

Pulsar Lensing Geometry

Siqi Liu^{1,3*}, Ue-Li Pen^{1,2,†}, J-P Macquart^{4,‡}, Walter Brisken^{5,§}, Adam Deller^{6,¶}

¹ *Canadian Institute for Theoretical Astrophysics, University of Toronto, M5S 3H8 Ontario, Canada*

² *Canadian Institute for Advanced Research, Program in Cosmology and Gravitation*

³ *Department of Astronomy and Astrophysics, University of Toronto, M5S 3H4, Ontario, Canada*

⁴ *ICRAR-Curtin University of Technology, Department of Imaging and Applied Physics, GPO Box U1978, Perth, Western Australia 6102, USA*

⁵ *National Radio Astronomy Observatory, P.O. Box 0, Socorro, NM 87801, USA*

⁶ *ASTRON, the Netherlands Institute for Radio Astronomy, Postbus 2, 7990 AA, Dwingeloo, The Netherlands*

18 May 2015

ABSTRACT

We analyze archival VLBI data of PSR B0834+06, concluding that for this example the plasma lenses can be precisely modelled using the (Pen & Levin 2014) inclined sheet model, resulting in two distinct lens planes. This data strongly favours the grazing sheet model over turbulence as the primary source of pulsar scattering. The simple 1-D structure of the lenses opens up the possibility of using interstellar lenses as precision probes for pulsar lens mapping, and new opportunities for removing scattering to improve pulsar timing. We describe the parameters and observable of this double screen system. While relative screen distances can in principle be accurately determined, a global conformal distance degeneracy exists which allows a rescaling of the absolute distance scale. This degeneracy is broken if the pulsar resides in a binary system, which is the case for most precision timing targets.

Key words: Pulsar

1 INTRODUCTION

Pulsars have long provided a rich source of astrophysical information due to their compact emission and predictable timing. One of the weakest measurements for most pulsars is their direct geometric distance. For some pulsars, timing parallax or VLBI parallax has resulted in direct distance determinations. For most pulsars, the distance is a major uncertainty for precision timing interpretations, including mass, moment of inertia, and gravitational wave direction (Boyle & Pen 2012).

Direct VLBI observation of PSR B0834+06 shows multiple images lensed by the interstellar plasma. Combining the angular positions and scintillation delays, the authors published the derived effective distance (Brisken et al. 2010) of approximately 1168 ± 23 pc for apexes whose time delays range from 0.1 ms to 0.4 ms, and 1121 ± 59 pc for 1 ms apexes. This represents a precise measurement compared to all other attempts to derive distances to this pulsar. This effective distance is a combination of pulsar-screen and earth-screen distances, and does not allow a separate deter-

mination of the individual distances. A binary pulsar system would in principle allow a breaking of this degeneracy (Pen & Levin 2014). One potential limitation is the precision to which the lensing model can be understood. In this paper, we demonstrate that the lensing screen consists of nearly parallel linear refractive structures, in two screens. The precise model confirms the one dimensional nature, and thus the small number of parameters that need to be measured to quantify the lensing screen.

2 LENSING

2.1 Archival data of B0834+06

Our analysis is based on the apex data selected from the secondary spectrum of pulsar B0834+06 in (Brisken et al. 2010), which was observed as part of a global VLBI project on 2005 November 12, with GBT, Arecibo, Lovell and Westerbork telescopes. The GB-AR and AR-WB baselines are close to orthogonal and of comparable lengths, resulting in relatively isotropic astrometric positions. Information from each identified apex includes delay τ , delay rate (differential frequency f_D), relative Right Ascension $\Delta\alpha$, relative declination $\Delta\delta$, error of $\Delta\alpha$ σ_α and error of $\Delta\delta$ σ_δ . Data of each apex are collected from four dual circular polarization 8 MHz wide sub-bands spanning the frequency range 310.5–

* E-mail: sqliu@cita.utoronto.ca

† E-mail: pen@cita.utoronto.ca

‡ E-mail: J.Macquart@curtin.edu.au

§ Email: wbrisken@aoc.nrao.edu

¶ E-mail: deller@astron.nl

342.5 MHz. As described in Briskeen et al. (2010), the inverse parabolic arclets were fitted to positions of their apexes, resulting in a catalog of apexes in each sub-band, each with delay and differential frequency. As previously described, the positions of the apexes appears constant across sub-bands. In this work, we first combine the apexes across sub-bands, resulting in a set of lenses. We focus on the southern group, which have negative differential frequency: this grouping appears as a likely candidate for a double lensing screen since two groups appear both in the VLBI angular positions, and the secondary spectra. We divide the apex data with negative differential frequency into two groups: in one group time delay ranges from 0.1 ms to 0.4 ms, which we call 0.4 ms group, and in the other group time delay at about 1 ms, which we call 1 ms group. The statistics of the positions of the points are: for 0.4 ms group, there are 10 apexes in the first two sub-bands, and 14 apexes in the last two sub-bands; for 1 ms group, there are 5, 6, 5 and 4 apexes in the four sub-bands subsequently, with center frequency $f_{\text{band}} = 314.5, 322.5, 330.5$ and 338.5 MHz.

Next we want to select the same apexes from four sub-bands. To match the same apexes in different sub-bands, we convert the differential frequency in different sub-bands to the one in 322.5 MHz, by $f_D/f_{\text{band}} \cdot 322.5$ MHz. We map a total of 9 apexes from the 0.4 ms group, and 5 apexes from the 1 ms group. This results in an estimation for the average value in $f = 322.5$ MHz and standard deviation among four sub-bands. They are listed in Table 1. The f_D and τ are the arithmetic mean value of the four sub-bands. The mean value of $\Delta\alpha$ and $\Delta\delta$ are the weighted mean. The angular offset of the scintillation image away from the central emission is calculated as $\theta^2 = (\Delta\alpha \cos(\delta))^2 + (\Delta\delta)^2 - \sigma_{\Delta\alpha}^2 - \sigma_{\Delta\delta}^2$.

The method we use to calculate the error of time delay τ , differential frequency f_D , $\Delta\alpha$ and $\Delta\delta$ are listed in Table 1, is by following equation:

$$\sigma_{\tau, f_D, \Delta\alpha, \Delta\delta}^2 = \frac{\sum_{i=1}^4 (x_i - \bar{x})^2}{n(n-1)}, \quad (1)$$

and $n = 4$ for four sub-bands.

2.2 One lens model

2.2.1 Distance of the lens

In the absence of a lens model, the fringe rate, delay and angular position cannot be uniquely related. To interpret the data, we adopt the lensing model of (Pen & Levin 2014). In this model, the lensing is due to projected fold caustics of a thin sheet closely aligned to the line of sight.

The relation of the distance of the pulsar D_p , the time delay τ with the angular offset θ , and the relation of velocity and the differential frequency f_D are described by the following equations:

$$\begin{aligned} \tau &= \frac{D_e \theta^2}{2c}, \\ f_D &= f \frac{d\tau}{dt}, \end{aligned} \quad (2)$$

where D_e is the effective distance. If we denote the distance of the pulsar D_p , the distance of the lens D_s , then the effective distance is equivalent to the distance of the lens placed at the middle point of the pulsar: $D_e = D_p D_s / (D_p - D_s)$.

We plot the values of θ vs square root of τ in Figure 1. A least square analysis of the effective distance follows:

$$\begin{aligned} k \sum_{i=1}^n \frac{1}{\sigma_{ki}^2} &= \sum_{i=1}^n \frac{\theta/\sqrt{\tau}}{\sigma_{ki}^2}, \\ \frac{1}{\sigma_k^2} &= \sum_{i=1}^n \frac{1}{\sigma_{ki}^2}, \end{aligned} \quad (3)$$

where k is denoted as the slope in Figure 1, and σ_{ki} as the error of the slope. $\sigma_{ki} = \sigma_{\theta i} / \sqrt{\tau}$, $\sigma_{\theta} = \sqrt{(\sigma_{\Delta\alpha} \cos(\delta))^2 + \sigma_{\Delta\delta}^2}$. $n = 9$ for 1 ms group, and $n = 4$ for 0.4 ms group.

From this $\theta-\sqrt{\tau}$ relation, the effective distance can be calculated as $D_e = 2c/k^2$. Thus, $D_{1e} = 1023 \pm 27$ for the 0.4 ms group, distance of lens 1, and $D_{2e} = 1281 \pm 82$ for the 1 ms group, distance of lens 2. The errors, and uncertainties on the error, precludes a definitive interpretation of the apparent difference in distance. At face value, this indicates that the lens 1 is closer to the pulsar, and we will use this as a basis for the model in this paper. We discuss consequences of alternate interpretations in section 2.4. We take $D_{1e} = 1023$ pc, combined with the VLBI measured distance of the pulsar 640 pc, the distance of lens 1 D_1 , where 0.4 ms scintillation points are refracted, is equal to 393.7 pc. For 1 ms apexes, the distance of lens 2 is equal to 426.7 pc. Lens 2 is closer to the pulsar, thus, the degeneracy of the distance of the screen is broken.

2.2.2 Angular positions of 0.4 ms group

We plot the observed relative angular positions in Figure 2. We fit a line to the angular positions of the 0.4 ms group, which has an positive angle of $\gamma = -25.2^\circ$ (east of the declination axis). We use this axis to define \parallel and define \perp by a 90° clockwise rotation from it.

We calculate θ from the $\theta-\sqrt{\tau}$ relation and observed τ . Because all of the θ here lie on the axis defined by γ on lens 1, so they are also denoted $\theta_{1\parallel}$ listed in the first column in Table 1. Then the calculated angular positions of the 0.4 ms group are calculated:

$$\begin{aligned} \Delta\alpha_C &= -\theta \sin\gamma / \cos\delta, \\ \Delta\delta_C &= -\theta \cos\gamma, \end{aligned} \quad (4)$$

which are marked out with the scatter points on the left side in Figure 2.

2.2.3 Discussion of one lens model

The 0.4 ms group lens solution appears consistent with the premise of the inclined sheet lensing model (Pen & Levin 2014).

The time in last column of Table 1, which we denote as t_0 , is calculated with $-2\tau f/f_D$, corresponds to the time required for the pulsar to intersect a given lens caustic. Knowing time delay τ , we can calculate the distance of the screen; knowing the angular position of point 5 and its differential frequency f_D , we can get the velocity of the pulsar; knowing the velocity and observation differential frequency, we can get the position of points 1–4.

| $\theta_{1\parallel}(\text{mas})$ | $f_D(\text{mHz})$ | $\tau(\text{ms})$ | $\Delta\alpha(\text{mas})$ | $\Delta\delta(\text{mas})$ | $t_0(\text{days})$ |
|-----------------------------------|-------------------|-------------------|----------------------------|----------------------------|--------------------|
| -8.26 | -12.9(2) | 0.0845(5) | 2.9(3) | -8.2(4) | -48.7 |
| -10.68 | -16.8(3) | 0.1412(9) | 3.9(6) | -10.6(4) | -62.8 |
| -12.32 | -18.9(2) | 0.188(2) | 5.1(6) | -10.6(6) | -74.2 |
| -13.41 | -20.4(5) | 0.222(3) | 5.6(3) | -11.73(8) | -81.4 |
| -13.82 | -21.2(6) | 0.236(2) | 5.1(4) | -12.6(5) | -83.3 |
| -14.59 | -22.3(5) | 0.2633(3) | 6.2(3) | -14.2(7) | -88.1 |
| -16.24 | -24.6(4) | 0.327(3) | 6.5(6) | -14.1(4) | -99.0 |
| -16.52 | -24.9(4) | 0.3378(3) | 8.3(4) | -14.4(8) | -101 |
| -17.39 | -26.1(4) | 0.3743(6) | 8.5(3) | -15.7(3) | -107 |
| ... | -35.1(5) | 0.950(2) | -15(1) | -21(1) | -202 |
| ... | -38.3(6) | 0.9763(9) | -15(1) | -20.7(3) | -190 |
| ... | -40.2(6) | 1.005(8) | -14(1) | -22.3(4) | -187 |
| ... | -41.3(5) | 1.037(3) | -11(1) | -19(3) | -188 |
| ... | -43.1(4) | 1.066(5) | -8(3) | -24(2) | -185 |

Table 1. 0.4 ms and 1 ms observation data and calculated data. The upper part of the table list the 0.4 ms group data, while the 1 ms group lie in the lower part of the table. Observation data include the differential frequency f_D , time delay τ from scintillation measurement (τ_1 for 0.4 ms group and τ_2 for 1 ms group); $\Delta\alpha$ and $\Delta\delta$ are from the VLBI measurement. The methods of how to calculate the error of time delay, differential frequency and the last column time are mentioned in Section 2.1.

2.3 Double lens model

2.3.1 Solving the double lens model

We denote the position of the pulsar point A, position of the lensed image on lens 2 point H, position of the lensed image on lens 1 point B, position of the observer point O, pedal from the pulsar to line HJ point J, the pedal from point H to line BD point F, and the pedal from point B to line HJ point G, for easier discussion. Because points 1–4 share the approximately same time delay with point 5, the lens where the image formed should be at the same distance away from us. The only reasonable position of screen (line HJ) that fits all these five points, marked with a solid line in Figure 2. That is unrealistic for the structure of the interstellar medium.

Therefore, we consider another model candidate: the double lens model. Respective calculation shows that the light is first refracted by lens 2 and then refracted by lens 1.

The first step is to calculate the position of J. We make an estimate of the distance of J by the 1 ms θ - $\sqrt{\tau}$ relation, and then we calculated the position of J by matching the time delay of point 2 and point 5. The result shows that lens 2 is 426.7 pc away from us. And its position is marked in Figure 2. Because J is the pedal to lens 2, we made a line that is perpendicular to AJ, the solid line in Figure 2 to denote lens 2.

The second step is to find the matched pairs of those two lenses. By trial and error, we found that the 5 points in 0.4 ms group that have the largest θ should be the candidates where lens 1 lie. These five matched lines are marked with dot dash lines in Figure 2 and their values are listed in the first two columns in Table 3. They are the located at a distance 393.7 pc away from us. Here we define three distances:

$$\begin{aligned}
 D_{p2} &= 640 \text{ pc} - 426.7 \text{ pc} = 213.3 \text{ pc}, \\
 D_{21} &= 426.7 \text{ pc} - 393.7 \text{ pc} = 33.0 \text{ pc}, \\
 D_1 &= 393.7 \text{ pc} - 0 \text{ pc} = 393.7 \text{ pc},
 \end{aligned} \tag{5}$$

where D_{p2} is the distance from the pulsar to lens 2, D_{21} is

the distance from lens 2 to lens 1, and D_1 is the distance from lens 1 to the observer.

Figure 3 and Figure 4 are examples of how light are being refracted on the first lens plane and the second lens plane. We specifically chose the point with $\theta_{1\parallel}$ equal to -17.39 mas, which is point 5 on lens 2 and point 6 on lens 1 as an example. We solve the solutions in double lens model by following equations:

$$\begin{aligned}
 \frac{JH}{D_{p2}} &= \frac{HG}{D_{21}}, \\
 \frac{FB}{D_{21}} &= \frac{BD}{D_1}.
 \end{aligned} \tag{6}$$

The solved positions are plotted in Figure 2, and respective time delays and differential frequencies are listed in Table 3. For the error of time delay τ in double lens model, we use the following equation:

$$\left(\frac{\sigma_{\tau_i}}{\tau_{2i}}\right)^2 = \left(\frac{\sigma_{\tau_{1i}}}{\tau_{1i}}\right)^2 + \left(\frac{\sigma_{\tau_{2i}}}{\tau_{2i}}\right)^2 + \left(\frac{\sigma_{\tau_{2j}}}{\tau_{2j}}\right)^2, \tag{7}$$

where τ_1 and σ_{τ_1} represent the time delay and its error from the 0.4 ms group on lens 1, τ_2 and σ_{τ_2} represent the time delay and its error from respective 1 ms group on lens 2. And τ_{2j} is the τ_2 for the nearest point in reference in Table 3 and $\sigma_{\tau_{2j}}$ is its error: for point $i = 1, 3$, $j = 2$; for point $i = 4$, $j = 5$.

For the error of differential frequency f_D , we use the following equation:

$$\left(\frac{\sigma_{f_i}}{f_{Di}}\right)^2 = \left(\frac{\sigma_{f_{D1}}}{f_{D1}}\right)^2 + \left(\frac{\sigma_{f_{D2}}}{f_{D2}}\right)^2 \tag{8}$$

where f_{D2} and $\sigma_{f_{D2}}$ are the differential frequency and its error of the point in the second row in Table 3. $i = 1, 3, 4, 5$ for the subscription.

2.3.2 Comparing calculated result in double lens model and observation

Comparing τ , we calculated time delay τ_M for these five points, and list the results in Table 3. For point 2 and 5, they fit perfectly because these are the two points that we

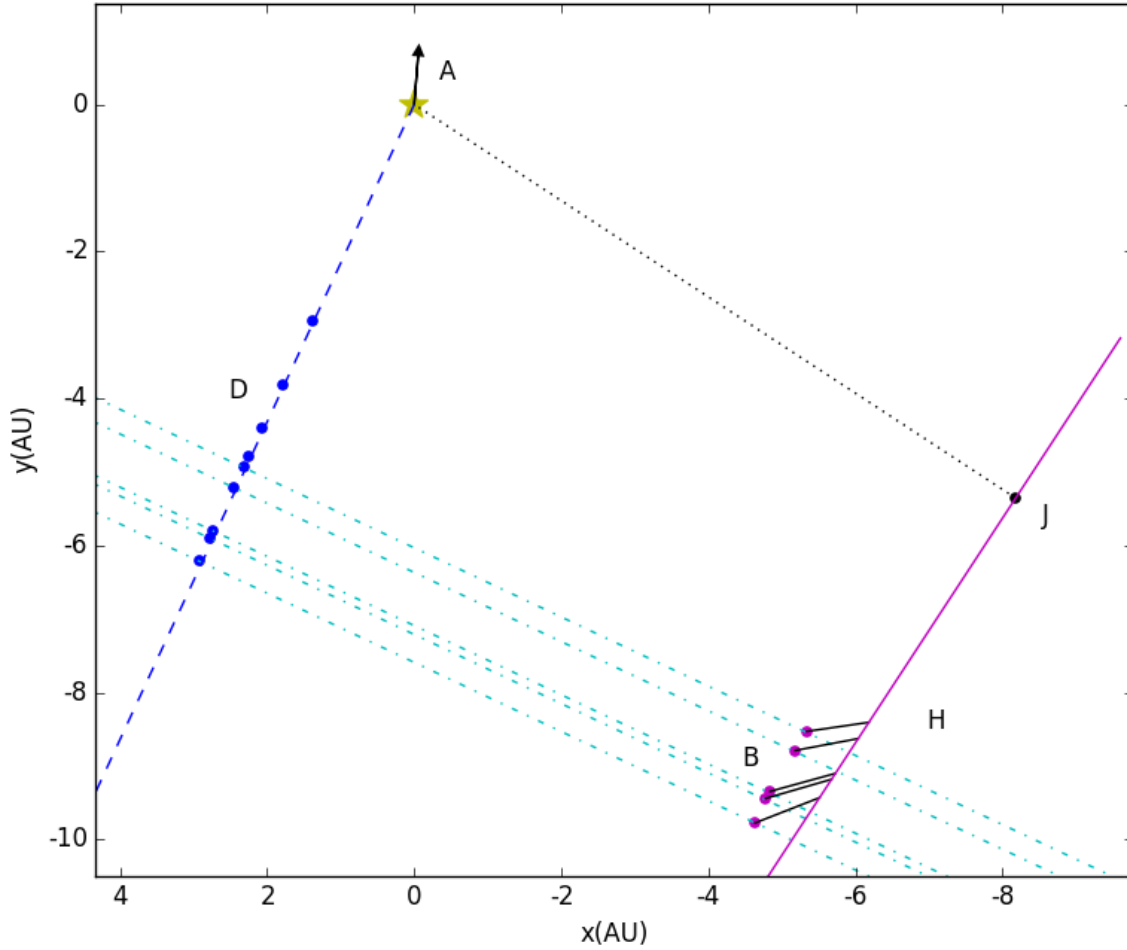


Figure 2. Observed and calculated angular positions of 0.4 ms and 1 ms data in double lens model. x axis and y axis are the relative distance to central emission of pulsar in Right Ascension direction and declination direction, on a 2D plane that is transverse to the line of sight. The position of the screen locate at 393.7 pc and 426.7 pc respectively. Scatter points on the left side, marked with letter D, are the calculated positions from the 0.4 ms apexes observation. Dash line is the fitted line of 0.4 ms apexes positions, with a angle $\gamma = -25.2^\circ$ east of north. The scatter points on the right side, marked with letter B, are the calculated lensed image on lens 1 from 1 ms group. The short solid line connects the lensed image one lens 1(observable) and lensed image 2(unobservable), marked with letter H. Long solid line is the fitted line of these positions. The dotted line on the top right side is vertical to the solid line, and the pedal is called J. Short light solid lines connect the observation points and the calculated positions in 1 ms group. Middle dot dash lines connect the 0.4 ms and 1 ms calculated positions with the same $\theta_{1||}$, which are denoted as lens 1. The proper motion of the pulsar is 192.4 km/s, with an angle $\epsilon = -4.34^\circ$ east of north, is marked with an arrow from the star, point A the position of the pulsar, at the top of the figure.

bring most pulsar timing array targets into the coherent timing regime, enabling arc minute localization of gravitational wave sources, lifting any potential source confusion.

Ultimately, the precision of the lensing results would be limited by the fidelity of the lensing model. In the inclined sheet model, the images move along fold caustics. The straightness of these caustics depends on the inclination angle, which in turn depends on the amplitude of the surface waves.

4 CONCLUSIONS

We have applied the (Pen & Levin 2014) inclined sheet model to archival apex data of PSR B0834+06. The data is well fit by two linear lensing screens, with nearly plane-parallel geometry. This appears a natural consequence of very smooth reconnection sheets, and are an unlikely outcome of ISM turbulence. These results, if extrapolated to multi-epoch observations of binary systems, might result in accurate distance determinations and opportunities for removing scattering induced timing errors.

5 ACKNOWLEDGEMENTS

We thank NSERC for support.

REFERENCES

- Boyle L., Pen U.-L., 2012, Phys. Rev. D, 86, 124028
- Briskin W. F., Macquart J.-P., Gao J. J., Rickett B. J., Coles W. A., Deller A. T., Tingay S. J., West C. J., 2010, ApJ, 708, 232
- Pen U.-L., Levin Y., 2014, MNRAS, 442, 3338
- Pen U.-L., Macquart J.-P., Deller A. T., Briskin W., 2014, MNRAS, 440, L36
- Walker M. A., Koopmans L. V. E., Stinebring D. R., van Straten W., 2008, MNRAS, 388, 1214