

Solar wind and geomagnetic activity influence on the equatorial plasma mass density at geosynchronous orbit

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Abstract

We consider two types of events, identified by decreases in D_{st} below a threshold value and increases in the the equatorial mass density at geosynchronous altitudes, ρ_{eq} , above a threshold value using the *Takahashi et al.* [2010] dataset. From the D_{st} events and 1-day averages, we find that there is a statistically weak and small amplitude difference between ρ_{eq} on the day of the event and the days before and after. For hourly averages, it is found that for 6 hours after onset, very few measurements of ρ_{eq} exist and that most of the contribution to the daily averages is from 6-24 hours after onset. For hourly averages, ρ_{eq} following the onset of a D_{st} event does not depend on the north-south component of the interplanetary magnetic field, B_z near onset time. The primary factor that determines if a post- D_{st} event onset peak in ρ_{eq} will occur in the hourly averages is elevated $F10.7$. From the ρ_{eq} events, we find a weak dependence on B_z after the onset of an event, with higher average B_z four hours after the event onset corresponding to larger ρ_{eq} values 24-36 hours after onset. In contrast, the average of B_z four hours prior to ρ_{eq} events does not have a statistically significant impact on ρ_{eq} after onset.

1 Introduction

The dependence of the plasmopause location on geomagnetic activity was observed as early as *Carpenter* [1966] and there are well-validated models of the plasmopause location as a function of L , MLT , and geomagnetic activity (*Lemaire* [1998], *Moldwin et al.* [2002], *O'Brien and Moldwin* [2003]). Models also exist for the plasmasphere density (*Gallagher et al.* [1988]; *Lemaire* [1998]) with similar parameterization.

The dependence of the plasma density outside of the nominal plasmopause location on geomagnetic and solar wind conditions has not been extensively studied until the past decade (*Takahashi et al.* [2006]; *Takahashi et al.* [2010]; *Denton et al.* [2016]). Earlier works that considered profiles of ion density found a significant dependence of the plasmopause location on geomagnetic activity (*Chappell et al.* [1970]; *Maynard and Grebowsky* [1977]; *Carpenter and Anderson* [1992]), but the relationship between geomagnetic activity and density outside of the plasmopause was not explicitly considered.

Takahashi et al. [2006] estimated magnetospheric mass density using measurements from the CRRES satellite during a 73-day period in 1991 when CRRES was in an elliptical, near equatorial orbit and primarily beyond the nominal plasmopause. They found that the local average ion mass, M , had some correlation with geomagnetic activity, with more negative hourly D_{st} values corresponding to higher M and 1.5- and 3-day averages of K_p corresponding to higher M (3-hour averages of K_p had little visual correlation with M). The mass density estimates were obtained from CRRES observations in the MLT range of [12:00, 18:00] with most observations at L between 5 and 7 R_E .

Denton et al. [2006] found a weak relationship between D_{st} and K_p and the field line distribution of plasma mass density, ρ , using observations of the first three toroidal Alfvén harmonic frequencies for L in the range of 6-8 R_E . A more pronounced peak in the distribution along a field line near the equator was found for lower D_{st} and larger K_p , but the equatorial value of ρ was found to be similar for D_{st} bins of [-142, -31] and [-31, 37] nT and K_p bins of [1.5, 3.4] and [3.4, 5.9].

Takahashi et al. [2010] developed a mass density dataset using measurements from the Space Environment Monitor instruments on the Geostationary Operational Environmental Satellites (GOES) satellites from 1980 through 1992, with most measurements in the range of $L = 6.8 \pm 0.2R_E$. The mass density was estimated using the the Alfvén wave velocity relationship, $V_A = B/\sqrt{\mu_0\rho}$, a magnetic field model, and a numerical solution to a wave equation with an ionospheric boundary and the assumption of a zero resistance ionosphere. The equatorial mass density, ρ_{eq} , was derived from the estimated mass density using a power law dependence on the geocentric distance to the field line of the observation, R , $\rho = \rho_{eq}(LR_E/R)^{1/2}$. They found a high correlation (~ 0.94) between 27-day averages of $F10.7$ and 27-day medians of ρ_{eq} .

Takahashi et al. [2010] noted that downward drops in the D_{st} index coincided with significant changes in ρ_{eq} for L near $6.8 R_E$. For five storms, two had ρ_{eq} spikes after D_{st} minimum, two had ρ_{eq} spikes before the drop, and one showed little change in ρ_{eq} . A key result was that when daily-averaged measurements were considered, the epoch average of D_{st} exhibited for storms with a minimum $D_{st} < -50$ nT showed an enhancement in ρ_{eq} the same day as minimum D_{st} in the epoch averages.

Yao et al. [2008] studied the relationship between D_{st} and the number density of low-energy O^+ in different regions (ring current and plasma sheet) using the TEAMS and ESA instruments on the low-altitude and high-inclination orbiting FAST satellite and found that the average N_{O^+} across the sampled L-shells (2-14) had a strong correlation (0.88) with the minimum D_{st} of the storm. Although the correlation at geosynchronous distances was not calculated, the data presented for four events show that the enhancements in N_{O^+} tended to appear near D_{st} minimum (± 2 hours).

The above results indicate that (1) the mass density in the at geosynchronous distances is best correlated with $F10.7$ and (2) there exists a much weaker statistical relationship between mass density and the geomagnetic activity indices K_p and D_{st} . In this work we consider the dependence of mass density estimates in the *Takahashi et al.* [2010] data set

and attempt to identify and statistically characterize the relationship between geomagnetic activity with mass density at geosynchronous distances. In addition, we attempt to identify solar wind processes that may drive the changes.

2 Data Preparation and Overview

The parameters ρ_{eq} and $F_{10.7}$ are from the dataset of *Takahashi et al.* [2010]; data are available from 1980 through 1991 from GOES 2, 3, 5, 6 and 7. All other parameters used are from *Kondrashov et al.* [2014], which has coverage from 1972 through 2013 and are on a 1-hour time grid. ρ_{eq} is the inferred equatorial mass density based on the 3rd harmonic torodial frequency of magnetic field measurements as described in *Takahashi et al.* [2010]. The smallest cadence for ρ_{eq} values is 10 minutes. To compute an hourly median over the same interval for which the solar wind parameters were averaged, the median of all ρ_{eq} values in a given hour window was used. GOES 6 had the longest span of available data and the primary results presented are for GOES 6 while measurements from other satellites were used for consistency checks.

The top panel of Figure 2 shows the long-term trends of $\log_{10}(\rho_{eq})$ from GOES 6 and $F_{10.7}$ computed using the median values in 27-day non-overlapping windows. A scatter plot of these two lines is shown in the lower panel of Figure 2 along with that for 1-day and 1-hour medians. The linear correlation for GOES 6 using all measurements is found to be 0.94, which is the same value documented by *Takahashi et al.* [2010] who used combined measurements from all satellites in the time interval of 1980 through 1991 using the constraint $0600 \leq MLT \leq 1200$, $K_{p3d} \geq 1.0$ and $D_{st} \geq -50$ nT. Our 27-day correlations for GOES 2, 5, and 7 are 0.81, 0.78, and 0.87, and the time range of coverage of measurements from these satellites is 1980-01-01/1983-05-16, 1983-01-01/1987-02-28, and 1983-05-27/1991-08-29, respectively.

3 Results

Two types of events are considered. The first is a drop in the D_{st} index below the threshold of -50 nT. These types of events are considered in order to determine if there is a well-defined ρ_{eq} dependence on geomagnetic activity/storms as indicated by the D_{st} index. It is expected that the plasmasphere will exhibit two responses near these threshold crossings. First is the movement of the plasmopause earthward (*Lemaire* [1998]) and the second is a possible change in density due to magnetospheric processes.

The second type of event is an increase in ρ_{eq} above a threshold of 20 amu/cm³. It is expected that these increases will typically occur after period of quiet geomagnetic activity but may also be the result of increases in solar wind driven geomagnetic activity. In this case, the plasmopause boundary may move outward past geosynchronous orbit and the increase will be due to measurements being made inside the plasmopause. A second possible reason for the increase could be due to magnetospheric processes.

3.1 D_{st} Events

Our first analysis uses GOES-6 data only from the time interval 1989-1991, which corresponds with the interval used in the epoch analysis of *Takahashi et al.* [2010]. In this interval, there were 75 D_{st} events where D_{st} remained below the -50 nT threshold for at least 12 hours. Only events with an onset time between 06:00 and 12:00 MLT were considered. The first hour of a local minimum in D_{st} following a threshold crossing defines the zero epoch time, and for each event, 4·24 hours were considered before and 8·24 hours after event onset. For multiple detected threshold crossings within 24 hours of each other, the first crossing was selected as the event because visual inspection showed the following crossings to typically be from brief downward deviations during the recovery period of a geomagnetic storm. To compute the epoch averages on a daily time scale, the median value of all available measurements for all events centered on a window of ± 12 hours of the epoch zero hour was computed, and these

averaging windows were shifted in increments of 24 hours. Similar results are obtained if we first reduce each event time series to have a 1-day cadence by computing medians for each event in 1-day bins and then compute the medians across events on each epoch day.

The stack plot in the upper panel of Figure 3 shows the epoch averages for the 75 events from 1989-1991. The minimum D_{st} median is -75 nT. Consistent with *Takahashi et al.* [2010], D_{st} events correspond to elevated ρ_{eq} before and after the event, although the magnitude of increase observed here is 4.5 amu/cm³ instead of the increase of ~ 10 amu/cm³ found in Figure 11 of *Takahashi et al.* [2010]. The vertical green bars in the ρ_{eq} plot show the number of ρ_{eq} values that were used to compute the medians and red error lines. Error lines for each parameter are the standard deviation of the values used in computing the median divided by the square root of the number of values. For the non- ρ_{eq} parameters, the number of measurements used in computing the medians is typically equal to the number of events, which is much larger than the number of ρ_{eq} values as all available measurements were used instead of restricting to only values where a ρ_{eq} value also existed.

In the lower panel of Figure 3, it is shown that the trends seen in the analysis of 1989-1991 also hold for the entire span of measurements from GOES 6 - a small elevation on the day of the event relative to the day before and day after, with median averages after the day of the event being slightly lower than that before.

To determine if the magnitude of the density medians in the top panel of Figure 3 on epoch days -1, 0, and 1 are statistically different from each other, we used a two-population bootstrap test. A bootstrap sample of the medians for each epoch day was created by sampling the values used to compute the median on each epoch day with replacement. The observed difference between the medians is shown in Table 1 along with the fraction of observations in 1000 bootstrap differences with a different sign than that observed.

If a statistically significant difference in the medians between two days existed, the fraction of bootstrap samples with a different sign is expected to be much smaller ($< 5\%$). In terms of a hypothesis test, we cannot reject the null hypothesis that the medians are the

Days	Difference (amu/cm ³)	%
-1 0	-4.48	13.90% ≥ 0
-1 1	-1.46	34.40% ≥ 0
1 0	-3.02	20.30% ≥ 0

Table 1: Results of test on of means of ρ_{eq} shown in the top panel of Figure 3 between days of threshold crossing near (day = 1 or -1) or on the day of a D_{st} event (day = 0).

same with a confidence higher than 100 minus the percentage value shown in the table; the highest confidence level of 86% is in the difference between day 0 and day -1.

A similar hypothesis test run for the bottom panel of Figure 3 yields a maximum confidence level of $\sim 85\%$.

These tests indicate that there is not a statistically significant difference in medians of daily-averaged ρ_{eq} on the day before and day after a threshold crossing in D_{st} . In addition, the observed differences are on the order of only 20% of the value of ρ_{eq} on the onset day, which is similar to the variation in the observations over the displayed epoch time interval.

We have also considered the hourly averages from the events shown in the bottom panel of Figure 3, which shown in the top panel of Figure 4. In the time period of elevated mass density within 6 hours of the start of the event, there are only ~ 4 samples per hourly interval and the peak value of ~ 70 amu/cm³ is associated with a single observation.

To increase the number of events, we have removed the constraint that the events must occur in the MLT range from 06:00-12:00 and the constraint that D_{st} remained below the threshold value for 12 hours or longer. The result is shown in the bottom panel of Figure 4. The two spikes (one upwards after onset, one downwards before onset) were tested for significant deviation from two baselines, one 24-12 hours before onset, one 16-48 hours after onset. The probability of the left spike coming from the same distribution as the left baseline is 1.06%, and 8.2% from the right baseline, whereas the probability of the right spike coming from the left baseline is 1.89% and 0.17% from the right baseline. This indicates that there is a statistically significant peak in ρ_{eq} within 6 hours of onset when hourly averages are considered. This result is consistent with previous observations of large mass density increases

in the plasma trough region during large geomagnetic storms (*Yao et al.* [2008]; *Takahashi et al.* [2010]).

To determine if there is an IMF B_z dependence, the epoch curves in the bottom panel of Figure 4 for ρ_{eq} and B_z were separated into two parts depending whether B_z was above or below its average four hours before and including the onset hour and four hours after and including the hour of onset; the result is shown in Figure 5. Any statistically significant difference at a 95% confidence level between the two parts is indicated by a green dot on the horizontal axis using the same bootstrap method described previously. (At this confidence level, false positives are expected to randomly occur 5% of the time, or 3.6 times for 72 values.) These two plots indicate that there is not a statistically significant dependence on IMF B_z around the time of the onset of a D_{st} event.

The process for creating Figure 5 was repeated except using the sort variable $F10.7$ instead of B_z , and the result is shown in Figure 6. The difference between the epoch curves is statistically significant at all times, as expected due to the high correlation between ρ_{eq} and $F10.7$, but elevated $F10.7$ is associated with a peak in the ρ_{eq} curve approximately 5 hours after onset. The values contributing to the peak were compared to a baseline of all values 16-48 hours after onset via a Wilcoxon rank-sum test and were found to have significantly different medians at the 99% significance level. This indicates the possibility that there must be a high base-line level of ρ_{eq} in order for an increase in ρ_{eq} to be observed for a D_{st} event.

3.2 ρ_{eq} Events

An alternative type of event can be defined that corresponds to large increases in ρ_{eq} . Figure 7 shows the epoch time series for events in which an increase of ρ_{eq} above 20 amu/cm² defines the start of the event. In finding ρ_{eq} events, missing values were replaced with linearly interpolated values. These events are associated with positive B_z and low values of D_{st} . Approximately 12 hours after onset, ρ_{eq} slowly decays over 40 hours to near pre-onset levels.

Figure 7 can be compared with Figure 9 of *Denton et al.* [2016], which shows daily

averages of ρ_{eq} from this dataset after the onset of quiet K_p intervals. The growth shown here is approximately linear and lasts ~ 12 –18 hours whereas the 10 quiet K_p events considered by *Denton et al.* [2016] showed growth that lasted for at least 48 hours.

Figure ?? shows the the events of Figure 7 when separated by B_z in the same manner as Figure 5. The largest difference occurs when the separation is performed after onset, with more positive B_z leading to a larger peak in ρ_{eq} than more negative B_z . The rate of change of ρ_{eq} also appears to be insensitive to B_z for approximately 12 hours prior to onset.

Figure 9 shows how the epoch average curve of ρ_{eq} in Figure 8 depends on $F10.7$. Statistically significant differences occur independent of if the separating factor the average of $F10.7$ before or after onset, as expected because the cadence of $F10.7$ is one day and its time scale of variation is longer than that of ρ_{eq} . Events with high $F10.7$ are associated with higher ρ_{eq} before and after event onset but similar peak values. As was the case for B_z , there is an approximately 12-hour time span of linear growth of ρ_{eq} near the time of onset but because the initial starting ρ_{eq} is higher for high $F10.7$, the associated growth rate is lower.

3.3 Summary and Conclusions

In recent years, various works have considered the influence of solar wind control and the geomagnetic activity relationship to the equatorial plasma mass density at geosynchronous orbit, generally using event studies and linear correlations (*Takahashi et al.* [2006]; *Denton et al.* [2006]; *Yao et al.* [2008]; *Takahashi et al.* [2010]; *Denton et al.* [2016]). The event studies show a possible relationship between ρ_{eq} and solar wind and geomagnetic activity indicators such as K_p and D_{st} while correlation studies have generally found a very weak relationship.

This work presents an analysis of two types of events intended to provide insight into the relationship between ρ_{eq} and solar wind and geomagnetic parameters by considering both large geomagnetic events and large ρ_{eq} events. The analysis attempts to address a major

difficulty with the identification of the cause in ρ_{eq} enhancements, which can occur for two reasons: (1) an extended period of positive B_z , which leads to low geomagnetic activity, plasmasphere refilling, and a possible expansion of the plasmasphere to geosynchronous altitudes and (2) strong negative B_z , which (a) leads to the inward motion of the higher density plasmasphere and leaves geosynchronous altitudes in the lower density plasma trough just outside of the plasmopause and (b) causes possible magnetospheric processes to enhance the mass density at geosynchronous altitudes.

On average, around the time of large geomagnetic events, there is a statistically weak enhancement of daily-averaged ρ_{eq} , but consideration of hourly average ρ_{eq} shows enhancements over a 6-hour window that contains few events. For the hourly-averaged measurements, the average of B_z around onset time did not have a significant impact on ρ_{eq} after onset.

It was shown that a key factor in determining if large ρ_{eq} enhancements will be observed is elevated $F10.7$. When geomagnetic events were separated according to if they occurred when $F10.7$ was low or high, significant enhancements in ρ_{eq} were only observed under high $F10.7$ conditions.

For ρ_{eq} , more positive B_z after onset lead to higher values of subsequent ρ_{eq} , which indicates that a substantial portion of the events may be due to the apparent mass refilling mechanism described by *Denton et al.* [2016]. Without additional information about the location of the plasmopause, it may not be possible to separate out the conditions required for enhancement due to enhanced magnetospheric activity from this mechanism.

References

- Carpenter, D. L., Whistler studies of the plasmopause in the magnetosphere: 1. Temporal variations in the position of the knee and some evidence on plasma motions near the knee, *J. Geophys. Res.*, *71*, 693–709, 1966.
- Carpenter, D. L., and R. R. Anderson, An ISEE/whistler model of equatorial electron density in the magnetosphere, *J. Geophys. Res.*, *97*, 1097–1108, 1992.
- Chappell, C. R., K. K. Harris, and G. W. Sharp, A study of the influence of magnetic activity on the location of the plasmopause as measured by OGO 5, *J. Geophys. Res.*, *75*, 50–56, 1970.
- Denton, R., Database of Input Parameters for Tsyganenko Magnetic Field Models, 2007.
- Denton, R. E., K. Takahashi, I. A. Galkin, P. A. Nsumei, X. Huang, B. W. Reinisch, R. R. Anderson, M. K. Sleeper, and W. J. Hughes, Distribution of density along magnetospheric field lines, *J. of Geophys. Res.*, *111*, 4213, 2006.
- Denton, R. E., K. Takahashi, J. Amoh, and H. J. Singer, Mass density at geostationary orbit and apparent mass refilling, *J. Geophys. Res.*, *121*, 2962–2975, 2016.
- Gallagher, D. L., P. D. Craven, and R. H. Comfort, An empirical model of the earth’s plasmasphere, *Advances in Space Research*, *8*, 15–21, 1988.
- Kondrashov, D., R. Denton, Y. Y. Shprits, and H. J. Singer, Reconstruction of gaps in the past history of solar wind parameters, *Geophys. Res. Lett.*, *41*, 2702–2707, 2014.
- Lemaire, J., *The Earth’s Plasmasphere*, Cambridge University Press, 1998.
- Maynard, N. C., and J. M. Grebowsky, The plasmopause revisited, *J. Geophys. Res.*, *82*, 1591–1600, 1977.

- 267 Moldwin, M. B., L. Downward, H. K. Rassoul, R. Amin, and R. R. Anderson, A new model
268 of the location of the plasmopause: CRRES results, *J. of Geophys. Res.*, *107*, 1339, 2002.
- 269 O'Brien, T. P., and M. B. Moldwin, Empirical plasmopause models from magnetic indices,
270 *Geophys. Res. Lett.*, *30*, 1–1, 2003.
- 271 Takahashi, K., R. E. Denton, R. R. Anderson, and W. J. Hughes, Mass density inferred from
272 toroidal wave frequencies and its comparison to electron density, *J. of Geophys. Res.*, *111*,
273 1201, 2006.
- 274 Takahashi, K., R. E. Denton, and H. J. Singer, Solar cycle variation of geosynchronous
275 plasma mass density derived from the frequency of standing Alfvén waves, *J. of Geophys.*
276 *Res.*, *115*, 7207, 2010.
- 277 Yao, Y., K. Seki, Y. Miyoshi, J. P. McFadden, E. J. Lund, and C. W. Carlson, Effect of solar
278 wind variation on low-energy O^+ populations in the magnetosphere during geomagnetic
279 storms: FAST observations, *J. of Geophys. Res.*, *113*, 2008.

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4 Figures

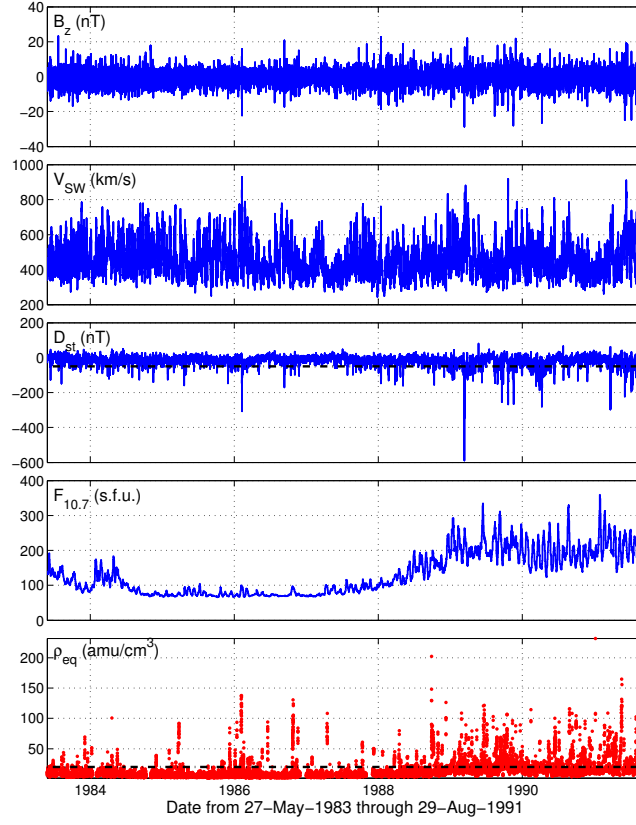


Figure 1: Overview of data used in this article. The top four panels show parameters from *Kondrashov et al.* [2014] and the bottom panel contains ρ_{eq} based on GOES 6 measurements from *Denton* [2007] after interpolation and averaging described in the text. Dashed horizontal lines in the D_{st} and ρ_{eq} panels indicate sample event cutoff thresholds of $D_{st} = -50$ nT and $\rho_{eq} = 20$ amu/cm³ considered in Section 3.

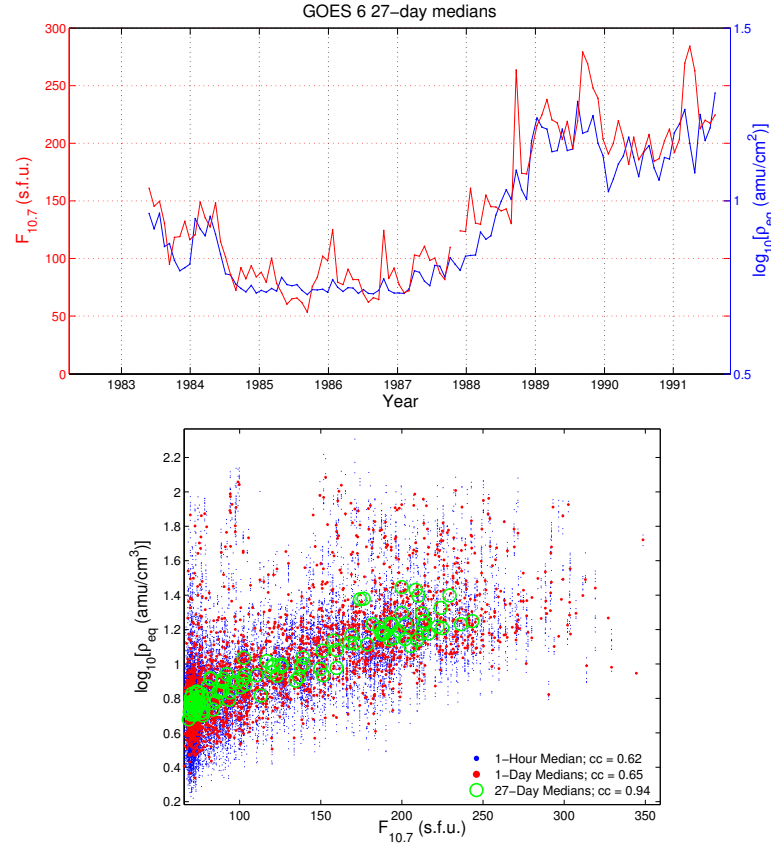


Figure 2: Top: 27-day non-overlapping medians of $F_{10.7}$ and $\log(\rho_{eq})$ from GOES 6. Bottom: Correlation between $\log(\rho_{eq})$ and $F_{10.7}$ using medians in non-overlapping hour, day, and 27-day windows.

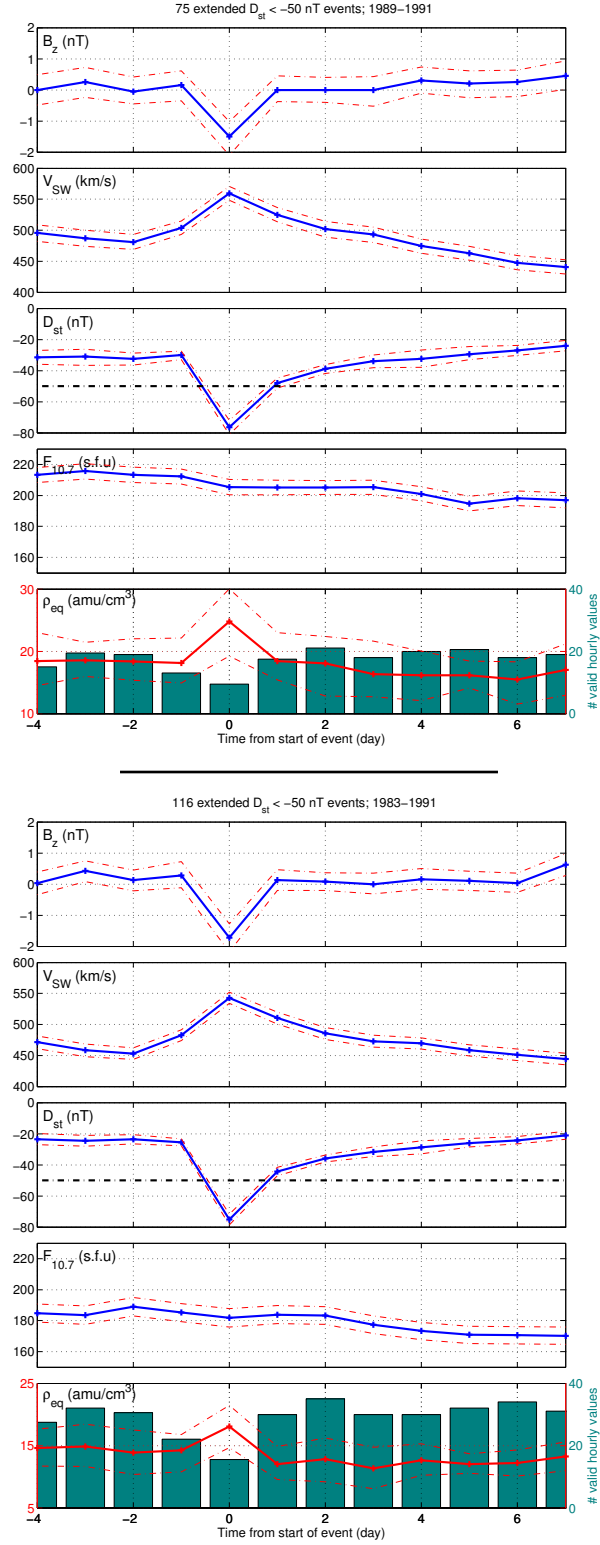
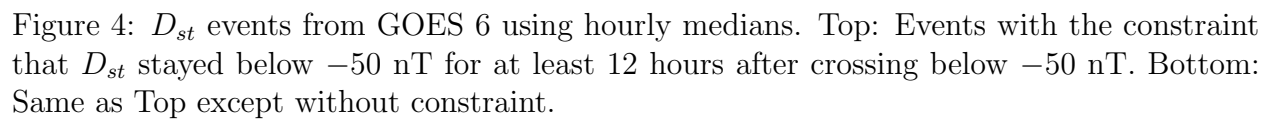


Figure 3: D_{st} events from GOES 6 using daily medians. Top: Events in the interval 1989-1991; compare to *Takahashi et al.* [2010] Figure 11. Bottom: Events in the interval 1983-1991.



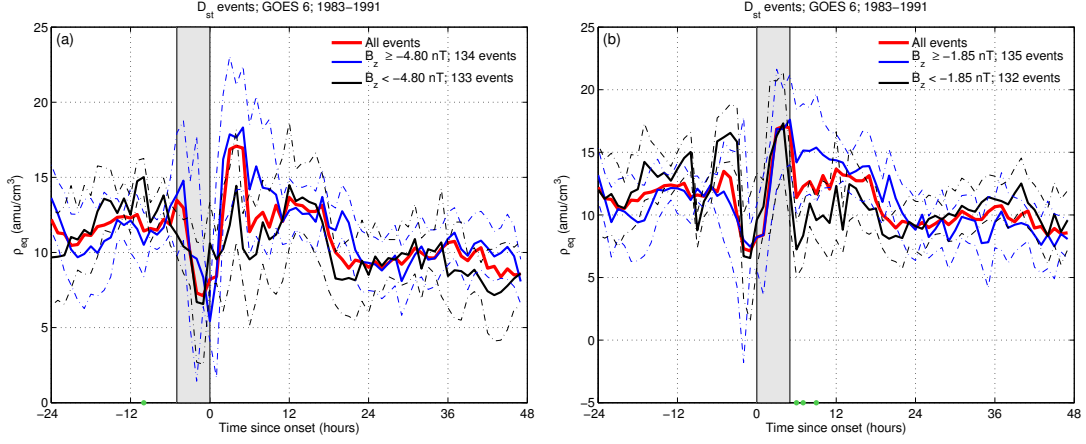


Figure 5: D_{st} events of Figure 4 separated by the average D_{st} value (a) at onset and four hours before, and (b) at onset and four hours after.

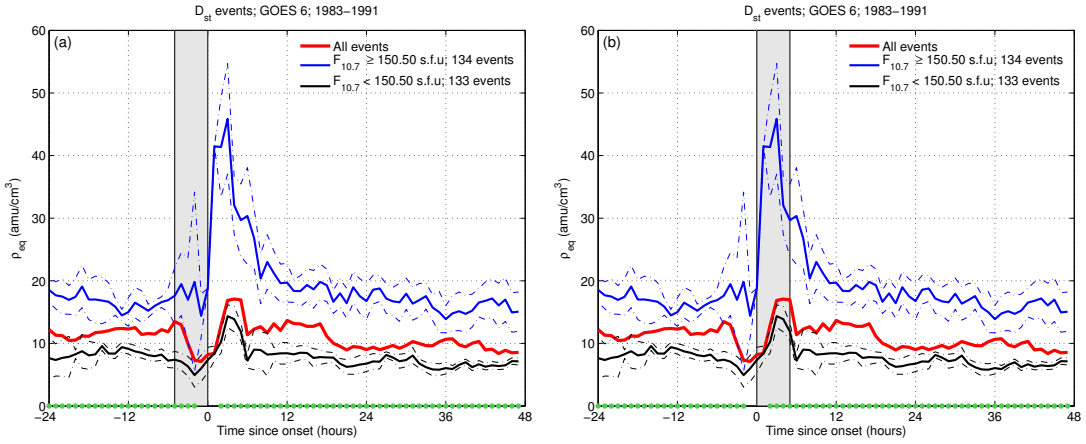


Figure 6: D_{st} events of Figure 4 separated by the average $F10.7$ value (a) at onset and four hours before, and (b) at onset and four hours after.

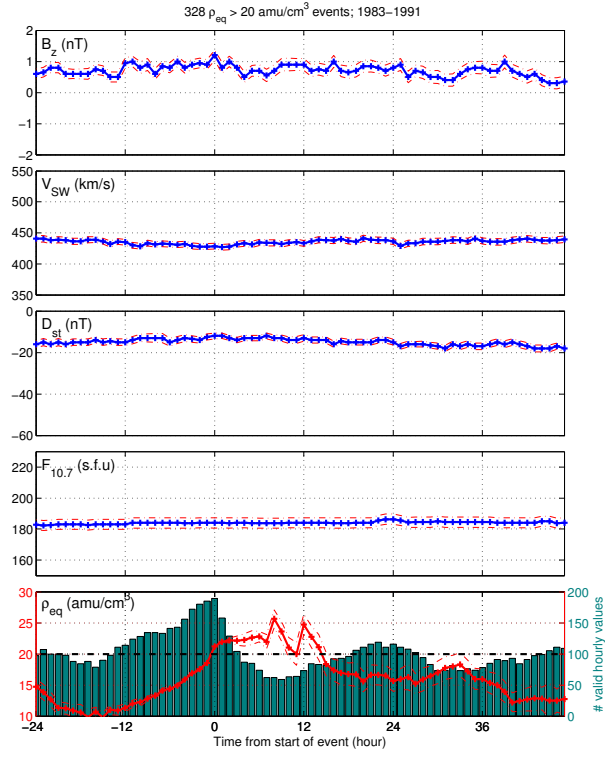


Figure 7: $\rho_{eq} > 20 \text{ amu/cm}^3$ events from GOES 6 using hourly medians.

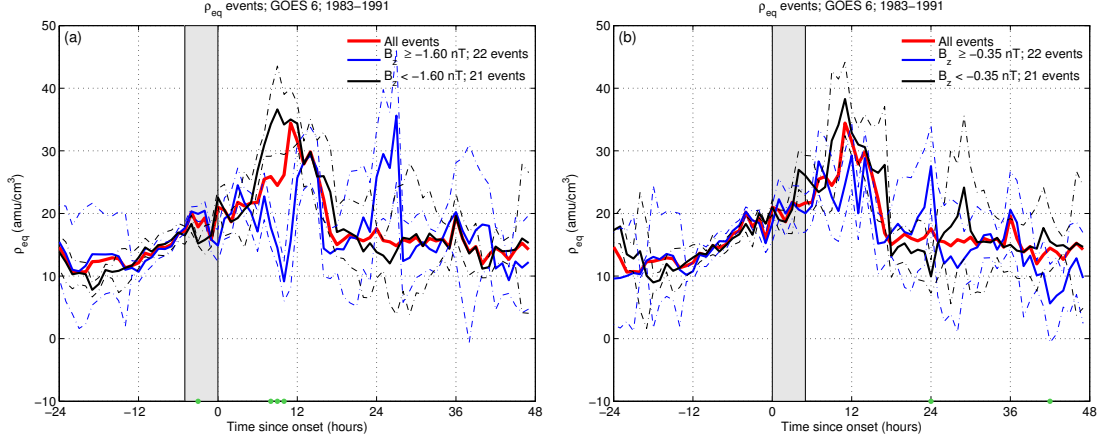


Figure 8: ρ_{eq} events of Figure 7 separated by the average B_z value (a) at onset and four hours before, and (b) at onset and four hours after.

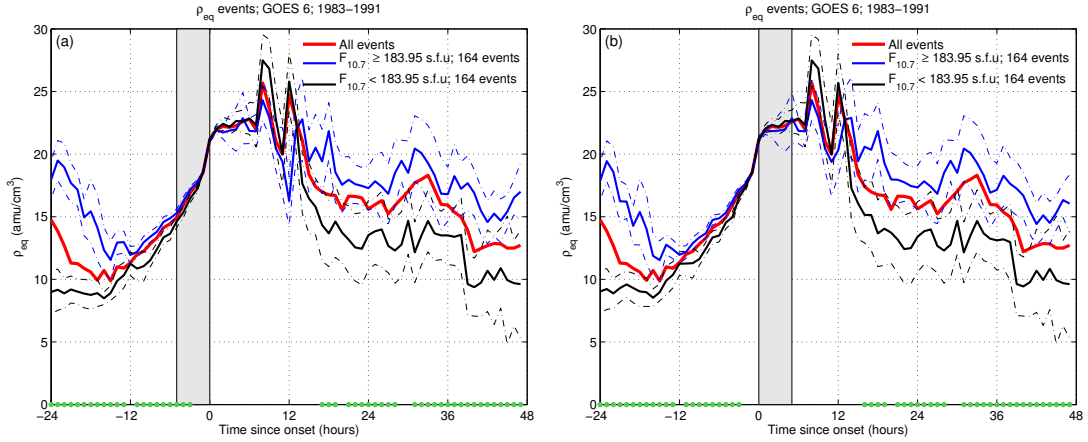


Figure 9: ρ_{eq} events of Figure 7 separated by the average $F10.7$ value (a) at onset and four hours before, and (b) at onset and four hours after.