

Notes on state-space formulation of PTA problem

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April 25, 2023

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1 Preamble

These notes collect some ideas around how to formulate the PTA data analysis as a state-space problem.

We will frequently reference previous works from A.Melatos, in particular Melatos 2018 (GR assignment UniMelb) and Melatos 2022 (private communication at UniMelb).

Geometric units ($c = \hbar = G = 1$) are used throughout, with the usual convention of Roman indices (i, j) for spatial indices. Note that i is used for both labelling tensor indices and as the imaginary number - I trust the difference to be obvious from the context.

2 Derivation of pulsar frequency modulation due to GW

We want to know how the pulse frequency from a pulsar is influenced by a passing gravitational wave (GW).

We will consider the pulse frequency as a photon with covariant 4-momentum p_μ ¹

¹ See Melatos 2022 for nuances and discussion around this construction of "pulse train as photon"

We take a gravitational plane wave that perturbs a background Minkowski spacetime as ²

$$g_{\mu\nu} = \eta_{\mu\nu} + H_{\mu\nu} e^{i(\Omega(\bar{n} \cdot \bar{x} - t) + \Phi_0)} \quad (1)$$

² c.f. Melatos 2018, Equation 1

where the GW has angular frequency Ω , propagates in the \bar{n} -direction and has a phase offset of Φ_0 . Note that we are free to choose our coordinate system such that Φ_0 is the GW phase at $t = 0$ *at the Earth*. The amplitude tensor $H_{\mu\nu}$ has zero temporal components ($H_{0\mu} = H_{\mu 0} = 0$) whilst the spatial part is

$$H_{ij} = h_+ e_{ij}^+(\bar{n}) + h_\times e_{ij}^\times(\bar{n}) \quad (2)$$

Note that we use a convention where the amplitudes $h_{+,\times}$ label the *constant* amplitudes of the gravitational plane wave. All time variation in the strain is a result of the exponential term. This is in contrast to the notation used by some authors ³ where the amplitudes include the time variability $h_{+,\times} = h_{+,\times}(t)$. The polarisation tensors $e_{ij}^{+,\times}$ are uniquely defined by the principal axes of the wave.

³ see e.g. arXiv:1003.0677

2.1 Setting up the problem

The frequency of a photon with 4-momentum p_μ recorded by an observer with 4-velocity u^μ is

$$\nu = p_\alpha u^\alpha \quad (3)$$

We consider both our emitter and receiver to be stationary, such that

$$u^\alpha|_{\text{emitter}} = u^\alpha|_{\text{receiver}} = (1, 0, 0, 0) \quad (4)$$

Consequently the frequency can be identified with the temporal component of the covariant 4-momentum,

$$\nu = p_t \quad (5)$$

The expression for the evolution of the pulse frequency as measured by the observer on Earth is then,

$$p_t(\tau)|_{\text{Earth}} = p_t(t_0)|_{\text{source}} + \int_{t=t_0}^{t=\tau} \dot{p}_t dt \quad (6)$$

where the overdot denotes a derivative w.r.t. t . Since the influence of the GW perturbation on \dot{p}_t is small, we can relate the source emission and receiver times as $\tau = t_0 + d$ and consider the photon trajectory to be an unperturbed path. ⁴

⁴ See also e.g. Maggiore who takes the same approach...

To complete our expression, we now just need to determine \dot{p}_t and integrate it.

2.2 Hamiltonian Mechanics

The Hamiltonian in covariant notation can be written as

$$H(x^\mu, p_\mu) = \frac{1}{2} g_{\mu\nu} p^\mu p^\nu, \quad (7)$$

which if we substitute in our expression for the perturbed metric is

$$H = \frac{1}{2}\eta_{\mu\nu}p^\mu p^\nu + \frac{1}{2}H_{ij}p^i p^j e^{i(\Omega(\bar{n}\cdot\bar{x}-t)+\Phi_0)} \quad (8)$$

Hamilton's equations are

$$\frac{dx^\mu}{d\lambda} = \frac{\partial H}{\partial p_\mu}, \quad \frac{dp_\mu}{d\lambda} = -\frac{\partial H}{\partial x^\mu} \quad (9)$$

for affine parameter λ . The derivative of the temporal component of the covariant momenta is then,

$$\frac{dp_t}{d\lambda} = -\frac{i\Omega}{2}H_{ij}p^i p^j e^{i(\Omega(\bar{n}\cdot\bar{x}-t)+\Phi_0)} \quad (10)$$

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Therefore the derivative w.r.t coordinate time t is,

$$\dot{p}_t = \frac{dp_t}{d\lambda} \left(\frac{dt}{d\lambda} \right)^{-1} = \frac{dp_t}{d\lambda} \left(\frac{1}{p^t} \right) \quad (11)$$

It will be useful to recognise that we can express the photon momentum in terms of the unit vector between the Earth and pulsar \bar{q}

$$p^\mu = \omega(1, -q^x, -q^y, -q^z) \quad (12)$$

where ω is the *constant* photon angular frequency. Given the small effect of the GW perturbation, at first order we can identify ω as either the frequency at source or observer (see Melatos 2022, PT16). Similarly, we can parameterize the spatial coordinates \bar{x} as,

$$\bar{x}(t) = -\bar{q}(t - \tau) \quad (13)$$

Note that when our integration variable $t = \tau$ (i.e. the photon is at Earth) then $\bar{x}(t = \tau) = 0$ which is what we expect given our chosen coordinate system. Also note that \dot{p}_t is entirely a function of the GW perturbation. In the Minkowski case the spacetime is stationary and so p_t should be conserved along the geodesic.

Bringing this all together

$$\dot{p}_t = \frac{dp_t}{d\lambda} \left(\frac{1}{p^t} \right) \quad (14)$$

$$= -\frac{i\Omega}{2}H_{ij}p^i p^j e^{i(\Omega(\bar{n}\cdot\bar{x}-t)+\Phi_0)} \left(\frac{1}{p^t} \right) \quad (15)$$

$$= \frac{-i\Omega\omega^2}{2}H_{ij}q^i q^j e^{i(\Omega(\bar{n}\cdot\bar{x}-t)+\Phi_0)} \left(\frac{1}{\omega} \right) \quad (16)$$

$$= \frac{-i\Omega\omega}{2}H_{ij}q^i q^j e^{i(\Omega(\bar{n}\cdot[-\bar{q}(t-\tau)]-t)+\Phi_0)} \quad (17)$$

$$= \frac{-i\Omega\omega}{2}H_{ij}q^i q^j e^{i(-\Omega t(1+\bar{n}\cdot\bar{q})+\Omega\tau\bar{n}\cdot\bar{q}+\Phi_0)} \quad (18)$$

$$= Ae^{i(\gamma t+B)} \quad (19)$$

⁵ This is equivalent to Melatos 2018, Eq 5 for the specific case of a GW propagating in the z-direction, with zero phase offset.

with

$$A = -\frac{i\Omega\omega}{2} H_{ij} q^i q^j \quad (20)$$

$$\gamma = -\Omega(1 + \bar{n} \cdot \bar{q}) \quad (21)$$

⁶ and

$$B = \Omega\tau\bar{n} \cdot \bar{q} + \Phi_0 \quad (22)$$

⁶ compare with Melatos 2022, PT11

2.3 Performing the integral

The frequency shift experienced by the observer relative to the source due to a GW is then

$$p_t(\tau)|_{\text{Earth}} - p_t(\tau - d)|_{\text{source}} = A \int_{t=\tau-d}^{t=\tau} e^{i(\gamma t + B)} dt \quad (23)$$

$$= \frac{-iA}{\gamma} e^{iB} \left(e^{i\gamma\tau} - e^{i\gamma(\tau-d)} \right) \quad (24)$$

$$= \frac{-iA}{\gamma} e^{i\Omega\tau\bar{n} \cdot \bar{q}} e^{i\Phi_0} e^{i\gamma\tau} \left(1 - e^{-i\gamma d} \right) \quad (25)$$

$$= \frac{-iA}{\gamma} e^{i\Omega\tau\bar{n} \cdot \bar{q}} e^{i\Phi_0} e^{i(-\Omega(1+\bar{n} \cdot \bar{q}))\tau} \left(1 - e^{-i\gamma d} \right) \quad (26)$$

$$= \frac{-iA}{\gamma} e^{i\Phi_0} e^{-i\Omega\tau} \left(1 - e^{-i\gamma d} \right) \quad (27)$$

$$= \frac{-iA}{\gamma} e^{i(-\Omega\tau + \Phi_0)} \left(1 - e^{-i\gamma d} \right) \quad (28)$$

2.4 Explicit expression and comparison with Melatos 22

We are now very close with making contact with the expression from Melatos 22. Let's be explicit and sub in our definitions for A and γ into 28

$$p_t(\tau)|_{\text{Earth}} - p_t(\tau - d)|_{\text{source}} = -i \frac{-\frac{i\Omega\omega}{2} H_{ij} q^i q^j}{-\Omega(1 + \bar{n} \cdot \bar{q})} e^{i(-\Omega\tau + \Phi_0)} \left(1 - e^{-i(-\Omega(1+\bar{n} \cdot \bar{q}))d} \right) \quad (29)$$

$$= \frac{\omega}{2} \frac{H_{ij} q^i q^j}{(1 + \bar{n} \cdot \bar{q})} e^{i(-\Omega\tau + \Phi_0)} \left(1 - e^{i\Omega(1+\bar{n} \cdot \bar{q})d} \right) \quad (30)$$

$$= \frac{\omega}{2} \frac{h_{ij}(\tau; \bar{x} = 0) q^i q^j}{(1 + \bar{n} \cdot \bar{q})} \left(1 - e^{i\Omega(1+\bar{n} \cdot \bar{q})d} \right) \quad (31)$$

which is exactly the expression derived in Melatos 2022.

2.5 The Φ_0 issue

There was previously some confusion over the presence of a single Φ_0 parameter that was shared between multiple pulsars. Does this cause $\tau = 0$ to be a special time when all pulsars are in phase and then go on to evolve to be out of phase? This is just due to some sloppy transcription of equations in another document. Ignoring

measurement noise, the measured photon frequency is related to the emitted photon frequency as

$$f_{\text{measured}} = f_{\text{emitted}} \left(1 - \frac{1}{2} \frac{h_{ij}(\tau; \bar{x} = 0) q^i q^j}{(1 + \bar{n} \cdot \bar{q})} \left(1 - e^{i\Omega(1 + \bar{n} \cdot \bar{q})d} \right) \right) \quad (32)$$

Now the pulsar frequencies remain in phase and $\tau = 0$ is no longer a special time. ⁷

⁷ This expression is slightly different from the (erroneous) one typed up in the manuscript! We have dropped a $1 + \bar{n} \cdot \bar{q}$ factor that was wrongly typed up in the other document

2.6 Trigonometric form

Lets express Eq. 33 as

$$f_{\text{measured}} = f_{\text{emitted}} g(\tau; \bar{\theta}) \quad (33)$$

where $\bar{\theta}$ is a set of parameters of the system. Lets write this in terms of the real part of the trig expansion. Let $a = -\Omega\tau + \Phi_0$ and $b = \Omega(1 + \bar{n} \cdot \bar{q})d$

$$g(\tau; \bar{\theta}) = 1 - \frac{1}{2} \frac{H_{ij} q^i q^j}{(1 + \bar{n} \cdot \bar{q})} e^{ia} (1 - e^{ib}) \quad (34)$$

$$= 1 - \frac{1}{2} \frac{H_{ij} q^i q^j}{(1 + \bar{n} \cdot \bar{q})} [\cos(a) + i \sin(a)] (1 - [\cos(b) + i \sin(b)]) \quad (35)$$

$$= 1 - \frac{1}{2} \frac{H_{ij} q^i q^j}{(1 + \bar{n} \cdot \bar{q})} (\cos(a) + i \sin(a) - \cos(a) \cos(b) - i \cos(a) \sin(b) - i \sin(a) \cos(b) + \sin(a) \sin(b)) \quad (36)$$

$$\stackrel{\text{Re}}{=} 1 - \frac{1}{2} \frac{H_{ij} q^i q^j}{(1 + \bar{n} \cdot \bar{q})} (\cos(a) - \cos(a) \cos(b) + \sin(a) \sin(b)) \quad (37)$$

$$= 1 - \frac{1}{2} \frac{H_{ij} q^i q^j}{(1 + \bar{n} \cdot \bar{q})} (\cos(a) - \cos(a + b)) \quad (38)$$

$$= 1 - \frac{1}{2} \frac{H_{ij} q^i q^j}{(1 + \bar{n} \cdot \bar{q})} [\cos(-\Omega\tau + \Phi_0) - \cos(-\Omega\tau + \Phi_0 + \Omega(1 + \bar{n} \cdot \bar{q})d)] \quad (39)$$