Izumi, R. M. S. Schofield, A. Effler, D. Sigg, P. Fritschel, M. Drago, *et al.*, Phys. Rev. D **99**, 042001 (2019), arXiv:1806.00532 [astro-ph.IM].

[27] P. Jung, S. H. Oh, Y.-M. Kim, E. J. Son, T. Yokozawa, T. Washimi, and J. J. Oh, Phys. Rev. D **106**, 042010 (2022).

[28] A. Marin, D. Shoemaker, and R. Weiss, "Subtracting narrow-band noise from ligo strain data in the third observing run," LIGo Technical Report LIGO-T970127-02-D.

[29] E. Cuoco, J. Powell, M. Cavaglià, K. Ackley, M. Bejger, C. Chatterjee, M. Coughlin, S. Coughlin, P. Easter, R. Essick, H. Gabbard, T. Gebhard, S. Ghosh, L. Haegel, A. Iess, D. Keitel, Z. Marka, S. Marka, F. Morawski, T. Nguyen, R. Ormiston, M. Puerrer, M. Razzano, K. Staats, G. Vajente, and D. Williams, arXiv e-prints , arXiv:2005.03745 (2020), arXiv:2005.03745 [astro-ph.HE].

[30] G. Vajente, Y. Huang, M. Isi, J. C. Driggers, J. S. Kissel, M. J. Szczepańczyk, and S. Vitale, Phys. Rev. D **101**, 042003 (2020), arXiv:1911.09083 [gr-qc].

[31] R. Ormiston, T. Nguyen, M. Coughlin, R. X. Adhikari, and E. Katsavounidis, arXiv e-prints , arXiv:2005.06534 (2020), arXiv:2005.06534 [astro-ph.IM].

[32] T. S. Yamamoto, S. Kuroyanagi, and G.-C. Liu, Phys. Rev. D **107**, 044032 (2023).

[33] S. J. Zhu, M. A. Papa, and S. Walsh, Phys. Rev. D **96**, 124007 (2017), arXiv:1707.05268 [gr-qc].

[34] D. Jones, L. Sun, J. Carlin, L. Dunn, M. Millhouse, H. Middleton, P. Meyers, P. Clearwater, D. Beniwal, L. Strang, A. Vargas, and A. Melatos, Phys. Rev. D **106**, 123011 (2022).

[35] S. Suvorova, L. Sun, A. Melatos, W. Moran, and R. J. Evans, Phys. Rev. D **93**, 123009 (2016), arXiv:1606.02412 [astro-ph.IM].

[36] The LIGO Scientific Collaboration, the Virgo Collaboration, the KAGRA Collaboration, R. Abbott, T. D. Abbott, F. Acernese, K. Ackley, C. Adams, N. Adhikari, R. X. Adhikari, *et al.*, arXiv e-prints , arXiv:2111.03606 (2021), arXiv:2111.03606 [gr-qc].

[37] "Gravitational Wave Open Science Center," `https://www.gw-openscience.org/`.

[38] P. Jaranowski, A. Królak, and B. F. Schutz, Phys. Rev. D **58**, 063001 (1998), arXiv:gr-qc/9804014 [gr-qc].

[39] D. M. Macleod, J. S. Areeda, S. B. Coughlin, T. J. Massinger, and A. L. Urban, SoftwareX **13**, 100657 (2021).

[40] B. Widrow, J. Glover, J. McCool, J. Kaunitz, C. Williams, R. Hearn, J. Zeidler, J. Eugene Dong, and R. Goodlin, Proceedings of the IEEE **63**, 1692 (1975).

[41] S. Haykin, *Adaptive Filter Theory*, 4th ed. (Prentice Hall, Upper Saddle River, New Jersey, 07458, 2002).

[42] A. Viterbi, IEEE Transactions on Information Theory **13**, 260 (1967).

[43] B. G. Quinn and E. J. Hannan, *The Estimation and Tracking of Frequency*, Cambridge Series in Statistical and Probabilistic Mathematics (Cambridge University Press, 2001).

[44] J. Cheer, V. Patel, and S. Fontana, The Journal of the Acoustical Society of America **145**, 3095 (2019), https://pubs.aip.org/asa/jasa/article-pdf/145/5/3095/15336560/3095_1_online.pdf.

[45] S. M. Kuo and J. Kunduru, Journal of the Franklin Institute **333**, 669 (1996).

[46] K. Kokkinakis, B. Azimi, Y. Hu, and D. R. Friedland, Trends in Amplification **16**, 102 (2012), pMID: 22923425, https://doi.org/10.1177/1084713812456906.

[47] S. S. Chowdhury, R. Hyder, M. S. B. Hafiz, and M. A. Haque, IEEE Journal of Biomedical and Health Informatics **22**, 450 (2018).

[48] Y. Gong and C. Cowan, in *2004 IEEE International Conference on Acoustics, Speech, and Signal Processing*, Vol. 2 (2004) pp. ii–825.

[49] A. Kar and M. Swamy, Signal Processing **131**, 422 (2017).

[50] A. Kar, A. Anand, and M. Swamy, Applied Acoustics **158**, 107043 (2020).

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