

Pulsar Lensing Geometry

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ABSTRACT

By analysing archival VLBI data of PSR B0834+06, we conclude that for this example the plasma lenses can be precisely modelled using the inclined sheet model (Pen & Levin 2014), resulting in two distinct lens planes. This data strongly favours the grazing sheet model over turbulence as the primary source of pulsar scattering. 7 observed apex parameters fit the model to percent accuracy. Comparison with VLBI proper motion results in a direct measure of the ionized ISM screen. The results are consistent with motions expected in the local galaxy. The simple 1-D structure of the lenses opens up the possibility of using interstellar lenses as precision probes for pulsar lens mapping, and new opportunities for removing scattering to improve pulsar timing. We describe the parameters and observables of this double screen system. While relative screen distances can in principle be accurately determined, a global conformal distance degeneracy exists which allows a rescaling of the absolute distance scale. This degeneracy is broken if the pulsar resides in a binary system, which is the case for most precision timing targets.

Key words: Pulsar

1 INTRODUCTION

Pulsars have long provided a rich source of astrophysical information due to their compact emission and predictable timing. One of the weakest measurements for most pulsars is their direct geometric distance. For some pulsars, timing parallax or VLBI parallax has resulted in direct distance determination. For most pulsars, the distance is a major uncertainty for precision timing interpretations, including mass, moment of inertia (Kramer et al. 2006; Lorimer & Kramer 2012), and gravitational wave direction (Boyle & Pen 2012).

Direct VLBI observation of PSR B0834+06 shows multiple images lensed by the interstellar plasma. Combining the angular positions and scintillation delays, the authors published the derived effective distance (Brisken et al. 2010) (hereafter B10) of approximately 1168 ± 23 pc for apexes on the main scattering axis. This represents a precise measure-

ment compared to all other attempts to derive distances to this pulsar. This effective distance is a combination of pulsar-screen and earth-screen distances, and does not allow a separate determination of the individual distances. A binary pulsar system would in principle allow a breaking of this degeneracy (Pen & Levin 2014). One potential limitation is the precision to which the lensing model can be understood.

In this paper, we examine the geometric nature of the lensing screens. In B10, VLBI astrometric mapping directly demonstrated the highly collinear nature of a single dominant lensing structure. First hints of single plane collinear dominated structure had been realized in Stinebring et al. (2001). While the nature of these structures are already mysterious, this puzzle is compounded by an offset group of lensed imaged called the 1ms group. The mysterious nature of lensing questions any conclusions drawn from scintillometry as a quantitative tool (Pen et al. 2014).

Using archival data we demonstrate in this paper that the lensing screen consists of nearly parallel linear refractive structures, in two screens. The precise model confirms the one dimensional nature, and thus the small number of parameters that quantify the lensing screen.

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2 LENSING

In this section we map the data onto the grazing incidence sheet model. The folded sheet model is qualitatively analogous to the reflection of a street lamp across a lake as seen from the opposite shore, as what is shown in Figure 1.

In the absence of waves, exactly one image forms at the point where the angle of incidence is equal to the angle of reflection. In the presence of waves, one generically sees a line of images above and below the unperturbed image. The grazing angle geometry simplifies the lensing geometry, effectively reducing it from a two dimensional problem to one dimension. The statistics of such reflections is sometimes called glitter, and has many solvable properties (Longuet-Higgins 1960). A similar effect occurs when the observer is below the surface. Two major distinctions arise: 1. the waves can deform the surface to create caustics in projection. Near caustics, Snell's law can lead to highly amplified refraction angles. 2. due to the odd image theorem, each caustic leads to two images. In practice, the surface could be caused by reconnection sheets (Braithwaite 2015), which have finite widths to regularize these singularities. Diffusive structures have Gaussian profiles, which was analysed in Pen & King (2012). The lensing details differ for convergent (underdense) *vs* divergent (overdense) lenses, first considered by Clegg et al. (1998).

The generic interstellar electron density is insufficient to deflect radio waves by the observed \sim mas bending angles. At grazing incidence, Snell's law results in an enhanced bending angle, which formally diverges. Magnetic discontinuities generically propagate as transverse surface waves, whose restoring force is the change in Alfvén speed on the two sides of the discontinuity. This completes the analogy to waves on a lake: for sufficiently inclined sheets the waves will appear to fold back onto themselves in projection on the sky. At each fold caustic, Snell's law diverges, leading to enhanced refractive lensing. The divergence is cut off by a finite width of the sheet. The generic consequence is a series of collinear images. Each projected fold of the wave results in two density caustics. Each density caustic leads to two geometric lensing images, for a total of 4 images for each wave inflection. The two geometric image in each caustic are separated by the characteristic width of the sheet, if this is smaller than the Fresnel scale, the two images become effectively indistinguishable.

A large number of sheets might intersect the line of sight to any pulsar. Only those sufficiently inclined would lead to caustic formation. Empirically, some pulsar scattering appears dominated by a single sheet, leading to the prominent inverted arclets (Stinebring et al. 2001).

2.1 Archival data of B0834+06

Our analysis is based on the apex data selected from the secondary spectrum of pulsar B0834+06 in B10, which was observed as part of a 300 MHz global VLBI project on 2005 November 12, with GBT (GB), Arecibo (AR), Lovell and Westerbork (WB) telescopes. The GB-AR and AR-WB baselines are close to orthogonal and of comparable lengths, resulting in relatively isotropic astrometric positions. Information from each identified apex includes delay τ , delay rate (differential frequency f_D), relative Right Ascension $\Delta\alpha$ and

relative declination $\Delta\delta$. Data of each apex are collected from four dual circular polarization 8 MHz wide sub-bands spanning the frequency range 310.5–342.5 MHz. As described in B10, the inverse parabolic arclets were fitted to positions of their apexes, resulting in a catalogue of apexes in each sub-band, each with delay and differential frequency. As previously described, the positions of the apexes appear constant across sub-bands. In this work, we first combine the apexes across sub-bands, resulting in a single set of images. We focus on the southern group with negative differential frequency: this grouping appears as a likely candidate for a double lensing screen. However, two groups (with positive/negative differential frequency) appear distinct in both the VLBI angular positions and the secondary spectra. We divide the apex data with negative f_D into two groups: in one group, time delay ranges from 0.1 ms to 0.4 ms, which we call 0.4 ms group, and in the other group, time delay at about 1 ms, which we call 1 ms group. In summary, the 0.4 ms group contains 10 apexes in the first two sub-bands, and 14 apexes in the last two sub-bands; the 1 ms group, contains 5, 6, 5 and 4 apexes in the four sub-bands subsequently, with center frequency of each band $f_{\text{band}} = 314.5, 322.5, 330.5$ and 338.5 MHz.

We select the equivalent apexes from four sub-bands. To match the same apexes in different sub-bands, we scale the differential frequency in different sub-bands to 322.5 MHz, by $f_D/f_{\text{band}} \cdot 322.5$ MHz. A total of 9 apexes from the 0.4 ms group and 5 apexes from the 1 ms group, were mapped by adjacent positions. This results in an estimation for the mean referenced frequency $f = 322.5$ MHz and a standard deviation among four sub-bands, listed in Table 1. The f_D , τ , $\Delta\alpha$ and $\Delta\delta$ are the mean values of n sub-bands, where n is the number of the sub-bands ($n = 3$ for points 4 to 6, nad 4 for the rest points), listed in Table 1.

We estimate the error of time delay τ , differential frequency f_D , $\Delta\alpha$ and $\Delta\delta$ listed in Table 1 from their band-to-band variance:

$$\sigma_{\tau, f_D, \Delta\alpha, \Delta\delta}^2 = \frac{1}{n} \sum_{i=1}^n \frac{(x_i - \bar{x})^2}{n-1}, \quad (1)$$

and n is the number of sub-bands. The outer $1/n$ accounts for the expected variance of a mean of n numbers.

2.2 One lens model

2.2.1 Distance to the lens

In the absence of a lens model, the fringe rate, delay and angular position cannot be uniquely related. To interpret the data, we adopt the lensing model of (Pen & Levin 2014). In this model, the lensing is due to projected fold caustics of a thin sheet closely aligned to the line of sight.

We define the *effective distance* D_e as

$$D_e \equiv \frac{2c\tau}{\theta^2}. \quad (2)$$

And the differential frequency is related to the rate of change of delay as $f_D = -f \frac{d\tau}{dt}$. The effective distance corresponds to the pulsar distance D_p , if the screen is exactly halfway. In general, $D_e = D_p D_s / (D_p - D_s)$ for a screen at D_s .

When estimating the angular offset of each apex, we subtract the expected noise bias: $\theta^2 = (\Delta\alpha \cos(\delta))^2 +$

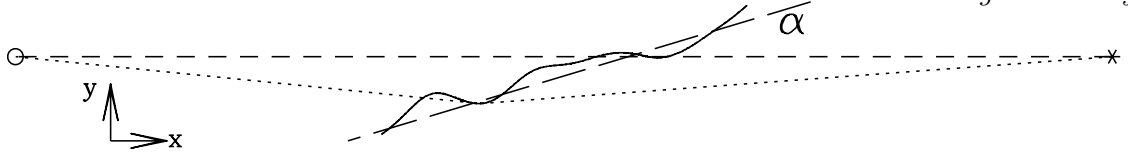


Figure 1. Lensing geometry. The earth is on the left and the pulsar is on the right. A section of the (crumpled) scattering sheet is shown in the middle as the solid line. It might be a plasma underdensity or overdensity of constance thickness. The dashed line shows the unperturbed light path. The dotted line shows the path of an iamge by a fold singularity of the projected sheet. The long-dashed line shows the position of the sheet in the absence of waves. The inclination between this sheet and the line of sight is α , in this plot $1/8$, chosen to exaggerate the effect for visual clarity and to reduce the computational cost. In practice, we envision values closer to 10×10^{-3} . See Pen & Levin (2014).

label	θ_{\parallel} (mas)	f_D (mHz)	τ (ms)	$\Delta\alpha$ (mas)	$\Delta\delta$ (mas)	t_0 (days)
1	-17.22	-26.1(4)	0.3743(6)	6.2	-11.9	-107
2	-16.36	-24.9(4)	0.3378(3)	8.0(4)	-14.5(8)	-101
3	-16.08	-24.6(4)	0.327(3)	7.2(6)	-13.9(4)	-99.0
4	-14.45	-22.3(5)	0.2633(3)	6.1(4)	-13.1(7)	-88.1
5	-13.68	-21.2(6)	0.236(2)	5.1(4)	-12.7(5)	-83.3
6	-13.27	-20.4(5)	0.222(3)	5.8(4)	-11.8(1)	-81.4
7	-12.21	-18.9(2)	0.188(2)	5.5(6)	-10.8(6)	-74.2
8	-10.58	-16.8(3)	0.1412(9)	3.9(6)	-10.0(4)	-62.8
9	-8.18	-12.9(2)	0.0845(5)	2.8(3)	-8.6(4)	-48.7
1'	...	-43.1(4)	1.066(5)	-8(3)	-24(2)	-185
2'	...	-41.3(5)	1.037(3)	-14(1)	-23(3)	-188
3'	...	-40.2(6)	1.005(8)	-14(1)	-22.3(5)	-187
4'	...	-38.3(6)	0.9763(9)	-14(1)	-20.6(3)	-190
5'	...	-35.1(5)	0.950(2)	-15(1)	-21(1)	-202

Table 1. 0.4 ms and 1 ms reduced apex data. The upper part of the table list the 0.4 ms group data, while the 1 ms group lie in the lower part of the table. Observation data include the differential frequency f_D , time delay τ (τ_1 for 0.4 ms group and τ_2 for 1 ms group); $\Delta\alpha$ and $\Delta\delta$ are from the VLBI measurement (there is only one matched position for point 1, thus no error). t_0 is the time at constant velocity for an apex to intersect the origin at constant speed along the main scattering parabola. More details in Section 2.1 and Section 2.2.2.

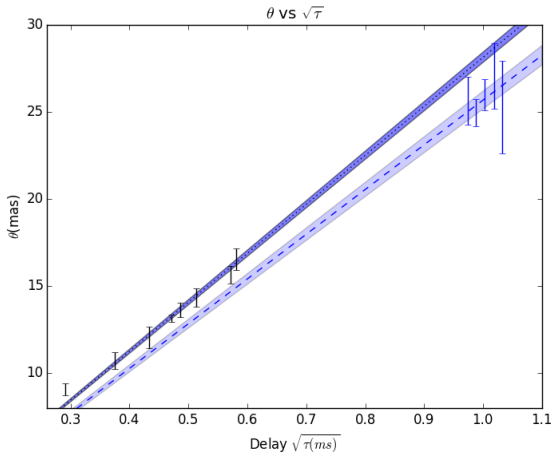


Figure 2. θ vs $\sqrt{\tau}$. Two separate lines through the origin were fitted to the points sampled among the 0.4 ms group and 1 ms group. The solid line is the fitted line of the 0.4 ms positions, where $D_{1e} = 1044$ pc with an error region of $\sigma_D = 22$ pc. The dashed lines are the fitted lines of the 1 ms position, where $D_{2e} = 1252$ pc with an error region $\sigma_D = 49$ pc.

$(\Delta\delta)^2 - \sigma_{\Delta\alpha}^2 - \sigma_{\Delta\delta}^2$. We plot the θ vs square root of τ in Figure 2. A least square fit to the distance results in

$D_{1e} = 1044 \pm 22$ pc for the 0.4 ms group, which we call lens 1 (point 1 is exempted for there is no error for that point, thus we cannot obtain the weighted mean for that point), and $D_{2e} = 1252 \pm 49$ pc for the 1 ms group, hereafter lens 2. The errors, and uncertainties on the error, precludes a definitive interpretation of the apparent difference in distance. At face value, this indicates that the lens 2 is closer to the pulsar, and we will use this as a basis for the model in this paper. The distances are slightly different from those derived in B10, which is partly due to a different subset of arclets analysed. We discuss consequences of alternate interpretations in Section 2.4. The pulsar distance was directly measured using VLBI parallax to be $D_p = 620 \pm 60$ pc. We take $D_{1e} = 1044$ pc, and the distance of lens 1 D_1 , where 0.4 ms group scintillation points are refracted, as 389 pc. Similarly, for 1 ms apexes, the distance of lens 2 is taken as 415 pc, slightly closer to the pulsar.

For the 0.4 ms group, we adopt the geometry from B10, assigning these points along line AD as shown in Figure 3 based solely on their delay, which is the best measured observable. The line AD is taken as a fixed angle of $\gamma = -25^\circ.2$ east of the declination axis. We use this axis to define \parallel and define \perp by a 90° clockwise rotation.

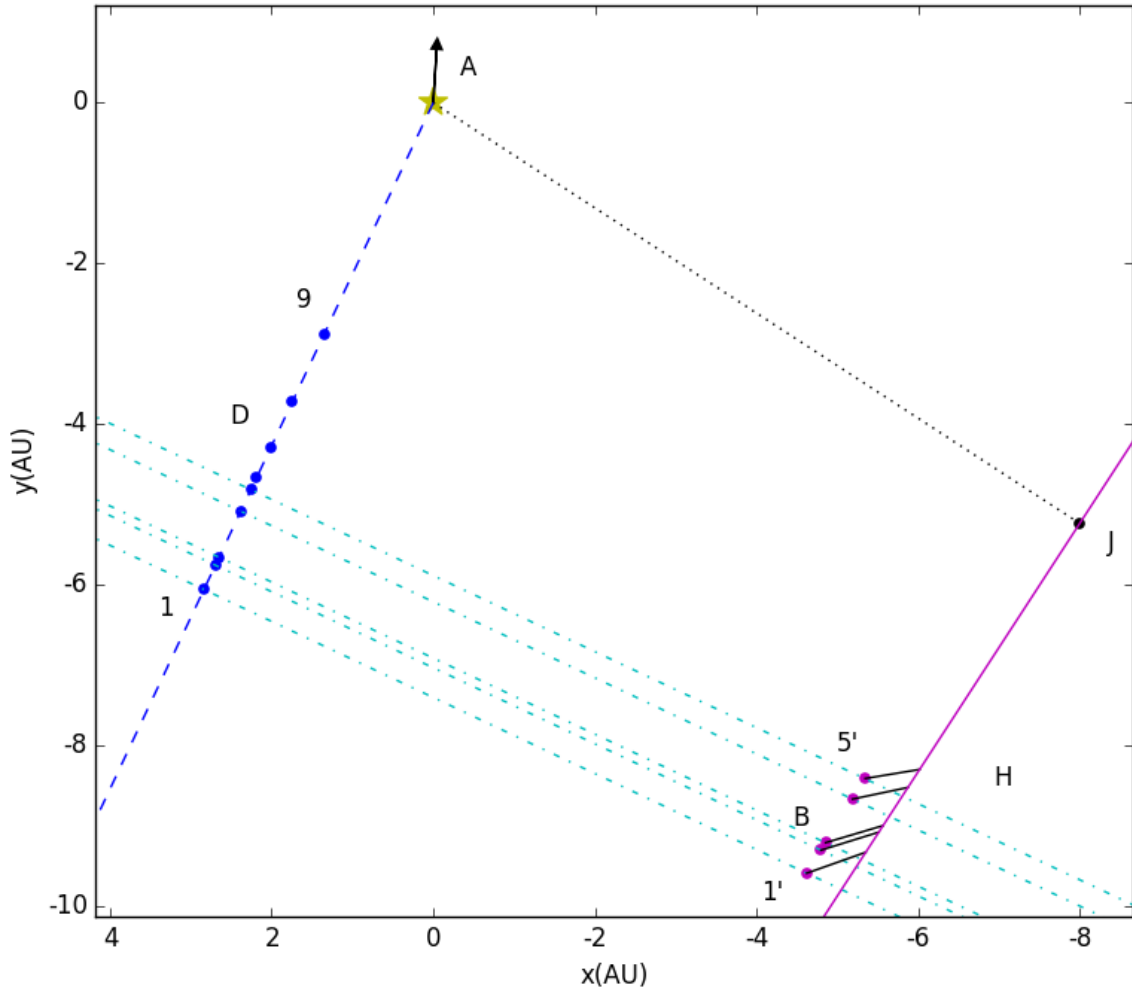


Figure 3. Angular positions of 0.4 ms and 1 ms group data in one lens model and double lens model. The axes represent the relative distance to the unrefracted pulsar in Right Ascension (calculated by $x = \Delta\alpha \cos(\delta)D$ and D represent the distance of the object to the observer) and declination (calculated by $y = \Delta\delta D$, with D defined previously) directions, on a 2D plane that is transverse to the line of sight. On the left side, the group points marked with letter D, are the derived positions from the time delays of 0.4 ms group in one lens model. 389 pc away from the observer, the dashed line aligns the 0.4 ms apexes positions, with a angle $\gamma = -25^\circ.2$ east of north. The points on the right side, mark the first and second refraction images in double lens model. The unobservable points denoted as letter H, are the calculated positions on lens 2 from 1 ms group; the observable points denoted as letter B, are the calculated positions on lens 1 from 0.4 ms group. These two piles of refracted images are connected by the short solid lines. The long solid line is the inferred geometry of the second lens; while the dot dash lines, constructed perpendicular to the AD scattering axis, denote the positions of lens 1. The dotted line on the top right is perpendicular to the solid line, intersecting at J. The relative motion pulsar-screen velocity is 185.3 km/s, with an angle $\epsilon = -3.7^\circ$ east of north, is marked with an arrow from the star, point A, at the top of the figure.

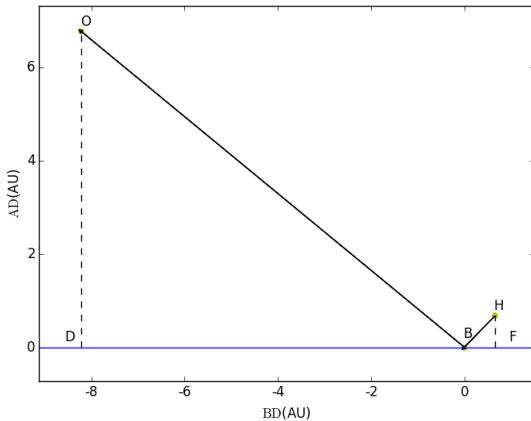


Figure 5. Refraction on lens 1. H is the lensed image on lens 2. B is the lensed image on lens 1. O is the position of the observer. $HF \perp DF$ and $OD \perp HJ$. As in the previous figure, the values are illustrated for point 1'.

2.3.2 Comparing model results in double lens model with observations

In order to compare τ , we calculate time delay τ_M for these five points, and list the results in Table 3. For points 4' and 1', they fit by construction since we use these to calculate the position of J; for the remaining three points, all of the results are within $3\text{-}\sigma$ of the observed time delays.

To compare differential frequency f_D , we need to calculate the velocity of the pulsar and the velocity of the lens. We take the lenses to be relative static, and solve for the velocity of the pulsar relative to the lens (in geocentric coordinate). The pulsar has two velocity components, and the two 1-D lenses effectively determine one component each. For $v_{||}$, we derive the velocity 172.4 ± 2.4 km/s from f_D of point 1 in 0.4 ms group. The direct observable is the time to crossing of each caustic, denoted t_0 in Table 1.

To calculate v_{\perp} , we choose the point 4', which has the smallest errorbar of differential frequency. With v_{\perp} to be 67.9 ± 2.8 km/s, we find $v_{\text{tot}} = 185.3 \pm 3.3$ km/s, with an angle $\epsilon = -3.7^\circ$, which is west of north. This velocity represents the pulsar's relative velocity to the earth, subtracting the screen's relative velocity to the earth. The model derived and observation velocities (heliocentric and in Local Standard of Rest) are listed in Table 2. The direction of the model velocity is marked on the top of the star in Figure 3.

With this velocity of the pulsar, we calculate the differential frequency f_M of points 5',3',2' and 1'. Results are listed in Table 3. The calculated results all lie within the $3\text{-}\sigma$ error intervals of the observed data.

The reduced χ^2 for time delay τ is 1.5 for 3 degrees of freedom and 2.2 for f_D for 4 degrees of freedom. This is consistent with the model.

Within this lensing model, we can test if the caustics are parallel. Using the lag error range of double lensed point 4 (the best constrained), we find a $1\text{-}\sigma$ allowed angle of 0.4 degrees from parallel with the whole lensing system. This lends support of a highly inclined sheet, probably aligned to better than 1%.

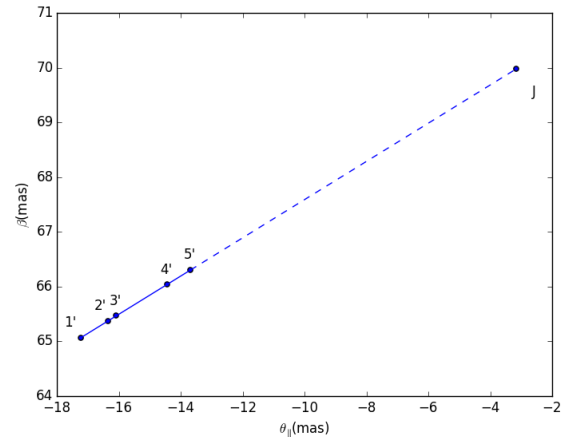


Figure 6. Deflection angle $\beta = \pi - \angle\text{AHB}$ on lens 2. Point J denotes the expected position to form a single refraction image, which was not observed. The small change in angle relative to the observed images precludes a finite refraction cut-off, since the data spans 10% bandwidth, with a 20% change in refractive strength. We propose a swallowtail caustic as the likely origin for the termination of the second lens sheet.

2.3.3 Discussion of double lens model

For 1 ms group, lens 2 only images a subset of the lens 1 images. This could happen if lens 1 screen is just under the critical inclination angle, such that only $3\text{-}\sigma$ waves lead to a fold caustic. If the lens 2 was at a critical angle, the chance of encountering a somewhat less inclined system is of order unity. More surprising is the absence of a single refraction image of the pulsar, which is expected at position J. This could happen if the maximum refraction angle is just below critical, such that only rays on the appropriately aligned double refraction can form images. We plot the refraction angle β in the direction that is transverse to the first lens plane in Figure 6. The data spans about 10% in frequency, making it unlikely that single lens image J would not be seen due to the larger required refraction angle. Instead, we speculate that the fold caustic could have formed near double lens image 1, and thus only intersections with the closer lens plane caustic south of image 1 are double lensed.

This is a generic outcome of a swallowtail catastrophe (Arnold 1990). In this picture, the sheet just starts folding near point 5'. North of point 5', no fold appears in projection. Far south of point 5', a full fold exhibits two caustics emanating from the fold cusp. Near the cusp the magnification is the superposition of two caustics, leading to enhanced lensing and higher likelihood of being observed.

We denote t_1 the time that the lensed image on lens 2 to move from point H to point J. From our calculation, 59 days earlier the single lensed pulsar image J would have appeared overlayed on point 5'; and 78 days before Nov 12, 2005, it overlayed on point 1'. The model predicts the presence of a single lensed image refracted at these points, in addition to the double lensed images.

Parameter	$\mu_{\alpha*}$ (mas/year)	μ_{δ} (mas/year)	μ_{l*} (mas/year)	μ_b (mas/year)	v_{l*} (km/s)	v_b (km/s)
model pulsar-screen velocity	-5.30 ± 1.11	61.97 ± 1.11	-56.45	22.23
VLBI pulsar proper motion	2.16 ± 0.19	51.64 ± 0.13	-46.69	28.02	-137.24	82.34
Screen motion	9.76	5.79	18.00	10.68

Table 2. Summary of velocities in double lens model. The velocities listed in equatorial coordinates are the relative velocity in heliocentric system, while the velocities in galactic coordinates are the relative velocities in LSR (Local Standard of Rest). $\mu_{\alpha*} = \Delta\alpha \cos(\delta)/t$ and $\mu_{l*} = \Delta l \cos(b)/t$, for we moved the center position from $(\Delta\alpha, \Delta\delta)$ or (l, b) to $(0,0)$. v_{l*} and v_b are the linear velocities relative to the LSR. The screen is only moving slowly (~ 21 km/s).

label	$\theta_{ }$ (mas)	τ_2 (ms)	σ_τ (ms)	τ_M (ms)	f_D (mHz)	σ_f (mHz)	f_M (mHz)	t_1 (days)
1'	-17.22	1.0663	0.0050	1.0663*	-43.08	0.84	-42.26	-78
2'	-16.36	1.0370	0.0059	1.0362	-41.27	0.88	-41.04	-73
3'	-16.08	1.005	0.011	1.027	-40.17	0.87	-40.64	-72
4'	-14.45	0.9763	0.00088	0.9763*	-38.31	0.64	-38.31†	-63
5'	-13.68	0.9495	0.0094	0.9550	-35.06	0.78	-37.21	-59

Table 3. Comparison of time delay τ and the differential frequency f_D of the observation and the model fitting result in the double lens model. $\theta_{||}$ denotes the angular offsets of the corresponding images at lens 1. The values with star symbols on them are the points that we use to calculate the position of J and the point with a † symbol is the point that we use to calculate the transverse velocity of the pulsar v_{\perp} . They agree with data by construction. The last column, t_1 is the time that the lensed image on lens 2 to move from point H to point J, which is also defined in Section 2.3.3.

2.4 Distance Degeneracies

With two lens screens, the number of observables increases: in principle one could observe both single reflection delays and angular positions, as well as the double reflection delay and angular position. Three distances are unknown, equal to the number of observables. Unfortunately, these measurements are degenerate, which can be seen as follows. From the two screens $i = 1, 2$, the two single deflection effective distance observables are $D_{ie} \equiv 2c\tau_i/\theta_i^2 = D_i^2(1/D_i + 1/D_{pi})$. A third observable effective distance is that of screen 2 using screen 1 as a lens, $D_{21e} = D_1^2(1/D_1 + 1/D_{21})$, within the triangle that is formed by lens 1, lens 2 and the observer. That is also algebraically derivable from the first two relations: $D_{21e} = D_{1e}D_{2e}/(D_{2e} - D_{1e})$.

In this archival data set, the direct single lens from the further plane at position J is missing. It would have been visible 59 days earlier. The difference in time delays to image J and the double reflection images would allow a direct determination of the effective distance to lens plane 2. Due to the close to 90 degree angle $\angle DAJ$ between lenses, the effect would be about a factor of 10 ill conditioned. With sufficiently precise VLBI imaging one could distinguish if the double refracted images are at position B (if lens 1 is closer to the observer) or position H (if lens 2 is closer to the observer). As described above, we interpret the effective distances to place screen 2 further away.

3 VLBI ASTROMETRY

PSR B0834+06 was observed 8 times with the Very Long Baseline Array (VLBA) between 2009 May and 2011 January. Four 16 MHz bands spread across the frequency range 1406 – 1692 MHz were sampled with 2 bit quantization in both circular polarizations, giving a total data rate of 512 Mbps per antenna. The primary phase calibrator was

J0831+0429, which is separated from the target by 2.1 degrees, but the target field also included an in-beam calibrator source J083646.4+061108, which is separated from PSR B0834+06 by only 5'. The cycle time between primary phase calibrator and target field was 5 minutes, and the total duration of each observation was 4 hours.

Standard astrometric data reduction techniques were applied (e.g., Deller et al. 2012, 2013), using a phase calibration solution interval of 4 minutes for the in-beam calibrator source J083646.4+061108. J083646.4+061108 is weak (flux density ~ 4 mJy) and its brightness varied on the level of tens of percent. The faintness leads to noisy solutions, and the variability indicates that source structure evolution (which would be translate to offsets in the fitted target position) could be present. Together, these two effects lead to reduce astrometric precision compared to that usually obtained with VLBI astrometry using in-beam calibration, and the results presented here could be improved upon if the observations were repeated using the wider bandwidths and higher sensitivity now available with the VLBA, potentially in conjunction with additional in-beam background sources.

While a straightforward fit to the astrometric observables yields a pulsar distance with a formal error $<1\%$, the reduced χ^2 of this fit is ~ 40 , indicating that the formal position errors greatly underestimate the true position errors, and that systematic effects such as the calibrator effects discussed above as well as residual ionospheric errors dominate. Accordingly, the astrometric parameters and their errors were instead obtained by bootstrap sampling (Efron & Tibshirani 1991). These results are presented in Table 4.

Table 4. Fitted and derived astrometric parameters for PSR B0834+06.

Reference right ascension (J2000) ^a	08:37:5.644606(9)
Reference declination (J2000) ^a	06:10:15.4047(1)
Position epoch (MJD)	55200
$\mu_{R.A.}$ (mas yr ⁻¹)	2.16(19)
μ_{Dec} (mas yr ⁻¹)	51.64(13)
Parallax (mas)	1.63(15)
Distance (pc)	620(60)
v_T (km s ⁻¹)	150(15)

^a The errors quoted here are from the astrometric fit only and do not include the ~ 1 mas position uncertainty transferred from the in-beam calibrator's absolute position.

4 DISCUSSIONS

4.1 Interpretation

The relative motion between pulsar and lens is directly measured by the differential frequency, and not sensitive to details of this model. B10 derived similar motions. This motion is in broad agreement with direct VLBI proper motion measurement, requiring the lens to be moving slowly compared to the pulsar proper motion or the LSR. The lens is ~ 200 pc above the galactic disk. Matter can either be in pressure equilibrium, or in free-fall, or some combination thereof. In free fall, one expects substantial motions. These data rule out retrograde or radially galactic orbits: the lens is co-rotating with the galaxy. In pressure equilibrium, gas rotates slower as its pressure scale height increases, which appears consistent with the observed slightly slower than co-rotating motion. The modest lens velocities also appear consistent with the general motion of the ISM, perhaps driven by galactic fountains (Shapiro & Field 1976) at these latitudes above the disk. In the inclined sheet model, the waves move at Alfvénic speed, but due to the high inclination, will move sub percent of this speed in projection on the sky, and completely negligible compared to other sources of motion.

Alternative models, for example, evaporating clouds (Walker & Wardle 1998) or strange matter (Pérez-García et al. 2013), do not make clear predictions. One would expect higher proper motions from these freely orbiting sources, larger future scintillation samples may constrain these models.

In order to incline one sheet randomly to better than 1% requires of order 10^4 randomly placed sheets, i.e. many per parsec. This sheet extends for ~ 10 AU in projection, corresponding to a physical scale greater than 1000 AU. These two numbers roughly agree, leading to a physical picture of magnetic domain boundaries every ~ 0.1 pc. B0834+06 has had noted arcs for multiple years, perhaps suggesting this dominant lens plane is larger than typical. One might expect to reach the end of the sheet within decades.

A generic prediction of the inclined sheets is a change in rotation measure across the scattering length. Over 1000 AU, one might expect a typical RM (rotation measure) change of 10^{-3} rad/m². At low frequencies, for example in LOFAR¹ or GMRT², the size of the scattering screen ex-

tends another order of magnitude in angular size, and RM changes increase to ~ 0.01 , which is plausibly measurable.

4.2 Possible Improvements

We discuss several strategies which can improve on the solution accuracy. The single biggest improvement would be to monitor over several months, as the pulsar crosses each individual lens, including both lensing systems. This allows a direct comparison of single lens to double lens arclets.

Angular resolution can be improved using longer baselines, for example adding a GMRT-GBT baseline doubles the resolution. Observing at multiple frequencies over a longer period allows for a more precise measurement: when the pulsar is between two lenses, the refraction angle β is small, and one expects to see the lensing at higher frequency, where the resolution is higher, and distances between lens positions can be measured to much higher accuracy.

Holographic techniques (Walker et al. 2008; Pen et al. 2014) may be able to measure delays, fringe rates, and VLBI positions substantially more accurately. Combining these techniques, the interstellar lensing could conceivably achieve distance measurements an order of magnitude better than the current published effective distance errors. This could bring most pulsar timing array targets into the coherent timing regime, enabling arc minute localization of gravitational wave sources, lifting any potential source confusion.

Ultimately, the precision of the lensing results would be limited by the fidelity of the lensing model. In the inclined sheet model, the images move along fold caustics. The straightness of these caustics depends on the inclination angle, which in turn depends on the amplitude of the surface waves. This analysis concludes a high degree of inclination.

5 CONCLUSIONS

We have applied the (Pen & Levin 2014) inclined sheet model to archival apex data of PSR B0834+06. The data is well fit by two linear lensing screens, with nearly plane-parallel geometry. The second screen provides a precision test with 10 observables (5 time delays and 5 differential frequencies) and 3 free parameters. The model fits the data to non-trivial percent accuracy on each of 7 data points. This natural consequence of very smooth reconnection sheets is an unlikely outcome of ISM turbulence. These results, if extrapolated to multi-epoch observations of binary systems, might result in accurate distance determinations and opportunities for removing scattering induced timing errors. This approach also opens the window to measuring precise transverse motions of the ionized ISM outside the galactic plane.

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¹ <http://www.lofar.org/>² <http://gmrt.ncra.tifr.res.in/>

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