

THE INFLUENCE OF SPIRAL ARMS ON ACTION-BASED DYNAMICAL MODELLING FOR THE MILKY WAY DISK

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ABSTRACT

RoadMapping is a well-tested and robust dynamical modelling machinery developed to constrain the Milky Way's (MW) gravitational potential by simultaneously fitting an axisymmetric parametrized potential model and an axisymmetric action-based stellar orbit distribution function (DF) to the discrete 6D phase-space measurements from stars in the Galactic disk. In this work we investigate if *RoadMapping*'s robustness is still given for data affected by non-axisymmetric spiral arms for which the axisymmetric modelling cannot account. We do this by drawing data from a N-body simulation snapshot of a MW-like spiral galaxy. We find that *RoadMapping* gives surprisingly successful potential constraints even in the presence of strong spiral arms and even when using the most simple assumption for an action-based DF, the quasi-isothermal DF (qDF). For data from large survey volumes *RoadMapping* finds models that average successfully over the spiral arms. As was to be expected the recovered models are slightly biased by the excess of stars in the spiral arms that feel stronger vertical gravitational forces. For all and even small survey volumes with radius $r_{\max} > 500$ kpc around the Sun the best fit model always reproduces the forces at the positions of the stars that entered the analysis very well. Survey volumes with $r_{\max} \leq 3$ kpc give already predictions for the overall gravitational potential as good as those from larger volumes. However, only volumes as large as $r_{\max} = 5$ kpc are able to recover the correct dark matter halo scale length. While small volumes ($r_{\max} < 2$ kpc) centered on spiral arms do not have a large predictive power, inter-arm regions appear smooth and axisymmetric enough such that the potentials derived from small volumes centered here can be extrapolated to regions outside of the survey volume more reliably. The stellar phase-space measurements from the upcoming Gaia data releases will cover a large proportion of the MW disk. As all survey volumes give good potential constraints where most of the stars are located, *RoadMapping* should not only be able to find a good average potential model for the MW from Gaia data. By breaking up and modelling the data in smaller volumes independently, *RoadMapping* will also be able to make correct measurements of non-axisymmetric structures in the MW's gravitational potential. [TO DO: Too long. Cut out 110 words.]

Keywords: Galaxy: disk — Galaxy: fundamental parameters — Galaxy: kinematics and dynamics — Galaxy: structure — [TO DO]

1. INTRODUCTION

The first step in learning more about the Milky Way's (MW) overall gravitational potential and orbit distribution function (DF)—which are crucial for understanding galaxy structure and formation—is to find the “best possible” axisymmetric model for the Galaxy. Given such a model the identification and characterization of non-axisymmetries like the bar, spiral arms or stellar streams in stellar phase-space (and chemical abundance) data would then become more straightforward.

Several approaches to constrain an axisymmetric potential and/or orbit DF have recently been put forward: Bovy & Rix (2013) and Piffl et al. (2014) fitted potential and DF simultaneously to stellar kinematics in the disk and got precise constraints on the overall potential; Sanders & Binney (2015) and ? investigated extended DFs for the disk and halo respectively (given a fiducial potential), that included in addition to the distribution in orbit space also metallicity of the star.

In this work we will continue our investigation of the *RoadMapping* approach (“Recovery of the Orbit Action Distribution of Mono-Abundance Populations and Potential Inference for our Galaxy”). The first application of

RoadMapping was done by Bovy & Rix (2013) and ?, hereafter Paper I, performed a detailed analysis of the strengths and weaknesses of the approach.

The idea behind *RoadMapping* is, that simple stellar populations in the MW disk—be it mono-abundance populations (??) TO DO (i.e., stars with the same [Fe/H] and $[\alpha/\mathrm{H}]$ Check) or maybe also mono-age populations [References: Maybe Martig??? Any Bovy reference about age???]—follow simple orbit DFs, like, e.g., the quasi-isothermal DF (qDF) by ?. Given an assumed gravitational potential one can calculate the orbits—or specifically the orbit's actions $\mathbf{J} = (J_R, J_\phi = L_z, J_z)$, which are integrals of motions—from the star's current phase-space position (\mathbf{x}, \mathbf{v}) . Only if the assumed gravitational potential is realistic, this orbital action distribution will follow a realistic orbit DF like the qDF. This allows one to simultaneously fit potential and orbit DF to observations.

Bovy & Rix (2013) employed this approach to measure the Milky Way's surface density profile within 1.1 kpc using 43 MAPs in the Galactic disk from the SEGUE survey Check Reference. Their potential model had only two Check free parameters (disk scale length and halo contribution to the radial force at the solar radius). To account for missing model flexibility they constrained the surface density for each MAP only at one best radius. The profile they derived in this fashion had a scale length of $R_s = 2.5$ kpc and was—in the regime $R > 6$ kpc

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Check—later confirmed by Piffl et al. (2014) using a different action-based procedure.

Given the success of this first application and in anticipation of the upcoming data releases from Gaia in 2016–2022 **Check**, Trick et al. 2016 improved the *RoadMapping* machinery and studied its strengths and breakdowns in detail, by investigating a large suite of mock data sets. Under the prerequisite of axisymmetric data and model, they found that *RoadMapping*'s modelling success is stable against minor misjudgements of DF or selection function, and that—if the true potential is not contained in the proposed family of model potentials—one can still find a good fit, given the limitations of the model. Paper I also found that measurement uncertainties of the order of those by the final Gaia data release should be good enough (within 3 kpc from the Sun) to allow for precise and unbiased modelling results.

[TO DO: Continue]

What was not investigated in Paper I, was the breakdown of axisymmetry in the data and what would happen if both axisymmetric distribution function and potential models would therefore not contain the true DF and potential anymore. This will be this work's subject of investigation.

Here we will specifically investigate the important question, if axisymmetric *RoadMapping* modelling can still give reliable constraints on the potential in the presence of non-axisymmetric spiral arms in the data.

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- Main question: Does axisymmetric *RoadMapping* modelling work in the presence of non-axisymmetric spiral arms?
- Consequences: Both potential and orbit DF are not axisymmetric, i.e., the fitted axisymmetric potential model and DF do per se not contain the truth.
- How to approach this: Use simulation by D'Onghia et al. 2013 and apply RM to it
- The potential model we use is chosen mostly for practical reasons and is not necessarily the optimal one for the simulation. Also, we use a single qDF as DF – because it is the simplest thing to do. Also independently of the non-axisymmetries the chosen models might deviate from the truth. Where we investigated deviations between model and truth in isolated test cases, here several assumptions break down simultaneously.
- Explain actions very shortly. $\mathbf{J} = (J_R, J_\phi = L_z, J_z)$ quantify oscillation in the coordinate directions (R, ϕ, z) . Are calculated from current phase-space position in a given potential Φ .
- Say that actions are conserved in an axisymmetric potential, but not in non-axisymmetric potentials. (Maybe the mean vertical action is conserved [TO DO: Reference].) It is therefore important to check, if our modelling works in a system where actions are not conserved.

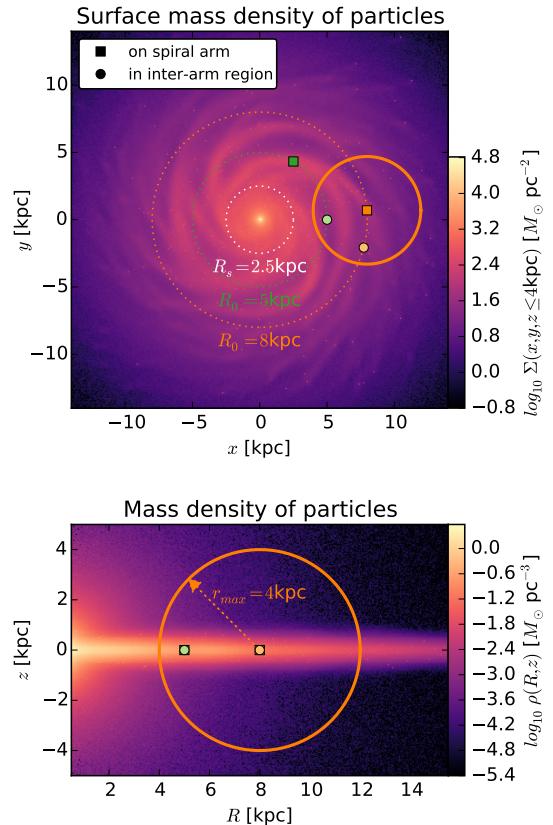


Figure 1. Simulation snapshot by D'Onghia et al. (2013). Shown are the surface mass density (in (x, y) plane, upper panel) and mass density (in (R, z) plane, lower panel) of the “star” particles belonging to disk, bulge and giant molecular clouds. (The dark matter halo in this simulation is static and analytic and not shown here.) Overplotted are the disk's scale length $R_s = 2.5$ kpc (see Section 2.1) and the radii at which we center our test survey volumes in this investigation, $R = 8$ and 5 kpc. The centers of the different survey volumes are marked with a square, if the survey volume is centered on a spiral arm, or with a circle, if the volume is centered on an inter-arm region. The yellow circle with radius $r_{\max} = 4$ kpc marks the survey volume in which we conduct the analysis discussed in detail in Section 4.1. [TO DO: overplot names of positions.]

2. DATA FROM A GALAXY SIMULATION

2.1. Description of the galaxy simulation snapshot

In this work we use stellar phase-space data drawn from high-resolution N-body simulation snapshot of a disk galaxy by D'Onghia et al. (2013) carried out with GADGET-3 code described in Springel et al. (2005), in which overdensities with properties similar to giant molecular clouds induced prominent spiral arms—and therefore non-axisymmetric sub-structure—via the swing amplification mechanism.

For details see D'Onghia et al. (2013), here we summarize the essential characteristics.

The simulation has a gravitationally evolving stellar disk within a static/rigid analytic dark matter halo.

The analytic halo follows a Hernquist (1990) profile

$$\rho_{\text{dm}}(r) = \frac{M_{\text{dm}}}{2\pi} \frac{a_{\text{dm}}}{r(r + a_{\text{dm}})^3} \quad (1)$$

with total halo mass $M_{\text{dm}} = 9.5 \cdot 10^{11} M_\odot$ and scale length $a_{\text{dm}} = 29$ kpc (Elena D'Onghia, private communication). [There is no way to check that this is indeed the case. What should I do?] [What does "total mass

computed at 160 kpc mean? What does concentration $c=9$ mean for a Hernquist halo? How does that fit together with $a=29$ kpc?

The disk consists of 10^8 “disk star” particles, each having a mass of $\sim 370 M_\odot$, and 1000 “giant molecular cloud” particles with mass $\sim 9.5 \cdot 10^5 M_\odot$. Initially the particles are distributed following an exponential disk profile with density

$$\rho_*(R, z) = \frac{M_*}{4\pi z_0 R_s^2} \operatorname{sech}^2\left(\frac{z}{z_0}\right) \exp\left(-\frac{R}{R_s}\right),$$

with $R_s = 2.5$ kpc and $z_0 = 0.1R_s$ [TO DO: Check in my own measurements] and total disk mass $M_* = 0.04 \cdot M_{\text{dm}} = 3.8 \cdot 10^{10} M_\odot$.

The bulge consists of 10^7 “bulge star” particles with mass $\sim 950 M_\odot$ and they are distributed following a spherical Hernquist profile analogous to Equation (1), with total mass $M_{\text{bulge}} = 0.01 \cdot M_{\text{dm}} = 9.5 \cdot 10^9 M_\odot$ and scale length $a_{\text{bulge}} = 0.1 \cdot R_s = 0.25$ kpc.

The simulation snapshot which we are using in this work has evolved under its own gravity for ~ 250 Myr, which corresponds to approximately one orbital period at $R \sim 8$ kpc. The mass density of simulation particles (without the DM halo) at this snapshot time is shown in Figure 1. Pronounced spiral arms have developed due to the “molecular cloud perturbers”, which can be seen in Figure 1 as small overdensities in the disk. The spherical bulge and very flattened disk are shown in the lower panel in Figure 1.

We have confirmed that the gravitational center of the particles corresponds to the coordinate origin.

2.2. Survey volume and data

The selection function of all-sky surveys like Gaia, that are only limited by the brightness of the tracers, are contiguous and—when ignoring anisotropic effects like dust obscuration—spherical in shape. For simplicity we will use spherical survey volumes centered on different vantage points and with sharp edges at a radius r_{max} around it (see also Section 3.1), which corresponds to a magnitude cut for stellar tracers all having the same luminosity. Figure 1 illustrates the different survey volume positions we use in this study: We use each a volume with $r_{\text{max}} = 1, 2, 3, 4$ or 5 kpc centered on a spiral arm and on an inter-arm region at both the equivalent of the solar radius in this simulation, $R = 8$ kpc, and at $R = 5$ kpc, where the spiral arms are more pronounced (see Figure 2) than at $R = 8$ kpc. From each volume we draw $N_* = 20,000$ random “disk star” particles from the simulation and use their phase-space positions ($\mathbf{x}_i, \mathbf{v}_i$) within the simulated galaxy’s restframe as data. To make the data sample more realistic, one would actually have to add measurement uncertainties, especially to the distances of the survey volumes central vantage point and the proper motions measured from there. We decided not to include measurement uncertainties: Firstly, their effect on *RoadMapping* modelling has been already investigated in Paper I and we found that the measurement uncertainties of the last data release of Gaia should be small enough to not disturb the modelling to much. Secondly, in this study we want to isolate and investigate the deviations of the data from the assumed potential and DF model and axisymmetry independently of other

Table 1
Best fit parameters of the DEHH-Pot (see Section [TO DO]), which we use as the global best fit symmetrized potential model.

circular velocity	$v_{\text{circ}}(R_\odot)$	222 km s ⁻¹
disk scale length	h_r	2.5 kpc
disk scale height	h_z	0.17 kpc
halo fraction	f_{halo}	0.54
halo scale length	a_{halo}	29 kpc
bulge mass	M_{bulge}	$0.95 \cdot 10^{10} M_\odot$
bulge scale length	a_{bulge}	0.25 kpc

effect.

[TO DO: Add a table that introduces names for the positions S8, S5, I8, I5, and lists their coordinates.]

2.3. Symmetrized potential model

For a galaxy with pronounced spiral arms an axisymmetric model matter distribution can per se not reproduce the true matter distribution globally. We derived an *overall best fit symmetrized* potential model from the distribution of particles to be able (i) to quantify the non-axisymmetries in the simulation snapshot better and (ii) to compare how close our axisymmetric *RoadMapping* results can get to it.

We derive this model by fitting axisymmetric analytical functions to the density distribution of each of the galaxy component’s particles: The bulge and halo follow each a Hernquist profile by construction (see Section [TO DO]). The disk with its spiral arms however does deviate strongly from its initial conditions in Equation ???. We chose the double exponential disk model from the *galpy* (galaxy python) library to fit the particle distribution in the disk. (This profile is better than the Miyamoto-Nagai disk in reproducing the overall radial density slope, but as it has not the advantage of allowing fast force calculations, we are not using it in the *RoadMapping* modelling.) The best fit parameters for this reference potential, to which we will refer as the DEHH-Potential (Double-Exponential disk + Hernquist halo + Hernquist bulge) in the remainder in this work, are given in Table 1. As can be seen in Figures [TO DO] in Section [TO DO] below, the DEHH-Pot fits the overall density distribution very well. Its density profile might be a little stepper around $z \sim 0$ than the actual particle distribution, but this should not affect the overall discussion, as the radial density, surface density and disk-to-halo ratio profiles are so well reproduced.

Globally our *RoadMapping* model can not be better than this *overall best fit symmetrized* DEHH-Pot, which acts therefore like a global upper limit to *RoadMapping*. We note that in deriving the DEHH-Pot we used the correct decomposition into bulge, disk and halo components, while *RoadMapping* only feels the composite gravitational effect. Given the flexibility of the potential model MNHH-Pot used in *RoadMapping* and the restriction of the modelling to a survey volume, it could however be possible that *RoadMapping* can do better in fitting the local potential inside the survey volume at the price of recovering the global potential parameters to less accuracy. We will discuss this in the following sections.

2.4. Quantifying influence of spiral arm

[TO DO: There is a short comment on that in D'Onghia 2013 as well]

[TO DO: Maybe give the different survey centroids names I8,S8,I5,S5, so it's easier to talk about them.]

[TO DO: Rewrite this section to actually reflect what I'm doing.]

Depending on size and position of the survey volume the spiral arms and inter-arm regions dominate the stellar distribution within the survey volume to different degrees. To quantify the strength of the spiral arm, we introduce the quantity

$$\kappa(x_j, y_j) \equiv \frac{\Sigma_{\text{disk},T}(x_j, y_j \mid |z| \leq 1.5 \text{ kpc})}{\Sigma_{\text{disk},S}(x_j, y_j \mid |z| \leq 1.5 \text{ kpc})} - 1 \quad (2)$$

where $\Sigma_{\text{disk},i}$ is the surface density of the disk component of the simulation snapshot ($i = T$) or of the symmetrized snapshot model DEHH-Pot in Section 2.3 ($i = S$) within the survey volume,

$$\Sigma_{\text{disk},i} \equiv \int_{-1.5 \text{ kpc}}^{1.5 \text{ kpc}} \rho_{\text{disk},i}(x_j, y_j, z) \, dz. \quad (3)$$

(x_j, y_j) is the centroid of a surface area element $(x_j \pm \delta, y_j \pm \delta)$ with $\delta = 0.25 \text{ kpc}$. $(x_c, y_c, z_c = 0)$ is the position of the survey volume's center within the simulation's cartesian coordinate system, $c \in \{A, B, C, D\}$. We consider all $n \simeq \pi r_{\max}^2 / 4\delta^2$ values of $\kappa(x_j, y_j)$ inside a given survey volume and calculate the statistics

$$\langle \kappa(r \leq r_{\max} \mid c) \rangle \equiv \frac{1}{n} \sum_j^n \kappa(x_j, y_j)$$

and $\Delta\kappa(r \leq r_{\max} \mid c) \equiv \max[\kappa(r \leq r_{\max} \mid c)] - \min(r \leq r_{\max} \mid c)$

$$\text{with } (x_j - x_c)^2 + (y_j - y_c)^2 \leq r_{\max}.$$

This gives us information if a spiral arm or an inter-arm region dominates the survey volume ($\langle \kappa \rangle > 0$ for spiral arms, $\langle \kappa \rangle < 0$ for inter-arm regions) and how large the contrast between spiral arms and inter-arm regions is ($\Delta\kappa$).

3. ROADMAPPING MODELLING

3.1. Likelihood

The data that goes into the modelling are the 6D position and velocity coordinates $(\mathbf{x}_i, \mathbf{v}_i)$ of N_* stars within the survey volume. For simplicity we use a purely spatial selection function $\text{sf}(\mathbf{x})$ of spherical shape,

$$\text{sf}(\mathbf{x}) \equiv \begin{cases} 1 & \text{if } |\mathbf{x} - \mathbf{x}_0| \leq r_{\max}, \\ 0 & \text{otherwise} \end{cases},$$

whose maximum radius r_{\max} defines the boundary of the survey volume and which is centred on $\mathbf{x}_0 \equiv (R_0, \phi_0, z_0 = 0)$. Given a parametrized potential model $\Phi(R, z)$ with parameters p_Φ , the i -th star is on an orbit characterized by the orbital actions

$$\mathbf{J}_i \equiv \mathbf{J}[\mathbf{x}_i, \mathbf{v}_i \mid p_\Phi].$$

The probability of stars to be on the orbit \mathbf{J}_i is proportional to a given orbit distribution function $\text{df}(\mathbf{J})$ with parameters p_{DF} ,

$$\text{df}(\mathbf{J}_i \mid p_{\text{DF}}) \equiv \text{df}(\mathbf{J}[\mathbf{x}_i, \mathbf{v}_i \mid p_\Phi] \mid p_{\text{DF}}) \equiv \text{df}(\mathbf{x}_i, \mathbf{v}_i \mid p_\Phi, p_{\text{DF}}),$$

where the latter equivalence arises from the Jacobian determinant between the angle-action coordinates $(\boldsymbol{\theta}, \mathbf{J})$ and cartesian phase-space coordinates (\mathbf{x}, \mathbf{v}) , which is $|\partial(\mathbf{x}, \mathbf{v})/\partial(\boldsymbol{\theta}, \mathbf{J})| = 1$ and therefore allows us to treat the df equivalently as a distribution of current phase-space coordinates or a distribution of orbital actions only, with uniform distribution in the angles $\boldsymbol{\theta}$. In some sense, the $\text{df}(\mathbf{J})$ describes how we expect a realistic stellar population in the MW disk to look like.

The joint likelihood of the i -th star being on an orbit \mathbf{J} in the potential Φ and being within the survey volume is therefore

$$\mathcal{L}_i \equiv \mathcal{L}(\mathbf{x}_i, \mathbf{v}_i) = \frac{\text{df}(\mathbf{x}_i, \mathbf{v}_i \mid p_\Phi, p_{\text{DF}}) \cdot \text{sf}(\mathbf{x}_i)}{\int \text{df}(\mathbf{x}, \mathbf{v} \mid p_\Phi, p_{\text{DF}}) \cdot \text{sf}(\mathbf{x}) \, d^3x \, d^3v}.$$

The details how we evaluate the likelihood normalisation numerically to sufficiently high enough precision are discussed in Paper I.⁴

In the scenario considered in this paper it can happen that there are a few (~ 1 in 20,000) stars entering the catalogue that are for some reason on rather extreme orbits, e.g., moving radially directly towards the center. These kinds of orbits do not belong to the set of orbits that we classically expect to make up a overall smooth galactic disk. To avoid that such single stars with very low likelihood mess up the modelling we introduce here a simple outlier model,

$$\mathcal{L}_i \longrightarrow \max(\mathcal{L}_i, \epsilon \cdot \text{median}(\mathcal{L})),$$

where $\epsilon = 0.001$ for $N_* = 20,000$ stars and $\text{median}(\mathcal{L})$ is the median of all the N_* stellar likelihoods \mathcal{L}_i with the given p_Φ and p_{DF} . This outlier model was not used in Paper I.

Following Paper I, we assume for now uninformative flat priors on the model parameters p_Φ and p_{DF} and find the maximum and width of the posterior probability function

$$\text{pdf}(p_\Phi, p_{\text{DF}} \mid \text{data}) \propto \prod_{i=1}^{N_*} \mathcal{L}_i \cdot \text{prior}(p_\Phi, p_{\text{DF}})$$

using a nested-grid approach and then explore the full shape of the pdf using a Monte Carlo Markov Chain (MCMC)⁵. Full details on this procedure are given in Paper I.

3.2. DF model

The most simple orbit distribution function exhibiting a disk-like structure may be the quasi-isothermal distribution function (qDF) introduced by Binney (2010) and Binney & McMillan (2011). It proved to be a successful model to describe the orbit distribution of individual mono-abundance populations (MAPs) in the Galactic disk (Bovy & Rix 2013; Ting et al. 2013), that seem to be isothermal in z -direction (i.e., “quasi-isothermal”). Modelling approaches trying to capture the overall disk distribution (Piffl et al. 2014; Sanders & Binney 2015) were describing the Galactic disk as a superposition of many qDFs. The qDF, which we already used in Paper

⁴ [TO DO: Write what exact numerical accuracy we use and check that it is actually good enough.]

⁵ [TO DO: Reference emcee]

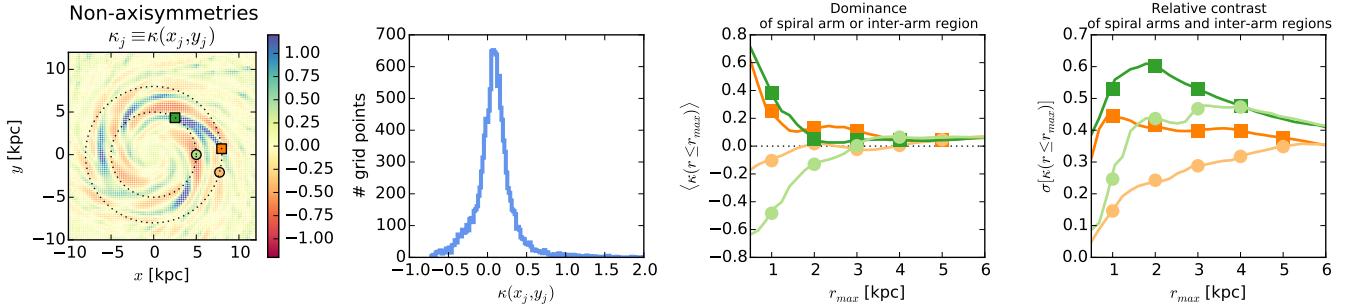


Figure 2. Contrast and dominance of the spiral arms. The left panel shows the values of κ_i calculated according to Equation 2 as described in Section 2.4 at regular grid points (x_i, y_i) with bin width 0.25 kpc. Marked are the centroids of the four test survey volumes of this study. The second panel shows a histogram over all $\kappa(x_j, y_j)$ values in the region $x, y \in [-14, 14]$ kpc. The two right panels show the dominance (with the mean $\langle \kappa(r \leq r_{max}) \rangle$ of all grid points within the survey volume as measure) and the relative contrast (with the standard deviation $\sigma[\kappa(r \leq r_{max})]$ as measure) of spiral arms and inter-arm regions inside a spherical survey volume with given centroid (colour-coded) and size (radius r_{max} shown on the x-axis). We chose two volumes in which the spiral arms dominate, and two in which an inter-arm region dominates. The dominance and contrast of spiral arms and inter-arm regions is stronger at $R = 5$ kpc than at $R = 8$ kpc. Also, inter-arm regions appear larger and smoother than spiral arms, as already inside a small volume centred on a spiral arm the contrast is quite large. The larger the volume the more does the overall effect of spiral arms and inter-arm regions average out. [TO DO: Jo suggested to look into implementing a tanh2 disk for galpy to not have the slight bias in mean(κ) at higher rmax. Maybe it's easy to implement. Maybe it isn't.]

I, has the functional form

$$\begin{aligned} qDF(\mathbf{J} | p_{\text{DF}}) \\ = f_{\sigma_R}(J_R, L_z | p_{\text{DF}}) \times f_{\sigma_z}(J_z, L_z | p_{\text{DF}}) \end{aligned} \quad (4)$$

where

$$\begin{aligned} f_{\sigma_R}(J_R, L_z | p_{\text{DF}}) = n \times \frac{\Omega}{\pi \sigma_R^2(R_g) \kappa} \exp\left(-\frac{\kappa J_R}{\sigma_R^2(R_g)}\right) \\ \times [1 + \tanh(L_z/L_0)] \end{aligned} \quad (5)$$

$$f_{\sigma_z}(J_z, L_z | p_{\text{DF}}) = \frac{\nu}{2\pi \sigma_z^2(R_g)} \exp\left(-\frac{\nu J_z}{\sigma_z^2(R_g)}\right) \quad (6)$$

(Binney & McMillan 2011). The guiding-center radius R_g , circular frequency Ω , radial/epicycle frequency κ and vertical frequency ν describe the near-circular orbit with given angular momentum L_z in a given potential. Counter-rotating orbits with $L_z < L_0$ are suppressed by the term $[1 + \tanh(L_z/L_0)]$ (with $L_0 \sim 10$ km s⁻¹ kpc). We set the radial stellar tracer density $n(R_g)$ and velocity dispersion profiles $\sigma_z(R_g)$ and $\sigma_R(R_g)$ to

$$n(R_g | p_{\text{DF}}) \propto \exp\left(-\frac{R_g}{h_R}\right) \quad (7)$$

$$\sigma_R(R_g | p_{\text{DF}}) = \sigma_{R,0} \times \exp\left(-\frac{R_g - R_\odot}{h_{\sigma,R}}\right) \quad (8)$$

$$\sigma_z(R_g | p_{\text{DF}}) = \sigma_{z,0} \times \exp\left(-\frac{R_g - R_\odot}{h_{\sigma,z}}\right). \quad (9)$$

The free model parameters of the qDF are therefore

$$p_{\text{DF}} \equiv \{\ln h_R, \ln \sigma_{R,0}, \ln \sigma_{z,0}, h_{\sigma,R}, h_{\sigma,z}\}.$$

Even though we do not have any stellar abundance or age information in the simulation snapshot we are going to investigate (see Section 2) and we therefore cannot define stellar sub-populations for which the assumption of such a simple model might be reasonable, we will still try to model the whole disk with a single qDF—to see how far we can get with the simplest possible model.

3.3. Potential model

[TO DO]

- Introduce MNHH potential model, explain that form of disk was mostly chosen to the closed form expression of Φ which allows for fast calculation.
- Mention action calculation and that we tested explicitly that fixing $\Delta=0.45$ and using staecel interpolation grid does not degrade the analysis
- Mention and reference galpy.

4. RESULTS

We run the *RoadMapping* modelling on several data sets drawn from different survey volumes in the simulation snapshot introduced in Section 2.1. In Section 4.1 we will analyze one of these *RoadMapping* analyses in detail, and in Section 4.2 we compare the different results for the different data sets.

4.1. An axisymmetric galaxy model from *RoadMapping*

In this section we will discuss the results of the *RoadMapping* analysis of a single data set. This fiducial data set has $N_* = 20,000$ stars that were drawn from the spherical volume with $r_{\text{max}} = 4$ kpc centered on a spiral arm at the “solar” radius $R = 8$ kpc shown as yellow sphere in Figure 1 (position S8). We use *RoadMapping* to fit a single qDF (see Section 3.2) and the potential model MNHH-Pot consisting of a Miyamoto-Nagai disk, a Hernquist halo and a (fixed) Hernquist bulge introduced in Section 3.3 to it. The best fit parameters are summarized in Table 2. The circular velocity $v_{\text{circ}}(R_\odot)$ and halo fraction f_{halo} are especially well recovered (compare to Table 1).

4.1.1. Recovering the stellar distribution

In Figure 3 we compare the configuration phase-space distribution of the 20,000 stars that entered the *RoadMapping* analysis as data with the best-fit stellar distribution, which is given by the action-based qDF and potential model found with *RoadMapping*. We note that

Table 2
Best-fit MNHH-Pot and qDF parameters as recovered from the *RoadMapping* analysis of a survey volume with $r_{\max} = 4$ kpc centered on a spiral arm at $R_0 = 8$ kpc (position S8).

circular velocity	$v_{\text{circ}}(R_\odot)$	$(223.0 \pm 0.1) \text{ km s}^{-1}$
Miyamoto-Nagai disk scale length	a_{disk}	$(3.62^{+0.06}_{-0.05}) \text{ kpc}$
Miyamoto-Nagai disk scale height	b_{disk}	$(0.26 \pm 0.02) \text{ kpc}$
halo fraction	f_{halo}	(0.53 ± 0.02)
halo scale length	a_{halo}	$(21 \pm 2) \text{ kpc}$
bulge mass	M_{bulge}	$0.95 \cdot 10^{10} M_\odot$ (fixed)
bulge scale length	a_{bulge}	0.25 kpc (fixed)
qDF tracer scale length	h_R	$3.34^{+0.05}_{-0.04} \text{ kpc}$
qDF radial velocity dispersion	$\sigma_{R,0}$	$15.91 \pm 0.08 \text{ km s}^{-1}$
qDF vertical velocity dispersion	$\sigma_{z,0}$	$14.0^{+0.2}_{-0.1} \text{ km s}^{-1}$
qDF radial velocity dispersion scale length	$h_{\sigma,R}$	$(4.6 \pm 0.5) \text{ kpc}$
qDF vertical velocity dispersion scale length	$h_{\sigma,z}$	$(5.65^{+0.06}_{-0.07}) \text{ kpc}$

the spiral arms introduce very strong non-axisymmetries in the data, both in the spatial and the velocity distribution (especially in v_R , where a lot more stars move outside than inside as compared to an axisymmetric model). We therefore compare in Figure 3 the data and fit separately for different spatial regions, $R >$ and < 8 kpc and $\phi <$ and $> 5^\circ$. In the region where the spiral arm dominates (blue) the best fit *RoadMapping* model is actually a very poor model. However, what the model underestimates in the spiral arm, it slightly overestimates in the other regions and is therefore indeed something like a good average model for the overall distribution. Especially the region at $R > 8$ kpc and $\phi < 5^\circ$, where neither the spiral arm nor the inter-arm regions dominate strongly, is very well described by the model. This is remarkable because we had no indication beforehand that a single qDF might be at all a good enough model to describe the overall stellar distribution in the simulation snapshot. But it obviously does—apart from the spiral arm of course, which we had no chance to capture anyway.

4.1.2. Recovering the potential

As shown in the previous section the best fit *RoadMapping* model seems to reproduce the stellar phase-space distribution quite well. But is the corresponding potential close to the true potential?

Figures 4 and 5 compare the true simulation snapshot potential and the axisymmetric DEHH-Pot from Table 1 with the best fit MNHH-Pot from the *RoadMapping* analysis. In particular, Figure 4 illustrates the overall matter density distribution, Figure 5 the rotation curve, surface density profile and disk-to-halo ratio and Figure 6 compares the true and recovered (median) gravitational forces at the position of each star in the data set. In Figure 5 we explicitly show the true simulation snapshot circular velocity both averaged over the whole $\Delta\phi = 2\pi$ (black lines), and over the survey volume's maximum radial extent $\phi_0 \pm \arcsin(r_{\max}/R_0)$ (grey lines).

The recovery of density, surface density, circular velocity curve and disk-to-halo ratio is especially good in the region where most of the stars are located, around $R \sim 6$ kpc and in the plane of the disk. In large regions inside the survey volume, and even outside, the density is recovered to within 15%. The circular velocity curve is even recovered to within 5%, which is also approximately the extent of perturbation that the spiral arms

cause with respect to a smooth rotation curve.

There is however a very small ($< 1.5\%$) underestimation of v_{circ} at larger radii. We suspect that this is a combination of a bias introduced by the spiral arms (see also discussion in Section [TO DO]) and the choice of potential model (as a similar bias shows up in the DEHH-Pot). The overall surface density profile is a bit overestimated ($\sim 15\%$) at smaller radii; the reason is clearly the local spiral arm at $R \sim 6$ kpc with its higher surface density and many stars entering the analysis, which bias the result and which was to be expected. At larger radii ($R \sim 10 - 12$ kpc) the fit of radial local density and surface density profile is less good, which we account to the Miyamoto-Nagai disk having a shallower radial profile than an exponential disk, and not enough stars in this outer regions to give good constraints. The much stronger bias in the disk-to-halo ratio at larger radii is the result of a misjudgement of the halo scale length. As we will see later (see Figure 8) we seem to need an even larger survey volume to have enough radial coverage to constrain the halo scale length properly. But again, where most of the stars are located, our *RoadMapping* model is a very good average model for the true galaxy.

The aspect of the potential to which the stellar orbits are actually sensitive to, are the gravitational forces. In Figure 6 we compare therefore the true force that each star in the data set feels ($F_{R,T}(*_i)$ and $F_{T,M}(*_i)$, calculated as sum of the individual contributions by each particle in the simulation and the analytic dark matter halo) with the force that the *RoadMapping* median model predicts for each star ($F_{R,M}(*_i)$ and $F_{z,M}(*_i)$). We scale the difference between truth and model by a typical radial or vertical force at the given radius for which we use

$$F_{R,\text{typ}}(R) \equiv v_{\text{circ},S}^2(R)/R \quad (10)$$

$$F_{z,\text{typ}}(R) \equiv F_{z,S}(R, z = h_r), \quad (11)$$

where $v_{\text{circ},S}$ and $F_{z,S}$ are the circular velocity and vertical force evaluated in the "true symmetric" DEHH-Pot in Table 1. As typical vertical force at a given radius we use the vertical force evaluated at the scale height $h_z = 0.17$ kpc of the disk, because we found that

$$\begin{aligned} &\langle |F_{z,S}(*_i)|(R) \rangle \\ &\lesssim \langle |F_{z,T}(*_i)|(R) \rangle \\ &\lesssim F_{z,S}(R, z = h_z) \\ &\lesssim \langle J_z(*_i) \times \Omega_z(*_i)/h_r \rangle(R), \end{aligned}$$

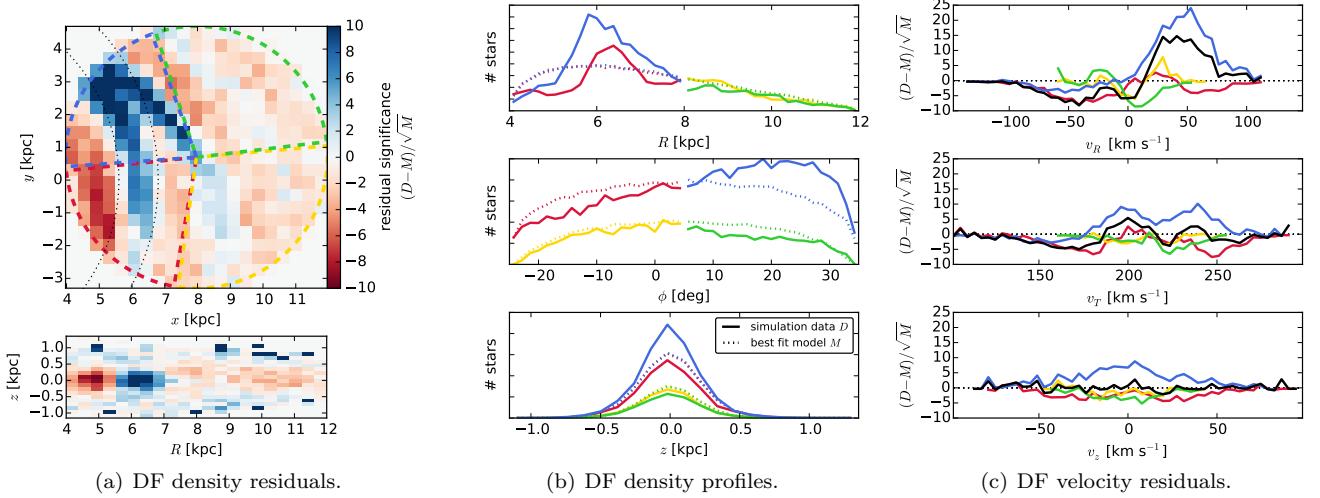


Figure 3. Comparison of the DF of stars in position-velocity space in the data set from the simulation snapshot (with $N_* = 20,000$ and $r_{\max} = 4$ kpc centered on position S8) and in the single-qDF best fit *RoadMapping* model from Table 2. Panel 3(a) shows the spatial density residual significance projected to the (x, y) and (R, z) plane. In the (x, y) panel the following regions are marked: $R < 8$ kpc, $\phi > 5^\circ$ (blue), $R < 8$ kpc, $\phi < 5^\circ$ (red), $R > 8$ kpc, $\phi > 5^\circ$ (green), $R > 8$ kpc, $\phi < 5^\circ$ (yellow). Panels 3(b) and 3(c) show the density profiles and velocity residual significance along each of the 6D phase-space coordinates separately for each of the four spatial regions. The blue region is very much dominated by the non-axisymmetric spiral arm. For the yellow region the axisymmetric single-qDF model is a good description. Overall the qDF is a good average axisymmetric model for the data. (In panel 3(a) we overplot the radii $R_{\text{spiral}} \in [5.6, 6.8]$ kpc as black dotted lines to mark the approximate extent of the stronger spiral arm, to compare it with Figure 7.) [TO DO: Move legend to R panel and make it bigger]

where, e.g., $\langle |F_{z,T}(*_i)|(R)\rangle$ denotes the mean of the absolute value of the true vertical force of all stars at radius R , and $J_z \times \Omega_z \sim \langle E_z \rangle$ is something like the mean vertical energy of a stellar orbit. Figure 6 shows therefore

$$\Delta F_R(*_i) \equiv \frac{|F_{R,M}(R_i, z_i)| - |F_{R,T}(x_i, y_i, z_i)|}{F_{R,\text{typ}}(R_i)} \quad (12)$$

$$\Delta F_z(*_i) \equiv \frac{|F_{z,M}(R_i, z_i)| - |F_{z,T}(x_i, y_i, z_i)|}{F_{z,\text{typ}}(R_i)} \quad (13)$$

$$(14)$$

for each star. The recovery is as expected: The true vertical force is stronger in the spiral arms (due to the higher surface density) and weaker in the inter-arm regions as compared to the axisymmetric best fit model. The radial force is well recovered where the majority of the stars are located, i.e. in the wide inter-arm regions and in the peaks of the spiral arms. Misjudgments happen in the wings of the spiral arms: The true radial force (i.e., the pull towards the galactic center) is stronger at the outer edge/leading side of the spiral arm because of the additional gravitational pull towards the massive spiral arm, and for the same reason weaker at the inner edge/trailing side. Overall the recovered *RoadMapping* model appears to be a good mean model, averaging over spiral arms and inter-arm regions.

4.1.3. Recovering the action distribution

In Sections 4.1.1 and 4.1.2 we have demonstrated the goodness of fit in the configuration space of the data, and of the recovered gravitational potential. What *RoadMapping* is actually fitting is however the distribution in action space. Figure 7 compares the data and the model action distribution given the best fit MNHH-Pot in Table 2. (We only use this axisymmetric potential to estimate the actions as they enter the fit, and do not

calculate the true actions in the true potential.)

We note that the radial and vertical action distribution fits indeed quite well; the axisymmetric model however contains much more stars on close to circular orbits ($J_R \sim 0, J_z \sim 0$) than the simulation does. In the data set there is an excess of stars in the galactic plane ($J_z \sim 0$) that have more eccentric orbits than the axisymmetric model would have. In Figure 3(a) we have marked the radial extent of the stronger spiral arm with dotted lines ($R_{\text{spiral}} \in [5.6, 6.8]$ kpc), and over-plotted the corresponding angular momenta $L_z = R_{\text{spiral}} \times v_{\text{circ}}(R_{\text{spiral}})$ in Figure 7 as a rough estimate where in action space we expect the stars of this spiral arm. It is again obvious that this spiral arm contains (i) more stars in general and (ii) more stars with eccentric orbits ($J_R > 0$) which are (iii) located mostly close to the plane ($J_z \sim 0$), as compared to the axisymmetric model. All of this confirms our expectations for orbits in a spiral arm.

This exercise of comparing the data and the model actions in a best-fit axisymmetric potential should be performed for a future application to data in the Milky Way as well. It could help to learn about the approximate orbits that stars move on in real spiral arms and to subsequently construct orbit distribution functions for spiral arms.

4.2. The influence of spiral arms in *RoadMapping* modelling

In the previous section we showed that for a large enough survey volume ($r_{\max} = 4$ kpc) *RoadMapping* can construct a good average axisymmetric potential (and DF) model for a galaxy with spiral arms. In the following we want to investigate how this modelling success depends on the position and the size of the survey volume within the galaxy and with respect to the spiral arms.

4.2.1. Test suite

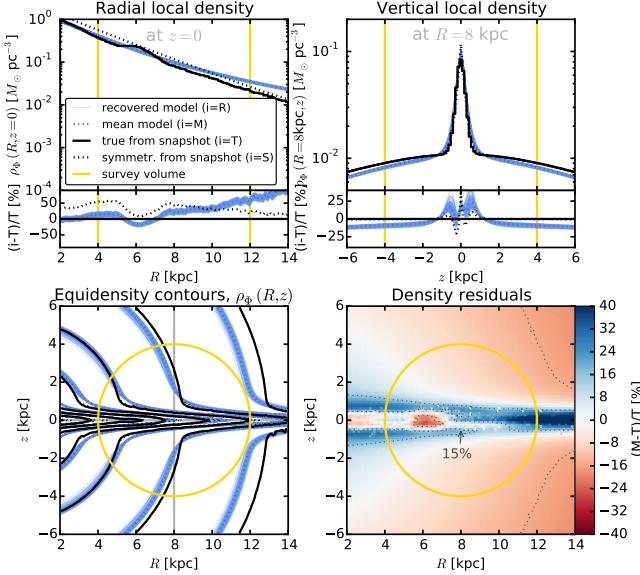


Figure 4. Comparison of the true density distribution $\rho_{\Phi,T}$ in the galaxy simulation snapshot (solid black line, averaged over ϕ) with the axisymmetric density distribution $\rho_{\Phi,R}$ recovered with *RoadMapping* (solid blue lines) from $N_* = 20,000$ stars in the survey volume with $r_{\max} = 4$ kpc (yellow line), as described in Section [TO DO]. The first two panels show density profiles along $(R, z = 0)$ and $(R = 8 \text{ kpc}, z)$, together with the relative differences between true and recovered ρ_{Φ} . The third panel displays equidenity contours of the matter distribution in the (R, z) plane. Overplotted are also the overall axisymmetric DEHH-Pot's $\rho_{\Phi,S}$ (dotted black line) (see Section 2.3) and the $\rho_{\Phi,M}$ of the recovered mean model in Table 2 (dotted blue line). The last panel shows the relative difference between the true $\rho_{\Phi,T}$ (averaged over all ϕ) and the recovered mean model $\rho_{\Phi,M}$. Over wide areas even outside of the survey volume the relative difference is less than 15%. At $R \gtrsim 8 \text{ kpc}$ and $z \sim 0$ it becomes apparent that the chosen potential model cannot perfectly capture the structure of the disk. [TO DO: Maybe show the profiles not at $R=8$ and $z=0$, but at $R=6$ and $z=h_z$ or so.] TO DO: Redo plot. Now I use the median of MCMC samples as best fit model.]

We center our test survey volumes at the positions marked in Figures 1 and 2 (see also Table [TO DO]) and consider volume sizes with $r_{\max} \in [0, 1, 2, 3, 4, 5]$ kpc for $R_0 = 8 \text{ kpc}$ and $r_{\max} \in [0, 1, 2, 3, 4]$ kpc for $R_0 = 5 \text{ kpc}$. As demonstrated in Figure 2 the spiral arm strength is very different in these four different test volumes. Each data set that we draw from the simulation contains a random selection of $N_* = 20,000$ stars inside this spherical volume and we fit a single qDF and MNHH-Pot to it.

4.2.2. Discussion of the model parameter recovery

Figure 8 shows the best fit parameters for the different *RoadMapping* analyses. We also overplot the parameters of the reference DEHH-Pot in Table 1. Overall the statistical random errors on the parameter recovery are very small for $N_* = 20,000$ and possible systematic errors dominate.

Let's first consider the parameters of the gravitational potential: All volumes recover $v_{\text{circ}}(R_\odot)$ within a few km s^{-1} ; in the largest volumes, where the circular velocity curve is probed over several kpc the estimate is the most accurate. The halo fraction f_{halo} of the radial force at the Sun is very well recovered, independent of the size of the volume. The estimate that we get for the best fit Miyamoto-Nagai disk scale height b_{disk} seems to be ap-

proximately independent of the size of the volume as well. We can even recover the true halo scale length a_{disk} , however only for a volume as large as $r_{\max} = 5 \text{ kpc}$. Smaller volumes that underestimate a_{halo} get slightly larger estimates for the disk scale length a_{disk} and the overall radial density slope is then probably closer to the truth, even if the individual parameters are not.

The right column of Figure 8 compares the recovered qDF parameters for the different survey volumes with the qDF parameters we got from fixing the potential model to the expHH-Pot and fitting the qDF only in a $r_{\max} = 5 \text{ kpc}$ volume. Even though the qDF parameters for small volumes are widely different for different positions within the galaxy, they all approach the values recovered with the expHH-Pot for larger volumes. There seems therefore to be indeed an overall best-fit qDF describing the average tracer distribution in the galaxy's disk. The only difference is in the $h_{\sigma,z}$ parameter, where the models with free MNHH-Pot recover a slightly larger value than the models using the known expHH-Pot. The reason is that the Miyamoto-Nagai disk has a different radial profile than the double exponential-disk (see Figure 4), which leads to a less steep radial decline in the vertical forces and therefore mean vertical orbital energies $\langle E_z \rangle \sim \nu \times J_z$ and therefore to a slightly longer $h_{\sigma,z}$ scale length. In general, volumes centered on spiral arms have larger velocity dispersion parameters $\sigma_{R,0}$ and $\sigma_{z,0}$ as compared to volumes at the same radius R but centered on an inter-arm region, and the volumes at $R = 5 \text{ kpc}$ with their stronger spiral arms have larger velocity dispersions than those at $R = 8 \text{ kpc}$ —which is what we expect. Most volumes recover similar tracer scale lengths h_R [TO DO: Value]; only the volume centered on the inter-arm region at $R = 8 \text{ kpc}$ recovers much longer h_R . This might be related to the fact, that most volumes have strong spiral arms at smaller radii, which leads to estimates of shorter tracer scale lengths, while volume I8 is located in a comparably smooth region of the galaxy.

4.2.3. Recovery of the gravitational forces depending on the size of the survey volume

In the previous section we found that the potential and qDF parameters recovered from different survey volumes can differ quite a lot. How good are the corresponding potential constraints in this case? To test this we calculate again $\Delta F_{R,*i}$ and $\Delta F_{z,*i}$ from Equations (12)-(13) at the position of each star $*_i$ in each data set. From the corresponding histograms we derive the median and the 15.87th and 84.13th percentiles and show them in the left panels two of Figure 9. We chose this diagnostic because the forces at the positions of the stars are the quantities of the potential to which our modelling is sensitive to.

In addition it is interesting to see, how good a recovered potential describes the overall gravitational potential of the galaxy, i.e. its "predictive power". We introduce another diagnostic, that uses a cylindrical grid centered on the respective positions in Table [TO DO], having always a radius of $r_{\max} = 5 \text{ kpc}$ and a height of $z = 1.5 \text{ kpc}$ both above and below the plane. (We throw out grid points close to the galactic center with $R < 0.125 \text{ kpc}$, however.) In the (x, y) plane the regular grid points have a distance of 0.25 kpc and in z they have a distance of 0.125 kpc to better sample the thin disk.

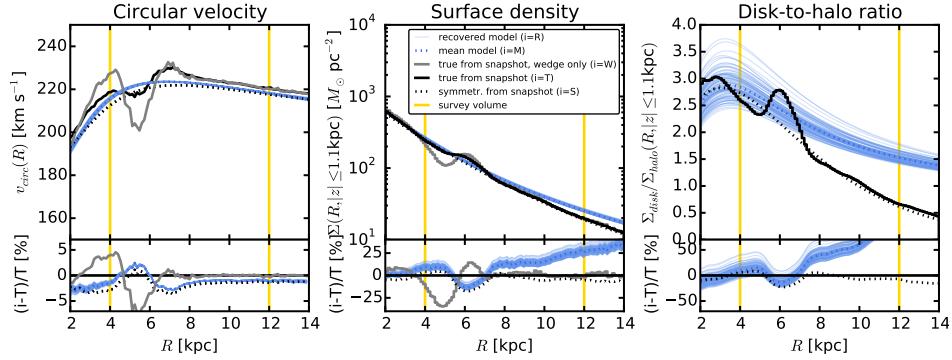


Figure 5. [TO DO: Remove MNHH symmetric potential from caption.] Comparison of the circular velocity curve, surface density profile within $|z| \leq 1.1$ kpc and disk-to-halo ratio of the surface density along R for the true potential of the galaxy simulation snapshot (solid black line) and the axisymmetric model potential recovered with *RoadMapping* (solid blue lines) (see Section [TO DO]). Overplotted are also the profiles of the symmetrized “true potential” (dotted black line) (see Section [TO DO] and Table 1) and the recovered mean model (dotted blue line) (see Table 2). The circular velocity curve is recovered to less than 5%, especially at larger radii. For the surface density and disk-to-halo ratio *RoadMapping* recovers the truth at radii $\lesssim 8$ kpc. The deviations at larger radii are connected to the discrepancies in the density in Figure [TO DO]. [TO DO: Add grey curve for disk to halo ratio.]

[comment: This was decided completely arbitrary. I guess we can think of a slightly different grid, that is more motivated.] We then evaluate at the position (x_j, y_j, z_j) of each regular grid point g_j the force residuals

$$\Delta F_R(g_j) \equiv \frac{|F_{R,M}(R_j, z_j)| - |F_{R,T}(x_j, y_j, z_j)|}{F_{R,\text{typ}}(R_j)} \quad (15)$$

$$\Delta F_z(g_j) \equiv \frac{|F_{z,M}(R_j, z_j)| - |F_{z,T}(x_j, y_j, z_j)|}{F_{z,\text{typ}}(R_j)}, \quad (16)$$

analogous to Equations (12)-(13). The two panels on the right in Figure 9 show the [15.87th, 84.13th] percentile range of the grid points’ distribution of $\Delta F(g_j)$.

The first and most important thing to notice in Figure 9 is, that we get very close to recovering the true forces at the positions of the stars in the survey volume, no matter how large the survey volume is.

Second, if we consider not the stars inside the survey volume, but a spatial average of the forces in a large volume ($r_{\text{max}} = 5$ kpc) independent of the survey volume, the radial forces are overall very well recovered, especially for large survey volumes. There is however a overestimation of $\sim 5 - 20\%$ in the vertical forces (depending on volume size and position) which is induced by the spiral arms (see explanation below).

Third, the constraints we get on the spatially averaged forces inside $r < 5$ kpc are almost as good for the survey volume of $r_{\text{max}} = 3$ kpc as compared to survey volumes of $r_{\text{max}} = 4$ or 5 kpc. If we had to decide between a $r_{\text{max}} = 3$ kpc volume with good data quality and a larger volume with worse data quality, we would loose nothing in terms of force recovery when using the smaller volume. (Only the halo scale length might not be as well constrained, see Figure 8).

Now let’s discuss the reason for the biases that we observe. The peak of the distribution in $\Delta F_R(*_i)$ and $\Delta F_R(g_j)$ is slightly biased towards an underestimation of $|F_{R,M}|$ in our *RoadMapping* models. We think the explanation is the following: Spiral arms are very thin. If a spiral arm crosses the observation volume both its leading side (at large radii) and its trailing side (at small radii) are also in the volume. Stars in the trailing side feel a lower gravitational pull towards the galaxy center

than they would if there was no spiral arm. Because there are in general more stars at smaller radii, the *RoadMapping* fit is slightly biased to reproduce in general slightly weaker radial forces (see also Figures 5 and 6).

The peak of $\Delta F_z(*_i)$ is approximately at 0, while the peak of $\Delta F_z(g_j)$ is strongly biased towards an overestimation of $|F_{z,M}|$. There are much more stars in the spiral arms than in the inter-arm regions, and the stars in the spiral arm feel stronger vertical forces because of the higher disk mass. *RoadMapping* finds a model that in general has much stronger vertical forces than expected for a smooth potential. While the actual vertical forces that the many stars in the spiral arms feel are very well recovered, it becomes obvious when looking at the grid points regularly distributed in space that the *RoadMapping* vertical forces are much stronger. As expected the overestimation is especially strong ($\sim 20\%$) for small survey volumes dominated by spiral arms, while small volumes dominated by an inter-arm region result in much better estimates for the spatially averaged $F_z(g_j)$ ($\sim 5\%$ bias). Large volumes lie somewhere in between (bias of $\sim 10\%$).

Because the stellar number asymmetry in the trailing vs. leading sides of spiral arms is much smaller than the stellar number asymmetry in spiral arm vs. inter-arm region, the bias is visible in the distribution of $\Delta F_{R,*_i}$ (F_R recovery biased only by a few stars \rightarrow bias for majority of stars visible) and not in $\Delta F_z(*_i)$ (majority of stars biases the fit \rightarrow we recover F_z for the majority of stars), but becomes really pronounced for $\Delta F_z g_j$ (the inter-arm regions dominate when averaging spatially \rightarrow large overestimation of F_z) and stays small for $\Delta F_R g_j$ (trailing and leading sides of spiral arms are similarly important when averaging spatially \rightarrow it becomes visible that the bias is actually not that big).

[TO DO: make consistent: i counts stars, j counts grid points.]

4.2.4. Influence of spiral arms on the recovery of the gravitational forces

The Figure 9 in Section 4.2.3 gave an overview over the accuracy of the gravitational forces recovery for our whole suite of data sets depending on the size of the different survey volumes. In the following we want to relate

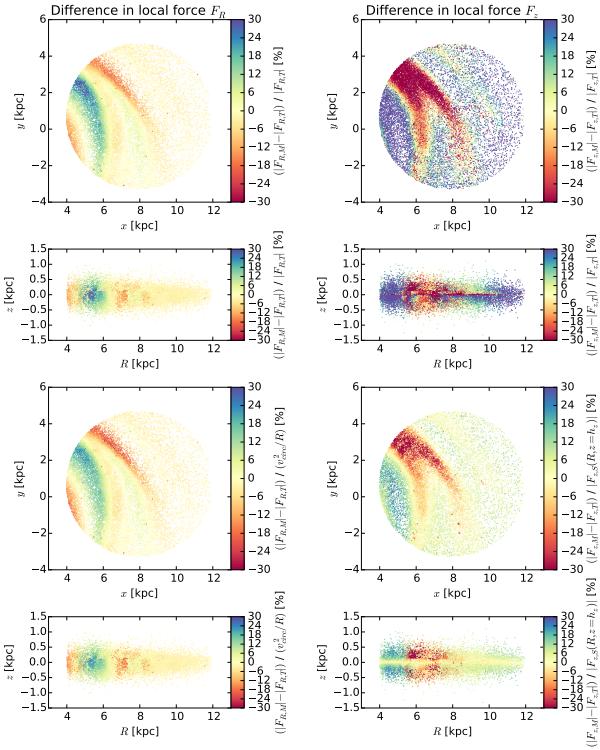


Figure 6. [TO DO: Rewrite. Figure was replaced.] Comparison of radial and vertical forces $F_R = -\partial\Phi/\partial R$ and $F_z = -\partial\Phi/\partial z$ of the true (upper two rows) or symmetrized (lower three rows, black and orange dotted lines) potential of the galaxy simulation snapshot with the potential model recovered with *RoadMapping* (solid blue lines in lower three rows). The upper two rows show the relative difference between the true potential force at the position of each star that entered the *RoadMapping* analysis, and the recovered potential force. The lower three rows compare contours of equal force and force profiles along lines of constant R and z . Overall the radial forces are very well recovered, which is related to the well recovery of the circular velocity curve in Figure ????. There are more problems with the vertical force, which is related to the higher surface densities in spiral arms and lower surface density in inter-arm regions, for which the axisymmetric model can only give a average solution. Also, the potential model is not the optimal model to describe the vertical density profile, which becomes also visible in the recovery of the vertical force close to the plane of the disk. [TO DO: Might be, that I have to re-calculate the forces for some of the stars. Write Test that tests if force $\gtrsim 10^{10}$ and then recalculates those forces.] [TO DO: Show only the lower panels.] TO DO: Redo plot. Now I use the median of MCMC samples as best fit model.]

the same quantities, i.e., local force recovery at the positions of the stars that entered the analysis (see Equations (12)-(13)) and power of a recovered model to predict the forces in a fixed large volume (see Equations (15)-(16)), to the actual strength and dominance of spiral arms in the respective survey volumes (see Section 2.4). In Figure 10 we find a connection between the local force recovery and the relative spiral contrast, and Figure 11 shows a trend that the "predictive power" of the recovered model depends on the dominance of a spiral arm in the survey volume. The former was to be expected, but the latter is a very interesting result.

Let's have a closer look at Figure 10, where we relate Equations (12)-(13) (see also Figure 9, left panels) to Equation [TO DO] (see also Figure 2, right panel). The average fraction of stars for which the recovery of radial or vertical force is less than 10% increases with

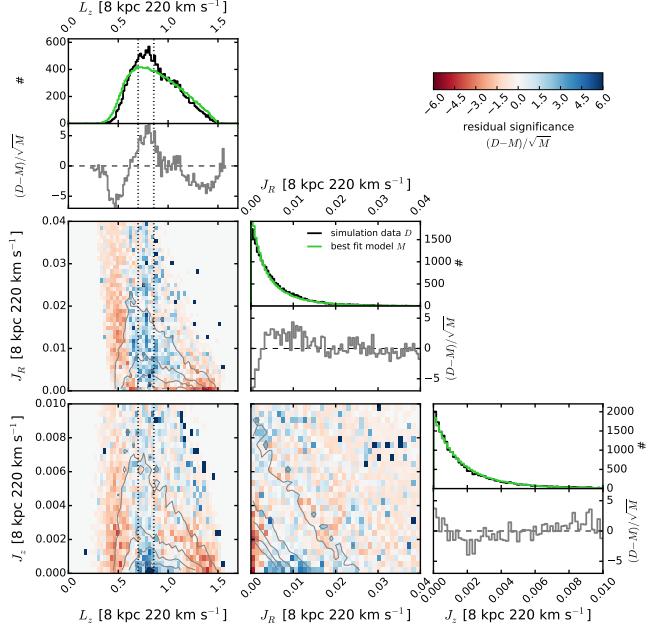


Figure 7. Comparison of the stellar action distribution of the data set D used in the analysis and the recovered axisymmetric stellar distribution M (see Figure 3 for the distribution in configuration space). All actions, of data set and best fit distribution, where calculated in the best fit MNHH-Pot in Table 2. The upper panel in each column compares the 1D histograms of the angular momentum L_z , the radial action J_R and the vertical action J_z of both data and best fit. The other 1D and 2D distributions show the residual significance $(D - M)/\sqrt{M}$. The 2D residuals are over-plotted with equidensity contours of the data D 's 2D action distribution. In Figure 3(a) we have marked the approximate radial extent of the stronger spiral arm with black dotted lines ($R_{\text{spiral}} \in [5.6, 6.8]$ kpc); the dotted lines in the L_z distributions in this figure correspond to $L_z = R_{\text{spiral}} \times v_{\text{circ}}(R_{\text{spiral}})$. This comparison gives a first impression of how the approximate action distribution in spiral arms might look like. An open task that the Galactic dynamical modelling community faces in the future, is to construct orbit distribution functions for spiral arms. [TO DO: Calculate reference data new. Now I use the median of MCMC samples as best fit model.] TO DO: Redo plot. Now I use the median of MCMC samples as best fit model.] [TO DO: Add in legend that simulation data was calculated in best fit potential too.] [TO DO: use the best fit and not the sym model to overplot the dotted lines]

decreasing spiral contrast $\sigma_\kappa \equiv \sigma[\kappa(r \leq r_{\max})]$ and is overall very similar (see the linear fits in Figure 10) for a given σ_κ , but there is scatter when considering individual analyses, which is stronger for the radial forces. The overall trend is however obvious and as expected: Volumes, in which the spatially steep surface density drop in the wings of a strong spiral arm is not balanced by larger areas with less perturbations ($\kappa_j \sim 0$), have both a large relative spiral contrast σ_κ and a large relative number of stars affected by the non-axisymmetric kinematics of the spiral arms. In this case there are less stars in the survey volume for which the axisymmetric *RoadMapping* model can properly predict the forces. We see again the trend, that volumes at $R_0 = 5$ kpc have a stronger spiral contrast and a worse forces recovery than the volumes centered at $R_0 = 8$ kpc. Volumes centered on inter-arm regions have, as expected, slightly lower σ_κ than volumes centered on spiral arms and therefore a better local forces recovery.

The result in Figure 11 is especially interesting. Here

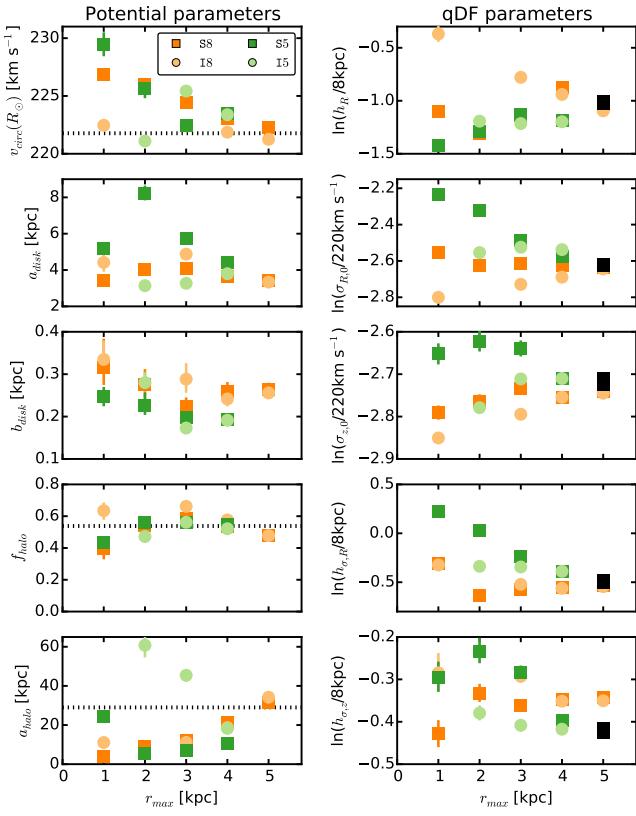


Figure 8. Overview over the model parameter estimates (MNHH-Pot parameters on the left, qDF parameters on the right) recovered with *RoadMapping* from data sets drawn from the simulation snapshot at different positions in the galaxy (colour-coded) and of different size (r_{\max} as indicated on the x -axis). The black dotted line shows the known model parameters from the DEHH-Pot in Table 1 (the Miyamoto-Nagai disk parameters a_{disk} and b_{disk} are related but not directly comparable to an exponential disk scale length and height). The black squares denote the qDF parameters we recovered by fixing the potential to the DEHH-Pot, centering a survey volume with $r_{\text{kpc}} = 5$ on the spiral arm at $R = 8$ kpc, and fitting the qDF only. A survey volume with a radial coverage as large as $r_{\max} = 5$ kpc is required to properly recover the "true" model parameters. [TO DO: Add legend to plot.] TO DO: Preliminary. Add more results. Make sure that the "true" potential parameters are the final version. [TO DO: Add additional y-axis on right side of plot with non-log values.]

we relate the volume averaged "predictive power" ((15)-(16) (see also Figure 9, right panels) to the dominance of the spiral arm in Equation [TO DO] (see also Figure 2, third panel). Firstly we note that $F_z(g_j)$ is recovered much worse than $F_R(g_j)$ and the reason is, as laid out in Section 4.2.3), that the reference cylinder contains some regions at high $|z|$ that is not probed by stars. The main result of this figure is the following: The less a volume is dominated by a spiral arm (smaller $\langle \kappa \rangle$), the better the model recovered from the corresponding volume can predict the forces in a large cylinder of $r_{\max} = 5$ kpc around the volumes midpoint. For large volumes where spiral arms and inter-arm regions average out ($\langle \kappa \rangle \sim 0$) this is true and as expected. For smaller volumes the picture is different: Small volumes centered on spiral arms ($\langle \kappa \rangle \gg 0$) contain a especially large fraction of stars in spiral arms that bias the potential estimates strongly. Small volumes centered on inter-arm regions ($\langle \kappa \rangle < 0$) seem to be smooth enough and close enough to the ax-

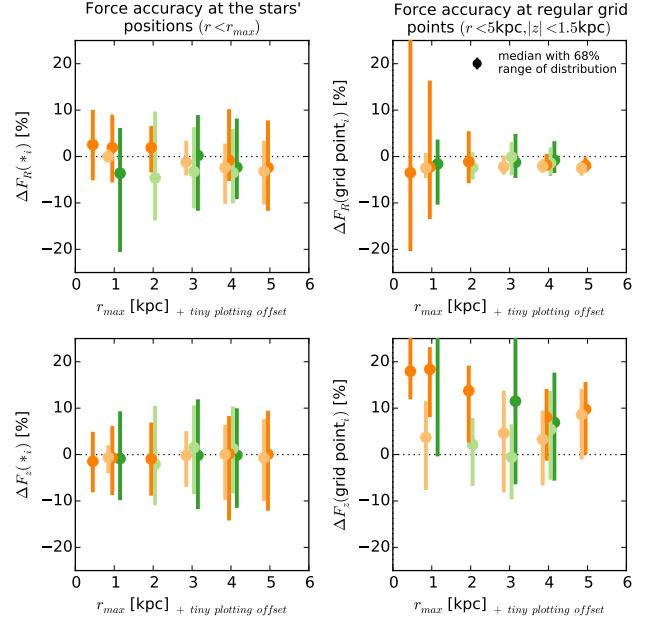


Figure 9. Accuracy of the radial (upper panels) and vertical (lower panels) gravitational forces recovered with *RoadMapping* from the suite of data sets introduced in Section ??, depending on the size r_{\max} of the survey volume. The left panels investigate how well *RoadMapping* recovers the forces for the ensemble of stars in the data set ($\Delta F_{R/z}(*_i)$ in Equations (12)-(13)) and the right panels test the "predictive power" of each of the data sets' best-fit potential by showing how well the forces are recovered at regular grid points in a large cylinder of $r_{\max} = 5$ kpc and height $|z| = 1.5$ kpc around the survey volumes center ($\Delta F_{R/z}(g_j)$ in Equations (15)-(16)). For each distribution of $\Delta F_{R/z}(*_i)$ and $\Delta F_{R/z}(g_j)$ we show here the median as dot with the [15.87th, 84.13th] percentile range as bar. We find that the forces are very well recovered at the positions of the stars independent of the size of the volume. The spiral arms introduce some biases in the overall recovered potential and we need at least a survey volume of $r_{\max} = 3$ kpc to get a potential with a reasonable "predictive power". [TO DO: Make sure all final analyses are in this plot.] [TO DO: use Delta F(gj) on right y-axis] [TO DO: Add legend to plot.] [TO DO: Jo suggested to make the right lower plot for slices at different heights, because it is very likely that most of the misjudgement is at high z. Alternatively we could use a different reference volume with that's not going up to 1.5kpc but maybe only up to 3-4 scale heights?]

isymmetric average model, that the potentials recovered from these volumes have real predictive power for much a much larger volume. This is an important result.

4.2.5. Circular velocity curves and surface density profiles

In this section we will illustrate even further that even though the recovered potential parameters are very different for different survey volumes, the actual potential parameter recovery is very good—and not just the gravitational forces on average. In Appendix [TO DO] in Figures 1-[TO DO] we show the circular velocity curves and surface density profiles within two times the disk scale height ($|z| \leq 2 \times h_z = 0.34$ kpc) for all data sets in this work and compare them with the actual matter distribution in the angular wedge $\phi_0 \pm \arcsin(r_{\max}/R_0)$ in which the corresponding survey volume lies. We also mark the approximate radial regions in which the stellar number density per radial bin $\Delta R = [TODO]$ is higher than average in each data set. Even though the curves vary extremely between the individual data sets, it becomes very obvious that it is indeed the regions in which

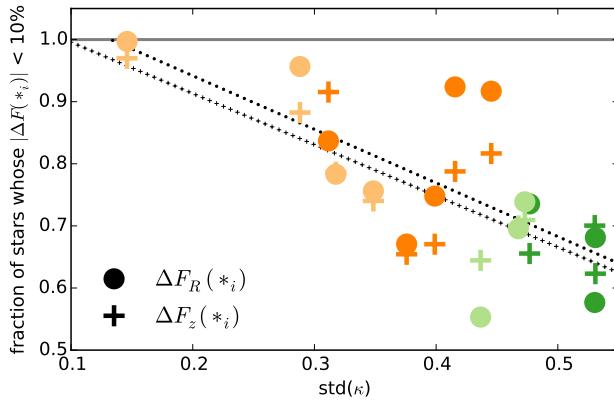


Figure 10. Influence of the spiral arm contrast within the survey volume on the recovery of the gravitational forces at the positions of the stars that entered the *RoadMapping* analysis. The relative spiral contrast on the x -axis was quantified as $\sigma[\kappa(r \leq r_{\max})]$ according to Equation [TO DO] in Section 2.4. On the y -axis the fraction of stars is shown for which the radial (circle symbols) and vertical (cross symbols) force residual calculated from Equations (12)-(13) is smaller than 10%. The dotted lines (small circles and crosses) are linear fits to the radial and vertical force residual fraction, respectively, and are guides to the eye that show the clear and expected trend that in volumes with smaller spiral arm contrast, where comparably fewer stars are located in spiral arms, the axisymmetric best-fit model can recover the true gravitational forces also for more stars. [TO DO: Make sure all final analyses are in this plot.] [TO DO: Legend with positions.] [TO DO: rewrite x-axis]

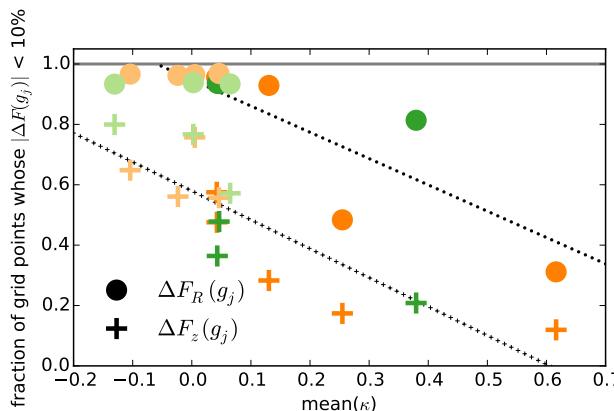


Figure 11. Influence of the spiral arm dominance within the survey volume on the "predictive power" of the models recovered from each survey volume. The spiral arm dominance on the x -axis was quantified as $\langle \kappa \rangle \equiv \langle \kappa(r \leq r_{\max}) \rangle$ from Equation [TO DO] in Section 2.4. The y -axis shows the fraction of a large and fixed cylindrical volume with $r_{\max} = 5$ kpc and $|z|1.5$ kpc centred around the respective survey volume positions given in Table [TO DO], i.e., number of regular grid points with radial (circle symbols) and vertical (cross symbols) force residuals smaller than 10%, i.e., how well the model can predict the potential in a large portion of the galaxy. The dotted lines (small circles and crosses) are guides to the eyes. This figure demonstrates that the "predictive power" of a model derived from a survey volume that is either large enough for spiral arms not to dominate the volume ($\langle \kappa \rangle \sim 0$, compare Figure 2), or that are derived from volumes centered on inter-arm regions ($\langle \kappa \rangle \lesssim 0$) are less biased by the kinematics in the spiral arms and have therefore a higher predictive power. [TO DO: Make sure all final analyses are in this plot.] [TO DO: Legend with positions.] [TO DO: rewrite x-axis]

the majority of stars is located that drives the *RoadMapping* fit. And that, no matter if this region is dominated by a spiral arm or inter-arm region and no matter if this region is only as small as $r_{\max} = 500$ kpc, *RoadMapping* is indeed constraining the local potential where most of the stars of the data set are located. Also, the constraints are not only most accurate but also most precise in these regions.

Only in the two volumes with $r_{\max} = [0.5, 1]$ kpc at position S8 *RoadMapping* has some difficulties fitting the circular velocity curve; the model expects a flat or falling rotation curve and is presented with a steeply rising rotation curve due to the spiral arm dominating the region. But given *RoadMapping* recovers a good average surface density profile and the circular velocity at least at the center of the small volume, the fit is still quite successful.

Given the sometimes extremely strong variations in circular velocity and surface density, and given the limitations of the smooth parametrized potential model, *RoadMapping* is very successful at recovering the local potential.

5. DISCUSSION

5.1. On the informativeness of an orbit distribution function

The qDF is surprisingly informative; even though there was absolutely no indication beforehand that the qDF would be actually a good model for the simulation snapshot; and even though small survey volumes dominated by non-axisymmetric spiral arms are definitely not well-described by the qDF. It appears however that a potential model that does not fit the gravitational forces acting on the stars must lead to such an unrealistic orbit/action distribution, that a fit with even such a simple orbit DF as the the qDF is impossible. This demonstrates once more how powerful the concept of an orbit DF is.

5.2. On the restrictiveness of the parametrized potential model

We do, on the one hand, use a bulge and halo model that reproduce the true bulge and halo better than we can hope to use one in reality for the MW. The fact that for all except of the largest volumes the true halo scale length is not remotely recovered (and some small volumes even have $f_{\text{halo}} = 0$), and that the contribution of the bulge to the overall potential is very small ([TO DO]) remedies this apparent advantage.

We use on the other hand a disk model, the Miyamoto-Nagai disk, that we chose purely for its convenient parametric form and of which we know that it is not a good model for the simulation snapshot, especially not for the radial density profile. As we saw in most figures in this work this might lead to biases in predicting the potential at radii where we have only a few or no stars. But because the spiral arms are such strong perturbations in the overall potential, a better disk model would not give much better results.

It appears that a potential model with a reasonable shape and flexibility (here: disk+bulge+halo structure with 5 free parameters) can do well enough in finding a good fit, both locally for small volumes and overall for large volumes.

5.3. Recovering non-axisymmetric structures from the potential from small volumes

[TO DO]

5.4. Precision of Gaia measurement errors

Considering measurement uncertainties of distances and proper motions, we found in Paper I that for a survey volume with $r_{\max} = 3$ kpc, distance uncertainties of $\pm 10\%$ and proper motion uncertainties of [TO DO] RoadMapping still gives unbiased parameter results. Even if the proper motion errors are not perfectly known.

At a distance of 3 kpc the measurement uncertainties of bright stars in Gaia lie even below these limits.

In paper I we focussed on recovering completely unbiased model parameters and found that RoadMapping is robust to moderate deviations of the model assumptions. In this work we released the condition that the model parameters itself had to be recovered accurately, but allowed RoadMapping to simply find an overall best fit for the data strongly affected by spiral arms—which was surprisingly successful in recovering the local potential even if the model parameters were not recovered.

We therefore presume that in reality we probably have an even larger margin of error before the measurement uncertainties mess up the constraints noticeably than we found in Paper I. And Gaia uncertainties are already below our formal error.

In addition, we found in this work that a volume of $r_{\max} = 3$ kpc should be already big enough to find an overall best fit axisymmetric model for the galaxy. At larger radii dust starts affecting the measurements.

Overall we should therefore be very well off by applying RoadMapping to Gaia data within $r_{\max} = 3$ kpc only.

5.5. Location of the Sun

[TO DO]

6. CONCLUSION

- [TO DO]

In Section ?? and Figure 8 we found that the potential and qDF parameters recovered from different survey volumes are widely different. In Sections 4.2.3 and 4.2.4 we found that the gravitational forces are however very well recovered on the average at the positions of the stars. The reason for this and possible biases were

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APPENDIX

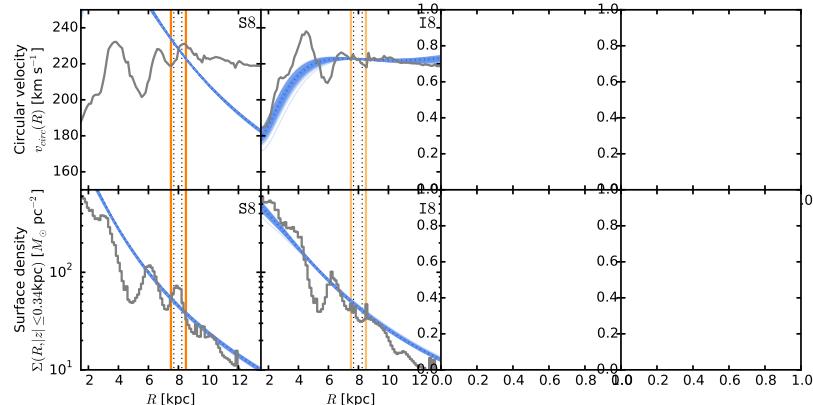


Figure 1. TO DO [TO DO: Make sure all final analyses are in this plot.]

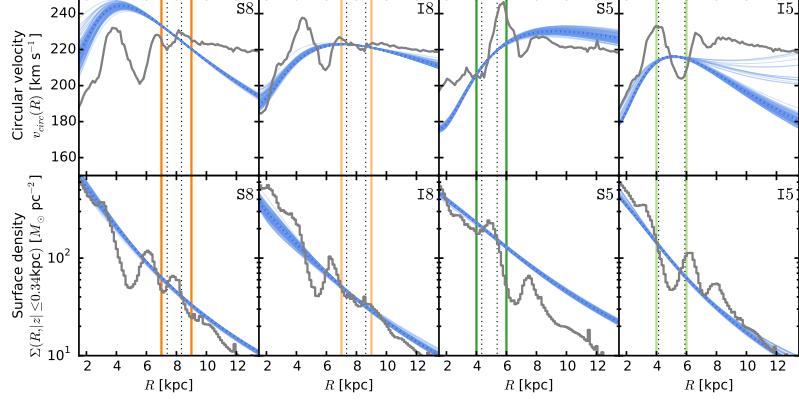


Figure 2. Same as Figure 1, but for all survey volumes with $r_{\max} = 1$ kpc. [TO DO: Make sure all final analyses are in this plot.]

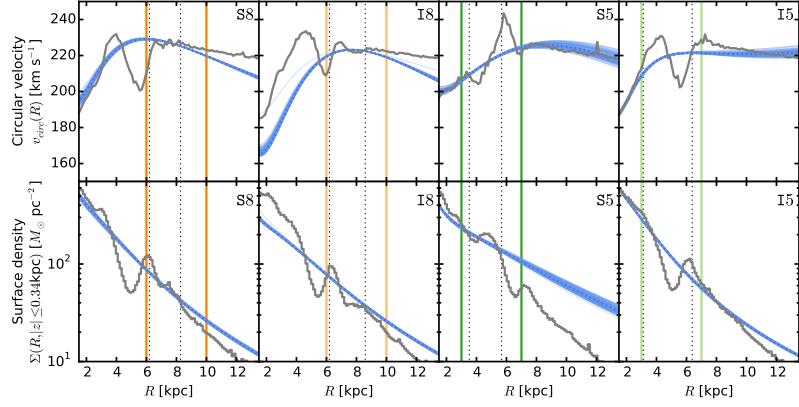


Figure 3. Same as Figure 1, but for all survey volumes with $r_{\max} = 2$ kpc. [TO DO: Make sure all final analyses are in this plot.]

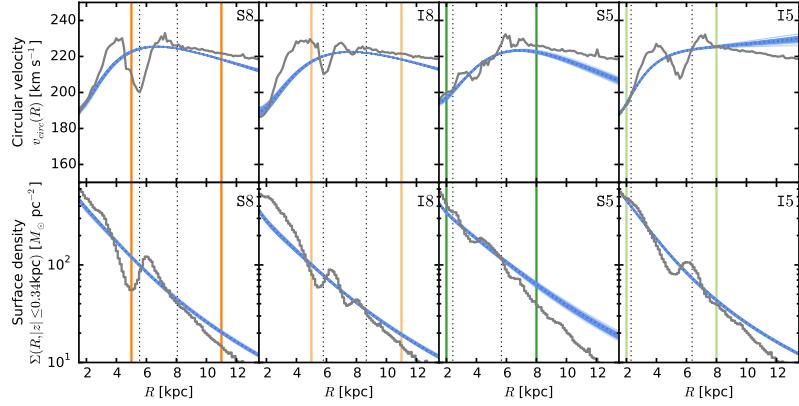


Figure 4. Same as Figure 1, but for all survey volumes with $r_{\max} = 3$ kpc. [TO DO: Make sure all final analyses are in this plot.]