

White Dwarfs as Dark Matter Detectors

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White dwarfs can serve as detectors for ultra-heavy dark matter states which interact to trigger type Ia supernovae. This was originally proposed in [?] and used to place bounds on primordial black holes. In this paper we extend the capability of white dwarf detectors to candidates with non-gravitational couplings, focusing on dark matter transits and collisions within the white dwarf. In particular, we provide a detailed analysis of the explosiveness for any heating mechanism in the white dwarf which releases high-energy standard model particles. We apply this mechanism to constrain Q-ball dark matter and model of dark nuclei in regions of parameter space fundamentally inaccessible to terrestrial-based experiments.

PACS numbers:

I. INTRODUCTION

The detection of ultra-heavy dark matter (DM) is an open problem which will ultimately require a confluence of astrophysical probes. For instance, DM masses above $\sim 10^{22}$ GeV will register fewer than 1 event per year in a typical terrestrial detector of size $\sim (100 \text{ m})^2$. Furthermore, the lack of conclusive signatures on a variety of experimental fronts has led many to seriously consider DM candidates far above the weak scale and their potential signatures [cite people](#). One possibility proposed by [?] is that ultra-heavy DM can trigger supernovae in sub-Chandrasekhar white dwarf (WD) stars by inducing runaway fusion. In this regard, white dwarfs can serve as detectors for ultra-heavy DM states.

White dwarfs are particularly suited to this task, as they are more susceptible to runaway fusion than are main-sequence stars. Runaway fusion requires that two criteria be met: a region within the star must be hot enough to support exothermic fusion reactions, and the rate at which energy is released by these reactions must dominate any cooling mechanisms that drain energy from the fusing region. The stellar medium of a WD has fewer cooling mechanisms available than does a non-degenerate star - WD cooling relies on thermal diffusion whereas main-sequence stars can also cool via thermal expansion. This is suppressed in a WD since its pressure is determined by electron degeneracy and is thus independent of temperature. The remaining primary cooling mechanism, diffusion, becomes less important over longer length scales and can be overcome by sufficiently heating a large enough region of the WD.

The necessary trigger for runaway fusion was initially computed in [?] and subsequently implemented in [?] to constrain primordial black holes, which can ignite WD stars via gravitational dynamical friction. In addition, the authors of [?] identify several other heating mechanisms involving DM which may be constrained in a similar manner. [Some of these have been explored in follow-up works...citations, citations.](#) In this work, we extend their analysis to DM candidates with generic non-gravitational couplings, focusing on DM transits through

the WD and DM-DM collisions within the WD which release energy in the form of standard model (SM) particles. More generally, we provide a detailed analysis of the explosive power and resulting constraints on any DM with interactions of this sort.

Concrete examples of DM candidates with interactions of this type include baryonic Q-balls found in supersymmetric extensions of the SM and [dark nuclei with high-order couplings to the SM](#). We are able to constrain these models in regions of parameter space fundamentally inaccessible to terrestrial experiments. However, it is important to note that any such DM constraints are by nature complimentary to terrestrial ones - it is more massive DM that is likely to trigger supernovae, and also more massive DM that have sufficiently low flux on Earth. What allows the WD to be effective in this regime is its enhanced surface area $\sim (4000 \text{ km})^2$ and exceptionally long lifetime $\sim \text{Gyr}$. In this sense, the WD is a large “space-time volume” detector. [We also have a “flux is too low” limit...should we include here?](#)

[Other ignition sources exist \(i.e., detector backgrounds\). How does this limit our constraints? Discuss the specific tests we apply \(WD lifetime, SN rate\) to derive our constraints. Possible long term addition: improve constraints by using non-type Ia sn rate instead of total rate. Should talk with SN people first.](#)

We begin in Section II by reviewing the explosion mechanism in a WD. In Section III, we parametrize the properties of DM necessary to trigger explosions through non-gravitational interactions in the case of a DM transit or collision in the star. The precise explosive power will be determined by a heating length, which is computed in Section IV for different possible interactions. Following this general discussion, we apply the WD detector to place constraints on ultra-heavy dark matter in Section V, including constraints on Q-balls as a concrete example. We conclude in Section VI.

II. WHITE DWARF RUNAWAY FUSION

In a WD, the *fusion temperature* T_f is a constant set by the energy required for ions to overcome their mutual Coulomb barriers $T_f \sim \text{MeV}$. WD cooling is set by the thermal diffusivity of photons and degenerate electrons. The two species dominate the cooling at different stellar densities, as determined precisely in [?]. Note that for a heated region of size R , this cooling rate scales as R^{-2} while the fusion rate scales as R^{-3} . Thus there is a critical *trigger size* λ_T below which diffusive cooling dominates the thermal evolution of a temperature peak, and above which the liberated fusion energy dominates [?]. Therefore, a region with temperature greater than T_f and size greater than λ_T in a given WD will launch a runaway fusion chain-reaction and result in a type Ia supernova. The numerical value of λ_T is highly sensitive to the WD density and has been analytically scaled for varying WD masses in [?]. As in [?], we restrict our attention to carbon-oxygen WDs in the upper mass range $0.7 - 1.4 M_\odot$ which correspond to a number density of nuclei $n_i \sim 10^{29} - 10^{32} \text{ cm}^{-3}$. Over this range, the trigger size is approximately $\lambda_T \sim 10^{-5} - 10^{-2} \text{ cm}$.

A general energy deposition event in the WD can be roughly characterized by two parameters: a temperature T and heating length L . This energy deposit will eventually take the form of a local peak in the WD temperature profile. We define T to be the characteristic temperature and L the characteristic length scale of this local peak as it initially appears. L is determined by the efficiency with which a given energy deposition mechanism interacts with the stellar medium. To demonstrate the significance of L , suppose that kinetic energy were transferred directly to neighboring ions via short-range elastic scatterers. These ions would thermalize over their collisional time scale resulting in a heating length L of order the ion mean free path. In the other extreme, suppose that a process produces a large number of electrons with energy just above the Fermi energy. These electrons have Pauli-suppressed interactions with the medium and will travel a long distance before their energy is scattered and thermalized, resulting in a large L - possibly of order the stellar radius.

We find that a heating event will destroy a given WD as long as the energy density *and* energy deposited are sufficiently large:

$$T \gtrsim T_f, \quad L \gtrsim \lambda_T. \quad (1)$$

Note that this condition determines if a region of size L and temperature T *immediately* initiates runaway fusion. Instead, it is more useful to determine whether a given energy deposit will *eventually* initiate runaway fusion. In particular, a heating event that results in $T \gg T_f$ and $L < \lambda_T$ is perhaps initially dominated by diffusive cooling but can potentially evolve into an “explosive” temperature profile capable of igniting the star. When diffusion dominates, the thermal evolution will approximately be given by the dilution of the fixed total energy of the initial



FIG. 1: Minimum energy required to trigger explosion in a WD, based on numerical results for λ_T [?].

region over the final volume. Taking this into account, we can express the explosion condition (1) in terms of an excess energy within the temperature peak (rather than the temperature itself):

$$E \gtrsim n_i T_f \max\{L, \lambda_T\}^3. \quad (2)$$

In fact, there is an absolute minimum explosion energy given by

$$E_{\text{boom}} \sim n_i T_f \lambda_T^3 \sim 10^{14} - 10^{21} \text{ GeV}, \quad (3)$$

where the trigger size λ_T varies over a range of WD densities. This is plotted in Figure 1. However, this energy is sufficient only if it thermalizes within the trigger size. If an energy is deposited on a length scale larger than λ_T , the explosion energy required is given by (2) and may be parametrically larger. Since we are concerned with processes that have a fixed deliverable energy (the energy of the incoming DM), the most explosive processes will be those that result in the most localized heating. As a result, understanding the heating length L for various processes is critical to assessing their explosive potential. This is done in Section IV.

III. DARK MATTER EXPLOSIVENESS

In this section, we parameterize the ability of generic DM-SM interactions to destroy a white dwarf. We focus on interactions of the general type shown in Figure III, where energetic SM particles are created in the star after either a DM particle scatters off of a WD constituent or two DM particles collide. The first of these allows stellar heating via *transit* events as the DM continually releases energy while traversing the star, and the second process leads to heating via point-like *collision* events. Note that these collisions are not necessarily particle-antiparticle annihilations, but could be a more general and complicated processes, like the collisions of heavy nuclei. The heating physics and event rates are parametrically different for transits and collisions, and thus so are



FIG. 2: General interaction between ultra-heavy DM and WD constituents. [add collision diagram, remove \$\epsilon_i\$ and \$n_i\$ labels](#)

the resulting constraints. We will treat these two types of events separately below.

A. DM Transit

A DM transit releases energy continually along the DM particle's trajectory, resulting in a cylindrical temperature profile with some transverse length scale L . Ignition will be determined by both L and the amount of energy deposited per length along the DM trajectory, so we parameterize this interaction in terms of a *linear energy deposit* (LED) χ :

$$\chi_d = \frac{1}{n} \frac{\partial E_{\text{medium}}}{\partial x}$$

where n is the number density of the stellar medium and $\partial E/\partial x$ is the energy deposited to the stellar medium per DM distance traveled. We will also find it useful to discuss the kinetic energy lost by the DM, which is described by a similar quantity known as the *stopping power* (SP) χ_s :

$$\chi_s = -\frac{1}{n} \frac{\partial E_{DM}}{\partial x}$$

It may often be that $\chi_d = \chi_s$ and we will consider this special case below, but we maintain the distinction in general as it is not required that these two rates be equal. For example, if the DM-SM interaction converts some of the mass energy of the stellar medium into kinetic energy of outgoing SM particles, then $\chi_d > \chi_s$ (which will indeed turn out to be true for Q-ball dark matter).

Nominally, the energy deposit and loss per length will depend on the stellar density n , the instantaneous DM energy E , and the details of the DM-SM interaction. However, over the range of WD densities we consider it is reasonable that these rates scale linearly in n . [does this need more justification?](#) And, since all DM particles will enter the WD with speeds of order the escape velocity of the star, which is roughly a constant $v_{\text{esc}} \approx 10^{-2}$ for all WDs, the initial energy of a transiting DM particle depends only on its mass. Thus χ_d and χ_s for DM

transiting a WD are functions only of m_{DM} and its particular SM interactions. It is through χ_d , χ_s , and m_{DM} that DM transits will constrain allowable DM models.

crust shielding?

Setting the total energy released during a transit of length $\max\{\lambda_T, L\}$ to be greater than the explosive energy (2), we find a lower bound on the interaction cross section sufficient to trigger runaway fusion:

$$N_i \sigma_{\epsilon, i} \gtrsim \begin{cases} \lambda_T^2 \left(\frac{T_f}{\epsilon} \right) & L < \lambda_T \\ \lambda_T^2 \left(\frac{L}{\lambda_T} \right)^2 \left(\frac{T_f}{\epsilon} \right) & L > \lambda_T \end{cases}. \quad (4)$$

Note that in deriving (4), we require that the DM transit time is less than the diffusion time corresponding to the explosive temperature profile. This results in the condition

$$\tau_d \gtrsim \frac{\max\{\lambda_T, L\}}{v_{\text{esc}}}, \quad (5)$$

where τ_d is the characteristic time for a region of size $\max\{\lambda_T, L\}$ and temperature T_f to diffuse $\mathcal{O}(1)$ of its heat, and v_{esc} is the DM transit velocity (set by the escape velocity of the WD). An estimate from the heat equation yields $\tau_d \sim \frac{\lambda_T^2}{\alpha}$, where α is the (temperature-dependent) diffusivity. [Show condition is true for all densities.](#) Similarly, we also require that the time to transfer energy ϵ out to the initial temperature profile of length L is less than the diffusion time scale τ_d . Both of these conditions are necessary to ensure that energy depositions from individual collisions add coherently into a temperature profile during the full transit length. While a heating event can still be explosive if either of these conditions are violated, the formalism of (4) no longer holds in this case.

In addition, we require that a DM with this interaction makes it through the crust of the WD of size of density n_c and size R_c :

$$N_i \sigma_{\epsilon, i} \lesssim \left(\frac{m_{DM} v_{\text{esc}}^2}{n_c R_c} \right) \frac{1}{\epsilon}. \quad (6)$$

Imposing (4), we find lower bound on m_{DM} such that the DM is able to both penetrate the crust *and* trigger an explosion:

$$m_{DM} \gtrsim T_f \max\{\lambda_T, L\}^2 \left(\frac{n_c R_c}{v_{\text{esc}}^2} \right) \quad (7)$$

If (7) is violated, then the DM interaction is either not strong enough to ignite the WD or it is too strong that it has no chance to transit the interior as it gets stopped in the crust. In the case of $\lambda_T > L$, this translates to a model-independent lower-bound on the DM mass $\sim 10^{27} - 10^{34}$ GeV. [used 1 km crust and same number density as interior.](#)

We now calculate the maximum possible DM mass that may transit a single WD within its lifetime \sim Gyr. The

expected transit rate for an ultra-heavy DM state is given by

$$\Gamma_{\text{DM}} = n_{\text{DM}} \sigma_g v \sim \frac{\rho_{\text{DM}}}{m_{\text{DM}}} \pi R_s^2 \left(\frac{v_{\text{esc}}}{v} \right)^2 v, \quad (8)$$

where $v \sim 10^{-3}$ is the virial velocity of DM and ρ_{DM} is the energy density of DM in the region of interest. σ_g denotes the capture cross section including a gravitational Sommerfeld enhancement. Considering a $1.25 M_{\odot}$ WD in the local dark matter halo $\rho_{\text{DM}} \sim 0.3 \text{ GeV/cm}^3$, we find that $m_{\text{DM}} \lesssim 10^{44} \text{ GeV} \sim 10^{20} \text{ g}$ will transit the WD at least once in a Gyr. If we instead consider (recently discovered) heavy white dwarfs in the galactic center $\rho_{\text{DM}} \sim 10^3 \text{ GeV/cm}^3$, this upper bound improves to $m_{\text{DM}} \lesssim 10^{48} \text{ GeV} \sim 10^{24} \text{ g}$. [Hard to put this into “scaling” notation.](#)

B. DM-DM Collisions

IV. HEATING LENGTH

For the DM interactions under consideration, the released energy must be efficiently transferred to the WD in order to trigger runaway fusion. In this section, we summarize the available modes of energy loss for different SM particles in a WD - detailed calculations of the various interactions are presented in Appendix A. For a given process, the heating length L is determined by the distance over which an incident particle of energy ϵ and any secondaries deposit $\mathcal{O}(1)$ of the initial energy and the energy transfer thermalizes ion constituents. As we are primarily concerned with depositing sufficient energy to ignite supernovae, we focus on high-energy particles $\epsilon \gg T_f \sim \text{MeV}$ which interact via the strong and electromagnetic forces: electrons, photons, muons, and hadrons. In this way, the WD may be thought of as a “particle detector” with electromagnetic and hadronic “calorimeter” components.

A. Hadrons

The end products of a hadronic shower will generally be hadrons of kinetic energy $1 - 10 \text{ MeV}$ which are incapable of inducing further nuclear disintegration.

B. Electrons and Photons

C. Heavy Charged Particles

V. CONSTRAINTS

With the heating lengths in hand for a wide variety of possible heating processes, we are now ready to place constraints on ultra-heavy dark matter models based on

observations of white dwarfs and supernovae. As discussed in [?], there are two ways we may exclude a dark matter candidate: (1) by directly observing white dwarfs that should have blown up already due to the presence of the candidate in sufficient abundance, and (2) by observing a supernovae rate incompatible with the expected rate due to the candidate detonating white dwarfs.

A. Q-balls

In various supersymmetric extensions of the standard model (SM), non-topological solitons called Q-balls can be produced in the early universe [? ?]. If these Q-balls were stable, they would comprise a component of the dark matter today. In gauge-mediated models with flat scalar potentials, the Q-ball mass and radius are given by

$$M_Q \sim m_F Q^{3/4}, \quad R_Q \sim m_F^{-1} Q^{1/4}, \quad (9)$$

where m_F is related to the scale of supersymmetry breaking (messenger scale). The condition $M_Q/Q < m_p$ ensures that the Q-ball is stable against decay to nucleons [?]. When an (electrically neutral) baryonic Q-ball interacts with a nucleon, it absorbs its baryonic charge as a minimum-energy configuration and induces the dissociation of the nucleon into free quarks. During this process, $\sim \text{GeV}$ of energy is released through the emission of 2-3 pions [?]. The cross section for this interaction is approximately geometric:

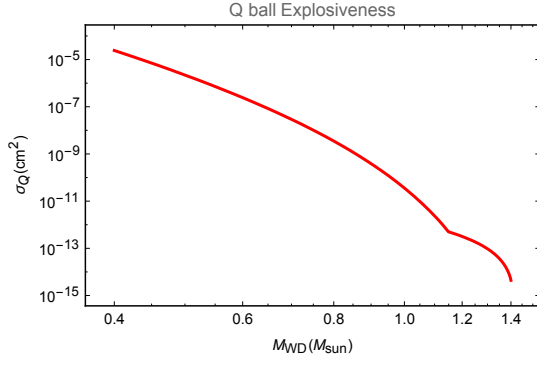
$$\sigma_Q \simeq \pi R_Q^2. \quad (10)$$

Note that a sufficiently massive Q-ball will become a black hole if the Q-ball radius is less than the Schwarzschild radius $R_Q \lesssim R_s \sim GM_Q$. In the model described above, this translates into a condition

$$m_F \left(\frac{M_{\text{pl}}}{m_F} \right)^3 \lesssim m_Q, \quad \left(\frac{M_{\text{pl}}}{m_F} \right)^4 \lesssim Q. \quad (11)$$

For Q-ball masses of this order, gravitational interactions become relevant.

We assume that for each Q-ball collision, there is equal probability to produce π^0, π^+ and π^- under the constraint of charge conservation. In the notation of Section III, this results in $N_{\pi} \sim 10$ pions released with $\epsilon \sim 500 \text{ MeV}$. The mean distance travelled by a relativistic particle before decaying is $d = \gamma v \tau$. For neutral pions $d_{\pi^0} \sim 10^{-5} \text{ cm}$ while for charged pions, $d_{\pi^{\pm}} \sim 10 \text{ m}$. Numerous experiments have studied the effects of 50–500 MeV pions incident upon complex nuclei targets such as carbon. It is found that there is roughly equal cross section of order $\mathcal{O}(100 \text{ mb})$ for a (neutral or charged) pion to either scatter elastically, scatter inelastically, or become absorbed with no final state pion [?]. Of these possibilities, pion absorption is the most relevant for energy loss as this will induce a hadronic shower.



Q-balls that transit a WD with sufficiently large cross section as given by (4) will ignite the star. Following the results of Section IV, we find that the relevant heating length L of the Q-ball interaction is simply. The resulting explosive power for Q-balls is plotted in Figure V A. If the Q-ball cross section is related to its mass and baryonic charge as in (9), we find that

$$m_Q \gtrsim 10^8 \text{ g} \left(\frac{m_F}{\text{TeV}} \right)^4, \quad Q \gtrsim 10^{38} \left(\frac{m_F}{\text{TeV}} \right)^4 \quad (12)$$

is capable of triggering runaway fusion in a heavy $\sim 1.25 M_\odot$ WD.

Need to implement SN rate and transit rate here...
Triangle Plot of Q-ball constraints vs. terrestrial.

VI. DISCUSSION

A. PARTICLE INTERACTIONS IN A WHITE DWARF

The interior of a WD is a complex environment (unless otherwise noted, we will assume a carbon-oxygen WD). Famously, the star is supported against collapse by electron degeneracy pressure with a characteristic Fermi energy

$$E_F = (3\pi^2 n_e)^{1/3} \sim 0.1 - 1 \text{ MeV} \quad (13)$$

where n_e is the number density of electrons. The nuclei are at an ambient temperature $T \sim \text{keV}$ and form a strongly-coupled plasma with coupling parameter

$$\Gamma \sim \frac{Z^2 \alpha}{n_i^{-1/3} T} \gg 1, \quad (14)$$

where n_i is the number density of nuclei. Here we provide a detailed analysis of the possible electromagnetic and strong interactions in a WD.

Coulomb Collisions

An incident particle of mass m , charge e , and velocity β scattering off a (stationary) target of mass M , charge

Ze with impact parameter b will transfer energy

$$E' = \frac{2Z^2 \alpha^2}{b^2 \beta^2 M}. \quad (15)$$

The differential cross section for the interaction is given by the Rutherford cross section

$$\frac{d\sigma}{dE'} = \frac{2\pi \alpha^2 Z^2}{M \beta^2} \frac{1}{E'^2}, \quad (16)$$

where we have assumed a sufficiently fast incident particle so that interactions are governed by single collisions with energy transfer E' [?]. Note that the differential cross section receives additional QED due to the incident spin, but for small energy transfers these corrections are negligible. It is straightforward to understand the parametric dependences of (16): there is increased likelihood to scatter for slowly moving incident particles undergoing “soft-scatters” against lighter targets. Therefore, one would expect that soft-scattering dominates the energy loss and that collisions with nuclei of mass M are suppressed by a factor $\mathcal{O}(\frac{Zm_e}{M})$ as compared to collisions with electrons. This is certainly true for incident charged particles in non-degenerate matter. However, both of these naive expectations turn out to be false when considering scattering off a degenerate species.

We first consider the energy loss of high-energy charged particles colliding with non-degenerate targets. In this case, the stopping power due to collisions with a number density n is given by:

$$-\frac{dE}{dx} = - \int dE' \left(\frac{d\sigma}{dE'} \right) n E' \quad (17)$$

$$\simeq - \frac{\pi n Z^2 \alpha^2}{M \beta^2} \log \left\{ \left(\frac{E_{\max}}{E_{\min}} \right)^2 + 1 \right\} \quad (18)$$

(17) must be integrated over all E' within the regime of validity for (16), thereby fixing the lower and upper bounds of the Coulomb logarithm. E_{kin} denotes the maximum possible energy transfer satisfying kinematic constraints:

$$E_{\text{kin}} = \frac{2M\beta^2\gamma^2}{1 + 2\gamma M/m + (M/m)^2}, \quad (19)$$

where $\gamma = (1 - \beta^2)^{-1/2}$ is the relativistic factor. Quantum mechanical uncertainty sets a limit to the accuracy that can be achieved in “aiming” an incident particle at a target $b > \frac{1}{\min\{M, m\}\beta\gamma}$. In terms of energy transfer, this translates to the condition

$$E' < E_{\text{quant}} = \frac{2Z^2 \alpha^2 \gamma^2 \min\{M, m\}^2}{M}. \quad (20)$$

In addition, the expression for the energy loss per scatter (15) is modified by relativistic considerations when the energy transfer is larger than the target mass [?]. For

our calculations, we take the maximum energy that an incident particle is able to transfer to be

$$E_{\max} = \min\{E_{\text{quant}}, E_{\text{kin}}, M\}. \quad (21)$$

On the other hand, the maximum impact parameter is set by Thomas-Fermi screening due to the degenerate electron gas:

$$\lambda_{\text{TF}} = \left(\frac{6\pi Z\alpha n_e}{E_F} \right)^{-1/2}, \quad (22)$$

This corresponds to a lower bound on the energy transfer

$$E' > E_{\text{sc}} = \frac{12\pi Z^3 \alpha^3 n_e}{\beta^2 M E_F}. \quad (23)$$

In addition, the lattice structure of ions in the WD introduces further complications [?]. The typical Coulomb lattice binding energy between ions of number density n is given by the electric potential

$$E_B \sim \frac{Z^2 \alpha}{n^{-1/3}} \sim 10^{-2} - 10^{-1} \text{ MeV} \quad (24)$$

For energy transfers greater than E_B , lattice effects can be safely ignored. On the other hand, momentum transfers from collisions below this threshold will lead to suppressed energy loss due to collective effects. For simplicity we set the lower bound on nuclei scattering (not applicable for electron targets) to be

$$E_{\min} = \max\{E_B, E_{\text{sc}}\} \quad (25)$$

so that we only account for energy transfers greater than the ionic lattice binding energy.

Now consider collisions with degenerate electrons. An incident particle transferring energy E' can only scatter those electrons within E' of the Fermi surface. We define a modified density of electrons $n_e(E')$ as:

$$n_e(E') = \begin{cases} \int_{E_F - E'}^{E_F} dE g(E) & E_F \geq E' \\ n_e & E_F \leq E' \end{cases}, \quad (26)$$

where $g(E)$ is the density of states per unit volume for a three-dimensional free electron gas. The effect of degeneracy can also be cast as a suppression of the differential cross section of order $\mathcal{O}(E'/E_F)$ whenever energy less than the Fermi energy is transferred. Therefore, unlike in the non-degenerate case, the energy loss due to soft-scatters are in fact *subdominant* to the contributions from rare hard-scatters. The stopping power is also sensitive to the integration bounds E_{\max} and E_{\min} as a power-law dependence unlike the logarithmic sensitivity (17) in the case of a non-degenerate target.

Compton Scattering

High-energy photons can efficiently deposit their energy to electrons in the WD via Compton scattering

$$k' = \frac{k}{1 + \frac{k}{m_e c^2} (1 - \cos \theta)}, \quad (27)$$

where k and θ denote the energy and angle of deflection for the incident photon, respectively. Clearly when $\theta \rightarrow 0$, none of the initial energy is transferred. However, for generic scattering angles and sufficiently energetic photons $k \gg m_e$, the energy transfer due to a single Compton interaction is typically $\mathcal{O}(1)$ of the initial energy. The Klein-Nishina cross section integrated over all scattering angles is given by

$$\sigma_{\text{comp}} \approx \frac{\pi \alpha^2}{2m_e k} \left\{ 1 + \log \left(\frac{4k^2}{m_e^2} \right) \right\}, \quad k \gg m_e, \quad (28)$$

evaluated in the rest frame of the target electron. Since the electrons in a WD are sufficiently relativistic, the Thomson scattering limit is never applicable. In addition, we ignore the effects of degeneracy when the incident photon transfers is highly energetic such that scatter is sufficient to knock electrons out of the Fermi sea.

Bremsstrahlung and Pair Production

Nuclear Interactions

Nuclear interactions can be either elastic or nonelastic. A single, backward elastic scatter would result in an incident particle losing virtually all of its energy if the incident and target masses are the same. However, we will be primarily concerned with light hadrons incident on relatively heavy nuclei, i.e. ping-pong balls bouncing around a sea of bowling balls. An elastic collision between an incident (non-relativistic) mass m and a heavy, stationary target mass M results in a final energy (averaged over all scattering angles)

$$\zeta \equiv \frac{E_f}{E_0} \approx \left(\frac{M}{m + M} \right)^2, \quad m < M. \quad (29)$$

where E_0 and E_f denote the initial and final energy of the incident particle. Therefore, a nucleon interacting with carbon nuclei transfers roughly 15% of its incident energy after each elastic collision.

The cross section for elastic scatters has considerable dependence on the incident particle species. For neutrons and protons, the nuclear elastic cross section is of order $\sigma_{\text{el}} \sim 100 \text{ mb}$ at energies greater than $\sim \text{MeV}$. However at lower energies, the elastic cross section effectively vanishes for protons but rises to $\mathcal{O}(\text{b})$ for neutrons. This can be understood physically as charged hadrons may only elastically scatter off nuclei if the incident energy is sufficient to overcome the mutual Coulomb barrier. Note that

the neutron capture cross section in carbon and oxygen is negligible compared to the elastic cross section. However, since the ions in a WD form a Coulomb lattice, elastic collisions cannot transfer energy less than (24) without exciting phonon modes. This limits the applicability of the energy loss function (29) to energies above \sim MeV for both neutrons and proton elastic collisions.

Next we consider nonelastic collisions, which are relevant for highly energetic incident particles. Namely, a particle of incident energy greater than the nuclear binding energy $E_{\text{nuc}} \sim 10$ MeV can interact to produce an $\mathcal{O}(1)$ number of energetic secondary hadrons (i.e. protons, neutrons, pions, etc.) emitted roughly in the direction of the primary particle. These secondary particles approximately split the initial total energy of the absorbed primary particle and can collide with other nuclei in the WD. The remnant nucleus will generally be left in an excited state and relax through the emission of slow $\mathcal{O}(10$ MeV) hadrons (“nuclear evaporation”) and photons [?]. A hadronic shower is the result of all such reactions caused by primary and secondary particles.

For simplicity, we take the nuclear mean free path for a hadron (proton, neutron, pion) to be constant in energy

$$l_{\text{had}} \sim \frac{1}{n\sigma_{\text{non}}}, \quad \sigma_{\text{non}} \sim 100 \text{ mb.} \quad (30)$$

Note that the shower length is only logarithmically sensitive to the initial energy or number of secondaries, and is thus primarily set by the nuclear mean free path. In this case, the cascade induced by a high-energy initial hadron is adequately described by a stopping power

$$\left(\frac{dE}{dx}\right)_{\text{had}} \sim \frac{E}{l_{\text{had}}}, \quad (31)$$

where we have dropped terms of logarithmic order. The shower will end once final-state hadrons reach a critical energy E_c - this is either set by the scale at which an

additional mechanism dominates the stopping power or the minimum energy required to induce further showers E_{nuc} . In this model, the total hadronic shower length X_{had} for an incident hadron of energy $E_0 > 10$ MeV is given by

$$X_{\text{had}} \sim l_{\text{had}} \log(E_0/E_c) \approx 10 l_{\text{had}}. \quad (32)$$

High-energy $k \gtrsim 10$ MeV photons can also interact with nuclei and induce a hadronic shower. There is considerable uncertainty in the evaluation of high-energy photonuclear cross sections, and a precise formulation is beyond the scope of this work. As a conservative estimate, we take a constant photonuclear mean free path and cross section:

$$l_{\gamma} \sim \frac{1}{n\sigma_{\gamma A}}, \quad \sigma_{\gamma A} \sim 1 \text{ mb.} \quad (33)$$

Similarly, charged leptons will lose energy via electronuclear interactions in which an incident electron or muon radiates a virtual photon that then interacts hadronically with a nearby nucleus. This radiative stopping power has been computed in [?] in terms of the photonuclear mean free path, and is of order

$$\left(\frac{dE}{dx}\right)_{eA} \sim E \left(\frac{\alpha}{l_{\gamma}}\right). \quad (34)$$

Acknowledgements

We are grateful to S. Rajendran for suggesting this project and for helpful conversations. We would also like to thank D. Grabowska, K. Harigaya, R. McGehee, and J. Wurtele for stimulating discussions.