

Suppressing low energy background scintillation signals in SBND

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Abstract

Scintillation light produced in large liquid argon detectors provide the opportunity to measure interesting low energy physics such as supernova neutrinos and nucleon decay. Several potential low energy backgrounds provide constant signals to the light detection system and could make detection of these signals impossible. In ^{39}Ar decays the background occurs at a high rate compared to the arrival of supernova neutrinos. Using the fast optical simulation library, the backgrounds as well as supernova neutrinos can be quickly simulated throughout the SBND detector to compare the ability of the light system to detect these signals. This document explores a simple set of cuts to remove the ^{39}Ar backgrounds while preserving the supernova signal for three different light system configurations.

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1 Scintillation from Low Energy Events in Liquid Argon Time Projection Chambers

Liquid argon time projection chambers make use of argon's chemical inertness by drifting charged particles through the detector to readout wires for distances as long as meters. Interactions inside the liquid argon volume produce an ion-electron pair, or result in an excitation of an argon atom. Upon recombination of the ion-electron pair or deexcitation of the argon atom scintillation light is produced. The amount of charge reaching the wires and the amount of light produced are inversely correlated as a function of the electric field magnitude. With an increase in field strength, the amount of scintillation light decreases, as the recombination of electron-ion pairs decreases. In LArTPCs the scintillation light could improve the detection of lower energy events which are harder to detect using charge.

While SBND is being constructed primarily as a beam-line neutrino detector, and R&D for future large liquid argon neutrino detectors (DUNE), the light detection system also provides a good opportunity to study low energy (sub 50 MeV) events. Galactic supernova neutrinos peak with an energy around 20 MeV, making the LDS the prime candidate to detect them, however the LDS is also being constantly bombarded by background from naturally occurring ^{39}Ar . We show that the constant low energy background originating from ^{39}Ar can be limited by setting thresholds on the PMTs as well as forming coincidences. The methods used to limit these backgrounds are explored further in the following sections. Sections 3 and 4.1 discuss the details of a simulation written to compare the ^{39}Ar background and supernova neutrinos as well as conclusions drawn from the simulation regarding removing the background signal.

2 ^{39}Ar and Supernova Neutrinos

^{39}Ar beta-decays with $Q = 0.565$ MeV and is created by cosmic ray interactions. The resulting activity of 1 Bq/kg means detectors like SBND will have a constant supply of low energy electrons being emitted throughout the detector (1). Figure 1 shows the beta-decay energy spectrum for ^{39}Ar , where the most common decay energy is around 200 keV. The emitted electrons ionise the liquid argon and produce scintillation photons, where the number produced is a function of the scintillation yield of the medium and the energy of the electron. Given SBND's expected electric field magnitude of 500 V/cm, the scintillation yield is taken to be 24,000 γ /MeV for low mass particles, meaning an average decay produces roughly 4,800 photons (2). The constant background of scintillation light from the ^{39}Ar decays could make it impossible to distinguish between background generated photons and other low energy signals seen in the detector, such as supernova neutrinos.

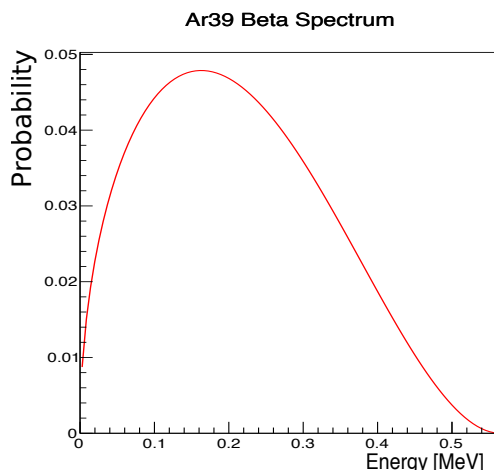


Figure 1: β -decay spectrum of ^{39}Ar .

In general, the ^{39}Ar events should not be a problem for beam neutrino events, as the energies of events coming from BNB or NuMI neutrino interactions will be much higher, and benefit from being in coincidence with the beam gate. Supernova neutrinos on the other hand, will be very few and have an energy of no more than 50 MeV (Figure 2). The spectrum assumes a supernova model with complete

Boltzmann neutrino transport and spherical symmetry (3). For a given supernova, detectors the size of MicroBooNE and SBND can expect on the order of 10 supernova neutrinos total at 1 kpc. This is compared to ^{39}Ar events, where in 1 second a detector of SBND's size can expect around 56,000 decays in a single TPC. Although the background rate is significantly higher, each supernova event interacts with roughly a factor 10 or more energy.

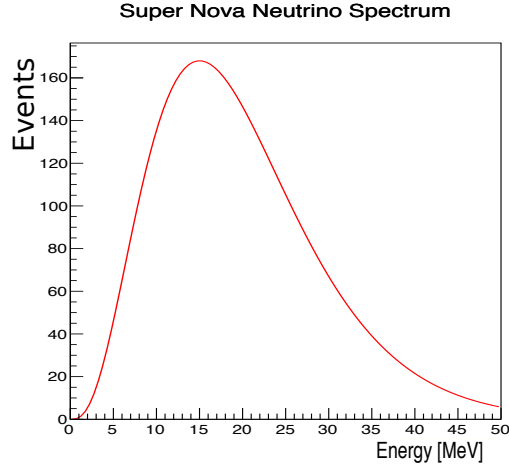


Figure 2: Expected energy spectrum of interacting neutrinos for galactic supernova neutrino flux (3).

3 Simulation Details

The simulation used here builds on work done in the tech note DocDB #1155, including great detail about the fast optical simulation, validation, and characteristics of scintillation light in SBND. The results presented here are an application of the simulation, using standalone code outside of the Larsoft framework.

3.1 Fast Simulation and Configurations Used

The background scintillation light produced by ^{39}Ar in a liquid argon TPC can be modelled by making use of fast optical simulation libraries. These libraries function as look-up tables for the “visibility” parameter, which is a probability that a photon originating at a specific location in the detector reaches a given PMT. This method is particularly useful, as tracking each generated photon individually through the detector is CPU and memory intensive. The visibility is tied to the light yield of the detector, where a higher light yield configuration results in a greater number of photons detected. The different configurations covered in this note are shown in Figure 3. The configurations include

- “No foils” - TPB-coated PMTs with opaque cathode
- “Cathode only” - TPB-coated PMTs with TPB-coated reflective foils on cathode
- “Full coverage” - TPB-coated PMTs with TPB-coated reflective foils on cathode and field cage

where adding further TPB-coated foils increases the light yield of the system.

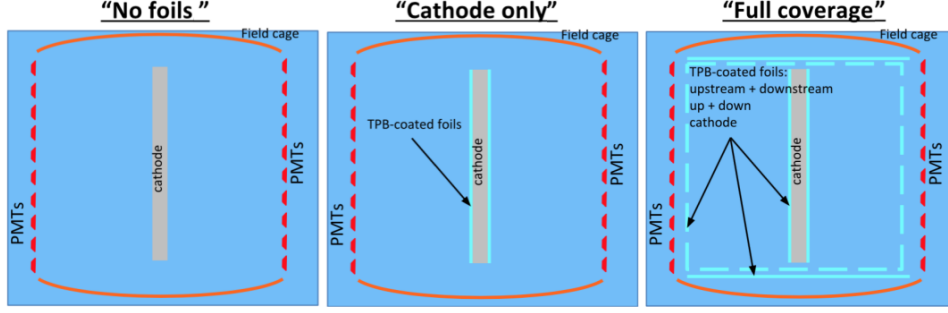


Figure 3: Light detection system configurations considered in this note. (Left) No TPB reflecting foils, only TPB on the PMTs. (Centre) Opaque cathode covered with TPB reflecting foil. (Right) Field cage and cathode covered with foils.

The libraries make use of voxelisation to discretise the location of a photon inside the detector. In the cases where TPB reflective foils are included, the library also distinguishes between direct and reflected visibility. In addition to the geometrical considerations, second order effects like Rayleigh scattering and a 50% efficiency for direct light incident on the PMT, due to TPB re-emission are already included. Exact details can be found in DodDB #1155.

3.2 Photons from Low Energy Events

The number of photons generated by an ionising particle can be expressed simply as the particle's energy deposition multiplied by a known scintillation yield, where the scintillation yield is a function of the electric field and deposition profile. Thus by including the visibility, the number of photons reaching the PMTs is calculated. This value is defined for each voxel-PMT combination. Lastly we apply the quantum efficiency for the PMTs to end up with the total number of detected photons on a PMT given an ionising particle deposits a certain amount of energy at a given location in the detector. This idea is expressed in the formula:

$$N_{\gamma} = \text{Poisson}[\text{Visibility} \times dE \times SY] \times QE \quad (1)$$

where SY is the scintillation yield and QE the PMT quantum efficiency. From before, given SBND's expected electric field magnitude of 500 V/cm, the scintillation yield is taken to be 24,000 γ /MeV for low mass particles. dE is the total energy the particle deposited into the medium, and Visibility a parameter extracted from the optical libraries. A typical ^{39}Ar decay will result in up to around 50 photoelectrons integrated over all PMTs, which is small compared to a beam neutrino event.

To model the background we calculate the expected number of decays expected in a fixed amount of time from the activity of ^{39}Ar , the mass of the liquid argon inside the detector, and the timing window:

$$N_{\text{decays}} = \text{Poisson}[\text{Activity} \times \text{time} \times \text{mass}] \quad (2)$$

where in this case the activity is given in Bq/kg. For example, taking the planned amount of liquid argon in one SBND TPC (56 tonne), a readout window of 1.2 ms and the activity of ^{39}Ar of 1 Bq/kg leads to about 67 events. The decays of each ^{39}Ar atom are treated as independent and are fluctuated by a Poisson distribution.

The visibility depends on several factors, such as the light yield of the system and the location inside the detector. With respect to the simulation, this means we select a detector configuration and randomly assign the decay a voxel within the detector volume. Finally the energy of the decay is generated from the ^{39}Ar decay spectrum. Together with the position, Equation 1 gives the number of photons generated per PMT.

The supernova neutrino events are simulated in a similar fashion to the ^{39}Ar events, where their location is randomly selected inside the detector and the energy spectrum is taken from 2. The photons of each supernova event are tracked like the photons from the ^{39}Ar decays.

Due to the small number of supernova events distributed over a few 10s of milliseconds, each event can be tracked separately in time. Whereas many ^{39}Ar events occur within several microseconds and may have overlapping signals. Section 3.3 discusses how timing is implemented into the simulation.

3.3 Simulating Timing

The arrival times of photons coming from events with the same energy and position can be substantially different, because photons generated by the event do not necessarily arrive at the same, due to transport effects. It is also likely that photons from two different events can arrive very close to one another in time. The fast optical simulation libraries contain a parameterised timing which accounts for effects of the propagation. Based on the wavelength of the photon and its distance, the time needed for the photon to travel from its point of origin to the PMT is calculated. The timing parameterisation is explained in detail in DocDB 1155. For the purpose of this simulation, the time each photon takes to propagate from the point of creation to the point of detection is defined as t_{trans} . Each photon is assigned a transport time value, where the typical time is on the order of several 10s of nanoseconds.

A larger variation of the arrival time comes from the scintillation time of the argon in the detector. As argon scintillation is governed by two separate time constants due to its singlet and triplet decay channels, this gives the emission time a spectrum between a few nanoseconds and several microseconds, seen in Figure 4. This is also calculated per photon and will be defined as t_{Scint} (2).

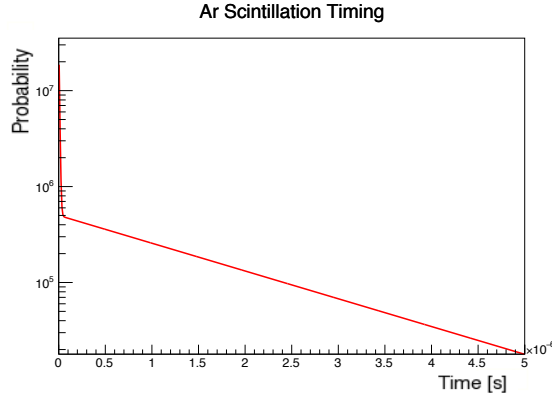


Figure 4: Scintillation timing for argon. Note the two components, where the fast component is the singlet decay channel composing roughly 25% of the light emitted, while the triplet decay channel composes roughly 75%.

The last component for the time comes from the time of the decay. The ^{39}Ar decays are assumed to be uniform in time, meaning t_{Decay} is pulled randomly over a predetermined time window. This is calculated per event, whereas the other two components are calculated per photon. Thus the recorded time of the photoelectron is a simple sum of these three terms:

$$t_{total} = t_{Decay} + t_{Scint} + t_{trans}$$

where t_{total} is calculated per photon arriving at the PMTs. The same procedure is used for both the photons generated by ^{39}Ar events and supernova neutrinos. By tracking the arrival times of all photons and on which PMT they arrive, we can study the effects of overlapping ^{39}Ar events on supernova neutrino events. It should be pointed out that the simulation makes certain simplifying assumptions, such as all energy released in the ^{39}Ar is instantly deposited inside the detector. Other simplifying assumptions include how the supernova neutrino interactions are modelled. As the intention of this simulation is not to test neutrino interaction Monte Carlo, the neutrino interactions are assumed to deposit all of their energy in a single voxel. This greatly simplifies the simulation, and should have little impact on results, as the variation in visibility between neighbouring voxels is rather small.

3.4 Validating the Simulation

In this section we use the simulation to produce “sanity” plots to cross-check if the simulation is properly reproducing the expected response of the system. From the average energy of a supernova neutrino and the average energy of a ^{39}Ar decay, the supernova neutrino will clearly produce a much larger signal - at least a factor of 10 more photons.

3.4.1 Independent Event Validation

Figure 5 compares single isolated events, which provide examples of expected topologies from different events for the cathode foils configuration. The Figure shows the SBND PMT array for a Q-value ^{39}Ar decay, a lower energy 5 MeV supernova neutrino, and an average value 22 MeV supernova neutrino, where each event was generated in the same voxel in the centre of the TPC. For the supernova events, PMTs much further away from the interaction point see a signal, while the overall signal for the individual ^{39}Ar decay is small across most PMTs.

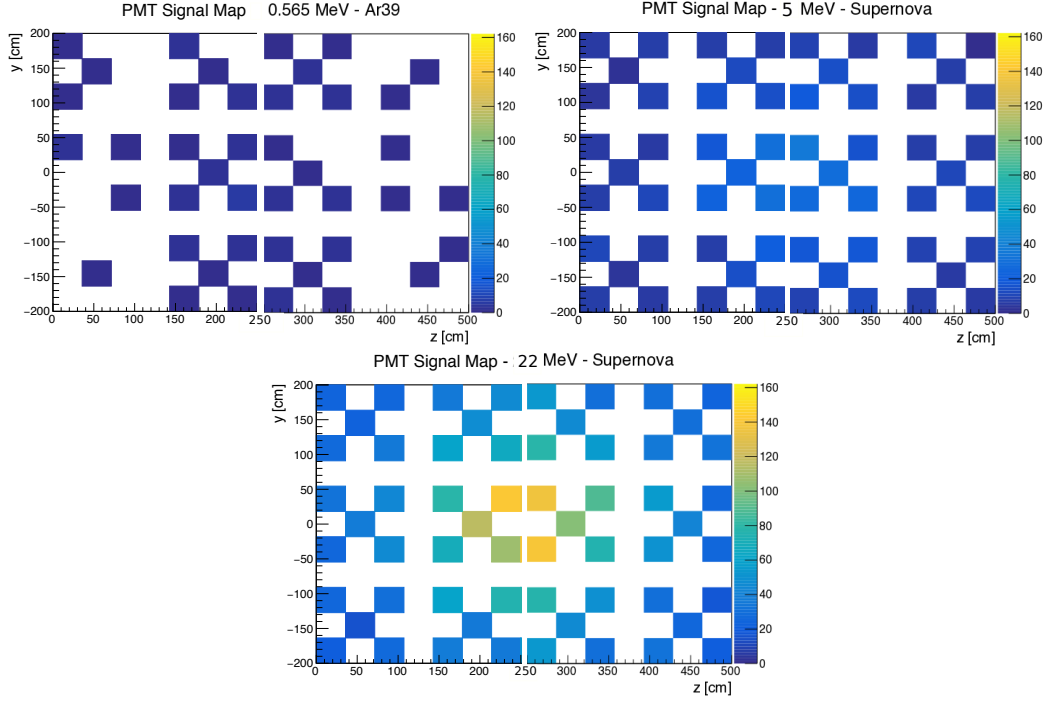


Figure 5: Comparison of single high-energy ^{39}Ar , 5 MeV SN, and 22 MeV SN events in middle of detector using 60 PMTs. Single high-energy ^{39}Ar expects about 50 hits, 5 MeV SN about 700 hits, and 22 MeV SN about 3000 hits. Shown are the results for the total light (VUV+visible) in the case of cathode reflective foils. Color (z-axis) corresponds to number of photoelectrons.

These plots demonstrate the spatial distribution of the photoelectrons from an event as well as their intensity. When comparing even the highest energy ^{39}Ar decays and a below average supernova neutrino, the single events can be rather clearly distinguished from one another, assuming a relatively uniform energy reconstruction throughout the detector (see Section 4.1).

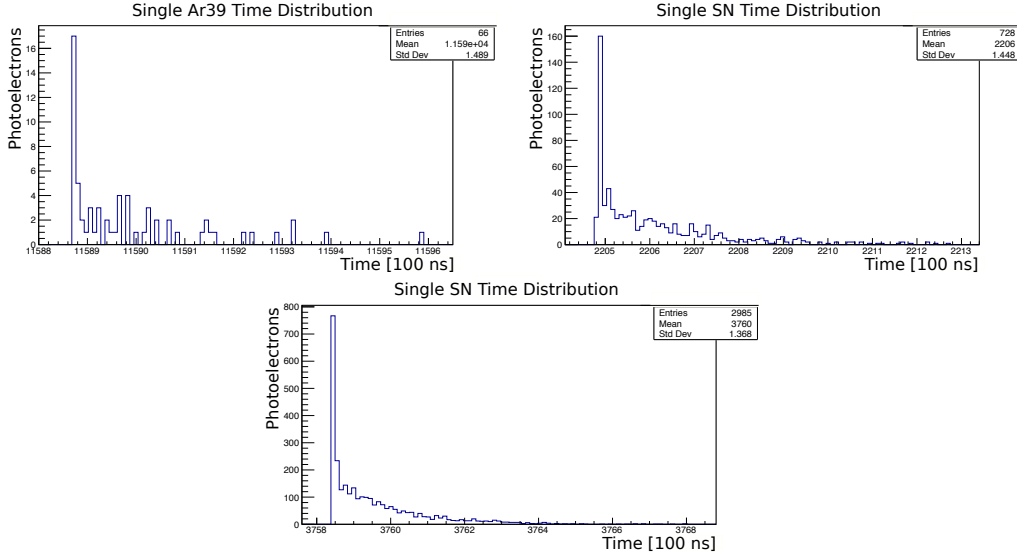


Figure 6: Three different signal time (t_{total}) distributions all produced in the centre of the detector for the cathode TPB reflective foils configuration. ^{39}Ar event with energy 0.565 MeV, supernova neutrino (up-right) with energy 5 MeV, and supernova neutrino (centre-bottom) with energy 20 MeV. Though the supernova neutrino events produce far more photons than the ^{39}Ar decay, the timing distributions are roughly the same (start-time offset is arbitrary).

We can also plot the photoelectron arrival times for the cathode TPB reflective foils configuration are shown in Figure 6. The factor 10 difference in energy between the single Q-value ^{39}Ar decay and a 5 MeV supernova neutrino results in roughly factor 10 increase in overall photoelectron count. The fast and slow light components are noticeable, with a sharp peak near the first events, which tapers off. The time over which a majority of the signal arrives is on the order of 0.8-1.0 μs for all cases (the start-time offset is arbitrary). This time is small compared to the supernova interaction rate. However given the amount of ^{39}Ar events and the comparatively long time constant, the background events could overlap in time and mimic a higher energy signature than expected from a single interaction. For this reason we will not treat them separately.

3.4.2 Time Dependant Validation

For the results shown in Section 4.3 and beyond, three different LDS configurations are considered, each including over 3×10^6 ^{39}Ar events (compared with 320,000 voxels). As the location of the event in the detector changes not only the visibility, but also the transport time, it becomes increasingly important to have a large sample of events which minimise statistical fluctuations. This is most relevant for events which only produce several photoelectrons, like ^{39}Ar .

We apply the timing method from Section 3.3 creating Figure 7 which shows one frame of the PMT readout for the ^{39}Ar dataset. This plot seeks to display how the structure of the decays are still visible, but that the long tails of the timing distributions will readily combine with the peaks of other decays close in time. Section 4.3 goes on to consider rejecting the ^{39}Ar events in-time compared to individual supernova events.

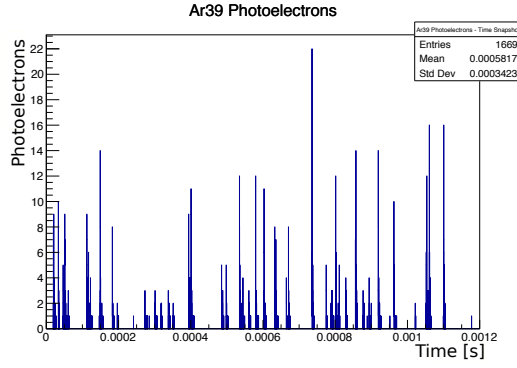


Figure 7: Number of photoelectrons vs time (t_{total}), integrated over all 60 PMTs. The binning is in 100 ns for a time period of 1.2 ms (readout window). This is for the full TPB reflective foil configuration.

For the large data sets, the distribution of detected photoelectrons in y (height) and z (beam-direction) was shown to be centred around the mid-way points: $y \approx 0$ and $z \approx 250$. This is because the PMTs are roughly evenly distributed in the y and z planes as well as the events being uniformly created. However when considering x (drift-direction), the light yield uniformity of the system makes a difference. Table 1 shows the average x position. The average distance from the photocathode decreases as the light yield of the system decreases. This can also be clearly seen in Figure 8 for each of the three cases, where the data points correspond to detected photoelectrons.

Configuration	Ave. X Position [cm]
Full Foils	95.14
Cathode	97.31
No Foils	146.4

Table 1: Average drift-direction position for detected photons given different configurations.

As Figures 8, 9 and 10 show the number of detected photoelectrons, these plots display how the location inside the detector plays a key role for the detection of photons. In Figure 8 we observe the behaviour for the no foils setup - the visibility decreases further away from the PMT plane, resulting in few events being detected far away from the PMTs. The addition of the cathode reflector foil causes an obvious enhancement at distances greater than 140 cm, (foil located at $x = 0$) as the light coming from that region has now been wavelength-shifted by the foils and propagates as visible light inside the detector. The addition of further reflective foils for the full foils configuration results in a more uniform distribution of light, where the trough disappears. In fact, as there are no foils behind the PMT plane, the number of photoelectrons detected is actually largest further away from the PMTs, where the photons are mostly likely to hit a TPB-covered foil.

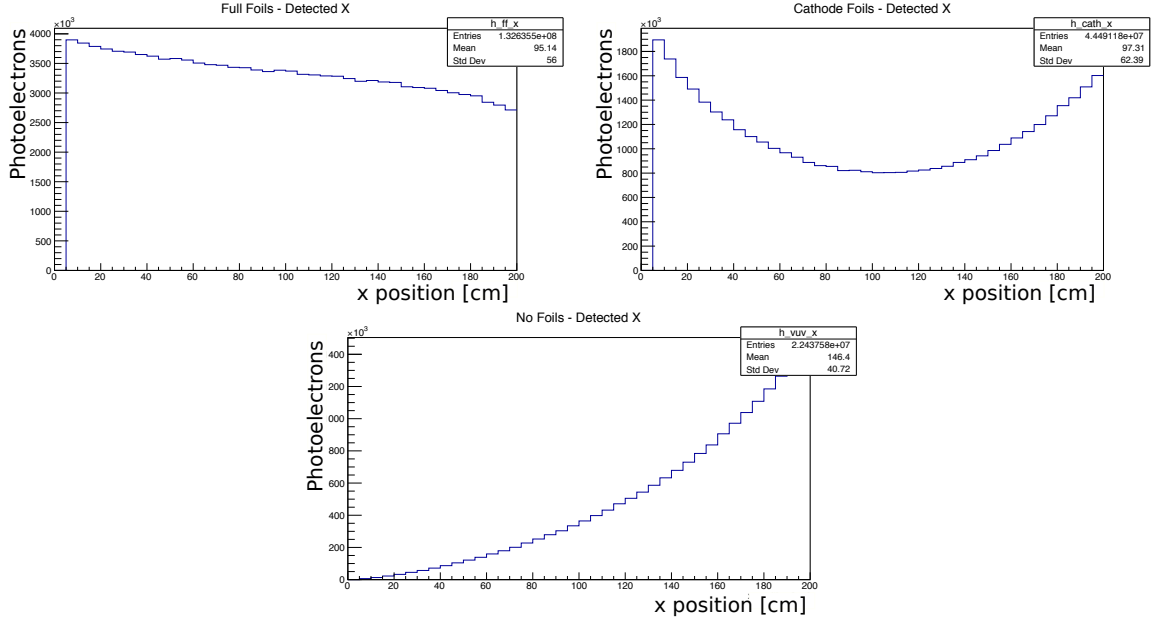


Figure 8: Detected photoelectrons versus x position in the detector.

In the y and z distributions, all three configurations have the values maximised around the centre of the detector, as would be expected. For the full foils case, we find that the events occurring near the edges of the detector have the highest number of detected photoelectrons. This is due to reflective foils close to the event origin, meaning events near the edges of the detector have a large visible light component. This is contrasted with the cathode and no foils case, where the centre of the detector sees a higher contribution. The shapes of these two cases are rather similar, where the photoelectrons produced at the edges contribute very little, and the two troughs match roughly the gap between the sets of PMTs in both the y and z directions. (See the first plot in Figure 5) This shows that the detection of low energy ^{39}Ar decays, in the cathode or no foils case are very sensitive to the configuration of the PMT array, and therefore the location inside the detector.

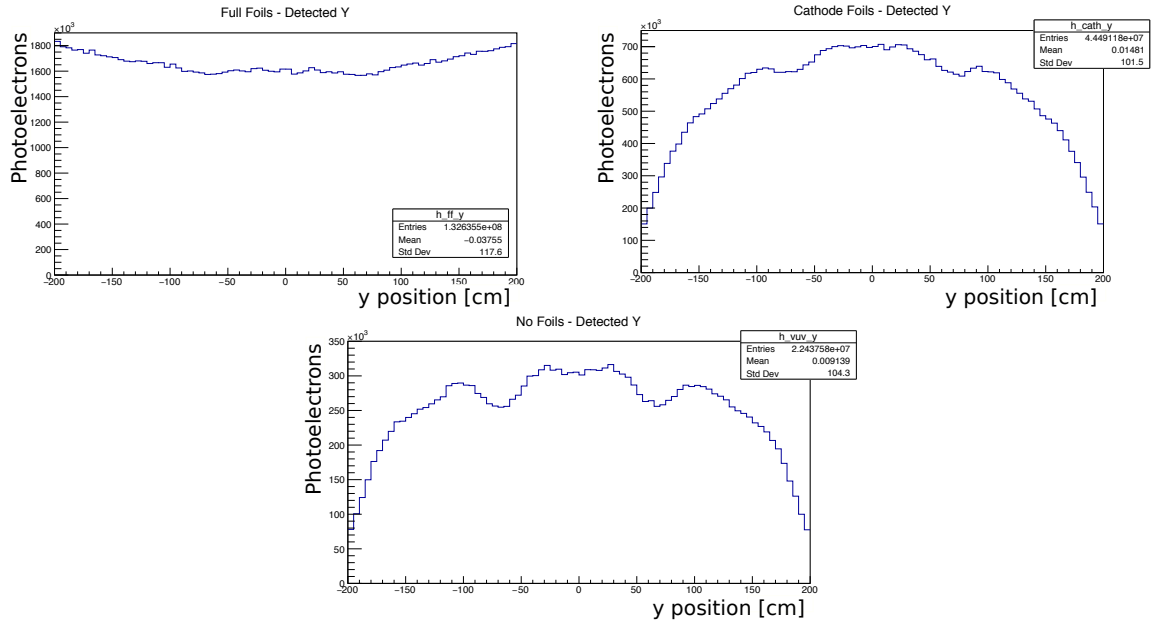


Figure 9: Detected photoelectrons versus y position in the detector.

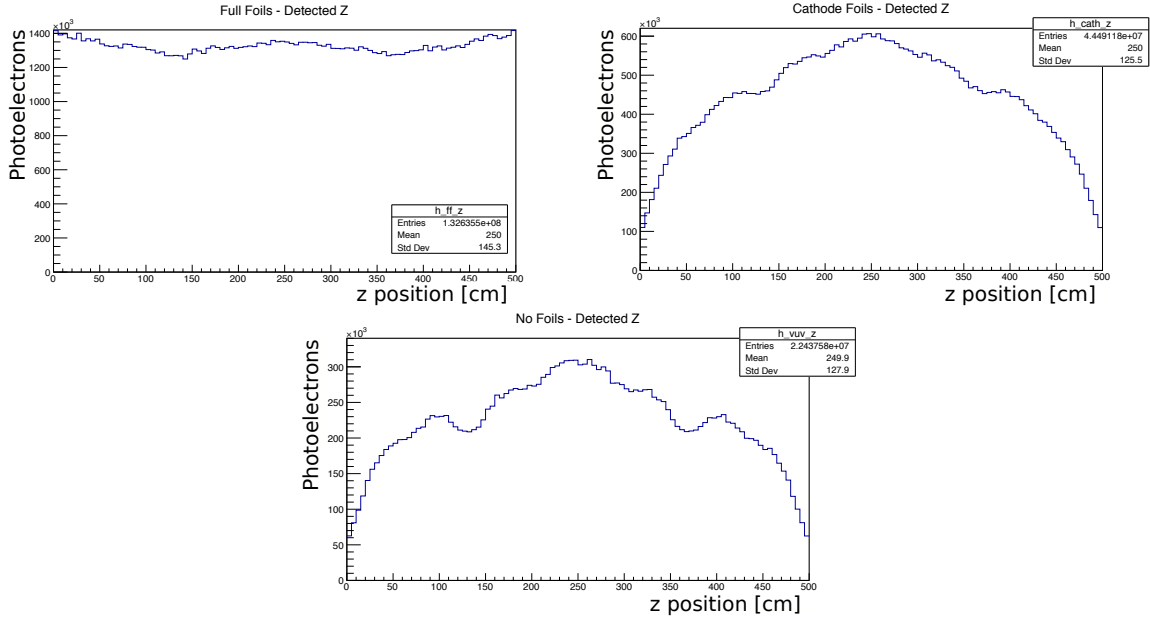


Figure 10: Detected photoelectrons versus z position in the detector

4 Limiting Background Signals

The number of expected supernova neutrino events for a liquid argon detector the size of SBND is small. In order to unambiguously observe them, a high trigger efficiency is necessary. This section outlines some ^{39}Ar background and simulated supernova neutrino events as described in Section 3. First we consider ^{39}Ar decays as independent in time, and we assume all of its photons arrive simultaneously. Despite being unphysical in that the ^{39}Ar decays could realistically overlap, this test is demonstrative of the different features of single ^{39}Ar and supernova events and how a basic minimum threshold trigger cuts the background spectrum. We then constrain single events to the first 100 nanoseconds of their signal, rather than the total light for each event. This is meant to correspond to the “fast light” arriving from an event in the usual trigger window. Finally we examine what happens when photons from different ^{39}Ar decays are allowed to overlap in time, and how a similar cut to the single events case can prevent triggering falsely on background.

4.1 Time Independent Cut Efficiencies

We first look at single events counting their total energies. This is the simplest case and further elements are added in the next sections. The format of the cut for this simple case is outlined in two parts:

- Variable individual PMT threshold
- Variable number of individual PMTs

where the individual PMT thresholds are uniform across all PMTs. An example cut could involve at least 3 PMTs each measuring 2+ photoelectrons apiece (giving a minimum of 6 photoelectrons). The supernova events are of higher energy than the ^{39}Ar , meaning the supernova events will typically produce more signal. Thus by increasing either the number of PMTs or the individual PMT thresholds, we expect to see only higher energy events.

Table 2 outlines the basic cuts and amount of the signal/background removed by the cut for both the full reflective foil and no reflective foil configurations. Taken from a sample size with statistical uncertainty of about 1.7%, the cuts for this initial study erred on the side of being overly conservative. While these cuts serve as preliminary guidelines, further studies outlined in Section 4.3 deal with considerably larger statistics. Several of these cuts are implemented in plots below. We do not apply the most restrictive cuts to the VUV only case as the amount of light is significantly reduced for this case.

No. PE	No. PMT	% ^{39}Ar	% ^{39}Ar VUV	% SN	% SN VUV
1	1	40.2	84.5	0.01	0.01
2	1	60.3	88.8	0.01	0.2
2	2	70.7	90.2	0.01	0.2
3	2	87.4	92.9	0.01	0.4
3	3	92.0	97.9	0.01	0.6
3	4	95.2	X	0.01	X
4	2	96.8	X	0.01	X
4	3	96.8	X	0.01	X
5	3	97.4	X	0.01	X

Table 2: % Reduction in signal and background compared for the full reflective foil and no reflective foils configurations. “VUV” indicates the no reflective foil case and “X” indicates that the cut is not relevant for this case.

By performing a simple cut of several photoelectrons incident over a few PMTs there is only a small reduction in supernova neutrino signal, but a significant reduction in the ^{39}Ar background. Figure 11 and 12 visualise how these cuts affect the ^{39}Ar and supernova neutrino energy spectra for the full foil and no foil cases.

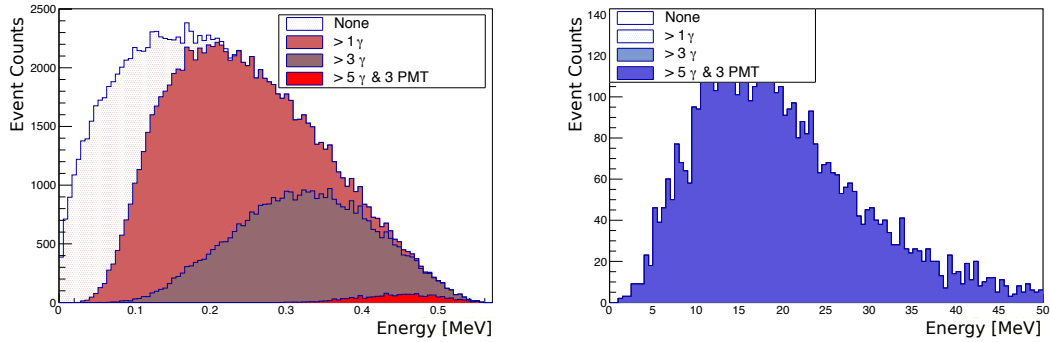


Figure 11: Energy spectra for ^{39}Ar and supernova neutrinos, assuming a large statistics sample for both, to observe how cuts on the number of photoelectrons per a number of PMTs modify the spectra. This case assumes a full coverage of reflective foils, and considers the contribution from both direct VUV and reflected visible light. The goal for the ^{39}Ar spectrum is to remove it completely through these cuts, and therefore almost disappears from the plot in the case of the most aggressive cuts. Legend shows cuts where “1 γ ” corresponds to the cut 1 photoelectron on 1 PMT, and “3 γ ” corresponds to an increase in the individual PMT threshold to 3 photoelectrons on 1 PMT.

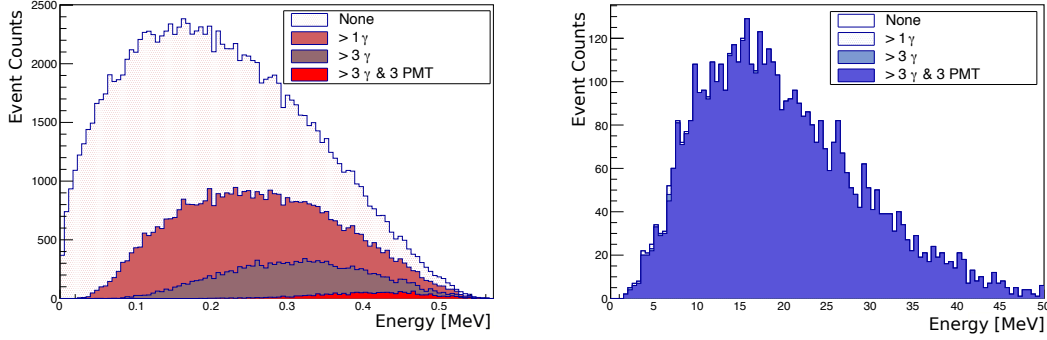


Figure 12: For the no reflective foils configuration - Energy spectra for ^{39}Ar and supernova neutrinos, assuming a large statistics sample for both, to observe how cuts on the number of photoelectrons per a number of PMTs modify the spectra. The goal for the ^{39}Ar spectrum is to remove it completely through these cuts, and therefore almost disappears from the plot in the case of the most aggressive cuts. Legend shows cuts where “ 1γ ” corresponds to the cut 1 photoelectron on 1 PMT, and “ 3γ ” corresponds to an increase in the individual PMT threshold to 3 photoelectrons on 1 PMT.

As the cuts become more stringent, the ^{39}Ar spectrum almost disappears, while the supernova spectrum is mostly preserved. Additionally, as the cuts require more photoelectrons, intuitively this limits all but the higher energy ^{39}Ar decays, which occur with a lower frequency. By requiring a signal on several different PMTs, the cuts place a geometry constraint on ^{39}Ar events occurring in low visibility voxels. Using the same cuts, we can look at the number of PMTs seeing at least one photoelectron per event in Figure 13. In a comparison between the two sets of plots: the higher light yield system preserves both the low energy supernova and low energy ^{39}Ar . However, as also indicated in Table 2, the no foils configuration begins to lose efficiency to increase purity, resulting in some of the lowest energy supernova events being cut.

Figure 13 shows how many PMTs see any signal from the scintillation events. In the full foils configuration, even the lowest energy supernova neutrinos are detected by over half of the PMTs. Conversely a majority of the ^{39}Ar events are detected on less than half of the PMTs. By using the cuts explored above, where a threshold on the individual PMTs is set, we see how the number of PMTs which now trigger decreases considerably. Setting the individual PMT thresholds to at least 2 photoelectrons causes the ^{39}Ar background to decrease exclusively to fewer than 5 PMTs. This means no single ^{39}Ar event produced sufficient light above the threshold (2γ) on more than 5 PMTs. Combining both the threshold cut and a cut on the number of PMTs above threshold gives more powerful rejection of the ^{39}Ar . The number of PMTs selected for the cut (3) is more of a tuning decision, and is investigated further in Section 4.3.

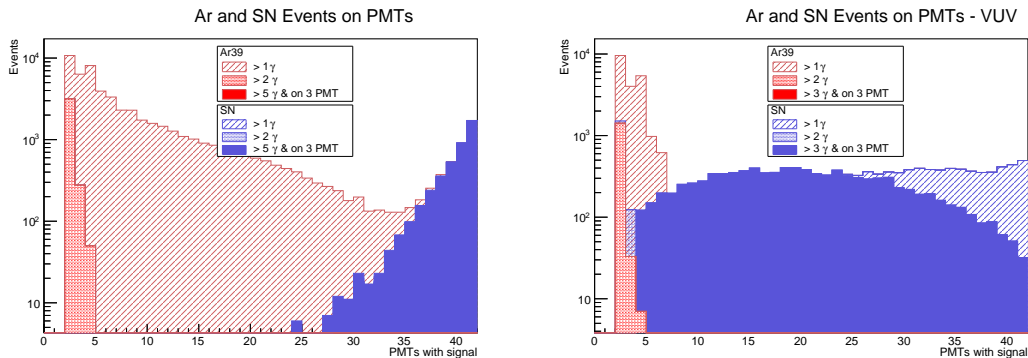


Figure 13: Number of PMTs seeing signal for a scintillation event, as an effect of cuts. “ $> 1\gamma$ ” indicates the threshold placed on each individual PMT and the number of PMTs is a multiplicity cut. Left: supernova spectrum remains unchanged for full foils case, where quickly ^{39}Ar events become limited to events on fewer than 5 PMTs, and eventually disappear. Right: without reflective-foils, some supernova events move from being visible on all PMTs to being visible on fewer, however all remain above maximum cut of 3 PMTs. Bin 0 has been omitted.

Considering the most stringent cut on Figure 13 of either a 3 or 5 PE threshold for the PMTs, as well as a multiplicity requirement of 3 PMTs exceeding the set threshold, the no foils configuration performs with over 95% efficiency. This is compared to the 100% efficiency achieved when the field cage foils are added. As for the ^{39}Ar , whether foils are included or not, roughly 100% of events are rejected. Therefore from these simple studies (no timing, all light collected), including foils seems to increase efficiency and have little change on the purity of the supernova signal.

4.2 Fast Light Cut Efficiencies

In the previous section we used the total integrated signal for an event, potentially including photoelectrons generated several microseconds after the first photoelectron was registered on the PMT. This is of course extremely difficult to implement in a detector. As an intermediate solution to gain insight into how the cuts would perform on a more realistic time scale, only the “fast” light of the single events were cut on. In this case, “fast” light is referring to light which is emitted due to the fast scintillation decay channel of Argon and occurs on the scale of 100 nanoseconds (see Figure 4). The 100 nanosecond time-scale is a result of instrumental effects (). The immediate results of these cuts are visualised on Figure 14, where the same events from Figure 5 have been reduced to only their fast light component.

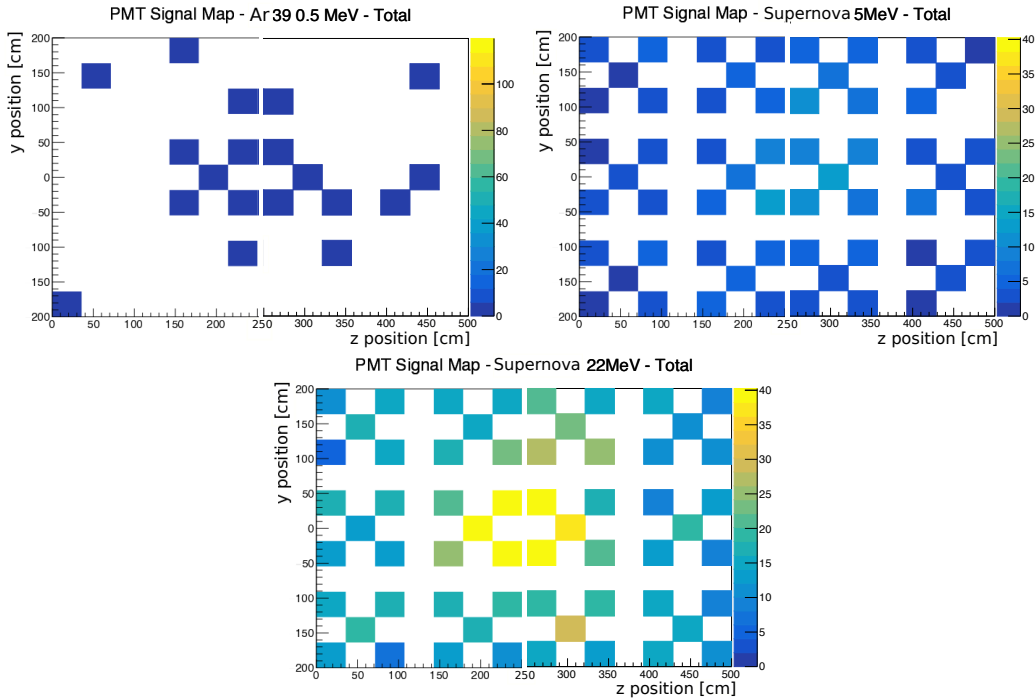


Figure 14: Comparison of single high-energy ^{39}Ar , 5 MeV SN, and 22 MeV SN events in the middle of the detector using 60 PMTs, for the “fast” light case, where the number of photoelectrons is cut by roughly 75%. Shown are the results for the total light (VUV+visible) in the case of full reflective foils. Color (z-axis) corresponds to number of photoelectrons.

Examining only the single ^{39}Ar decay, taking the light emitted in only the first 100 nanoseconds seems to effectively reduce the number of PMTs seeing signal by about 25%. This is to be expected, as the fast light comprises about 25% of the total scintillation light released. From this result, we would largely expect the required cut on the photoelectrons to go down, as the cut on multiple PMTs becomes more effective and the total signal accepted per event is decreased. Table 3 shows the updated cuts and efficiencies when considering only the fast light.

No. PE	No. PMT	% ^{39}Ar	% ^{39}Ar VUV	% SN	% SN VUV
1	1	60.2	84.5	0.01	0.5
2	1	97.3	98.8	0.01	0.8
2	2	73.7	93.2	0.01	5.2
3	2	97.4	99.9	0.01	13.1
3	3	98.0	99.9	0.01	15.2
3	4	99.2	X	0.01	X
4	2	99.8	X	0.01	X
4	3	99.8	X	0.01	X
5	3	99.9	X	0.04	X

Table 3: % Reduction in signal and background compared for the reflective foil and no reflective foils configurations, when considering the fast light. “VUV” indicates the no reflective foil case and “X” indicates that the cut is not relevant for this case.

Previously we explored the effect of the cuts and the reductions of the energy spectra for ^{39}Ar and supernova events, for both the full foil and no foil (vuv only) cases. These are shown as Figures 15 and 16 respectively. The events appear to behave as expected - lower light yield means lower trigger rate. However the smaller amount of light leads some losses to the supernova neutrino signal.

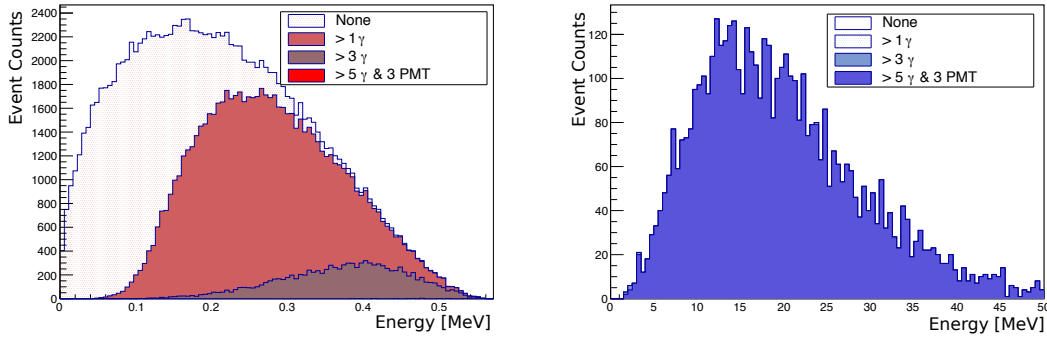


Figure 15: Energy spectra for ^{39}Ar and supernova neutrinos, assuming a large statistics sample for both, to observe how cuts on the number of photoelectrons per a number of PMTs modify the spectra. This case assumes reflective foils, considers the contribution from both direct VUV and reflected visible light, and replicates measuring only the “fast” light component by accepting only photoelectrons measured in the first 100 nanoseconds.

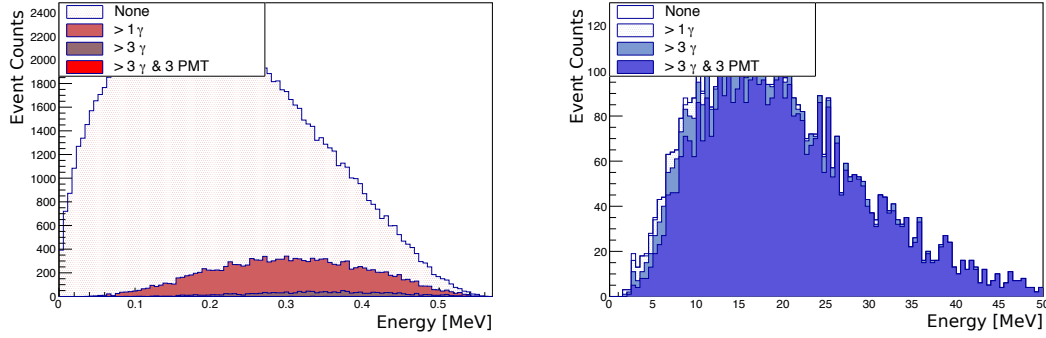


Figure 16: Energy spectra for ^{39}Ar and supernova neutrinos for the no reflective foils configuration, assuming a large statistics sample for both, to observe how cuts on the number of photoelectrons per a number of PMTs modify the spectra. The goal is to completely remove the ^{39}Ar spectrum through these cuts. This is for only the first 100 nanoseconds, to replicate measuring only the “fast” light component.

Although the method of considering the fast light is a step towards a more physical description of the system, ^{39}Ar decays occur with such frequency that the possibility exists for photoelectrons from two or more ^{39}Ar decays to occur on the same PMT at the same time. By allowing the photoelectrons from different decays to coincide, the system is more accurately described, but also requires a slight modification of the cuts previously used.

4.3 Cut efficiencies with full timing simulation

In this section we employ a setup where we take into account the full timing simulation of multiple ^{39}Ar decays and how this modifies the cuts shown in the previous sections. As seen in Figure 7, multiple background decays close in time are capable of causing a trigger when using cuts designed for single events. To prevent this, the trigger should include a timing window during which the threshold must be reached to open a trigger gate. In Section 4.2 we have seen that the supernova neutrino events produce sufficient fast light to trigger the system in this way.

4.3.1 ^{39}Ar cut efficiency

The trigger is now constructed of two stages. The first step being a check of several PMTs seeing a signal above threshold within a small amount of time. If this is fulfilled then a second, longer time window is opened to check for additional photons on each of those PMTs. The number of unique PMTs seeing a signal must occur within 10 ns and each PMTs threshold met or exceeded within 100 ns. Both the PMT thresholds and the number of PMTs were varied to explore the relationship of the ^{39}Ar trigger rate. Because photoelectrons are not tied specifically to a scintillation event, the efficiency of rejection is expressed as a rate - the frequency with which ^{39}Ar decays successfully exceed the thresholds set. The supernova events continue to be considered individually and expressed as an efficiency. Note that after a trigger, the system cannot trigger again until the 100 nanoseconds window for the previous trigger closes.

The performance of the cuts for different LDS configurations are shown in Figure 17, where the colour indicates the intensity of the signal and white bins mean these cut configurations have been excluded or not tested.

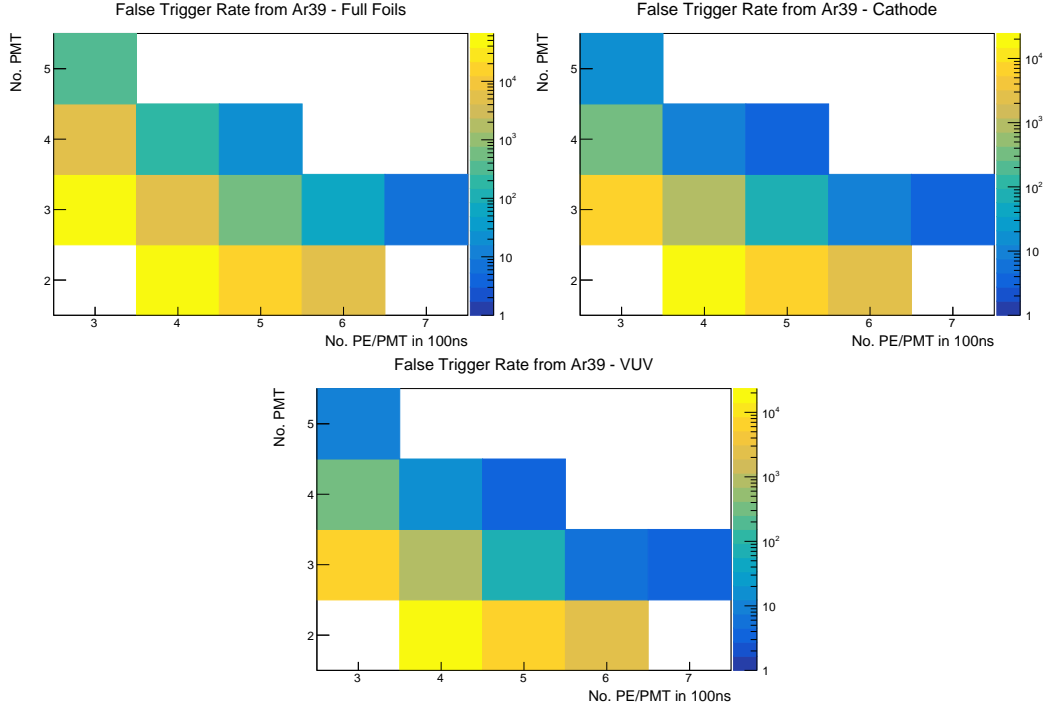


Figure 17: Comparison of effective cuts to reduce ^{39}Ar trigger rate for different LDS configurations. The colour indicates the rate of triggers, and as such should be minimised. Therefore in this case, the most effective cuts are identified by the bins with the lowest value (most blue). Note different colour scales.

Figure 17 shows how simple cuts are able to reduce the background signal with timing included, however the threshold has increased from the single event case. Uniform scaling is not applied across the three plots, as the large range makes distinguishing different colours more difficult. For this reason, attention should be given to the scales to the rights of the plots. The representation shown in Figure 18 allows the different LDS configurations to be compared more easily for a given number of PMTs in 10 nanoseconds. The errors considered here are purely statistical, as the data sample generated creates over 3×10^6 ^{39}Ar decays, properly sampling the energy distribution, as well as the 320,000 voxels of the detector.

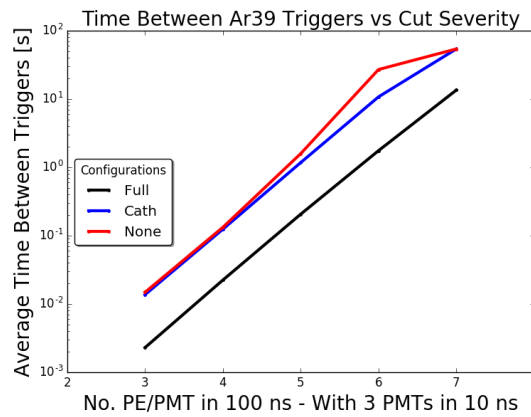


Figure 18: Average time between triggers due to ^{39}Ar photoelectrons as a function of the number of the individual PMT thresholds. The plot shows the rejection ability and optimum threshold for three different light system configurations given a requirement of 3 different PMTs in 10 nanoseconds.

We also consider the case where the PMT multiplicity cut is increased and the PMT thresholds held constant. In Figure 19 we notice that the average time between triggers increases, as expected from Figure 17. At a sufficiently high PMT multiplicity for the shown 2 or 3 photoelectrons per PMT cases, we can achieve a comparable level of rejection to that of Figure 18. While the 2 photoelectron threshold

remains an effective option for the cathode and no foils configurations, the average time between triggers for the full foil configuration is likely too short. Therefore the full foil configuration would benefit most from requiring a 3 photoelectron per PMT threshold, increasing the time between triggers by two orders of magnitude.

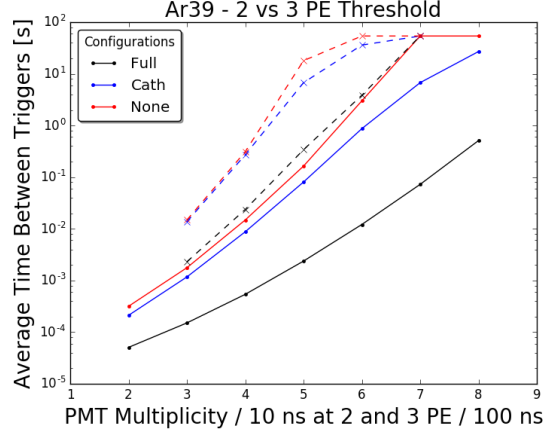


Figure 19: Comparison of PMT multiplicity cuts to increase average time between triggers induced by ^{39}Ar . Dotted lines represent the case where the PMT thresholds are each set to 3 photoelectrons, and the solid lines 2 photoelectrons. The colour of the line indicates which LDS configuration.

4.3.2 ^{39}Ar cut efficiency - expanded cut time window

The effect of increasing the time from 10 to 100 nanoseconds was also tested, as one of the strictest elements of the cuts are that multiple PMTs need to see a photoelectron within 10 nanoseconds of each other. We expect more ^{39}Ar events to pass the cuts, but also potentially allows the later-arriving visible light to contribute more. Indeed in Figure 20 we see that the overall rates have increased for most of the cut configurations, however the higher threshold values selected before continue to remove the background to rates below 1 Hz.

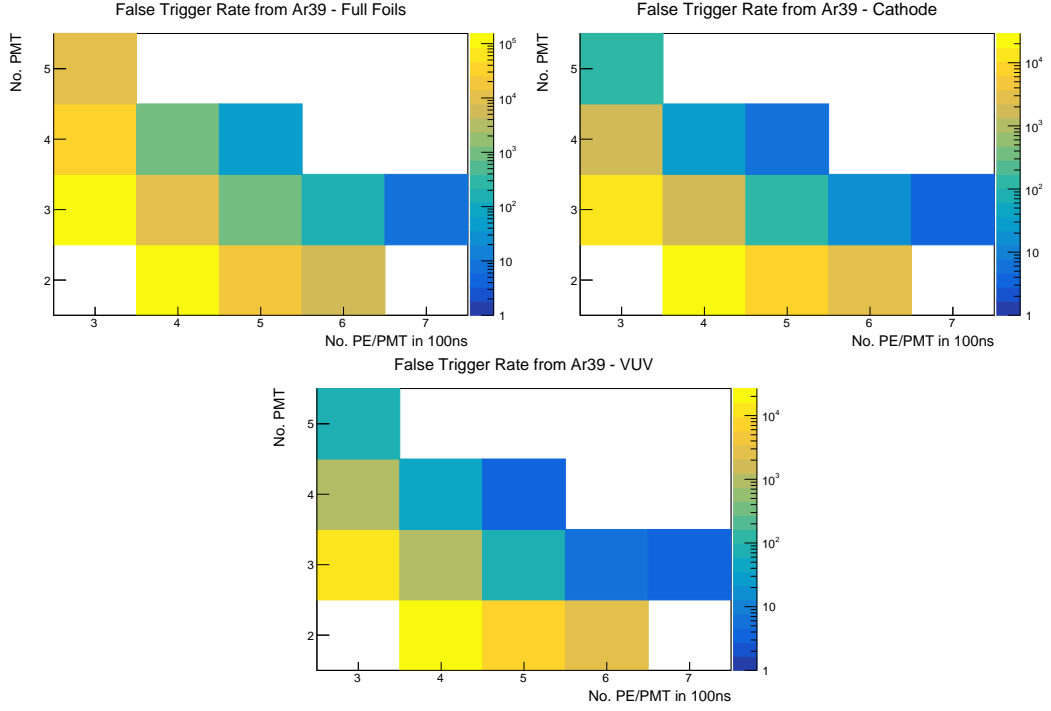


Figure 20: Comparison of effective cuts to reduce ^{39}Ar trigger rate for different LDS configurations. Time between PMTs seeing signal increased from 10 to 100 ns. The colour indicates the rate of triggers, and as such should be minimised. Therefore in this case, the most effective cuts are identified by the bins with the lowest value (most blue). Note that the plots are not normalised to one another.

Figure 21 takes the 3 PMT cut for both 10 and 100 nanoseconds to compare how increasing the PEs removes the background. The 100 nanosecond case (dashed) is always greater than the respective 10 nanosecond case (solid), aside from a threshold of 7 for the full foils configuration. These values are within 0.0185 Hz of each other and is within any expected error.

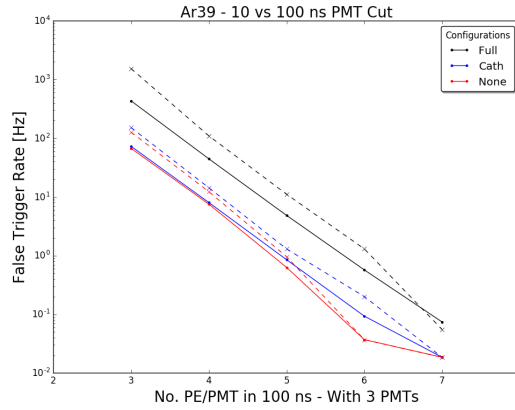


Figure 21: False trigger rate from ^{39}Ar given three different LDS configurations, where solid lines indicate a 10 nanosecond cut on 3 PMTs, while the dashed line has a time window of 100 nanoseconds.

4.3.3 Supernova cut efficiency

The second consideration when selecting the trigger threshold is to configure the threshold such that the supernova ν signal efficiency is kept as high as reasonably possible. For the same cuts demonstrated in Figure 17, the % losses of supernova events are shown in Figure 22.

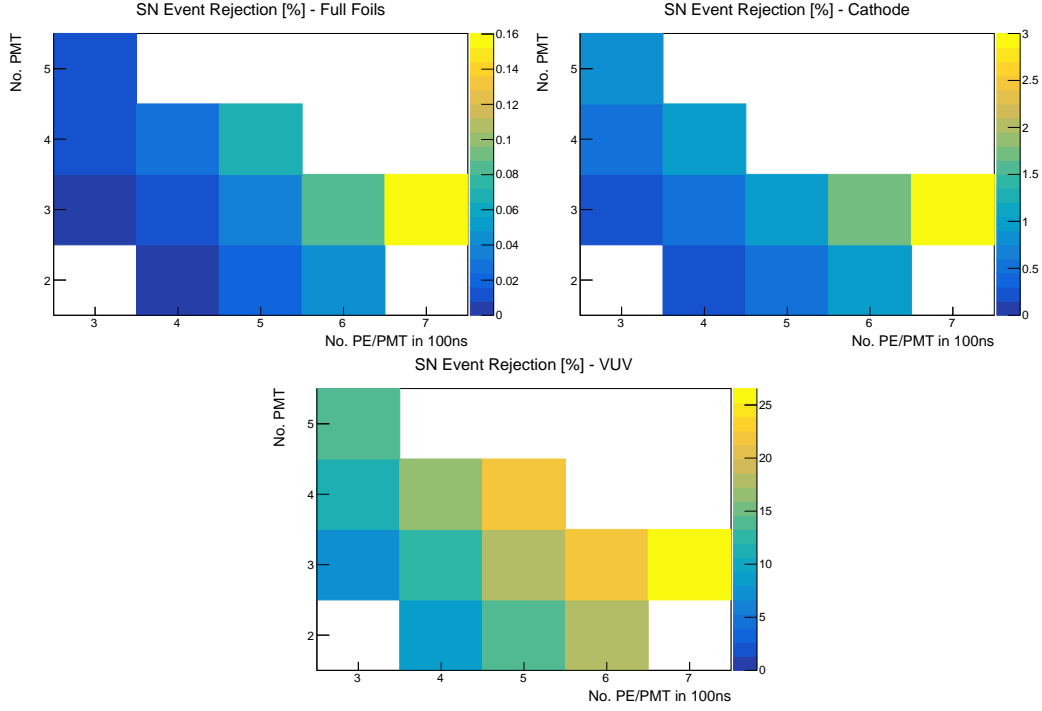


Figure 22: Comparison of thresholds on supernova detection efficiency for different LDS configurations. Colour indicates 1 - Efficiency, and as such should be minimised. Therefore in this case, the most effective cuts are identified by the bins with the lowest value (most blue). Note that the plots are not normalised to one another.

The results from Figure 22 echo results from before, where increasing the threshold or number of PMTs required to trigger begins rejecting supernova events along with the ^{39}Ar . As the plots are not normalised across the different LDS configurations, it is therefore worth noting how little an impact changing thresholds have on the supernova signal given the full foils case - which sees maximum losses of around 0.16%. Comparatively the cathode case has a maximum loss of around 4% and the VUV case up to 27%. Considering the 3 PMT case, a comparison between the full foils, cathode foils, and no foils case for the supernova efficiency is shown in Figure 23.

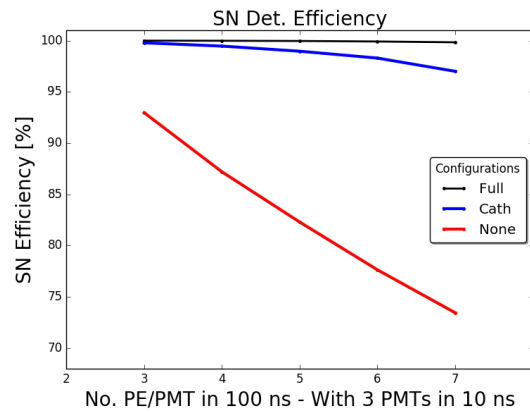


Figure 23: Efficiency of threshold detecting photoelectrons produced from a supernova neutrino event as a function of the number of photoelectrons required on a PMT, given a 3 PMT case.

Much like in the single event case, the higher light yield systems perform better in terms of efficiency, despite having a higher background photon rate. The cathode and full foil configuration efficiencies remain above 90% even for higher cuts, while the efficiency dips for the no foil case. This is due to the lack of uniformity in the LDS, such that events produced further from the photocathode have very few photons detected.

As for the ^{39}Ar , we can consider how the supernova detection efficiency is modified by increasing the PMT multiplicity instead of the PMT thresholds. Figure 24 displays that also for this case where the PMT multiplicity is increased, the efficiencies for both the full and cathode foils cases do not decrease below 95% (for either a 2 or 3 photoelectron threshold). The no foils configuration however, has a roughly 5% higher efficiency at 3 photoelectron threshold, and almost 10% at 2 photoelectrons.

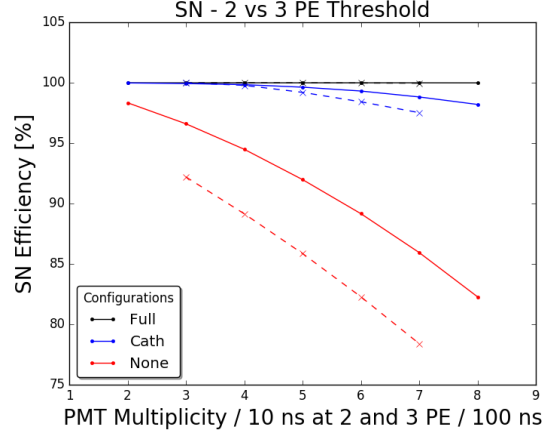


Figure 24: Efficiency of detecting supernova neutrino given a PMT multiplicity cut at a selected PMT threshold. Dotted lines correspond to a threshold of 3 photoelectrons per PMT and solid lines 2 photoelectrons per PMT. The colours of the lines represent different light detection system configurations.

While using a lower threshold per PMT increases the resolution of the system, having the PMT multiplicity cut with the 10 nanosecond time window could make the cut more sensitive to electronics jitters. To make this cut more robust in the future, the results should also be investigated when the timing window is widened to 50 or 100 ns.

Considering both results from the ^{39}Ar study and the supernova study, it clearly displays that the cuts oppositely impact the desirable outcome. Assuming that equal weight can be given to the supernova efficiency, and rejecting the ^{39}Ar background, then the sets of Figures 17 and 22 can be considered together to find the optimal bin for minimising the background and maximising the signal.

Results are shown on a single plot for the 3 PMT case where the detection efficiency for supernova ν and the inverse of the trigger rate (average time between triggers) are shown in Figure 25, where the solid line is for the left axis and the dashed line the right axis. The colour scheme is as before with black being the full foils configuration, blue cathode foils, and red no foils. As was shown in the previous plots, when increasing the cuts both the supernova ν detection efficiency decreases and the average time between triggers increases.

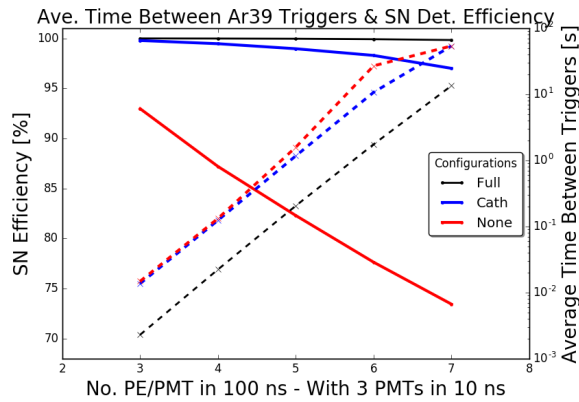


Figure 25: For the 3 PMT case, the supernova ν trigger efficiency (solid lines) and the time between triggers (dashed lines) are displayed simultaneously with their respective trends for increasing the number of photoelectrons per PMT needed to trigger.

As the supernova efficiency for events in the full foils configuration is consistently better than 99% even at the highest thresholds, this configuration provides several values where the background triggers are nearly zero. This is similar to the cathode configuration, where losses in efficiency are rather small - up to 3%. Comparatively the VUV only case, where the losses in efficiency are much higher than cases where reflective foils are included, up to 27%. Despite this, the most effective thresholds are the highest, as this effectively reduces the background signal to zero, with minimum efficiency across all configurations at 73%. The 7 PE/PMT see the lowest background, with 5 PMT and 3 PE/PMT performing only slightly worse. The final trigger value will most likely be influenced by the capacities of the DAQ system and as such will need to be revisited in the future.

5 Conclusion

We have used the simulation to study light system responses in SBND to a constant background of low energy ionising events generated by ^{39}Ar , which could interfere with acquisition of important low energy events, such as supernova neutrinos. Despite the background having a high frequency compared to supernova neutrino events, employing simple cuts based on the arrival times of photons isolated the supernova neutrino events with high efficiency. If complete rejection of the background is the priority, the system lacking reflective foils sees the supernova neutrino detection efficiency drop below 80%. The higher light yield configurations, where some reflective foils are implemented, see significantly better efficiency. Concluding that including TPB coated reflective foils improve overall the detection of low energy events.

References

- [1] P. Benetti *et al.* *Measurement of the specific activity of ^{39}Ar in natural argon.* Nucl. Instrum. Methods A, 2008. arXiv:astro-ph/060313v2
- [2] T. Heindl *et al.* *Table-top setup for investigating the scintillation properties of liquid argon.* JINST, 2011.
- [3] Irene Tamborra *et al.* *High-resolution supernova neutrino spectra represented by a simple fit.* Phys.Rev. D 86, 2012.
- [4] Diego Garcia-Gamez *et al.* *Light detection system simulations for SNBD.* SBND DocDB 1155, 2016.