

Notes 124(1) : Spiral Galaxy

The basic equation is :

$$\nabla \cdot \underline{g} = 4\pi G(R - \omega T) \quad (1)$$

where:

$$\underline{g} = cR \underline{J} \quad (2)$$

where \underline{J} is angular momentum. The velocity curve of the galaxy is :



The Newtonian region A is described by :

$$\nabla \cdot \underline{g} = 4\pi G R = 4\pi G \rho \quad (3)$$

i.e.

$$\nabla \cdot \underline{J} = 4\pi c R G R \quad (4)$$

In order to describe the region B, we use

$$R = \omega T \quad (5)$$

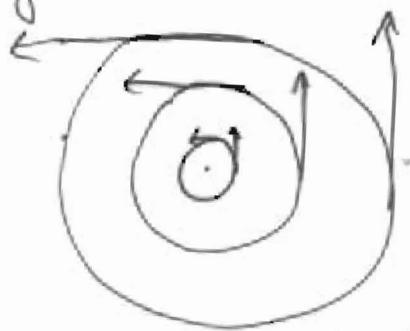
so

$$\nabla \cdot \underline{g} = \nabla \cdot \underline{J} = 0 \quad (6)$$

Therefore:

$$\underline{J} = -J_y \underline{i} + J_x \underline{j} \quad (7)$$

and look is follows:



This is beginning to look like a spiral galaxy or a whirlpool.

If the angular momentum is constant:

$$|\underline{J}| = (J_x^2 + J_y^2)^{1/2} \quad - (8)$$

= constant,

i.e.

$$J = mr\upsilon = mr^2\omega \quad - (9)$$

= constant

and

$$\upsilon = \omega r \quad - (10)$$

where

$$\omega = \frac{d\theta}{dt} \quad - (11)$$

If

$$\omega r = \text{constant} \quad - (12)$$

the

$$\boxed{\upsilon = \text{constant}} \quad - (13)$$

so regia B.

We have:

$$J = mr\upsilon \quad - (14)$$

3) and if v is constant, \mathcal{I} is proportional to r .
The angular velocity is:

$$\frac{d\theta}{dt} = \frac{mv}{r} \approx \frac{\text{constant}}{r} - (15)$$

So: $\theta = \frac{\text{constant}}{r} \int_0^T dt - (16)$

$$\boxed{\theta = \frac{\text{constant} \cdot T}{r}} - (17)$$

which is a spiral or a parabola.

124(2) : Angular Momentum as a Constant of Motion

Consider a particle of mass m moving in a central force field described by the potential U . If the latter depends only on the distance of the particle from the force centre and not on orientation, the angular momentum of the system is conserved:

$$\underline{J} = \underline{r} \times \underline{p} = \text{constant} \quad - (1)$$

(J.B. Marion and S.D. Tinkham, "Classical Dynamics", HCB 1988, 3rd ed., page 246).

The Lagrangian analysis of this problem defines the Lagrangian in plane polar coordinates r and θ as follows:

$$L = \frac{1}{2} m (\dot{r}^2 + r^2 \dot{\theta}^2) - U(r) \quad - (2)$$

The angular momentum conjugate to the coordinate θ is conserved:

$$\dot{p}_\theta = \frac{\partial L}{\partial \dot{\theta}} = 0 = \frac{d}{dt} \frac{\partial L}{\partial \dot{\theta}} \quad - (3)$$

i.e.

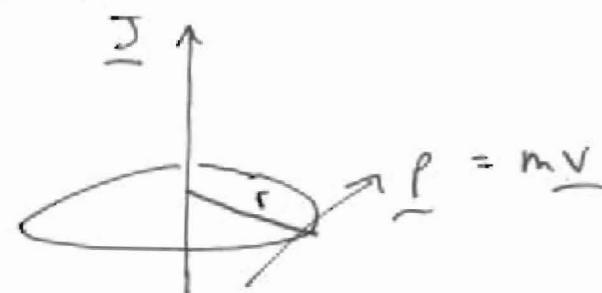
$$J = p_\theta = \frac{\partial L}{\partial \dot{\theta}} = mr^2 \dot{\theta} = \text{constant} \quad - (4)$$

The angular velocity is defined by:

$$\omega = \dot{\theta} = \frac{d\theta}{dt} \quad - (5)$$

so

$$\boxed{J = mr^2 \omega = mr\dot{\theta} = \text{constant}} \quad - (6)$$



The angular momentum \mathbf{J} is the first integral of the motion. The infinitesimal area dA is:

$$dA = \frac{1}{2} r^2 d\theta \quad - (7)$$

and the areal velocity is:

$$\frac{dA}{dt} = \frac{1}{2} r^2 \frac{d\theta}{dt} = \frac{1}{2} r^2 \omega \quad - (8)$$

$$\boxed{\frac{dA}{dt} = \frac{\mathbf{J}}{2m} = \text{constant}} \quad - (9)$$

This is Kepler's second law of planetary motion.

This is true for all central-force motion, not just for an inverse square law.

If the plane of, for example, a galaxy is a $X-Y$, the angular momentum is:

$$\underline{\mathbf{J}} = m r^2 \omega \underline{\mathbf{k}} \quad - (10)$$

where:

$$\underline{\mathbf{r}} = r_x \underline{\mathbf{i}} + r_y \underline{\mathbf{j}}, \quad - (11)$$

$$r = |\underline{\mathbf{r}}| = \left(r_x^2 + r_y^2 \right)^{1/2}, \quad - (12)$$

$$\text{so } \nabla \cdot \underline{\mathbf{J}} = \frac{dJ}{dz} = 0 \quad - (13)$$

$$\text{and } \underline{\mathbf{p}} = p_x \underline{\mathbf{i}} + p_y \underline{\mathbf{k}}. \quad - (14)$$

The ECE equation of motion is:

$$\nabla \cdot \underline{\mathbf{g}} = c^2 (R - \omega T) \quad - (15)$$

3) where:

$$\underline{\underline{g}} = c^2 \left(T^{01} \underline{i} + T^{02} \underline{j} + T^{03} \underline{k} \right) \quad - (16)$$

Now integrate over volume:

$$T^{0i} = \int T^{0i} dV = VT^{0i} \quad - (17)$$

$$i = 1, 2, 3.$$

and define the angular momentum tensor:

$$\underline{\underline{J}} = \frac{c}{k} T^{0i} \quad - (18)$$

Therefore:

$$\boxed{\underline{\underline{g}} = \frac{ck}{V} \underline{\underline{J}}} \quad - (19)$$

Units check

$$g = m s^{-2}; \quad c = m s^{-1}, \quad k = m kg m^{-1}, \quad \checkmark$$

$$J = kg m^2 s^{-1}, \quad V = m^3$$

Therefore:

$$\underline{\nabla} \cdot \underline{\underline{g}} = \frac{ck}{V} \underline{\nabla} \cdot \underline{\underline{J}} = 4\pi G \rho \quad - (20)$$

where: $G = \frac{c^2 k}{8\pi}$ $- (21)$

Therefore

$$\boxed{\underline{\nabla} \cdot \underline{\underline{J}} = \rho = \frac{1}{2} V c \rho} \quad - (22)$$

4) and:

$$\nabla \cdot \underline{J} = \frac{Vc}{R} (R - \omega T) \quad -(23)$$

If it assumed that the mass density is

$$\rho = \frac{m}{V} \quad -(24)$$

then:

$$\nabla \cdot \underline{J} = \frac{1}{2} mc \quad -(25)$$

Denote: $\nabla (R - \omega T) := (R - \omega T)' \quad -(26)$

then:

$$\begin{aligned} \nabla \cdot \underline{J} &= \frac{c}{R} (R - \omega T)' \\ &= \frac{1}{2} mc \end{aligned} \quad -(27)$$

and

$$m = \frac{2}{R} (R - \omega T)' \quad -(28)$$

Therefore mass is proportional to $(R - \omega T)'.$

The fundamental dynamics of the Newton and Coulomb laws are given by eq. (25), where mass m acts as a source of the field $\underline{J}.$

5) Central Motion

In central motion:

$$\nabla \cdot \underline{J} = 0 \quad - (29)$$

and so:

$$R = cT \quad - (30)$$

is a general law of central motion. Therefore Kepler's second law is $\underline{\text{Eq. } (20)}$. Eq. (30) means that:

$$m = 0 \quad - (31)$$

meaning out there is no mass outside the plane of the orbit.

Kepler's First and Third Laws

These depend specifically on the inverse square law of Newton. If Newtonian dynamics are defined by:

$$\omega \rightarrow 0 \quad - (32)$$

(spin correction goes to zero), then the first and third laws are given by:

$$\nabla \cdot \underline{J} = \frac{c}{k} R' \quad - (33)$$

$$R' = R^{10} + R^{20} + R^{30} \quad - (34)$$

124(3) : Relation Between Torsion and Spin Concentration for General Central Orbits.

The general relation between force law and orbit is given from a Lagrangian analysis by eq. (7.20) of Maria and Thornton. This can be expressed as :

$$F(r) = -mv^2 \left(\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) + \frac{1}{r} \right) \quad (1)$$

where m is the mass of the attracting particle. The potential energy is : -(2)

$$U(r) = - \int F dr = m \int v^2 \left(\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) + \frac{1}{r} \right) dr$$

and the gravitational potential is :

$$\boxed{\Gamma = \frac{1}{m} U(r) = \int v^2 \left(\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) + \frac{1}{r} \right) dr} \quad (3)$$

Maria and Thornton write eq. (1) as :

$$\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) + \frac{1}{r} = - \frac{mr^2}{J^2} F(r) \quad (4)$$

where

$$J = mr^2\omega = mr^2\dot{\theta} = \text{constant.} \quad (5)$$

Therefore :

$$F(r) = -r^2 \omega^2 m \left(\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) + \frac{1}{r} \right). \quad (6)$$

Define:

$$v = r\omega \quad (7)$$

Let the orbital linear velocity. The angular velocity is:

$$\omega = \frac{d\theta}{dt} \quad (8)$$

In ECE theory: (9)

$$\frac{d}{dt} = -\nabla \Phi + \underline{\omega} \underline{\Phi} = c^2 T$$

Eq. (9) gives a relation between the torsion T and the spin connection $\underline{\omega}$ for a given $\underline{\Phi}$, defined by eq. (3). From Kepler's second law:

$$\frac{dA}{dt} = \frac{1}{2} r^2 \omega = \text{constant} \quad (10)$$

Therefore:

$$v = \frac{2}{r} \frac{dA}{dt} \quad (11)$$

Therefore $v = \frac{k}{r} \quad (12)$

where:

$$k = 2 \frac{dA}{dt} = \text{constant} \quad (13)$$

for all central orbits.

∴ Therefore in eq. (3) :

$$\underline{\Phi} = \int \frac{k^2}{r^2} \left(\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) + \frac{1}{r} \right) dr \quad - (14)$$

where : $R = 2 \frac{dA}{dt}$. $\quad - (15)$

Since R is a constant :

$$\boxed{\underline{\Phi} = R^2 \int \left(\frac{1}{r^2} \left(\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) + \frac{1}{r} \right) dr \right)} \quad - (16)$$

Newtonian orbits

These are given by :

$$\frac{1}{r} = \frac{1}{d} (1 + \epsilon \cos \theta) \quad - (17)$$

(M&T eq. (7.41)), and by :

$$\boxed{\underline{\Phi} = -\frac{Gm}{r}} \quad - (18)$$

(M&T eq. (4.6)). The Newtonian orbits are

particular case of eq. (16) (M&T page 250)

Log Spiral orbits

$$\text{These are : } r = R \exp(\alpha \theta) \quad - (19)$$

and $\underline{\Phi} = R^2 (1 + \alpha^2) \int \frac{1}{r^2} dr$

$$\boxed{\underline{\Phi} = -\frac{1}{2} R^2 (1 + \alpha^2) \frac{1}{r^2}} \quad - (20)$$

24 (4): Kepler's Equation
 In astronomy we need the function $\theta(t)$, so the direction of an orbiting object may be found at any time. Planetary motion for example is described by:

$$\frac{d}{r} = 1 + e \cos \theta \quad - (1)$$

i.e. the equation of a conic section with one focus at the origin (Marsden & Thornton eq. (7.41)). Here e is the eccentricity and $2L$ is the latus rectum. Define the period of the orbit as T . This is the time taken for the radius vector to sweep out the entire area πab of an elliptical orbit. By Kepler's second law, the area $(\pi ab / T)t$ is swept out (M & T page 261). Thus:

$$dA = \frac{\pi ab}{T} t \quad - (2)$$

$$dA = \frac{1}{2} r^2 d\theta \quad - (3)$$

where

$$\text{If } \theta = 0 \text{ at } t = 0: \quad \frac{\pi ab}{T} t = \frac{1}{2} \int_0^\theta r^2 d\theta \quad - (4)$$

$$\text{where, from eq. (1): } r = \frac{d}{1 + e \cos \theta} \quad - (5)$$

$$\text{so } \frac{\pi ab}{T} t = \frac{d^2}{2} \int_0^\theta \frac{d\theta}{(1 + e \cos \theta)^2} \quad - (6)$$

Now we:

$$ab = \alpha^3 (1-\epsilon^2)^{-3/2} \quad - (7)$$

So:

$$\boxed{\frac{2\pi t}{\tau} = 2 \tan^{-1} \left(\left(\frac{1-\epsilon}{1+\epsilon} \right)^{1/2} \tan \frac{\theta}{2} - \frac{\epsilon (1-\epsilon^2)^{1/2} \sin \theta}{1+\epsilon \cos \theta} \right)} \quad - (8)$$

(Match eq. (7.53)).

This equation must be inverted to give θ as a function of t . It may be possible to use Maxima to do this. An approximate result is:

$$\theta(t) = \frac{2\pi t}{\tau} + 2\epsilon \sin \left(\frac{2\pi t}{\tau} \right) + \frac{5}{4} \epsilon^2 \sin \left(\frac{4\pi t}{\tau} \right) + \frac{1}{12} \epsilon^3 \left(13 \sin \left(\frac{6\pi t}{\tau} \right) - 3 \sin \left(\frac{2\pi t}{\tau} \right) \right) + \dots \quad - (9)$$

This problem was also solved by Kepler's equation:

$$\boxed{\frac{2\pi t}{\tau} = \phi - \epsilon \sin \phi} \quad - (10)$$

$$\text{where : } \boxed{\tan \frac{\theta}{2} = \left(\frac{1+\epsilon}{1-\epsilon} \right)^{1/2} \tan \frac{\phi}{2}} - (11)$$

so ϕ is found as a function of t from eq. (10), by inversion, and ϕ is related to θ by eq. (11).

Velocity as a Function of the Radius Vector

This is found from Kepler's equation using:

$$v^2 = \dot{x}^2 + \dot{y}^2 - (12)$$

$$= a^2 \dot{\phi}^2 (1 - \epsilon^2 \cos^2 \phi) - (13)$$

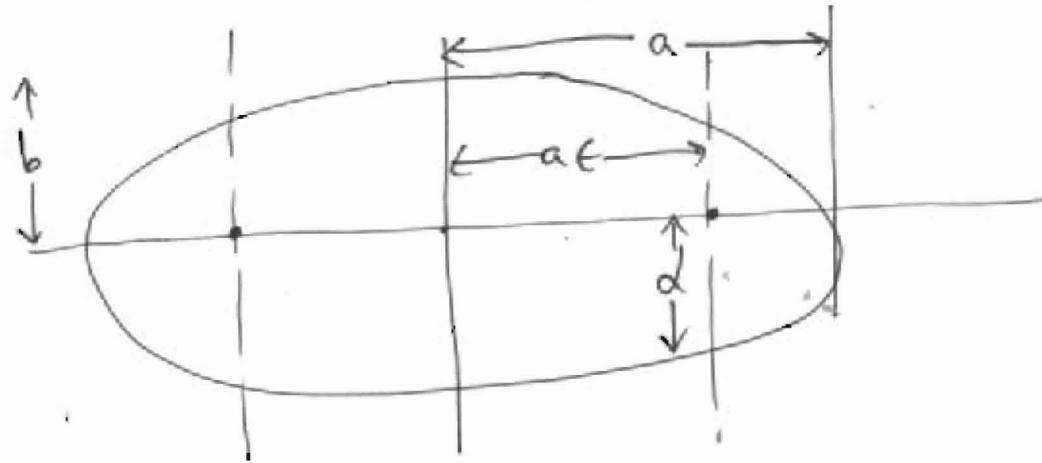
(m & T eq. (7.68)). So:

$$\boxed{v^2 = \frac{k}{\mu} \left(\frac{2}{r} - \frac{1}{a} \right)} - (14)$$

Kepler's third law is:

$$\boxed{\tau^2 = \frac{4\pi^2 \mu}{k} a^3} - (15)$$

The inverse sq. law is: $F = -\frac{k}{r^2}$
 $k = m_1 m_2 G$



The angular velocity is defined by:

$$\omega = \frac{d\theta}{dt} = \frac{J}{mr^2} = \frac{\text{constant}}{r^2} \quad (1)$$

So :

$$\theta = \text{constant} \int \frac{1}{r^2} dt \quad (2)$$

In their chapter seven, Maria and Thonta give :

$$\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) + \frac{1}{r} = - \frac{mr^2}{J^2} F(r) \quad (3)$$

which is valid for any central force law $F(r)$. This gives plenty of scope for applications.

From eq. (3) :

$$\frac{d\theta}{dt} = \frac{J}{mr^2} = - J F(r) \left(\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) + \frac{1}{r} \right)^{-1} \quad (4)$$

where

$$J = mr^2 \omega = \text{constant of motion} \quad (5)$$

So the orbit for any central force law is :

$$\theta = - J \int F(r) \left(\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) + \frac{1}{r} \right)^{-1} dt \quad (6)$$

The Inverse Square Force Law

This is : $F(r) = - \frac{k}{r^2} = - \frac{m_1 m_2 G}{r^2} \quad (7)$

2) Eq. (7) gives:

$$\frac{1}{r} = \frac{1}{a} (1 + \epsilon \cos \theta) \quad (8)$$

where a and ϵ are the half major axis and eccentricity of an ellipse. So the dependence of $1/r$ on θ is given by eq. (8) for an inverse square law.

If the orbit is a circle:

$$\epsilon = 0 \quad (9)$$

so:

$$\theta = - \int r F(r) dt \quad (10)$$

$$= Jk \int \frac{1}{r} dt$$

$$\boxed{\theta = \frac{Jkt}{r}} \quad (11)$$

because r is constant in time for a circular orbit.

More generally, r is a function of t in eq. (6). This can be seen from the fact that θ is a function of t in eq. (8). The dependence of r on t for the inverse square law is given by Kepler's equation:

$$v^2 = \frac{k}{m} \left(\frac{2}{r} - \frac{1}{a} \right) \quad (12)$$

where

$$\underline{v} = \frac{dx}{dt} \underline{i} + \frac{dy}{dt} \underline{j} \quad (13)$$

$$\underline{v}^2 = \dot{x}^2 + \dot{y}^2 \quad (14)$$

So :

$$\frac{dr}{dt} = \left(\frac{k}{m} \left(\frac{2}{r} - \frac{1}{a} \right) \right)^{1/2}$$

$$\frac{dt}{dr} = \left(\frac{k}{m} \left(\frac{2}{r} - \frac{1}{a} \right) \right)^{-1/2} \quad - (15)$$

$$t = \int \left(\frac{k}{m} \left(\frac{2}{r} - \frac{1}{a} \right) \right)^{-1/2} dr \quad - (16)$$

- 1) The dependence of t on r may be obtained from eq. (16). From eqns. (6) and (15) :

$$\theta = - J \int F(r) \left(\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) + \frac{1}{r} \right)^{-1} \left(\frac{k}{m} \left(\frac{2}{r} - \frac{1}{a} \right) \right)^{-1/2} dr$$

\\$ - (17)

- 2) and the dependence of θ on r may be obtained from this equation.
- 3) The dependence of r on t may be found by inverting the results of eq. (16) and parameterizing numerically.
- 4) The simplest method is to use eq. (2) and eq. (16). The latter is integrated by Maxima to give t as a function of r , and the result inverted to give r as a function of t .

4) Log Spiral Orbit

This is worked out completely by M & T, pp. 250 and 251. If a particle moves on a log spiral:

$$r = k \exp(\alpha\theta) \quad (1)$$

where k and α are constants. Eq. (1) gives:

$$\mathbf{F}(r) = -\frac{J^2}{mr^3} (1 + \alpha^2) \hat{r} \quad (2)$$

and $\theta = \frac{1}{2\alpha} \log \left(\frac{2\alpha J t}{mk^2} + c \right) \quad (3)$

wh $r = \left(\frac{2\alpha J t}{m} + k^2 c \right)^{1/2} \quad (4)$

Animate eqns. (1), (3) and (4).

These were first discussed by Roger Cotes (1682 - 1716) and the orbits are Cotes' Spirals.



4(6): Effect of Constant Spacetime Torsion on the Newtonian Attraction of Stars in a Galaxy

In general, the total energy in a central orbit of any kind is constant:

$$E = T + U = \text{constant} \quad - (1)$$

$$E = \frac{1}{2} m r^2 + \frac{1}{2} \frac{J^2}{mr^2} + U(r) \quad - (2)$$

Therefore:

$$\frac{dr}{dt} = \left(\frac{2}{m} (E - U) - \frac{J^2}{m^2 r^2} \right)^{1/2} \quad - (3)$$

where E and J are constants. The inverse square potential is:

$$U = m \bar{V} = - \frac{6m^2}{r} \quad - (4)$$

Therefore, θ in M & T chapter 7:

$$\theta = \int \frac{(J/r^2) dr}{2m(E - U(r) - \frac{J^2}{2mr^2})^{1/2}} \quad - (5)$$

One may animate θ as a function of r .

The effective potential is:

$$\bar{V}(r) = U(r) + \frac{J^2}{2mr^2} \quad - (6)$$

$$\bar{V}(r) = - \frac{6m^2}{r} + \frac{J^2}{2mr^2}$$

2)

and consists of the attractive and centrifugal repulsion, i.e. respectively negative and positive.

The Effect of Constant Spacetime Torsion

From previous work it is known that torsion, when integrated over volume, is proportional to angular momentum. Therefore in a galaxy, there is an additional repulsive potential due to constant spacetime torsion:

$$V(\text{torsion}) = \frac{J_T^2}{2mr^3} \quad - (7)$$

The extra repulsive force is:

$$\begin{aligned} F(\text{torsion}) &= -\frac{dV(\text{torsion})}{dr} \\ &= +\frac{J_T^2}{mr^3} \end{aligned} \quad - (8)$$

From problem (7.22) of m & T , if:

$$J_0^2 := J_T^2 (1 + \alpha_0) \quad - (9)$$

then this is produced by eq. (8) is a logarithmic spiral:

$$r = k \exp(\alpha_0 \theta) \quad - (10)$$

Therefore this simple model produces the main features of a spiral galaxy. The Newtonian curve of θ versus r is given by eq. (5), and the spiral orbits by eq. (10):

$$\theta = \frac{1}{\alpha_0} \log_e \frac{r}{R} - (11)$$

Eq. (5) gives:

$$\cos \theta = \frac{1}{E} \left(\frac{d}{r} - 1 \right)$$

$$\theta = \cos^{-1} \left(\frac{1}{E} \left(\frac{d}{r} - 1 \right) \right) - (12)$$

Animations can be made of θ as a function of r for eqns. (11) and (12).

The arms of a whirlpool galaxy are logarithmic spirals. We have:

$$\frac{dr}{d\theta} = br - (13)$$

$$\text{and } v = \frac{dr}{d\theta} \frac{d\theta}{dt} = \omega \frac{dr}{d\theta} = b \omega r \\ = \text{constant} - (14)$$

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1) Note 124(8): Logarithmic spiral orbit

Consider the lagrangian: $L = \frac{1}{2} m (\dot{r}^2 + r^2 \dot{\theta}^2)$ — (1)

$$L = \frac{1}{2} m (\dot{r}^2 + r^2 \dot{\theta}^2)$$

and the Euler-Lagrange equation: $\frac{d}{dt} \frac{\partial L}{\partial \dot{r}} = \frac{d}{dt} \frac{\partial L}{\partial \dot{\theta}}$ — (2)

$$\frac{d}{dt} \frac{\partial L}{\partial \dot{r}} = \frac{d}{dt} \frac{\partial L}{\partial \dot{\theta}} = m \ddot{r} + r \dot{\theta}^2 = m \ddot{r} - m r \ddot{\theta} \quad (3)$$

Thus: $\frac{\partial L}{\partial \dot{r}} = m \ddot{r}$, $\frac{d}{dt} \frac{\partial L}{\partial \dot{\theta}} = m \ddot{\theta}$ — (4)

$$\frac{\partial L}{\partial \dot{r}} = -m r \dot{\theta}^2 \quad (5)$$

Thus: $\dot{r} = -r \dot{\theta}^2$ — (6)

Making the transformation variable $u = \frac{1}{r}$ — (7)
and using the method of Narita and Tanaka page

$$249: \dot{r} = -\frac{J^2 u^2}{m^3} \frac{du}{d\theta} \quad (7)$$

$$\dot{r} = -\frac{J^2 u^2}{m^3} \frac{du}{d\theta} \quad (8)$$

$$\dot{\theta}^2 = \frac{J^2}{m^3} u^3 \quad (9)$$

The centripetal force is:

2) Using eq. (8)

$$F = -\frac{dU}{dr} = m r \dot{\theta}^2 = \frac{J^2}{m r^3} - (9)$$

which is the centrifugal force outwards.

Therefore

$$mr^2 \frac{-dU}{dr} = 0 - (10)$$

and the orbital equation is

$$\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) = -\frac{1}{r} - (11)$$

This is satisfied by ~~the~~ log spiral

$$r = e^{\theta} - (12)$$

We have:

$$r(t) = \left(\frac{2Jt}{m} + C \right)^{1/2} - (13)$$

and

$$\theta(t) = \frac{1}{2} \log_e \left(\frac{2Jt}{m} + C \right) - (14)$$

giving the evolution of a whirlpool galaxy.

3) The velocity of a star is given by the formula

$$V^2 = r^2 + \dot{r}^2 - (15)$$

 and may be calculated from eqs. (13) and
 (14).

Animations Eqs. (13) to (15) may be animated. If

the velocity V becomes constant, then it is

$$r^2 + \dot{r}^2 = \text{constant} - (16)$$

and it is seen that eqs. (13), (14) and (16)
 give an equation which

gives the trajectory as resolved to a spiral
 time along the arms of the log spiral

At same critical time the spiral arms of the log spiral
 become straightened out as observed

in figure 1.1. It is observed that the spiral arms of the log spiral become straightened out as observed

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124(9): Velocity Curve of a Log Spiral orbit.

The velocity of the star on a logarithmic spiral orbit is:

$$v^2 = \dot{r}^2 + r^2 \dot{\theta}^2 \quad - (1)$$

where:

$$\dot{\theta} = \frac{J}{mr_0^2} e^{-2\theta\alpha} \quad - (2)$$

(Mazda and Thornton, eq. (7.23), page 250),

and:

$$\dot{r} = \frac{d_0 r_0 J}{m r^3} e^{d\theta} \quad - (3)$$

(Mazda and Thornton, eq. (7.28), page 251).

The log spiral is:

$$r = r_0 e^{d\theta} \quad - (4)$$

so

$$\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) = \frac{d^2}{r^2} \quad - (5)$$

Therefore:

$$v^2 = \left(\frac{d_0 r_0 J}{m} \right)^2 \frac{e^{2d\theta}}{r^4} + \left(\frac{J}{mr_0^2} \right)^2 r^2 e^{-4d\theta}$$

- (6)

2)

Define:

$$A = \frac{\alpha \cdot r_0 J}{m}, \quad B = \frac{J}{mr_0^2} - (7)$$

Re:

$$\nu^2 = \left(\frac{Ar e^{d\theta}}{r^3} \right)^2 + \left(Br e^{-2d\theta} \right)^2 \quad - (8)$$

Limits

1) As $r \rightarrow \infty$ $\nu \rightarrow Br e^{-d\theta}$ - (9)

2) As $r \rightarrow 0$ $\nu \rightarrow \infty$ - (10)

$\nu \rightarrow \infty$ - (11)

In R limit (9):

$$\nu \rightarrow \frac{J r e^{-d\theta}}{mr_0^2} - (12)$$

Using eq. (4):

$$\nu \rightarrow \frac{J}{mr_0} = \text{constant}$$

for any r the graph of ν against r is a plateau as observed experimentally.

1) 124(10): The Potential Energy Generated by Spacetime Torsion.

Consider a star moving along the radial direction \underline{e}_r , wth velocity:

$$\underline{v}_r = r \underline{\dot{e}}_r - (1)$$

By Newton's first law it would move along this direction permanently unless acted upon by an external force, \underline{F} . This force produces work done on the star:

$$\bar{W}_{12} = \int_1^2 \underline{F} \cdot d\underline{r}, - (2)$$

work which transforms the star from condition 1 to 2. The kinetic energy generated by \underline{v} of eq. (1) is:

$$T = \frac{1}{2} m \dot{r}^2. - (3)$$

The torsion of spacetime does work on the star while keeping T constant. The potential energy transferred to the star from spacetime is defined by:

$$\int_1^2 \underline{F} \cdot d\underline{r} = U_1 - U_2 - (4)$$

so:

$$\underline{F} = -\nabla U - (5)$$

This potential energy moves the star in a direction transverse to \underline{e}_r , producing the

2) Velocity: $\underline{v}_\theta = r \dot{\theta} \underline{e}_\theta$ - (6)

This process may be thought of as a particle moving outward from the centre of a rotating platform.

sparking
rotation

star moving
at uniform
velocity

The total velocity is $\underline{v} = \dot{r} \underline{e}_r + r \dot{\theta} \underline{e}_\theta$ - (7).

and the total energy is $E = \frac{1}{2} m (\dot{r}^2 + r^2 \dot{\theta}^2)$ - (8)

$E = T + U$ - (9)

where $T = \frac{1}{2} m \dot{r}^2$ - (10)

$U = \frac{1}{2} m r^2 \dot{\theta}^2$ - (11)

The lagrangian is $L = T - U = \frac{1}{2} m (\dot{r}^2 - r^2 \dot{\theta}^2)$ - (12)

3)

The Euler-Lagrange equation is:

$$\frac{d}{dt} \frac{\partial L}{\partial \dot{x}_i} = \frac{d}{dt} \frac{\partial L}{\partial \dot{\theta}} - (13)$$

and

$$\frac{d}{dt} \frac{\partial L}{\partial \dot{\theta}} = \frac{d}{dt} \frac{\partial L}{\partial \dot{\theta}} - (14)$$

The motion takes place in a plane, so:

$$\boxed{\begin{aligned} \ddot{\theta} &= \frac{d\dot{\theta}}{d\theta} = \text{constant} - (15) \\ &= mr^2 \dot{\theta} \end{aligned}}$$

The conservation of angular momentum:

$$\frac{d\dot{\theta}}{dt} = 0 - (16)$$

and the conservation of energy:

$$\frac{dE}{dt} = 0 - (17)$$

The potential energy is given by:

$$U = \frac{1}{2} mr^2 \dot{\theta}^2 = \frac{1}{2} \frac{I}{mr^2} \dot{\theta}^2 - (18)$$

and

$$\boxed{F = \frac{\dot{\theta}^2}{2mr^3}} - (19)$$

1) 124(1): For Law and const due to Constant Tension

The Euler Lagrange equation is: $\frac{d}{dt} \frac{dL}{dr} = \frac{d}{dr} \frac{dL}{dt} - (1)$

$$\frac{dL}{dr} = \frac{1}{2} m (r^2 - r^2 \theta^2) - (2)$$

wt $L = \frac{1}{2} m (r^2 - r^2 \theta^2)$ - (2)

$$so: F = m r \dot{\theta} = - \frac{du}{dr} - (3)$$

with $W_{12} = \int_1^2 F \cdot dr - (4)$

$$and F = - \nabla u - (5)$$

$$\text{Hence: } \boxed{\int_1^2 F \cdot dr = u_1 - u_2} - (6)$$

The work done is the change of potential energy:

$$W_{12} = u_2 - u_1 - (7)$$

Work is done by the star if $u_1 > u_2$. The

$$\text{fact is: } \boxed{\int_1^2 F \cdot dr = - \int_1^2 du = u_1 - u_2}$$

- (8)

2) If F is positive valued the face is repulsive by convention. Therefore a repulsive face means that U_1 is greater than U_2 by convention. If F is negative valued the face is attractive by convention, so U_2 is greater than U_1 by convention. The initial state of potential is chosen such that:

$$U_1 = 0 \quad (9)$$

$$\text{and the final state is: } U_2 = \frac{1}{2} m r^2 \theta^2 \quad (10)$$

$$\text{Therefore: } U_2 > U_1 \quad (11)$$

$$\boxed{W_{12} > 0} \quad (12)$$

and:

this means that work has been done on the star by the force of spacetime. By convention, the face is negative valued, meaning that the star is attracted by the whirling spacetime.

$$\text{Therefore: } F = m r \omega^2 = -m r \dot{\theta}^2 \quad (13)$$

$$\text{Now make the charge: } u = \frac{1}{r} \quad (14)$$

$$\text{so } \frac{du}{d\theta} = -\frac{1}{r^2} \frac{dr}{d\theta} = -\frac{1}{r^2} \frac{dr}{dt} \frac{dt}{d\theta} = -\frac{1}{r^2} \frac{i}{\dot{\theta}} \quad (15)$$

$$3) \text{ So } \frac{du}{d\theta} = -\frac{m}{J} i - (16)$$

$$\frac{d^2 u}{d\theta^2} = -\frac{m^2}{J^2} r^2 i - (17)$$

and

$$\boxed{F(r) = -\frac{J^2}{m^2 r^2} \frac{d}{d\theta^2} \left(\frac{1}{r} \right)} - (18)$$

If the orbit is a logarithmic spiral:

$$r = r_0 \exp(\alpha\theta) - (19)$$

$$\text{and } \frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) = \frac{d^2}{d\theta^2} \frac{1}{r} - (20)$$

$$\text{so } F(r) = -\frac{d^2 J^2}{m^2 r^3} - (21)$$

$$\text{so potential is: } u(\theta) = \frac{d^2 J^2}{m r^3} - (22)$$

$$\text{so } v^2 = \left(\frac{d^2 J^2}{m r^2} + \frac{J^2 \theta^2}{r^2} \right)^2 - (23)$$

The velocity is:

$$v^2 = \left(\frac{d^2 J^2}{m r^2} \right)^2 + \left(\frac{J^2 \theta^2}{r^2} \right)^2$$

$$\boxed{v^2 = \left(\frac{A e^{d\theta}}{r^2} \right)^2 + \left(B r e^{-d\theta} \right)^2}$$

$$v \rightarrow \infty \quad B r e^{-d\theta} = \frac{J}{m r_0}$$

$$- (24)$$

1) 124(12): Force Law for a Logarithmic Spiral
(Mava and Thorntor page 25°)

In this development consider a star moving with velocity
 in a plane. Let ω be due to the star by a

potential: $U(r) = -\frac{J^2}{2m} \left(1 + d^2\right)^{-\frac{1}{2}}$ — (1)

so: $F(r) = -\frac{J^2}{2m} \left(1 + d^2\right)^{-\frac{3}{2}}$ — (2)

The orbit is the logarithmic spiral:

$$r = r_0 \exp(\alpha\theta) — (3)$$

Eq. (1) is the potential energy due to constant

spacetime torsion

We have:

$$\theta(t) = \frac{1}{2d} \log e \left(\frac{2d J t}{m r_0^2} + c \right) — (4)$$

$$r(t) = \left(\frac{2d J t}{m} + r_0^2 c \right)^{1/2} — (5)$$

Animate eqns. (4) and (5)

$$\dot{\theta} = \frac{J}{m r^2} — (6)$$

$$\text{Also: } \dot{r} = \frac{d J}{m r} — (7)$$

2)

$$\sqrt{r^2 + \frac{r^2}{n^2}} = \frac{r^2 \omega^2}{J} - (8)$$

In the case of elliptical orbits, the angular momentum can further be written as $J = (1+d^2)^{1/2} r \omega$, where d is the semi-latus rectum. Substituting this value of J in eqn (8) we get the following result:

$\sqrt{r^2 + \frac{r^2}{n^2}} = \frac{(1+d^2)^{1/2} r \omega^2}{r^2}$

So:

$$J = m v r (1+d^2)^{-1/2} - (9)$$

In gravitation, m and v are constants. So, J is also a constant. This shows that the angular momentum is conserved.

Here J , d , r and C are constants. Eq.

(a) shows that the orbit evolves to that of a circle of radius $(1+d^2)^{-1/2} r$. In a circular orbit both r and ω are constant.

Let:

$$R = (1+d^2)^{-1/2} r - (10)$$

and so

$$\begin{aligned} J &= m v R \\ &= \text{constant} \end{aligned} - (11)$$

Re basic equation: $M \& \vec{r}$ eqn (7.21):

$$\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) + \frac{1}{r^2} = - \frac{m r^2 F(r)}{J^2} - (12)$$

124(13): Resonance in the Inverse Square Law of Gravitation

Attraction

Start with the equation:

$$\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) + \frac{1}{r^2} = -\frac{mr^3}{J^2} F(r) \quad (1)$$

Kepler's laws and Newtonian dynamics are given by:

$$F(r) = -\frac{k}{r^2} \quad (2)$$

which is the Inverse Square law of gravitational attraction.

Therefore:

$$\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) + \frac{1}{r^2} = \left(\frac{mk}{J^2} \right) \quad (3)$$

Now introduce an oscillatory structure into the familiar inverse square law:

$$F(r) = -\frac{k}{r^2} \cos(\omega\theta) \quad (4)$$

Then:

$$\boxed{\frac{d^2}{d\theta^2} \left(\frac{1}{r} \right) + \frac{1}{r^2} = \left(\frac{mk}{J^2} \right) \cos(\omega\theta)} \quad (5)$$

This is a D'Alembert Euler resonance equation.

At resonance:

2)

$$\frac{1}{r} = \frac{m k \cos(\alpha\theta)}{J^2(d^2 - 1)^{1/2}} \quad (6)$$

so if $d^2 = 1$, $d = \pm 1$ — (7)

then $\frac{1}{r} \rightarrow 0$ — (8)

and the force $F(r)$ becomes negative infinite.
 the mass m is attracted to the mass M by an infinite
 force, and the system implodes.
Application to counter gravitation
 when $\cos(\alpha\theta) < 0$ — (9)

$$F(r) > 0 \quad - (10)$$

then it is repulsive. At the resonance point (6) the object m is repelled from M .
 this is an example of resonant counter-gravitation,
 where the system explodes.