

Relationship between solar wind, D_{st} , and plasmasphere mass density on one-hour time scales

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Abstract

This paper looks at how various magnetosphere conditions behave around the onset of geomagnetic events. It confirms results from previous papers that a sudden drop in D_{st} correlates with a spike in equatorial mass density, while adding a level of depth and specification to the applicability of their results. It compares data at 1-hour averages to data at 1-day averages, and briefly to data at 27-day averages to examine if the trend holds at varying time scales and finds that it does, so long as the data is pre-processed to linearly remove a large time scale $F_{10.7}$ dependence. By then looking at an hourly average of variables from 24 hours before event onset to 48 hours after, short timescale trends can be discerned.

1 Introduction

Takahashi et al. [2006], show how trends in storm dependence on mass/density only appear in longer timescales. Looking at K_p vs mass (amu) in the range of 6 to 7 R_E , a trend shows up in the 1.5 day averages that doesn't appear in the 3 hour averages.

Denton et al. [2006] show how D_{st} affects the distribution of plasma density along different magnetic latitudes, and specifically along the same field lines as looked at in later papers (6-8 R_E). Though this shows that the trends for density may differ between field lines, it's mentioned mostly as a point for future research as the data used in this paper is already adjusted for one field line, as described by Takahashi et al. [2010].

Yao et al. [2008] looks at the differences in how D_{st} correlates with number density for different ions in different regions (ring current and plasma sheet), but still finds a general correlation during each of the four storms selected.

Takahashi et al. [2010] state that spikes in the Disturbance Storm Time (D_{st}) index coincide with significant changes in ρ_{eq} at an L-shell of 6.8 R_E . For five storms over a 20 day period two had ρ_{eq} spikes after the D_{st} drop, two had ρ_{eq} spikes before the drop,

and one showed little change in ρ_{eq} . They then show an epoch analysis where ρ_{eq} is seen to spike the day of a D_{st} drop, using a daily average of 30 minute ρ_{eq} and one hour D_{st} measurements.

2 Data Preparation

The parameters ρ_{eq} and $F_{10.7}$ used in this work are from the data associated with Denton [2007], with data available from 1980-1991; all other variables are from Kondrashov et al. [2014] over the time range of 1972-2013, which are on a 1-hour time grid. In cases where ρ_{eq} was available from multiple Geostationary Operational Environmental Satellites (GOES) satellites at the same time, the value from GOES 7 was selected to avoid overlapping data points. ρ_{eq} is the inferred mass density based on the 3rd harmonic frequency of magnetic field measurements. The smallest cadence for ρ_{eq} values is 10 minutes. To compute an hourly average over the same time range as the solar wind parameters, the median of all values in a given hourly range was used. Fill values were used for hours when no measurements were available.

In this work events are considered when D_{st} or ρ_{eq} crosses a threshold value. In finding events, all fill values were replaced with linearly interpolated values.

To check for potential bias in the data as used by our analysis, we checked how the data availability varied with hour, shown in Figure 1. We also looked at

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how storm conditions affected data availability. Both a two sample t-test for difference in means and a Wilcoxon rank sum test showed significance at the 1% level that D_{st} during periods with available data was of a different distribution than D_{st} during periods missing data. When testing for other significant D_{st} differences, the Wilcoxon test was significant while the t-test was not in pre-noon vs post-noon distributions, as well as $\rho_{eq} > 40$ vs $\rho_{eq} < 40$ distributions indicating perhaps an increase in variability without an increase in mean.

While the GOES Satellites tend to keep an average geostationary distance of around $L = 6.8R_E$ shown in Takahashi et al. [2010], the actual plasmopause location varies significantly, as shown in O'Brien and Moldwin [2003]. This means that during periods of large D_{st} the plasmopause may be far from the point of measurement of ρ_{eq} creating a discord in the correlation. Gallagher et al. [2000] provides a model for ρ_{eq} at a range of L-shells and a brief discussion of the elemental contributions.

3 Results

Figure 2 shows values of solar wind averages and mass density used in this study. We will briefly compare our results to other published results to verify our methods and handling of data.

3.1 D_{st} Events

Two event indicators are looked at in this study. The first is looking for a drop in DST below the threshold of $-50nT$ specified in Takahashi et al. [2010], dubbed the "onset", and then considering the timeframe a event until D_{st} passes back above the $-50nT$ threshold. This method finds 669 such periods between May 1983 and August 1991 with an average duration of 9 hours and a median duration of 3 hours. A subset of longer duration events will be looked at later. Figure 2 shows the average values of B_Z , V_{SW} , D_{st} , $F_{10.7}$, and Mass Density for a period of 24 hours before and 48 hours after a event onset, as well as a plot of mass density with the found event times highlighted.

This figure shows a definite spike in the Z component of the magnetic field, as well as the defined drop in DST , but no obvious change in mass density at an hourly timescale. This points to an issue with only looking at long-timescale trends between density and D_{st} , and allows for the possibility that other factors are influencing the long term correlation since there's no obvious connection on a short timescale. One pos-

sibility is that, as suggested in Takahashi et al. [2010], $F_{10.7}$ plays a significant role in driving long term density values which biases the long term correlation of density and D_{st} .

3.2 Mass Density Events

Figure 4 shows this same algorithm, but looking for a rise in mass density over a value of $40g/cm^3$. This results in 130 events with a mean duration of 32 hours and a median duration of 17 hours, marked in red on the left figure.

This shows that when using all data for mass density derived events, almost no significant changes can be seen around event onset.

4 More specific events

It's hypothesized that progressively picking more specific event criteria will allow for the possibility of more significant results, at the expense of more bias in the selection process and potentially less overall usefulness of the results. That said, the predictability of extreme events is of definite interest, so an attempt has been made to find some reproducible method of prediction. Looking at events that last longer than 12 hours and events with an onset threshold greater than $70g/cm^3$ results in the left and right sides of Figure 5 respectively.

Neither of these seem to indicate anything too significant, so looking at D_{st} events instead to look for something that causes a significant change in Mass Density results in Figure 6. This shows that by either looking only at D_{st} events that last longer than an hour (left) or at events where the onset condition is $D_{st} < -80nT$, a spike in mass density is seen, but also a definite lack of data availability to the point where that spike may be coming from less than five of the total 143 events.

Unfortunately there are no events in this time frame that are longer than 12 hours with D_{st} minima lower than $-80nT$ that have existing mass density data around onset, so an analysis of this particular relationship can't be made.

4.1 Change in ρ_{eq}

If instead of looking for events based on a threshold of ρ_{eq} , we instead look for a certain amount of change in ρ_{eq} as the basis for event onset, we get Figure 7.

Both raw change and percent change were done in case of bias towards high or low ρ_{eq} periods, respec-

tively. This shows a distinct lack of correlation in either case to a change in D_{st} . The percent change in ρ_{eq} does, however, seem to be preconditioned by changes in B_z and $F_{10.7}$.

5 $F_{10.7}$ dependence

In an effort to analyze the dependence of ρ_{eq} on $F_{10.7}$, a few tests were performed. Takahashi et al. [2010] mention a strong correlation between the two. The long term correlation could be a bias for D_{st} 's effects, so a linear model was created, recreating mass density purely from $F_{10.7}$ in the form of $\rho_{eq}(t) = A * F_{10.7}(t)$. This re-created ρ_{eq} shows around a 45% correlation with the actual ρ_{eq} , suggesting a strong influence, while doing the same procedure with DST shows only a 20-25% correlation.

Taking this re-created data set and subtracting it from the original should remove the $F_{10.7}$ dependence from the data, and allow for a less biased analysis of the relationship between D_{st} and ρ_{eq} . Figure 8 shows the stack plot for the reduced $F_{10.7}$, where a more distinct peak after event onset can be seen. It also shows what that removed trend looks like, of the form $\rho_{eq} - \rho_{eq, F_{10.7}}$.

Similarly, looking at the 1 day averaged data with and without the $F_{10.7}$ dependence, Figure 9 shows that the results of Takahashi et al. [2010] hold only once the dependence has been removed. The 27 day trend couldn't be determined as not enough valid data existed to test it.

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6 Figures

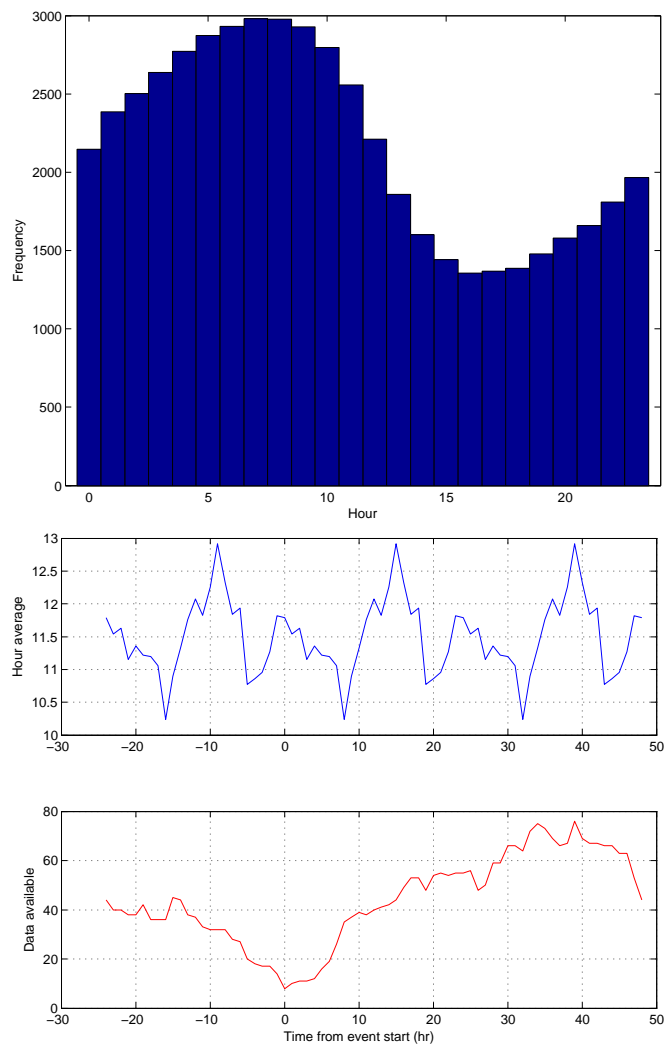


Figure 1: (a) Number of NaN points per hour of observation in the total data set (b) Data availability relative to average event hour in events where $D_{st} < -80nT$

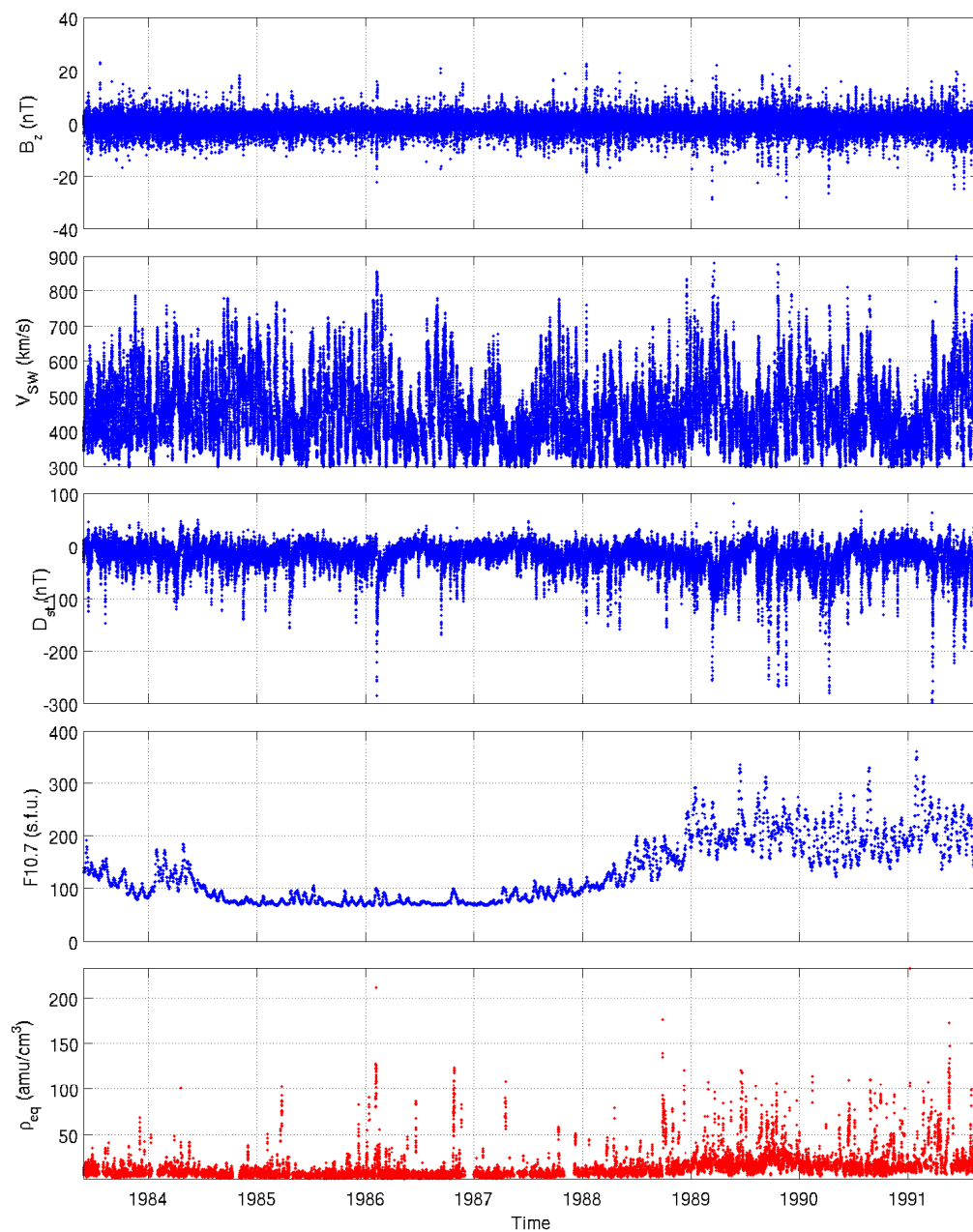


Figure 2: Overview of data used in this study. The top four panels are from Kondrashov et al. [2014] and the bottom panel is based on Denton [2007] after interpolation and averaging described in the text.

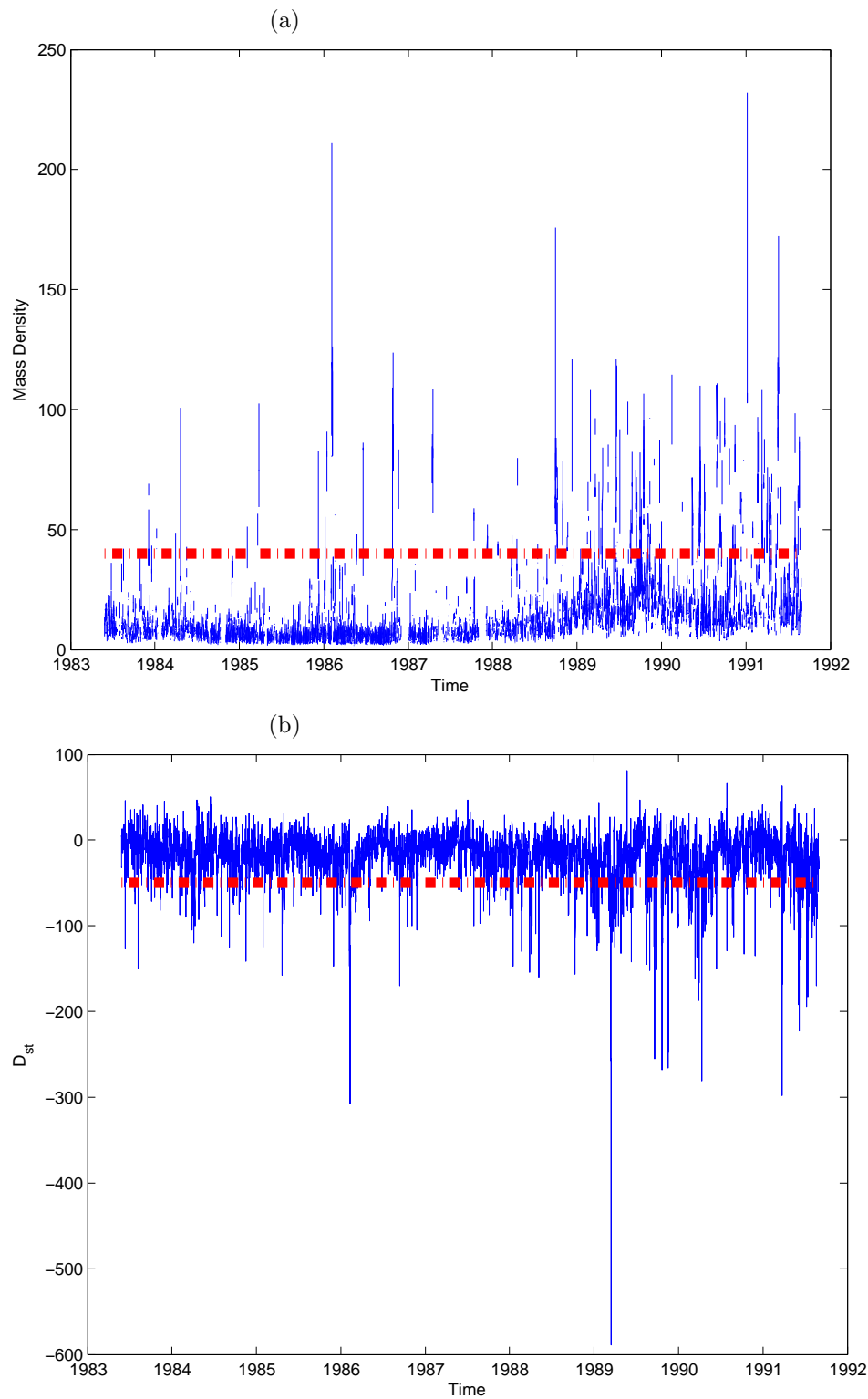


Figure 3: Averages around time of (a) 669 ρ_{eq} and (b) 130 D_{st} events along with thresholds used to define an event. The number of values used in the averages for mass density is shown in the bottom panel (a) and (b).

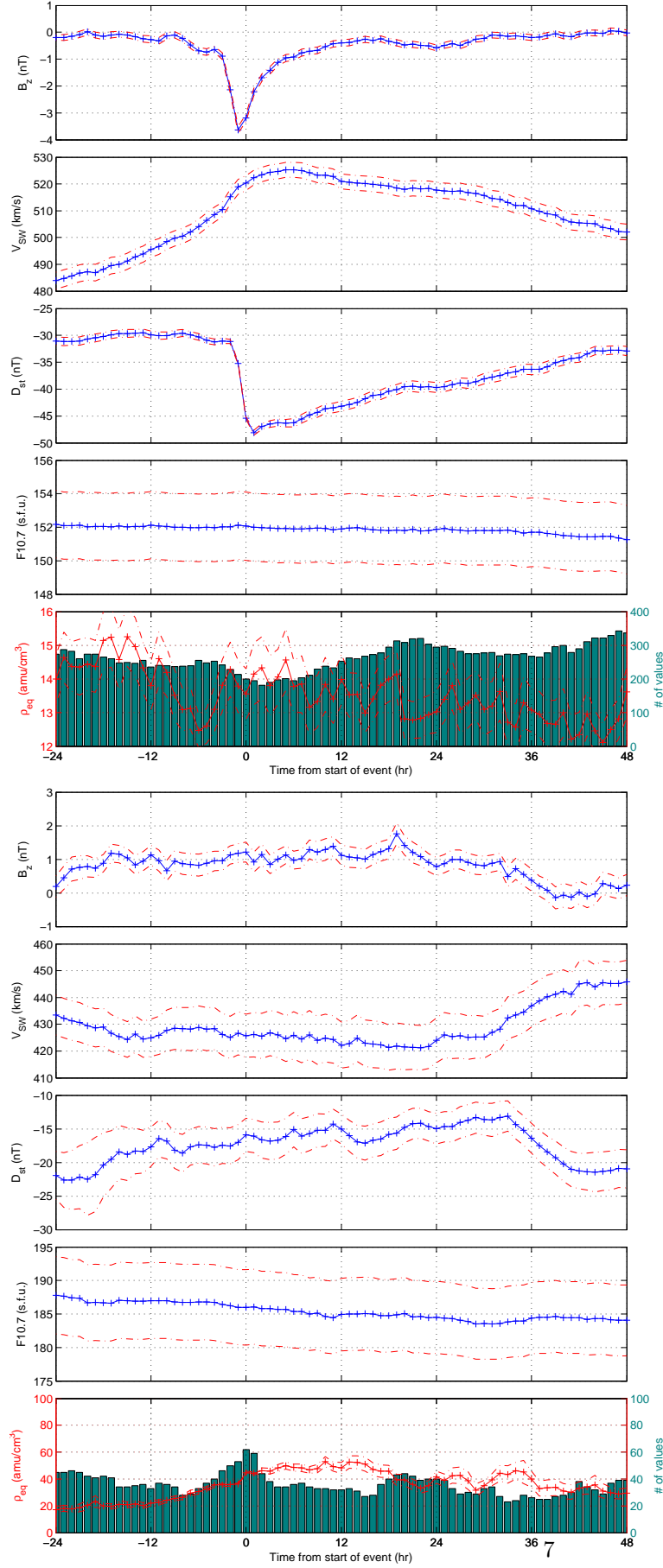


Figure 4: (a) Average of solar wind and near-Earth measurements around the time D_{st} crossed be-

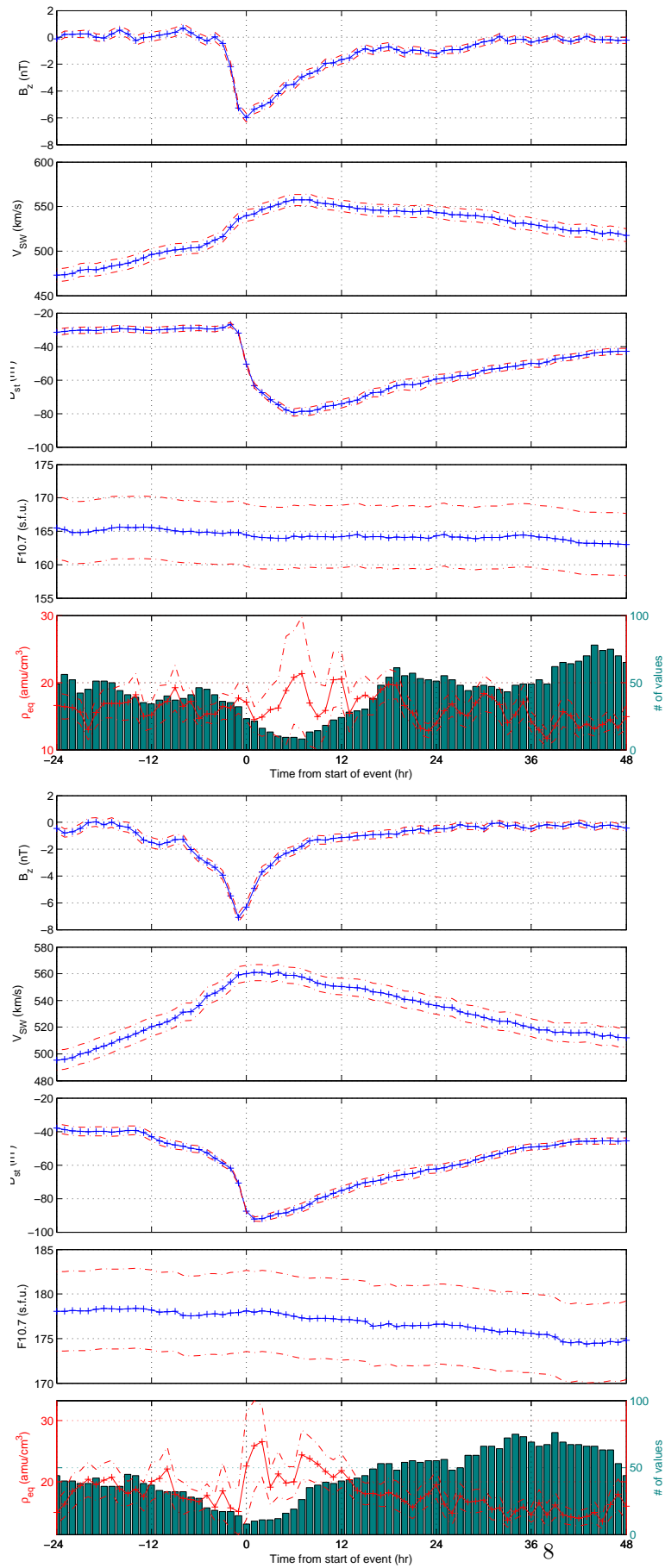


Figure 5: (a) Same as ??(a) except with constraint that D_{st} stayed below -40 nT for at least 12 hours. (b)

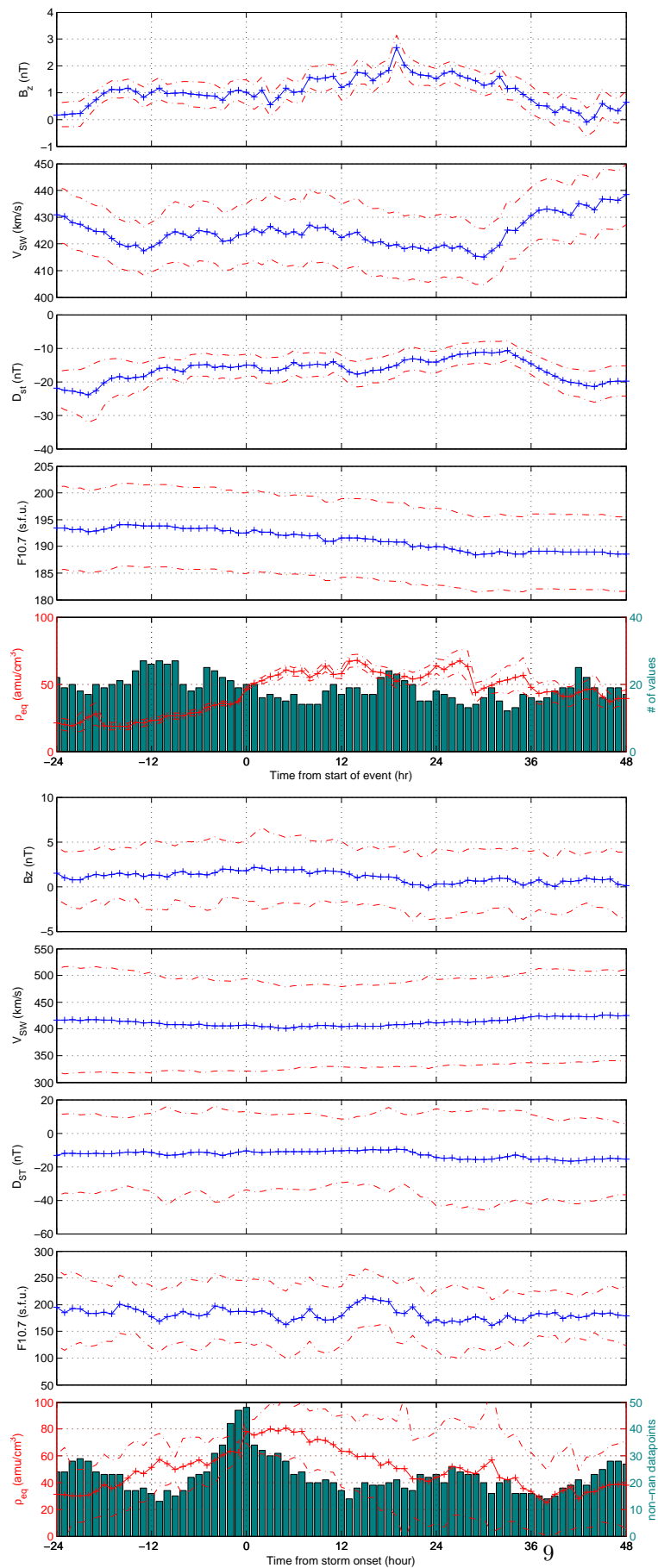


Figure 6: (a) Same as ??(a) except with constraint that ρ_{eq} stayed above 50 nT for at least 12 hours. (b)

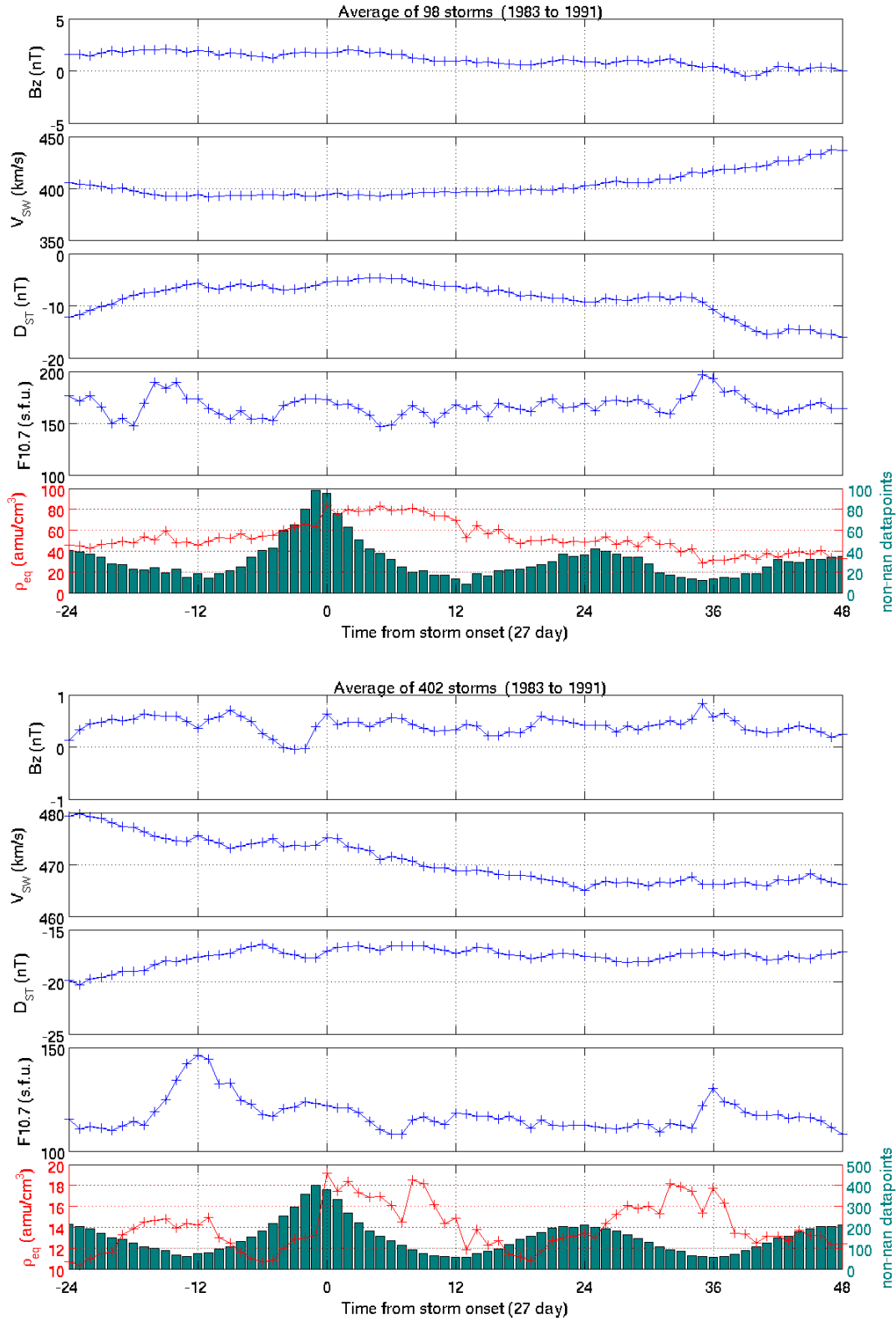


Figure 7: $\text{diff}(\rho_{eq}) > 20 \frac{\text{amu}}{\text{hour}}$, $\text{diff}(\rho_{eq}) > 30 \frac{\%}{\text{hour}}$

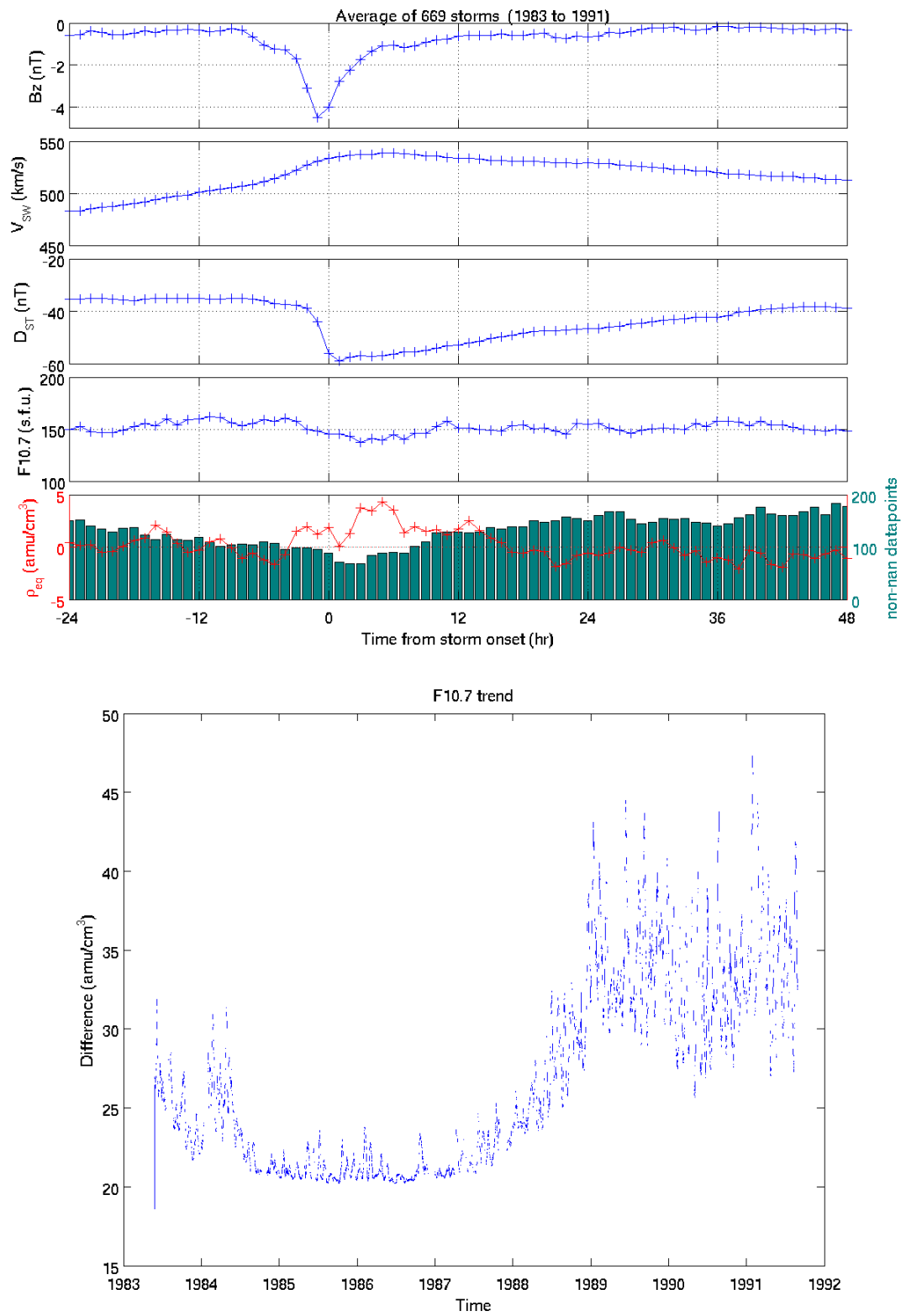


Figure 8: ρ_{eq} with $F_{10.7}$ removed, and Difference from original

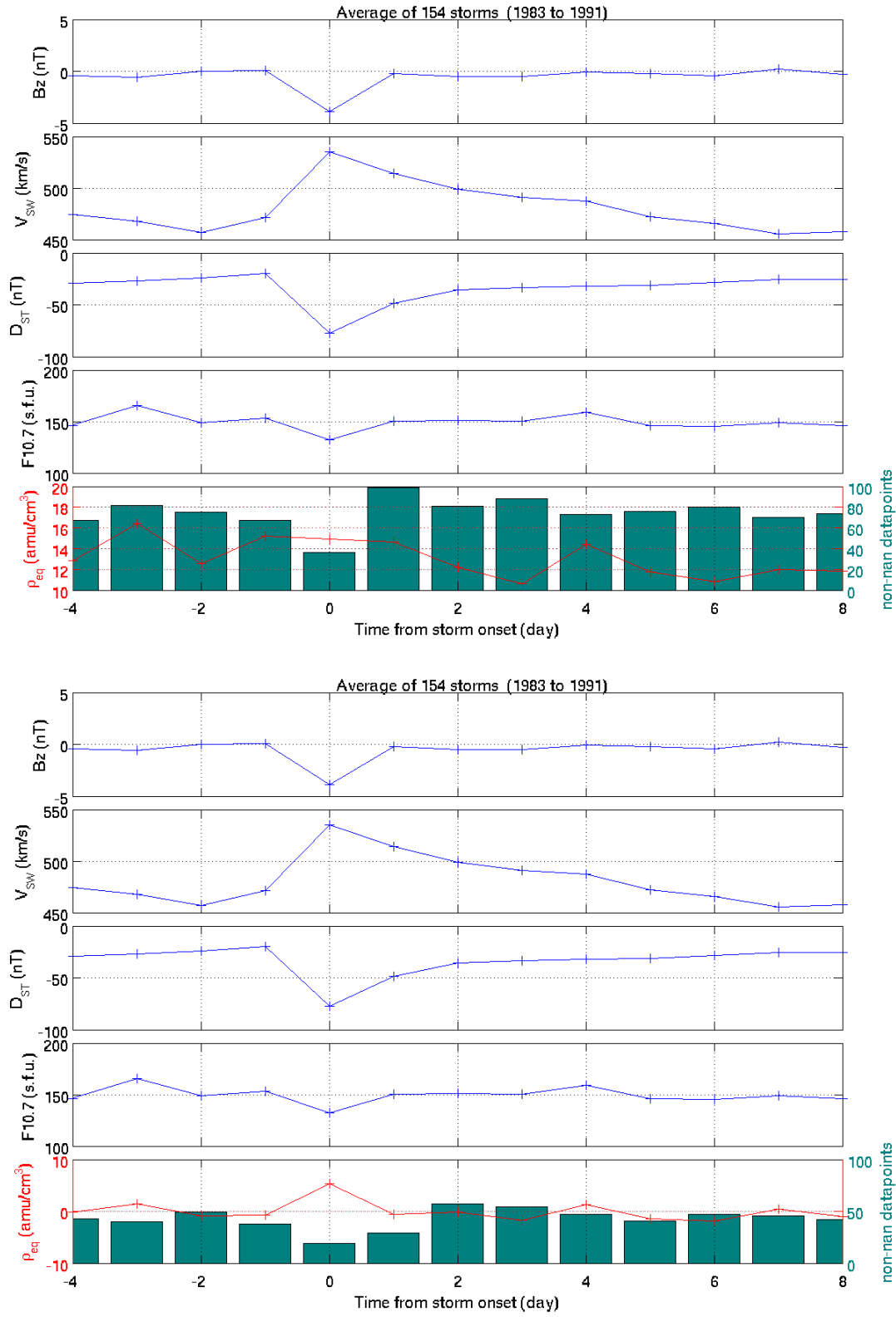


Figure 9: Original ρ_{eq} , and with $F_{10.7}$ dependence removed, both at 1-day averages